First studies of the diffuse X-ray emission in the Large Magellanic Cloud with eROSITA

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ABSTRACT

Context. In the first months after its launch in July 2019, the extended Roentgen Survey with an Imaging Telescope Array (eROSITA) on board Spektrum-Roentgen-Gamma (Spektr-RG, SRG) performed long-exposure observations in the regions around supernova (SN) 1987A and supernova remnant (SNR) N132D in the Large Magellanic Cloud (LMC).

Aims. We analysed the distribution and the spectrum of the diffuse X-ray emission in the observed fields to determine the physical properties of the hot phase of the interstellar medium (ISM).

Methods. Spectral extraction regions were defined using the Voronoi tessellation method. The spectra were fit with a combination of thermal and non-thermal emission models. The eROSITA data are complemented by newly derived column density maps for the Milky Way and the LMC, 888 MHz radio continuum map from the Australian Square Kilometer Array Pathfinder (ASKAP), and optical images of the Magellanic Cloud Emission Line Survey (MCELS).

Results. We detect significant emission from thermal plasma with \(kT = 0.2\) keV in all the regions. There is also an additional higher-temperature emission component from a plasma with \(kT = 0.7\) keV. The surface brightness of this component is one order of magnitude lower than that of the lower-temperature component. In addition, non-thermal X-ray emission is significantly detected in the superbubble 30 Dor C. The absorbing column density \(N_H\) in the LMC derived from the analysis of the X-ray spectra taken with eROSITA is consistent with the \(N_H\) obtained from the emission of the cold medium over the entire area. Neon abundance is enhanced in the regions in and around 30 Dor and SN 1987A, indicating that the ISM has been chemically enriched by the young stellar population. In the centre of 30 Dor, there are two bright extended X-ray sources, which coincide with the stellar cluster RMC 136 and the Wolf-Rayet stars RMC 139 and RMC 140. For both regions the emission is best modelled with a high-temperature (\(kT > 1\) keV) non-equilibrium ionisation plasma emission and a non-thermal component with a photon index of \(\Gamma = 1.3\). In addition, we detect an extended X-ray source at the position of the optical supernova remnant (SNR) candidate J0529–7004 with thermal emission, and thus confirm its classification as an SNR.

Conclusions. Using data from the early observations of the regions around SN 1987A and SNR N132D with eROSITA we confirm that there is thermal interstellar plasma in the entire observed field. eROSITA with its large field of view and high sensitivity at lower X-ray energies allows us for the first time to carry out a detailed study of the ISM at high energies consistently over a large region in the LMC. We thus measure the properties of the interstellar plasma and the distribution of non-thermal particles and derive the column density of the cold matter on the line of sight.


1. Introduction

The Large Magellanic Cloud (LMC) is the largest satellite galaxy of the Milky Way at a distance of 50 kpc (e.g. de Grijs et al. 2014). It is a gas-rich dwarf galaxy with an almost face-on planar disk and many interesting asymmetric features, such as spiral arms, off-centre bar, and a shell-like structure (de Vaucouleurs 1955). Older stellar populations dominate the LMC mass and are either smoothly distributed in the disk or form a homogeneous stellar halo (Borissova et al. 2004, 2006). Young stars, on the contrary, are mainly found in the spiral arms (Younsi et al. 2019). The stellar populations indicate that the last major star formation in the Magellanic Clouds (MCs) occurred 200±50 Myr ago (e.g. Joshi & Panchal 2019, and references therein).

Optical emission line and radio images reveal that the LMC has a large number of bright H ii regions with a wide distribution...
in the disk. The most distinct object among them is the giant H\(_\text{ii}\) region 30 Doradus (30 Dor), also called the Tarantula nebula. It is located north of the eastern end of the stellar bar of the LMC and is host to the massive super-star cluster RMC 136.

The distribution of cold gas as seen in H\(_i\) emission (Mathewson et al. 1974; Kim et al. 2003) also reveals interesting structures in the eastern part of the LMC. As found by Luks & Rohlfs (1992), there are two distinct H\(_i\) components in the LMC, the L- and D-components. The D-component covers most of the LMC and is located in the plane of the galaxy disk. The L-component on the other hand is much more localised and has radial velocities that are \(\sim 30 \text{–} 60 \text{ km s}^{-1}\) lower than the D-component. The H\(_i\) gas distribution suggests that there was a close encounter between the LMC and the Small Magellanic Cloud (SMC) about 150 – 200 Myr ago (Fujimoto & Noguchi 1990). Fukui et al. (2017) revisited the H\(_i\) gas in the LMC and reported that the D- and L-components show a complementary distribution. They also found a third component, called the I-component, with radial velocities between the L- and D-component. The I-component was most likely formed by interactions between the two major components. In addition, Fukui et al. (2017) found a correlation between a large elongated complex of molecular clouds seen in CO called the CO ridge (Fukui et al. 1999; Mizuno et al. 2001; Fukui et al. 2008) and the likely interaction region. The CO ridge is located west of 30 Doradus and extends towards the south.

At high energies, the LMC is known to host bright supernova remnants (SNRs, e.g. N132D) and is a unique place where we can study the evolution of a supernova (SN) into an SN (SN 1987A). South of 30 Dor, east of the CO ridge, a large diffuse structure called the X-ray spur was observed in X-rays in the ROSAT survey of the LMC (Blondiau et al. 1997; Points et al. 2001). The X-ray spur is located south of 30 Dor and the supergiant shell LMC 2 (LMC-SGS 2) and seems to coincide with the regions where the L- and D-components of H\(_i\) collide with each other. The analysis of X-ray data taken with XMM-Newton has shown that the X-ray spur was most likely caused by the H\(_i\) collision (Knies et al. 2021). While in 30 Dor and in regions around it the interstellar medium (ISM) was heated by the stellar winds and supernovae of massive stars, there is no indication of stellar heating in the X-ray spur.

In this paper we present the analysis of the diffuse X-ray emission around 30 Dor and SNR N132D in the LMC based on the first observations with the extended ROentgent Survey with an Imaging Telescope Array (eROSITA; Merloni et al. 2012; Predehl et al. 2021), taken during the early commissioning and calibration phase of the mission. eROSITA is a new X-ray telescope, which was launched on July 13, 2019, on board the Spektrum-Roentgen-Gamma (SpektR-G, SRG) spacecraft. eROSITA consists of seven telescope modules (TMs), each equipped with a CCD detector. With its large field of view (FOV) with a diameter of \(\sim 1\) degree and high sensitivity in the energy band up to 10 keV, eROSITA allows us to study the distribution and the spectral properties of large extended X-ray emission for the first time. We supplement our study with the latest generation of radio survey data from the Australian Square Kilometer Array Pathfinder (ASKAP, Johnston et al. 2008; Hotan et al. 2021) and optical emission line images of the Magellanic Cloud Emission Line Survey (MCeLS, Smith et al. 2004). The studies of the point sources in the same observations and of SN 1987A are presented in separate papers by (Haberl et al., in prep.) and Maitra et al. (2021), respectively. We report on the study of the diffuse X-ray emission in the entire LMC, which will be based on the eROSITA All-Sky Survey (eRASS) data in future publications.

2. Data

2.1. eROSITA

During the journey to Lagrange point L2, where SRG has been operating in a stable orbit ever since, commissioning tests and calibration of the instruments were carried out, along with performance verification observations. The first-light observation of eROSITA was pointed at SN 1987A next to 30 Dor, while several calibration observations were performed around SNR N132D. We used the data from eROSITA observations in the LMC obtained during the commissioning and calibration phases (see Table 1). The observations were carried out in the pointing or the field-scan mode and were roughly pointed at PSR J0540–6919, SN 1987A, and SNR N132D, some of them with offsets of 20′ for calibration purposes. The data were processed with the eROSITA Science Analysis Software System (eSASS, Brunner et al. 2021). The eSASS pre-processing pipeline produces energy calibrated event files, which can then be used for the analysis. For the Early Data Release (EDR) we used the data of the processing version c001.

2.2. N\(_i\)\(_h\) map

X-rays interact via photoelectric absorption with baryonic matter distributed along the line of sight. The absorption cross section is a function of energy \(\propto E^{-3}\) (Wilms et al. 2000), implying that the low-energy portion of the total X-ray spectrum is strongly affected by photoelectric absorption. Hydrogen and helium are the most abundant elements in space. Their abundance ratio is still very close to the value defined by the primordial nucleosynthesis. Because of hydrogen’s dominance in abundance, the number of atomic and molecular hydrogen is used as a proxy for the total number of interstellar and intergalactic gas atoms.

The amount of neutral atomic hydrogen can be measured via the H\(_I\) 21cm line emission. Here, we make use of the HI4PI survey (HI4PI Collaboration et al. 2016)\(^1\), which is based on the data of the Effelsberg-Bonn HI Survey (EBHIS, Winkel et al. 2016) and the third revision of the Galactic All-Sky Survey (GASS, Kalberla & Haud 2015). HI4PI quantifies the to-

\(^1\) https://www.astro.uni-bonn.de/hisurvey/AllSky_profiles/
Fig. 2. Spatial distribution the Milky Way’s \( N_{\text{HI}} \) column density for \( V_{\text{LSR}} = -155 \) km/s to +80 km/s above a threshold of \( N_{\text{HI}} = 3 \times 10^{19} \) cm\(^{-2}\) with a maximum of \( 1 \times 10^{21} \) cm\(^{-2}\). Superimposed as contours is the spatial distribution of the Milky Way gas with excess optical extinction derived from the \( N_{\text{HI}}/A_V \) ratio (see text for details). The contours correspond to \( N_{\text{HI}} = 5 \times 10^{19} \) cm\(^{-2}\) and to 1, 2, and 3 \( \times 10^{20} \) cm\(^{-2}\). The LMC is located at \( l = 280^\circ, b = -33^\circ \), the SMC at \( l = 302^\circ, b = -44^\circ \). The yellow ellipse indicates the area observed with eROSITA.

Table 1. List of used eROSITA observations.

<table>
<thead>
<tr>
<th>ID</th>
<th>Target</th>
<th>Observation Type</th>
<th>Instruments</th>
<th>Start Date</th>
<th>Pointing Exposure [ks]</th>
</tr>
</thead>
<tbody>
<tr>
<td>700016</td>
<td>SN 1987A</td>
<td>Comissioning</td>
<td>TM3, 4</td>
<td>15-09-20</td>
<td>21</td>
</tr>
<tr>
<td>700161</td>
<td>SN 1987A</td>
<td>First Light</td>
<td>TM1, 2, 3, 4, 5, 6, 7</td>
<td>18-10-20</td>
<td>80</td>
</tr>
<tr>
<td>700156</td>
<td>SNR N132D</td>
<td>Calibration</td>
<td>TM5, 6−7</td>
<td>10-10-20</td>
<td>48</td>
</tr>
<tr>
<td>700179</td>
<td>SNR N132D</td>
<td>Calibration</td>
<td>TM1, 2, 3, 4, 5, 6, 7</td>
<td>22-11-20</td>
<td>60</td>
</tr>
<tr>
<td>700182</td>
<td>SNR N132D</td>
<td>Calibration</td>
<td>TM1, 2, 3, 4, 5, 6, 7</td>
<td>27-11-20</td>
<td>40</td>
</tr>
<tr>
<td>700183</td>
<td>SNR N132D</td>
<td>Calibration</td>
<td>TM1, 2, 3, 4, 5, 6, 7</td>
<td>25-11-20</td>
<td>40</td>
</tr>
<tr>
<td>700184</td>
<td>SNR N132D</td>
<td>Calibration</td>
<td>TM1, 2, 3, 4, 5, 6, 7</td>
<td>23-11-20</td>
<td>40</td>
</tr>
<tr>
<td>700185</td>
<td>SNR N132D</td>
<td>Calibration</td>
<td>TM1, 2, 3, 4, 5, 6, 7</td>
<td>25-11-20</td>
<td>40</td>
</tr>
<tr>
<td>700205</td>
<td>PSR J0540−6919</td>
<td>ART-XC Calibration</td>
<td>TM1, 2, 3, 4, 5, 6, 7</td>
<td>07-12-20</td>
<td>20</td>
</tr>
<tr>
<td>700206</td>
<td>PSR J0540−6919</td>
<td>ART-XC Calibration</td>
<td>TM1, 2, 3, 4, 5, 6, 7</td>
<td>07-12-20</td>
<td>30</td>
</tr>
<tr>
<td>700207</td>
<td>PSR J0540−6919</td>
<td>ART-XC Calibration</td>
<td>TM1, 2, 3, 4, 5, 6, 7</td>
<td>07-12-20</td>
<td>36</td>
</tr>
</tbody>
</table>

Regions with significantly low values for \( N_{\text{HI}}/A_V \) spatially mark molecular gas.

With the aim of getting a reliable estimate of the foreground \( N_{\text{HI,Gal}} \), because it modifies the soft X-ray emission from the LMC by photoelectric absorption, we cross-correlated the HI4PI data with the interstellar reddening \( E_{B-V} \) data from Schlegel et al. (1998). We applied the correction \( E_{B-V} \) (true) = 0.884 \( E_{B-V} \) (Schlafly & Finkbeiner 2011). According to Cardelli et al. (1989) and Weingartner & Draine (2001) the visual extinction of the diffuse interstellar medium is \( A_V = 3.1 E_{B-V} \), which is adopted in the following. We evaluated \( N_{\text{HI}}/A_V \) across an area of about 8000 deg\(^2\), which is sufficiently large to distinguish between the Magellanic Cloud System and the Milky Way halo gas.

For the Milky Way gas we find a value of \( N_{\text{HI}}/A_V = (2.04 \pm 0.15) \times 10^{21} \) cm\(^{-2}\) mag\(^{-1}\), which is compatible with Güver & Özel (2009) who determined \( (2.21 \pm 0.09) \times 10^{21} \) cm\(^{-2}\) mag\(^{-1}\). To obtain a map of the distribution of the excess extinction regions, we calculated the mean \( N_{\text{HI}}/A_V \) for the Milky Way halo gas by applying a spatial mask to exclude the H\( \alpha \) emission of the Magellanic Cloud System. Excess extinction regions, which deviate significantly from the field-averaged median value, are shown by the contour lines in Fig. 2, while the H\( \alpha \) data allow us to distinguish between the Milky Way and the Magellanic Cloud System emission. It is not possible to separate the optical extinction data. We do not find any evidence for additional neutral large-scale structures in the immediate vicinity of our field observed with eROSITA, indicating that there are no significant amounts of molecular gas in the Milky Way Galaxy in this area of the sky. The LMC X-ray radiation is dominantly modulated only by its intrinsic H\( \alpha \) distribution (see Sect. 6.4). On smaller angular scales additional soft X-ray shadows are expected because of the three dimensional structure of neutral and molecular interstellar medium that belongs to the LMC itself.

3. X-ray analysis

3.1. Images

eROSITA data were analysed using the eSASS version eSASSusers_201009. We created exposure-corrected mosaic images using the data in the broad band of 0.2 – 10.0 keV and in the following sub-bands: 0.2 – 0.5 keV, 0.5 – 1.0 keV, 1.0 – 2.0 keV, and 2.0 – 10.0 keV. To create the images, event files for all seven telescope modules (TMs) of eROSITA were binned with a bin size of 160 pixels, yielding counts images with a pixel size of 8\( '' \). These counts images and corresponding exposure maps were created for each observation and energy band. After combining the images and exposure maps into large mosaics of all observations, the mosaics of the counts images were divided by the
Fig. 3. Exposure-corrected mosaic image of the SN 1987A and SNR N132D regions (left) in three-colour presentation (red: 0.2 – 0.5 keV, green = 0.5 – 1.0 keV, blue = 1.0 – 2.0 keV) and a mosaic of exposure maps of all observations (right) in the entire energy range of 0.2 – 10.0 keV shown in linear scale in the range of 0 ks to 110 ks.

respective exposure-map mosaics to create exposure-corrected mosaic images. Figure 3 shows the mosaic images of all data in Table 1 below 2.0 keV and an exposure map of all observations in the broad band of 0.2 to 10.0 keV. To create the mosaic images we applied a lower cut for the exposure of 1σ below the mean value to mask out the outer areas of the FOVs in which the photon statistics were low.

In Wolter Type-I telescopes like eROSITA, single reflections of photons from nearby X-ray sources outside the FOV on the second hyperboloid mirror shells can also reach the detectors. Even though baffles in front of the telescopes can reduce the effect, it will result in stray light in the data and contaminate both the images and the spectra. In our case single reflections from photons from SNR N132D can become significant in observations to the east. We therefore carefully inspected the images for possible stray light, which would have been visible as an enhanced arc-like structure in one-half of the FOV of each observation on the side close to N132D. Fortunately, it caused no visible effect in the images. For the spectral analysis, for which we used the data of observations 700156, 700161, and 700179, any contamination by additional emission from N132D will only be visible in the spectra extracted from 700161. As no enhancement in emission or change in spectral parameters was found in the spectral fits in regions that might be affected by stray light from N132D, we assume, also for the spectral analysis, that the effect is negligible.

For the study of the diffuse emission, the standard eSASS source detection routine (Brunner et al., 2021) was applied to all observations. Source detection was performed in the 0.2 – 2.3 keV band for each observations using all available TMs. First, the sliding-box detection task erbox is run in local mode creating a preliminary list of sources, which is used to create a background map using the task erbackmap. Next, erbox is run in global mode using the background map. This second source list is used to create an updated background map. These two files are then used as input for the task emldet, which carries out point spread function fitting to the sources, yielding a final source catalogue with information such as the position, positional error, and detection likelihood. In the next step, all point sources are removed from the event file of each observation. We also manually defined regions for additional emission from very bright sources, which cause extended emission, and stray light from the X-ray binary LMC X-1, which is located close the FOV and causes narrow arc-like features at the southern edge of the FOVs in the eastern pointings, seen as blue emission in Fig. 3. These regions were also excluded from the data. We thus obtained ‘cheesed’ event files with holes at the position of the removed sources. Images were created from the cheesed event files for each observation in the same way as for the full mosaic. In Fig. 4 we show the cheesed mosaic images in the same energy bands as in Fig. 3.

3.2. Spectral analysis

Using the event files from which the point sources have been excluded, we extracted spectra in regions that were defined based on the photon statistics using the Voronoi tessellation algorithm (Cappellari & Copin 2003). For the spectral analysis, we focus on observations with SN 1987A and SNR N132D at or close to the aimpoints, which have high exposure (>20 ks, 700156, 700161, 700179). Voronoi tessellation was not performed on the entire data of each observation at once, since otherwise ob-
have shown that the diffuse X-ray emission in the ISM of normal galaxies requires at least two thermal plasma components with different temperature: $kT_1 = 0.2$ keV consistently in most cases, most likely emission from the hot ISM in equilibrium and from unresolved stellar sources, and $kT_2 > 0.5$ keV from regions, which experienced recent heating (i.e. H II regions, superbubbles, and SNRs) or also include unresolved binaries. For the plasma emission we tried both the collisional ionisation equilibrium model vapec\(^2\) and non-equilibrium ionisation model vne1 (Borkowski et al. 2001, and references therein). Since the lower-temperature component has an ionisation timescale $\tau = n_e t$ of $\sim 10^{13}$ s cm\(^{-3}\), and is thus consistent with collisional ionisation equilibrium, while that of the higher-temperature component is not well constrained, we decided to use the results obtained with two vapec model components for further discussion.

In addition, some SNRs and a few superbubbles are also known to emit significant non-thermal X-ray emission, with 30 Doradus C (30 Dor C) located close to SN 1987A being one of the only two non-thermal superbubbles known so far in the Local Group of galaxies. We therefore included a power-law component to model the emission in regions like 30 Dor C and to verify if non-thermal emission can also be detected outside the known non-thermal sources. For the foreground absorption, we included two components, one for the absorbing column in the Milky Way $N_{H,\text{Gal}}$, which is fixed to the value from the newly calculated map of the Galactic $N_{H,\text{cold,Gal}}$ (see Sect. 2.2), and another component for the LMC with $0.5 \times$ solar abundances. For the $N_{H}$, we use the model tbvarabs (Wilms et al. 2000).

For the fits we first set all element abundances to the average value of $0.5 \times$ the solar values (Russell & Dopita 1992). We assumed the element abundances reported by Anders & Grevesse (1989). In the next step we freed the abundances of O, Ne, and Mg in the hot thermal plasma emission component one by one since there is significant emission from thermal plasma up to 2.0 keV. We verified whether the fit improves with the new free parameter using the F-test statistics. If the change was not significant, the parameter was set back to $0.5 \times$ solar. Figure 5 shows the distribution of the values for $\chi^2$/d.o.f.

In Fig. 6 we show four example spectra with the best-fit models. The upper left spectrum is taken in the western wing of the giant H II region 30 Dor, while the upper right panel shows the spectrum of the central region of 30 Dor. The lower left spectrum was extracted in the western part of the non-thermal shell of 30 Dor C, the superbubble located south-west of 30 Dor. In addition, we also show one of the spectra taken in a less active region between SN 1987A and SNR N132D, where no bright Hα structures are observed. The best-fit parameters for the spectra are listed in Table 2.

4. Massive stars in 30 Doradus

The 30 Doradus region is known to host a large number of very massive stars (Feast et al. 1960), including the compact cluster of young massive stars RMC 136 in the centre. It contains the Wolf-Rayet star R136a1, which is one of the most massive stars detected to date (Crowther et al. 2010). Compact clusters of young massive stars are expected to be very efficient cosmic ray accelerators (for a review see Bykov 2014; Bykov et al. 2020). In X-rays bright diffuse emission has been detected from the central region in addition to point sources, which were resolved with Chandra (Townsend et al. 2006a,b).

\(^2\) http://atomdb.org/
With eROSITA, two bright sources are detected in the central region, one at the position of the star cluster RMC 136 and another source at the position of the Wolf-Rayet stars RMC 139 and RMC 140. We extracted the spectra of these two extended sources. The emission in both regions is well reproduced with thermal emission from a non-equilibrium ionisation plasma (\textit{vnei}) and non-thermal emission (\textit{powerlaw}). The temperature of the plasma is $kT > 1$ keV, and thus higher than in the surrounding regions with diffuse X-ray emission. The photon index of the power-law component is low ($\Gamma = 1.3$), indicating a hard X-ray spectrum (Table 3). Using the eROSITA spectrum we obtain a significantly enhanced abundance for Mg in both regions and for Si for RMC 136. Other elements are not well constrained due to poor photon statistics for the emission lines.

### 5. Supernova remnants

In the part of the LMC that was observed with eROSITA in the early phase, there are several known SNRs, which can all be confirmed in the eROSITA data (Fig. 7). In addition, there are two sources that are known to be SNR candidates (one radio and one optical source). The analysis of the eROSITA emission of these objects is presented in the following.

#### 5.1. SNR candidate J0529–7004

SNR candidate J0529–7004 is an arc-like structure seen in optical emission-line images with a high flux ratio of $[S\text{II}] / H\alpha = 1$. Therefore, Yew et al. (2021) have suggested it as an SNR candidate. An X-ray source was detected with ROSAT at its position (Haberl & Pietsch 1999). With eROSITA the position was observed in observation 700185. There is faint extended X-ray emission, as can be seen in the three-colour image in Fig. 8. We extracted spectra at the position of the optical SNR candidate.
Table 2. Fit parameters.

<table>
<thead>
<tr>
<th>Model</th>
<th>Parameter</th>
<th>30 Dor west</th>
<th>30 Dor centre</th>
<th>30 Dor C west</th>
<th>Between SNRs</th>
</tr>
</thead>
<tbody>
<tr>
<td>tbvarabs</td>
<td>$N_{H,LMC}$ [$10^{22}$ cm$^{-2}$]</td>
<td>0.32 (0.28 – 0.37)</td>
<td>0.55 (0.48 – 0.61)</td>
<td>0.54 (0.49 – 0.61)</td>
<td>0.083 (0.074 – 0.096)</td>
</tr>
<tr>
<td>vapec1</td>
<td>$kT_1$ [keV]</td>
<td>0.21 (0.19 – 0.22)</td>
<td>0.20 (0.18 – 0.22)</td>
<td>0.17 (0.14 – 0.22)</td>
<td>0.22 (0.21 – 0.23)</td>
</tr>
<tr>
<td>vapec1</td>
<td>$n$ [$10^{-3}$ cm$^{-3}$]</td>
<td>2.5 (1.9 – 3.6)</td>
<td>3.0 (1.7 – 5.1)</td>
<td>0.78 (0.29 – 2.0)</td>
<td>0.62 (0.57 – 0.69)</td>
</tr>
<tr>
<td>vapec2</td>
<td>$kT_2$ [keV]</td>
<td>0.73 (0.67 – 0.78)</td>
<td>0.40 (0.36 – 0.46)</td>
<td>0.62 (0.28 – 1.1)</td>
<td>0.81 (0.75 – 0.85)</td>
</tr>
<tr>
<td>vapec2</td>
<td>Ne (solar)</td>
<td>2.8 (2.1 – 3.5)</td>
<td>1.7 (1.2 – 2.7)</td>
<td>0.5 (N/A)</td>
<td>1.7 (0.5 – 2.8)</td>
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<tr>
<td>vapec2</td>
<td>Mg (solar)</td>
<td>1.0 (0.78 – 1.3)</td>
<td>1.3 (0.94 – 1.9)</td>
<td>0.5 (N/A)</td>
<td>0.5 (N/A)</td>
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<tr>
<td>powerlaw</td>
<td>photon index</td>
<td>&lt;3.5</td>
<td>8.7 (6.7 – 11.)</td>
<td>32. (31. – 34.)</td>
<td>&lt;1.6</td>
</tr>
<tr>
<td>powerlaw</td>
<td>$n$ [$10^{-5}$ cm$^{-3}$]</td>
<td>2.2 (2.0 – 2.5)</td>
<td>5.0 (3.0 – 7.5)</td>
<td>0.53 (0.13 – 2.0)</td>
<td>0.83 (0.73 – 0.94)</td>
</tr>
<tr>
<td>$\chi^2$</td>
<td></td>
<td>373</td>
<td>197</td>
<td>322</td>
<td>403</td>
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<tr>
<td>d.o.f.</td>
<td></td>
<td>295</td>
<td>175</td>
<td>267</td>
<td>328</td>
</tr>
</tbody>
</table>

Notes. The ranges in brackets are 90% confidence intervals.

Fig. 7. eROSITA mosaic images in three colours (same as in Fig. 3, left) and ASKAP 888 MHz image (Pennock et al. 2021, accepted). Known SNRs are shown in white, while the SNR candidates J0529–7004 and J0538–6921 are shown in magenta and cyan, respectively.

Table 3. Fit parameters for the emission at RMC 136 and RMC 139/140.

<table>
<thead>
<tr>
<th>Model</th>
<th>Parameter</th>
<th>RMC 136</th>
<th>RMC 139/140</th>
</tr>
</thead>
<tbody>
<tr>
<td>vnei</td>
<td>$kT$ [keV]</td>
<td>1.9 (1.4 – 2.9)</td>
<td>1.3 (1.1 – 1.7)</td>
</tr>
<tr>
<td>vnei</td>
<td>$\tau$ [$10^{11}$ s cm$^{-3}$]</td>
<td>0.9 (0.6 – 1.5)</td>
<td>1.4 (0.9 – 2.3)</td>
</tr>
<tr>
<td>vnei</td>
<td>Mg (solar)</td>
<td>3.9 (2.7 – 5.6)</td>
<td>3.5 (2.5 – 4.6)</td>
</tr>
<tr>
<td>vnei</td>
<td>Si (solar)</td>
<td>3.6 (2.1 – 5.4)</td>
<td>0.5 (N/A)</td>
</tr>
<tr>
<td>powerlaw</td>
<td>photon index</td>
<td>1.3 (1.1 – 2.4)</td>
<td>1.3 (1.2 – 1.5)</td>
</tr>
<tr>
<td>$\chi^2$</td>
<td></td>
<td>164</td>
<td>200</td>
</tr>
<tr>
<td>d.o.f.</td>
<td></td>
<td>97</td>
<td>104</td>
</tr>
</tbody>
</table>

Notes. The ranges in brackets are 90% confidence intervals.

Fig. 8. Three-colour eROSITA image of SNR J0529–7004 (red: 0.2 – 0.5 keV, green = 0.5 – 1.0 keV, blue = 1.0 – 2.0 keV) with contours of [S ii] emission (MCELS).

The local background was taken from a nearby region with no significantly enhanced emission, and was subtracted from the source spectrum. The eROSITA spectra are shown in Fig. 9 (upper panel). First, we fit the spectra with one thermal plasma model vapec or vnei, which did not yield a good fit. We therefore fit the spectrum again with a two-vapec model and obtain the best-fit with $\chi^2 = 132$ and 61 degrees of freedom (d.o.f.) with the following parameters: $kT_1 = 0.20$ (0.13 – 0.27) keV, $kT_2 = 0.68$ (0.62 – 0.74) keV. The foreground column density tends to zero with an upper limit of $N_H = 1.6 \times 10^{20}$ cm$^{-2}$. The best-fit model yields a flux (absorbed) of $F_X(0.2 – 10.0$ keV $) = (1.4 \pm 1.0) \times 10^{-13}$ erg s$^{-1}$ cm$^{-2}$. Due to poor photon statistics, element abundances could not be determined and were all set to 1.0.
with a flux (absorbed) of \( F_X(0.2 - 10.0 \text{ keV}) = (6 \pm 2) \times 10^{-14} \text{ erg s}^{-1} \text{ cm}^{-2} \). Due to the lack of a clear structure in X-rays and the very faint flux, this object cannot be confirmed as an SNR. Since this source is located in a region with enhanced diffuse emission both in the images of the optical line emission (see Fig. 10, lower right) and those in X-rays, it is difficult to identify a possible SNR. In addition, the diffuse emission in the optical and X-rays suggest that the region is filled with ionised gas and was most likely heated by interstellar shocks in the past. In such a region the SNR would be expanding in a medium with a density that is lower than in an unshocked medium, which would explain the lack of obvious X-ray and optical emission from the SNR.

6. Discussion

We created maps for all fit parameters by filling the extraction regions with the parameter values of the best-fit models. The parameter maps for the normalisation per arcmin\(^2\) for the three source emission components are shown in Fig. 10. Normalisation in XSPEC for the thermal model vapec is defined as

\[
norm_{\text{vapec}} = \frac{1}{10^{14} \times (4 \pi D_{\text{LMC}}^2)^{\frac{3}{2}}} \int n_e n_H dV \ [\text{cm}^{-5}],
\]

with \( n_e \) and \( n_H \) being the electron and hydrogen densities, respectively. For the power-law model the normalisation is \( \norm_{\text{powerlaw}} = \text{number of photons keV}^{-1} \text{ cm}^{-2} \text{ s}^{-1} \) at 1 keV. For the images in Fig. 10 the normalisation of the fit was divided by the size of the extraction region \( A \) in arcmin\(^2\).

6.1. Thermal component

In all regions there is significant emission from thermal plasma. We calculated the average temperature and its standard deviation from the parameter maps for \( kT_1 \) and \( kT_2 \), in which each pixel is filled with the best-fit parameter values. We thus obtain a mean temperature of \( kT_1 = 0.22 \) keV with a standard deviation of \( \sigma_{kT1} = 0.02 \) keV for the lower-temperature component and \( kT_2 = 0.74 \) keV and \( \sigma_{kT2} = 0.10 \) keV for the higher-temperature component. These temperature values are consistent with the results of the analysis of XMM-Newton observations of south-eastern parts of the LMC (regions around 30 Dor, in the X-ray spur, and west of them) by Knies et al. (2021). As can be seen in Fig. 10 (upper panels), the normalisation per area of the lower-temperature component (vapec1) is one order of magnitude higher than that of the higher-temperature component (vapec2).

From magnitude and colour measurements of stars in the LMC, the thickness of the disk was determined to be \( d = 3 \pm 1 \) kpc at the position of SNR N132D by Subramanian & Subramaniam (2009) and Rubele et al. (2012). The volume \( V \) of the plasma in each of the spectral analysis regions can be written as

\[
V \approx \frac{A \pi^2}{180} D_{\text{LMC}}^2 d,
\]

where \( A \) is the area of the extraction region in arcmin\(^2\), \( D_{\text{LMC}} = 50 \) kpc is the distance to the LMC, and \( d = 3 \) kpc is the thickness of the disk. By assuming a homogeneous depth of the volume, introducing a filling factor \( f \) for the plasma, and using the relation \( n_e = 1.2 n_H \), we can write

\[
n_H \approx \left( \frac{1.2 \times 10^{14} \times (4 \pi D_{\text{LMC}}^2) \norm_{\text{vapec}}}{fV} \right)^{\frac{1}{3}} = \left( \frac{5.6 \times 10^{22} \times \norm_{\text{vapec}}/A}{\pi f d} \right)^{\frac{1}{3}}.
\]
With the mean value of $\text{norm}_{\text{vapec1}}/A = (1.6 \pm 1.0) \times 10^{-5}$ cm$^{-2}$ arcmin$^{-2}$ (fit uncertainty of $\sim 10^{-4}$ and standard deviation of $0.4 \times 10^{-5}$) for the brighter, lower-temperature $\text{vapec1}$ component in the western regions around SNR N132D, we get a mean density of $n_{\text{H1}} = (3.1 \pm 1.5) \times 10^{-4} f^{-1/2}$ cm$^{-3}$. We assume that the emitting plasma can be regarded as an ideal gas with the pressure

$$P/k_B = (n_e + 1.1 n_{\text{H1}}) T = 2.3 n_{\text{H1}} T,$$

where $k_B$ is the Boltzman constant. The temperature $T_1$ of $\text{vapec1}$ and the hydrogen number density $n_{\text{H1}}$ derived from the normalisation yield a mean pressure of $P_1/k_B = (1.9 \pm 1.1) \times 10^3 f^2$ cm$^{-3}$K in the western regions. This pressure is consistent with the thermal pressure in the disk of the Milky Way (Cox 2005).

The higher normalisations of the thermal emission models in the eastern regions around 30 Dor compared to the regions around SNR N132D imply that the emission is brighter in and around 30 Dor and in the south-east than it is towards the west. These are the regions that also show bright line emission in the optical, as can be confirmed in the images of the Magellanic Clouds Emission Line Survey (MCELS, see Fig. 10, lower right, Smith et al. 2004). In the eROSITA mosaic images (Fig. 3) the emission in these regions appears to be fainter than or as faint as in the west. At the same time, the regions also appear green in the three-colour presentation, suggesting that the observed variation in surface brightness and colour is due to higher absorption. This will be discussed further in Sect. 6.4.

### 6.2. Non-thermal component

Non-thermal emission is observed significantly in the superbubble 30 Dor C (see Fig. 10, lower left, and Fig. 6, lower left). The photon index obtained from the eROSITA data is $\Gamma = 2.5 \pm 0.5$, which is in agreement with earlier studies (e.g. Kavanagh et al. 2015, 2019, and references therein). Recently, in a combined XMM-Newton and NuSTAR analysis, Lopez et al. (2020) showed that the complete shell of 30 Dor C is detected up to 20 keV. The authors found that contrary to prior studies, their region D (part of 30 Dor C West here) requires a thermal component of $kT = 0.86 \pm 0.01$ keV, which is consistent with the eROSITA fit with the temperature of 0.62 (0.28 – 1.1) keV given in Table 2. The photon indices of a power-law component obtained by eROSITA for the entire 30 Dor C West region is $\Gamma = 2.4$ (2.3 – 2.6), which is somewhat higher than the value of $2.12 \pm 0.02$ obtained for region D, but consistent with $2.39 \pm 0.03$ obtained for region C, which were both covered by our 30 Dor C West region. SNR MCSNR J0536–6913, which was found earlier with XMM-Newton (Kavanagh et al. 2015; Babazaki et al.

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**Fig. 10.** Parameter maps of the normalisation per arcmin$^2$, illustrating the surface brightness of the diffuse emission, for the two thermal components ($\text{vapec1, vapec2}$) and the powerlaw component. The images are shown in log-scale. The letters W, X, Y, Z in the labels correspond to the exponents of the marks of the colour scale. The red plus sign indicates the position of the star cluster RMC 136, the yellow cross that of SNR/PWN N157B. The lower right panel shows H$\alpha$ (red), [S ii] (green), and [O iii] (blue) images of MCELS in three-colour presentation.
2018) and is apparent in the eROSITA image (Figure 7, left), was not detected above 8 keV in the NuSTAR observation.

Synchrotron radiation of very high-energy electrons and positrons is a likely source of a substantial amount of the observed non-thermal X-rays. The high-energy leptonic component is accelerated in supernova remnants, superbubbles, and pulsar wind nebulae, and can be produced by inelastic collisions of cosmic ray nuclei. The high-energy leptons producing the synchrotron X-rays simultaneously up-scatter photons from ambient radiation field. The inverse-Compton radiation from multi-TeV leptons producing the observed X-ray synchrotron apparently contribute to the TeV gamma-ray emission detected from the pulsar wind nebula (PWN) N157B and the superbubble 30 Dor C by the ground-based Cherenkov gamma-ray telescope H.E.S.S. (see e.g. H. E. S. S. Collaboration et al. 2015). Observations with H.E.S.S. yielded a luminosity of \( \sim 0.9 \times 10^{35} \text{ erg s}^{-1} \) for 30 Dor C for the 1 – 10 TeV gamma-ray band with a power-law distribution of a photon index 2.6 ± 0.2, which is consistent with the photon index of X-rays detected by eROSITA in 30 Dor C west (see Table 2). The observed TeV emission apart from the inverse-Compton mechanism mentioned earlier could originate from the decay of neutral pions produced in the inelastic collisions of relativistic nuclei with the ambient matter (the so-called hadronic scenario). While the expected fluxes in the leptonic scenario can explain the observed fluxes of X-rays and TeV gamma-rays, as well as the measured widths of the non-thermal X-ray filaments (Kavanagh et al. 2019), the hadronic origin of the observed gamma-rays cannot be excluded (see e.g. H. E. S. S. Collaboration et al. 2015).

There is also significant non-thermal emission in the central region of 30 Dor (see Fig. 6, upper right and Table 2). This non-thermal component can indicate some contamination by emission from the composite SNR with a PWN N157B, located south-west of the region (yellow cross in Fig. 10, lower left). The emission of N157B itself was removed when point and point-like sources were cut out from the event files. The X-ray spectrum of N157B consists of thermal emission of the SNR and the dominant non-thermal power-law emission from the pulsar and the PWN with a photon index of \( \Gamma = 2.29 \pm 0.05 \), while the pulsar shows a power-law emission with \( \Gamma = 1.73 \pm 0.11 \) (Chen et al. 2006). The high lower limit of \( \Gamma > 2.5 \) for the photon index for the non-thermal emission detected inside 30 Dor suggests that it is not caused by emission from N157B. It should be noted, however, that the gamma-ray spectrum of N157B detected by H.E.S.S. (H. E. S. S. Collaboration et al. 2012) can be described well by a power-law component with a photon index \( \Gamma = 2.8 \pm 0.2 \) in the energy range between 600 GeV and 12 TeV.

Since non-thermal emission is also found to the north of the super-star cluster RMC 136 (red plus sign) and possibly in a larger region towards the east, it might also be caused by particles that were accelerated in the shocks of the winds of massive stars inside RMC 136 and are diffusing inside and probably also outside of the nebula. Diffuse non-thermal X-ray emission was detected with \textit{Chandra} from the most massive Galactic star cluster Westerlund 1 (Muno et al. 2006). They estimated the X-ray luminosity to be about \( 3 \times 10^{44} \text{ erg s}^{-1} \), and found that while the photon index of the power-law X-ray component is \( \Gamma = 2.7 \pm 0.2 \) within the circle of 1′ radius around the cluster, the spectrum gets harder with \( \Gamma = 2.1 \pm 0.2 \) in the annulus between 1′ and 2′ and even \( \Gamma = 1.7 \pm 0.1 \) between 2′ and 3.5′. Possible mechanisms of the origin of the non-thermal emission in the clusters of young massive stars were discussed by Bykov (2014). Further studies will be performed in the future when additional data of the eROSITA All-Sky Survey is available.

6.3. Element abundances

The eROSITA spectra suggest that the element abundance of neon is enhanced in the regions around 30 Dor (Table 2, Fig. 11). The mean value of the Ne abundance around 30 Dor is 2.25 times solar, with a mean value of the 90% error of 1.16. In the regions around N132D the mean Ne abundance is 1.93 and the mean

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value of the 90% error is 0.70. In total, a free parameter for the Ne abundance improved the fit in 235 regions, while for 151 regions, the fit was consistent with a Ne abundance of 0.5 times solar.

There are also line transitions of oxygen and magnesium in the energy range of 0.4 – 2.0 keV, in which eROSITA has the highest response. However, the O and Mg lines are contaminated by background emission (solar-wind charge exchange lines and Al Kα line at 1.486 keV from the detector, respectively), which were included in the background model as free components. Therefore, the O and Mg abundances are not well constrained. The O abundance was fixed to 0.5 times solar in all regions since freeing the parameter did not improve the fit. In the case of Mg, in the east, the mean value is 1.62 with a mean 90% error of 0.76 (the Mg abundance was free in 70 out of 247 regions), and in the west the mean abundance is 1.59 with a mean 90% error of 0.35 (free in only 4 out of 139 regions). Therefore, O is consistent with 0.5 times the solar value in all regions, and there is no clear gradient seen in the Mg abundance.

To understand the possible origin of Ne, we calculated and plotted the star formation rate in these two areas in LMC (marked with boxes in Fig. 11) based on the star formation rates of Harris & Zaritsky (2009) (Fig. 12). In the eastern region (labelled ‘1’), star formation has been ongoing over the last 10$^6$ – 10$^7$ years, producing a large number of young massive stars that have created the large and complex HII regions. Most likely, the ISM has also been enriched by stellar winds and supernovae of these massive stars.

### 6.4. Foreground absorption in the LMC

The fits show that the foreground column density $N_{\text{HI}}$, in the LMC is higher around 30 Dor (see Fig. 13) than in the west, also indicated by the green colour of the X-ray image. The values obtained from the analysis of the X-ray spectrum taken with eROSITA is in very good agreement with the $N_{\text{HI}}$ in the LMC obtained from the distribution of cold matter (Sect. 2.2), as can be seen in Fig. 13. This agreement between the values derived from the X-ray analysis and measured directly from the cold medium seen over a large area of $>1 \degree$ is remarkable, and shows the advantage of the large FOV and the high sensitivity of eROSITA in the energy band of 0.2 – 2.0 keV.

Together with the high normalisation (and thus intrinsic flux) of the emission in the same regions, the enhanced $N_{\text{HI}}$ indicates that the hot X-ray emitting plasma is located behind a high-density region. The same stars that created the giant H II region 30 Dor and the other H II structures in its surroundings are the likely sources of the hot interstellar plasma. This result corroborates nicely the scenario of the collision of large HI structures in the LMC for the origin of 30 Dor and the star-forming regions south of 30 Dor presented by Fukui et al. (2017) and recently confirmed by Knies et al. (2021) based on a multi-wavelength study of the south-eastern part of the LMC, in particular the X-ray spur. A large HI component, called the L-component encountered the HI component in the disk of the LMC (D-component) in the past, starting roughly at the position where 30 Dor is observed now. The collision continued south of 30 Dor where young massive stars and, even further to the south, a long ridge of dust and CO clouds along with star-forming regions are found next to the X-ray spur. The L-component is now located in front of the disk of the LMC at the position of 30 Dor, which corresponds to the eastern part of the eROSITA observations that we analysed, and absorbs the emission from 30 Dor and from the sources and the hot interstellar plasma around it.

### 7. Summary

The first-light observation of eROSITA was pointed at SN 1987A and has provided us with an impressive X-ray view of a large number of sources and complex diffuse emission in the LMC in a large 1$\degree$ diameter field. This field includes several prominent objects: the giant H II-region 30 Dor, the non-thermal superbubble 30 Dor C, and bright SNRs and pulsars. At a distance of about 1.5$\degree$ from SN 1987A there is SNR N132D, which is the brightest X-ray SNR in the LMC. As one of the X-ray calibration sources, SNR N132D and its surroundings have also been observed in the early phase of the eROSITA mission multiple times. In this paper we have presented the study of the diffuse X-ray emission in the LMC detected in these early eROSITA observations.

To understand the properties of the hot interstellar plasma and the processes that form it, we have been studying SNRs and superbubbles in the Magellanic Clouds using XMM-Newton and Chandra (e.g. Sasaki et al. 2011; Kavanagh et al. 2012; Warth et al. 2014; Kavanagh et al. 2015, 2019). Due to the much smaller FOV of these X-ray observatories, our studies have so far had to focus on a small number of selected objects. The large FOV of eROSITA and the high sensitivity at energies below 2 keV combined with the much better spatial and spectral resolution than ROSAT, which was the last X-ray observatory that performed a survey of the entire LMC (for a study of the diffuse emission, see Sasaki et al. 2002), makes eROSITA the perfect instrument with which to study the hot ISM in our neighbour galaxy. Thanks to the early commissioning and calibration observations, we obtained a set of long exposures with eROSITA towards the most interesting regions of the LMC.

We have analysed the data of 11 eROSITA observations by eliminating all point and point-like sources and focusing on the diffuse emission. Using a tessellation algorithm, we divided the data into small regions with a signal-to-noise ratio of $>100$ and $>50$ around SN 1987A and SNR N132D, respectively. The spectra of the diffuse emission in these regions were analysed assuming a combination of a two-component thermal plasma and one non-thermal emission model. Maps of the parameter values obtained for the best-fit model have been created and studied. The results of the spectral analysis are as follows:

- We detect emission from thermal plasma in the ISM in the LMC in all regions. There is dominant emission from a low-temperature component with $kT = 0.2$ keV and another lower-brightness thermal component with $kT \approx 0.7$ keV. The interstellar density and pressure derived from the parameters of the major lower-temperature component are consistent with values measured in the Milky Way.
- The emission from the higher-temperature thermal component is stronger in the environment of 30 Dor, suggesting that the young stellar population has caused recent heating.
- In these regions, the element abundances also seem to be enhanced, as indicated by the high Ne abundance. Most likely, the ISM has been enriched by the stellar winds and supernovae of massive stars.
- In addition, significant non-thermal emission is confirmed in the superbubble 30 Dor C. There are indications of the presence of non-thermal emission also east of 30 Dor, which requires further investigation.
- The X-ray spectral analysis yields an absorbing column density $N_H$ in the LMC, which is surprisingly consistent with the column density derived from the measurements of H I and the gas-to-dust ratio at low energies. This is the first time that the
foreground column density has been determined through X-ray spectroscopy over such a large contiguous field and is in agreement with direct measurements from the cold interstellar medium.

- We analysed the spectra of the massive stellar cluster RMC 136 and the emission from the Wolf-Rayet stars RMC 139 and RMC 140. The emission is nicely reproduced by a model consisting of emission from a non-ionisation equilibrium plasma with \( kT > 1 \) keV and \( \tau = n_e \tau_i \approx 10^{3.1} \) s cm\(^{-1}\) and a power-law component with \( \Gamma \approx 1.3 \).

- Based on eROSITA image and spectroscopy, we confirm SNR J529–7004 as a new SNR in the LMC.

As we have been anticipating during the years of preparations for the eROSITA mission, eROSITA is the perfect telescope for studying the ISM at high energies. Currently, eROSITA is carrying out a total of eight all-sky surveys over a period of four years. We will extend the study of the ISM in the LMC based on the eROSITA all-sky survey data. In addition, the eROSITA all-sky survey will also allow us to study the hot phase of the ISM in the LMC.

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References


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Fig. 13. \( N_{\text{H, LMC}} \) in the LMC obtained from the fit of the eROSITA spectra (left) vs. \( N_{\text{H, LMC}} \) directly calculated from cold H\(_{\text{I}}\) and dust (right) in cm\(^{-2}\). Both images are shown in log-scale using the same lower and upper cuts. The contours of \( N_{\text{H, LMC}} \) are plotted in both panels.
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