

Euclid Quick Data Release (Q1)

A first view of the star-forming main sequence in the Euclid Deep Fields

Euclid Collaboration: A. Enia^{1,2}, L. Pozzetti², M. Bolzonella², L. Bisigello³, W. G. Hartley⁴, C. Saulder^{5,6}, E. Daddi⁷, M. Siudek^{8,9}, G. Zamorani², P. Cassata^{10,3}, F. Gentile^{11,2}, L. Wang^{12,13}, G. Rodighiero^{10,3}, V. Allevato¹⁴, P. Corcho-Caballero¹³, H. Domínguez Sánchez¹⁵, C. Tortora¹⁴, M. Baes¹⁶, Abdurro'uf¹⁷, A. Nersesian^{18,16}, L. Spinoglio¹⁹, J. Schaye²⁰, Y. Ascasibar^{21,22}, D. Scott²³, E. Duran-Camacho^{24,8}, S. Quai^{25,2}, M. Talia^{25,2}, Z. Mao², N. Aghanim²⁶, B. Altieri²⁷, A. Amara²⁸, S. Andreon²⁹, N. Auricchio², H. Aussel⁷, C. Baccigalupi^{30,31,32,33}, M. Baldi^{1,2,34}, A. Balestra³, S. Bardelli², A. Basset³⁵, P. Battaglia², R. Bender^{5,6}, A. Biviano^{31,30}, A. Bonchi³⁶, E. Branchini^{37,38,29}, M. Brescia^{39,14}, J. Brinchmann^{40,41}, S. Camera^{42,43,44}, G. Cañas-Herrera^{45,46,20}, V. Capobianco⁴⁴, C. Carbone⁴⁷, J. Carretero^{48,49}, S. Casas⁵⁰, F. J. Castander^{9,51}, M. Castellano⁵², G. Castignani², S. Cavuoti^{14,53}, K. C. Chambers⁵⁴, A. Cimatti⁵⁵, C. Colodro-Conde²⁴, G. Congedo⁵⁶, C. J. Conselice⁵⁷, L. Conversi^{58,27}, Y. Copin⁵⁹, F. Courbin^{60,61}, H. M. Courtois⁶², M. Cropper⁶³, A. Da Silva^{64,65}, H. Degaudenzi⁴, G. De Lucia³¹, A. M. Di Giorgio¹⁹, C. Dolding⁶³, H. Dole²⁶, F. Dubath⁴, C. A. J. Duncan⁵⁷, X. Dupac²⁷, S. Dusini⁶⁶, A. Ealet⁵⁹, S. Escoffier⁶⁷, M. Fabricius^{5,6}, M. Farina¹⁹, R. Farinelli², F. Faustini^{52,36}, S. Ferriol⁵⁹, F. Finelli^{2,68}, S. Fotopoulou⁶⁹, M. Frailis³¹, E. Franceschi², P. Franzetti⁴⁷, S. Galeotta³¹, K. George⁶, B. Gillis⁵⁶, C. Giocoli^{2,34}, P. Gómez-Alvarez^{70,27}, J. Gracia-Carpio⁵, B. R. Granett²⁹, A. Grazian³, F. Grupp^{5,6}, L. Guzzo^{71,29,72}, S. Gwyn⁷³, S. V. H. Haugan⁷⁴, J. Hoar²⁷, W. Holmes⁷⁵, I. M. Hook⁷⁶, F. Hormuth⁷⁷, A. Hornstrup^{78,79}, P. Hudelot⁸⁰, K. Jahnke⁸¹, M. Jhabvala⁸², B. Joachimi⁸³, E. Keihänen⁸⁴, S. Kermiche⁶⁷, A. Kiessling⁷⁵, B. Kubik⁵⁹, M. Kümmel⁶, M. Kunz⁸⁵, H. Kurki-Suonio^{86,87}, Q. Le Boulc'h⁸⁸, A. M. C. Le Brun⁸⁹, D. Le Mignant⁹⁰, S. Ligi⁴⁴, P. B. Lilje⁷⁴, V. Lindholm^{86,87}, I. Lloro⁹¹, G. Mainetti⁸⁸, D. Maino^{71,47,72}, E. Maiorano², O. Mansutti³¹, O. Marggraf⁹², M. Martinelli^{52,93}, N. Martinet⁹⁰, F. Marulli^{25,2,34}, R. Massey⁹⁴, D. C. Masters⁹⁵, S. Maurogordato⁹⁶, E. Medinaceli², S. Mei^{97,98}, M. Melchior⁹⁹, Y. Mellier^{100,80}, M. Meneghetti^{2,34}, E. Merlin⁵², G. Meylan¹⁰¹, A. Mora¹⁰², M. Moresco^{25,2}, L. Moscardini^{25,2,34}, R. Nakajima⁹², C. Neissner^{103,49}, S.-M. Niemi⁴⁵, J. W. Nightingale¹⁰⁴, C. Padilla¹⁰³, S. Paltani⁴, F. Pasian³¹, K. Pedersen¹⁰⁵, W. J. Percival^{106,107,108}, V. Pettorino⁴⁵, S. Pires⁷, G. Polenta³⁶, M. Poncet³⁵, L. A. Popa¹⁰⁹, F. Raison⁵, R. Rebolo^{110,111,24}, A. Renzi^{10,66}, J. Rhodes⁷⁵, G. Riccio¹⁴, E. Romelli³¹, M. Roncarelli², E. Rossetti¹, B. Rusholme¹¹², R. Saglia^{6,5}, Z. Sakr^{113,114,115}, A. G. Sánchez⁵, D. Sapone¹¹⁶, B. Sartoris^{6,31}, J. A. Schewtschenko⁵⁶, M. Schirmer⁸¹, P. Schneider⁹², T. Schrabback¹¹⁷, M. Scodeggio⁴⁷, A. Secroun⁶⁷, G. Seidel⁸¹, S. Serrano^{51,118,9}, P. Simon⁹², C. Sirignano^{10,66}, G. Sirri³⁴, J. Skottfelt¹¹⁹, L. Stanco⁶⁶, J. Steinwagner⁵, C. Surace⁹⁰, P. Tallada-Crespi^{48,49}, A. N. Taylor⁵⁶, H. I. Teplitz⁹⁵, I. Tereno^{64,120}, S. Toft^{121,122}, R. Toledo-Moreo¹²³, F. Torradeflot^{49,48}, I. Tutusaus¹¹⁴, L. Valenziano^{2,68}, J. Valiviita^{86,87}, T. Vassallo^{6,31}, G. Verdoes Kleijn¹³, A. Veropalumbo^{29,38,37}, Y. Wang⁹⁵, J. Weller^{6,5}, A. Zacchei^{31,30}, F. M. Zerbi²⁹, I. A. Zinchenko⁶, E. Zucca², M. Ballardini^{124,125,2}, E. Bozzo⁴, C. Burigana^{126,68}, R. Cabanac¹¹⁴, A. Cappi^{2,96}, D. Di Ferdinando³⁴, J. A. Escartin Vigo⁵, L. Gabarra¹²⁷, M. Huertas-Company^{24,8,128,129}, J. Martín-Fleitas¹⁰², S. Matthew⁵⁶, N. Mauri^{55,34}, R. B. Metcalf^{25,2}, A. Pezzotta⁵, M. Pöntinen⁸⁶, C. Porciani⁹², I. Risso¹³⁰, V. Scottez^{100,131}, M. Sereno^{2,34}, M. Tenti³⁴, M. Viel^{30,31,33,32,132}, M. Wiesmann⁷⁴, Y. Akrami^{133,134}, I. T. Andika^{135,136}, S. Anselmi^{66,10,137}, M. Archidiacono^{71,72}, F. Atrio-Barandela¹³⁸, C. Benoist⁹⁶, K. Benson⁶³, P. Bergamini^{71,2}, D. Bertacca^{10,3,66}, M. Bethermin¹³⁹, A. Blanchard¹¹⁴, L. Blot^{140,137}, H. Böhringer^{5,141,142}, S. Borgani^{143,30,31,32,132}, A. S. Borlaff^{144,145}, M. L. Brown⁵⁷, S. Bruton¹⁴⁶, A. Calabro⁵², B. Camacho Quevedo^{51,9}, F. Caro⁵², C. S. Carvalho¹²⁰, T. Castro^{31,32,30,132}, F. Cogato^{25,2}, T. Contini¹¹⁴, A. R. Cooray¹⁴⁷, O. Cucciati², S. Davini³⁸, F. De Paolis^{148,149,150}, G. Desprez¹³, A. Díaz-Sánchez¹⁵¹, J. J. Diaz²⁴, S. Di Domizio^{37,38}, J. M. Diego¹⁵, P.-A. Duc¹³⁹, Y. Fang⁶, A. G. Ferrari³⁴, P. G. Ferreira¹²⁷, A. Finoguenov⁸⁶, A. Fontana⁵², F. Fontanot^{31,30}, A. Franco^{149,148,150}, K. Ganga⁹⁷, J. García-Bellido¹³³, T. Gasparetto³¹, V. Gautard¹¹, E. Gaztanaga^{9,51,152}, F. Giacomini³⁴, F. Gianotti², G. Gozaliasl^{153,86}, M. Guidi^{1,2}, C. M. Gutierrez¹⁵⁴, A. Hall⁵⁶, S. Hemmati¹¹², C. Hernández-Monteagudo^{111,24}, H. Hildebrandt¹⁵⁵, J. Hjorth¹⁰⁵, J. J. E. Kajava^{156,157}, Y. Kang⁴, V. Kansal^{158,159}, D. Karagiannis^{124,160}, K. Kiiveri⁸⁴, C. C. Kirkpatrick⁸⁴

S. Kruk²⁷, J. Le Graet⁶⁷, L. Legrand^{161, 162}, M. Lembo^{124, 125}, F. Lepori¹⁶³, G. Leroy^{164, 94}, G. F. Lesci^{25, 2}, J. Lesgourgues⁵⁰, L. Leuzzi^{25, 2}, T. I. Liaudat¹⁶⁵, A. Loureiro^{166, 167}, J. Macias-Perez¹⁶⁸, G. Maggio³¹, M. Magliocchetti¹⁹, E. A. Magnier⁵⁴, C. Mancini⁴⁷, F. Mannucci¹⁶⁹, R. Maoli^{170, 52}, C. J. A. P. Martins^{171, 40}, L. Maurin²⁶, M. Miluzio^{27, 172}, P. Monaco^{143, 31, 32, 30}, C. Moretti^{33, 132, 31, 30, 32}, G. Morgante², C. Murray⁹⁷, K. Naidoo¹⁵², A. Navarro-Alsina⁹², S. Nesseris¹³³, F. Passalacqua^{10, 66}, K. Paterson⁸¹, L. Patrizii³⁴, A. Pisani^{67, 173}, D. Potter¹⁶³, M. Radovich³, P.-F. Rocci²⁶, S. Sacquegna^{148, 149, 150}, M. Sahlén¹⁷⁴, D. B. Sanders⁵⁴, E. Sarpa^{33, 132, 32}, C. Scarlata¹⁷⁵, A. Schneider¹⁶³, M. Schultheis⁹⁶, D. Sciotti^{52, 93}, E. Sellentin^{176, 20}, F. Shankar¹⁷⁷, L. C. Smith¹⁷⁸, S. A. Stanford¹⁷⁹, K. Tanidis¹²⁷, G. Testera³⁸, R. Teyssier¹⁷³, S. Tosi^{37, 130}, A. Troja^{10, 66}, M. Tucci⁴, C. Valieri³⁴, A. Venhola¹⁸⁰, D. Vergani², G. Verza¹⁸¹, P. Vielzeuf⁶⁷, N. A. Walton¹⁷⁸, and E. Soubrie²⁶

(Affiliations can be found after the references)

May 14, 2025

ABSTRACT

The star-forming main sequence (SFMS) is a tight relation observed between stellar masses and star formation rates (SFR) in a population of galaxies. This relation is observed at different redshifts, in various morphological, and environmental domains, and is key to understanding the underlying relations between a galaxy budget of cold gas and its stellar content. *Euclid* Quick Data Release 1 (Q1) gives us the opportunity to investigate this fundamental relation in galaxy formation and evolution. We complement the *Euclid* release with public IRAC observations of the *Euclid* Deep Fields, improving the quality of recovered photometric redshifts, stellar masses, and SFRs, as is shown both with simulations and a comparison with available spectroscopic redshifts. From Q1 data alone, we recover more than ~ 30 k galaxies with $\log_{10}(M_*/M_\odot) > 11$, giving a precise constraint of the SFMS at the high-mass end. We investigated the SFMS, in a redshift interval between 0.2 and 3.0, comparing our results with the existing literature and fitting them with a parameterisation taking into account the presence of a bending of the relation at the high-mass end, depending on the bending mass, M_0 . We find good agreement with previous results in terms of M_0 values, and an increasing trend for the relation scatter at higher stellar masses. We also investigate the distribution of physical (e.g. dust absorption, A_V , and formation age) and morphological properties (e.g., Sérsic index and radius) in the SFR–stellar mass plane, and their relation with the SFMS. These results highlight the potential of *Euclid* in studying the fundamental scaling relations that regulate galaxy formation and evolution in anticipation of the forthcoming Data Release 1.

Key words. Galaxies: evolution; Galaxies: formation; Galaxies: fundamental parameters; Galaxies: statistics

1. Introduction

The star-forming main-sequence (SFMS) is a relation between stellar masses (M_*) and star-formation rates (SFRs) that is observed for star-forming galaxies (SFGs). It has been extensively studied in the last few decades (Brinchmann et al. 2004; Daddi et al. 2007; Elbaz et al. 2011), with investigations into its slope, normalisation, scatter, and evolution over time (see Speagle et al. 2014; Popesso et al. 2023, and references therein). The SFMS is observed across different redshifts and is already in place by $z \sim 6$ (e.g., Clarke et al. 2024; Cole et al. 2025). It hosts the majority of star formation at each epoch (Rodighiero et al. 2011), suggesting that galaxies spend most of their lifetimes on the SFMS, undergoing secular evolution. The tightness of the relation, with a typical scatter between 0.2 dex and 0.4 dex, implies its universality as the main mode of galaxy growth (Schreiber et al. 2015; Donnari et al. 2019; Matthee & Schaye 2019).

This relation emerges from the interplay between the stellar content of galaxies and their cold gas reservoirs (i.e., the so-called Kennicutt–Schmidt relation from Schmidt 1959 and Kennicutt 1998b, and the (resolved or integrated) molecular gas main-sequence, see e.g. Lin et al. 2019; Morselli et al. 2020; Ellison et al. 2021; Baker et al. 2023), and it has been shown to hold at sub-kiloparsec scales too (e.g., Wuyts et al. 2013; Hsieh et al. 2017; Lin et al. 2017; Abdurro’uf & Akiyama 2017; Ellison et al. 2018; Enia et al. 2020; Baker et al. 2022).

At masses of $\log_{10}(M_*/M_\odot) > 10$ at $0 < z < 1$ and at the high-mass end $\log_{10}(M_*/M_\odot) > 11$ at $z \sim 2$, the relation appears to exhibit a deviation from the linear trend, the so-called bending

of the SFMS (Whitaker et al. 2014; Schreiber et al. 2015; Tomczak et al. 2016; Popesso et al. 2019; Leja et al. 2022; Daddi et al. 2022a; Leroy et al. 2024; Wang et al. 2024). The bending traces changes in cold-gas accretion (Kereš et al. 2005; Dekel & Birnboim 2006) and availability for star formation processes, and could be a consequence of the reduced availability of cold gas in halos entering the hot accretion mode phase (Daddi et al. 2022a), or feedback from active galactic nuclei (Fabian 2012), or both (Bower et al. 2017). Additionally, the reactivation of star formation in the disks of galaxies that are approaching quiescence or have already been quenched may also contribute to the bending of the SFMS (Mancini et al. 2019). This turnover mass can be linked with the host halo mass quenching threshold (Yang et al. 2007; Behroozi et al. 2019; Leslie et al. 2020; Popesso et al. 2023), defining the transition between an environment favourable to star formation to a regime where these processes are suppressed.

The *Euclid* Quick Release Q1 (2025) is the first release of *Euclid* survey data, corresponding to a single reference observing sequence (ROS, see *Euclid* Collaboration: Scaramella et al. 2022) of the *Euclid* Deep Fields (EDFs). This is a homogeneous view of a large area of the extragalactic sky (~ 63 deg²) from optical to near-infrared (NIR), complemented with observations of the same fields at $3.6 \mu\text{m}$ and $4.5 \mu\text{m}$ with the Infrared Array Camera (IRAC, Fazio et al. 2004) on *Spitzer* (Werner et al. 2004). These fields have the potential to become the best-studied extragalactic fields of the coming decades. In this work, we illustrate the results obtained with the data and products of Q1 for the SFMS, investigating its evolution up to $z = 3$, and the distribution of physical and morphological parameters along the SFMS,

* e-mail: andrea.enia@unibo.it

as well as validating these results with the existing literature, a first demonstration of the potential of *Euclid* to investigate scaling relations and the baryon cycle.

This paper is structured as follows. In Sect. 2, we describe the data released for Q1. In Sect. 3, we describe the methods used to recover the photometric redshifts and physical parameters (PPs). In the same section, we validate the results, reporting the performance of our methods on simulations and the available sub-sample of spectroscopic redshifts and $H\alpha$ -estimated SFRs. In Sect. 4, we report the results for the SFMS. In Sect. 5, we present our conclusions and perspectives for the upcoming Data Release 1.

Throughout this paper we adopt a flat Lambda cold dark matter (Λ CDM) cosmology with $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_m = 0.3$, and $\Omega_\Lambda = 0.7$, and assume a [Chabrier \(2003\)](#) initial mass function (IMF). All magnitudes are given in the AB photometric system ([Oke & Gunn 1983](#)).

2. Data

A detailed description of the Q1 data release is presented in [Euclid Collaboration: Aussel et al. \(2025\)](#), [Euclid Collaboration: McCracken et al. \(2025\)](#), [Euclid Collaboration: Polenta et al. \(2025\)](#), and [Euclid Collaboration: Romelli et al. \(2025\)](#). A summary of the scientific objectives of the mission can be found in [Euclid Collaboration: Mellier et al. \(2025\)](#). In short, for Q1 *Euclid* observed $\sim 63 \text{ deg}^2$ of the extragalactic sky, divided into EDF-North (EDF-N), EDF-Fornax (EDF-F), and EDF-South (EDF-S), in four photometric bands, one in the visible (I_E , [Euclid Collaboration: Cropper et al. \(2025\)](#)), and three in the NIR (Y_E , J_E , and H_E , [Euclid Collaboration: Jahnke et al. \(2025\)](#)). These observations are complemented by ground-based observations carried out with multiple instruments to cover the wavelength range between $0.3 \mu\text{m}$ and $1.8 \mu\text{m}$ by the Ultraviolet Near-Infrared Optical Northern Survey (UNIONS, [Gwyn et al. \(2025\)](#)) and the Dark Energy Survey (DES, [Flaugher et al. \(2015\)](#); [Dark Energy Survey Collaboration et al. \(2016\)](#)).

In order to obtain robust results and improve the quality of the recovered photometric redshifts and PPs (see Sect. 3.3), we also added to the *Euclid* photometry two available IRAC bands, at $3.6 \mu\text{m}$ and $4.5 \mu\text{m}$, covering all the EDFs ([Euclid Collaboration: Moneti et al. \(2022\)](#); [Euclid Collaboration: McPartland et al. \(2025\)](#)). More details on how IRAC photometry is measured can be found in [Euclid Collaboration: Bisigello et al. \(2025\)](#).

In Table 1 we report the filters used in this work, with the observed 10σ depths for an extended source in an aperture that is twice the full width at half maximum (FWHM, i.e., the one with the lowest resolution between the optical and *Euclid* bands, see [Euclid Collaboration: Romelli et al. \(2025\)](#), for further details). For this work, we started from the available *Euclid* catalogues and applied a series of selections to make our analysis more robust, removing compact or low-quality sources. These selections are:

- SPURIOUS_FLAG = 0;
- DET_QUALITY_FLAG < 4;
- MUMAX_MINUS_MAG > -2.6.

For further details on the meaning of the flags, see [Euclid Collaboration: Tucci et al. \(2025\)](#). This selection skims the sample from stars and compact objects such as quasi-stellar objects (QSOs).

We further cleaned our sample from these two classes of objects using the classification probability of the Q1 data products, imposing the following criteria:

Table 1. Filters used in this work, with associated observed depths.

Band	λ_{eff} [μm]	EDF-N	EDF-F	EDF-S
$u_{\text{CFHT/MegaCam}}$	0.372	23.39		
g_{HSC}	0.480	24.87		
$r_{\text{CFHT/MegaCam}}$	0.640	24.01		
$i_{\text{PAN-STARRS}}$	0.755	23.07		
z_{HSC}	0.891	23.35		
g_{Decam}	0.473		24.65	24.72
r_{Decam}	0.642		24.33	24.37
i_{Decam}	0.784		23.76	23.78
z_{Decam}	0.926		23.06	23.12
VIS/ I_E	0.715	24.75	24.70	24.74
NISP/ Y_E	1.085	23.16	23.10	23.15
NISP/ J_E	1.375	23.31	23.24	23.30
NISP/ H_E	1.773	23.24	23.19	23.24
IRAC1	3.550	24.05	24.05	23.15
IRAC2	4.493	23.95	23.95	23.05

Notes. Reported magnitudes are the 10σ observed median depths of the observing tiles for an extended source in a $2\times\text{FWHM}$ diameter aperture. For IRAC values see [Euclid Collaboration: Moneti et al. \(2022\)](#) and [Euclid Collaboration: McPartland et al. \(2025\)](#).

- PROB_QSO < 0.86;
- PROB_STAR < 0.10.

See Sect. 4 of [Euclid Collaboration: Tucci et al. \(2025\)](#) for further details, and also [Euclid Collaboration: Matamoro Zatarain et al. \(2025\)](#) about the classification thresholds.

Finally, we benefited from the results of the morphological analysis for Q1 ([Euclid Collaboration: Walmsley et al. \(2025\)](#); [Euclid Collaboration: Quilley et al. \(2025\)](#)), applying another set of cuts related to the morphological parameters and the size of the source. We kept sources with:

- $q > 0.05$;
- $0.01a < R_e < 2a$.

q is the Sérsic axis ratio, R_e the Sérsic radius, and a the isophote semi-major axis, in units of VIS pixels. For further information see Sect. 4 and the left panel of Fig. 3 in [Euclid Collaboration: Quilley et al. \(2025\)](#). These further remove diffraction spikes, cosmic rays, or stars that survived the cuts described above.

The last cut that we applied was in magnitude, in order to work with a mass-complete sample, limiting our analysis to sources with observed $H_E < 24$, corresponding to an average measured signal-to-noise ratio of five, measured from the $2\times\text{FWHM}$ aperture photometry. Our final sample is composed of 8 090 074 sources.

3. Physical properties and validation

We refer the reader to [Euclid Collaboration: Tucci et al. \(2025\)](#) for a description of how the Q1 data were processed to infer photometric redshifts and PPs. In this section, we briefly summarise the procedure.

Due to the large number of sources detected (of the order of tens of millions) in the EDFs, machine learning (ML) methods have been developed to speed up the computational process while achieving a comparable performance of template fitting methods (see e.g., [Euclid Collaboration: Desprez et al. \(2020\)](#); [Euclid Collaboration: Enia et al. \(2024\)](#)). Data products produced for Q1 have been obtained with a nearest-neighbours (NNs) algorithm, nnpz, which finds a k number of NNs (30 in the *Euclid*

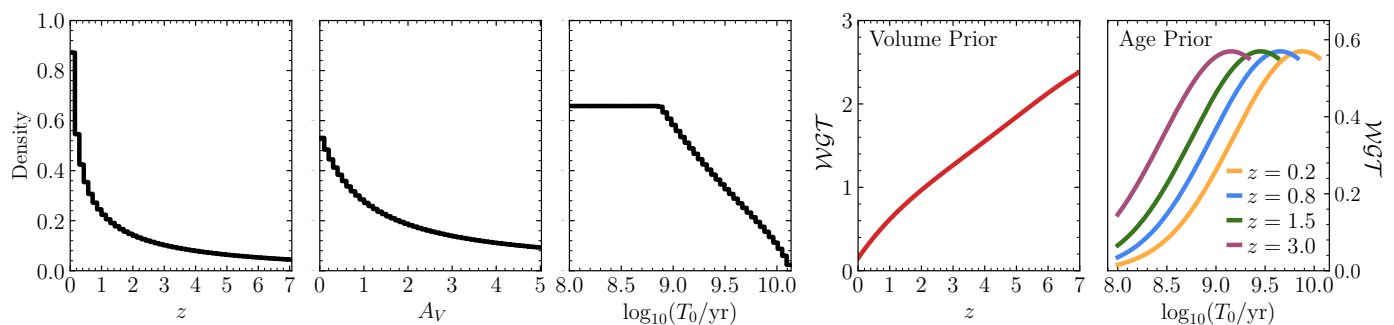


Fig. 1. *Left panels:* Distribution of redshifts, z , attenuations, A_V , and ages, T_0 for the reference sample used in this work. *Right panels:* Two priors adopted in this work – volume and age – as multiplicative weights on the reference galaxies, depending on their redshifts or ages.

pipeline, 80 for this work) in the template space (i.e., magnitude and colour) for each target galaxy from a reference sample, and infers the photometric redshifts and PPs from those.

The reference sample differs from the one used to produce the Q1 data products, the reason for which is explained below. It was built from a grid of templates (Bruzual & Charlot 2003, 2016 version¹), using the MILES stellar library, which adopts the Kroupa (2001) IMF, which we therefore converted to Chabrier (2003) taking into account a difference of about 0.03 dex for PPs. The models were built with exponentially delayed star formation histories:

$$\text{SFR}(t) \propto (t - T_0) e^{-(t-T_0)/\tau}, \quad (1)$$

drawn from a Halton (1964) grid² in a six-dimensional space with the following free parameters:

- redshifts: $0 < z < 7$;
- ages: $8 < \log_{10}(T_0/\text{yr}) < 10.138$;
- e-folding timescale: $8 < \log_{10}(\tau/\text{yr}) < 10.5$;
- ionisation parameter: $-4 < \log_{10}(U) < -2$;
- metallicities: $0.1 < Z/Z_{\odot} < 2$.

For dust attenuation, we generated models with both Calzetti et al. (2000) and SMC (Prevot et al. 1984) laws, with V -band attenuation, A_V , between 0 and 5. Stellar masses and SFRs were inferred from the amplitude of the observed spectral energy distribution (SED), as a scaling parameter recovered by nnpz. Metallicities and ionisation parameters are distributed uniformly within the given ranges; the same for τ , but on a logarithmic scale. The redshifts are distributed with a linear scaling in $(1+z)$ steps, while ages and A_V scale logarithmically from the lowest value to the highest. For these last three parameters, their distribution in the reference sample is shown in black in the first three panels of Fig. 1. We then generated the noise-free observed-frame photometry associated with each model. In the end, the reference sample consists of 1 490 150 objects.

nnpz then finds the 80 NNs for every target galaxy, based on the observed magnitudes and colours. Each of these NNs will have its own weight – measured from the χ^2 distance between the reference and the target – and scaling parameter from the SED amplitude. By combining them, we measure the median (or the mode) of the distribution of NNs, which ultimately are the inferred photo- z s, PPs, and absolute magnitudes of the target galaxies.

¹ http://www.bruzual.org/bc03/Updated_version_2016/

² A method to generate a quasi-random grid, which is evenly distributed across the parameter space, minimising the presence of gaps and clusters of points.

Unlike what has been done in the *Euclid* pipeline, we added two more photometric points to the reference sample, accounting for the IRAC1 and IRAC2 channels (see Sect. 2). Moreover, having access to the nnpz results – that is, the set of NNs for each target galaxy – we can impose whatever physically motivated condition (i.e. a prior) directly on the NNs, either by measuring the output, z_{phot} or PPs, only on the set of NNs that satisfy the condition, or by differently re-weighting the NNs in the reference sample in order to penalise unphysical solutions.

As is reported in Sect. 6.1.1 of *Euclid Collaboration: Tucci et al. (2025)*, the Q1 pipeline results (obtained without the application of any condition to the NNs) contain an artificially high number of low- z galaxies with extremely young ages – $\log_{10}(T_0/\text{yr})$ starts from 7 in the pipeline – observed at the peak of their star-forming activity, and thus at the limit of specific star-formation rate (sSFR) inherent to parametric models ($\log_{10} \text{sSFR}/\text{yr}^{-1} \simeq -7.8$, see Fig. 8 of Ciesla et al. 2017). The resulting redshift distribution is ultimately skewed towards non-physically high number counts at low z , creating an artefact straight line at the sSFR saturation limit (see Fig. 14 in *Euclid Collaboration: Tucci et al. 2025*). In principle, this issue could be addressed in the Q1 pipeline data product by imposing an age prior on the NNs; for example, ignoring those with ages < 0.1 Gyr. This would significantly reduce the impact of artefacts, but at the cost of reducing the Q1 sample by about 40%, excluding from the sample all sources without even a single NN that satisfies the prior. Losing these sources would introduce systematic biases into our analysis. To avoid this, we took some precautions that deviate from what has been done in the pipeline, in order to reduce the impact of those artefact sources without losing a significant fraction of the sample: the boundaries for ages and A_V mentioned above used to generate the reference sample for this work are different from those reported in *Euclid Collaboration: Tucci et al. (2025)*, with $0 < A_V < 3$ and $7 < \log_{10}(T_0/\text{yr}) < 10.138$. We also increased the k number of NNs to 80. Finally, we imposed both a volume and an age prior on the NNs in the reference sample.

With the volume prior, we increased the weights of NNs at higher redshifts compared to the ones at lower redshifts, where the volume of the Universe sampled by the survey is smaller and fewer galaxies are expected to be observed. This prior was implemented as a multiplicative weight assigned to each object in the reference sample, and only depends on the redshift as

$$\text{WGT}(z) \propto \frac{dV_c(z)}{dz}, \quad (2)$$

where $dV_c(z)$ is the comoving volume shell in the interval $[z, z + dz]$. This prior is shown in red in the centre-right panel of Fig. 1.

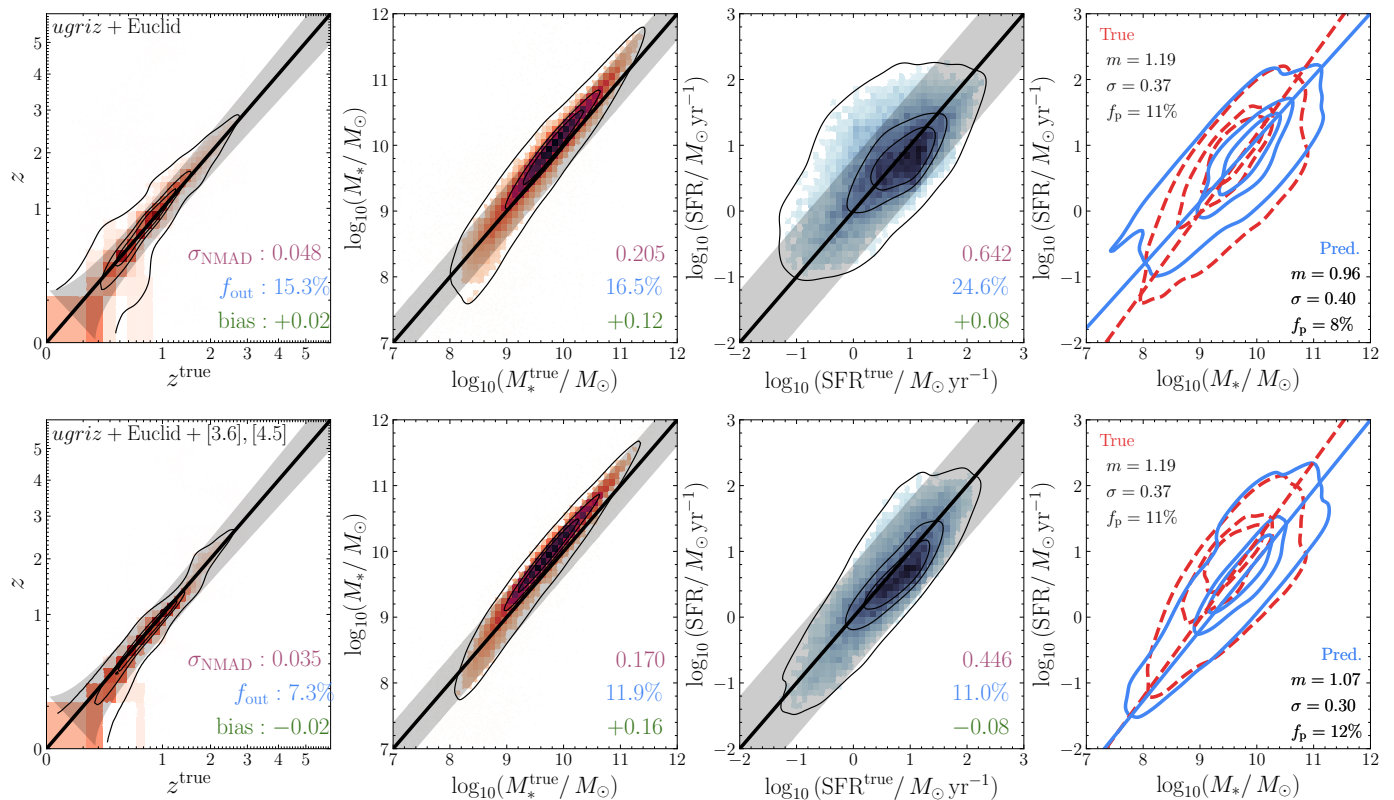


Fig. 2. Results for the *nnpz* run on the 70k simulated FS2 galaxies, without (*top panels*) and with the first two WISE filters (*bottom panels*). The black line is the 1:1 relation; the shaded area is the region beyond which a prediction is an outlier. In every panel, the contours are the area containing 86%, 39% (corresponding to the 2σ and 1σ levels for a 2D histogram), and 20% of the sample. For SFMS (right panel), the true distribution is reported in red (dashed) and the predicted one in blue (solid). The lines are the orthogonal distance regression (ODR) best fit to the distribution with passive galaxies removed. The reported metrics are NMAD (purple), the outlier fraction, f_{out} (blue), and the bias (green) for the photometric redshifts and physical parameters, and the slope, m , scatter, σ , and fraction of passive galaxies, f_p , for the SFMS, all defined in Sect. 3.1.

The age prior is once again a multiplicative weight, $\mathcal{W}_{GT}(T_0)$, to apply to NNs. This takes into account the fact that younger galaxies could be observed at higher redshifts, while this possibility should be reduced at low redshifts. In building the prior, we looked at the distribution of ages in different redshift bins in the 2 deg^2 of the Cosmic Evolution Survey (COSMOS, Weaver et al. 2022), finding that these can be modelled as normal distributions with the peak age decreasing while moving at a higher redshift. This weight is then constructed as a truncated normal distribution centred on two thirds the age of the Universe at any given z , with width 0.7 in $\log_{10}(T_0/\text{yr})$. A weight equal to zero is assigned for those sources with an age greater than the age of the Universe at the given z . The shape of the age prior is reported in the rightmost panel of Fig. 1 for four indicative redshift values.

This procedure makes us sensitive to the bulk of the population of SFGs, while reducing our ability to properly identify and describe outliers (e.g., the starburst galaxies, as the presence of a star-forming burst is not directly accounted for in the reference sample). Given the main scopes of this work, this is acceptable, since it has been found that the exponentially delayed τ model is an accurate description of the star formation histories of the main population of SFMS galaxies, at least at $z < 2$ (see e.g., Speagle et al. 2014; Ciesla et al. 2017). Recently, studies focussing on non-parametric models of star formation histories have found how these models could better recover the complex events that arise during the evolution of a galaxy (Iyer et al. 2019; Leja et al. 2019; Baes et al. 2020). In the future, for Data Release 1, mod-

els will better account for starburst galaxies, with the possibility of exploring other regions of the parameter space, even accounting for complex star formation histories (see for example Euclid Collaboration: Corcho-Caballero et al. 2025).

We validated our results both on the available state-of-the-art simulations adapted to reproduce as closely as possible *Euclid* observations (i.e., the Flagship2 simulation, FS2, see Euclid Collaboration: Castander et al. 2025), and on a compilation of the available spectroscopic redshifts and $H\alpha$ measured SFRs in the EDFs. Although the latter is fundamental to assess the quality of recovered parameters with what can be interpreted as the closest possible thing to a ‘ground truth’ value, the former is necessary to put a degree of confidence in all the other recovered quantities that have no such ‘ground truth’ to compare with.

3.1. Metrics for quality assessment

The metrics used to quantify the quality of the results are defined differently when referring to redshifts or PPs. We refer the reader to Euclid Collaboration: Enia et al. (2024) for a full discussion of thresholds and catastrophic outlier definitions; here, we simply report the definitions.

We first defined a set of true values, z_{test} and y_{test} (on a logarithmic scale for PPs), to confront with the predicted values, z_{pred} and y_{pred} . We then defined the normalised median absolute

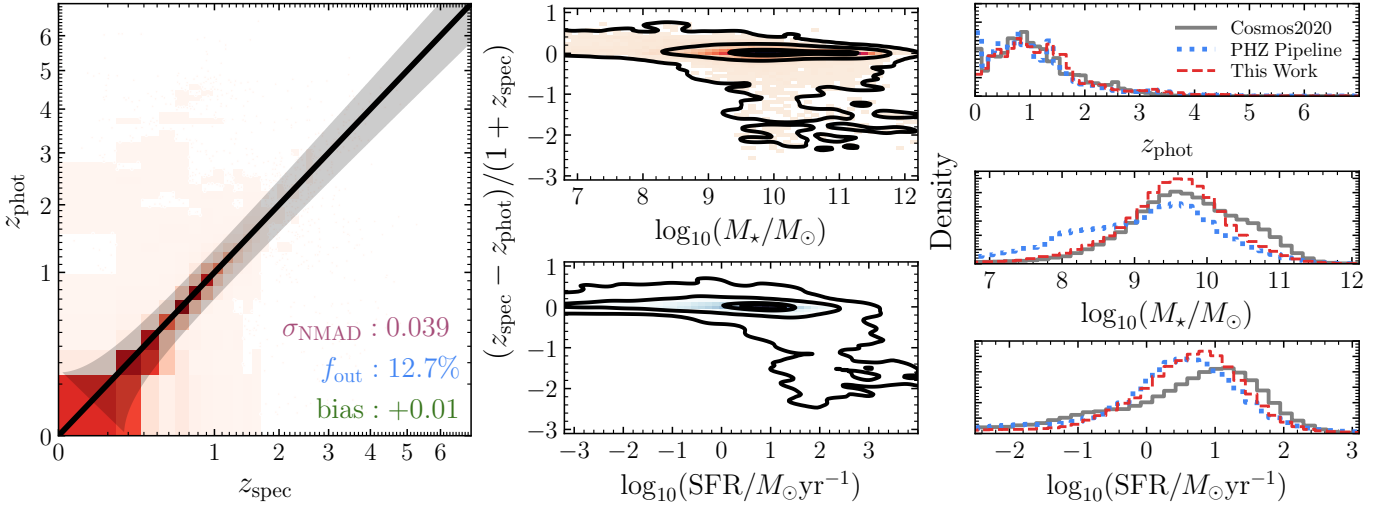


Fig. 3. *Left:* Comparison between the measured photometric redshifts and a compilation of all the available and reliable spectroscopic redshifts, colour-coded by the density of objects in each bin. The black line is the 1:1 relation, while the shaded grey area is the region beyond which a prediction is considered to be an outlier. The reported metrics are NMAD (purple), the outlier fraction, f_{out} (blue), and the bias (green), all defined in Sect. 3.1. *Middle:* Normalised redshift difference as a function of stellar masses (top) and SFRs (bottom). The black contours are at the same levels as in Fig. 2. *Right:* Normalised distributions of photometric redshifts (top), stellar masses (centre), and SFRs (bottom) for our full sample (dashed red lines), compared with the results coming from the Q1 data products (PHZ, in dotted blue lines), and COSMOS2020 (in solid grey lines), at the same magnitude cuts applied in this work (see Sect. 2).

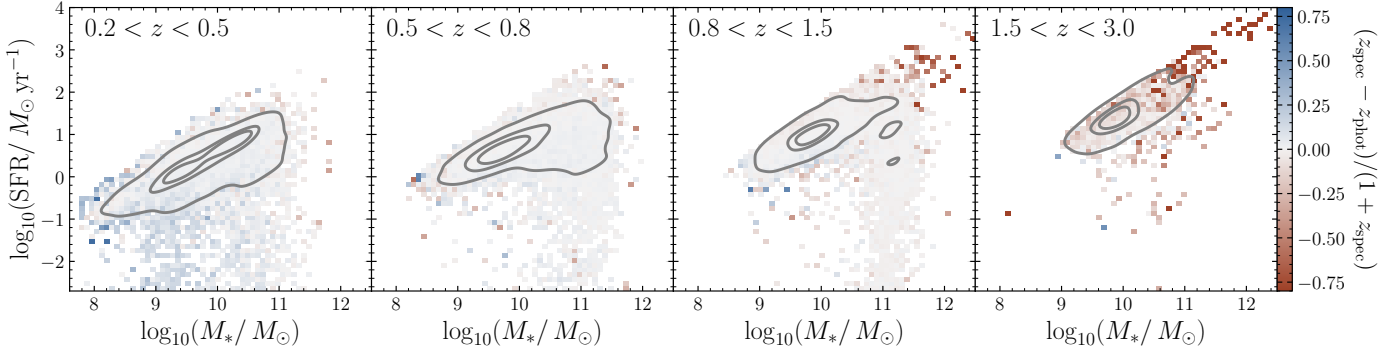


Fig. 4. M_* -SFR plane at four different redshift intervals, only for the sources with a reliable z_{spec} described in Sect. 3.3, colour-coded by the median normalised redshift difference in each bin. To give an idea of the density of objects in each redshift interval, we superimposed the contours in grey, with the same levels as in Fig. 2.

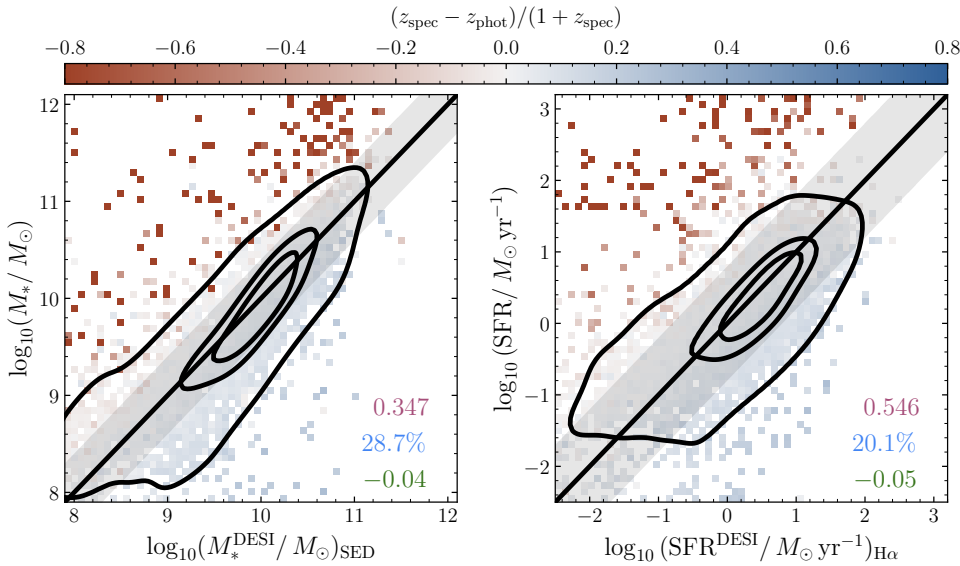


Fig. 5. Comparison between the stellar masses and SFRs measured in this work and a sample of objects from DESI with stellar masses from SED fitting (*left*) and SFRs from $H\alpha$ (*right*), colour-coded as a function of the median normalised redshift difference in each bin. The black line is the 1:1 relation, while the shaded area is the region beyond which a prediction is an outlier. The contours are the same as in Fig. 2. The reported metrics are defined in Sect. 3.1.

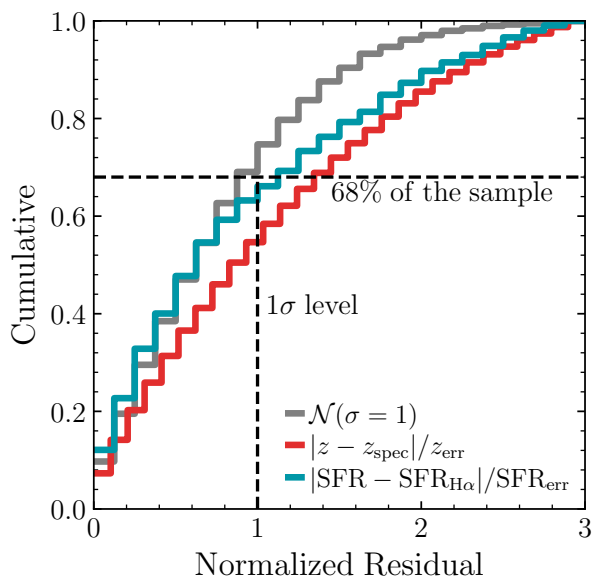


Fig. 6. Cumulative distribution of the normalised residuals for photometric redshifts (red) and SFRs (blue). For comparison, we also show in grey the cumulative of a normal distribution with $\sigma = 1$. Dashed black lines highlight the position where 68% of the distributions are, and the 1σ level for the residuals.

deviation as

$$\text{NMAD} = 1.48 \times \text{median} \begin{cases} \frac{|z_{\text{pred}} - z_{\text{test}}|}{1 + z_{\text{test}}} - b, & \text{for redshifts,} \\ |y_{\text{pred}} - y_{\text{test}}| - b, & \text{for PPs,} \end{cases} \quad (3)$$

with b being the model bias (see below).

Then, we defined the outlier fraction,

$$f_{\text{out}} = \begin{cases} \frac{|z_{\text{pred}} - z_{\text{test}}|}{1 + z_{\text{test}}} > 0.15, & \text{for redshifts,} \\ |y_{\text{pred}} - y_{\text{test}}| > t_{\text{out}}, & \text{for PPs,} \end{cases} \quad (4)$$

with $t_{\text{out}} = 0.4$ for stellar masses and $t_{\text{out}} = 0.8$ for SFRs.

Finally, we defined the bias,

$$b = \text{median} \begin{cases} \left(\frac{z_{\text{pred}} - z_{\text{test}}}{1 + z_{\text{test}}} \right), & \text{for redshifts,} \\ (y_{\text{pred}} - y_{\text{test}}), & \text{for PPs.} \end{cases} \quad (5)$$

3.2. Validation on simulations

We randomly selected about 70 k sources from a complete octant of the FS2 simulation (Euclid Collaboration: Castander et al. 2025). We checked that the selection does not skew the ground-truth values with respect to the full distribution. These mock sources are, by construction, distributed in the redshift range $0 < z < 3$, with $6.4 < \log_{10}(M_*/M_{\odot}) < 12$ and $-4 < \log_{10}(\text{SFR}/M_{\odot} \text{ yr}^{-1}) < 3.1$. We then followed the same procedure described in Sect. 2 of Euclid Collaboration: Enia et al. (2024) to produce a realistic catalogue of mock sources that reproduces as closely as possible what is observed in Q1. We perturbed the intrinsic fluxes with the noise level found in Q1 in the same set of filters (see Table 1), and cut at observed magnitudes $H_E < 24$. Instead of IRAC, in the FS2 simulation these wavelength ranges are covered by mock observations with WISE bands, W1 and W2 (Wright et al. 2010). Given the similar range

covered, we ensured that the resulting performance is comparable if not exactly the same.

We then ran `nnpz` with the same reference sample described in Sect. 3, one run with the WISE filters and one without, and produced the recovered results as the median of the 80 NNs. The results are shown in Fig. 2, where we report the recovered versus true relation without (top panel) and with (bottom panel) the two WISE filters. These results follow closely what has already been described in Euclid Collaboration: Enia et al. (2024), at least at the order-of-magnitude level. It is immediately noticeable how the addition of the two filters at $3.6 \mu\text{m}$ and $4.5 \mu\text{m}$ improves parameter estimation, with photo- z NMADs and the outlier fraction decreasing from 0.048 to 0.035 and from 15% to 7%, respectively. The same holds for stellar masses (NMAD decreasing from 0.206 to 0.170 and the outlier fraction from 17% to 12%) and especially for SFRs (NMAD decreasing from 0.643 to 0.446 and the outlier fraction from 25% to 11%). As was expected, the recovery of SFMS improves with respect to the case without the two filters in NIR, with a much better recovery of the slope and normalisation of the SFMS relation.

3.3. Photometric redshifts and PPs validation

Quality assessment for the redshifts was performed by looking at how they compare with respect to the observed spectroscopic ones. The EDFs cover regions of the sky where there is a plethora of coverage from other spectroscopic surveys. In total, we successfully matched 63 504 galaxies with a reliable spectroscopic redshift; that is, with a redshift quality flag of 3 or 4 (see description in Sect. 5.2 of Euclid Collaboration: Tucci et al. 2025). These z_{spec} values are from: the Dark Energy Spectroscopic Instrument (DESI, DESI Collaboration et al. 2016, 2024); the 16th Data Release of the Sloan Digital Sky Survey (SDSS, Ahumada et al. 2020); the 2MASS Redshift Survey (2MRS, Huchra et al. 2012); the PRISM Multi-object Survey (PRIMUS, Coil et al. 2011); the Australian Dark Energy Survey (OzDES, Yuan et al. 2015; Childress et al. 2017; Lidman et al. 2020); 3dHST (Brammer et al. 2012); the 2-degree Field Galaxy Redshift Survey (2dFGRS, Colless et al. 2001); the 6-degree Field Galaxy Redshift Survey (6dFGS, Jones et al. 2009); the MOSFIRE Deep Evolution Field Survey (MOSDEF, Kriek et al. 2015); the VANDEL ESO public spectroscopic survey (Pentericci et al. 2018; Talia et al. 2023); the JWST Advanced Deep Extragalactic Survey DR3 (JADES, D'Eugenio et al. 2025); the 2-degree Field Lensing Survey (2dFLens, Blake et al. 2016); and the VIMOS VLT deep survey (VVDS, Le Fèvre et al. 2005).

The results are reported in Fig. 3. In the left panel, we compare the photometric and spectroscopic redshifts for the subset of reliable spectroscopic redshifts in our sample. These are almost equally divided between EDF-F and EDF-N, with only a select number (123) of objects in EDF-S. The trend we find using *Euclid* real data is similar to what is shown in the previous section, with a non-negligible improvement when adding the two IRAC bands. To put it in quantitative terms, compared to the same analysis performed without the two IRAC bands (not shown in Fig. 3) the NMAD decreases from 0.06 to 0.04 and the fraction of outliers decreases from 26% to 13%. In the central panels, we show the normalised difference between spectroscopic and photometric redshifts, as a function of the stellar masses and the SFRs. We do not observe any troubling systematic trend, with the exception of the (expected) behaviour in which galaxies mistakenly placed at higher redshift – mostly those with $(z_{\text{spec}} - z_{\text{phot}})/(1 + z_{\text{spec}}) > -1$ – are found with an incorrect higher SFRs, so some caution must be taken when dealing with those

high- z galaxies, or with $\log_{10}(\text{SFR}/M_{\odot} \text{ yr}^{-1}) > 2.4$. In terms of observed bias, we find a negligible value of +0.01, well below the typical uncertainties associated with photo- z s (see below).

When looking at the full sample of objects – not just spectroscopic ones – there are no ground-truth values to compare with, but we can still investigate how our results agree with the full distribution of photometric redshifts, stellar masses, and SFRs, especially when compared with other surveys. This is done in the right panels in Fig. 3, where we compare our results for the full sample (the dashed red lines) with what is observed in the LePhare run from the classic catalog of COSMOS 2020 (the solid grey lines), and with the inferred PPs from the Q1 data products (PHZ, the dotted blue lines), both at the same magnitude cut applied in this work (i.e. $H_E < 24$). The first two distributions in redshift are comparable, with the main differences observed in a slightly lower fraction of $0.6 < z < 1.2$ objects in our sample (and conversely, a few more $z > 3$ galaxies), while the Q1 data products exhibit a significantly greater number of $z < 0.2$ objects. As for the stellar masses, our results improve with respect to the almost flat at $8 < \log_{10}(M_*/M_{\odot}) < 9$ distributions of Q1 data products; however, we find fewer objects in the $10.2 < \log_{10}(M_*/M_{\odot}) < 11.2$ range and conversely more in the $9.2 < \log_{10}(M_*/M_{\odot}) < 10.2$ range with respect to COSMOS. The biggest differences are observed in terms of SFRs, with more galaxies in our sample in the $0 < \log_{10}(\text{SFR}/M_{\odot} \text{ yr}^{-1}) < 1$ range with respect to COSMOS, and conversely slightly fewer galaxies per SFR bin in our sample at $1 < \log_{10}(\text{SFR}/M_{\odot} \text{ yr}^{-1}) < 2$.

In Fig. 4, we report all this information in the M_* –SFR plane, where the SFMS is observed. For the sub-sample of sources with a reliable spectroscopic redshift, we show the median value of the normalised redshift difference in each bin, with red colours highlighting sources mistakenly placed at a higher redshift, and blue colours the opposite. While the latter catastrophic outliers are a small issue only visible in the first redshift interval (with $0.2 < z < 0.5$), the former becomes more and more prominent at higher redshifts (i.e., $z > 1.5$), introducing a non-negligible bias in the estimates of M_* and SFR in the highest mass and SFR regimes, with these skewed towards higher values due to the wrong redshift attribution.

We compare the measured PPs with the ones in the value-added catalogue of PPs in DESI (Siudek et al. 2024), although for a limited number of sources (~ 5 k) in the interval $0 < z < 0.6$. In particular, we confront the SFRs with the ones obtained from both $H\alpha$ and $H\beta$ line measurement – not exactly ‘ground truth’ values, but close – while the stellar masses are compared to their SED fitting results, even though, given how it has been performed on fewer and shallower filters (see Table 1 in Siudek et al. 2024), we expect our stellar masses to be more reliable. The IMF is the same as the one adopted for this work (Chabrier 2003), and SFRs were obtained from $H\alpha$ following Kennicutt (1998a). The results are shown in Fig. 5, where the result for each PP is colour-coded as a function of the median value of $(z_{\text{spec}} - z_{\text{phot}})/(1 + z_{\text{spec}})$ in each bin. Despite the limited sample, the performance is in line with what is expected from the simulations. It is immediately clear how the vast majority of catastrophic outliers (i.e. where PPs fall outside of the defined thresholds) are a consequence of sources with an incorrect photometric redshift estimate. Again, the observed bias in stellar masses and SFR is still below the typical uncertainties associated with each PPs.

Finally, we used these SFRs from $H\alpha$ and the reliable spectroscopic redshift sample to place some constraints on the estimated uncertainties in the photo- z s and SFRs. Uncertainties were

measured from the 16th and 84th percentiles of the weighted distribution of NNs (see Sect. 3). To account for the possibility of an under- (or over-) estimation of those uncertainties, we looked at the cumulative distribution of $(z_{\text{phot}} - z_{\text{spec}})/z_{\text{err}}$ – and similarly $(\text{SFR} - \text{SFR}_{H\alpha})/\text{SFR}_{\text{err}}$ – where we expect 68% of these to fall below 1 if the uncertainties are well estimated. In contrast, an underestimate would lead to fewer sources within the 1σ limit, and the opposite would be true for an overestimate. These cumulative distributions (‘normalised residuals’, red for redshifts and blue for SFRs) are reported in Fig. 6. We find that our uncertainties are slightly underestimated for redshifts (1σ level reached for 55% of the sample) and almost spot on for SFRs. We estimate the underestimation of the uncertainties of the photometric redshifts to be a factor of about 1.5.

We take the mode of the distribution of uncertainties as the typical values for each parameter, which are 0.05 for redshifts, 0.11 for stellar masses, and 0.08 for SFRs (on a logarithmic scale). The typical uncertainty on SFRs are used in the following section for the fit of the SFMS.

4. Results

We started from the sample described in Sect. 2, and limited our analysis to the redshift range between 0.2 and 3.0. This is motivated by the need to obtain a statistically reliable sample in terms of mass completeness and the quality of the recovered photometric redshifts and physical properties. In Euclid Collaboration: Enia et al. (2024), it is shown how the main source of biases in the analysis of SFMS during cosmic time – apart from the inherent dispersions in determining the correct PPs – arises from the photo- z estimation, whereby typically some low- z (i.e. $z < 0.5$) objects are placed at high z (up to around 10% of catastrophic outliers) with increased stellar masses and SFRs. The net effect is a steepening of the SFMS at lower redshifts, and the opposite at higher z (see Fig. 11 of Euclid Collaboration: Enia et al. 2024). This is also observed in the validation tests that we performed with simulations and reported in Sect. 3.3 (see Fig. 4).

We performed our subsequent analysis in the following redshift bins: [0.2, 0.5], [0.5, 0.8], [0.8, 1.5], [1.5, 3.0], with the exception of the morphological analysis, where we stopped at $z = 1.5$, since for higher redshifts the quality of the recovered morphological parameters is limited by the sizes of the sources reaching the resolution limit of the survey. Based on the results shown in Fig. 4, we can place a certain degree of confidence in the highest-mass end of the first two redshift intervals, while greater caution is required for the last two. In each redshift bin, we limited our analysis – and the reported values – to $\log_{10}(M_*/M_{\odot}) < 11.5$.

4.1. Mass completeness

We estimated the mass completeness of the sample limited to $H_E < 24$ following the method described in Pozzetti et al. (2010). We selected galaxies close to the limiting magnitude of our sample; that is, those with $23 < H_E < 24$ (with identical results of choosing the faintest 20% as in Pozzetti et al. 2010), and measured for each galaxy

$$\log_{10}(M_{\text{lim}}/M_{\odot}) = \log_{10}(M_*/M_{\odot}) - 0.4(H_E - 24), \quad (6)$$

representing the mass the galaxies would have at the limiting magnitude. We then measured the 95th percentile of the distribution of $M_{\text{lim}}(z)$ for each redshift bin. This is reported in Fig. 7.

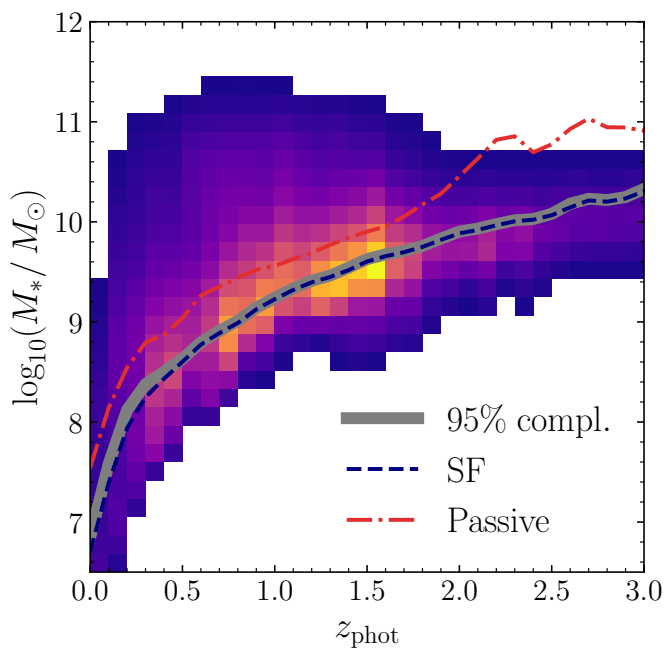


Fig. 7. The stellar mass completeness for the sample used in this work (Pozzetti et al. 2010), with a cut at $H_E < 24$, shown as the density of galaxies in the photometric redshifts vs. stellar masses plane. The solid grey line is the 95% stellar mass completeness limit, red dashed-dotted line for passive galaxies, and blue dashed line for SFGs.

The classification into star-forming (dashed blue line) and passive galaxies (dash-dotted red line) was done with the selection criteria based on the $NUV-r^+-J$ diagram (as is explained below).

The sample is around 95% complete for M_* , which increases from $\log_{10}(M_*/M_\odot) \sim 7$ (at $z \sim 0.1$) to ~ 9 (at $z \sim 1$), and increases from $\log_{10}(M_*/M_\odot) \sim 9.5$ to $\log_{10}(M_*/M_\odot) \sim 10$, while going to higher redshifts, from $z = 1.5$ to 2.5 . For passive galaxies, the limit is 0.3 – 0.4 dex higher up to $z \sim 1.8$, and about 0.7 dex higher at $z > 2$. Up to $z = 3$, there are 4 654 430 galaxies over the 95% stellar mass completeness limit.

4.2. Star-forming and passive galaxy classification

Colour-based classifications of galaxies use the principle of separating red and blue galaxies and distinguish between dusty and intrinsically red ones (e.g. the $U - V$ versus $V - J$ colour diagram Labbé et al. 2005; Wuyts et al. 2007; Williams et al. 2010). In this work, we use the $NUV - r^+$, $r^+ - J$ colours (Ilbert et al. 2010), whereby star-forming and passive galaxies are discriminated based on their absolute NUV , r^+ , and J magnitudes, with quiescent galaxies satisfying the following relations:

$$\begin{aligned} NUV - r^+ &> 3(r^+ - J) + 1; \\ NUV - r^+ &> 3.1. \end{aligned} \quad (7)$$

This combination of criteria works in a similar fashion to the UVJ diagram, but it has been widely used in recent literature as it is more sensitive to recent star formation via the $NUV - r^+$ colour, which separates passive (redder) and star-forming (bluer) galaxies (Ilbert et al. 2013; Arnouts et al. 2013). Truly passive galaxies are then distinguished from dusty, star-forming objects via the $r^+ - J$ colour, ensuring a proper separation between the two different populations.

The colours $NUV - r^+$ versus $r^+ - J$ of our sample are shown in Fig. 8, colour-coded as a function of the number density of objects with certain colours (top panels) and the median T_0/τ (middle panels) and sSFR in each bin (bottom panels), in the four different redshift bins. Only objects whose mass is higher than the $M_{\text{lim}}(z)$ value at the z_{min} of the redshift bin are shown, to account for sample incompleteness (see Fig. 7). To be consistent with the selection criteria reported above, the adopted absolute magnitude, J , is not the one estimated in the rest frame *Euclid* J_E filter but in the $J_{\text{UltraVISTA}}$ filter (Euclid Collaboration: Schirmer et al. 2022); similarly, we adopted the r^+ band as in Ilbert et al. (2010).

We recover the well-known increase in the fraction of passive galaxies while going to later cosmic times, highlighting the assembly of the population of quiescent galaxies observed in the Universe. When checking the colour diagrams against the measured values of sSFR, we notice the presence of a small fraction of objects ($\sim 3\%$) outside the boundaries for quiescent galaxies in the $NUV-r^+-J$ diagram, but with a median value of $\log_{10}(\text{sSFR}/\text{yr}^{-1}) < -11$ (and $T_0/\tau > 8$); that is, in a region of the M_* -SFR plane where we would expect only quiescent galaxies. This is a small, negligible number of interlopers, which is reassuring with regard to the goodness of the classification based on the rest-frame colours.

4.3. The SFR- M_* relation in the EDFs

The M_* -SFR plane for the sample of SFGs selected with the $NUV-r^+-J$ colours is shown in Fig. 9, in four redshift bins, colour-coded in terms of the logarithmic density of objects per bin. The SFMS is observed up to the highest redshift bin ($z < 3$). In the figure, we report three previously published SFMS relations for comparison: Eq. (15) of Popesso et al. (2023, black dashed line), that is, a comprehensive compilation of 27 literature SFMS relations fitted to the same functional form, in the mass range $8.5 < \log_{10}(M_*/M_\odot) < 11.5$; Schreiber et al. (2015, dotted grey line), obtained from ~ 10 k galaxies with the deepest *Herschel* observations of the GOODS and CANDELS-*Herschel* programme, with $9.5 < \log_{10}(M_*/M_\odot) < 11.5$; and Daddi et al. (2022a, dash-dotted blue line), obtained from a stacking analysis of ~ 400 k colour-selected SFGs in COSMOS (Delvecchio et al. 2021), here too in the mass range $8.5 < \log_{10}(M_*/M_\odot) < 11.5$.

All these SFMS forms include the presence of a bending of the relation at the high-mass end. Having enough statistics in terms of the number of objects per bin, we can significantly constrain the deviation from the linear form at high masses. For example, in each redshift interval Delvecchio et al. (2021) and Daddi et al. (2022a) found no more than 838 SFGs at $\log_{10}(M_*/M_\odot) > 11$, and in the same mass range the stacking analysis in Schreiber et al. (2015) has < 15 galaxies per bin. Analogously, most of the studies in the compilation of Popesso et al. (2023) stop at $\log_{10}(M_*/M_\odot) = 11$, and those that extend further never reach more than 10^3 galaxies per bin in comparable redshift ranges. Due to the large area observed in Q1, the minimum number of galaxies we have at $\log_{10}(M_*/M_\odot) > 11$ is 875, in the $[0.2, 0.5]$ redshift bin. This number increases to 7640 (for $0.5 < z < 0.8$), 18951 (for $0.8 < z < 1.5$), and 8165 (for $1.5 < z < 3.0$), one or two orders of magnitude higher than the former reported statistics. These numbers are more than enough to subdivide the $\log_{10}(M_*/M_\odot) > 11$ region into two different bins at $z > 0.8$, to better constrain the part of the SFMS where the SFR appears to saturate. The same reasoning, scaled by an order of magnitude, applies if we consider $\log_{10}(M_*/M_\odot) > 10.5$.

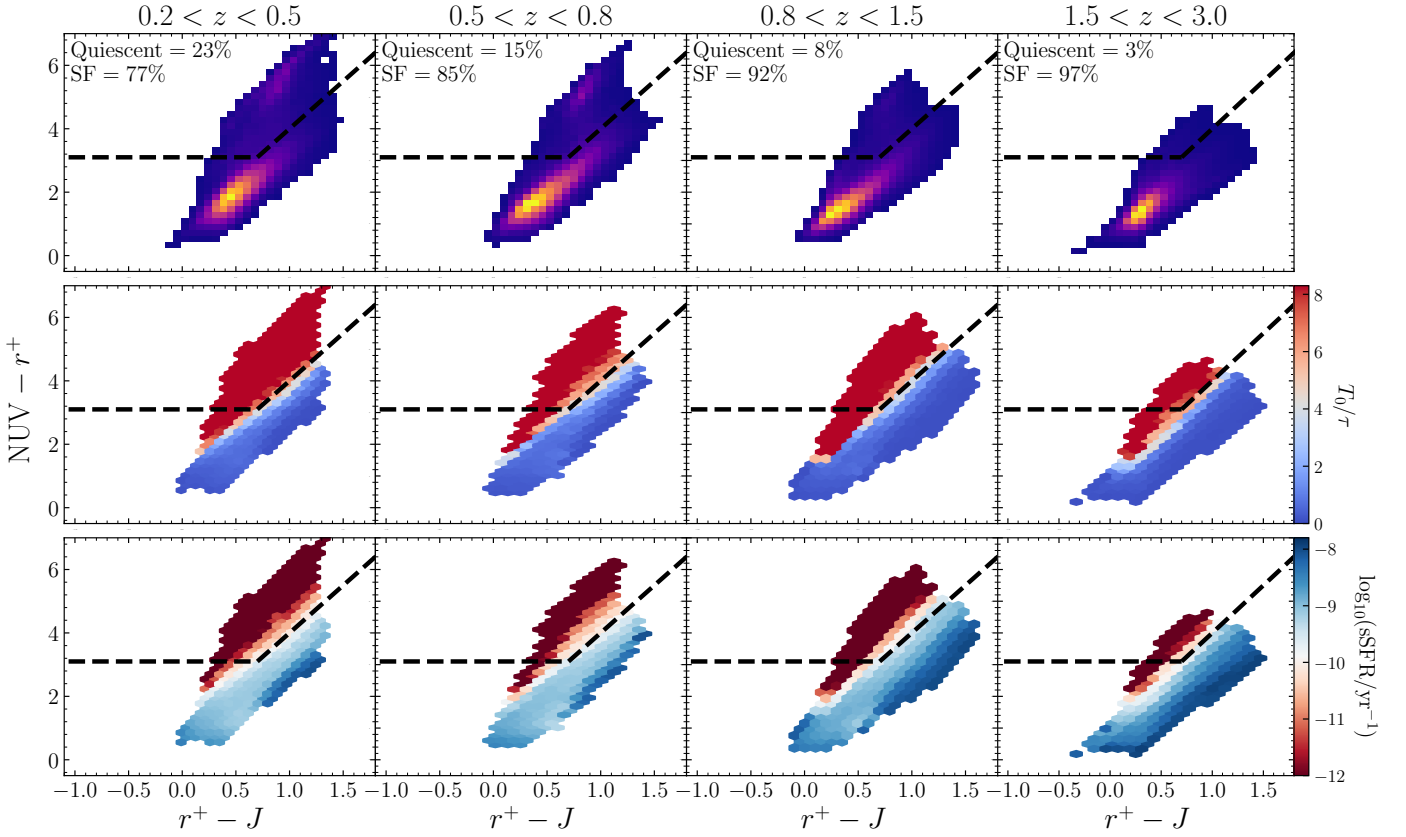


Fig. 8. $\text{NUV} - r^+$ versus $r^+ - J$ rest-frame colours for the sources above the 95% mass completeness limit, in four redshift bins, colour-coded with the density of objects (*upper panels*), the ratio between T_0 and τ (*middle panels*), and the sSFR (*bottom panels*). Dash-dotted lines divide the regions between star-forming and quiescent galaxies, following the criteria presented in [Ilbert et al. \(2013\)](#).

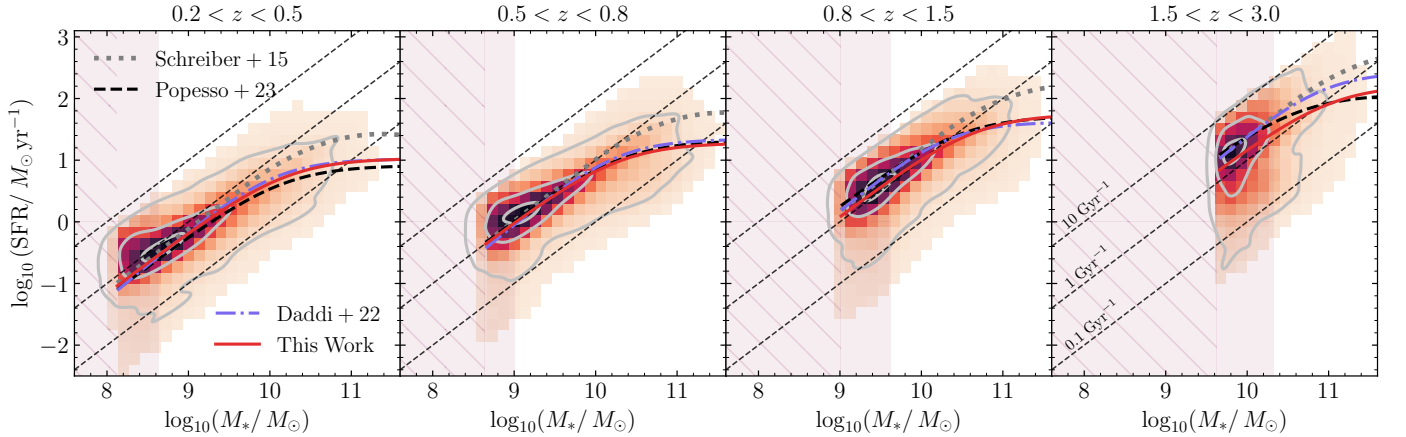


Fig. 9. SFR- M_* relation for SFGs, selected with the $\text{NUV} - r^+$ versus $r^+ - J$ condition, coloured by the density of objects in each redshift bin. Within the coloured bins is 99.9% of the sample of SFGs, while the grey contours are the lines enclosing 95%, 50%, and 10% of it. The red line is the fitted SFMS presented in this work, starting from the completeness limit mass at the start of the z -bin. The other SFMSs are from [Popesso et al. \(2023\)](#), (dashed black line), [Schreiber et al. \(2015\)](#), (dotted grey line), and [Daddi et al. \(2022a\)](#), (dashed-dotted blue line), all evaluated at the median redshift of each redshift bin (or to the closest reported value). The shaded purple regions highlight where the mass is below the completeness limit at the start (hatched) and at the end of the redshift bin. Dashed grey lines highlight the region of constant sSFR.

Our fit to the observed data is the red line in Fig. 9. Overall, the fits in [Popesso et al. \(2023\)](#) and [Daddi et al. \(2022a\)](#) closely resemble our results, while [Schreiber et al. \(2015\)](#) find systematically higher SFRs at the highest mass end. The functional form that we fit these data to, first proposed by [Lee et al. \(2015\)](#), is the same as in Eq. (15) of [Popesso et al. \(2023\)](#) or Eq. (1) of [Daddi](#)

[et al. \(2022a\)](#),

$$\text{SFR} = \frac{\text{SFR}_{\text{max}}}{1 + (M_0/M_*)^\gamma}; \quad (8)$$

that is, a parameterisation where the SFR is linked to the stellar mass through three parameters – the bending mass after which the relation deviates from the linear behaviour (M_0), the maxi-

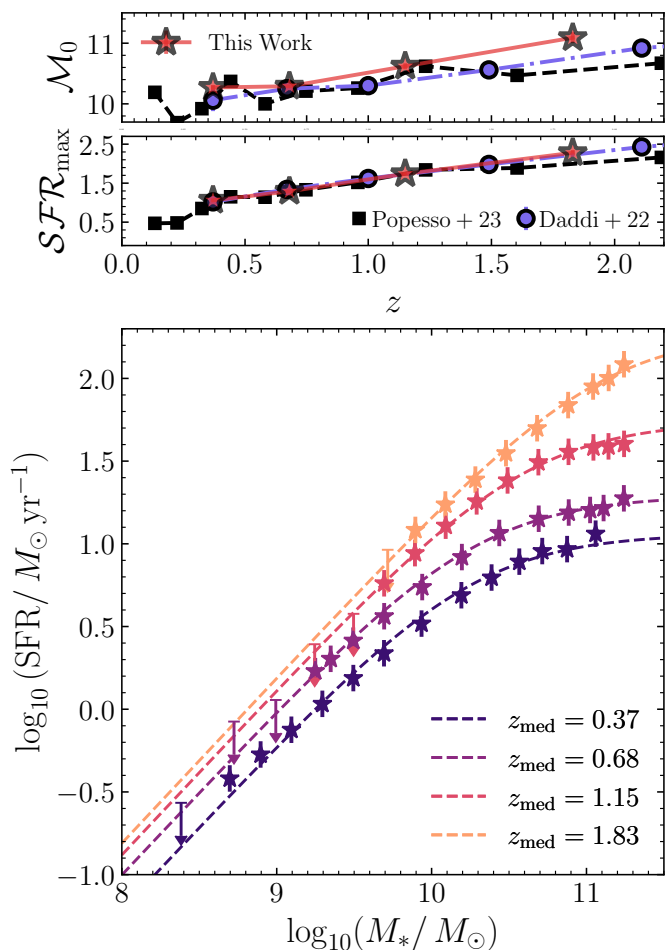


Fig. 10. *Lower panel:* Our fit of the SFMS, in the four redshift bins (stars with error bars and arrows) and best-fit model (dashed lines). The reported redshift is the median z of the bin. *Upper panels:* Best-fit (red stars) and reported (squares for Popesso et al. 2023; circles for Daddi et al. 2022a) values for the bending mass and maximum SFR, defined as $M_0 \equiv \log_{10}(M_0/M_\odot)$ and $SFR_{\max} \equiv \log_{10}(SFR_{\max}/M_\odot \text{ yr}^{-1})$, as a function of redshift.

imum SFR for $M_* \gg M_0$ (SFR_{\max}), and the slope of the linear relation when $M_* \ll M_0$ (γ). These parameters have been shown to be directly linked to fundamental properties in models of gas accretion (e.g. the bending mass M_0 is directly linked to the ratio between M_{shock} and M_{stream} , see Daddi et al. 2022a,b). For each fit, we kept γ fixed at 1.0, as it has been shown to be a representative value around which almost every fit converges at different redshifts (see Lee et al. 2015; Popesso et al. 2023; Daddi et al. 2022a). This also makes it possible to make a direct comparison with these works in terms of bending mass and the maximum SFR.

When fitting the SFMS, we must take into account the fact that for each redshift bin there are two possible values for the mass completeness limit, depending on whether consider the mass at the lower limit of the redshift bin, z_{start} , or at the higher limit, z_{end} . Depending on the width of the redshift bin, these two masses can differ significantly, of the order of 0.4 – 0.6 dex (see Fig. 9). When $M_* > M_{\text{lim}}(z_{\text{end}})$, we are dealing with a complete sample; in this case, we fit the points measured as the median stellar masses and SFRs of the distribution of SFGs in each bin of mass, with the associated uncertainty as the quadrature sum of the standard deviation of the median in each bin

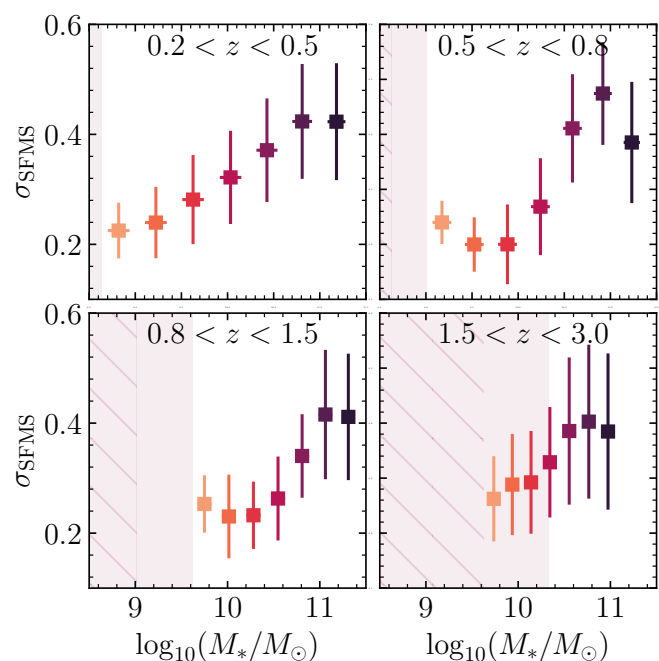


Fig. 11. Scatter of the SFR– M_* relation, σ_{SFMS} , in different mass bins. The purple shaded regions highlight where the mass is below the completeness limit at the start (hatched) and end of the redshift bin.

Table 2. Fitted values for $\log_{10}(M_0/M_\odot)$, $\log_{10}(SFR_{\max}/M_\odot \text{ yr}^{-1})$, and the relation scatter σ_{SFMS} for each redshift bin, with the median redshift in bin reported in the first column.

z_{med}	$\log_{10}(M_0/M_\odot)$	$\log_{10}(SFR_{\max}/M_\odot \text{ yr}^{-1})$	σ_{SFMS}
0.37	10.27 ± 0.04	1.06 ± 0.03	0.34 ± 0.11
0.68	10.29 ± 0.05	1.29 ± 0.03	0.27 ± 0.09
1.15	10.62 ± 0.05	1.74 ± 0.04	0.28 ± 0.08
1.83	11.10 ± 0.09	2.28 ± 0.07	0.40 ± 0.12

and the typical uncertainty on the SFRs (see Sect. 3.3). These are shown as stars in the bottom panel of Fig. 10, coloured as a function of the redshift bin to which they belong. When $M_* < M_{\text{lim}}(z_{\text{start}})$ we consider the sample to be incomplete, and these galaxies are excluded from the fit. In the mass bins where $M_{\text{lim}}(z_{\text{start}}) < M_* < M_{\text{lim}}(z_{\text{end}})$, we are preferentially missing lower-mass galaxies, which tend to have lower SFRs. As a result, we treat the SFR data points in these bins as upper limits, as is indicated by the arrows in the bottom panel of Fig. 10.

The SFR–mass bins were obtained by binning the distribution of stellar mass to uniformly cover the stellar masses space with a similar number of sources to make these statistically significant. The bins start from where the stellar mass is higher than the 95% completeness limit at the lower limit of the redshift bin; that is, $\log_{10}(M_*/M_\odot) \sim 8.1$ at $z = 0.37$, $\log_{10}(M_*/M_\odot) \sim 8.6$ at $z = 0.67$, $\log_{10}(M_*/M_\odot) \sim 9.0$ at $z = 1.15$, and $\log_{10}(M_*/M_\odot) \sim 9.5$ at $z = 1.83$. The results of the fit are reported in Table 2, and shown in Fig. 10 as dashed lines colour-coded by redshift bin. The bending mass and maximum SFR are reported as a function of redshift in the upper panels of Fig. 10, as red stars and solid lines. We also show, for comparison, the same parameters found in Daddi et al. (2022a, as blue circles and dashed dotted line) and Popesso et al. (2023, black edged squares and dashed line) as a function of z . As an internal check, we also fitted the same points with a linear relation (e.g., Eq. 10 in Popesso et al.

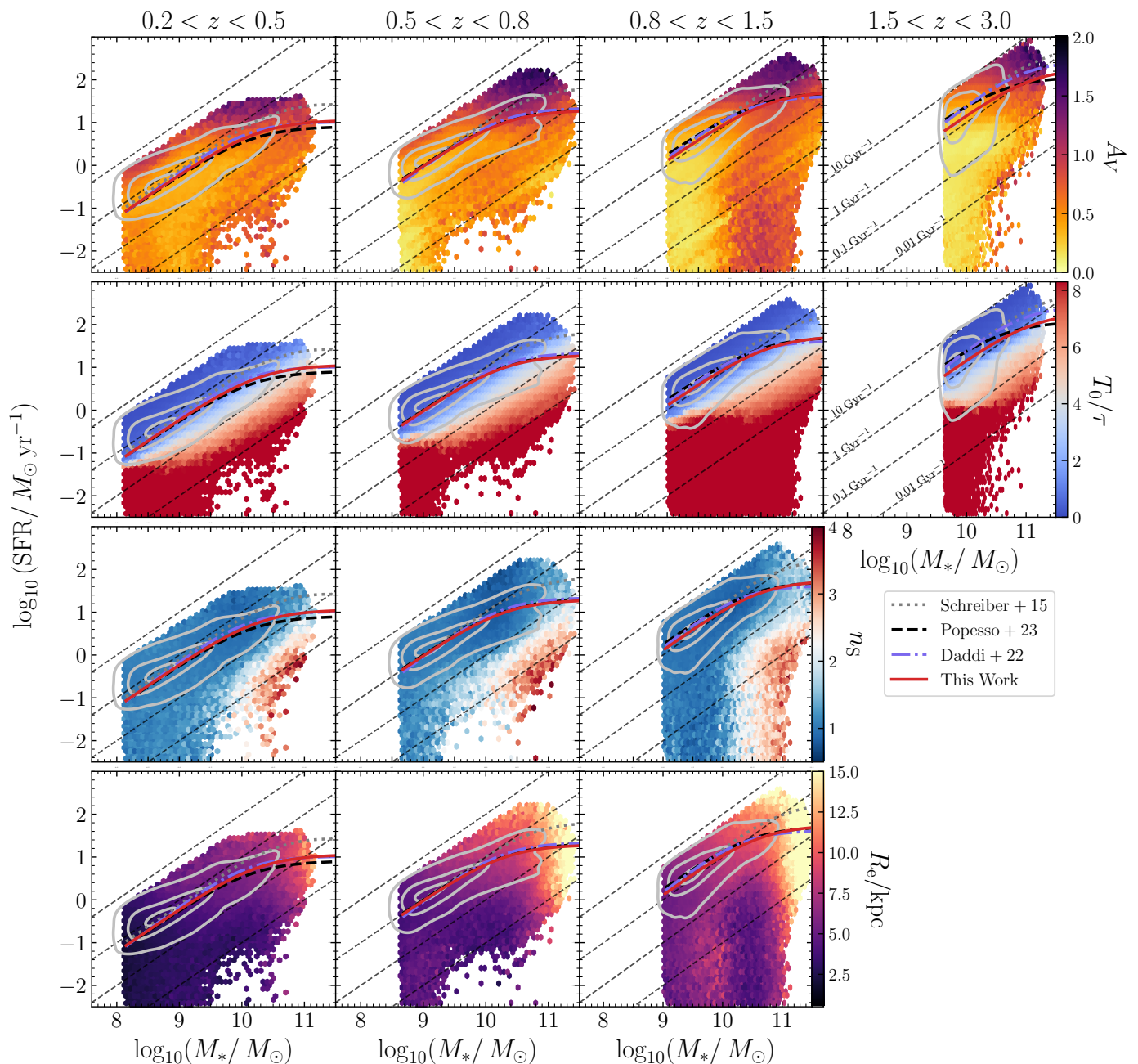


Fig. 12. SFR– M_* relation, colour-coded by the median A_V (top), T_0/τ (middle-top), Sérsic index (middle-bottom), and Sérsic radius (bottom) of objects in each bin. A bin is coloured only when the number of sources falling within the bin is higher than 20. The contours in grey contain 90%, 50%, and 10% of the full sample. The literature SFMS are the same as the ones shown in Fig. 9 and reported in Sect. 4. Dashed lines highlight the loci of equal sSFR.

2023), obtaining worse χ^2 for each redshift bin with respect to the relation including the bending at the high-mass end.

Our recovered bending masses M_0 increase with redshift. We find results similar to those of Daddi et al. (2022a), with the main difference in the [0.8, 1.5] bin, with a higher $\log_{10}(M_0/M_\odot) = 10.62 \pm 0.05$ (similar to what was found in Popesso et al. 2023). In the [1.5, 3.0] bin, our bending mass is more in line with Daddi et al. (2022a) at $\log_{10}(M_0/M_\odot) = 11.10 \pm 0.09$, (0.6 dex higher than the one in Popesso et al. 2023). However, for this redshift bin, we must again warn the reader about the caveats highlighted in Sect. 3.3 and Fig. 4; that is, at $11 < \log_{10}(M_*/M_\odot) < 11.5$ any wrong redshift attribution will introduce an important source of systematic biases, skewing the fitted bins at higher values of

M_* and SFR, and therefore the recovered M_0 and SFR_{max} . In any case, the observed agreement and evolution with redshift are remarkable, when taking into account the fact that our SFRs have been evaluated with photometry ranging from the g band (the u band is available only in one field) up to $4.5 \mu\text{m}$, and thus without properly accounting for dust-obscured star formation processes, which are disentangled from quiescence only when photometry from the far-IR up to radio is available.

Trying to find the best-fit parameters of the SFMS with the functional form in Eq. (8), we also measured the scatter in the relation σ_{SFMS} , fitting the difference between the observed set of galaxy stellar masses and SFRs and the model as a normal distribution, $\mathcal{N}(\sigma_{\text{SFMS}})$. We find the same tight scatter, $\sigma_{\text{SFMS}} \approx 0.3$,

observed in previous studies (e.g. Clarke et al. 2024, for a SED-fitting derived SFMS), consistent in all redshift intervals within the uncertainties (Table 2). We also investigated whether there is any correlation between σ_{SFMS} and stellar masses. The topic is still highly debated, from both theoretical and observational points of view. Simulations find either a mass dependence (e.g. Ceverino et al. 2018; Katsianis et al. 2019; Legrand et al. 2022) or no dependence at all (e.g. Donnari et al. 2019; Pallottini & Ferrara 2023). The same is true for observational studies, with reported cases of no trend (e.g. Schreiber et al. 2015; Shivaie et al. 2015; Clarke et al. 2024), a monotonically increasing trend (Popesso et al. 2019; Sherman et al. 2021; Cole et al. 2025), or even a decrease with stellar mass (Willett et al. 2015). In each redshift interval, we measured the SFMS scatter in seven different mass bins as was previously done, starting from the mass limit completeness at the end of the interval (or at the start, for $1.5 < z < 3.0$) in order to reduce any systematic effect arising from mass incompleteness. To account for the uncertainties in redshift, stellar mass, and SFR, we measured σ_{SFMS} in each mass bin ten times, each time on a sample generated perturbing the best-fit redshifts and PPs with the measured uncertainties. We then took every measure of σ_{SFMS} as the median of all iterations, with associated uncertainties as the 16th and 84th percentile of the distributions. The results are reported in Fig. 11. Even though caution must be taken given the various sources of uncertainties, we observe an increase in σ_{SFMS} with stellar mass, or at least an overall increasing trend.

The SFMS dependence on the absorption, A_V , and the ratio between the age and the e-folding time, τ , is shown in the first two rows of Fig. 12, where the colours correspond to the median of A_V and T_0/τ values of the sources within each bin, with the SFMS contours superimposed in grey. In this case, we have not limited the sample to SFGs only (i.e. we have not removed the sources identified as quiescent in the NUV- r^+ - J diagram), highlighting a bin when the number of sources within the bin is greater than 20. In this way, we can explore less populated regions of the M_* -SFR plane, where passive galaxies are expected to appear. The location of galaxies in the plane below a certain limiting sSFR (e.g. 0.01 Gyr^{-1}) should not be considered as absolute, but rather as indicative that those objects are passive galaxies and could lie anywhere below that line. As was expected, maximum extinction values ($A_V > 2.0$) are associated with the most massive and star-forming galaxies of the SFMS, with $1 < \text{sSFR}/\text{Gyr}^{-1} < 10$, while low or zero dust extinction values are observed at lower sSFRs. The typical A_V of SFMS galaxies decreases slightly from $A_V \sim 1.0$ to $A_V \sim 0.5$ with increasing redshift. Looking at the distribution of fitted A_V , we recover only a small fraction of objects ($\sim 1\%$) with a high extinction of $A_V > 2.5$, with the A_V distribution peaking at $0.5 < A_V < 1.0$. This could be due to the particular limitations and caveats of the recovered sample (see Sect. 3), especially in the range of wavelengths covered. Lower sSFR galaxies are associated with stellar populations with older ages, in particular at low redshift, and high age/ τ ratios, while younger ages are found in the upper part of the SFMS and at high redshift (Nersisyan et al. 2025).

The morphological parameters of all the *Euclid*-detected sources were measured by running the SourceExtractor++ code (Bertin et al. 2020; Kümmel et al. 2022), fitting the detections as two-dimensional Sérsic profiles (see Euclid Collaboration: Romelli et al. 2025 and Euclid Collaboration: Quilley et al. 2025, for further details). Here, we focus on the Sérsic radius, R_e , and the Sérsic index, n_S . In order to be more conservative, following what was found in Euclid Collaboration: Quilley

et al. (2025, see their Fig. 2), in addition to the cuts described in Sect. 2, we also removed the sources whose fit did not converge on a solution. In these cases, the Sérsic axis ratio is exactly equal to one, and the Sérsic index $n_S < 0.302$ or $n_S > 5.45$. These cuts reduce our sample for morphological analysis by 17%, leaving a total of 6 702 811 sources.

In the two bottom rows of Fig. 12, we report the SFMS, colour-coded for the median values of n_S (middle-bottom row) and R_e (bottom row), converted from arcseconds to kiloparsecs for the predicted photo- z . We note that the higher values of $n_S > 3$ are associated with the most massive, low-sSFR objects, and the radii, R_e , increase with stellar mass at all redshifts. The combination of the limited volume of Universe sampled at low z with the increasing difficulty of properly identifying smaller galaxies at higher redshifts due to the fixed *Euclid* resolution (roughly translating into a limit in measured effective radiuses, $R_{e,\text{lim}} \approx 2 - 3 \text{ kpc}$, in the redshift interval [0.5, 3.0]) results in an apparent decrease in R_e with cosmic time, as the distributions of galaxies falling within each stellar mass-SFR bin suffer from this size incompleteness.

In all redshift intervals, the concurrent presence of the SFMS and the prominence of exponential disks with $n_S \approx 1$ (in blue) is immediately visible, while at lower sSFR, close to the limit where measuring SFRs becomes difficult, we note a crowding of sources with de Vaucouleurs profiles of $n_S \approx 4$ (in red). Between those, there is an intermediate region of profiles with $2 < n_S < 3$ (the white strip) at about $\text{sSFR} \sim 0.05 \text{ Gyr}^{-1}$, which indicates an evolutionary trend in the structural parameters while moving in the stellar mass-SFR plane, in agreement with what has been observed previously in other samples (Kauffmann et al. 2003; Wuyts et al. 2011; Whitaker et al. 2015; McPartland et al. 2019; Cooke et al. 2023; Martorano et al. 2025). This kind of transition is observed only at the medium-to-high-mass end $\log_{10}(M_*/M_\odot) > 10$, and appears to show only a weak redshift dependence. Interestingly, the mass over which the sequence of passive $n_S \approx 4$ profiles is observed is almost coincident with the fitted bending mass, M_0 , a trend that has already been observed in the literature (e.g. Leslie et al. 2020).

Euclid combines a large sample area with a uniform multi-wavelength view, which has great potential for environmental studies. Despite that, here we do not focus on how the environment shapes SFMS, as this is the main scope of other Q1 works (Euclid Collaboration: Cleland et al. 2025; Euclid Collaboration: Mai et al. 2025). The soon-to-be-available DR1 observations of the EDFs are expected to have at least ten ROSSs, which corresponds to 1.25 magnitude-deeper data. Combining the DR1 data, along with other multi-wavelength observations and the *Euclid* spectroscopy sample, will allow for a thorough investigation of the environmental effects on the SFMS.

5. Conclusions

The *Euclid* Q1 data release, with its first look at the EDFs, is already a good test of the capabilities of the mission to investigate the formation and evolution of galaxies, especially in terms of the enhanced statistics of such a large area of the extragalactic sky. In this work, we have investigated the relation between stellar masses and SFRs, the SFMS, and how it relates to the other well-constrained PPs and morphological parameters, comparing the results with the existing literature, as a fair validation of the *Euclid* results and a first exploration of the mission capabilities to return a statistically robust sample. Even a single ROS of the EDFs is able to yield reliable measurements for the photometric redshifts, stellar masses, and SFRs – which in this work have

been improved due to the addition of the first two channels of IRAC – for more than eight million galaxies at a magnitude cut of $H_E < 24$. In particular, more than ~ 30 k galaxies are found with $\log_{10}(M_*/M_\odot) > 11$, a substantial improvement over all the other extragalactic surveys limited to a few square degrees of the sky.

The recovered scaling relation between stellar mass and SFR is consistent with what is known from previous studies on the SFMS. The data show a similar tight dispersion, between 0.26 and 0.40 dex, with an increase in σ_{SFMS} occurring at higher masses, and are in excellent agreement with the functional forms reported in Poppo et al. (2023), Daddi et al. (2022a), and Schreiber et al. (2015). These works fit SFMS with a non-linear term that translates into a bending of the relation, becoming more pronounced at high masses. We fit the relation with the same parameterisation – that is, linking the SFR directly to the bending mass, M_0 , and the SFR maximum, SFR_{max} – recovering the same statistical agreement with the previous studies. M_0 increases almost monotonically with redshift, starting from $\log_{10}(M_0/M_*) \simeq 10.3$ at $z = 0.37$, up to 11.1 at cosmic noon. Our results are in accord with the presence of this reduction in SFR in massive galaxies, with a better χ^2 with respect to a linear one. At the same time, when correlating our results with the catalogue of morphological properties, we recover the well-known bimodality between exponential and smaller star-forming disks and de Vaucouleurs profiles for passive and bigger galaxies.

Acknowledgements. AnEn, LuPo, MiBo, EmDa, ViAl, FaGe, SaQu, MaTa, MaSc, and BeGr acknowledge support from the ELSA project. "ELSA: Euclid Legacy Science Advanced analysis tools" (Grant Agreement no. 101135203) is funded by the European Union. Views and opinions expressed are however those of the author(s) only and do not necessarily reflect those of the European Union or Innovate UK. Neither the European Union nor the granting authority can be held responsible for them. UK participation is funded through the UK HORIZON guarantee scheme under Innovate UK grant 10093177. AnEn, MiBo, StCa, ElZu acknowledge support from the INAF MiniGrant 2023 "ADIEU: Anomaly Detections In Euclid". This work has been partially funded by Italian MUR Premiale MITIC 2017. The Euclid Consortium acknowledges the European Space Agency and a number of agencies and institutes that have supported the development of *Euclid*, in particular the Agenzia Spaziale Italiana, the Austrian Forschungsförderungsgesellschaft funded through BMK, the Belgian Science Policy, the Canadian Euclid Consortium, the Deutsches Zentrum für Luft- und Raumfahrt, the DTU Space and the Niels Bohr Institute in Denmark, the French Centre National d'Etudes Spatiales, the Fundação para a Ciência e a Tecnologia, the Hungarian Academy of Sciences, the Ministerio de Ciencia, Innovación y Universidades, the National Aeronautics and Space Administration, the National Astronomical Observatory of Japan, the Nederlandse Onderzoekschool Voor Astronomie, the Norwegian Space Agency, the Research Council of Finland, the Romanian Space Agency, the State Secretariat for Education, Research, and Innovation (SERI) at the Swiss Space Office (SSO), and the United Kingdom Space Agency. A complete and detailed list is available on the *Euclid* web site (www.euclid-ec.org). This work has made use of the *Euclid* Quick Release Q1 data from the *Euclid* mission of the European Space Agency (ESA), 2025, <https://doi.org/10.57780/esa-2853f3b>. This work has made use of CosmoHub, developed by PIC (maintained by IFAE and CIEMAT) in collaboration with ICE-CSIC. CosmoHub received funding from the Spanish government (MCIN/AEI/10.13039/501100011033), the EU NextGeneration/PRTR (PRTR-C17.I1), and the Generalitat de Catalunya. Based on data from UNIONS, a scientific collaboration using three Hawaii-based telescopes: CFHT, Pan-STARRS, and Subaru www.skysurvey.cc. Based on data from the Dark Energy Camera (DECam) on the Blanco 4-m Telescope at CTIO in Chile <https://www.darkenergysurvey.org>. In preparation for this work, we used the following codes for Python: Numpy (Harris et al. 2020), Scipy (Virtanen et al. 2020), Pandas (Wes McKinney 2010), mnpz (Tanaka et al. 2018),

References

Abdurro'uf & Akiyama, M. 2017, MNRAS, 469, 2806
 Ahumada, R., Allende Prieto, C., Almeida, A., et al. 2020, ApJS, 249, 3
 Arnouts, S., Le Floch, E., Chevillard, J., et al. 2013, A&A, 558, A67
 Baes, M., Nersesian, A., Casasola, V., et al. 2020, A&A, 641, A119

Baker, W. M., Maiolino, R., Belfiore, F., et al. 2023, MNRAS, 518, 4767
 Baker, W. M., Maiolino, R., Bluck, A. F. L., et al. 2022, MNRAS, 510, 3622
 Behroozi, P., Wechsler, R. H., Hearin, A. P., & Conroy, C. 2019, MNRAS, 488, 3143
 Bertin, E., Schefer, M., Apostolakis, N., et al. 2020, in Astronomical Society of the Pacific Conference Series, Vol. 527, Astronomical Data Analysis Software and Systems XXIX, ed. R. Pizzo, E. R. Deul, J. D. Mol, J. de Plaa, & H. Verkouter, 461
 Blake, C., Amon, A., Childress, M., et al. 2016, MNRAS, 462, 4240
 Bower, R. G., Schaye, J., Frenk, C. S., et al. 2017, MNRAS, 465, 32
 Brammer, G. B., Sánchez-Janssen, R., Labbé, I., et al. 2012, ApJ, 758, L17
 Brinchmann, J., Charlot, S., White, S. D. M., et al. 2004, MNRAS, 351, 1151
 Bruzual, G. & Charlot, S. 2003, MNRAS, 344, 1000
 Calzetti, D., Armus, L., Bohlin, R. C., et al. 2000, ApJ, 533, 682
 Ceverino, D., Klessen, R. S., & Glover, S. C. O. 2018, MNRAS, 480, 4842
 Chabrier, G. 2003, PASP, 115, 763
 Childress, M. J., Lidman, C., Davis, T. M., et al. 2017, MNRAS, 472, 273
 Ciesla, L., Elbaz, D., & Fensch, J. 2017, A&A, 608, A41
 Clarke, L., Shapley, A. E., Sanders, R. L., et al. 2024, ApJ, 977, 133
 Coil, A. L., Blanton, M. R., Burles, S. M., et al. 2011, ApJ, 741, 8
 Cole, J. W., Papovich, C., Finkelstein, S. L., et al. 2025, ApJ, 979, 193
 Colls, M., Dalton, G., Maddox, S., et al. 2001, MNRAS, 328, 1039
 Cooke, K. C., Kartaltepe, J. S., Rose, C., et al. 2023, ApJ, 942, 49
 Daddi, E., Delvecchio, I., Dimauro, P., et al. 2022a, A&A, 661, L7
 Daddi, E., Dickinson, M., Morrison, G., et al. 2007, ApJ, 670, 156
 Daddi, E., Rich, R. M., Valentino, F., et al. 2022b, ApJ, 926, L21
 Dark Energy Survey Collaboration, Abbott, T., Abdalla, F. B., et al. 2016, MNRAS, 460, 1270
 Dekel, A. & Birnboim, Y. 2006, MNRAS, 368, 2
 Delvecchio, I., Daddi, E., Sargent, M. T., et al. 2021, A&A, 647, A123
 DESI Collaboration, Adame, A. G., Aguilar, J., et al. 2024, AJ, 168, 58
 DESI Collaboration, Aghamousa, A., Aguilar, J., et al. 2016, arXiv:1611.00036
 D'Eugenio, F., Cameron, A. J., Scholtz, J., et al. 2025, ApJS, 277, 4
 Donnari, M., Pillepich, A., Nelson, D., et al. 2019, MNRAS, 485, 4817
 Elbaz, D., Dickinson, M., Hwang, H. S., et al. 2011, A&A, 533, A119
 Ellison, S. L., Lin, L., Thorp, M. D., et al. 2021, MNRAS, 501, 4777
 Ellison, S. L., Sánchez, S. F., Ibarra-Medel, H., et al. 2018, MNRAS, 474, 2039
 Enia, A., Rodighiero, G., Morselli, L., et al. 2020, MNRAS, 493, 4107
 Euclid Collaboration: Aussel, H., Tereno, I., Schirmer, M., et al. 2025, A&A, submitted (Euclid Q1 SI), arXiv:2503.15302
 Euclid Collaboration: Bisigello, L., Rodighiero, G., Fotopoulou, S., et al. 2025, A&A, submitted (Euclid Q1 SI), arXiv:2503.15323
 Euclid Collaboration: Castander, F., Fosalba, P., Stadel, J., et al. 2025, A&A, 697, A5
 Euclid Collaboration: Cleland, C., Mei, S., De Lucia, G., et al. 2025, A&A, submitted (Euclid Q1 SI), arXiv:2503.15313
 Euclid Collaboration: Corcho-Caballero, P., Ascasibar, Y., Verdoes Kleijn, G., et al. 2025, A&A, submitted (Euclid Q1 SI), arXiv:2503.15315
 Euclid Collaboration: Cropper, M., Al-Bahlawan, A., Amiaux, J., et al. 2025, A&A, 697, A2
 Euclid Collaboration: Desprez, G., Paltani, S., Coupon, J., et al. 2020, A&A, 644, A31
 Euclid Collaboration: Enia, A., Bolzonella, M., Pozzetti, L., et al. 2024, A&A, 691, A175
 Euclid Collaboration: Jahnke, K., Gillard, W., Schirmer, M., et al. 2025, A&A, 697, A3
 Euclid Collaboration: Mai, N., Mei, S., Cleland, C., et al. 2025, A&A, submitted (Euclid Q1 SI), arXiv:2503.15331
 Euclid Collaboration: Matamoros Zatarain, T., Fotopoulou, S., Ricci, F., et al. 2025, A&A, submitted (Euclid Q1 SI), arXiv:2503.15320
 Euclid Collaboration: McCracken, H. J., Benson, K., Dolding, C., et al. 2025, A&A, submitted (Euclid Q1 SI), arXiv:2503.15303
 Euclid Collaboration: McPartland, C. J. R., Zalesky, L., Weaver, J. R., et al. 2025, A&A, 695, A259
 Euclid Collaboration: Mellier, Y., Abdurro'uf, Acevedo Barroso, J., et al. 2025, A&A, 697, A1
 Euclid Collaboration: Moneti, A., McCracken, H. J., Shuntov, M., et al. 2022, A&A, 658, A126
 Euclid Collaboration: Polenta, G., Frailis, M., Alavi, A., et al. 2025, A&A, submitted (Euclid Q1 SI), arXiv:2503.15304
 Euclid Collaboration: Quilley, L., Damjanov, I., de Lapparent, V., et al. 2025, A&A, submitted (Euclid Q1 SI), arXiv:2503.15309
 Euclid Collaboration: Romelli, E., Kümmel, M., Dole, H., et al. 2025, A&A, submitted (Euclid Q1 SI), arXiv:2503.15305
 Euclid Collaboration: Scaramella, R., Amiaux, J., Mellier, Y., et al. 2022, A&A, 662, A112
 Euclid Collaboration: Schirmer, M., Jahnke, K., Seidel, G., et al. 2022, A&A, 662, A92
 Euclid Collaboration: Tucci, M., Paltani, S., Hartley, W. G., et al. 2025, A&A in press (Euclid Q1 SI), arXiv:2503.15306

- Euclid Collaboration: Walmsley, M., Huertas-Company, M., Quilley, L., et al. 2025, A&A, submitted (Euclid Q1 SI), arXiv:2503.15310
- Euclid Quick Release Q1. 2025. <https://doi.org/10.57780/esa-2853f3b>
- Fabian, A. C. 2012, ARA&A, 50, 455
- Fazio, G. G., Hora, J. L., Allen, L. E., et al. 2004, ApJS, 154, 10
- Flaugher, B., Diehl, H. T., Honscheid, K., et al. 2015, AJ, 150, 150
- Gwyn, S., McConnachie, A. W., Cuillandre, J.-C., et al. 2025, arXiv:2503.13783
- Halton, J. H. 1964, Commun. ACM, 7, 701–702
- Harris, C. R., Millman, K. J., van der Walt, S. J., et al. 2020, Nature, 585, 357
- Hsieh, B. C., Lin, L., Lin, J. H., et al. 2017, ApJ, 851, L24
- Huchra, J. P., Macri, L. M., Masters, K. L., et al. 2012, ApJS, 199, 26
- Ilbert, O., McCracken, H. J., Le Fèvre, O., et al. 2013, A&A, 556, A55
- Ilbert, O., Salvato, M., Le Floch, E., et al. 2010, ApJ, 709, 644
- Iyer, K. G., Gawiser, E., Faber, S. M., et al. 2019, ApJ, 879, 116
- Jones, D. H., Read, M. A., Saunders, W., et al. 2009, MNRAS, 399, 683
- Katsianis, A., Zheng, X., Gonzalez, V., et al. 2019, ApJ, 879, 11
- Kauffmann, G., Heckman, T. M., White, S. D. M., et al. 2003, MNRAS, 341, 54
- Kennicutt, Robert C., J. 1998a, ApJ, 498, 541
- Kennicutt, Jr., R. C. 1998b, ARA&A, 36, 189
- Kereš, D., Katz, N., Weinberg, D. H., & Davé, R. 2005, MNRAS, 363, 2
- Kriek, M., Shapley, A. E., Reddy, N. A., et al. 2015, ApJS, 218, 15
- Kroupa, P. 2001, MNRAS, 322, 231
- Kümmel, M., Vassallo, T., Dabin, C., & Gracia Carpio, J. 2022, in Astronomical Society of the Pacific Conference Series, Vol. 532, Astronomical Data Analysis Software and Systems XXX, ed. J. E. Ruiz, F. Pierfederici, & P. Teuben, 329
- Labbé, I., Huang, J., Franx, M., et al. 2005, ApJ, 624, L81
- Le Fèvre, O., Vettolani, G., Garilli, B., et al. 2005, A&A, 439, 845
- Lee, N., Sanders, D. B., Casey, C. M., et al. 2015, ApJ, 801, 80
- Legrand, L., Hutter, A., Dayal, P., et al. 2022, MNRAS, 509, 595
- Leja, J., Carnall, A. C., Johnson, B. D., Conroy, C., & Speagle, J. S. 2019, ApJ, 876, 3
- Leja, J., Speagle, J. S., Ting, Y.-S., et al. 2022, ApJ, 936, 165
- Leroy, L., Elbaz, D., Magnelli, B., et al. 2024, A&A, 691, A248
- Leslie, S. K., Schinnerer, E., Liu, D., et al. 2020, ApJ, 899, 58
- Lidman, C., Tucker, B. E., Davis, T. M., et al. 2020, MNRAS, 496, 19
- Lin, L., Belfiore, F., Pan, H.-A., et al. 2017, ApJ, 851, 18
- Lin, L., Pan, H.-A., Ellison, S. L., et al. 2019, ApJ, 884, L33
- Mancini, C., Daddi, E., Juneau, S., et al. 2019, MNRAS, 489, 1265
- Martorano, M., van der Wel, A., Baes, M., et al. 2025, A&A, 694, A76
- Matthee, J. & Schaye, J. 2019, MNRAS, 484, 915
- McPartland, C., Sanders, D. B., Kewley, L. J., & Leslie, S. K. 2019, MNRAS, 482, L129
- Morselli, L., Rodighiero, G., Enia, A., et al. 2020, MNRAS, 496, 4606
- Nersesian, A., van der Wel, A., Gallazzi, A. R., et al. 2025, A&A, 695, A86
- Oke, J. B. & Gunn, J. E. 1983, ApJ, 266, 713
- Pallottini, A. & Ferrara, A. 2023, A&A, 677, L4
- Pentericci, L., McLure, R. J., Franzetti, P., Garilli, B., & the VANDELS team. 2018, arXiv:1811.05298
- Popesso, P., Concas, A., Cresci, G., et al. 2023, MNRAS, 519, 1526
- Popesso, P., Concas, A., Morselli, L., et al. 2019, MNRAS, 483, 3213
- Pozzetti, L., Bolzonella, M., Zucca, E., et al. 2010, A&A, 523, A13
- Prevot, M. L., Lequeux, J., Maurice, E., Prevot, L., & Rocca-Volmerange, B. 1984, A&A, 132, 389
- Rodighiero, G., Daddi, E., Baronchelli, I., et al. 2011, ApJ, 739, L40
- Schmidt, M. 1959, ApJ, 129, 243
- Schreiber, C., Pannella, M., Elbaz, D., et al. 2015, A&A, 575, A74
- Sherman, S., Jogee, S., Florez, J., et al. 2021, MNRAS, 505, 947
- Shivaei, I., Reddy, N. A., Shapley, A. E., et al. 2015, ApJ, 815, 98
- Siudek, M., Pucha, R., Mezcuca, M., et al. 2024, A&A, 691, A308
- Speagle, J. S., Steinhardt, C. L., Capak, P. L., & Silverman, J. D. 2014, ApJS, 214, 15
- Talia, M., Schreiber, C., Garilli, B., et al. 2023, A&A, 678, A25
- Tanaka, M., Coupon, J., Hsieh, B.-C., et al. 2018, PASJ, 70, S9
- Tomczak, A. R., Quadri, R. F., Tran, K.-V. H., et al. 2016, ApJ, 817, 118
- Virtanen, P., Gommers, R., Oliphant, T. E., et al. 2020, Nature Methods, 17, 261
- Wang, T.-M., Magnelli, B., Schinnerer, E., et al. 2024, A&A, 681, A110
- Weaver, J. R., Kauffmann, O. B., Ilbert, O., et al. 2022, ApJS, 258, 11
- Werner, M. W., Roellig, T. L., Low, F. J., et al. 2004, ApJS, 154, 1
- Wes McKinney. 2010, in Proceedings of the 9th Python in Science Conference, ed. Stéfan van der Walt & Jarrod Millman, 56–61
- Whitaker, K. E., Franx, M., Bezanson, R., et al. 2015, ApJ, 811, L12
- Whitaker, K. E., Franx, M., Leja, J., et al. 2014, ApJ, 795, 104
- Willett, K. W., Schawinski, K., Simmons, B. D., et al. 2015, MNRAS, 449, 820
- Williams, R. J., Quadri, R. F., Franx, M., et al. 2010, ApJ, 713, 738
- Wright, E. L., Eisenhardt, P. R. M., Mainzer, A. K., et al. 2010, AJ, 140, 1868
- Wuyts, S., Förster Schreiber, N. M., Nelson, E. J., et al. 2013, ApJ, 779, 135
- Wuyts, S., Förster Schreiber, N. M., van der Wel, A., et al. 2011, ApJ, 742, 96
- Wuyts, S., Labbé, I., Franx, M., et al. 2007, ApJ, 655, 51
- Yang, X., Mo, H. J., van den Bosch, F. C., et al. 2007, ApJ, 671, 153
- Yuan, F., Lidman, C., Davis, T. M., et al. 2015, MNRAS, 452, 3047
- 1 Dipartimento di Fisica e Astronomia, Università di Bologna, Via Gobetti 93/2, 40129 Bologna, Italy
 - 2 INAF-Osservatorio di Astrofisica e Scienza dello Spazio di Bologna, Via Piero Gobetti 93/3, 40129 Bologna, Italy
 - 3 INAF-Osservatorio Astronomico di Padova, Via dell'Osservatorio 5, 35122 Padova, Italy
 - 4 Department of Astronomy, University of Geneva, ch. d'Ecogia 16, 1290 Versoix, Switzerland
 - 5 Max Planck Institute for Extraterrestrial Physics, Giessenbachstr. 1, 85748 Garching, Germany
 - 6 Universitäts-Sternwarte München, Fakultät für Physik, Ludwig-Maximilians-Universität München, Scheinerstrasse 1, 81679 München, Germany
 - 7 Université Paris-Saclay, Université Paris Cité, CEA, CNRS, AIM, 91191, Gif-sur-Yvette, France
 - 8 Instituto de Astrofísica de Canarias (IAC); Departamento de Astrofísica, Universidad de La Laguna (ULL), 38200, La Laguna, Tenerife, Spain
 - 9 Institute of Space Sciences (ICE, CSIC), Campus UAB, Carrer de Can Magrans, s/n, 08193 Barcelona, Spain
 - 10 Dipartimento di Fisica e Astronomia "G. Galilei", Università di Padova, Via Marzolo 8, 35131 Padova, Italy
 - 11 CEA Saclay, DFR/IRFU, Service d'Astrophysique, Bat. 709, 91191 Gif-sur-Yvette, France
 - 12 SRON Netherlands Institute for Space Research, Landleven 12, 9747 AD, Groningen, The Netherlands
 - 13 Kapteyn Astronomical Institute, University of Groningen, PO Box 800, 9700 AV Groningen, The Netherlands
 - 14 INAF-Osservatorio Astronomico di Capodimonte, Via Moiarriello 16, 80131 Napoli, Italy
 - 15 Instituto de Física de Cantabria, Edificio Juan Jordá, Avenida de los Castros, 39005 Santander, Spain
 - 16 Sterrenkundig Observatorium, Universiteit Gent, Krijgslaan 281 S9, 9000 Gent, Belgium
 - 17 Johns Hopkins University 3400 North Charles Street Baltimore, MD 21218, USA
 - 18 STAR Institute, University of Liège, Quartier Agora, Allée du six Août 19c, 4000 Liège, Belgium
 - 19 INAF-Istituto di Astrofisica e Planetologia Spaziali, via del Fosso del Cavaliere, 100, 00100 Roma, Italy
 - 20 Leiden Observatory, Leiden University, Einsteinweg 55, 2333 CC Leiden, The Netherlands
 - 21 Departamento de Física Teórica, Facultad de Ciencias, Universidad Autónoma de Madrid, 28049 Cantoblanco, Madrid, Spain
 - 22 Centro de Investigación Avanzada en Física Fundamental (CIAFF), Facultad de Ciencias, Universidad Autónoma de Madrid, 28049 Madrid, Spain
 - 23 Department of Physics and Astronomy, University of British Columbia, Vancouver, BC V6T 1Z1, Canada
 - 24 Instituto de Astrofísica de Canarias, Vía Láctea, 38205 La Laguna, Tenerife, Spain
 - 25 Dipartimento di Fisica e Astronomia "Augusto Righi" - Alma Mater Studiorum Università di Bologna, via Piero Gobetti 93/2, 40129 Bologna, Italy
 - 26 Université Paris-Saclay, CNRS, Institut d'astrophysique spatiale, 91405, Orsay, France
 - 27 ESAC/ESA, Camino Bajo del Castillo, s/n., Urb. Villafranca del Castillo, 28692 Villanueva de la Cañada, Madrid, Spain
 - 28 School of Mathematics and Physics, University of Surrey, Guildford, Surrey, GU2 7XH, UK
 - 29 INAF-Osservatorio Astronomico di Brera, Via Brera 28, 20122 Milano, Italy
 - 30 IFPU, Institute for Fundamental Physics of the Universe, via Beirut 2, 34151 Trieste, Italy
 - 31 INAF-Osservatorio Astronomico di Trieste, Via G. B. Tiepolo 11, 34143 Trieste, Italy
 - 32 INFN, Sezione di Trieste, Via Valerio 2, 34127 Trieste TS, Italy

- 33 SISSA, International School for Advanced Studies, Via Bonomea 265, 34136 Trieste TS, Italy
- 34 INFN-Sezione di Bologna, Viale Bertini Pichat 6/2, 40127 Bologna, Italy
- 35 Centre National d'Etudes Spatiales – Centre spatial de Toulouse, 18 avenue Edouard Belin, 31401 Toulouse Cedex 9, France
- 36 Space Science Data Center, Italian Space Agency, via del Politecnico snc, 00133 Roma, Italy
- 37 Dipartimento di Fisica, Università di Genova, Via Dodecaneso 33, 16146, Genova, Italy
- 38 INFN-Sezione di Genova, Via Dodecaneso 33, 16146, Genova, Italy
- 39 Department of Physics "E. Pancini", University Federico II, Via Cinthia 6, 80126, Napoli, Italy
- 40 Instituto de Astrofísica e Ciências do Espaço, Universidade do Porto, CAUP, Rua das Estrelas, PT4150-762 Porto, Portugal
- 41 Faculdade de Ciências da Universidade do Porto, Rua do Campo de Alegre, 4150-007 Porto, Portugal
- 42 Dipartimento di Fisica, Università degli Studi di Torino, Via P. Giuria 1, 10125 Torino, Italy
- 43 INFN-Sezione di Torino, Via P. Giuria 1, 10125 Torino, Italy
- 44 INAF-Osservatorio Astrofisico di Torino, Via Osservatorio 20, 10025 Pino Torinese (TO), Italy
- 45 European Space Agency/ESTEC, Keplerlaan 1, 2201 AZ Noordwijk, The Netherlands
- 46 Institute Lorentz, Leiden University, Niels Bohrweg 2, 2333 CA Leiden, The Netherlands
- 47 INAF-IASF Milano, Via Alfonso Corti 12, 20133 Milano, Italy
- 48 Centro de Investigaciones Energéticas, Medioambientales y Tecnológicas (CIEMAT), Avenida Complutense 40, 28040 Madrid, Spain
- 49 Port d'Informació Científica, Campus UAB, C. Albareda s/n, 08193 Bellaterra (Barcelona), Spain
- 50 Institute for Theoretical Particle Physics and Cosmology (TTK), RWTH Aachen University, 52056 Aachen, Germany
- 51 Institut d'Estudis Espacials de Catalunya (IEEC), Edifici RDIT, Campus UPC, 08860 Castelldefels, Barcelona, Spain
- 52 INAF-Osservatorio Astronomico di Roma, Via Frascati 33, 00078 Monteporzio Catone, Italy
- 53 INFN section of Naples, Via Cinthia 6, 80126, Napoli, Italy
- 54 Institute for Astronomy, University of Hawaii, 2680 Woodlawn Drive, Honolulu, HI 96822, USA
- 55 Dipartimento di Fisica e Astronomia "Augusto Righi" - Alma Mater Studiorum Università di Bologna, Viale Bertini Pichat 6/2, 40127 Bologna, Italy
- 56 Institute for Astronomy, University of Edinburgh, Royal Observatory, Blackford Hill, Edinburgh EH9 3HJ, UK
- 57 Jodrell Bank Centre for Astrophysics, Department of Physics and Astronomy, University of Manchester, Oxford Road, Manchester M13 9PL, UK
- 58 European Space Agency/ESRIN, Largo Galileo Galilei 1, 00044 Frascati, Roma, Italy
- 59 Université Claude Bernard Lyon 1, CNRS/IN2P3, IP2I Lyon, UMR 5822, Villeurbanne, F-69100, France
- 60 Institut de Ciències del Cosmos (ICCUB), Universitat de Barcelona (IEEC-UB), Martí i Franquès 1, 08028 Barcelona, Spain
- 61 Institució Catalana de Recerca i Estudis Avançats (ICREA), Passeig de Lluís Companys 23, 08010 Barcelona, Spain
- 62 UCB Lyon 1, CNRS/IN2P3, IUF, IP2I Lyon, 4 rue Enrico Fermi, 69622 Villeurbanne, France
- 63 Mullard Space Science Laboratory, University College London, Holmbury St Mary, Dorking, Surrey RH5 6NT, UK
- 64 Departamento de Física, Faculdade de Ciências, Universidade de Lisboa, Edifício C8, Campo Grande, PT1749-016 Lisboa, Portugal
- 65 Instituto de Astrofísica e Ciências do Espaço, Faculdade de Ciências, Universidade de Lisboa, Campo Grande, 1749-016 Lisboa, Portugal
- 66 INFN-Padova, Via Marzolo 8, 35131 Padova, Italy
- 67 Aix-Marseille Université, CNRS/IN2P3, CPPM, Marseille, France
- 68 INFN-Bologna, Via Irnerio 46, 40126 Bologna, Italy
- 69 School of Physics, HH Wills Physics Laboratory, University of Bristol, Tyndall Avenue, Bristol, BS8 1TL, UK
- 70 FRACTAL S.L.N.E., calle Tulipán 2, Portal 13 1A, 28231, Las Rozas de Madrid, Spain
- 71 Dipartimento di Fisica "Aldo Pontremoli", Università degli Studi di Milano, Via Celoria 16, 20133 Milano, Italy
- 72 INFN-Sezione di Milano, Via Celoria 16, 20133 Milano, Italy
- 73 NRC Herzberg, 5071 West Saanich Rd, Victoria, BC V9E 2E7, Canada
- 74 Institute of Theoretical Astrophysics, University of Oslo, P.O. Box 1029 Blindern, 0315 Oslo, Norway
- 75 Jet Propulsion Laboratory, California Institute of Technology, 4800 Oak Grove Drive, Pasadena, CA, 91109, USA
- 76 Department of Physics, Lancaster University, Lancaster, LA1 4YB, UK
- 77 Felix Hormuth Engineering, Goethestr. 17, 69181 Leimen, Germany
- 78 Technical University of Denmark, Elektrovej 327, 2800 Kgs. Lyngby, Denmark
- 79 Cosmic Dawn Center (DAWN), Denmark
- 80 Institut d'Astrophysique de Paris, UMR 7095, CNRS, and Sorbonne Université, 98 bis boulevard Arago, 75014 Paris, France
- 81 Max-Planck-Institut für Astronomie, Königstuhl 17, 69117 Heidelberg, Germany
- 82 NASA Goddard Space Flight Center, Greenbelt, MD 20771, USA
- 83 Department of Physics and Astronomy, University College London, Gower Street, London WC1E 6BT, UK
- 84 Department of Physics and Helsinki Institute of Physics, Gustaf Hällströmin katu 2, 00014 University of Helsinki, Finland
- 85 Université de Genève, Département de Physique Théorique and Centre for Astroparticle Physics, 24 quai Ernest-Ansermet, CH-1211 Genève 4, Switzerland
- 86 Department of Physics, P.O. Box 64, 00014 University of Helsinki, Finland
- 87 Helsinki Institute of Physics, Gustaf Hällströmin katu 2, University of Helsinki, Helsinki, Finland
- 88 Centre de Calcul de l'IN2P3/CNRS, 21 avenue Pierre de Coubertin 69627 Villeurbanne Cedex, France
- 89 Laboratoire d'étude de l'Univers et des phénomènes eXtremes, Observatoire de Paris, Université PSL, Sorbonne Université, CNRS, 92190 Meudon, France
- 90 Aix-Marseille Université, CNRS, CNES, LAM, Marseille, France
- 91 SKA Observatory, Jodrell Bank, Lower Withington, Macclesfield, Cheshire SK11 9FT, UK
- 92 Universität Bonn, Argelander-Institut für Astronomie, Auf dem Hügel 71, 53121 Bonn, Germany
- 93 INFN-Sezione di Roma, Piazzale Aldo Moro, 2 - c/o Dipartimento di Fisica, Edificio G. Marconi, 00185 Roma, Italy
- 94 Department of Physics, Institute for Computational Cosmology, Durham University, South Road, Durham, DH1 3LE, UK
- 95 Infrared Processing and Analysis Center, California Institute of Technology, Pasadena, CA 91125, USA
- 96 Université Côte d'Azur, Observatoire de la Côte d'Azur, CNRS, Laboratoire Lagrange, Bd de l'Observatoire, CS 34229, 06304 Nice cedex 4, France
- 97 Université Paris Cité, CNRS, Astroparticule et Cosmologie, 75013 Paris, France
- 98 CNRS-UCB International Research Laboratory, Centre Pierre Binétruy, IRL2007, CPB-IN2P3, Berkeley, USA
- 99 University of Applied Sciences and Arts of Northwestern Switzerland, School of Engineering, 5210 Windisch, Switzerland
- 100 Institut d'Astrophysique de Paris, 98bis Boulevard Arago, 75014, Paris, France
- 101 Institute of Physics, Laboratory of Astrophysics, Ecole Polytechnique Fédérale de Lausanne (EPFL), Observatoire de Sauverny, 1290 Versoix, Switzerland
- 102 Aurora Technology for European Space Agency (ESA), Camino bajo del Castillo, s/n, Urbanización Villafranca del Castillo, Villanueva de la Cañada, 28692 Madrid, Spain

- ¹⁰³ Institut de Física d'Altes Energies (IFAE), The Barcelona Institute of Science and Technology, Campus UAB, 08193 Bellaterra (Barcelona), Spain
- ¹⁰⁴ School of Mathematics, Statistics and Physics, Newcastle University, Herschel Building, Newcastle-upon-Tyne, NE1 7RU, UK
- ¹⁰⁵ DARK, Niels Bohr Institute, University of Copenhagen, Jagtvej 155, 2200 Copenhagen, Denmark
- ¹⁰⁶ Waterloo Centre for Astrophysics, University of Waterloo, Waterloo, Ontario N2L 3G1, Canada
- ¹⁰⁷ Department of Physics and Astronomy, University of Waterloo, Waterloo, Ontario N2L 3G1, Canada
- ¹⁰⁸ Perimeter Institute for Theoretical Physics, Waterloo, Ontario N2L 2Y5, Canada
- ¹⁰⁹ Institute of Space Science, Str. Atomistilor, nr. 409 Măgurele, Ilfov, 077125, Romania
- ¹¹⁰ Consejo Superior de Investigaciones Científicas, Calle Serrano 117, 28006 Madrid, Spain
- ¹¹¹ Universidad de La Laguna, Departamento de Astrofísica, 38206 La Laguna, Tenerife, Spain
- ¹¹² Caltech/IPAC, 1200 E. California Blvd., Pasadena, CA 91125, USA
- ¹¹³ Institut für Theoretische Physik, University of Heidelberg, Philosophenweg 16, 69120 Heidelberg, Germany
- ¹¹⁴ Institut de Recherche en Astrophysique et Planétologie (IRAP), Université de Toulouse, CNRS, UPS, CNES, 14 Av. Edouard Belin, 31400 Toulouse, France
- ¹¹⁵ Université St Joseph; Faculty of Sciences, Beirut, Lebanon
- ¹¹⁶ Departamento de Física, FCFM, Universidad de Chile, Blanco Encalada 2008, Santiago, Chile
- ¹¹⁷ Universität Innsbruck, Institut für Astro- und Teilchenphysik, Technikerstr. 25/8, 6020 Innsbruck, Austria
- ¹¹⁸ Satlantis, University Science Park, Sede Bld 48940, Leioa-Bilbao, Spain
- ¹¹⁹ Centre for Electronic Imaging, Open University, Walton Hall, Milton Keynes, MK7 6AA, UK
- ¹²⁰ Instituto de Astrofísica e Ciências do Espaço, Faculdade de Ciências, Universidade de Lisboa, Tapada da Ajuda, 1349-018 Lisboa, Portugal
- ¹²¹ Cosmic Dawn Center (DAWN)
- ¹²² Niels Bohr Institute, University of Copenhagen, Jagtvej 128, 2200 Copenhagen, Denmark
- ¹²³ Universidad Politécnica de Cartagena, Departamento de Electrónica y Tecnología de Computadoras, Plaza del Hospital 1, 30202 Cartagena, Spain
- ¹²⁴ Dipartimento di Fisica e Scienze della Terra, Università degli Studi di Ferrara, Via Giuseppe Saragat 1, 44122 Ferrara, Italy
- ¹²⁵ Istituto Nazionale di Fisica Nucleare, Sezione di Ferrara, Via Giuseppe Saragat 1, 44122 Ferrara, Italy
- ¹²⁶ INAF, Istituto di Radioastronomia, Via Piero Gobetti 101, 40129 Bologna, Italy
- ¹²⁷ Department of Physics, Oxford University, Keble Road, Oxford OX1 3RH, UK
- ¹²⁸ Université PSL, Observatoire de Paris, Sorbonne Université, CNRS, LERMA, 75014, Paris, France
- ¹²⁹ Université Paris-Cité, 5 Rue Thomas Mann, 75013, Paris, France
- ¹³⁰ INAF-Osservatorio Astronomico di Brera, Via Brera 28, 20122 Milano, Italy, and INFN-Sezione di Genova, Via Dodecaneso 33, 16146, Genova, Italy
- ¹³¹ ICL, Junia, Université Catholique de Lille, LITL, 59000 Lille, France
- ¹³² ICSC - Centro Nazionale di Ricerca in High Performance Computing, Big Data e Quantum Computing, Via Magnanelli 2, Bologna, Italy
- ¹³³ Instituto de Física Teórica UAM-CSIC, Campus de Cantoblanco, 28049 Madrid, Spain
- ¹³⁴ CERCA/ISO, Department of Physics, Case Western Reserve University, 10900 Euclid Avenue, Cleveland, OH 44106, USA
- ¹³⁵ Technical University of Munich, TUM School of Natural Sciences, Physics Department, James-Frank-Str. 1, 85748 Garching, Germany
- ¹³⁶ Max-Planck-Institut für Astrophysik, Karl-Schwarzschild-Str. 1, 85748 Garching, Germany
- ¹³⁷ Laboratoire Univers et Théorie, Observatoire de Paris, Université PSL, Université Paris Cité, CNRS, 92190 Meudon, France
- ¹³⁸ Departamento de Física Fundamental. Universidad de Salamanca. Plaza de la Merced s/n. 37008 Salamanca, Spain
- ¹³⁹ Université de Strasbourg, CNRS, Observatoire astronomique de Strasbourg, UMR 7550, 67000 Strasbourg, France
- ¹⁴⁰ Center for Data-Driven Discovery, Kavli IPMU (WPI), UTIAS, The University of Tokyo, Kashiwa, Chiba 277-8583, Japan
- ¹⁴¹ Ludwig-Maximilians-University, Schellingstrasse 4, 80799 Munich, Germany
- ¹⁴² Max-Planck-Institut für Physik, Boltzmannstr. 8, 85748 Garching, Germany
- ¹⁴³ Dipartimento di Fisica - Sezione di Astronomia, Università di Trieste, Via Tiepolo 11, 34131 Trieste, Italy
- ¹⁴⁴ NASA Ames Research Center, Moffett Field, CA 94035, USA
- ¹⁴⁵ Bay Area Environmental Research Institute, Moffett Field, California 94035, USA
- ¹⁴⁶ California Institute of Technology, 1200 E California Blvd, Pasadena, CA 91125, USA
- ¹⁴⁷ Department of Physics & Astronomy, University of California Irvine, Irvine CA 92697, USA
- ¹⁴⁸ Department of Mathematics and Physics E. De Giorgi, University of Salento, Via per Arnesano, CP-193, 73100, Lecce, Italy
- ¹⁴⁹ INFN, Sezione di Lecce, Via per Arnesano, CP-193, 73100, Lecce, Italy
- ¹⁵⁰ INFN-Sezione di Lecce, c/o Dipartimento Matematica e Fisica, Via per Arnesano, 73100, Lecce, Italy
- ¹⁵¹ Departamento Física Aplicada, Universidad Politécnica de Cartagena, Campus Muralla del Mar, 30202 Cartagena, Murcia, Spain
- ¹⁵² Institute of Cosmology and Gravitation, University of Portsmouth, Portsmouth PO1 3FX, UK
- ¹⁵³ Department of Computer Science, Aalto University, PO Box 15400, Espoo, FI-00 076, Finland
- ¹⁵⁴ Instituto de Astrofísica de Canarias, c/ Via Lactea s/n, La Laguna 38200, Spain. Departamento de Astrofísica de la Universidad de La Laguna, Avda. Francisco Sanchez, La Laguna, 38200, Spain
- ¹⁵⁵ Ruhr University Bochum, Faculty of Physics and Astronomy, Astronomical Institute (AIRUB), German Centre for Cosmological Lensing (GCCL), 44780 Bochum, Germany
- ¹⁵⁶ Department of Physics and Astronomy, Vesilinnantie 5, 20014 University of Turku, Finland
- ¹⁵⁷ Serco for European Space Agency (ESA), Camino bajo del Castillo, s/n, Urbanizacion Villafranca del Castillo, Villanueva de la Cañada, 28692 Madrid, Spain
- ¹⁵⁸ ARC Centre of Excellence for Dark Matter Particle Physics, Melbourne, Australia
- ¹⁵⁹ Centre for Astrophysics & Supercomputing, Swinburne University of Technology, Hawthorn, Victoria 3122, Australia
- ¹⁶⁰ Department of Physics and Astronomy, University of the Western Cape, Bellville, Cape Town, 7535, South Africa
- ¹⁶¹ DAMTP, Centre for Mathematical Sciences, Wilberforce Road, Cambridge CB3 0WA, UK
- ¹⁶² Kavli Institute for Cosmology Cambridge, Madingley Road, Cambridge, CB3 0HA, UK
- ¹⁶³ Department of Astrophysics, University of Zurich, Winterthurerstrasse 190, 8057 Zurich, Switzerland
- ¹⁶⁴ Department of Physics, Centre for Extragalactic Astronomy, Durham University, South Road, Durham, DH1 3LE, UK
- ¹⁶⁵ IRFU, CEA, Université Paris-Saclay 91191 Gif-sur-Yvette Cedex, France
- ¹⁶⁶ Oskar Klein Centre for Cosmoparticle Physics, Department of Physics, Stockholm University, Stockholm, SE-106 91, Sweden
- ¹⁶⁷ Astrophysics Group, Blackett Laboratory, Imperial College London, London SW7 2AZ, UK
- ¹⁶⁸ Univ. Grenoble Alpes, CNRS, Grenoble INP, LPSC-IN2P3, 53, Avenue des Martyrs, 38000, Grenoble, France
- ¹⁶⁹ INAF-Osservatorio Astrofisico di Arcetri, Largo E. Fermi 5, 50125, Firenze, Italy

- ¹⁷⁰ Dipartimento di Fisica, Sapienza Università di Roma, Piazzale Aldo Moro 2, 00185 Roma, Italy
- ¹⁷¹ Centro de Astrofísica da Universidade do Porto, Rua das Estrelas, 4150-762 Porto, Portugal
- ¹⁷² HE Space for European Space Agency (ESA), Camino bajo del Castillo, s/n, Urbanización Villafranca del Castillo, Villanueva de la Cañada, 28692 Madrid, Spain
- ¹⁷³ Department of Astrophysical Sciences, Peyton Hall, Princeton University, Princeton, NJ 08544, USA
- ¹⁷⁴ Theoretical astrophysics, Department of Physics and Astronomy, Uppsala University, Box 515, 751 20 Uppsala, Sweden
- ¹⁷⁵ Minnesota Institute for Astrophysics, University of Minnesota, 116 Church St SE, Minneapolis, MN 55455, USA
- ¹⁷⁶ Mathematical Institute, University of Leiden, Einsteinweg 55, 2333 CA Leiden, The Netherlands
- ¹⁷⁷ School of Physics & Astronomy, University of Southampton, Highfield Campus, Southampton SO17 1BJ, UK
- ¹⁷⁸ Institute of Astronomy, University of Cambridge, Madingley Road, Cambridge CB3 0HA, UK
- ¹⁷⁹ Department of Physics and Astronomy, University of California, Davis, CA 95616, USA
- ¹⁸⁰ Space physics and astronomy research unit, University of Oulu, Pentti Kaiteran katu 1, FI-90014 Oulu, Finland
- ¹⁸¹ Center for Computational Astrophysics, Flatiron Institute, 162 5th Avenue, 10010, New York, NY, USA