VVVX near-IR photometry for 99 low-mass stars in the Gaia EDR3 Catalog of Nearby Stars

A. Mejías1, D. Minniti1,2, J. Alonso-García3,4, J. C. Beamín5, R. K. Saito6, and E. Solano7,8

1 Departamento de Ciencias Físicas, Universidad Andrés Bello, Av. Fernández Concha 700, Santiago, Chile
e-mail: andreac.mejiasa@gmail.com
2 Vatican Observatory, V00120 Vatican City State, Italy
3 Centro de Astronomía (CITEVA), Universidad de Antofagasta, Av. Angamos 601, Antofagasta, Chile
4 Millennium Institute of Astrophysics, Nuncio Monseñor Sotero Sanz 100, Of. 104, Providencia, Santiago, Chile
5 Núcleo de Astroquímica y Astrofísica, Instituto de Ciencias Químicas Aplicadas, Facultad de Ingeniería, Universidad Autónoma
de Chile, Av. Pedro de Valdivia 425, Santiago 7300912, Chile
6 Departamento de Física, Universidad Federal de Santa Catarina, Trindade 88040-900, Florianópolis, SC, Brazil
7 Departamento de Astrofísica, Centro de Astrobiología (CSIC-INTA), ESAC Campus, Camino Bajo del Castillo s/n, E-28692 Villanueva de la Cañada, Madrid, Spain
8 Spanish Virtual Observatory, E-28692 Villanueva de Cañada, Madrid, Spain

Received 12 July, 2021 ; accepted 13 December, 2021

ABSTRACT

Context. Red dwarf stars, which represent 75% of stars in the Milky Way, can be studied in great detail in the solar neighborhood, where the sample is more complete.

Aims. We intend to better characterize red dwarf candidates selected from the Gaia Catalog of Nearby Stars (GCNS) using optical and near-infrared multi-filter photometry from the Vista Variables in the Vía Láctea eXtended (VVVX) Survey, the DECam Plane Survey (DECaPS), the Panoramic Survey Telescope and Rapid Response System (Pan-STARRS), and the Wide-field Infrared Survey Explorer (WISE).

Methods. We performed a cross-matching procedure among the positions of a color-selected sample of M dwarfs in the VVVX Survey and the Gaia Early Data Release 3 sub-catalog of nearby stars. We explored their stellar parameters and spectral types using the Virtual Observatory SED Analyzer (VOSA). Radii were also obtained from the computed luminosities and T eff using the Stefan-Boltzmann equation. Masses and ages were computed for some of the objects using evolutionary tracks and isochrones. Additional mass estimations were obtained with the M - M relation. We then validated our results for the stellar parameters of two of our objects with spectra obtained with the TripleSpec instrument at the SOAR telescope, as well as those of our total amount of stars through a direct comparison with an independent sample from the literature. We revised the objects in our sample and compared their proper motion vectors with other sources within 30" to identify possible companions and probed their renormalized unit weight error (RUWE) values to identify unresolved companions.

Results. We present a catalog of physical parameters for 99 low-mass objects with distances from 43.2 to 111.3 pc. Effective temperatures range from 2500 to 3400 K, with the majority of stars in the sample compatible with the status of M4 dwarfs. We obtained a good agreement between the stellar parameters computed with VOSA and the estimations from observed spectra, also when comparing with an independent sample from the literature. The distribution of masses obtained with VOSA is concentrated toward the very low-mass regime. Eight objects present values of RUWE ≥ 1.4 and seven are consistent with being part of a binary system.

Key words. Stellar Astrophysics – Low-mass stars – Galactic plane – Virtual Observatory Tools

1. Introduction

Low-mass stars, also known as red dwarfs or M dwarfs, are the lowest mass stars that are capable of fusing hydrogen in their cores (0.075 < M< 0.6), with spectral types between M0 to M9. They comprise around 75% of the objects in the Milky Way, and ~ 50% of these stars may harbor terrestrial exoplanets, with one in seven M dwarfs hosting Earth-sized planets within the habitable zone (Kopparapu 2013; Dressing & Charbonneau 2015). Their low masses and faint luminosities maintain the habitable zone, defined as the region where water can exist in a liquid state on a nearby exoplanet’s surface (Dole 1964; Hart 1979; Kasting et al. 1993). In addition, when compared to Sun-like stars or earlier types, the smaller masses and radii of red dwarfs favor the detection and characterization of smaller and less massive exoplanets via transit and radial velocity methods, since the light curves and radial velocity signals are more significant for Earth-sized planets around red dwarfs compared with the signals produced by the same planet around a larger star (Feliz et al. 2021; Muirhead et al. 2011).

Therefore, planets that are orbiting within the habitable zone are more easily discovered and characterized, making M dwarfs primary targets for exoplanet search missions. It is then particularly interesting to identify and characterize M dwarfs that are close to the Sun, since their stellar parameters will be better constrained, allowing us to better determine the properties of an exoplanet orbiting them.

Article number, page 1 of 10
Over the last twenty years, several efforts have been made to study samples of red dwarfs in the vicinity of the Sun. In the work of Reid et al. (2004), the New Luyten Two Tenths (NLTT; Luyten 1980) proper motion catalog was cross-matched with the Two Micron All-Sky Survey (2MASS) Second Incremental release to produce a catalog of 1913 stars, of which 1910 have astrometric, photometric, or spectroscopic distances (or a combination of the three), and 815 of those stars are likely to be within 20 pc of the Sun. This resulting catalog from the NLTT follow-up observations is composed of M dwarfs, spanning spectral types between M2 and M7. Large-scale surveys have also allowed for the identification of very low-mass stars and substellar objects as brown dwarfs. In 2008, Phan-Bao et al. (2008) reported nearby M and L dwarfs within 30 pc at low Galactic latitudes (|b| < 15°) over 4800 square degrees in the Deep Near-Infrared Southern Sky Survey (DENIS) database. Another example is given by the work of Thompson et al. (2013), where the authors presented 42 low-mass stars and brown dwarf candidates, including 15 M dwarfs with Wide-Field Infrared Survey Explorer (WISE) satellite photometry.

Other catalogs such as the All-Sky Catalog of Bright M dwarfs from Lépine & Gaidos (2011) provide information on dates, including 15 M dwarfs with Wide-Field Infrared Survey Explorer (WISE) satellite photometry. In this work, we additionally provide deep optical photometry in the $grizY$ bands from the DECaPs survey (DeCaPs; Schlafly et al. 2018) and the Panoramic Survey Telescope And Rapid Response System Survey (Pan-STARRS; Chambers et al. 2016). Due to the importance of precise characterizations of the sources to accurately determine the properties of any possibly orbiting exoplanets, we computed effective temperatures and spectral types and probed their astrometric parameters to identify objects that could be in binary systems.

In this work, we additionally provide deep optical photometry in the $grizY$ bands from the DECaPs survey (DeCaPs; Schlafly et al. 2018) and the Panoramic Survey Telescope And Rapid Response System Survey (Pan-STARRS; Chambers et al. 2016). Due to the importance of precise characterizations of the sources to accurately determine the properties of any possibly orbiting exoplanets, we computed effective temperatures and spectral types and probed their astrometric parameters to identify objects that could be in binary systems.

The paper is organized as follows. In Section 2, we describe the different surveys where we obtained the multi-wavelength photometry to select low-mass stars in the Galactic plane. In Section 3, we provide a summary of the selection method for low-mass stars in the VVVX survey. We then present the methods that we use to determine their stellar parameters and spectral types from photometry only, along with a further characterization from the spectroscopic data for two of our objects in Section 4. The results of our computations and their validations can be found in Section 5, followed by a discussion in Section 6, and the main conclusions of our work are presented in Section 7.

2. Data

In this work, we explored a region of the southern Galactic disk with the VVVX survey (Minniti 2018), an extension of the VVV survey that maps the midplane and bulge of the Milky Way in three near-infrared photometric bands ($JHK_s$). The VVVX survey fills the gap between the VVV and VISTA Hemisphere Survey (VHS) and expands the coverage of the sky from 562 deg$^2$ in VVV to ~ 1700 deg$^2$. It also extends the VVV time-baseline enabling proper motion measurements of $\lesssim 0.3$ mas yr$^{-1}$ in the
optically obscured regions where Gaia is limited by extinction (Smith et al. 2018). Our region of study corresponds to an area of 120° x 9°, consisting of two stripes of 83 x 2 tiles each, where one tile corresponds to a field of view of 1.501 deg².

A VVVX photometric catalog is being produced (Alonso-García et al., in prep.) following the point spread function (PSF) fitting techniques used to create the VVV PSF photometric catalog, as described in Alonso-García et al. (2018). We used a preliminary version of this new catalog to identify the sources and their near-infrared photometry in our region of interest.

Astrometric information was obtained from the Second Data Release (DR2) of the Gaia mission and was updated with the new Early Third Data Release (EDR3; see Table 1).

Optical photometry data were obtained from the DECaPS Survey (Schlafly et al. 2018) for those regions between 240° > l > 5°, and from Pan-STARRS1 (Chambers et al. 2016) for 240° > l > 230°. The DECaPS Survey is a five-band optical and near-infrared survey of the southern Galactic plane with the Dark Energy Camera (DECam) at Cerro Tololo. The total footprint of this survey covers ~ 1000 deg². On the other hand, the Pan-STARRS1 (PS1; Chambers et al. 2016) is the first part of Pan-STARRS that has already been completed and forms the basis for both Data Releases 1 and 2 (DR1 and DR2). The PS1 survey used the 1.8 Pan-STARRS Telescope #1 in Hawaii and its 1.4 Gigapixel camera to image the sky in five broadband filters grizY.

Finally, we added the WISE magnitudes reported in the Gaia Catalog of Nearby Stars when available to obtain infrared photometry at longer wavelengths than the ones of the VVVX survey. WISE mapped the entire sky at the beginning of 2010 and continued through early 2011 in four mid-infrared bandpasses centered at wavelengths of 3.4, 4.6, 12, and 22 µm namely W1 to W4, respectively (Wright et al. 2010). WISE’s two shortest wavebands, W1 and W2, are the most valuable in searching for cold stars and brown dwarfs because these objects emit a substantial amount of their flux at these wavelengths, which are the ones used in this work. These values and the near-infrared photometry data of our objects can be found in Tables 2 and 3.

3. Selection of low-mass objects

The 99 objects included in this work are a sub-sample of red dwarfs candidates that were color-selected in the VVVX survey disk region. The procedure for the selection of these stars is briefly described below, and a full description of our selection method will be presented in Mejías et al. (in prep).

First, we performed a cross-match between the positions reported for the VVVX sources with DECaPS or Pan-STARRS depending on the region of the VVVX footprint, and then with the Gaia DR2 data to obtain astrometric information as parallaxes and proper motions. The specific steps done for the selection of red dwarf candidates are as follows: 1) We applied a quality criteria to the photometry, selecting only sources with σr < 0.1. 2) From this high-quality data, we selected all sources with a parallax value larger than 1.0 mas. 3) To select our low-mass stars candidates, we applied near-IR and optical color cuts, and 4) we used the reduced proper motion value to discriminate possible red giant contaminants.

To determine the color cuts, we explored a spectrophotically confirmed sample of M dwarfs from the work of West et al. (2011). We determined the mean (J − H), (H − Ks) and (J − Ks) colors and defined our color cuts to be within 1-σ from the mean values (Eqs. 1-3):

\[ 0.414 < J - H < 0.695, \]
\[ 0.058 < H - Ks < 0.504, \]
\[ 0.621 < J - Ks < 1.051. \]

On the other hand, we based our optical color cuts in those reported in the work of Kowalski et al. (2009) (Eqs. 4-6):

\[ i - z \geq 0.3, \]
\[ r - i \geq 0.6, \]
\[ r - i \geq 1.35 + [(i - z) - 0.319] + 0.6. \]

In a color-selected sample of M dwarfs, red giant stars are the main contaminants. We carried out a reduced proper motion selection to identify possible giant stars that could be present in the sample. We calculated the J magnitude reduced proper motion for the objects in the sample with the following definition:

\[ H_J = J + 5 \log_{10} \mu \]

We used the same criteria adopted in Rojas-Ayala et al. (2014), where dwarf stars are all the objects with a reduced proper motion greater than the result given by the Eq. 8:

\[ H_J^{\text{dwarf}} > H_J^* = 68.5/(J - Ks) - 50.7. \]

The total sample contains N ~ 3.7 × 10⁵ objects, updated with the outcome of the EDR3. Distances were also updated to the EDR3 distances computed from a Bayesian approach by Bailer-Jones et al. (2021).

The recent Early Data Release of the Gaia Mission includes a smaller catalog of 331,312 objects that are within 100 pc of the Sun (Gaia Collaboration et al. 2021b). We performed a positional cross-matching with a 1” search radius between our sample and the Gaia Catalog of Nearby Stars with and found 99 objects, shown in Fig. 1.

For the 99 low-mass objects of this work, distance values typically do not change between DR2 and EDR3 when they are closer than ~ 80 pc. For larger distances, EDR3 values indicate that objects are closer by 1.4% and more precise (31% smaller errors). On the other side, proper motions are 0.2% higher with 55% smaller errors, corresponding to an important improvement of the astrometric parameters for the objects in our sample.

4. Stellar parameters

4.1. Effective temperatures, log(g), radii, masses, and ages

Determinations of astrophysical parameters is fundamental to understanding the properties of stars. We used the Virtual Observatory SED Analyzer (VOSA: Bayo et al. 2008) to estimate effective temperature and surface gravity, assuming [Fe/H] = 0.0 for our objects, which is a consistent value for stars in the solar neighborhood (Reid et al. 2001; Haywood 2001; Casagrande et al. 2011), together with estimations of mass and radius. These estimations are done by fitting theoretical models from a library that is available in the VOSA database.

The procedure is summarized as follows: we first provided the photometry of our objects in the optical bandpasses of the DECaPS and Panstarrs Surveys (grizY), along with the near-infrared bands of the VVVX Survey (JHKS). We additionally added WISE photometry when available.
Then we used the VOSA to construct the spectral energy distribution of each object, which allows the user to select the collection of theoretical models to be used in the fitting. The fit is made through a chi-square test between the model and the SED of the object (see Fig. 2).

In this work, we also made use of the BT-Settl CIFIST grid of models (Allard et al. 2011), where a solar metallicity value is adopted, together with the following range in effective temperature and surface gravity:

\[ 1200 \, \text{K} \leq T_{\text{eff}} \leq 4500 \, \text{K}, \]
\[ 4.5 \leq \log g \leq 5.5. \]

The radii of our objects are calculated using the Stefan-Boltzmann equation, making use of the luminosities and effective temperatures previously computed by VOSA. We also estimated the masses and ages for our objects with the isochrones and evolutionary tracks available in the VOSA. For each object, VOSA uses the theoretical isochrones and evolutionary tracks appropriate for the model that best fits the observed photometry. In our case, since we used the BT-Settl CIFIST models to fit the spectral energy distributions of the objects of this work, VOSA computes the ages and masses with the Baraffe et al. (2015) isochrones and evolutionary tracks. To do so, the values of effective temperature and luminosities obtained from the previous chi-square fit are used again, but this time as starting points for interpolating the theoretical isochrones and evolutionary tracks.

### 4.2. Spectral types

To compute the spectral types of these objects, we fitted empirical stellar spectra from the spectral library of Kesseli et al. (2017). Again, we used VOSA to do a chi-square test to obtain the best-fit models and estimate spectral types for the objects. This library is created using spectra from the Sloan Digital Sky Survey’s Baryon Oscillation Spectroscopic Survey (BOSS) and includes spectral templates covering spectral types from O5 through L3, they are binned by metallicity from \(-2.0\) dex through \(+1.0\) dex, and are separated into main-sequence dwarf stars and giants. Metallicity bins are extended down through the spectral sub-type M8, and the wavelength coverage is from 3650 to 10200 angstroms at a resolution of \(R \sim 2000\) (Kesseli et al. 2017).
4.3. TSPEC/SOAR Spectra

We obtained near-infrared (NIR) low-resolution spectra for two of the objects with the TripleSpec instrument mounted on the SOAR 4.1m telescope located at Cerro Pachón, in Chile. We observed in the usual ABBA pattern and obtained telluric standard AOV stars right after the observations of the source at a similar airmass. For the object EDR3 5931456492784696064, we obtained four 90 second exposures, and the telluric star was HIP 81692; for object EDR3 6062219761378864128, we obtained eight 120 second exposures, and the observed telluric was HIP 64582.

For the data reduction for TripleSpec spectra, we used the IDL spextools package (Cushing et al. 2004). This includes wavelength calibration, flat fielding, sky subtraction, and extraction of the individual A and B positions. We then combined all the exposures for each object using the xcombspec task. We adopted a robust weighted mean to combine the spectra. The final step was to correct the instrumental response function and telluric correction and merge the different orders. In order to do this, we used the xtellcor software (Vacca et al. 2003) and xmergeorders task.

In Figure 3, we show the spectra compared with spectral templates of ±1 sub-spectral type in 0.5 spectral type steps. These templates were collected from the public SpeX Prism Library1 and come from Burgasser et al. (2004) and Kirkpatrick et al. (2010). The spectral type of the best-fit spectrum was assigned to our objects and then we derived the effective temperature from the stellar color and effective temperature sequence from Pecaut & Mamajek (2013).

5. Results

The stellar parameters were computed using VOSA (see Section 4). The characterized sample is composed of a majority of M4 stars ($\bar{\epsilon} = 4.3$ and $\sigma = 1.5$), with a mean effective temperature of $\sim 3000$ K and a standard deviation of $\sim 178.61$ K. Since the BT-Settl theoretical model is fixed at solar metallicity, our sample is assumed to have $[Fe/H] = 0.0$ dex. The distributions of effective temperatures and spectral types are shown in Figure 4.

We obtained a mean value for the mass of $0.11 \, M_\odot$ with a minimum of 0.075 $M_\odot$ and a maximum of 0.27 $M_\odot$ for 61 objects of our sample. With respect to age, we report only one object (EDR3 5352517097624883584) with a measured age of $4.87^{+6.67}_{-1.82}$ and $2.84$ Gyr.

We also used the $M_Ks - M_\ast$ relation from the work of Mann et al. (2019), and obtained masses for 95 objects, since the remaining four did not lie between the limits of the relation. The results are similar to those computed with VOSA ($x_{\text{Mass}} = 0.16\, M_\odot$; $\text{Mass}_{\text{min}} = 0.085\, M_\odot$; $\text{Mass}_{\text{max}} = 0.31\, M_\odot$), but the distribution is different, given that VOSA results are much more concentrated towards lower masses (see Fig. 6).

5.1. Comparison between stellar parameters obtained with VOSA and TSPEC spectra

From Fig. 3, we selected the best-fit spectral template for our objects and assigned spectral types to then estimate the effective temperatures with the stellar color and effective temperature sequence from Pecaut & Mamajek (2013). We determined the best fit for both objects by comparing the spectral template with the observed spectra by eye, focusing on the $K_\ast$ band.

For the first star plotted in Fig. 3, VOSA assigned a spectral type of M5.0V with an effective temperature of $3000 \pm 50$ K. From the spectral templates, we observed that the best fit is obtained with the M4.5V (shown in red). This indicates an effective temperature around $3060$ K, which is in good agreement with the temperature estimated from VOSA.

For our other object, we obtained a spectral type of M7.0V with VOSA, but we assigned a spectral type of M6.5V following the spectral template that best fits the spectrum (2MASS J04360273+1547536 in red). Again, from the Pecaut & Mamajek (2013) sequence, we can estimate an effective temperature for this object of $2740$ K, which is in good agreement with the temperature estimated with VOSA which is $2800 \pm 50$ K.

5.2. Comparison with an independent sample from literature

To test our results, we sought a complete independent sample of stars with similar spectral types and similar photometry that also have computed stellar parameters. We found 11 stars to form the basis of a comparison, drawn from the works of Houdebine

---

1 http://pono.ucsd.edu/~adam/browndwarfs/spexprism/
et al. (2019) and Kuznetsov et al. (2019), of which 6 have stellar parameters computed in both works.

We applied our procedure (described in Sects. 4.1 and 4.2) to again compute the stellar parameters for these stars. Our estimated values and combined errors are in good agreement, with all the latter within 3-sigma except for three stars. Those have the highest values of error reported in the literature and, in particular, the star that is furthest from 3-sigma has stellar parameters calculated in both papers, when used as a comparison, only shows this high value of combined error when considering the error reported in Kuznetsov et al. (2019), which is higher than the one in Houdebine et al. (2019) (89 K and 38 K, respectively). Furthermore, we note that there is no clear trend that our method over- or underestimates the values for effective temperatures, since objects can be found both above and below the one-to-one relation marked as a gray line in Fig. 5.

6. Discussion

Color-magnitude diagrams (CMDs) with different combinations of photometric bands for the objects in this sample and the entire GCNS are presented in Fig. 8. In the left panel, we present the CMD for the $G$, $B_P$ and $R_P$ bands of Gaia. The white dwarf sequence is located at the bottom left part, at a bluer ($B_P - R_P$) color and fainter magnitudes than the rest of the GCNS sample ($0.0 \leq (B_P - R_P) \leq 2.0$ ; $G > 12.5$). The core H burning main sequence stars lie in the darkest region of the CMD, and our stars are concentrated in this region.

Some objects are shown at the faint end of the CMD with bluer ($B_P - R_P$) color ($2.0 \leq (B_P - R_P) \leq 4.0$ ; $G > 19.0$). It is important to note that, as mentioned in Smart et al. (2019), the Gaia
**Fig. 7.** Proper motion diagram for the 99 sources detected in the VVVX footprint (red stars) with respect to the proper motions in right ascension and declination coordinates of nearby objects within 30″ (blue dots).

$Bp$ band is intrinsically imprecise due to the very low flux of these objects at this wavelength range. We were not able to compute masses for these objects with the $M_{Ks} - M_*$ relation since it is only valid for red dwarfs between $4.0 < M_{Ks} < 11.0$. Only one has a mass estimated with VOSA which is $0.079M_\odot$, with spectral type L0. On the other hand, at redder colors above the main sequence, there is the object EDR3 6062219761378864128, for which we did not identify any signs of binarity when probing its astrometric information and its RUWE. We observed this object with TSPEC in the near-infrared and obtained a spectral type of M7.0 with 2800 K.

The central panel of Fig. 8 shows the near-infrared CMD. The white dwarf sequence is not visible as these objects are too faint in these bands. The main sequence is shown around $J-K_s < 5.0$ and $G-K_s < 7.5$ and our objects are located at the faintest of this region. The white dwarf sequence is found along with a brown dwarf candidate companion previously known from Reylé (2018). We flagged all these objects as BIN in Table 4, where we present our complete results for the stellar parameters of the 99 objects of the sample. For five of our BIN candidates, the co-moving companion is not included in our sample of low-mass stars, but is present in the GCNS – except one that is neither found in our sample nor in the GCNS. The estimated distance for this system is a little bit larger than 100 pc (103.6 pc), and since the objects included in the Gaia Catalog of Nearby Stars were only required to have a non-zero probability to be within 100 pc, it is possible that only one of the components of the system was included in the catalog. We present their positions and Gaia photometry in Table 5.

For stars with unresolved companions, a single-source astrometric model performs poorly. Such stars can be easily detected with the reduced $\chi^2$ statistics, or the renormalized unit weight error (RUWE; Belokurov et al. 2020). We defined that objects that are further that a single-source fitting from the astrometric solution are those with a RUWE $\geq 1.4$. We marked them in Fig. 9 as unfilled black triangles, and used the flag $HR$ in Table 4.

In our sample, six sources have a high RUWE value, implying the possibility of unresolved companions. In addition, we identified two objects that were previously classified as binary systems and present at the same time, a high RUWE value, pointing again to unresolved companions. We propose that these systems could be triple systems, with one component that is resolved and two more that could not be identified as individual objects. We marked them as (Multiple system) in Table 4.

We searched for our objects in the SIMBAD Astronomical Database (Wenger et al. 2000) and found 3 objects. The first one is EDR3 5254655030851879040, which is included in the work of Gagné et al. (2018) as an M5 that forms part of the Volans-Carina stellar association, which is located at a distance of ~ 75-100 pc with an estimated age of ~ 90 Myr. The other two objects are EDR3 5257810972787599360 and EDR3 5409811755160786048, and both are classified as brown dwarf candidates in the ultra-cool dwarf catalog from Gaia DR2 data from Reylé (2018).

In addition, we performed a search around 1″ of the coordinates of the objects in the Vizier database (Ochsenbein et al. 2000). We highlight 14 of our objects that were present in catalogs of wide binaries (Tian et al. 2020; Mugrauer & Michel 2020; Hartman & Lépine 2020), ultra-cool dwarfs candidates in Gaia (Reylé 2018), and large-amplitude variables in Gaia (Mowlavi et al. 2021), together with catalogs with ultraviolet photometry (Yershov 2014) and lensed Gaia systems (Delchambre et al. 2019). This references are indicated in Table 4.

**7. Conclusions**

We identified a sample of 99 nearby red dwarf stars within 100 pc from the Sun from the Gaia EDR3 that are in the VVVX footprint along with the external VVVX fields of the Galactic disk. To properly characterize their physical properties and spectral types, we first collected (in addition to their near-infrared $JHK_s$ magnitudes) the optical photometry in the $grizY$ bands from the DECaPS and Pan-STARRS surveys and mid-IR data from the WISE Survey when available (see Table 1 for astrometric information and Tables 2 and 3 for photometry).

The spectral energy distribution of these objects was constructed from photometry by the Virtual Observatory SED Analyzer (VOSA) and then compared with the BT-Settl CIFIST (Allard et al. 2011) theoretical models to estimate their stellar parameters and with the empirical spectral library from Kesseli et al. (2017) to obtain their spectral types. We obtained low-resolution spectra with the TSPEC instrument mounted on the SOAR telescope for two objects of the sample and we estimated...
stellar parameters by comparing these spectra with other spectral templates within ± 1 sub-spectral type. Spectral types classified using spectral templates are, for both cases, 0.5 sub-spectral types earlier. Temperatures from the Pecaut & Mamajek (2013) sequence are in agreement with the errors.

We additionally tested our procedure with a complete independent sample obtained from the works of Houdebine et al. (2019) and Kuznetsov et al. (2019) and compared the stellar parameters reported by them with ours for 11 objects with similar spectral types and photometry. Values for effective temperatures are in agreement, with a nearly one-to-one relation. Combined errors are also calculated and all are within 3-sigma, except for three stars that present the highest errors in the literature.

We also used the Baraffe et al. (2015) isochrones and evolutionary models to determine masses and ages. We only report the age of one object to be 4.87+0.66−0.67 Gyr, along with three additional upper limits for the age of 3.0, 1.82, and 2.84. We also computed masses with the $M_{K_s} - M_*$ relation from Mann et al. (2019) and found that the distribution of masses with VOSA are more concentrated to the very low-mass regime, while results from the $M_{K_s} - M_*$ relation are more spread out in terms of mass.

We revised the 99 sources in the images of DECaPS in Aladin, together with their astrometric information from Gaia EDR3 with the aim to identify possible stellar or sub-stellar companions. Seven objects were classified as binaries based on their consistency with regard to the parallaxes and proper motions with nearby objects. Two of these binary systems also have RUWE values that are far from unity (RUWE > 1.4), indicating an additional component that could be unresolved. The final stellar parameters are presented in Table 4.

Acknowledgements. We are very grateful to the anonymous referee for the useful comments and suggestions, which helped improve this paper. This work is based on data products from VVVX Survey observations made with the VISTA telescope at the ESO Paranal Observatory under program ID 198.B-2004. AM is funded by the National Agency for Research and Development (ANID) / Scholarship Program / DOCTORADO BECAS CHILE / 2021 - 21210535. DM is supported by the BASAL Center for Astrophysics and Associated Technologies (CATA) through grant AFB-170002, and by FONDECYT Regular grant No. 1170121. J.A.-G. acknowledges support from FONDECYT Regular 1201490 and ANID - Millenium Science Initiative Program - ICN12_009 awarded to the Millenium Institute of Astrophysics MAS. JCB acknowledges support from FONDECYT postdoctorado 3180716. R.K.S. acknowledges support from CNPq/Brazil through project 305902/2019-9. This research has made use of the TOPCAT\(^2\) tool (Taylor 2005), VOSA\(^3\), developed under the Spanish Virtual Observatory project supported by the Spanish MINECO through grant AyA2017-84089. This research has made use of "Aladin Sky Atlas" developed at CDS, Strasbourg Observatory, France. This research has also made use of the SIMBAD database, operated at CDS, Strasbourg, France. This work used data from the SpeX Prism Library service developed by the Spanish Virtual Observatory in the framework of the IAU Commission G5 Working Group: Spectral Stellar Libraries. This research has made use of the VizieR catalog access tool.

\(^2\) http://www.starlink.ac.uk/topcat/
\(^3\) http://svo2.cab.inta-csic.es/theory/vosa/
CDS, Strasbourg, France (DOI:10.26093/cds/vizier). The original description of the VizieR service was published in 2000, A&AS 143, 23. The authors also acknowledge J.B. Pullen.
Appendix A: Spectral energy distributions of the 99 low-mass objects

Spectral energy distributions with the model fit of each object included in this work are available online at https://github.com/andreamejiasb/SEDs_VVVX-GCNS_stars