The multichord stellar occultation on 2019 October 22 by the trans-Neptunian object (84922) 2003 VS₂


(Affiliations can be found after the references)

Received date / Accepted date

ABSTRACT

Context. Stellar occultations have become one of the best techniques to gather information about the physical properties of trans-Neptunian objects (TNOs), which are critical objects for understanding the origin and evolution of our Solar System.

Aims. The purpose of this work is to determine, with better accuracy, the physical characteristics of the TNO (84922) 2003 VS₂ through the analysis of the multichord stellar occultation on 2019 October 22 and photometric data collected afterward.

Methods. We predicted, observed, and analyzed the multichord stellar occultation of the Second Gaia Data Release (Gaia DR2) source 3449076721168026624 (m₂ = 14.1 mag) by the plutino object 2003 VS₂ on 2019 October 22. We performed aperture photometry on the images collected and derived the times when the star disappeared and reappeared from the observing sites that reported a positive detection. We fit the extremities of such positive chords to an ellipse using a Monte Carlo method. We also carried out photometric observations to derive the rotational light curve amplitude and rotational phase of 2003 VS₂ during the stellar occultation. Combining the results and assuming a triaxial shape, we derived the 3D shape of 2003 VS₂.

Results. Out of the 39 observatories involved in the observational campaign, 12 sites, located in Bulgaria (one), Romania (ten), and Serbia (one), reported a positive detection; this makes it one of the best observed stellar occultations by a TNO so far. Considering the rotational phase of 2003 VS₂ during the stellar occultation and the rotational light curve amplitude derived (Δm = 0.264 ± 0.017 mag), we obtained a mean area-equivalent diameter of Dₘₐᵉ = 545 ± 13 km and a geometric albedo of 0.134 ± 0.010. By combining the rotational light curve information with the stellar occultation results, we derived the best triaxial shape for 2003 VS₂, which has semiaxes a = 339 ± 5 km, b = 235 ± 6 km, and c = 226 ± 8 km. The derived aspect angle of 2003 VS₂ is θ = 59° ± 2° or its supplementary θ = 121° ± 2°, depending on the north-pole position of the TNO. The spherical-volume equivalent diameter is Dᵥₑₐᵉ = 524 ± 7 km. If we consider large albedo patches on its surface, the semi-major axis of the ellipsoid could be ~10 km smaller. These results are compatible with the previous ones determined from the single-chord 2013 and four-chord 2014 stellar occultations and with the effective diameter and albedo derived from Herschel and Spitzer data. They provide evidence that 2003 VS₂’s 3D shape is not compatible with a homogeneous triaxial body in hydrostatic equilibrium, but it might be a differentiated body and/or might be sustaining some stress. No secondary features related to rings or material orbiting around 2003 VS₂ were detected.

Key words. Kuiper belt objects: individual: 2003 VS₂ – Methods: observational – Techniques: photometric

1. Introduction

Trans-Neptunian objects (TNOs) offer a unique opportunity to better understand the origins and the chemical, dynamical, and collisional evolution of the outer Solar System. Possibly dispatched to further distances from the Sun than that of Neptune due to gravitational perturbations after their formation (Gomes et al. 2005; Levison et al. 2008), their global composition has been virtually unaffected by solar irradiation, keeping it very similar to that of the primitive nebula (Morbidelli et al. 2008).

In the last decade, stellar occultations by TNOs have proved to be one of the best techniques to determine the size and shape of these objects, to show features such as satellites (Siccafoose et al. 2019) or rings (Braga-Ribas et al. 2014; Ortiz et al. 2015, 2017), and to reveal possible atmospheres (Hubbard et al. 1988; Stern & Trafford 2008; Meza et al. 2019). If we combine this technique with light-reflection measurements, we can also derive their geometric albedo. Furthermore, we can calculate their density if we assume hydrostatic equilibrium or if the body is part of a binary system or has a satellite, in which case we can obtain its mass.

Although this technique has been increasingly used in the past few years, it is still challenging to predict and obtain positive results from multichord stellar occultations (Ortiz et al. 2020), mainly due to the uncertainties in the orbits of TNOs. The Second Gaia Data Release (Gaia DR2; Gaia Collaboration et al. 2016b, 2018) eases the process by providing peerless accurate positions and proper motions of more than one billion sources. However, the large orbital periods of TNOs and the short time
span of the observations result in non-negligible uncertainties in
their orbital elements, making astrometric observations close to
the stellar occultation date indispensable for a successful event
prediction.

The object (84922) 2003 VS₂ orbits in the 3:2 mean motion
resonance (MMR) with Neptune, which makes it a plutino. This
TNO presents a double-peaked rotational light curve with an
accurately determined rotation period of 7.41753 ± 0.00001 h
(Santos-Sanz et al. 2017). No satellites nor secondary features
have been discovered so far around 2003 VS₂. Its near-infrared
spectrum reveals the presence of exposed water ice (Barkume
et al. 2008) which, according to Monnert et al. (2012), might
explain the increase in 2003 VS₂’s albedo from the typical value
of 0.07 for the plutinos. The orbital elements and the most rele-
vant physical characteristics of 2003 VS₂ can be found in Table
[1]

A recent work was published with results from a four-chord
and two single-chord stellar occultations by 2003 VS₂ in 2013
and 2014 (Benedetti-Rossi et al. 2019). In that work, the au-
tors reconstructed the 3D shape of 2003 VS₂ by combining the
data from the multichord stellar occultation and the rotational
light curve obtained. The principal semiaxes that provided the
best fit had values of $a = 313.8 \pm 7.1$ km, $b = 265.5^{+8.4}_{-8.8}$ km, and
$c = 247.3^{+26.0}_{-26.3}$ km, with these measurements being inconsistent
with a Jacobi ellipsoid (Chandrasekhar 1987). This solution was
highly affected by the rotational light curve amplitude derived
in that paper, which was 0.141 ± 0.009 mag and significantly
smaller than the previously moderately large ones reported by
Ortiz et al. (2006), Sheppard (2007), and Thirouin et al. (2010).
In this work we show that such a rotational light curve ampli-
tude cannot be correct based on new data and other reasoning.
Benedetti-Rossi et al. (2019) also reported the presence of a put-
tative secondary feature detected from one of the observing sta-
tions but, without further data, they could not discard the pos-
sibility that it could be due to a companion star or instrumen-
tal effects.

In this we work we present the results of the multichord stellar
occultation by 2003 VS₂ on 2019 October 22. A total of 39 ob-
servations were involved in the campaign, from which 12 posi-
tive chords were obtained, which is a considerable improvement
with respect to the previous occultations by this TNO. Com-
bining the collected data from the stellar occultation and en-
suing photometric measurements, we derived the 3D shape of
2003 VS₂.

2. Observations

This section presents the observations carried out to improve the
prediction of the stellar occultation, the observations of the ac-
tual event, and the observations performed shortly afterward to
obtain the rotational light curve of 2003 VS₂.

2.1. Occultation predictions

The stellar occultation on 2019 October 22 was singled out from
the systematic searches for TNO occultation candidate stars
carried out by the European Research Council (ERC) Lucky
Star¹ project collaboration (Desmurs et al. 2015, 2018). The
Lucky Star’s NIMA² ephemeris thus gave the initial predic-
tion³ for the 2019 October 22 stellar occultation, and the can-
didate star was identified in the Gaia DR2 catalog (source ID:
3449076221168026496; UCAC4 identifier 616-023624). Rele-
vant information about the star, such as its coordinates, proper
motions, parallax, and their uncertainties, as well as the star’s $G,
B, V,$ and $K$ magnitudes, can be found in Table 2.

To reduce the uncertainty on 2003 VS₂’s orbit and narrow
down the predicted location of the shadow path, we performed
two observing runs with two different telescopes a few days be-
fore the event. The first observing run was carried out on 2019
October 5, with the charge-coupled device (CCD) Andor ikon-
L camera of the 1.5-m telescope at the Sierra Nevada Observ-
atory (OSN) in Granada, Spain. This camera provides a field
of view (FOV) of 792 × 792, with an image scale of 0.232
pixel⁻¹. The 15 images of this set were acquired in 2x2 binning
mode, with no filter and an integration time of 400 s. The aver-
age seeing was 1′.84. Bias and flat-field frames were taken for
standard calibration, which was done afterward, following the
steps in Fernández-Valenzuela et al. (2016).

The second run was taken on 2019 October 8, with the IO:O
camera of the Liverpool 2-m Telescope at the Roque de los
Muchachos Observatory in La Palma, Spain. This instrument has
a FOV of 10′×10′ and an image scale of 0.′15 pixel⁻¹. The ten
images of this set were acquired in 2x2 binning mode, with the
Sloan R-filter and 300 s of integration time. The average seeing
was 1′.92. Bias and flat-field frames were also obtained for stan-
dard calibration.

With the OSN data, the obtained offsets with respect to the
Jet Propulsion Laboratory (JPL) #30 orbit were (−360 ± 36) mas
in right ascension (RA) and (+4 ± 25) mas in declination (Dec).
The Liverpool data yielded offsets of (−368 ± 12) mas in RA
and (+24 ± 11) mas in Dec.

We made two different predictions using the OSN and the
Liverpool data, although the result was roughly identical (the
Liverpool prediction being ~ 20 mas north of the OSN). We
used the predicted shadow path obtained by updating 2003 VS₂’s
ehmeris from JPL with the Liverpool data (shown in Fig. 1),
as it was taken closer to the event date and with better seeing.

2.2. Stellar occultation observations

On 2019 October 22, 39 observational stations spread across
11 countries (including both professional telescopes and ama-
teur observers, Fig. 1) were all set to observe Gaia DR2 source
3449076221168026496 around the predicted occultation time.
Out of the total participating stations, 12 reported a positive
detection, 14 reported a negative detection (among which two were
very close to the shadow path and provide constraints to the
shape fitting), and 13 could not observe due to bad weather or
technical difficulties. Detailed information about all the partici-
patant observatories, divided between positive detections, negative
detections, and observations with technical problems or overcast,
can be found in Tables 3, 4, and A.1, respectively. The detailed
instrumental settings at the sites with positive detections are also
given in Table 3. All the teams that could observe collected se-
ries of flexible image transport system (FITTS) images, except for
one site in Romania that took data in tagged image file format
(TIFF) and required a different analysis, as explained in Sect.
3.1. The time span of the collected images includes several min-
utes before and after the stellar occultation.

In this kind of event, it is crucial to save, within the images’
header, the individual acquisition time and all participating sta-
tions must be synchronized. Clock synchronizations were used

¹ The fourth U.S. Naval Observatory CCD Astrograph Catalog.
² https://lesia.obspm.fr/lucky-star/index.php
³ Numerical Integration of the Motion of an Asteroid.
Table 1: Orbital elements and physical characteristics of 2003 VS$_2$.  

<table>
<thead>
<tr>
<th></th>
<th>(au)</th>
<th>(au)</th>
<th>(°)</th>
<th>(°)</th>
<th>(h)</th>
<th>(mas)</th>
<th>(mas)</th>
<th>(km)</th>
<th>(km)</th>
<th></th>
</tr>
</thead>
<tbody>
<tr>
<td>$a$</td>
<td>39.249</td>
<td>36.386</td>
<td>0.073</td>
<td>14.83</td>
<td>4.14±0.07</td>
<td>7.41753±0.00001</td>
<td>0.141±0.009</td>
<td>553$^{+36}_{-33}$</td>
<td>548.3$^{+29.5}_{-44.6}$</td>
<td>0.131$^{+0.024}_{-0.013}$</td>
</tr>
</tbody>
</table>

Notes. Orbital elements of 2003 VS$_2$ from the Jet Propulsion Laboratory (JPL) Small-Body Database Browser (https://ssd.jpl.nasa.gov/sslcb.cgi); $a$, semi-major axis; $q$, perihelion distance; $e$, eccentricity; and $i$, inclination. Absolute magnitude $H_i$ is from Santos-Sanz et al. (2017). The rotational light curve amplitude ($\Delta m$), spherical volume-equivalent diameter ($D_{\text{area eq}}$), and albedo ($p_V$) are from Benedetti-Rossi et al. (2019). The mean area-equivalent diameter ($D_{\text{area eq}}$) was calculated from the values given in Benedetti-Rossi et al. (2019) and taking into account the rotational phase using Eq. [1]. Mommert et al. (2012) obtained an area-equivalent diameter of $D = 523\pm34.3$ km and ($b$) an albedo of $p_V = 0.147\pm0.033$.

Table 2: Main information of the occulted star (Gaia DR2 3449076721168026496$^a$; UCAC4 identifier 616-023624$^a$).

<table>
<thead>
<tr>
<th>RA (ICRF$^a$)</th>
<th>Dec (ICRF$^a$)</th>
<th>pmRA</th>
<th>pmDec</th>
<th>$\pi$</th>
<th>$G$</th>
<th>$B$</th>
<th>$V$</th>
<th>$K$</th>
</tr>
</thead>
<tbody>
<tr>
<td>05h 30m 38.0442s</td>
<td>+33° 07′ 24″</td>
<td>0.0401</td>
<td>+33° 07′ 74″</td>
<td>0.0330</td>
<td>-0.34±0.06</td>
<td>-1.76±0.05</td>
<td>0.2084±0.0353</td>
<td>14.1625</td>
</tr>
</tbody>
</table>

Notes. Star coordinates in RA and Dec were propagated to the occultation epoch (2458779.35870) and respective errors (errRA, errDec), proper motion in RA and Dec and respective errors, absolute stellar parallax with error, and magnitude $G$ from Gaia DR2 (Gaia Collaboration et al. 2016b$^{b}$, 2018$^{b}$). Magnitudes $B, V, K$ are from the Naval Observatory Merged Astrometric Dataset (NOMAD) catalog (Huchra et al. 2003; Lindegren et al. 2016). We used several apertures and annuli and chose those that resulted in the least scatter in the photometry. We show the resulting normalized flux (the blended flux of the occulted star plus TNO divided by its mean value outside the occultation moment) versus the time from the positive sites, which show deep drops in flux around the predicted occultation time. We note that 2003 VS$_2$ is too faint to be seen with the available equipment, so the star’s flux drops to zero during the occultation. We derived the error bars in Fig. 2$^c$ from Poisson noise calculations, but the results were scaled so that their standard deviation matches that of the data outside the main drop in occultation (column seven in Table 5$^d$). We do not see secondary flux drops other than the one corresponding to the main body, which could indicate that there is not a sufficiently wide and dense ring orbiting around 2003 VS$_2$ that could have produced detectable flux drops; however, due to the small signal-to-noise ratio (S/N), we cannot discard it. We took special care when performing the photometry of the data from the places that missed the occultation but were very close to the shadow path, as they put constraints to 2003 VS$_2$’s final shape.

2.3. Rotational light curve observations

Given that 2003 VS$_2$ has a double-peaked rotational light curve and is large enough to allow a hydrostatic equilibrium shape (see a detailed explanation in Benedetti-Rossi et al. 2019), we can assume that the body is a triaxial ellipsoid (Chandrasekhar 1987). If this were the case, one would need at least three body projections at different rotational phases to correctly derive the actual 3D shape of 2003 VS$_2$.

However, we can overcome this by combining the occultation information with time-series photometric data taken closer to the occultation date. To do so, we observed 2003 VS$_2$ for two consecutive nights, two days after the event (2019 October 24 and 25), with the 1.23-m telescope at the Calar Alto Observatory in Almeria (Spain). This telescope is equipped with a 4k x 4k CCD DLR-MKIII camera, which provides a FOV of 21.5′×21.5′.
Table 3: Details of the observing stations of the 2003 VS₃ multichord stellar occultation on 2019 October 22 with positive detection.

<table>
<thead>
<tr>
<th>Chord number</th>
<th>Site name</th>
<th>Location</th>
<th>Latitude (N)</th>
<th>Longitude (E)</th>
<th>Telescope Aperture (cm)</th>
<th>Detector/Instrument</th>
<th>Exposure Time Cycle Time (s)</th>
<th>Observers</th>
</tr>
</thead>
<tbody>
<tr>
<td>1</td>
<td>ROASTERR-1 Observatory (1) Romania L04</td>
<td>46° 49' 15''/6 23° 35' 47''</td>
<td>391°</td>
<td>30</td>
<td>KAF8300</td>
<td>3</td>
<td>5.18</td>
<td>L. Hudin</td>
</tr>
<tr>
<td>2</td>
<td>ROASTERR-1 Observatory (2) Romania L04</td>
<td>46° 49' 15''/6 23° 35' 47''</td>
<td>391°</td>
<td>28</td>
<td>ATIK-414ex</td>
<td>2</td>
<td>3.01</td>
<td>M. Boaca V. Inceu</td>
</tr>
<tr>
<td>3</td>
<td>Romania –</td>
<td>46° 42' 37''/557 23° 35' 35''</td>
<td>391°</td>
<td>30.5</td>
<td>SBIG STT-1603ME</td>
<td>2</td>
<td>3</td>
<td>V. Turcu</td>
</tr>
<tr>
<td>4</td>
<td>Romania –</td>
<td>46° 42' 37''/554 23° 35' 35''</td>
<td>783.405</td>
<td>40.6</td>
<td>SBIG STT-1603ME</td>
<td>2</td>
<td>3</td>
<td>D. Moldovan L. Mircea</td>
</tr>
<tr>
<td>5</td>
<td>Romania –</td>
<td>45° 42' 11''/94 21° 26' 12''</td>
<td>92°</td>
<td>27.94</td>
<td>ZWO ASI 224 CMOS</td>
<td>3</td>
<td>5.14</td>
<td>L. Stoian A. Juravle</td>
</tr>
<tr>
<td>6</td>
<td>Berthelot Observatory Romania L54</td>
<td>45° 36' 59''/4 22° 53' 19''</td>
<td>400</td>
<td>36.83</td>
<td>SBIG STL11000M CCD</td>
<td>2</td>
<td>8.64</td>
<td>A. Senedcu A. Sonka</td>
</tr>
<tr>
<td>7</td>
<td>Stardust Observatory Romania L13</td>
<td>45° 38' 30''/7 25° 37' 19''</td>
<td>597</td>
<td>20.32</td>
<td>CCD Atik 383L</td>
<td>3</td>
<td>10.28</td>
<td>L. Curelalu</td>
</tr>
<tr>
<td>8</td>
<td>Romania –</td>
<td>44° 42' 24''/5 23° 47' 19''</td>
<td>109°</td>
<td>33.6</td>
<td>ASI 1600</td>
<td>3</td>
<td>3.01</td>
<td>M. Predatu</td>
</tr>
<tr>
<td>9</td>
<td>Stardreams Observatory Romania L16</td>
<td>45° 12' 13''/3 26° 02' 44''</td>
<td>382°</td>
<td>20</td>
<td>ATIK 460ex</td>
<td>4</td>
<td>5.52</td>
<td>R. Gherase</td>
</tr>
<tr>
<td>10</td>
<td>Astronomical Station Vidojevica Serbia C89</td>
<td>43° 08' 24''/6 21° 33' 20''</td>
<td>1150</td>
<td>140</td>
<td>Andor ikonL</td>
<td>0.8</td>
<td>1.55</td>
<td>O. Ilic</td>
</tr>
<tr>
<td>11</td>
<td>Belogradchik Observatory Bulgaria –</td>
<td>43° 37' 22''/7 22° 40' 30''</td>
<td>650</td>
<td>60</td>
<td>FLI PL9000</td>
<td>6</td>
<td>6.86</td>
<td>E. Semkov</td>
</tr>
<tr>
<td>12</td>
<td>Romania –</td>
<td>44° 19' 19''/14 25° 59' 8''</td>
<td>70°</td>
<td>35.5</td>
<td>ASI 1600</td>
<td>2</td>
<td>2.50</td>
<td>M. Teodorescu</td>
</tr>
</tbody>
</table>

Notes. The sites are sorted by their distance to the center of the predicted shadow path.

Altitudes are from Google Earth. The cycle time is the sum of the exposure time plus dead time.

with an image scale of 0′′.32 pixel⁻¹. This wide FOV allowed us to keep the same stellar field both nights, making it possible to choose the same reference stars during the run to minimize systematic photometric errors.

A total of 143 images were acquired in 2×2 binning mode and with an integration time of 400s, using no filter to maximize the S/N. The moonshine was 14% for the first night and 7% for the second night. The average seeing at Calar Alto from the Differential Image Motion monitor was 1′′.6 for the second night; unfortunately, the Calar Alto seeing tracker did not work for the first night. The average measured full width at half maximum (FWHM) was ∼ 2.5 pixels (≈ 1′′.6) for the first night and ∼ 2.3 pixels (≈ 1′′.5) for the second. The S/N was between 30 and 70 during the first night and between 20 and 80 during the second night. We also took bias and flat-field frames each night for standard calibration, which was performed using our own specific routines written in IDL.

**3. Data reduction**

**3.1. Time synchronization and extraction**

Robust clock synchronization is essential to have the absolute acquisition time written on each frame header to obtain the actual star disappearance and reappearance times from each site and, hence, the projected chords’ relative position. In this regard, some of the images collected required a different treatment to obtain their absolute acquisition time, as we detail below. We saw no evidence for wrong time synchronization for the remaining positive chords, given their relative positions (see Fig. 5a).

**Chords 3 and 4**

There was an intentionally applied 1 s difference between the images acquired from the two telescopes at this site, aiming to cover the whole event avoiding the dead time of the CCDs. Since the two telescopes used the same instrument and took FITS im-
Table 4: Details of the observing stations of the 2003 VS$_2$ multichord stellar occultation on 2019 October 22 with negative detection.

<table>
<thead>
<tr>
<th>Site name</th>
<th>Location</th>
<th>Latitude (N)</th>
<th>Longitude (E)</th>
<th>Telescope aperture (cm)</th>
<th>Observers</th>
</tr>
</thead>
<tbody>
<tr>
<td>Bremthal</td>
<td>Germany</td>
<td>50° 08' 17.4''</td>
<td>08° 21' 50.4''</td>
<td>25.4</td>
<td>O. Klös</td>
</tr>
<tr>
<td>PDlink Observatory</td>
<td>Cadca, Slovakia</td>
<td>49° 24' 15.2''</td>
<td>18° 42' 09.47''</td>
<td>40</td>
<td>P. Delincak</td>
</tr>
<tr>
<td>M. Suhora Observatory</td>
<td>Poland</td>
<td>49° 34' 09''</td>
<td>20° 04' 03''</td>
<td>60</td>
<td>M. Drozdz</td>
</tr>
<tr>
<td>Skalnate Pleso</td>
<td>Slovakia</td>
<td>49° 11' 21''</td>
<td>20° 14' 02''</td>
<td>130</td>
<td>M. Husárik</td>
</tr>
<tr>
<td>Massa, Italy</td>
<td></td>
<td>44° 01' 16''</td>
<td>10° 07' 56''</td>
<td>30</td>
<td>P. Baruffetti</td>
</tr>
<tr>
<td>Gothard Observatory</td>
<td>Szombathely, Hungary</td>
<td>47° 15' 29''</td>
<td>16° 36' 15.6''</td>
<td>80</td>
<td>G. M. Szabó</td>
</tr>
<tr>
<td>Cmi Vrh Observatory</td>
<td>Slovenia</td>
<td>45° 56' 45''</td>
<td>14° 04' 16''</td>
<td>60</td>
<td>J. Skvarč</td>
</tr>
<tr>
<td>Monte Agliale Observatory</td>
<td>Borgo a Mozzano, Italy</td>
<td>43° 59' 43''</td>
<td>10° 30' 53''</td>
<td>50</td>
<td>F. Ciabattari</td>
</tr>
<tr>
<td>Stazione Osservativa di Basovizza</td>
<td>Osservatorio Astronomico di Trieste, Italy</td>
<td>45° 38' 33''</td>
<td>13° 52' 23''</td>
<td>35</td>
<td>P. Di Marcantonio</td>
</tr>
<tr>
<td>Piszkestető Station</td>
<td>Konkoly Obs., Hungary</td>
<td>47° 55' 6''</td>
<td>19° 53' 41''</td>
<td>100</td>
<td>A. Pal</td>
</tr>
<tr>
<td>Konkoly Observatory</td>
<td>Normafa, Hungary</td>
<td>47° 29' 59''</td>
<td>18° 57' 51''</td>
<td>60</td>
<td>A. Pal</td>
</tr>
<tr>
<td>Vratnik, Croatia</td>
<td>Romania</td>
<td>49° 24' 15''</td>
<td>18° 42' 09''</td>
<td>20</td>
<td>H. Mikuž</td>
</tr>
<tr>
<td>Mount Etna Observatory</td>
<td>Osservatorio Astrofisico di Catania, Italy</td>
<td>37° 41' 05''</td>
<td>14° 58' 04''</td>
<td>91</td>
<td>J. Alonso-Santiago</td>
</tr>
</tbody>
</table>

Notes. The sites are sorted by their distance to the center of the predicted shadow path.

(a) Altitudes are from Google Earth.

ages with the same exposure time, we merged the data from these two telescopes to obtain one single chord. By doing this, we doubled the sampling of the light curve and, therefore, we obtained the immersion and emersion times from this site with a smaller uncertainty. Hereafter we refer to this as a single chord (chord 3').

Chord 8

Although this chord nearly overlapped with chord 9, its center was shifted by more than 12 s with respect to the linear fit of the rest of the centers (see Fig. 2 and Table 5). Later tests performed a month after the occultation event showed a difference of more than one minute between the acquisition system and a synchronized clock, but we were not able to determine the exact time offset during the occultation event. All of this suggested that the time synchronization at this site was not correctly applied, so we could only use the chord length and not the absolute time of the chord.

Chord 12

The images from this site were recorded in a data cube in TIFF format. We converted them to FITS format using the Planetary
3.2. Occultation times and chord lengths

From every site that reported a positive detection, the stellar occultation starting and ending times (disappearance and reappearance of the star, respectively) were determined by fitting each light curve to a sharp-edge square-well model convolved by the following: (1) Fresnel diffraction by a point-like source at the distance of 2003 VS$_2$ from the observer, (2) the CCD bandwidth, (3) the finite stellar diameter projected at the object’s distance, in kilometers, and (4) the finite integration time (see, e.g., Wide-mann et al. [2009] Braga-Ribas et al. [2013]).

During the stellar occultation, 2003 VS$_2$ was at a geocentric distance of $\Delta = 36.1$ au$=5.40\times10^6$ km. This implies that for a typical wavelength of $\lambda=0.65$ $\mu$m, the Fresnel scale $F=\sqrt{\Delta D}/2$ has a value of 1.33 km.

We estimated the projected stellar diameter using the formulae from van Belle (1999) and its $B$, $V$, and $K$ apparent magnitudes, obtained from the NOMAD catalog (Zacharias et al. 2004 Table[2]). The diameters obtained are 0.86 km if considering a super giant or 0.91 km if considering a main sequence star.

Finally, we converted the shortest integration time among the positive observations (0.8 s, chord 10) into the distance in the sky-plane traveled by 2003 VS$_2$ between two adjacent data points; since the velocity of 2003 VS$_2$’s shadow path was 13.01 km s$^{-1}$, that distance was 10.41 km. As a result, our light curves are mainly dominated by the integration times and not by Fresnel diffraction or the stellar diameter.

The fitting procedure then searches for the times of disappearance and reappearance of the star that minimize a classical $\chi^2$ function, as explained in Sicardy et al. (2011 supplementary information). The uncertainty bars were estimated by varying the occultation times to increase $\chi^2$ to $\chi^2+1$. Since these uncertainty bars depend on the fit of the photometry and as the typical errors in the photometry are very small, the uncertainty in the derived time of ingress and egress is also small, that is to say a small fraction of the integration time. Figure[3] shows an example of the best fit to one of the stellar occultation light curves (chord 2). All the derived disappearance and reappearance times and chord lengths are listed in Table[5].

3.3. Rotational light curve

We performed aperture photometry on 2003 VS$_2$ and 16 reference stars with good photometric behavior using the routines and procedures mentioned in Fernández-Valenzuela et al. [2016]. We tested different synthetic aperture radii and sky annuli to maximize the S/N of the object while minimizing the dispersion of the residuals to the rotational light curve fit. We chose different aperture parameters for each observing day, selecting the same reference stars, and then combined the best photometry results.

The obtained rotational light curve is shown in Fig[4] and the data used are available online. We folded the relative flux of 2003 VS$_2$ versus time using its well-known rotation period of $7.41753 \pm 0.00001$ h (Santos-Sanz et al. 2017), and we used a fourth-order Fourier function to fit the folded data. The obtained peak-to-peak amplitude of the rotational light curve was $\Delta m = 0.264 \pm 0.017$ mag; the nominal value is the one that best fits the data in terms of minimization of the sum of squared residuals and the uncertainty is given as the standard deviation of a Monte Carlo distribution. This amplitude is larger but consistent within the error bars with other values found in the literature ($0.23 \pm 0.07$ mag, $0.21 \pm 0.02$ mag, $0.21 \pm 0.03$ mag; Ortiz et al. [2006] Sheppard [2007] Thirouin et al. [2010], respectively). However, the greatest disagreement appears when com-

Fig. 2: Normalized light curves from the positive detections of the stellar occultation by 2003 VS$_2$ on 2019 October 22 and the two closest negatives. The relative flux of the occulted star with respect to the comparison chosen stars is plotted against time, given in seconds after 2019 October 22 20:40:00 UT. The uncertainty bars of the flux were plotted for all the chords, although some have the size of the points and are not visible. The light curves have been displaced in flux for better visualization and they follow the same order as in Table[3]. Chords 0 and 13 (top and bottom chords, in gray) correspond to the negative detections from observers H. Mikuž and V. Dumitrescu, respectively (see Table[4]). Light curves plotted in blue required special consideration regarding time synchronization, see Sect. [5]. Chord 3’ is the result of merging chords 3 and 4, see the text for details.

Imaging PreProcessor (PIPP), but then all of the images had the same time written on the header. The observer provided us with the acquisition times of the images, but their accuracy only reached the order of seconds (as this is the accuracy reached by the software used to save the images). We used the mid-integration times for the light curve and then increased the uncertainties of the immersion and emersion times by the exposure time (two seconds) since it was not possible to know if the acquisition times we have corresponded to the beginning, middle, or end of the images’ integration.
Table 5: Star disappearance and reappearance UT times on 2019 October 22, chord lengths, time shifts, and dispersion ($\sigma$) of the light curves outside the occultation.

<table>
<thead>
<tr>
<th>Chord number</th>
<th>Disappearance (s)</th>
<th>Reappearance (s)</th>
<th>Chord size (s)</th>
<th>Chord Size (km)</th>
<th>Shift (s)$^a$</th>
<th>$\sigma$</th>
</tr>
</thead>
<tbody>
<tr>
<td>1</td>
<td>20:41:34.800 ± 1.155</td>
<td>20:42:14.995 ± 0.225</td>
<td>40.2 ± 1.2</td>
<td>523 ± 16</td>
<td>+0.7</td>
<td>0.06</td>
</tr>
<tr>
<td>2</td>
<td>20:41:35.54 ± 0.46</td>
<td>20:42:14.47 ± 0.55</td>
<td>38.9 ± 0.7</td>
<td>506 ± 9</td>
<td>+0.6</td>
<td>0.06</td>
</tr>
<tr>
<td>3$^b$</td>
<td>20:41:35.609 ± 0.109</td>
<td>20:42:16.585 ± 0.077</td>
<td>40.98 ± 0.13</td>
<td>533.1 ± 1.7</td>
<td>+0.03</td>
<td>0.06</td>
</tr>
<tr>
<td>5</td>
<td>20:41:44.375 ± 0.525</td>
<td>20:42:24.95 ± 0.40</td>
<td>46.0 ± 0.7</td>
<td>528 ± 9</td>
<td>-0.06</td>
<td>0.11</td>
</tr>
<tr>
<td>6</td>
<td>20:41:37.325 ± 1.100</td>
<td>20:42:24.062 ± 1.287</td>
<td>46.7 ± 1.7</td>
<td>608 ± 22</td>
<td>+1.6</td>
<td>0.08</td>
</tr>
<tr>
<td>7</td>
<td>20:41:35.420 ± 0.425</td>
<td>20:42:17.411 ± 0.382</td>
<td>42.0 ± 0.6</td>
<td>546 ± 8</td>
<td>+0.07</td>
<td>0.05</td>
</tr>
<tr>
<td>8</td>
<td>20:42:02.05 ± 0.73</td>
<td>20:42:37.76 ± 0.76</td>
<td>35.7 ± 1.1</td>
<td>464 ± 14</td>
<td>-0.14</td>
<td>0.23</td>
</tr>
<tr>
<td>9</td>
<td>20:41:37.491 ± 0.590</td>
<td>20:42:14.897 ± 0.160</td>
<td>37.4 ± 0.6</td>
<td>487 ± 8</td>
<td>+1.1</td>
<td>0.06</td>
</tr>
<tr>
<td>10</td>
<td>20:41:57.914 ± 0.110</td>
<td>20:42:31.944 ± 0.136</td>
<td>34.03 ± 0.17</td>
<td>443 ± 2</td>
<td>+0.3</td>
<td>0.02</td>
</tr>
<tr>
<td>11</td>
<td>20:41:54.60 ± 0.15</td>
<td>20:42:28.76 ± 0.12</td>
<td>34.16 ± 0.19</td>
<td>444 ± 2</td>
<td>-0.7</td>
<td>0.02</td>
</tr>
</tbody>
</table>

Notes. Chords’ shifts are explained in Sect. 4.1.

$^a$A positive shift means a shift toward the east, and a negative one indicates a shift toward the west.

$^b$Chord 3’ is the result of merging chords 3 and 4 in Table 3, see Sect. 3.1 for details.

$^c$We have added two seconds to the uncertainty bar given by the square-well fit for chord 12, as explained in Sect. 3.1.

We calculated the rotational phase of 2003 VS$_2$ at the time of the stellar occultation$^5$ which was 0.32 with respect to the absolute brightness maximum (see the vertical black-dashed line in Fig. 4). We did not correct the Julian dates for light travel time since this source of error is negligible due to the closeness in time of all data. The rotational phase obtained implies that the apparent surface area of 2003 VS$_2$ was near its minimum during the occultation event. We note that the rotational phase at the time of the occultation is not influenced by the amplitude of the rotational light curve, but by the rotation period. Given the precision of the rotation period and the time span between the data taken for the occultation event and the rotational light curve, the error calculating the rotational phase is negligible.

4. Results analysis

4.1. Limb fitting

The shape generally considered for a TNO’s limb fitting is an ellipse (see, e.g., [Ortiz et al. 2017; Benedetti-Rossi et al. 2019], even though the object might not project a perfect regular elliptical shape. In this work we do not have enough positive chords to account for topographic features or deviations from a true elliptical shape, so we can only use the simplest and most general model for the limb fitting.

Five adjustable parameters characterize the considered ellipse: the body’s center coordinates relative to the star in the sky plane ($f_x$, $g_y$); the apparent semi-major axis $a'$; the apparent semi-minor axis $b'$; and the tilt angle of the ellipse $PA$, which is the position angle of the semi-minor axis from celestial north and positive to the west. If we assume the Gaia DR2 star position to be correct, the coordinates ($f_x$, $g_y$) give the offsets in RA and Dec, respectively, to be applied to the object’s adopted ephemeris.

We considered two configurations of the positive chords for the limb fitting. For the first configuration, we kept the relative positions of the original chords but had to correct the absolute

$^5$To do this, we considered 2458779.36248264 to be the Julian date of the occultation event, as it was the closest recorded date to the average value of the mid-occultation times from all sites. However, choosing a different Julian date close to the stellar occultation would not change the results given the object’s rotation period.
time of chord 8 (see Sect. 3.1). In this regard, we performed a weighted least-squares polynomial fit of degree one to the centers of the chords in Table 5, not using chord 8 for that fit. We then shifted chord 8 until its center laid on said linear regression and thus performed the elliptical fitting to the chords’ extremities (Fig. 5a). For the second configuration, we aligned the centers of all the chords using the aforementioned linear regression and then performed the elliptical limb fitting to the chords’ extremities (see Fig. 5b).

The alignment of the centers is a condition for the parallel chords of an ellipse, and the needed shifts are not within the centers’ uncertainties. Although synchronization via NTP servers provides theoretical accuracies of 0.01 s, uncertainties of up to tenths of a second have been reported to arise from using different operating systems (Barry et al. 2015), camera software, or even due to delays in the shutter opening. However, adding a nominal average error to all the extremities would be unwise since that error should be the same for the ingress and egress of each chord, and we would be overestimating the error. On the other hand, the necessity of correcting unexplained time shifts has also been reported (Elliot et al. 2010, Braga-Ribas et al. 2013). So by aligning the chords, we account for possible systematic timing errors and small topography that could have decentralized the chords.

We obtained the best elliptical fit to the extremities of the chords via minimization of the sum of squared residuals

\[ \sum (d_i - d_{\text{fit}})^2 \]

with \( d_i \) being the shortest distance between the chord’s extreme and the evaluated ellipse following the direction of the chord. The uncertainties of the elliptical parameters were determined using a Monte Carlo method. To do this, we generated

10 000 random sets of extremities of the chords for each configuration, sampled from within the uncertainty bars, and then searched for the ellipse that minimizes the aforementioned equation. The Monte Carlo distributions obtained for each of the ellipse parameters are included in Appendix B. In Table 6, we show the nominal value and the standard deviation of each distribution.

The instantaneous area-equivalent diameter of 2003 VS₂ is also given in Table 6. Given that the rotational phase of 2003 VS₂ was near its brightness minimum during the stellar occultation (see Fig. 5), the projected area was also near its minimum, and thus this diameter is a lower limit for the real equivalent diameter of 2003 VS₂. If we take the rotational phase and the rotational light curve amplitude into account, we can derive the mean area-equivalent diameter as follows:

\[ D_{\text{mean}} = D_{\text{occ}} \times \frac{10^{0.4(m_{\text{occ}} - m_{\text{mean}})}}{5}, \]

where \( m_{\text{occ}} \) is the relative magnitude of 2003 VS₂ during the stellar occultation (\( m_{\text{occ}} = 0.104 \pm 0.010 \) and \( m_{\text{mean}} = 0 \), see Fig. 4), and \( D_{\text{occ}} \) is its instantaneous area-equivalent diameter.

The derived mean area-equivalent diameter of 545 ±13 km is slightly greater than the value obtained from Herschel radiometric data (\( D = 523^{+35}_{-43} \) km, Mommer et al. 2012), but it is in agreement within the error bars. It is, however, slightly smaller than the area-equivalent diameter of 553±56 km calculated with the data from Benedetti-Rossi et al. (2013) and applying Eq. 1 to take the rotational phase at the occultation time into account, yet it is still compatible within the uncertainty.

### 4.2. Geometric albedo

The geometric albedo at the V band of 2003 VS₂ at the time of the stellar occultation can be derived from the following:

\[ p_V = 10^{0.4m_{\text{Sun}} - H_{\text{occ}}} / (4\pi), \]

with \( m_{\text{Sun}} \) being the V magnitude of the Sun (\( m_{\text{Sun}} = -26.74 \) mag), \( H_{\text{occ}} \) being the instantaneous absolute magnitude of 2003 VS₂ in the V band at the time of the stellar occultation, and \( A \) being the projected area of 2003 VS₂ during the event in astronomical units squared.

To obtain \( H_{V,\text{occ}} \) we corrected 2003 VS₂’s rotationally averaged \( H_{V} (4.14 \pm 0.07 \text{ mag}; \text{Alvarez-Candal et al. } 2016) \) and private communication) by adding the theoretical relative magnitude of 2003 VS₂ during the stellar occultation, which is given by the Fourier fit to the rotational light curve and has a value of 0.104 ± 0.010 mag at the 1σ level.

The obtained geometric albedo of 0.134 ± 0.010 (Table 6) is slightly smaller but compatible with the one derived from the combination of Herschel thermal measurements and Spitzer data (0.147^{+0.062}_{-0.045}; Lellouch et al. 2013). It is also in agreement with the last derived albedo from Benedetti-Rossi et al. (2019) (0.131^{+0.024}_{-0.012}; see Table 6). If we use the same absolute magnitude as Lellouch et al. (2013) (\( H_V = 4.11 \pm 0.38 \) mag) to calculate the albedo, it has a value of 0.14 ± 0.05, which is still smaller but compatible within the error bars.

### 4.3. Size and shape

The double-peaked rotational light curve of 2003 VS₂ suggests that it is either a rotating triaxial ellipsoid or an oblate spheroid with a significant irregularity or a large albedo variation along its surface; this second case is unlikely, especially considering 2003 VS₂’s large light curve amplitude. In addition, the two
<table>
<thead>
<tr>
<th>Parameter</th>
<th>Original chords</th>
<th>Aligned chords</th>
</tr>
</thead>
<tbody>
<tr>
<td>Semi-major axis $a'$ (km)</td>
<td>292 ± 3</td>
<td>293 ± 3</td>
</tr>
<tr>
<td>Semi-minor axis $b'$ (km)</td>
<td>231 ± 6</td>
<td>230 ± 6</td>
</tr>
<tr>
<td>Oblateness $\epsilon' = (a' - b')/a'$</td>
<td>0.21 ± 0.03</td>
<td>0.21 ± 0.03</td>
</tr>
<tr>
<td>Tilt angle (°)</td>
<td>-11 ± 2</td>
<td>-6 ± 2</td>
</tr>
<tr>
<td>$(f_c, g_c)$ (km)$^b$</td>
<td>(-1010 ± 2, 770 ± 5)</td>
<td>(-1007 ± 2, 769 ± 5)</td>
</tr>
<tr>
<td>$(f_c, g_c)$ (mas)$^b$</td>
<td>(-38.57 ± 0.09, 29.40 ± 0.19)</td>
<td>(-38.43 ± 0.09, 29.37 ± 0.17)</td>
</tr>
<tr>
<td>Area-equivalent diameter (km)</td>
<td>519 ± 10</td>
<td>519 ± 10</td>
</tr>
<tr>
<td>Mean area-equivalent diameter (km)</td>
<td>545 ± 13</td>
<td>545 ± 13</td>
</tr>
<tr>
<td>Geometric albedo $p_V$</td>
<td>0.134 ± 0.010</td>
<td>0.134 ± 0.010</td>
</tr>
</tbody>
</table>

Notes. The nominal values were obtained by minimization of a least squares function and the uncertainties are the standard deviation of the Monte Carlo distributions, see Sect. 4.1. The corresponding distributions can be seen in appendix B. The uncertainty bars of the equivalent diameters and albedo were obtained analytically from their respective mathematical expressions.

$^a$ Chord 8 is the only chord that has been shifted for this case, see Sect. 4.1.

$^b$ Offsets with respect to the JPL #33 orbit.

Fig. 5: Elliptical fit to the chords of the stellar occultation for the two considered configurations: (a) original distribution of the chords, with chord 8 (dashed line) already shifted; and (b) final distribution of the chords after aligning their centers using a least squares linear fit. In both plots, the positive chords are shown in solid blue and the negative chords are in dotted blue; uncertainties of the star’s disappearance time are shown in green and those of the reappearance time are in red; and the black dots show the center of the chords. From top to bottom, the chords follow the same order as in Table 3. The limiting negative chord in the north corresponds to observer H. Mikuž and the chord limiting in the south corresponds to observer V. Dumitrescu, see Table 4. The black arrows in the top left of each plot show the direction of the shadow motion. The best elliptical fit to the extremities of the chords is shown in black and all the ellipses from the Monte Carlo distribution are plotted in gray, see Sect. 4.1 for details.

Maxima and minima per rotation cycle are different, which usually indicates a triaxial shape with some albedo spots or small topographic features, since we do not expect a perfectly symmetrical and homogeneous body orbiting in space. This is the case for many observed TNOs with large rotational light curve amplitudes, such as Haumea (Lacerda 2010), Varuna (Fernández-Valenzuela et al. 2019), and 2008 OG$_{19}$ (Fernández-Valenzuela et al. 2016). On the other hand, if not triaxial, the rotation period of 2003 VS$_2$ would be 3.6 h, which is probably too fast for a TNO (no TNO is known to rotate that fast). Hence, assuming a triaxial ellipsoidal shape for 2003 VS$_2$, we searched for the axes $a > b > c$ of a triaxial body, rotating around its $c$ axis, which would give the observed elliptical projection during the stellar occultation (Fig. 5), while also showing an amplitude of $\Delta m = 0.264 \pm 0.017$ mag on its rotational light curve.

To do this, we used the procedures described in Gendzwill & Stauffer (1981) to generate ellipsoids characterized by three semi-major axes and three orientation angles with respect to the Cartesian system, and to then project these ellipsoids into a plane perpendicular to the line of sight in order to compare that projection with the instantaneous limb of 2003 VS$_2$ during the stellar occultation (Fig. 5). An additional constraint to the model is the observed rotational light curve amplitude of 2003 VS$_2$ (Fig. 4). To include it, we implemented the well-known relation between the peak-to-peak variation of the rotational light curve of a small
body and its three principal semiaxes \( (\text{Binzel et al. 1989}) \):

\[
\Delta m = 2.5 \log \left( \frac{a}{b} \right) - 1.25 \log \left( \frac{a^2 \cos^2(\theta) + c^2 \sin^2(\theta)}{b^2 \cos^2(\theta) + c^2 \sin^2(\theta)} \right),
\]

where \( \theta \) is the polar aspect viewing angle, that is, the angle between the rotation axis (c, in this case) and the line of sight.

For each of the ellipsoids obtained via Monte Carlo for the limb fitting (Sect. 4.3, Fig. 5), we searched for the ellipsoid that would give the most similar projection in terms of least squares minimization. Although the limb fit obtained in Sect. 4.3 has virtually the same values for both considered chord configurations, we decided to search for the corresponding ellipsoid of both solutions to see if the slight difference in the tilt angle would give different 3D shapes. The obtained distributions for the ellipsoid's semiaxes, aspect angle, and the corresponding rotational light curve amplitude derived from Eq. (3) are shown in appendix B. In Table 7 we show the best 3D fit, which is the one corresponding to the best limb fit in Table 6; also the uncertainty bars are given as the standard deviation of the Monte Carlo distributions. According to the results, the obtained triaxial ellipsoid would produce a variation of 0.24 mag during its rotation so, as stated before, the remaining observed rotational light curve amplitude might be due to albedo spots or topographic features that cannot be studied with the available data; we would either need more chords from the stellar occultation to study topographic features, or rotational light curves obtained with different filters to study albedo variations. The differences in the rotational light curve amplitude due to albedo spots would be on the order of a few cents of a magnitude; this is the range of variability due to albedo variegations observed in large TNOs such as Varuna, Haumea, 2008 O\(\text{G}_{\text{I}}\). Considering this and given that a model of a triaxial body without albedo spots to explain the light curve amplitude is a simplification, we also computed the ellipsoid accepting solutions that give a rotational light curve amplitude of 0.06 mag less than the observed value, so as to allow for large albedo spots. These results are also included in Table 7 and appendix B. Finally, the spherical volume equivalent diameter \( (D_{\text{V}_{\text{eq}}}) \) derived from these parameters is also presented in Table 7.

Both obtained solutions are almost compatible with the one presented in Benedetti-Rossi et al. (2019) within the errors. The differences might arise from the fact that, in the previous work, the calculations were simplified by considering that 2003 VS\(\text{Z} \) was on its brightness maximum during the 2014 occultation, when the actual rotational phase at the time of the occultation was +0.07 with respect to the brightness maximum. Moreover, the rotational light curve amplitude they used was smaller than the one used in this work. However, we checked that the two amplitudes cannot be explained simultaneously by a change in the aspect angle.

We derived the ecliptic coordinates of the pole \( (\lambda_p, \beta_p) \) by simultaneously minimizing the difference between the observed rotational light curve amplitude and the one derived using Eq. (3) and the difference between the aspect angle \( \delta \) derived in Sect. 4.3 and the one derived via the following:

\[
\delta = \frac{\pi}{2} - \arcsin \left[ \sin \beta_e \sin \beta_p + \cos \beta_e \cos \beta_p \cos (\lambda_e - \lambda_p) \right],
\]

with \( (\lambda_e, \beta_e) \) being the ecliptic coordinates of the sub-Earth point in the object-centered reference frame at the time of this work’s stellar occultation. The pole’s ecliptic coordinates \( (\lambda_p, \beta_p) \) derived are \( (228^\circ, 39^\circ) \) (for a more detailed explanation, see Fernández-Valenzuela 2022 and references therein). Combining the pole’s coordinates, 2003 VS\(\text{Z} \)’s ephemeris, and the 3D axes obtained in Sect. 4.3 we obtained the theoretical variability of the rotational light curve amplitude through the years and compared it to the values in the literature. The results are plotted in Fig. 6. It can be seen that the previously published results were obtained during the minimum of the rotational light curve amplitude, but it has been increasing since 2005. The published value found in Benedetti-Rossi et al. (2019) is much lower than theoretically expected, but the remaining measurements do agree with the model. We suspect that the rotational light curve in Benedetti-Rossi et al. (2019) may have been contaminated by a background star or had some unidentified technical problem as of yet or a reduction artifact because there are no physical scenarios that can explain a sudden decrease in the amplitude, except perhaps a sudden brightening due to dust release through sublimation activity or through a collision, but this is very unlikely.

The projection of the ellipsoids obtained from the original-chords’ configuration at the rotational phase in which the stellar occultation in Benedetti-Rossi et al. (2019) happened \( (+0.07 \text{ with respect to the brightness maximum}) \) gives ellipses with semiaxes of \( a \times b = 305 \pm 4 \times 230 \pm 6 \text{ km} \) (for the observed rotational light curve amplitude) and \( a \times b = 302 \pm 4 \times 230 \pm 6 \text{ km} \) (for a minimum rotational light curve amplitude of 0.18 mag). We note that they are in agreement with the limb fit reported in Benedetti-Rossi et al. (2019).

To conclude, we can compare these results with the theoretical axes ratios of a triaxial ellipsoid with the rotation period of 2003 VS\(\text{Z} \) and a homogeneous density to see if the derived shape is compatible with a hydrostatic-equilibrium figure. Using the formalism from Chandrasekhar (1987), we searched for the theoretical values of the axes ratios \( a/b \) and \( b/c \) (solid and dashed lines in Fig. 7) respectively of a triaxial ellipsoid with a rotation period of 7.41753 ± 0.00001 h and homogeneous densities between 600 and 1000 kg m\(^{-3}\). We also plotted in Fig. 7 the axes ratios’ bands derived from the stellar occultation (their values are presented in Table 7). In the figure, one can see that there is an overlapping region between the theoretical and observed axes ratio \( a/b \), but that is not the case for the axes ratio \( b/c \). Therefore, we can conclude that the derived shape of 2003 VS\(\text{Z} \) is not consistent with the hydrostatic-equilibrium figure of a homogeneous body with the rotation period of 2003 VS\(\text{Z} \) for any density...
value. This result suggests that, similar to the case of Haumea (Ortiz et al. 2017), we might need to consider granular physics to explain the body’s shape, because in this case, a differentiated body is less plausible due to 2003 VS$_2$’s smaller size. In fact, the differentiation obtained by Loveless et al. (2022) for icy bodies with radii larger than 200 km, though not negligible, is so small that it would probably not produce a significant deviation from the equilibrium shape for a homogeneous body of the size of 2003 VS$_2$. However, that model excludes or simplifies some of the physical and chemical processes and considers pure spherical bodies, so we cannot discard some differentiation. The possibility that 2003 VS$_2$ might be sustaining some stress seems more plausible but, in reality, we can only speculate that the deviation of its shape from that of pure hydrostatic equilibrium could be due to one or a combination of both scenarios.

5. Conclusions

On 2019 October 22, we observed the stellar occultation of the GAIA source 3449076721168026496 ($m_V = 14.1$ mag) caused by the plutino (84922) 2003 VS$_2$. Out of the 39 participating observing sites, 12 reported a positive detection, located in Bulgaria, Romania, and Serbia. Two positive chords were combined forming a single one, giving 11 effective positive chords. This is one of the best observed stellar occultations by a TNO so far.

The projected shape of 2003 VS$_2$ was fitted to an ellipse considering two configurations of the positive chords. For the first one, we fit the extremities of the original chords obtained from the stellar occultation after shifting only one chord due to a problem with its absolute time. For the second case, we shifted all the positive chords to align their centers. The best solution for the instantaneous limb of 2003 VS$_2$ was obtained by minimizing the sum of squared residuals, and the uncertainties were derived via a Monte Carlo method and correspond to the standard deviation of the Monte Carlo distributions shown in appendix B for the two chords’ configurations considered for the limb fitting.

Chord 8 is the only chord that has been shifted for this case, see Sect. 4.1.

Theoretical axes ratios $a/b$ (solid blue line) and $b/c$ (dashed red line) of a 3D body in hydrostatic equilibrium, rotating with a period of $7.41753 \pm 0.00001$ h, for different densities (Chandrasekhar 1987). The axes ratios from the observed data are plotted as colored bands for both of the considered cases. The theoretical $b/c$ ratios from hydrostatic equilibrium do not agree with the observations for any possible density of a homogeneous body because there is no intersection of the dashed red line with the pink band.

Table 7: Results of the 3D fit of 2003 VS$_2$.

<table>
<thead>
<tr>
<th>Parameter</th>
<th>Original chords$^a$</th>
<th>Shifted chord</th>
<th>Original chords$^a$</th>
<th>Shifted chords</th>
</tr>
</thead>
<tbody>
<tr>
<td>$a$ (km)</td>
<td>339 ± 5</td>
<td>338 ± 5</td>
<td>327 ± 5</td>
<td>330 ± 5</td>
</tr>
<tr>
<td>$b$ (km)</td>
<td>235 ± 6</td>
<td>233 ± 7</td>
<td>233 ± 6</td>
<td>235 ± 9</td>
</tr>
<tr>
<td>$c$ (km)</td>
<td>226 ± 8</td>
<td>227 ± 7</td>
<td>228 ± 8</td>
<td>222 ± 8</td>
</tr>
<tr>
<td>Aspect angle $\theta$ (°)</td>
<td>59 ± 2</td>
<td>57 ± 3</td>
<td>51 ± 2</td>
<td>51 ± 6</td>
</tr>
<tr>
<td>Rotational light curve amplitude (mag)</td>
<td>0.2497 ± 0.0008</td>
<td>0.2489 ± 0.0009</td>
<td>0.187 ± 0.003</td>
<td>0.184 ± 0.003</td>
</tr>
<tr>
<td>$a/b$</td>
<td>1.44 ± 0.06</td>
<td>1.45 ± 0.06</td>
<td>1.41 ± 0.06</td>
<td>1.40 ± 0.07</td>
</tr>
<tr>
<td>$b/c$</td>
<td>1.040 ± 0.017</td>
<td>1.022 ± 0.018</td>
<td>1.02 ± 0.02</td>
<td>1.06 ± 0.03</td>
</tr>
<tr>
<td>Spherical volume equivalent diameter (km)</td>
<td>524 ± 7</td>
<td>523 ± 7</td>
<td>517 ± 7</td>
<td>516 ± 8</td>
</tr>
</tbody>
</table>

Notes. We present the ellipsoids obtained using the observed rotational light curve amplitude ($\Delta m = 0.264 \pm 0.017$ mag) and the ones allowing for bigger albedo spots, see Sect. 4.3. The nominal values of the parameters are the ones obtained from the optimal ellipses in Table 6. The uncertainties are given as the standard deviation of the Monte Carlo distributions shown in appendix B for the two chords’ configurations considered for the limb fitting.

(1)
tary $\theta_2 = 121^\circ \pm 2^\circ$, depending on the considered sense of rotation. If we allow for a 0.1 mag variability due to albedo spots in the 3D model, the resulting ellipsoid has a semi-major axis $a \sim 10$ km smaller and fully compatible with the projected shape in Benedetti-Rossi et al. (2019). These values give a spherical-volume equivalent diameter of $D_{eq} = 524 \pm 7$ km. This solution is not compatible with a homogeneous body in hydrostatic equilibrium rotating with the known period of 2003 VS$_2$, requiring differentiation or, most likely, an internal structure that can sustain stress to some degree. Finally, we found no evidence of a dense ring or debris material orbiting around 2003 VS$_2$ of the type seen in Chariklo (Braga-Ribas et al. [2014]).

Acknowledgements. We acknowledge financial support from the State Agency for Research of the Slovak MCIU through the “Center of Excellence Severo Ochoa” award to the Instituto de Astrofísica de Andalucía (SEV-2017-0709). Funding from Spanish projects PID2020-112798GB-I00 from AEI and Proyecto de Excelencia of the Junta de Andalucía P2020-01309 is acknowledged. Part of the research leading to these results has received funding from the European Research Council under the European Community’s H2020 (2014-2020) ERC Grant Agreement no. 660416 “LUCKY STAR”. M.V.L. acknowledges funding from Spanish project AYA2017-89637-R (FEDER/MICINN). P.S.S. acknowledges financial support by the Spanish grant AYA-RTI2018-098657-J-I00 “LEOSBNAF”. Part of the work of M.P. was financed by a grant of the Romanian National Authority for Scientific Research and Innovation, CNCS - UEFISCDI, PNII-IP1-TE-2019-1504. E.F.-V. acknowledges financial support from the Florida Space Institute and the Space Research Initiative. The following authors acknowledge the respective CNPq grants: F.B-R 309578/2017-5; B.E.M. 150612/2020-6; RV-M 304544/2017-5; 401903/2016-8; J.I.B.C. 308150/2016-3 and 310683/2019-6; M.A. 427700/2018-3; 310683/2017-3 and 473002/2013-2. D.I. and O.V. acknowledge funding provided by the Ministry of Education, Science, and Technological Development of the Republic of Serbia (contracts 451-03-9/2021-14/200104, 451-03-9/2021-14/200002). D.I. acknowledges the support of the Alexander von Humboldt Foundation. M.H. thanks the Slovak Academy of Sciences (VEGA No. 200152) for financial support of the Alexander von Humboldt Foundation. M.H. thanks the Slovak Academy of Sciences (VEGA No. 200152) for financial support of the Alexander von Humboldt Foundation. M.H. thanks the Slovak Academy of Sciences (VEGA No. 200152) for financial support of the Alexander von Humboldt Foundation.

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Appendix A: Observing sites involved in the occultation campaign that could not observe

In this appendix we include the list of contacted observing sites that could not observe due to bad weather or technical problems.

Table A.1: Details of the observing sites involved in the 2003 VS$_2$ 2019 October 22 occultation that could not observe.

<table>
<thead>
<tr>
<th>Site name Location</th>
<th>Latitude (N) Longitude (E)</th>
<th>Telescope aperture (cm)</th>
<th>Observers</th>
<th>Problem</th>
</tr>
</thead>
<tbody>
<tr>
<td>Galați Observatory (x2) Romania C73</td>
<td>45° 25' 7''</td>
<td>40</td>
<td>O. Tercu</td>
<td>Bad Weather</td>
</tr>
<tr>
<td>Romania</td>
<td>28° 1' 57''</td>
<td>20</td>
<td>A. M. Stoian, G. Neagu, D. Zlat</td>
<td></td>
</tr>
<tr>
<td>Astroclubul București (x2) Romania (mobile telescope) –</td>
<td>–</td>
<td>40</td>
<td>D. Berteșteanu, M. H. Naiman, Z. Deak</td>
<td></td>
</tr>
<tr>
<td>Romania (mobile telescope) –</td>
<td>–</td>
<td>20</td>
<td>B. Dumitru</td>
<td>Technical problems</td>
</tr>
<tr>
<td>Bârlad Observatory Romania L22</td>
<td>46° 13' 54''</td>
<td>20</td>
<td>C. Vantdevara</td>
<td>Bad weather</td>
</tr>
<tr>
<td>G.V.Schiaparelli Observatory Varese, Italy 204</td>
<td>45° 52' 04''</td>
<td>84</td>
<td>L. Buzzi</td>
<td>Bad weather</td>
</tr>
<tr>
<td>San Marcello Pistoiese Italy 104</td>
<td>44° 3' 46''</td>
<td>60</td>
<td>P. Bacci</td>
<td>Technical problems</td>
</tr>
<tr>
<td>ISON Uzhgorod Observatory Derenivka, Ukraine K99</td>
<td>48° 33' 48''</td>
<td>40</td>
<td>V. Kudak</td>
<td>Technical problems</td>
</tr>
<tr>
<td>Odessa-Mayakaki Ukraine 583</td>
<td>46° 23' 50''</td>
<td>80</td>
<td>Y. Krugly</td>
<td>I. Belskaya, Bad weather</td>
</tr>
<tr>
<td>Zalischy Ukraina L18</td>
<td>48° 50' 53''</td>
<td>30</td>
<td>Y. Krugly</td>
<td>I. Belskaya, Bad weather</td>
</tr>
<tr>
<td>Uzhgorod-Derenivka, Ukraine K99</td>
<td>48° 33' 48''</td>
<td>40</td>
<td>Y. Krugly</td>
<td>I. Belskaya, Bad weather</td>
</tr>
<tr>
<td>Kyiv-Comet Station Ukraine 585</td>
<td>50° 17' 52''</td>
<td>70</td>
<td>Y. Krugly</td>
<td>I. Belskaya, Bad weather</td>
</tr>
<tr>
<td>Kharkiv-Chuguev Station Ukraine 121</td>
<td>49° 38' 34''</td>
<td>70</td>
<td>Y. Krugly</td>
<td>I. Belskaya, Bad weather</td>
</tr>
<tr>
<td>Hvar Observatory Hvar, Croatia</td>
<td>43° 10' 42''</td>
<td>106</td>
<td>D. Ruzdjak</td>
<td>S. Cikota, Bad weather</td>
</tr>
</tbody>
</table>

Appendix B: Monte Carlo distributions

In this appendix we present the obtained Monte Carlo distributions for the limb fitting and the 3D fitting of 2003 VS$_2$ discussed in Sects. 4.1 and 4.3 respectively, for the two considered distributions of the positive chords.
Fig. B.1: Monte Carlo distributions for the semiaxes $a$ and $b$ (in km), the tilt angle (in degrees), and the coordinates $(x, y)$ of the center of the ellipse, from the elliptical fit to the unshifted chords (with only chord 8 shifted) in Fig. 5a, see Sect. 4.1. The vertical red lines show the value of the parameters for the best elliptical fit via minimization of the sum of squared residuals, see Table 6.

Fig. B.2: Monte Carlo distributions for the semiaxes $a$ and $b$ (in km), the tilt angle (in degrees), and the coordinates $(x, y)$ of the center of the ellipse, from the elliptical fit to the aligned chords in Fig. 5b, see Sect. 4.1. The vertical red lines show the value of the parameters for the best elliptical fit via minimization of the sum of squared residuals, see Table 6.

Fig. B.3: Monte Carlo distributions for the semiaxes $a$, $b$, and $c$ (in km) of a triaxial ellipsoid compatible with the occultation observation and observed rotational light curve amplitude, as well as distributions of the aspect angle and derived rotational light curve amplitude obtained via Eq. 3, for the case of the original chords with only chord 8 shifted; see Sects. 4.1 and 4.3. The vertical red lines show the ellipsoidal values corresponding to the best elliptical fit, see Table 7.
Fig. B.4: Monte Carlo distributions for the semiaxes $a$, $b$, and $c$ (in km) of a triaxial ellipsoid compatible with the occultation observation and observed rotational light curve of 2003 VS$_2$, as well as distributions of the aspect angle and derived rotational light curve amplitude obtained via Eq. 3, for the case of all the chords shifted; see Sects. 4.1 and 4.3. The vertical red lines show the ellipsoidal values corresponding to the best elliptical fit, see Table 7.

Fig. B.5: Monte Carlo distributions of the axes ratios $a/b$ and $b/c$ of the fitted ellipsoid for the considered case of (a) the chords unshifted with only chord 8 shifted and (b) the aligned chords, for the observed rotational light curve amplitude. The vertical red lines show the axes ratios derived from the best elliptical fit, see Table 7.

Fig. B.6: Monte Carlo distributions for the semiaxes $a$, $b$, and $c$ (in km) of a triaxial ellipsoid compatible with the occultation observation and allowing a minimum rotational light curve amplitude of 0.18 mag, as well as distributions of the aspect angle and derived rotational light curve amplitude obtained via Eq. 3, for the case of the original chords with only chord 8 shifted; see Sects. 4.1 and 4.3. The vertical red lines show the ellipsoidal values corresponding to the best elliptical fit, see Table 7.
Fig. B.7: Monte Carlo distributions for the semiaxes a, b, and c (in km) of a triaxial ellipsoid compatible with the occultation observation and allowing a minimum rotational light curve amplitude of 0.18 mag, as well as distributions of the aspect angle and derived rotational light curve amplitude obtained via Eq. 3 for the case of all the chords shifted; see Sects. 4.1 and 4.3. The vertical red lines show the ellipsoidal values corresponding to the best elliptical fit, see Table 7.

Fig. B.8: Monte Carlo distributions of the axes ratios $a/b$ and $b/c$ of the fitted ellipsoid for the considered case of (a) the chords unshifted with only chord 8 shifted and (b) the aligned chords, if we allow a minimum rotational light curve amplitude of 0.18 mag. The vertical red lines show the axes ratios derived from the best elliptical fit, see Table 7.