

A new hot core in the outer Galaxy: Impact of metallicity on the formation of complex organic molecules

Youxin Wang^{1,*,***}, Arnaud Belloche¹, and Robin T. Garrod²

¹ Max-Planck-Institut für Radioastronomie, Auf dem Hügel 69, 53121 Bonn, Germany

² Departments of Chemistry and Astronomy, University of Virginia, Charlottesville, VA 22904, USA

Received 28 November 2025 / Accepted 20 April 2026

ABSTRACT

Context. Many complex organic molecules (COMs) in star-forming regions are believed to form on dust grains. We thus expect both the reduced metallicity and dust-to-gas ratio in the outer Galaxy to have an impact on the chemical composition of these regions.

Aims. We investigate if certain COMs are more sensitive than others to metallicity by measuring the chemical composition of hot cores in the outer Galaxy.

Methods. We used the interferometer NOEMA to perform an imaging spectral line survey of G135.27+2.79, a hot core candidate located at a galactocentric distance of 13.1 kpc. We derived the rotational temperatures and column densities of the detected molecules while assuming local thermodynamic equilibrium and compared the chemical composition of G135.27+2.79 to other sources and to the predictions of the three-phase astrochemical code MAGICKAL.

Results. G135.27+2.79 hosts three continuum cores, labeled MM1, MM2, and MM3. Toward MM1, we detected 28 molecules, including 12 COMs, and several of their less abundant isotopologs. Most species trace a hot, compact region, confirming MM1 as a hot core (the third one identified in the outer Galaxy). MM1 drives a bipolar CO outflow. COMs show a velocity gradient along the outflow axis but opposite to that of CO, which may be related to a wide-angle disk wind. The chemical composition of MM1 correlates rather well with that of the inner and outer Galaxy hot cores G31.41+0.31 and WB89-789 SMM1, but its molecular abundances relative to methanol lie in between, which may reflect the influence of metallicity on COM formation. The model results agree reasonably well, though with a few notable exceptions, with the COM abundances of MM1 relative to methanol and with the abundance ratios between MM1 and G31.41+0.31. Sensitivity to the reduced metallicity and dust-to-gas ratio varies between molecules, with carbon chains and nitriles most negatively affected. The lower dust-to-gas ratio leads to slower adsorption under low-metallicity conditions so that more carbon is locked up into CO in the gas. Slow adsorption means that CO is hydrogenated more efficiently on grains, enhancing CO-related COM abundances above expectations.

Conclusions. These results demonstrate that metallicity has a significant impact on the formation of COMs. A larger source sample is needed to investigate the robustness of the deviations noted between the model and the observations for a few species.

Key words. astrochemistry – line: identification – stars: formation – ISM: molecules

1. Introduction

Complex organic molecules (COMs) are defined as carbon-bearing molecules containing six or more atoms (Herbst & van Dishoeck 2009). These molecules are important tracers of the chemical complexity in the interstellar medium (ISM). COMs have been detected in a wide variety of astrophysical environments, ranging from warm inner envelopes surrounding protostars (e.g., Jørgensen et al. 2016; Belloche et al. 2020; Chen et al. 2023) to protoplanetary disks (e.g., Walsh et al. 2013; Lee et al. 2019; Fadul et al. 2025) and even to extragalactic environments, such as the center of the nearby starburst galaxy NGC 253 (Martín et al. 2021; Bouvier et al. 2025; Huang et al. 2025). How these species are formed has been the focus of much attention and debate (e.g., Jørgensen et al. 2020). COMs are generally thought to form and evolve in both the gas phase and the solid phase, the latter consisting of the surfaces and ice mantles of dust grains (e.g., Hasegawa et al. 1992;

van Dishoeck & Blake 1998; Garrod et al. 2008). Observations over the past decade, particularly the detection of gas-phase COMs at low temperatures in prestellar cores and protostellar envelopes (e.g., Bacmann et al. 2012; Vastel et al. 2014; Bouscasse et al. 2024), and more recently the identification of several COMs in interstellar ices using the James Webb Space Telescope (JWST; e.g., Rachid et al. 2022; Yang et al. 2022; Rocha et al. 2024) suggest that the solid phase plays a major role in the production of COMs.

Hot molecular cores associated with high-mass star-forming regions ($M > 8 M_{\odot}$; (Walmsley 1995)) and their low-mass counterparts, known as hot corinos (Ceccarelli 2004), serve as excellent laboratories for investigating the chemical complexity of the ISM. Hot cores are compact (< 0.1 pc), dense ($n > 10^6$ cm⁻³) regions that form around nascent high-mass protostars. They represent a transient yet chemically rich phase during the early stages of high-mass star formation (e.g., Kurtz et al. 2000; Cesaroni 2005). As a protostar forms, it begins to heat its surrounding envelope. When the temperature rises above ~ 100 K, thermal desorption of the ice mantles coating interstellar dust grains occurs, releasing a wide variety of COMs into the gas phase (Jørgensen et al. 2020). The advent of broadband instrumentation, especially in interferometers over the past

* Corresponding author: yxwang@mpifr-bonn.mpg.de

** Member of the International Max Planck Research School (IMPRS) for Astronomy & Astrophysics (<https://blog.mpi-fr-bonn.mpg.de/imprs/>) at the Universities of Bonn and Cologne.

decade, has significantly enhanced the feasibility of conducting unbiased spectral line surveys of hot cores and hot corinos, thereby advancing our understanding of COM formation. Notable examples include the Exploring molecular complexity with ALMA (EMoCA)/Re-exploring Molecular Complexity with ALMA (ReMoCA) project, which targeted hot cores in Sgr B2(N) in the Galactic center region (e.g., Belloche et al. 2016, 2019, 2025; Bonfand et al. 2017; Busch et al. 2024); the G31.41+0.31 Unbiased ALMA sPectral Observational Survey (GUAPOS) project, focused on the hot core G31.41+0.31 in the inner Galaxy (e.g., Mininni et al. 2020, 2023); and the ALMA Protostellar Interferometric Line Survey (PILS) project, which studied the low-mass protostellar system IRAS 16293-2422 in the solar neighborhood (e.g., Jørgensen et al. 2016; Calcutt et al. 2018; Jørgensen et al. 2018). These surveys have enhanced our understanding of chemical complexity in the inner Galaxy and, surprisingly, revealed that the chemical compositions of Sgr B2(N2), G31.41+0.31, and IRAS 16293-2422B are remarkably similar when oxygen- and nitrogen-bearing COMs are normalized separately (Jørgensen et al. 2020; Belloche et al. 2025). This raises the question whether this similarity still holds beyond the solar circle, where the metallicity is reduced. Our goal is to investigate if metallicity has an impact on the prevalence of COMs in star-forming regions of our Galaxy.

The outer Galaxy refers to regions beyond the solar circle, with Galactocentric distances, D_{gc} , greater than 8.15 kpc (Reid et al. 2019). The disk of our Galaxy has a metallicity gradient between 5 and 17 kpc, as measured in various elements (e.g., C, N, O, Fe) that exhibit different radial trends (e.g., Luck & Lambert 2011). For instance, at distances of 12–17 kpc from the center, the metallicity traced by Fe is about 1.7 to 3.4 times lower than at the solar circle. Several studies of star-forming regions in the outer Galaxy have investigated the impact of metallicity on molecular abundances (Bernal et al. 2021; Fontani et al. 2022a,b; Koelemay & Ziurys 2025). These studies generally report that the abundances of simple molecules and (cold) CH_3OH , when normalized to H_2 , remain comparable to those in the inner Galaxy. However, these abundances depend critically on the adopted H_2 column density, which is typically inferred from dust continuum emission and therefore sensitive to assumptions about the gas-to-dust ratio (a value of 100 was used in these papers). However, the gas-to-dust ratio is expected to increase with galactocentric distance (Giannetti et al. 2017). Indeed, when accounting for this effect, Fontani et al. (2024) and Gigli et al. (2025) reported decreasing molecular abundances toward larger galactocentric radii. Reporting abundances relative to another molecule, such as CH_3OH , allows us to avoid the uncertainties related to the derivation of H_2 column densities.

Given that many COMs form on dust grains and contain multiple heavy atoms (Garrod & Herbst 2006), their formation is likely influenced by the metallicity and dust-to-gas ratio. Consequently, the chemical composition of the ice mantles of dust grains (and the gas phase, after desorption) in the outer Galaxy may differ from that in the inner Galaxy. In other words, the outer Galaxy serves as an ideal laboratory to study the impact of metallicity and the dust-to-gas ratio on the formation of COMs (Shimonishi 2024). Moreover, studies in the outer Galaxy are critical for bridging our understanding between the normal-metallicity environment in the inner Galaxy and the low-metallicity environment observed in nearby dwarf galaxies such as the Large Magellanic Cloud (LMC) and Small Magellanic Cloud (SMC). Recently, six hot cores have been identified in the LMC (Shimonishi et al. 2016, 2020; Sewilo et al. 2018, 2022; Hamedani Golshan et al. 2024) and two in the SMC

(Shimonishi et al. 2023). However, only two hot cores, WB89-789 SMM1 and Sh 2-283-1a SMM1, have been detected in the outer Galaxy so far (Shimonishi et al. 2021; Ikeda et al. 2026). To better understand the impact of metallicity on the formation of COMs, two actions are essential: building a larger sample of hot cores in the outer Galaxy observed through unbiased spectral line surveys and comparing their chemical composition with the predictions of astrochemical models.

This study focuses on the search for hot cores embedded in the outer Galaxy, and we investigate the impact of metallicity on the formation of COMs. To this end, we first conducted single-dish telescope observations to identify promising hot core candidates. These candidates were then followed up at high angular resolution with an unbiased spectral line imaging survey using the Northern Extended Millimeter Array (NOEMA) with the goal of characterizing their physical structure and chemical composition. In this work, we present and analyze NOEMA observations of G135.27+2.79, the hot core candidate of our sample that was observed first. The source selection, the observational setup, and the data reduction are described in Sect. 2. Section 3 presents an analysis of the observations and a comparison to other sources. Section 4 compares the chemical composition derived for G135.27+2.79 to the predictions of the chemical model MAGICKAL. The results are discussed in Sect. 5, and Sect. 6 presents our conclusions.

2. Observations and data reduction

2.1. Source selection

Our sample was constructed by identifying candidate high-mass star-forming regions in the outer Galaxy based on the presence of water or Class II methanol masers, which are indicators of young stellar objects (YSOs; Kurtz et al. 2000). Class II methanol masers, in particular, are known to trace exclusively high-mass YSOs (e.g., Minier et al. 2003; Ellingsen 2006). We focused on sources located in the second and third Galactic quadrants with heliocentric distances of less than 10 kpc, enabling comparison with well-studied regions in the inner Galaxy. We selected all sources fulfilling these criteria beyond a galactic radius, D_{gc} , of 12 kpc from the Effelsberg water maser surveys of IRAS sources (Wouterloot et al. 1988, 1993), the Parkes/ATCA/MERLIN methanol multi-beam survey (Green et al. 2017), the 6.7 GHz methanol maser catalog (Pestalozzi et al. 2005), and the Bar and Spiral Structure Legacy (BeSSeL) Very Long Baseline Interferometry survey (Reid et al. 2019, and reference therein). Moreover, we kept only sources with a maser position that was known sufficiently accurately compared to the beams of the IRAM 30m telescope at 3 mm /1.3 mm and the Atacama Pathfinder Experiment (APEX) telescope at 1.3 mm, either from interferometric observations or from cross-scans in the case of Effelsberg. We also included an IRAS source at $D_{\text{gc}} = 15.8$ kpc with no maser detection but detected in H_2CO (Blair et al. 2008). Our final sample comprises 19 sources spanning D_{gc} from 12 to 17 kpc. The seven targets in the second quadrant were observed with the IRAM 30m telescope and the twelve sources in the third quadrant with the APEX telescope. These observations achieved deep integrations, with rms sensitivities at a spectral resolution of $\sim 1 \text{ km s}^{-1}$ of about 5 mK for the APEX observations, and ranging from 5.9 to 7.1 mK and from 8.1 to 18.2 mK for the IRAM 30 m observations at 3 mm and 1.3 mm, respectively. This enabled a sensitive search for high-temperature rotational emission of COMs. Cold methanol emission was detected in all sources, but hints of the presence

of a hot core were found in only four targets (weak emission of CH_3CN or torsionally excited CH_3OH).

The first of our hot core candidates that we subsequently observed at higher angular resolution, G135.27+2.79 (also known as WB89-437, IRAS 02395+6244), is a luminous far-infrared source detected with IRAS (Wouterloot & Brand 1989). The distance to G135.27+2.79 was determined through trigonometric parallax measurements of its associated H_2O maser, yielding a heliocentric distance of 6.0 ± 0.2 kpc (Hachisuka et al. 2009), and a Galactocentric distance of 13.1 kpc (assuming a Galactocentric distance of the Sun of 8.15 kpc, Reid et al. 2019). The trigonometric distance is shorter than the kinematic distance assumed earlier by Wouterloot & Brand (1989). Adjusted to the new distance, the IRAS luminosity of $7.1 \times 10^4 L_\odot$ reported by Wouterloot & Brand (1989) becomes $3.1 \times 10^4 L_\odot$, which corresponds to a B0 star if it is a single object (Panagia 1973). At this distance, the expected metallicity for G135.27+2.79 is $[\text{Fe}/\text{H}] \approx -0.3$ dex, based on the Galactic radial abundance gradient $\frac{d[\text{Fe}/\text{H}]}{dD_{\text{gc}}} = -0.061$ dex kpc $^{-1}$ (Luck & Lambert 2011). This corresponds to a metallicity $Z \sim 0.5 Z_\odot$ (see Sect. 4 for the corresponding elemental abundances). The source is associated with several H_2O maser spots (Wouterloot et al. 1993; Hachisuka et al. 2009), strong CO wing emission tracing an outflow (Wouterloot & Brand 1996), but no radio continuum emission was detected, suggesting that it is extremely young or that it consists of a cluster of lower mass stars (Rudolph et al. 1996). Additionally, several simple C-bearing species (e.g., HCN, C_4H , and $\text{c-C}_3\text{H}_2$) and the cold ($T_{\text{ex}} < 20$ K) emission of two COMs (CH_3CCH and CH_3OH) were detected toward this source (Bernal et al. 2021; Fontani et al. 2022a,b; Gigli et al. 2025; Koelemay & Ziurys 2025).

2.2. NOEMA observations

We conducted a spectral line survey toward G135.27+2.79 using the NOEMA interferometer¹ in the frequency range 215.1–246.4 GHz. The observations were carried out in the C configuration with nine antennas on 2023 October 5, December 3, and December 5. The phase center was $(\alpha, \delta)_{(J2000)} = 02^{\text{h}}43^{\text{m}}28^{\text{s}}.57, +62^\circ57'08''.4$, which is the water maser position from the BeSSeL survey (Hachisuka et al. 2009; Reid et al. 2019). The half-power beam width (HPBW) of the primary beam of the 15 m antennas is about $23''$ at 1.3 mm.

The spectral line survey was divided into two spectral setups, S1 and S2. Each setup was observed in two polarizations and delivered four spectral windows of bandwidth 3.9 GHz each, two per sideband (LO, LI, UI, UO). The on-source integration time was 6 h for S1 and 5.2 h for S2. The baselines of the array ranged from 19.5 m to 328.0 m for S1 and from 17.7 m to 368.0 m for S2. The observations were performed with a channel spacing of 250 kHz using the PolyFiX correlator. The spectra were smoothed to a resolution of 500 kHz ($0.6\text{--}0.7$ km s $^{-1}$). Bandpass calibration was performed using observations of 3C84, and flux calibration was performed using observations of LKHA101 and MWC349. The maximum recoverable scale was about $8''$, which was calculated using Eq. (3.28) in Cortés et al. (2025). Phase and amplitude calibration was performed using J0244+6228 and 0224+671. Details about the observational setups are given in Table A.1. Following Maury et al. (2019), we assume an absolute flux calibration uncertainty of 15%.

2.3. Data reduction

The data were calibrated with GILDAS/CLIC² and then imaged with IMAGER version 4.4³. We ran the IMAGER pipeline step by step to perform self-calibration, imaging, and splitting of the continuum and line emission. To improve the data quality, the self-calibration was conducted in SURVEY mode without a line catalog. Within spectral windows S1-UI and S2-LI, the CO line is very strong, dominates the integrated flux, and is spatially extended, making the SURVEY mode inappropriate for self-calibration. Therefore, we split each of these windows into two sub-windows and applied self-calibration separately: in SPLIT mode for the sub-window containing the CO line and in SURVEY mode for the other one. The splitting of the line and continuum emission was performed with the C-SCM method of IMAGER. The size (HPBW) of the synthesized beam, the corresponding linear size, and the rms noise level in the final cubes are given in Table A.2. The noise level of each spectral window corresponds to the median of the noise levels measured in all channel maps using the procedure *go noise* in GREG/GILDAS².

3. Results and analysis

3.1. Continuum emission

Although the *uv* coverage and array configurations differ between the two NOEMA setups, the resulting continuum emission maps are largely consistent in structure. Figure 1 presents the continuum maps of spectral windows S1-UI-CO and S2-LI-CO, which have a similar angular resolution (HPBW $\sim 0.8''$). While slight variations in morphology are evident, the overall structure remains comparable: each map reveals three continuum peaks, which we named MM1, MM2, and MM3. MM1 and MM2 are located at equatorial offsets ($0.23''$, $0.15''$) and ($4.3''$, $4.5''$) from the phase center, respectively. These positions were determined by averaging the results from all sub-bands using 2D Gaussian fits. Determining the peak position of MM3 with a Gaussian fit is challenging due to its proximity to MM1. Instead, we constructed the continuum profile along the direction going through the peak positions of redshifted and blueshifted SiO components detected around MM3 (see Sect. 3.3.3 for details) and identified the corresponding peak continuum position of MM3 as $(-2.1'', 1.7'')$.

To investigate the nature of the detected emission in MM1, the in-band spectral index, α , was analyzed across the eight sub-bands. α was determined by fitting the flux densities to the formula $F_\nu \propto \nu^\alpha$ using the *scipy* function *curve_fit*. The fluxes F_ν in all sub-bands were extracted using an aperture with a diameter of $1.8''$, which corresponds to approximately twice the beam size of the lowest spatial resolution (S1-LO). Figure E.1 shows the result of the in-band spectral index fitting. We find $\alpha = 3.3 \pm 1.4$, indicating that the continuum emission is dominated by dust emission at a scale of $1.8''$. We then obtain $\beta = \alpha - 2 = 1.3$, which is similar to what Bonfand et al. (2017) derived for the hot core Sgr B2(N2). Such a flat index suggests a dust opacity spectrum intermediate between models of dust grains of Ossenkopf & Henning (1994) without ice mantles and with thin ice mantles.

The optical depth τ at the peak position of MM1 was estimated using the following equation (Bonfand et al. 2017):

$$\tau = -\ln \left(1 - \frac{F_\nu^{\text{beam}}}{\Omega_{\text{beam}} B_\nu(T_d)} \right), \quad (1)$$

² <https://www.iram.fr/IRAMFR/GILDAS/>

³ <https://imager.oas.u-bordeaux.fr/>

¹ <https://iram-institute.org/science-portal/noema/>

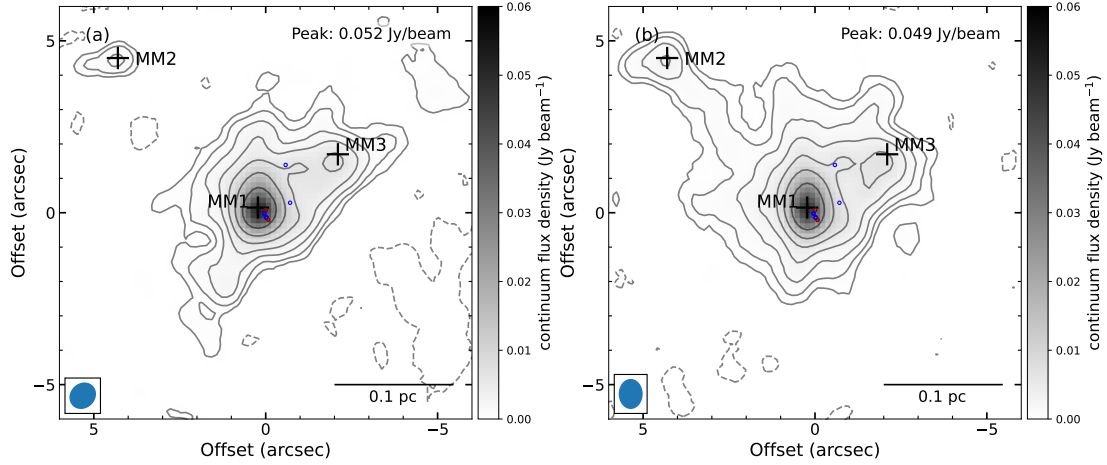


Fig. 1. Continuum maps of G135.27 at 233 GHz (a) and 229 GHz (b) in gray scale, extracted from setups S1-UI-CO and S2-LI-CO, respectively. Contour levels start at 3σ and increase by a factor of 2, where $\sigma = 0.13$ and 0.12 mJy beam $^{-1}$, respectively. Negative contours at -3σ are shown as dashed lines. The black crosses mark three continuum peaks of G135.27, labeled MM1, MM2, and MM3 (see Sect. 3.1), with the peak flux density of MM1 indicated in the upper right corner. The empty circles denote the positions of H₂O masers (Hachisuka et al. 2009), with blue and red indicating maser spots blue- and redshifted with respect to the source systemic velocity of -71.7 km s $^{-1}$, respectively (Blair et al. 2008). The beam size (HPBW) is displayed as an ellipse in the bottom left corner. The offsets are defined with respect to the phase center (see Sect. 2.2). The maps are not corrected for primary-beam attenuation.

with $\Omega_{\text{beam}} = \frac{\pi}{4 \ln 2} \times \text{HPBW}_{\text{maj}} \times \text{HPBW}_{\text{min}}$ as the solid angle of the synthesized beam, T_d as the dust temperature, and $B_\nu(T_d)$ as the Planck function. The term F_ν^{beam} is the peak continuum flux density in Jansky per beam, corrected for primary beam attenuation. Assuming that dust and gas are well coupled, we adopted a dust temperature (T_d) of 75 K. This value corresponds to the rotational temperature (T_{rot}) of CH₃CCH toward the continuum peak position, determined from our population diagram analysis (see Sect. 3.4.3 and Fig. H.2). We used the temperature derived from CH₃CCH because, as with the dust continuum emission, this molecule traces more extended structures than other COMs. The opacity calculations for all sub-bands yield values lower than 0.04, confirming that the dust continuum emission at the scale traced with NOEMA in this source is optically thin.

3.2. H₂ column density and total gas mass of MM1

The H₂ column density of MM1, N_{H_2} , was estimated from the millimeter dust continuum emission. Since this emission is optically thin at the observed frequencies (Sect. 3.1), N_{H_2} was calculated using the following equation:

$$N_{\text{H}_2} = \frac{F_\nu \xi_d}{\Omega \kappa_\nu B_\nu(T_d) \mu_{\text{H}_2} m_{\text{H}}}, \quad (2)$$

where F_ν is the peak continuum flux density, Ω is the beam solid angle, κ_ν is the dust mass opacity (in square centimeter per gram of dust) at frequency ν , ξ_d is the gas-to-dust ratio, μ_{H_2} is the mean molecular weight per H₂ (2.8, Kauffmann et al. 2008), and m_{H} is the hydrogen mass. We assumed $\kappa_\nu = \kappa_0(\nu/\nu_0)^\beta$, with $\kappa_0 = 1.6$ cm² g⁻¹ at 1.3 mm as Bonfand et al. (2017), since we have a similar value of β (see Sect. 3.1).

To estimate the gas-to-dust ratio, ξ_d , we used the C¹⁸O and dust emission integrated over an aperture of 1.8'' in diameter. Here again, we used CH₃CCH to estimate the temperature at a diameter of 1.8'' (which is similar to the C¹⁸O emission size; see Fig. C.2), assuming that the rotational temperatures of CH₃CCH and C¹⁸O, as well as the dust temperature, are the same under local thermodynamic equilibrium (LTE) conditions. Employing

the CH₃CCH spectrum averaged over a ring of inner and outer diameters of 0.9'' and 2.7'', respectively, we derived a temperature of 63 K from its population diagram (see Sect. 3.4.3 and Fig. H.2). The gas mass surface density was derived from the C¹⁸O column density, since the C¹⁸O line is optically thin, with an opacity of ~ 0.2 according to our Weeds model (see Sect. 3.4). The total column density N_{tot} of C¹⁸O was calculated using the following equation:

$$N_{\text{tot}} = \frac{8\pi k_B \nu^2 W}{hc^3 A_{ul} b g_u} Q \exp\left(-\frac{E_u}{k_B T_{\text{rot}}}\right), \quad (3)$$

where k_B is the Boltzmann constant, W is the integrated intensity over the full line profile after correction for primary beam attenuation, h is the Planck constant, c is the speed of light, A_{ul} is the Einstein coefficient for spontaneous emission, b is the beam filling factor with a value of one for C¹⁸O, g_u is the degeneracy of the upper level, Q is the partition function, and E_u is the upper-level energy. This yielded a C¹⁸O column density of 3.3×10^{16} cm⁻², with the integrated intensity measured from the spectrum averaged over an aperture of diameter 1.8''. We applied Eq. (1) of Giannetti et al. (2017), which assumes a CO abundance of 9.5×10^{-5} relative to H₂ in the solar neighborhood, to derive the total gas mass surface density from C¹⁸O. Following their approach, we adopted a C/H gradient of -0.08 dex kpc⁻¹ from Luck & Lambert (2011) and used the ¹⁶O/¹⁸O isotopic ratio relation, $^{16}\text{O}/^{18}\text{O} = (58.8 \pm 11.8) D_{\text{gc}} + (37.1 \pm 82.6)$, from Wilson & Rood (1994). The dust mass surface density was calculated as $\Sigma_{\text{dust}} = \frac{S_\nu / \Omega}{\kappa_\nu B_\nu(T_d)}$, with S_ν the flux density integrated over an aperture of diameter 1.8'' and Ω the solid angle of this aperture. This yields a gas-to-dust mass ratio of 303 ± 66^4 . The Galactocentric gradient of gas-to-dust ratio determined by Giannetti et al. (2017) yields a ratio of ~ 380 at a Galactocentric distance of 13.1 kpc, but the dispersion is large. Our result is consistent with this estimate within the uncertainties, and we thus adopted a

⁴ Using a smaller aperture equal to the beam size gives a gas-to-dust ratio of 335 ± 73 , which is consistent with this result within the uncertainties.

value ξ_d of 303 for the rest of the analysis. This yields a peak H_2 column density of $(1.1 \pm 0.1) \times 10^{24} \text{ cm}^{-2}$, without accounting for the calibration uncertainty. Assuming that the source diameter is equal to the beam size, we obtained a mean gas number density of $n_{H_2} = 1.5 \times 10^7 \text{ cm}^{-3}$. Assuming a typical density profile with a slope of -1.5 , the gas number density at $0.3''$, corresponding to the typical size of the COM emission (see Sect. 3.4.2), is $\sim 8 \times 10^7 \text{ cm}^{-3}$.

Furthermore, the gas mass of the envelope within a diameter of $3.0''$ can be estimated using the following equation:

$$M_{\text{gas}} = \frac{S_\nu D^2 \xi_d}{\kappa_\nu B_\nu(T_d)}, \quad (4)$$

where S_ν is the flux density integrated over the aperture and D is the heliocentric distance to G135.27. Assuming a radial power-law profile for the temperature, we obtained an exponent of -0.25 using the two CH_3CCH temperatures mentioned above (75 K for $0.9''$ and 63 K for $1.8''$), yielding a temperature of 55 K at a diameter of $3.0''$ by extrapolation. Assuming the same temperature for the dust, we obtained a total gas mass of $49 M_\odot$ within a radius of 9000 au.

3.3. Kinematics

3.3.1. Outflow morphology

Wouterloot & Brand (1996) reported strong line wings in CO 1–0 and 2–1 toward G135.27 based on IRAM 30m observations, suggesting the presence of an outflow. Follow-up millimeter interferometric observations with BIMA, using a synthesized beam of $21.8'' \times 14.9''$, confirmed this outflow, revealing blue- and redshifted lobes oriented toward the northeast and southwest, respectively, as well as a region of overlap between the lobes near the dust continuum peak (Brand et al. 2001). Our NOEMA data allowed us to explore the outflow morphology at a higher angular resolution.

Figure 2 presents the redshifted and blueshifted intensity maps of CO and ^{13}CO integrated over the ranges listed in Table F.1 (see Appendix F for details), overlaid on the 1.3 mm continuum emission. Blueshifted emission and redshifted emission are predominantly observed in the north-east and south-west, respectively. The outflow axis is oriented at a position angle of approximately 68° east from north, which we defined by the longest extent of the blueshifted emission and passing through the continuum peak position of MM1. Notably, there is a region of overlap between the blue- and redshifted emissions near the center of MM1, possibly suggesting a wide opening angle of the southeastern lobe along the line-of-sight, maybe similar to the large opening angle of the redshifted lobe seen in the plane of the sky. The outflow direction is consistent with that reported in Brand et al. (2001). Multiple H_2O maser features were reported around MM1 at angular separations of approximately 0.1 – $0.2''$, $0.8''$, and $1.5''$ from MM1, spanning a velocity range from -76 to -62 km s^{-1} (see Fig. 2 and Table 3 in Hachisuka et al. 2009). Their proximity to MM1 and location within the overlap region of the blue- and redshifted emission suggest that they are associated with the outflow driven by MM1.

Figure 3a displays the position–velocity (PV) diagram of CO 2–1 taken along the outflow axis defined in Fig. 2. It reveals two components. The blueshifted component that extends up to about $16''$ from MM1 in Fig. 3 has projected velocities that reach $\sim 27 \text{ km s}^{-1}$ at that distance. The second component is compact and has much higher red- and blueshifted projected velocities

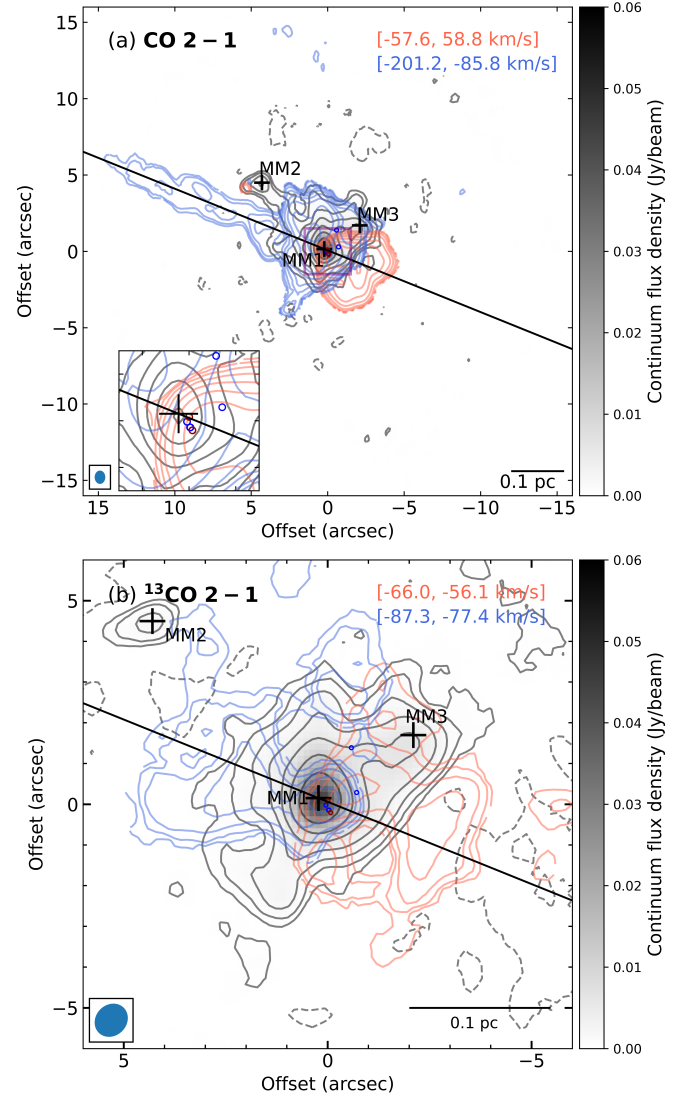


Fig. 2. Morphology of the outflow traced by CO (a) and ^{13}CO (b). In each panel, the blue and red contours represent the blue- and redshifted emission, respectively. The integration ranges indicated in the top-right corner. The blue and red contours start at 3σ and then increase by a factor of two, where $\sigma = 34.1$ (CO, red), 34.0 (CO, blue) in panel a, and $10.2 \text{ mJy beam}^{-1} \text{ km s}^{-1}$ for both red and blue contours in panel b. The solid line represents the outflow axis. The black contours show the 1.3 mm continuum emission as in Fig. 1, and the symbols are the same as in that figure. In panel a, a zoomed-in view of size $1'' \times 1''$ around MM1 (within the purple square) is shown in the bottom-left corner to better display the positions of the H_2O masers. The maps are not corrected for primary-beam attenuation.

(up to about 100 km s^{-1}). It is detected up to $\sim -2.5''$ and $\sim 0.3''$ toward the southwest and northeast of MM1, respectively. We estimated the dynamical timescales of these two velocity components using the relation

$$t_{\text{dyn}} = \frac{R_{\text{proj}}}{V_{\text{proj}}} \times \frac{\cos(i)}{\sin(i)}, \quad (5)$$

where R_{proj} and V_{proj} are the projected length and line of sight gas velocity, respectively, and i is the inclination angle of the outflow with respect to the line of sight. The blue- and redshifted lobes exhibit a distinct bipolar outflow structure with a small overlap near the center (see Fig. 2), indicating that the outflow axis is inclined at an intermediate angle to the line of sight.

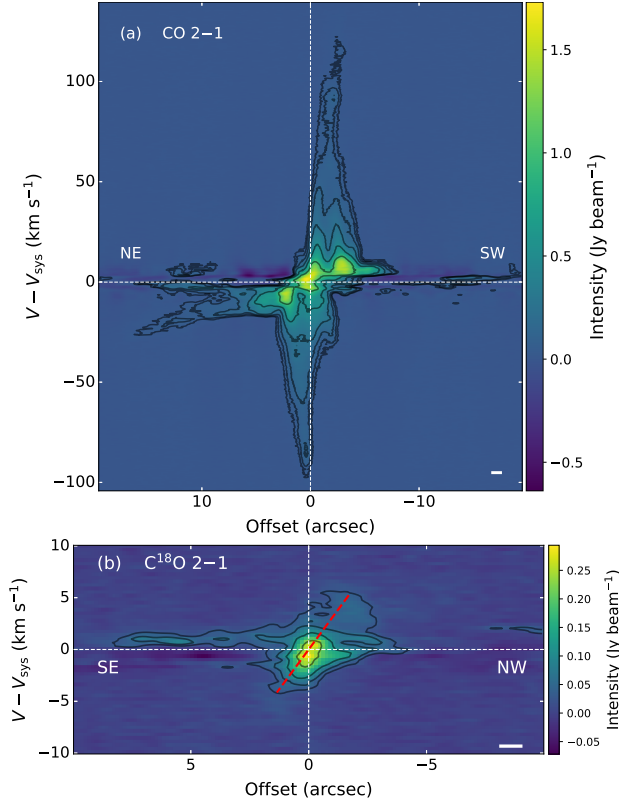


Fig. 3. Position–velocity diagrams of CO (panel a) and C^{18}O (panel b) extracted along the outflow axis (northeast to southwest) defined in Fig. 2 and perpendicularly to it (southeast to northwest), respectively. The y axis shows the velocity with respect to the systemic velocity of the source, displayed by the horizontal dashed line. The vertical white line represents the continuum peak position of G135.27 MM1. The horizontal bar in the bottom-left corner of each panel displays the beam size along the direction of the cut. The dashed red line in panel b highlights the velocity gradient traced by C^{18}O .

Table 1. Dynamical timescales of the outflow components assuming an inclination angle of 45° .

Component	R_{proj} (")	V_{proj} (km s^{-1})	t_{dyn} (yr)
Compact-red	2.5	122	5.8×10^2
Compact-blue	0.3	97	88
Extended-red	8	5	4.6×10^4
Extended-blue	16	27	1.7×10^4

We thus adopted an inclination angle of 45° . The dynamical timescales for the compact and extended velocity components are 88–580 yr and 17 000–46 000 yr, respectively (see Table 1). These distinct timescales likely reflect different episodes of mass ejection: the compact component traces a recent, high-velocity outflow event occurring close to the protostar, whereas the extended component corresponds to slower, more dispersed material ejected during an earlier, sustained phase. The dynamical timescale of the extended component is comparable to those measured in Galactic high-mass protostellar objects (Beuther et al. 2002; Maud et al. 2015), as well as in the outer Galaxy source Sh2–283–1a SMM1 ($D_{\text{gc}} = 15.7$ kpc, Ikeda et al. 2025).

Compared to CO, C^{18}O usually traces circumstellar envelopes (see, e.g., Hsu et al. 2025). Its PV diagram extracted

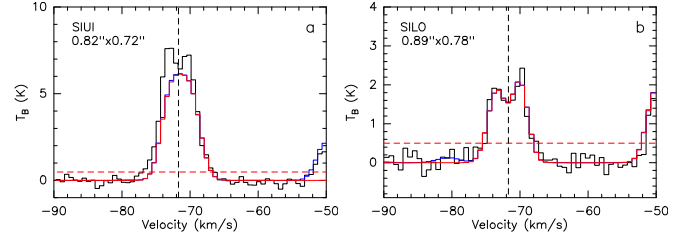


Fig. 4. Spectra of a $\text{CH}_3\text{OH } v_l = 0$ ($\nu_{\text{rest}}=232.783$ GHz, $E_u=447$ K) and b $\text{CH}_3\text{OCHO } v_l = 0$ ($\nu_{\text{rest}}=218.281$ GHz, $E_u=100$ K) toward the G135.27 MM1 continuum peak position. In each panel, the velocity axis refers to the rest frequency of the transition. The vertical and horizontal dashed lines indicate the systemic velocity and the 3σ noise level, respectively. The blue spectrum shows the LTE model incorporating all identified molecular species, while the red spectrum displays the LTE model including only CH_3OH (panel a) and CH_3OCHO (panel b). The observational setup and spectral window are labeled in the top-left corner of each panel, along with the corresponding HPBW.

perpendicularly to the CO outflow axis shows a clear velocity gradient from the bottom left to the top right, with redshifted and blueshifted emission on opposite sides of the protostar (Fig. 3b). This indicates systematic motion of the gas, possibly due to rotation within the circumstellar envelope (e.g., Tobin et al. 2012; Hsu et al. 2025).

3.3.2. Kinematics traced by CH_3OH and CH_3OCHO

The spectra of CH_3OH and CH_3OCHO toward MM1 reveal two distinct velocity components, as illustrated in Fig. 4, which shows one rotational transition of each molecule, selected for the absence of contamination from other species. The spectra of most other COMs can also be fitted with two velocity components (see details in Sect. 3.4.1 and Table B.1). Small differences in the peak velocities of the two components reported in Table B.1 among different molecules are smaller than the channel width and likely arise from the limited spectral resolution. In contrast, most simple molecules exhibit only one velocity component, indicating a complex kinematic structure within MM1.

Figure 5 presents maps of the red- and blueshifted emission of the transitions shown in Fig. 4. Additionally, the redshifted emission peaks of all CH_3OH and CH_3OCHO transitions that have a peak signal-to-noise ratio higher than three and are not contaminated by other species are marked with red crosses, while their blueshifted emission peaks are indicated by blue crosses. The redshifted and blueshifted components of CH_3OH and CH_3OCHO do not peak at the same position. They are shifted along a direction roughly parallel to the outflow axis, with the redshifted component closer to MM1 and the blueshifted component further away toward the southwest ($\sim 0.1''$ or 600 au from MM1). This velocity gradient is also visible in the moment 1 maps shown in Fig. G.1. Its direction is surprisingly opposite to the orientation of the outflow. This anti-alignment implies that the kinematics traced by CH_3OH and CH_3OCHO are not dominated by the outflow. A similar anti-alignment has also been detected on small scales (~ 300 au) in the NGC 1333 IRAS 4A2 protostar (De Simone et al. 2024). We discuss the possible origin of this opposite velocity gradient in Sect. 5.2.

3.3.3. A second outflow traced by SiO

Notably, SiO serves as another powerful tracer of protostellar outflows or jets because its abundance is significantly enhanced

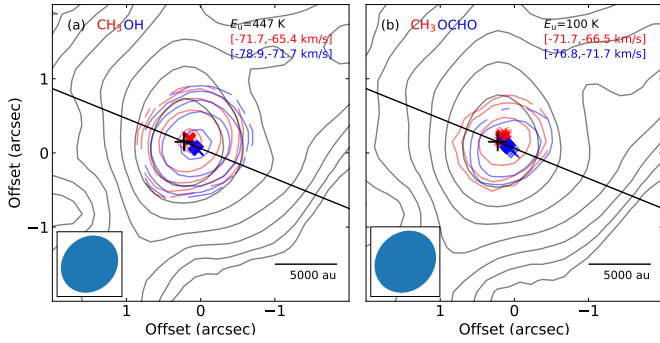


Fig. 5. Blueshifted (blue contours) and redshifted (red contours) emission of the transitions shown in Fig. 4. Contours start from 3σ and increase by a factor of two, where $\sigma = 10.9$ (CH₃OH, red) and 11.6 (CH₃OH, blue) in panel a, and 9.9 (CH₃OCHO, red) and 9.8 mJy beam⁻¹ km s⁻¹ (CH₃OCHO, blue) in panel b. The redshifted and blueshifted emission peaks of all selected transitions of CH₃OH and CH₃OCHO are marked with red and blue crosses, respectively. The integration ranges are indicated in the top-right corner of each panel. The continuum beam size is indicated as an ellipse in the bottom-left corner of each panel. The black line represents the outflow axis traced by the large-scale CO emission. The black cross indicates the continuum peak position, while the black contours represent the continuum emission starting at the 3σ level and increasing by factors of two.

in shocked gas due to the sputtering of silicon from dust grains (e.g., [Bachiller & Pérez Gutiérrez 1997](#); [Bally 2016](#); [Liu et al. 2025](#)). We present the red- and blueshifted emission of SiO in Fig. 6a. Notably, an additional bipolar outflow centered on MM3 was identified through high-velocity SiO emission relative to the systemic velocity of MM3 (-70.7 km s⁻¹), which was determined from the CN and HC₃N spectra toward the MM3 position. Both molecules exhibit a local maximum in their emission at this position. The position angle of this outflow is about 66° east from north, which is similar to that of MM1.

Figures 6b-c display SiO ($J = 5-4$) spectra, extracted at the peaks of the blueshifted and redshifted emission, respectively. The SiO line shows a symmetric shift in the blue and red directions relative to the systemic velocity of MM3, with peak emission occurring at $|V_{\text{peak}} - V_{\text{sys}}| \sim 64$ km s⁻¹. The spectral characteristics of the high-velocity SiO emission closely resemble those of extremely high velocity (EHV) outflows, which are typically associated with the earliest stages of protostellar evolution and exhibit velocities exceeding 30 km s⁻¹ (e.g., [Bachiller 1996](#); [Tafalla et al. 2010](#); [Tychoniec et al. 2019](#)). This strongly suggests that MM3 is a YSO driving an outflow. Part of the blueshifted and redshifted CO emission seen around MM3 in Fig. 2 may result from this second outflow, which would in part explain the complex morphology of the CO emission.

At lower velocity, only redshifted SiO emission is detected, with no corresponding blueshifted emission (see magenta contours in Fig. 6a). Its distribution shows three distinct features. The central feature near MM1 and the top-right feature near MM3 are both elongated along their respective outflow directions, suggesting that they originate from these outflows. Surprisingly, the third feature consists of two spots located in the southeast of MM1, at a position angle of about 130 deg east from north and apparently not related to MM1's outflow. Spectra from the farthest spot ($3.3''$, $-2.5''$) reveal lines from other molecules, including CO, CS, SO, SO₂, and tentative detections of H₂CO and CH₃OH. The presence of sulfur-bearing species (SO and SO₂), as well as SiO, is commonly linked to shock activity (e.g., [Jiménez-Serra et al. 2008](#); [Tychoniec et al. 2021](#);

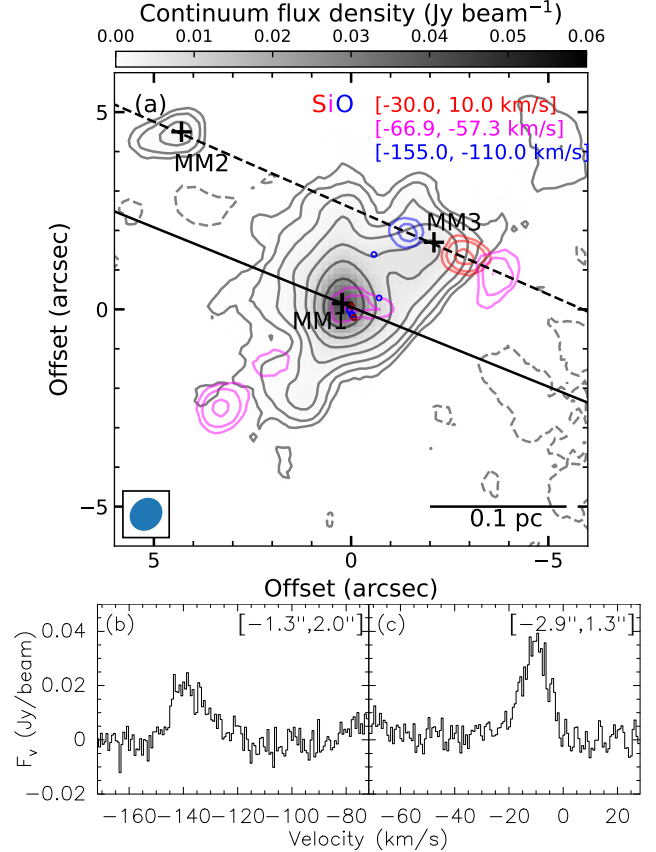


Fig. 6. Panel a: extremely high velocity blueshifted (blue contours) and redshifted (red contours) lobes of the outflow associated with MM3 traced by SiO 5–4. The lower-velocity redshifted emission is shown in magenta. The integration range for each component is indicated in the top-right corner. The contours start at 3σ and increase by a factor of two, where $\sigma = 27.5$ (red), 13.5 (magenta), and 29.1 (blue) mJy beam⁻¹ km s⁻¹. The black contour levels of the continuum emission and the symbols are the same as in Fig. 1. The solid and dashed lines mark the axes of the outflow associated with MM1 and MM3, respectively. The SiO spectra extracted from the peak positions of the blue- and redshifted emission are shown in panels b and c respectively. The positional offsets relative to the phase center are indicated in the top-right corner of each panel.

[Artur de la Villarmois et al. 2022](#)). CH₃OH is also efficiently released into the gas phase via sputtering and desorption during shock events ([Suutarinen et al. 2014](#)). Therefore, the chemical composition of the stronger spot suggests that it results from some shock activity, the origin of which is still unknown.

3.4. Spectral line modeling with Weeds

3.4.1. Method of modeling

To identify the molecular lines in the observed spectra and derive column densities and excitation temperatures, we used the software Weeds ([Maret et al. 2011](#)). Weeds⁵, a part of the GILDAS/CLASS software, enables the generation of synthetic spectra under the LTE assumption. The calculation of synthetic spectra relies on spectroscopic data obtained from a local database. In this work, we used the same database as [Belloche et al. \(2025\)](#), which includes spectroscopic data from the

⁵ <https://www.iram.fr/IRAMFR/GILDAS/doc/html/weeds-html/weeds.html>

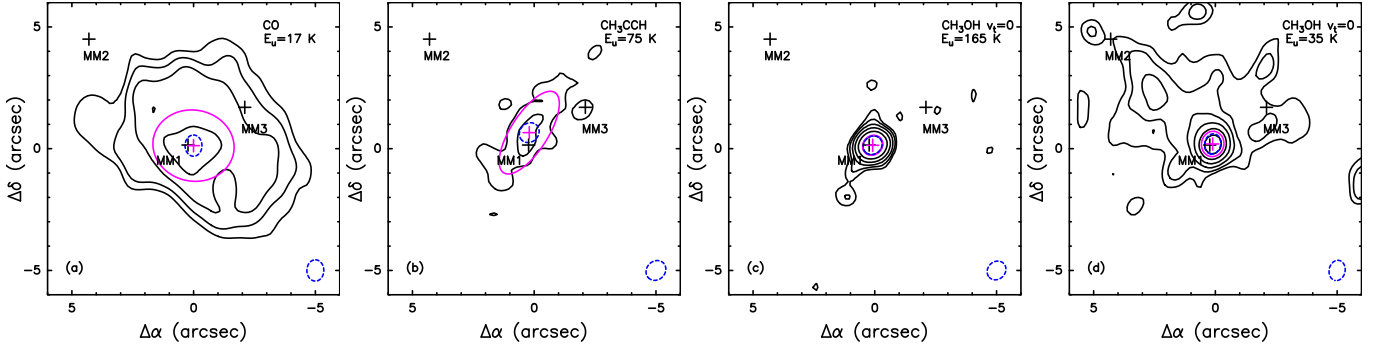


Fig. 7. Integrated intensity maps of CO 2–1 (a), CH₃CCH 13₀–12₀ (b), CH₃OH 10_{2,9}–9_{3,6} (c), and CH₃OH 5_{0,5}–4_{0,4} (d). The values of the noise level, σ , are 80.7, 11.6, 14.3, and 16.5 mJy beam^{−1} km s^{−1}, respectively. The contours start at 3σ and increase by a factor of two at each step. In each panel, the molecule name and the upper-level energy of the line are shown in the upper right corner, the black crosses indicate the continuum peak positions, the magenta cross and ellipse represent the peak position and FWHM of the 2D Gaussian fit, and the blue dashed ellipses (at the continuum peak and in the bottom right corner) represent the beam size.

Cologne Database for Molecular Spectroscopy (CDMS⁶, Müller et al. 2001, 2005; Endres et al. 2016), the Jet Propulsion Laboratory (JPL⁷) spectroscopy database (Pickett et al. 1998), and the Lille Spectroscopic Database (LSD⁸, Motiyenko & Margulès 2025), as well as additional predictions contributed by spectroscopists. For CH₂DOH, we used a spectroscopic entry provided by L. Coudert that is based on Coudert et al. (2014) rather than the JPL entry used by Belloche et al. (2025), which has numerous issues at higher frequencies. For CO, H₂S, HDO, DCN, and ³³SO₂, which were not modeled in Belloche et al. (2025), we used entries from the CDMS database, or from JPL in the case of HDO.

Practically, the modeling in Weeds requires five input parameters for each species: line width, velocity offset relative to the systemic velocity, V_{sys} , of the source, temperature, column density, and emission size. The line width and velocity offset relative to V_{sys} are determined by performing one-dimensional Gaussian fits to optically thin and unblended transitions of each molecule. The same line width and velocity offset are then adopted for all transitions of that molecule. Almost all simple species show a single velocity component. CH₃CCH and CH₂DOH clearly show only one velocity component, while most COMs exhibit two components toward the continuum peak position of MM1. For COMs with weak lines, it is difficult to determine how many velocity components they trace due to low signal-to-noise ratios. Therefore, we adopted two velocity components for most COMs but a single velocity component for most simple species, CH₃CCH, CH₂DOH, and COMs with weak lines. The emission sizes were derived as described in Sect. 3.4.2. Rotational temperatures and column densities were initially derived using population diagrams (see Sect. 3.4.3) and subsequently refined through visual inspection.

We modeled each species separately and then added up the contributions of all detected species to create the final synthetic spectrum. This approach works well for optically thin lines that overlap in frequency or for lines (optically thick or thin) that come from different regions within the beam. However, it is not accurate for optically thick lines that overlap in frequency and originate from the same region or along the same line of sight. In such cases, the synthetic spectrum can overestimate the actual

line flux, particularly for CH₃OH spectral lines with low upper-level energies in this work. The emission of vibrationally excited states of a given molecule were modeled independently of the vibrational ground state and the emission of isotopologs of a given molecule were also modeled separately.

3.4.2. Sizes of the COM emission

To determine the emission size of each spectral line, we fit its integrated intensity map with a 2D Gaussian function. We selected only the spectral lines that have a peak temperature above 3σ and are not contaminated by other species. Figures 7a, 7b, and 7c present examples of CO ($E_u = 17$ K), CH₃CCH ($E_u = 75$ K), and CH₃OH ($E_u = 165$ K), respectively, representing three spatial scale groups across all molecules. The fitted emission sizes of CO and CH₃CCH are larger than the beam, indicating well-resolved emission, while CH₃OH shows an emission size only slightly exceeding the beam, suggesting marginally resolved emission. From the fitted sizes, we calculated the deconvolved size of the emission region for all molecules. For CH₃OH $v_t = 0$, the median deconvolved size is about $0.3''$, but the size decreases with increasing upper level energy, from $\sim 0.6''$ at low energy to $\sim 0.1''$ at $E_u \sim 800$ K (top panels of Fig. C.1). However, the deconvolved sizes also depend on the line opacity, as illustrated in the bottom panels of Fig. C.1. All mean sizes larger than $0.4''$ in Fig. C.1c correspond to lines with an opacity higher than 1 (see Fig. C.1f). The Gaussian fitting may be biased by the line saturation, suggesting that the large sizes of the low-energy transitions are most likely overestimated.

Figure C.2 shows the deconvolved size for all detected molecules. In this work, we adopt a fixed deconvolved size of $0.3''$ for all COMs (except CH₃CCH and CH₃CHO) in our Weeds modeling, as most COMs are not fully resolved and exhibit similar morphology. CH₃CCH and CH₃CHO trace more extended structures, as was reported in other sources (e.g., Fayolle et al. 2015; Busch et al. 2022), and we used their median deconvolved sizes for the Weeds modeling ($1.9''$ and $0.8''$, respectively).

3.4.3. Population diagrams

Population diagrams are constructed to determine the rotational temperature, T_{rot} , and total column density, N , for each molecule under the LTE assumption. We used the following equation

⁶ <https://cdms.astro.uni-koeln.de/>

⁷ <https://spec.jpl.nasa.gov/>

⁸ <https://lsd.univ-lille.fr/>

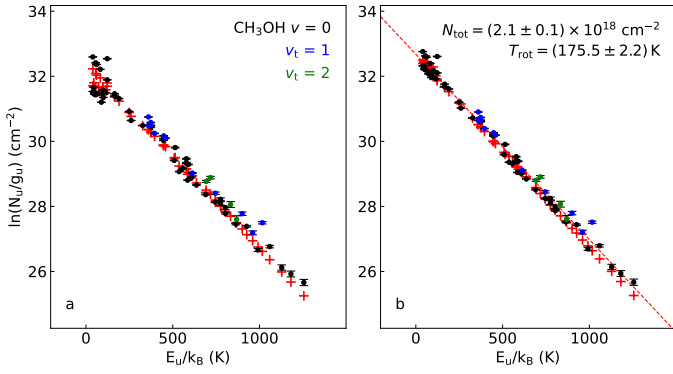


Fig. 8. Population diagram for CH₃OH towards the continuum peak position of G135.27 MM1. Panel a shows the observed data points (see the legend in the upper-right corner) and synthetic data points (in red). In panel b, the opacity correction was applied to both the observed and synthetic populations and the observed populations were corrected for contamination by other molecules. The red line represents a linear fit to the observed data points in linear-logarithmic space. The fit results are indicated in panel b.

(Goldsmith & Langer 1999):

$$\ln \frac{N_u}{g_u} = \ln \frac{8\pi k_B \nu^2 W}{hc^3 A_{ul} b g_u} = \ln \frac{N_{\text{tot}}}{Q} - \frac{E_u}{k_B T_{\text{rot}}}, \quad (6)$$

where N_u is the column density of the upper level and $b = \frac{\theta_s^2}{\theta_s^2 + H P B W^2}$ is the beam filling factor, with θ_s as the emission size.

In our population diagram analysis, we selected transitions that were minimally contaminated by lines from other species and had a peak signal-to-noise ratio greater than 3σ . In cases where a transition was partially contaminated, the contribution from the contaminating species was subtracted from the measured integrated intensities, based on the full LTE model. Each population diagram was also corrected for optical depth following the method described in Goldsmith & Langer (1999). The opacities are obtained from our best-fit Weeds model. However, the optical depth correction may not be sufficient to mitigate the effects for transitions with extremely high opacities, because gradients of excitation temperature along the line of sight are not accounted for in the population diagram method. Therefore, transitions with opacity higher than 2 were excluded from the population diagram analysis. Figure 8 illustrates the population diagram for CH₃OH, showing both the observed and modeled data points. In the right panel, deviations between the observed and synthesized data points are generally small after applying the two corrections, indicating that the chosen input parameters for the Weeds modeling are reliable and the LTE assumption is valid. The population diagrams of other molecules can be found in Fig. H.1. Derived column densities and rotational temperatures from the population diagrams are summarized in Table B.1.

This approach is applicable to all detected COMs except CH₃COOH, C₂H₅CN, and *a*-(CH₂OH)₂, as the transitions of these three molecules are weak, with only a few exhibiting a signal-to-noise ratio of approximately three, and confined to a narrow range of upper level energies. This limitation makes it challenging to derive reliable rotational temperatures and column densities using the population diagram method for these three molecules. Here, for CH₃COOH and *a*-(CH₂OH)₂, we adopted the same temperature as derived for CH₃OH (175 K),

while a temperature of 190 K from CH₃CN was used for C₂H₅CN. The population diagram method was also applied to simple molecules that exhibit multiple transitions spanning a wide range of upper level energies, including HDO, OCS, SO₂ and ³⁴SO₂, H₂CS, H₂CO, HNCO, HC₃N, CH₂CO, and *t*-HCOOH. For other simple molecules for which the population diagram method could not be used, we adopted the rotational temperature of CH₃CCH because CH₃CCH and these simple molecules trace similarly extended structures. For several isotopologs, such as H₂¹³CO, we assumed the same temperature as derived for their main isotopolog.

3.4.4. Weeds modeling result

Based on our Weeds modeling, we systematically identified the molecular emission lines detected toward the continuum peak position of MM1. Figure D.1 presents the observed spectrum in black, zoomed-in along the y axis in order to highlight the weak lines, overlaid with the Weeds synthetic spectrum of all identified molecules in dark red. The unzoomed spectrum is shown in Fig. D.2. In total, we identified 485 spectral lines with signal-to-noise ratios exceeding 3, from 28 molecules, several of their less abundant ¹³C, ¹⁸O, ¹⁷O, ³⁴S, and ³³S isotopologs, as well as one deuterated COM (CH₂DOH). The detected molecules span a broad range of chemical complexity, from simple diatomic species to COMs (12 in total) containing up to ten atoms. Additionally, three molecules (¹³C³⁴S, OC³³S, and CH₂(OH)CHO) are considered tentatively detected due to their low signal-to-noise ratios or partial line blending. The list of detected molecules is presented in Table 2. The input parameters adopted for the Weeds model of each species as well as the number of detected lines are detailed in Table B.1. Numerous lines with a peak signal-to-noise ratio higher than 3 remain unidentified in the spectrum of MM1 (see Table J.1). We consider it likely that most of these unidentified lines correspond to rotational transitions in vibrationally excited states of molecules whose ground state and some lower vibrational states are included in our model, but for which spectroscopic predictions of higher states are still missing in the spectroscopic databases.

We applied the same Weeds modeling procedure to identify the spectral lines detected toward the continuum peak positions of MM2 and MM3. Toward MM2, only eight species were detected, while 20 species were identified toward MM3. All identified molecules are simple species, except for CH₃OH detected toward both MM2 and MM3 and CH₃CCH toward MM3 (see Tables I.1 and I.2). The input parameters of the Weeds models of MM2 and MM3 are detailed in Tables B.2 and B.3, respectively. Although CH₃OH is detected toward MM2 and MM3, its emission does not peak at their continuum peak positions, as can be seen in Fig. 7d. The emission of low-energy transitions of methanol traces extended structures across the cores, suggesting that it arises from larger scales rather than being directly associated with MM2 and MM3 themselves. In contrast, the CH₃CCH emission has a local peak close to the continuum position of MM3 (see Fig. 7b), indicating that it traces the dense core material directly associated with MM3.

We note that uncertainties in the derived N_{tot} arise from the adopted parameters in the Weeds modeling. They were estimated by varying one parameter at a time while keeping the other four input parameters fixed. Under this assumption, uncertainties are below 10% for both N_{tot} and the line width, and about 20% for T_{rot} . The emission sizes of individual molecules range from 0.2'' to 0.33'', corresponding to an uncertainty of about 40% for each species (see Fig. C.2). Adopting a fixed value of 0.3'' may introduce an uncertainty of up to a factor of two

Table 2. Molecules identified toward G135.27 MM1.

2 atoms			3 atoms		4 atoms	5 atoms	6 atoms	7 atoms	8 atoms	9 atoms	10 atoms
CN	SO	CS	H ₂ S	OC ³⁴ S	H ₂ CS	CH ₂ CO	CH ₃ OH	CH ₃ CHO	CH ₃ COOH	CH ₃ OCH ₃	CH ₃ C(O)CH ₃
CO	³⁴ SO	¹³ CS	HDO	OC ³³ S ^(a)	H ₂ CO	HC ₃ N	¹³ CH ₃ OH	CH ₃ CCH	CH ₃ OCHO	C ₂ H ₅ OH	a-(CH ₂ OH) ₂
¹³ CO	³³ SO	C ³⁴ S	DCN	SO ₂	H ₂ ¹³ CO	t-HCOOH	CH ₂ DOH		CH ₂ (OH)CHO ^(a)	C ₂ H ₅ CN	
C ¹⁸ O	S ¹⁸ O	C ³³ S	OCS	³⁴ SO ₂	HNCO		CH ₃ CN				
C ¹⁷ O	SiO	¹³ C ³⁴ S ^(a)	O ¹³ CS	³³ SO ₂			NH ₂ CHO				

Notes. ^(a)Tentative detection.

in N_{tot} , which dominates the total error budget. Since both the COMs and CH₃OH were analyzed using the same method and beam size and exhibit similar spatial distributions, systematic uncertainties related to the line width and flux calibration should cancel out in the abundance ratios. For species with comparable emission sizes, beam-dilution effects are also reduced, whereas smaller sizes may still introduce uncertainties of up to a factor of two. Consequently, the column density ratios are less affected by systematic uncertainties than the column densities themselves.

3.5. Comparison to galactic hot cores

Figure 9 presents a comparison of the abundances of COMs relative to CH₃OH between G135.27 MM1 and other galactic hot cores. The reference species selected for this comparison is CH₃OH, as it is considered a parent molecule in the synthesis of more COMs (e.g., Herbst & van Dishoeck 2009). We have selected two hot cores the chemical composition of which has been studied in detail in the past: G31.41+0.31 (hereafter G31.41) and WB89-789 SMM1 (hereafter WB89-789). G31.41 is located in the inner Galaxy at a Galactocentric distance of 5.0 kpc, as determined via trigonometric parallax measurements from the BeSSeL Survey (Immer et al. 2019). Its chemical composition was derived in the frame of the GUAPOS project (Mininni et al. 2020, 2023; Colzi et al. 2021; García de la Concepción et al. 2022; López-Gallifa et al. 2024). We adopted a CH₃OH column density of $1.0 \times 10^{19} \text{ cm}^{-2}$ for G31.41, derived from the best-fit model of high-energy ($\sim 300\text{--}800 \text{ K}$) CH₃OH $v_t = 1$ transitions (see Table 2 in Mininni et al. 2023). The chemical composition of WB89-789, the first hot core identified in the outer Galaxy, was published in Shimonishi et al. (2021). We adopted a Galactocentric distance of 17.1 kpc, based on the BeSSeL Distance Calculator V2⁹, rather than the 19.0 kpc used in Shimonishi et al. (2021), which was derived from optical spectroscopy of a K3 III star (Brand & Wouterloot 2007). Based on the Galactic gradient reported by Yan et al. (2023), we assumed a ¹²C/¹³C isotopic ratio of 100 at 17.1 kpc to estimate the CH₃OH column density from that of its isotopolog ¹³CH₃OH, rather than the value of 150 at 19 kpc adopted by Shimonishi et al. (2021). We obtained a CH₃OH column density of $2.9 \times 10^{17} \text{ cm}^{-2}$ for WB89-789.

Overall, the COM abundances (relative to methanol) in G135.27 MM1 show a good correlation with those measured in G31.41 and WB89-789, although a noticeable scatter is present. To quantify this correlation, we calculated the Pearson correlation coefficients between the relative abundances of COMs in G135.27 MM1 and those in the two sources, obtaining $r = 0.80$ and $r = 0.75$ for G31.41 and WB89-789, respectively, both with p -values less than 0.01. Considering that the typical uncertainty

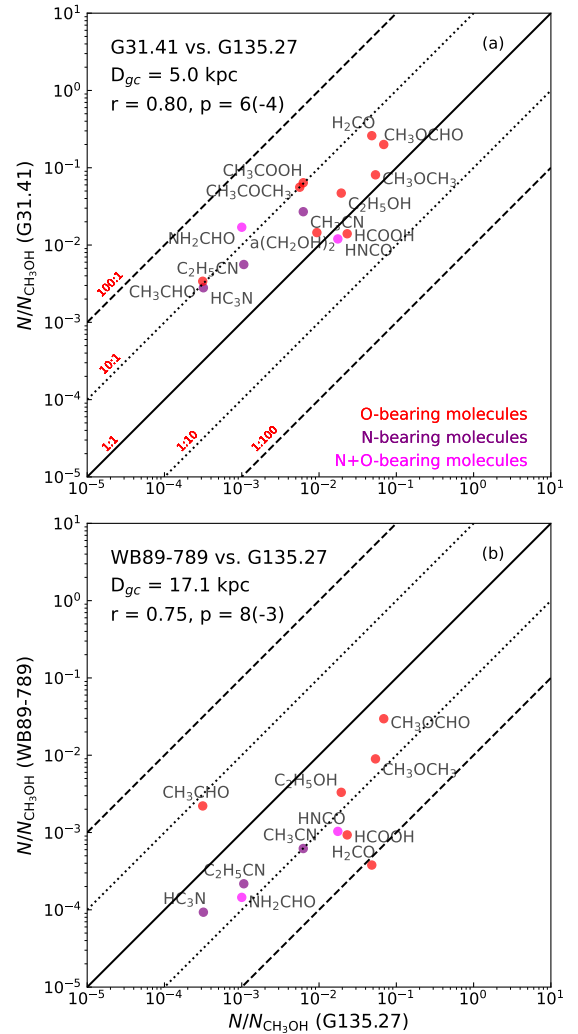


Fig. 9. Correlation plots of column densities normalized to methanol in G135.27 MM1 compared to G31.41 (panel a) and WB89-789 (panel b). Oxygen-, nitrogen-, and nitrogen+oxygen-bearing molecules are marked in red, purple, and magenta, respectively. Dashed and dotted lines indicate deviations by factors of 100 and 10, respectively. The Pearson correlation coefficient (r) and P -value for each comparison are shown in the top-left corner of each panel.

in the abundance ratios is on the order of 20–30% (corresponding to ~ 0.1 dex in logarithmic scale, and at most a factor of two in some cases), the correlations remain statistically significant within the uncertainties. Moreover, our analysis reveals no systematic difference between N-bearing and O-bearing COMs among these three hot cores.

⁹ <https://cloud.mpi-fr-bonn.mpg.de/index.php/s/LCdPgRJmMfj6XJB?opendetails=>

The abundances relative to methanol of most COMs in G135.27 MM1 are lower than those in G31.41 but higher than those in WB89-789, suggesting that G135.27 MM1 occupies an intermediate position in terms of chemical richness. Given that G135.27 MM1 is located between G31.41 in the inner Galaxy and WB89-789 in the outer Galaxy, its metallicity is likely intermediate as well. This trend may reflect the influence of metallicity on the formation and abundance of COMs.

One exception to this trend is CH_3CHO , which displays a different pattern: its abundance relative to methanol is similar in G31.41 and WB89-789, both higher than in G135.27 MM1. This is likely because CH_3CHO forms significantly ($\sim 40\%$) via gas-phase reactions ($\text{C}_2\text{H}_5 + \text{O} \rightarrow \text{CH}_3\text{CHO} + \text{H}$) during the warm-up stage (Garrod et al. 2022). This gas-phase formation pathway makes the abundance of CH_3CHO less directly dependent on the initial ice composition and metallicity, and more sensitive to the evolving physical conditions and chemical kinetics of the protostellar environment.

3.6. Comparison to extragalactic hot cores

Apart from those discovered in the Milky Way, hot cores have also been identified in nearby galaxies, namely the LMC and SMC. The metallicity of the LMC is reported to be lower than the solar value by a factor of two to three (e.g., Choudhury et al. 2016, 2021). Shimonishi et al. (2021) reported the first extragalactic hot core in the LMC, ST11, identified through the detection of a hot SO_2 component (190 K), but no COMs were detected, not even CH_3OH . Subsequent studies have revealed additional hot cores in the LMC, including N113-A1 and N113-B3 (Sewilo et al. 2018), ST16 (Shimonishi et al. 2020), N105-2A and N105-2B (Sewilo et al. 2022), and N132-14A (Hamedani Golshan et al. 2024). In the SMC, which exhibits an even lower metallicity than the LMC, with values estimated to be 4 to 10 times lower than the solar metallicity (e.g., Choudhury et al. 2018), Shimonishi et al. (2023) reported the detection of two hot cores, S07 and S09. Similar to ST11 in the LMC, the two hot cores in the SMC are identified through the presence of hot SO_2 emission (171 K for S07 and 146 K for S09). In these sources, however, CH_3OH primarily traces colder gas at temperatures of about 40 K, distributed over extended regions of size 0.2–0.3 pc, rather than the compact hot core material.

Figure 10 presents a comparison of the abundances of COMs (CH_3CN , CH_3OCH_3 , and CH_3OCHO) relative to CH_3OH between G135.27 MM1, WB89-789, and the LMC hot cores where hot methanol has been detected. Overall, the abundances of CH_3CN , CH_3OCH_3 , and CH_3OCHO relative to CH_3OH show no clear trends among hot cores in these low-metallicity environments, with the exception of WB89-789. The values in G135.27 MM1 are comparable to those observed in the LMC hot cores. This similarity is reasonable given that the expected metallicity of G135.27 ($\sim 0.5 Z_\odot$) is close to that of the LMC ($\sim 0.4 Z_\odot$, Choudhury et al. 2021). For WB89-789, the abundances of all three COMs relative to CH_3OH are markedly lower than those observed in other sources. The chemical composition of WB89-789 was derived by Shimonishi et al. (2021) assuming a beam fill factor of 1. However, the emission of the COMs identified by Shimonishi et al. (2021) is barely resolved, with fitted sizes (not deconvolved) that are close to the beam size. This suggests that much smaller sizes should be assumed to model the COM emission and derive their proper column density in WB89-789. Reducing the emission size results in higher column densities, but also higher opacities, which may have an impact on the balance between optically thin and thick lines specific to each

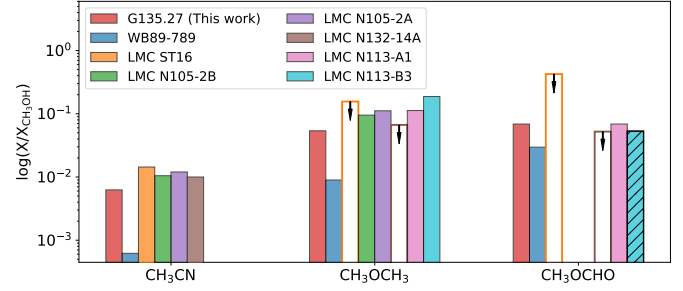


Fig. 10. Comparison of the abundances of COMs relative to CH_3OH between G135.27 MM1, WB89-789, and hot cores in the LMC. The LMC hot cores are ST16 (Shimonishi et al. 2020), N113-A1 and N113-B3 (Sewilo et al. 2018), N105-2A and N105-2B (Sewilo et al. 2022), and N132-14A (Hamedani Golshan et al. 2024). Unfilled bars with an arrow indicate upper limits, while hatched bars denote tentative detections.

molecule. A reanalysis of the ALMA data using smaller sizes is needed before we can draw any firm conclusion concerning WB89-789.

Naturally, we needed to consider whether the difference in spatial scales could lead to variations in the derived abundances. The COM emission extends over ~ 5000 au in G135.27 MM1, and ~ 0.1 pc in the LMC hot cores. Shimonishi et al. (2021) determined COM abundances in WB89-789 at both 5000 au and 0.1 pc scales and found that, while absolute abundances vary somewhat with scale, the ratios of COMs to CH_3OH for these three species remain similar within a factor of two. This suggests that the difference in spatial scales between the observations of G135.27 and the LMC hot cores does not significantly affect the comparison shown in Fig. 10, and that the relative abundances can be meaningfully compared despite the different scales. Nevertheless, variations in environmental parameters, such as cosmic-ray ionization rates and UV radiation fields, may also influence the chemistry and contribute to differences between the outer Galaxy and the LMC.

4. Astrochemical modeling

To interpret the COM abundances detected toward the hot core in G135.27 MM1, we simulate hot-core chemistry under both local- and outer-Galaxy conditions using the astrochemical code MAGICKAL (Model for Astrophysical Gas and Ice Chemical Kinetics and Layering, Garrod et al. 2022). MAGICKAL is a three-phase (gas phase, grain/ice-surface and bulk-ice mantle) chemical kinetics code with an extensive chemical network intended for the study of COM chemistry. Along with standard gas-phase reactions and diffusive grain-surface chemistry, the model includes the nondiffusive surface and ice-mantle chemical treatment introduced by Jin & Garrod (2020) and further developed by Garrod et al. (2022). Here, MAGICKAL is applied in single-point mode, using a two-stage physical treatment that has been used in various past studies (e.g., Garrod et al. 2008; Belloche et al. 2019; Garrod et al. 2022; Shope et al. 2024; Belloche et al. 2025). Stage 1 consists of an isothermal collapse ($T_{\text{gas}} = 10$ K) from initial gas density $n_{\text{H}} = 3 \times 10^3 \text{ cm}^{-3}$, reaching a final value of $2 \times 10^8 \text{ cm}^{-3}$ in ~ 1 Myr. During the collapse, the visual extinction evolves in proportion with $n_{\text{H}}^{2/3}$, starting at a value $A_{\text{V,init}} = 2$ or 3 mag and taking a maximum value of 500 mag. While the gas temperature is held steady, the dust temperature evolves with visual extinction according to the expression of Hocuk et al. (2017), resulting in an initial value

$T_{\text{dust}} \approx 12.7$ K or 11.8 K, and final value of 5.5 K. In stage 2, the gas density is held steady, while the gas and dust temperatures rise in concert to a final value of 400 K. As in past work, three timescales for the warm up stage are considered in this model: *fast* (5×10^4 yr), *medium* (2×10^5 yr) and *slow* (1×10^6 yr). The same canonical interstellar cosmic-ray ionization rate of $\zeta_0 = 1.3 \times 10^{-17} \text{ s}^{-1}$ is used in all models, throughout stages 1 and 2.

The adoption of two $A_{V,\text{init}}$ values in stage 1 is intended to allow for a modest exploration of the effects of external UV while the core is still relatively diffuse; as noted by Garrod et al. (2022), certain COMs are more strongly affected by UV-driven processing of the early ices formed on grains, making the choice of initial visual extinction potentially important for those species. Furthermore, it is unclear how the initial visual extinction profile in a nascent hot core may vary between local- and outer-Galaxy conditions; however, the higher gas-to-dust ratio in the outer Galaxy (see below) would suggest lower extinction values than within the solar circle, all else being equal.

Using the same physical setups in each, the models are run under chemical conditions appropriate to the local galactocentric radius (8.2 kpc) and that of G135.27 (13 kpc). In practice, this means the adoption of different initial elemental abundances of C, N, and O, and the gas-to-dust ratio by mass. Local-Galaxy initial abundances are $n(\text{C}^+) = 1.40 \times 10^{-4} n_{\text{H}}$, $n(\text{N}) = 2.14 \times 10^{-5} n_{\text{H}}$, and $n(\text{O}) = 3.20 \times 10^{-4} n_{\text{H}}$ (Jin et al. 2022). These are scaled according to Luck & Lambert (2011), giving values at 13 kpc of $n(\text{C}^+) = 5.47 \times 10^{-5} n_{\text{H}}$, $n(\text{N}) = 1.20 \times 10^{-5} n_{\text{H}}$, and $n(\text{O}) = 1.66 \times 10^{-4} n_{\text{H}}$. These correspond to the initial forms and values used for these elements at the beginning of stage 1 in the chemical models in the low-metallicity case. We note that C, N, and O abundances are also available from the studies by Arellano-Córdova et al. (2020); Méndez-Delgado et al. (2022), who derive total elemental abundances in H II regions from ionic abundances, using ionization correction factors. Luck & Lambert (2011), on the other hand, determine elemental abundances from photospheric absorption lines of Cepheids, providing direct total abundances without ionization corrections. This, combined with the larger sample size provided by Luck and Lambert (~400 versus fewer than 50 in the other studies), made that study a more favorable choice for our modeling purposes.

Under local conditions, a gas-to-dust ratio by mass of 100 is adopted, which is scaled to 303 at 13 kpc for G135.27 (see Sect. 3.2). Making the usual assumption of a representative grain radius of 0.1 μm , this leads to a gas-to-dust ratio by number of 7.57×10^{11} locally, or 2.88×10^{12} at 13 kpc.

The chemical network used here is an updated version of that used by Garrod et al. (2022), and includes a more extensive gas-phase radical chemistry, as well as the adoption of more accurate quantum tunneling treatments for certain surface reactions. The network is described in more detail by Bonfand et al. (in prep.).

4.1. Chemical model predictions for low- versus normal-metallicity

Figure 11 illustrates the ratio R of molecular abundances (relative to methanol) of low- and normal-metallicity hot-core models for a set of 96 molecules. The results are shown for the initial visual extinction values and warm-up timescales introduced in the previous paragraphs. This comparison reveals how molecular abundances respond to changes in metallicity across the Milky Way. For each panel, molecules that have a significantly lower abundance relative to methanol ($R < 1/3$) in the low-metallicity models whatever the warm-up timescale are highlighted in red

(e.g., $\text{C}_2\text{H}_5\text{CN}$, $\text{C}_2\text{H}_5\text{OH}$, $\text{CH}_3\text{C}(\text{O})\text{CH}_3$, CH_3COOH), indicating strong sensitivity to metallicity. In contrast, molecules with R values within a factor of two of unity across all timescales ($1/2 < R < 2$) are shown in blue, suggesting that their formation pathways are not more sensitive to metallicity variations than methanol (e.g., CH_3OCHO , NH_2CHO). Molecules with abundances relative to methanol significantly higher ($R > 3$) in the low-metallicity models whatever the warm-up timescale are highlighted in green (e.g., $\alpha\text{-(CH}_2\text{OH)}_2$, $\text{CH}_2(\text{OH})\text{CHO}$). Since methanol itself shows a moderate decrease of its abundance relative to H_2 (less than a factor of three) in the low-metallicity models, $R > 3$ indicates that the abundances of these molecules relative to H_2 are even enhanced at low metallicity. The other molecules that do not fit into these three categories are displayed in black.

Most of the simpler species shown in Fig. 11, such as water, CO, CO_2 , and NH_3 (which are major ice components) are fairly stable in their peak warm-up stage gas-phase abundances with respect to methanol, at low- Z versus normal- Z values. These species also show minimal variation as a function of warm-up timescale or with different initial visual extinctions. Since each of these species is mostly preserved in the solid phase until temperatures are high enough for them to desorb substantially, it is instructive to consider their abundances in the ices at the end of the cold collapse stage. In the models, much of the production of COMs occurs during this earlier period.

In absolute terms, water ice, for example, falls by a factor ~1.4 (1.6) for $A_{V,\text{init}} = 2$ mag ($A_{V,\text{init}} = 3$ mag) in the 13 kpc model versus the local-Galaxy model, which is a little less than the drop-off in the total oxygen budget of 1.93 $\times$. This means that some other major O-bearing species must have fallen by a factor greater than this value. Indeed, CO ice at the end of the cold collapse stage is a factor ~4.1 (3.5) lower in the 13 kpc model, while CO_2 ice falls by a factor ~4.6 (4.0). These decreases are greater than the individual drop-off factors in either the oxygen (1.93 $\times$) or carbon budget (2.559 $\times$). For ammonia, NH_3 , the main solid-phase carrier of nitrogen in the models, the absolute ice abundance falls by a factor ~1.8 (1.9) for $A_{V,\text{init}} = 2$ mag ($A_{V,\text{init}} = 3$ mag) in the 13 kpc model versus the local-Galaxy model, which is much more consistent with the drop in nitrogen budget of 1.778 $\times$.

On a simple analysis, one might expect that the decline in CO abundance would follow the reduction in budget of the least abundant of either the elemental C or O; this is because CO is efficiently formed in the gas phase and is quite chemically stable, allowing it to incorporate a large fraction of the overall budget of each element (any grain-surface conversion to CO_2 or other species notwithstanding). Regarding oxygen, the slight overabundance of water ice may be explained by the lower CO and CO_2 abundances themselves; the fact that there is a reduction in the carbon budget reduces the possible abundances of these species, allowing water to incorporate a larger proportion of the reduced overall oxygen budget. The fact that the final CO (and CO_2) ice abundance is depressed beyond the factor associated with the lower carbon budget indicates that there is another factor at work.

Methane is typically the other major carbon-carrying species in the ice mantles (e.g., McClure et al. 2023), and its absolute abundance falls by a factor 7.5 (9.5) in the 13 kpc models versus the local-Galaxy models. It is thus disproportionately reduced (by ~2.9–3.7 $\times$) beyond the expected reduction associated with lower metallicity. This much lower solid-phase CH_4 abundance, and the associated lower abundances of other hydrocarbons, including radicals (such as CH_3) and larger, stable species (such

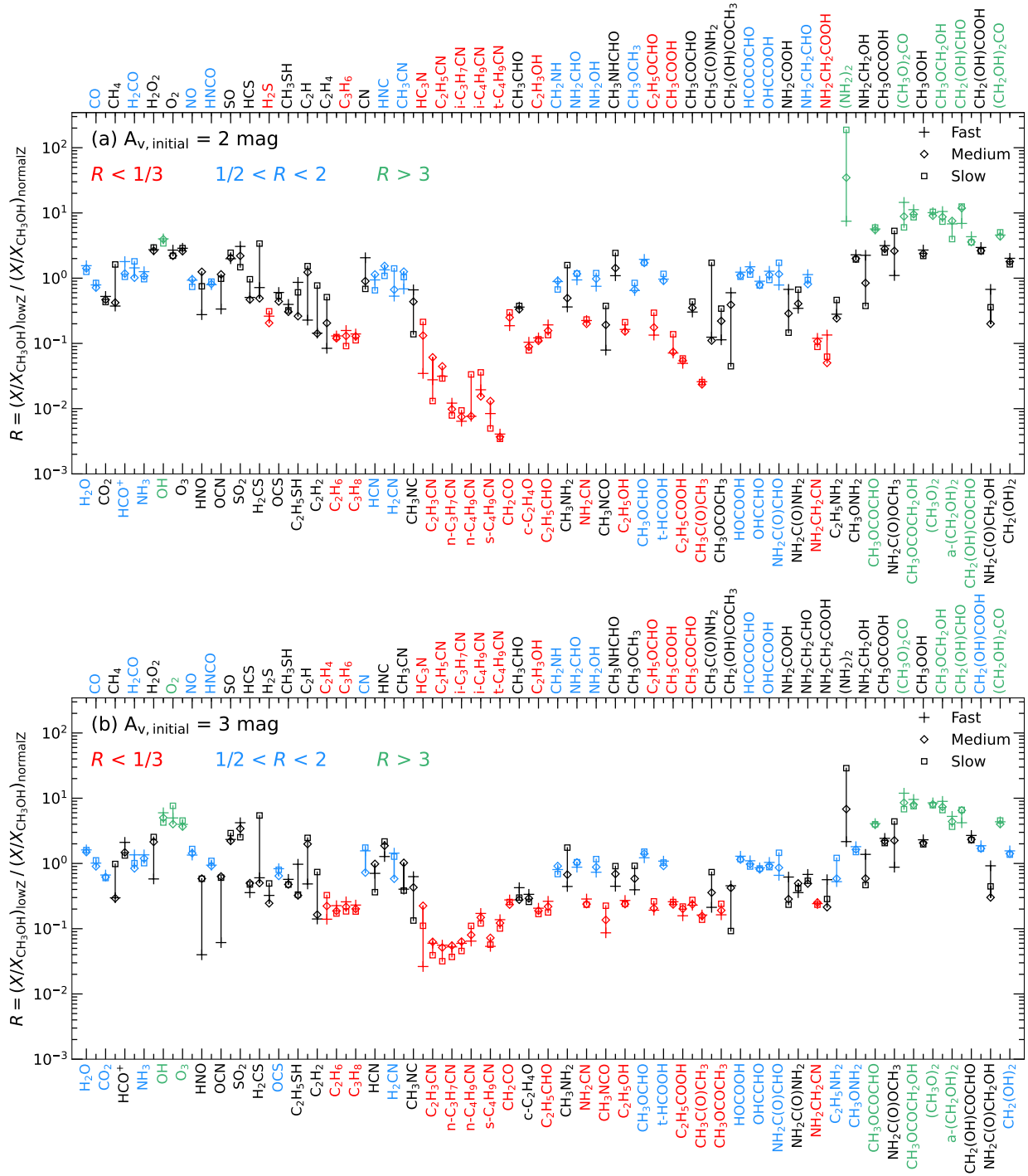


Fig. 11. Ratio R of molecular abundances (relative to methanol) between low- and normal-metallicity hot-core models for a sample of 96 molecules, shown for three warm-up timescales: 5×10^4 yr (fast), 2×10^5 yr (medium), and 1×10^6 yr (slow). Panels a and b show models with an initial visual extinction of 2 and 3 mag, respectively. Molecules with R lower than one-third at all three timescales are highlighted in red, while those with R higher than three at all three timescales are highlighted in green, and molecules with R within a factor of two of unity ($1/2 < R < 2$) at all three timescales are highlighted in blue. The other molecules are displayed in black.

as C_2H_6) leads to lower abundances of certain more complex organics that happen to be directly dependent on the addition of methyl or larger radical groups to other radicals, or on the hydrogenation of unsaturated hydrocarbon species (originating in the gas phase – see below) whose abundances are also depressed. This effect manifests in the gas-phase behaviors shown in Fig. 11, wherein ethane (C_2H_6), propane (C_3H_8), and several nitriles

(e.g., ethyl cyanide, $\text{C}_2\text{H}_5\text{CN}$) show declines of around an order of magnitude or more in their ratios with methanol, between the low- Z and normal- Z models.

Methanol itself, with which all other species are compared in Fig. 11, is less affected in its collapse-stage solid-phase abundance than CO, falling by a factor ~ 2.3 (2.5). The peak gas-phase methanol abundances during the warm-up stage are similarly

affected, although there is a degree of variation between values from the different combinations of warm-up timescale and $A_{V,\text{init}}$ that does not exhibit a clear pattern. The fact that the production of both solid- and gas-phase methanol is closely related to the conversion of grain-surface CO means that the latter appears very stable in Fig. 11.

In fact, while the final CO ice abundance in the collapse stage is only lower by a factor of a few in the low-Z case, during much of this stage the surface coverage of CO on the grains is five to ten times lower (only eventually rising further when gas densities become much greater). This lower coverage is caused by the $\sim 3\times$ lower rate of adsorption onto the dust grains associated with the larger gas-to-dust ratio. This has the effect that there is less reactive material on the surface at any one time in the low-Z model with which H may react; most of the other material on the surface of the grain/ice is water or CO₂, which are essentially inert to reactions with H atoms. Although the adsorption of gas-phase atomic H is also slowed down on a per-atom basis in the low-Z case, the gas-phase abundance of H undergoes a commensurate rise, so that its total rate of adsorption is largely unaffected during the period when freeze-out is strongest. This drives up the surface population of atomic hydrogen, such that the CO that is yet present on the surface tends to be driven to a further degree of hydrogenation. This favors the production of species such as glycolaldehyde, CH₂(OH)CHO, and ethylene glycol, a-(CH₂OH)₂, along with other species whose production is driven by nondiffusive reactions between CO-product radicals such as HCO, CH₃O, and CH₂OH. Such species, clustered toward the right of Fig. 11, show substantial increases in their peak gas-phase abundances relative to methanol during the warm-up stage. The abundance gains across these various species add up so as to correspond to the excessive fall in CO and CO₂ ice abundances.

Curiously, the peak gas-phase abundance of methyl formate, CH₃OCHO, is not so strongly enhanced in the low-Z case. This likely reflects a relative scarcity of H-addition/abstraction reactions in the network for that species, as compared with glycolaldehyde. Our network includes reactions and rates taken from [Álvarez-Barcia et al. \(2018\)](#), who calculated mechanisms for the addition or abstraction of atomic H with species including methyl formate and glycolaldehyde, including the interconversion of the latter with ethylene glycol, via one of two intermediate radicals. However, their calculations did not go as far in the consideration of the methyl formate conversion, and indeed our network does not include such interconversion steps. Thus, while CH₃OCH₂OH is formed on the grains in stage 1, similarly to a-(CH₂OH)₂, it cannot be interconverted to CH₃OCHO. Glycolaldehyde abundance in stage 2 benefits from the latter back conversion from ethylene glycol, serving to increase its abundance substantially. Therefore, the apparent lack of enhancement in methyl formate in the low-Z models may be an artifact of the network under the extreme conditions sampled in these models.

While the grain surfaces during the collapse stage host the production of various molecules, including methane (via hydrogenation of atomic C) and various COMs, the low-Z conditions also produce important changes in the behavior of the gas-phase chemistry that is relevant to those species. Crucially, the much lower rate of freeze-out of gas-phase species onto the grains in the low-Z case provides more time for atomic carbon to be locked up into, primarily, gas-phase CO (although, as described above, it is also more rapidly converted into larger species once it reaches the grains). As a result, less atomic C is available to be directly accreted onto the dust grains, resulting in the lower methane ice abundances. The production of solid-phase methane

is indeed seen more generally in models to be formed early in the evolution of the dust-grain ices, while the availability of free carbon in the gas phase is high, prior to the strong growth of CO (see [Garrod & Pauly 2011](#)).

Exacerbating the effects of slower freeze-out is the lower C:O ratio in the low-Z models, which may be just as important to the ultimate abundances of various organic molecules as the actual decline in either C or O elemental abundances would be. In the normal-Z case, the initial C:O ratio is 0.44, versus 0.33 at low-Z. In the models, this leaves much less carbon available for incorporation into pure hydrocarbon forms due to incorporation into CO.

Differences between the $A_{V,\text{init}} = 2$ mag and $A_{V,\text{init}} = 3$ mag cases are most clearly manifest in the abundances of the nitrile species, with lower values with respect to methanol seen in the $A_{V,\text{init}} = 2$ mag case. It should be noted that the choice of comparison with respect to methanol, which is less affected by lower Z due to the robustness of the CO behavior, tends to overemphasize the effects of lower metallicity on the nitriles. However, the ratios for those species are in some cases around an order of magnitude lower for $A_{V,\text{init}} = 2$ mag. This is mainly related to the higher degree of conversion of neutral atomic C to C⁺ at this extinction, leaving less neutral carbon that can be converted into neutral carbon chains, such as C₃, that contribute directly to, for example, ethyl cyanide (C₂H₅CN) on the dust grains. Again, this carbon is preferentially converted into CO in the gas.

Within each panel in Fig. 11, there exists a degree of variation for individual species as a function of warm-up timescale. This indicates that those species that do not vary strongly in this way owe much of their production to the earlier, cold stage. In the most extreme case, that of hydrazine, (NH₂)₂, cosmic-ray induced UV photolysis of solid-phase ammonia is enhanced by longer timescales, prior to desorption of the ice mantles.

4.2. Comparison of model predictions to G135.27 MM1

Figure 12 compares the measured abundances (relative to methanol) obtained for G135.27 MM1 in Sect. 3.4 to the model predictions for different warm-up timescales. Unfilled bars with arrows represent observational upper limits, while red, blue, green, and black bars correspond to species exhibiting different abundance patterns with respect to metallicity variations in the models, as illustrated in Fig. 11. With the cosmic-ray ionization rate and final gas density held fixed, the key parameters distinguishing the models are the warm-up timescale and the initial visual extinction. To quantify the overall agreement with the observations, we compute a matching parameter (the root mean square of the ratios between the model predictions and the observed abundances) for each of the six models. The resulting matching parameters are shown in Fig. K.1. Some of the detected species are not included in our matching calculations, as they are not present in the chemical network used in the model grid. CH₂(OH)CHO is excluded because it is only tentatively detected and we treat it as an upper limit, and all models appear to overestimate it. CN and CO are excluded, as they primarily trace more extended cold gas rather than the hot core environment. We also exclude the sulfur-bearing species, since the dominant form of sulfur in dense clouds is uncertain, making the chemistry poorly constrained. In total, 16 species have been included in the matching routine, comprising 11 of the 13 detected COMs (excluding CH₃CCH and CH₂(OH)CHO) and five simple molecules (H₂CO, HC₃N, HNCO, t-HCOOH, and CH₂CO). As illustrated in Fig. K.1, compared to the models with $A_{V,\text{init}} = 2$ mag, those with $A_{V,\text{init}} = 3$ mag show significantly

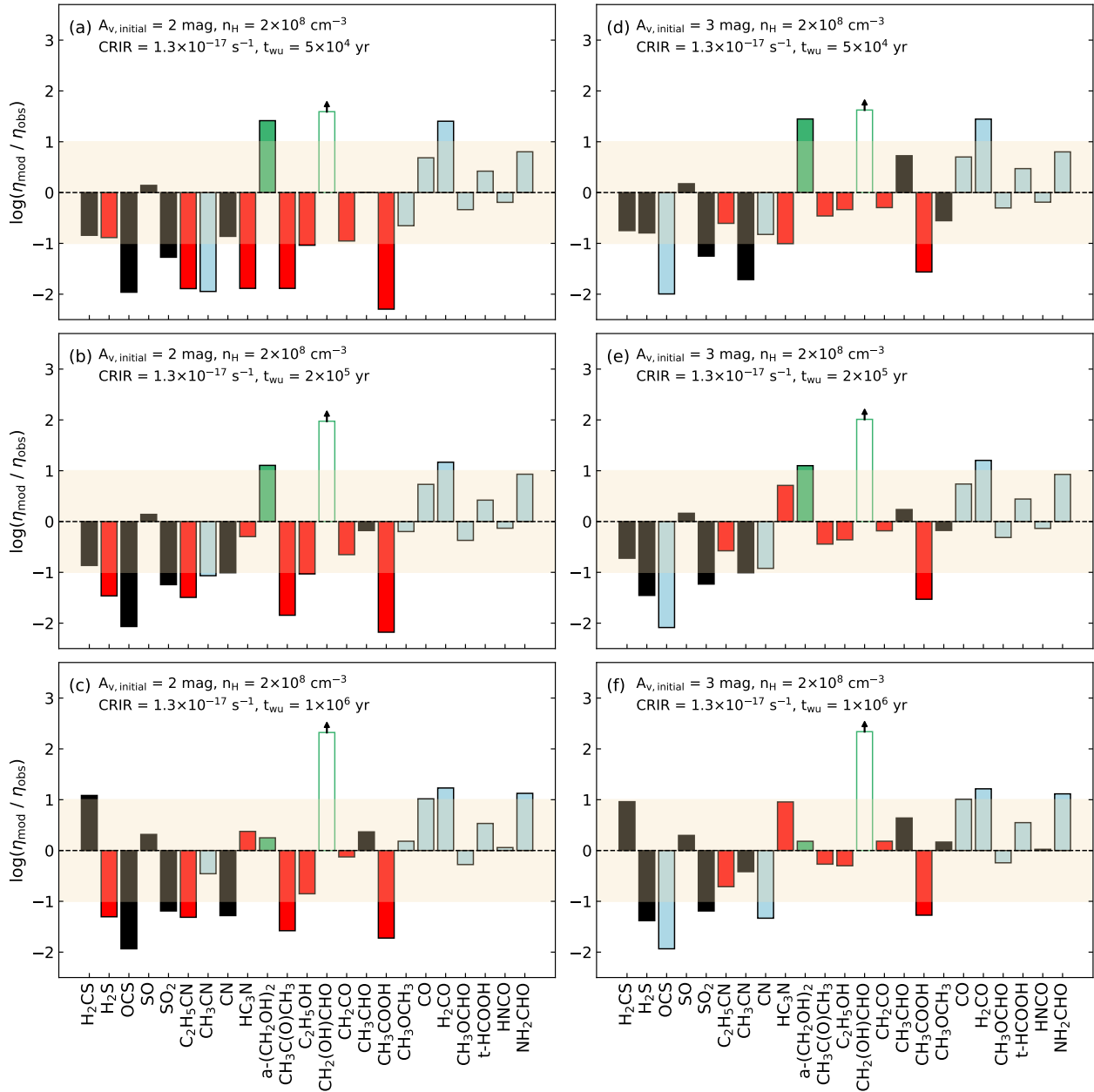


Fig. 12. Comparison of the model abundances to observational results for different initial visual extinctions, $A_{V, \text{init}} = 2$ mag (left panels) and 3 mag (right panels), at varying warm-up timescales: fast (top panels), medium (middle panels), and slow (bottom panels). Bars indicate the number of orders of magnitude by which each modeled abundance ratio exceeds (positive bars) or falls below (negative bars) the corresponding observed ratio. The shaded region denotes agreement within one order of magnitude. Unfilled bars with an arrow represent cases where the comparison is based on an observational upper limit. The color coding is the same as in Fig. 11.

better consistency with the observed abundances, whatever the warm-up timescale. The slow warm-up timescale produces the smallest dispersion compared to the other timescales, regardless of whether $A_{V, \text{init}} = 2$ mag or 3 mag is adopted. Therefore, the best-fit model for G135.27 MM1 is provided by the one adopting a slow warm-up timescale of $t_{\text{wu}} = 1 \times 10^6$ yr, and an initial visual extinction of $A_{V, \text{init}} = 3$ mag, which is shown in Fig. 12f.

In general, most common hot-core species are reasonably well reproduced by the best-fit model, i.e., are within one order of magnitude of the observed values (as marked by the shaded region). COMs that are more sensitive to metallicity than methanol and are marked as red bars ($\text{C}_2\text{H}_5\text{CN}$, $\text{C}_2\text{H}_5\text{OH}$,

$\text{CH}_3\text{C(O)CH}_3$, and CH_3COOH) are all somewhat underproduced in the model, although their abundances are mostly within one order of magnitude of the observed values. CH_3OCHO and NH_2CHO , which are both similarly sensitive to metallicity as methanol (blue bars), show opposite trends in the best-fit model: CH_3OCHO is slightly underproduced, while NH_2CHO is overproduced by about one order of magnitude. $\text{a-(CH}_2\text{OH)}_2$, whose predicted abundance is enhanced by the effect of metallicity, is consistent with the observations. Additionally, COMs shown in black bars are consistent with the model. On the other hand, simple molecules show larger discrepancies, with some exhibiting significant deviation relative to the observations. For example,

three of the five S-bearing species are underproduced by more than one order of magnitude. However, the initial S budget is uncertain, and the initial value used in the models ($8 \times 10^{-8} n_{\text{H}}$) may be a substantial underestimate.

5. Discussion

5.1. Nature of G135.27's continuum sources

We used the LTE spectral modeling tool Weeds to analyze the continuum-subtracted spectra of the three continuum sources detected toward G135.27, and derive their chemical compositions. These three dense cores, which we called MM1, MM2, and MM3, exhibit distinct chemical and physical properties that reflect their differing evolutionary stages. Toward MM1, we identified 13 COMs, 12 firmly and one tentatively ($\text{CH}_2(\text{OH})\text{CHO}$). They range from relatively simple species such as CH_3OH (six atoms) to larger molecules such as CH_3COCH_3 (ten atoms). Hundreds of spectral lines associated with these COMs were detected. The detection of such a wide variety of COMs, along with their numerous transitions, reflects a chemically rich environment in MM1 (see Sect. 3.4.4 and Fig. D.1). The derived temperatures span a range from 65 to 190 K, with a mean value of approximately 124 K and a median value of 120 K, based on a population diagram analysis and synthetic spectrum fitting with Weeds (see Table B.1). Both the mean and median rotational temperatures exceed the typical thermal desorption thresholds of COMs from dust grain surfaces (e.g., Garrod & Herbst 2006; Jørgensen et al. 2020). Additionally, the emission sizes are generally on the order of $0.3''$, corresponding to a radius of approximately 900 au at the distance of G135.27 (see Sect. 3.4.2). The combination of compact emission, high rotational temperatures, and the detection of a rich inventory of COMs indicate that MM1 is a hot core, making it the third such object discovered in the outer Galaxy after WB89-789 identified by Shimonishi et al. (2021) and Sh 2-283-1a recently identified by Ikeda et al. (2026). Additionally, MM1 drives a bipolar outflow traced by CO and ^{13}CO . The central overlap of the lobes and their extent in the plane of the sky suggests a wide opening angle along the line of sight (Fig. 2). High-velocity components reveal both a compact, recent ejection close to the protostar and an extended, slower component from an earlier phase, with nearby H_2O masers further supporting ongoing outflow activity. However, only one SiO feature is detected near MM1 (Fig. 6): it is redshifted with low velocities and elongated along the outflow direction. WB89-789, where CO was not observed, shows a symmetric SiO distribution with secondary peaks tracing shocked gas from bipolar outflows.

In contrast, the continuum sources MM2 and MM3 exhibit much simpler chemical signatures compared to MM1. Only the COM CH_3OH is detected toward both cores, with CH_3CCH additionally identified in MM3. Assuming that the COM emission scales with the continuum peak flux density yields COM line intensities in MM2 and MM3 that would have been detected in our dataset. Therefore, the non-detection of COMs in MM2 and MM3 is unlikely to be due to a lack of sensitivity, unless they harbor COM emission over a much smaller size than MM1. As demonstrated in Sect. 3.4.4, CH_3OH appears to originate from larger-scale gas rather than being directly associated with these cores. Although the CH_3CCH emission in MM3 peaks near the continuum position, indicating that it traces the dense core material directly associated with MM3, CH_3CCH commonly traces less warm and more extended gas than other COMs (e.g., Chen et al. 2025). Additionally, rotational temperatures

are below 100 K in both cores, with H_2CO yielding 41 K for MM2 and 89 K for MM3. These simpler chemical signatures and lower temperatures suggest that MM2 and MM3 are at an earlier evolutionary stage, prior to the onset of the hot core chemistry characteristic of MM1. This interpretation is further supported by the narrower line widths measured toward MM2 and MM3 (~ 1.5 and 2 km s^{-1} , respectively) compared to MM1 (4 km s^{-1}), which reflects a lower level of turbulence or dynamical activity within these cores. Moreover, the CO and ^{13}CO spectra toward MM2 exhibit redshifted absorption against the continuum, that is an inverse P Cygni profile, providing evidence of infall motions (e.g., Di Francesco et al. 2001; Mottram et al. 2013; Evans et al. 2015). The presence of such an infall feature suggests that MM2 is a dynamically young core likely undergoing gravitational collapse. On the other hand, no outflow activity is detected toward MM2 in any of the observed molecular tracers, including CO and SiO. This absence of outflow signatures indicates that MM2 may be at the prestellar stage. For MM3, the detection of EHV SiO emission reveals the presence of a powerful outflow. Such EHV SiO emission is typically associated with the early protostellar stage of star formation (see Sect. 3.3.3). These characteristics, including cold gas, limited chemical complexity, and active outflow processes, indicate that MM3 is a young, embedded protostellar object.

5.2. Origin of the velocity gradient traced by COMs

As mentioned in Sect. 3.3.2, COMs (CH_3OH and CH_3OCHO) trace a velocity gradient that is aligned with the outflow axis but opposite in direction to the large-scale CO outflow. To understand the physical processes responsible for the observed opposite velocity gradient, it is essential to consider the formation and release mechanisms of CH_3OH and CH_3OCHO . These two species are closely linked, both theoretically and observationally (Simons et al. 2020; Qin et al. 2022). CH_3OH forms on icy dust grain surfaces through successive hydrogenation of CO (e.g., Watanabe & Kouchi 2002), while CH_3OCHO is primarily produced via surface reactions between HCO and CH_3O (Simons et al. 2020). Both molecules are generally believed to form on grain surfaces at relatively low temperatures (e.g., Garrod et al. 2008; Jin & Garrod 2020), and are subsequently released into the gas phase as the grains are heated by the central protostar. In addition to thermal desorption, shock-related heating and sputtering may also play a significant role in some environments, such as the quiescent Galactic Center cloud G+0.693–0.027 (e.g., Requena-Torres et al. 2008; Zeng et al. 2018) and the shocked region located in the outflow of G31.41 (López-Gallifa et al. 2025). Grain processing caused by the passage of a shock driven by an outflow can lead to nonthermal desorption, providing an alternative pathway for injecting CH_3OH and CH_3OCHO into the gas phase (e.g., Flower et al. 2010; Busch et al. 2024).

A similar opposite velocity gradient has been observed in the low-mass protostar NGC 1333 IRAS 4A2 where CH_3OH and SiO show a velocity gradient on a scale of $\sim 150 \text{ au}$ along the outflow axis with a direction opposite to the kinematics of the outflow on larger scales. To interpret this phenomenon, De Simone et al. (2024) proposed three possible explanations: (1) rotating disk or inner envelope, (2) jet or outflow precessing or with an axis in the plane of the sky, and (3) wide-angle disk wind. They favored the last scenario.

In the case of G135.27, the C^{18}O PV diagram extracted perpendicular to the outflow axis shows a velocity gradient consistent with rotation in the protostellar envelope (Fig. 3b). In contrast, the gradients traced by CH_3OH and CH_3OCHO are

aligned with the outflow axis and therefore cannot be explained by rotation. Their orientation opposite to the large-scale CO outflow also rules out the possibility that they trace the outflow itself. Instead, these features could be explained by interaction with a wide-angle wind, where shocks driven into the envelope release COMs from dust grains and generate the observed inverse gradients, consistent with the interpretation proposed for NGC 1333 IRAS 4A2. The shock-sensitive nature of CH_3OH and CH_3OCHO further supports this scenario.

However, the spatial scales traced by the COM velocity gradient differ between G135.27 and NGC 1333 IRAS 4A2. In G135.27, the offset between blue- and redshifted COM emission indicates that the opposite velocity gradient extends over $0.1''$ (600 au), a size significantly larger than the extent of the gradient reported in NGC 1333 IRAS 4A2 (~ 150 au). This difference in scale may result from the mass difference between these two star-forming regions: NGC1333 IRAS 4A2 is a low-mass protostar with an envelope mass of $2.3 M_\odot$ and a luminosity of $4.7 L_\odot$ (Belloche et al. 2020), while G135.27 is a protostar with an envelope mass of $49 M_\odot$ (see Sect. 3.2) and a luminosity of $3.1 \times 10^4 L_\odot$ (see Sect. 2.1), which qualify it as a high-mass star-forming region. The typical disk radii in low-mass star-forming regions are smaller than 50 au (Long et al. 2019; Maury et al. 2019), while disk radii in high-mass star-forming regions can reach up to 1000 au, such as the 858 au disk around IRAS 20126+4104 (Chen et al. 2016). This comparison supports the idea that a wide-angle disk wind in G135.27 can extend to ~ 600 au, further reinforcing the interpretation that the opposite velocity gradient traced by CH_3OH and CH_3OCHO is associated with such a wind.

5.3. Metallicity-dependent behavior of classes of COMs

Figure 13 shows the fraction and number of COMs exhibiting different abundance patterns in response to metallicity variations for N-, O-, and O+N-bearing COMs, using our chemical models. These statistics build directly on the molecular-level behavior illustrated in Fig. 11. The distribution of molecules with respect to metallicity sensitivity varies markedly among classes of COMs. Across both $A_{V,\text{init}} = 2$ mag and 3 mag, N-bearing COMs exhibit the highest fraction of species whose abundances are more sensitive to metallicity than methanol ($R < 1/3$), with roughly 70% of all N-bearing COMs affected. In contrast, O-bearing COMs show a smaller fraction (about one third) of species more sensitive to metallicity than methanol. Only 8% O+N-bearing COMs are more sensitive to metallicity than methanol. For molecules with metallicity sensitivity comparable to CH_3OH ($1/2 < R < 2$), the trend is reversed: the fraction increases from N-bearing to O-bearing, and then to O+N-bearing COMs, also supporting that N-bearing COMs are the most sensitive to metallicity variations. Additionally, only O-bearing COMs show enhancement at low metallicity.

Stronger variation among the N-bearing species is associated in the models with molecules within that group whose formation has a direct dependence on radicals related to methane and other hydrocarbons. These species' abundances are diminished as the result of the more complete conversion of gas-phase carbon into CO, associated with the slower freeze-out caused by a greater gas-to-dust ratio.

The formation mechanisms of O-bearing COM species are typically more dependent on CO-related chemistry, involving radicals such as HCO and $\text{CH}_3\text{O}/\text{CH}_2\text{OH}$ that are closely tied to CO hydrogenation. These radicals are strongly enhanced, beyond

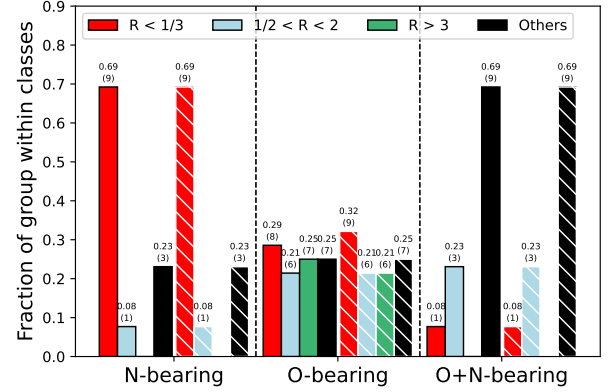


Fig. 13. Fraction and number of molecules exhibiting different abundance patterns with respect to metallicity variations for N-, O-, and O+N-bearing COMs, based on chemical modeling results. Their values are shown at the top of each bar. The color coding is the same as in Fig. 11. Solid bars and hatched bars denote the fractions for initial visual extinctions of 2 and 3 mag, respectively.

any enhancement in the abundance of CO itself, due to elevated surface hydrogenation, which allows their product COMs to maintain or exceed the ratios (versus methanol) achieved under low-Z conditions. However, there remain some O-bearing species, such as ethanol ($\text{C}_2\text{H}_5\text{OH}$), that have influential formation mechanisms not directly related to CO chemistry and are more reliant on the addition of methyl (CH_3) or larger alkyl radicals to other species. These species tend to be somewhat diminished in the same way as various N-bearing COMs.

Those molecules bearing both O and N tend to be among that group of species in the models that are more reliant on CO chemistry, so that, in spite of the lower abundance of elemental nitrogen, almost all of them are able to avoid drastic reductions in their ratios with methanol, compared with the normal-Z case that is they avoid falling into the $R < 1/3$ category.

5.4. Impact of metallicity on the COM chemical composition

We compared in Fig. 11 the outcomes of the chemical model under low and normal metallicities. In order to test if the behavior of the models with respect to metallicity is corroborated by the observations, we need to compare the chemical composition of G135.27 to the composition of a hot core in the inner Galaxy. We chose the hot core G31.41 because its detailed chemical composition has been derived (Mininni et al. 2020, 2023; Colzi et al. 2021; García de la Concepción et al. 2022; López-Gallifa et al. 2024) and it has a similar luminosity and envelope mass ($4.4 \times 10^4 L_\odot$ and $70 M_\odot$, respectively; Mininni et al. 2020) as G135.27. Figure 14a shows the observational ratios between G135.27 and G31.41 for the molecules detected in both sources, and Fig. 14b displays the corresponding modeled abundance ratios between reduced (appropriate for a Galactocentric distance of 13 kpc) and normal metallicity.

For most molecules, the observational ratios follow the qualitative trends predicted by the models, largely independent of the adopted model timescale. Nevertheless, quantitative differences of up to one to two orders of magnitude are found for several species. Overall, these results point to a general consistency with metallicity-dependent trends predicted by the models, while leaving room for significant deviations at the level of individual molecules. However, several species deviate from this general behavior. In particular, $\text{CH}_2(\text{OH})\text{CHO}$, CH_3OCHO , and $\text{a}-(\text{CH}_2\text{OH})_2$ exhibit the opposite trend to the models, showing

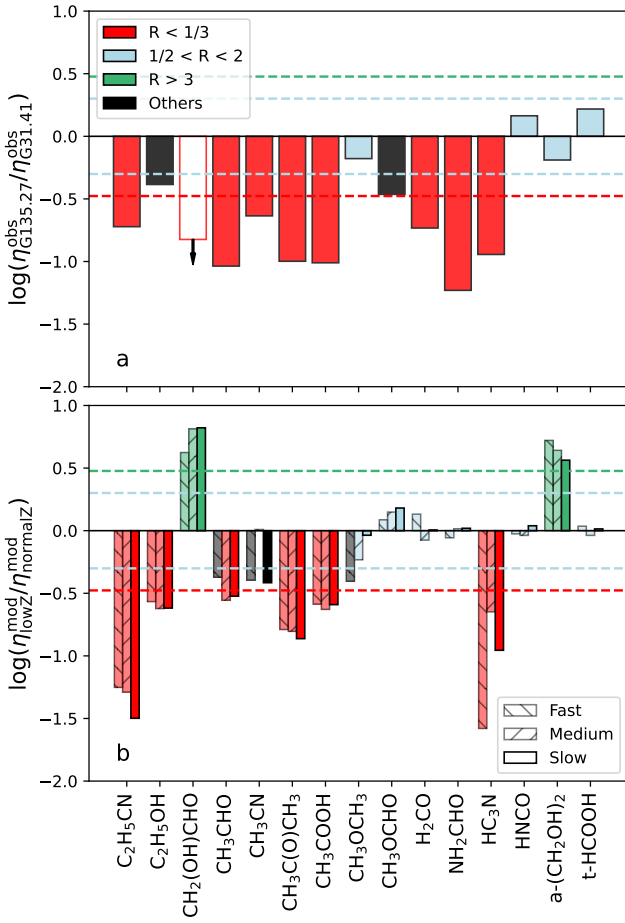


Fig. 14. Panel a: observed abundance ratios, R , between G135.27 MM1 and G31.41. Panel b: abundance ratios, R , of models ($A_{\text{V,init}} = 3$ mag) with low and normal metallicity. In both panels, the ratios correspond to $\eta = X/X_{\text{CH}_3\text{OH}}$. Molecules with R lower than one-third (dashed red line) or higher than three (dashed green line) are highlighted in red or green, respectively, while those with R within a factor of two of unity ($1/2 < R < 2$; dashed blue lines) are highlighted in blue. The other molecules are shown in black. The unfilled bar with arrow in panel a represents an observational upper limit for G135.27 MM1. In panel b, the solid bars correspond to the best-fit model of G135.27 MM1 ($t_{\text{w}} = 1 \times 10^6$ yr), while the hatched bars represent the ratios for the other two timescales.

lower relative abundances in the low- Z environment than in the normal- Z case, while the models predict the opposite for all modeled timescales. NH_2CHO also deviates strongly from the model predictions: while the model predicts a similar sensitivity to metallicity as methanol, the molecule is more than one order of magnitude less abundant (relative to methanol) in G135.27 than in G31.41. These four species contain, or are strongly related to, grain-surface chemistry involving the HCO radical. The reactions involving this species are strongly influenced by the atomic-H abundance on the grains, which is sensitive to the dust temperature. The use of only a single-point model, with an equally low minimum value of T_{dust} for both the local and outer-galaxy setups, could lead to an excess of HCO related species. A more dedicated evaluation of T_{dust} for outer-galaxy sources, based on lower dust content (and thus lower $A_{\text{V,init}}$) would be warranted, along with the use of multi-dimensional models of hot-core chemo-dynamical evolution (see Barger et al. 2021; Bonfand et al. in prep.). It is also useful to compare the behavior of these HCO group-bearing COMs with the behavior of HCO itself, as observed by Gigli et al. (2025) using data

from the CHEMOUT survey. Those authors found that HCO gas-phase abundances in cold cores scale with measured C/H ratios in the 8.34–24 kpc range of galactocentric radii. Methanol also shows a similar C/H dependence in those observations, indicating that HCO abundances with respect to methanol should be stable as a function of galactocentric radius (within the expressed range). Assuming that a meaningful comparison can be made between HCO-bearing COMs in hot gas and HCO in cold cores, our modeled abundances of methyl formate (CH_3OCHO) and formamide (NH_2CHO) appear to be more or less consistent with the CHEMOUT results. The other two species, glycolaldehyde ($\text{CH}_2(\text{OH})\text{CHO}$) and ethylene glycol ($\text{a}-(\text{CH}_2\text{OH})_2$), are still overproduced with respect to those data. However, for the reasons noted above, regarding dust temperature, it is unclear whether the abundance of cold HCO in the gas will provide an accurate reflection of COM production efficiencies on grains, or their gas-phase abundances at later times. CH_3CN and $\text{C}_2\text{H}_5\text{OH}$ are assigned different colors in Fig. 14a and b, but the agreement between the models and the observations is in fact relatively good with a ratio close to 1/3.

These results suggest that metallicity plays a role in shaping the abundances of certain COMs. However, other physical parameters—such as density structure, temperature evolution, and dynamical history—may also influence the chemical composition. Additionally, it should be noted that the metallicity of G31.41 is likely higher than solar, approximately $1.5 Z_{\odot}$ if we take the Galactic metallicity gradient of Luck & Lambert (2011). This implies that the current model, which assumes solar metallicity, may not be a good reference for the comparison shown in Fig. 14, and a revised model adjusted to the expected metallicity of G31.41 should be considered to provide more accurate predictions.

Overall, molecules that are predicted by our best-fit model in Fig. 14b to be more sensitive to metallicity relative to CH_3OH are generally confirmed to behave in this manner by the observational results, while only about half of the molecules that exhibit similar metallicity sensitivity to CH_3OH in the models do behave in this manner in the observations. One possibility is that our predictions focus primarily on the impact of metallicity, without accounting for other processes such as shocks, enhanced cosmic ray ionization, or local UV radiation, all of which can significantly alter molecular abundances. Another possibility is that the comparison is limited to a single source, making it difficult to draw statistically robust conclusions. Expanding the analysis to a larger sample of hot cores in the outer Galaxy will be essential to better constrain the role of metallicity relative to other physical and chemical effects.

6. Conclusions

We have carried out an imaging spectral line survey with NOEMA to probe the chemical composition of G135.27+2.79, a hot core candidate in the outer Galaxy ($D_{\text{gc}} = 13.1$ kpc). At the angular resolution (~ 5000 au) of the survey, G135.27 consists of three continuum cores, which we called MM1, MM2, and MM3. We derived the molecular composition of these cores assuming LTE. The main results of our analysis are the following:

1. We detected 28 molecules toward MM1 and several of their less abundant isotopologs, and we tentatively detected three additional molecules. Among these detected molecules, 13 (including a tentative detection of $\text{CH}_2(\text{OH})\text{CHO}$) are COMs, most of which trace a hot and compact region. This confirms G135.27 MM1 as a bona fide hot core, the third one identified in the outer Galaxy;

2. MM1 hosts a large-scale bipolar outflow traced by CO and ^{13}CO . Interestingly, COMs such as CH_3OH and CH_3OCHO trace a velocity gradient along the outflow axis at a scale of 600 au but with a direction opposite to that of the large-scale CO outflow. This feature may be linked to a wide-angle disk wind;
3. Only simple molecules trace MM2. Its CO spectrum shows an inverse P-Cygni profile but no sign of an outflow. MM2 may thus be a collapsing prestellar core. MM3 shows emission of simple molecules and extended CH_3CCH emission accompanied by a bipolar outflow traced by high-velocity SiO, suggesting that it is a young protostar.

To assess the impact of a reduced metallicity and dust-to-gas ratio in the outer Galaxy on hot-core chemistry, we compared the composition of MM1 with that of two other hot cores: G31.41 in the inner Galaxy and WB89-789 further out in the outer Galaxy. The COM abundances relative to CH_3OH in MM1 show a good correlation with those measured in G31.41 and WB89-789, but with values intermediate between those of G31.41 and WB89-789. This may reflect the influence of metallicity on the formation of COMs.

We ran the astrochemical model MAGICKAL to simulate the abundances of ~ 100 molecules under low-metallicity ($D_{\text{gc}}=13$ kpc) and normal-metallicity ($D_{\text{gc}}=8.2$ kpc) conditions. Slower adsorption onto grains due to the lower dust-to-gas ratio at low metallicity produces a more complete gas-phase conversion of the limited carbon available into CO, rather than it being adsorbed onto the grains to form methane or a related species. This behavior is further enhanced by the C:O ratio falling in the low-Z case. Larger hydrocarbon species, including radicals, are also depressed both in the gas phase and on the grains due to the reduced budget of free carbon. Thus, while other species may suffer, CO – and those species that may form from it via grain-surface chemistry – is less strongly affected by the decline in elemental C or O. Furthermore, the slow uptake of material onto the grains leads to lower coverages of CO, leading to more effective conversion to COMs of increasing degrees of saturation. COMs with production routes related to atomic C and/or pure hydrocarbon chemistry are therefore more likely to suffer diminished abundances at low-Z, including larger N-bearing species. The abundances of O-bearing COMs are often less diminished or even enhanced (with respect to methanol) under low-Z conditions.

The low-metallicity model that best fits the COM abundances of G135.27 MM1 relative to methanol corresponds to a slow warm-up timescale of $t_{\text{wu}} = 1 \times 10^6$ yr and an initial visual extinction of $A_{V,\text{init}} = 3$ mag. The abundances of most common hot-core species are reproduced within one order of magnitude, but $\text{CH}_2(\text{OH})\text{CHO}$, CH_3COOH , and NH_2CHO deviate more significantly.

In order to more carefully evaluate the impact of metallicity on the chemical composition of hot cores, we compared the abundance ratios of G135.27 MM1 and G31.41 to the abundance ratios of the low-Z and normal-Z models. Overall, the model reproduces well the behavior of most investigated molecules, demonstrating that it catches the influence of metallicity (and gas-to-dust ratio) on the chemistry relatively well. However, the model predicts opposite trends for $\text{CH}_2(\text{OH})\text{CHO}$, $\alpha\text{-(CH}_2\text{OH)}_2$, CH_3OCHO , and NH_2CHO compared to the observations, suggesting that their chemical networks require more attention or that the treatment of extinction and dust temperature need to be more tailored to specific local conditions.

Overall, these comparisons indicate that metallicity has a significant impact on the formation of COMs. Expanding the

source sample will be essential to assessing the robustness of the observed deviations, while incorporating additional processes (e.g., shocks, cosmic rays, UV radiation) into the models may further improve the agreement between models and observations.

Data availability

Appendices E–K are available at Zenodo (DOI: [10.5281/zenodo.17779972](https://doi.org/10.5281/zenodo.17779972)).

Acknowledgements. We gratefully thank Karl M. Menten, whose wisdom, guidance, and encouragement inspired this research. Although he passed away before the completion of this article, his influence is deeply felt throughout this work. This paper is dedicated to his memory. We thank the IRAM staff for carrying out the NOEMA observations of project S23AD and Michael Bremer for his support regarding the calibration of the data. IRAM is supported by INSU/CNRS (France), MPG (Germany) and IGN (Spain). We thank L. Coudert and S. Spezzano for providing spectroscopic predictions for CH_2DOH . Part of this work has been carried out within the Collaborative Research Center 1601 (SFB 1601 sub-project A1) funded by the Deutsche Forschungsgemeinschaft (DFG, German Research Foundation) – project ID 500700252. YXW received financial support for this research from the International Max Planck Research School (IMPRS) for Astronomy and Astrophysics at the Universities of Bonn and Cologne. RTG thanks the National Science Foundation for funding through the Astronomy & Astrophysics program (grant number 2206516). RTG also thanks the Alexander von Humboldt Foundation for funding a renewed research stay in Germany in Spring 2025.

References

- Álvarez-Barcia, S., Russ, P., Kästner, J., & Lamberts, T. 2018, *MNRAS*, **479**, 2007
- Arellano-Córdova, K. Z., Esteban, C., García-Rojas, J., & Méndez-Delgado, J. E. 2020, *MNRAS*, **496**, 1051
- Artur de la Villarmois, E., Guzmán, V. V., Jørgensen, J. K., et al. 2022, *A&A*, **667**, A20
- Bachiller, R. 1996, *ARA&A*, **34**, 111
- Bachiller, R., & Pérez Gutiérrez, M. 1997, *ApJ*, **487**, L93
- Bacmann, A., Taquet, V., Faure, A., Kahane, C., & Ceccarelli, C. 2012, *A&A*, **541**, L12
- Bally, J. 2016, *ARA&A*, **54**, 491
- Barger, C. J., Lam, K. H., Li, Z.-Y., et al. 2021, *A&A*, **651**, A43
- Belloche, A., Müller, H. S. P., Garrod, R. T., & Menten, K. M. 2016, *A&A*, **587**, A91
- Belloche, A., Garrod, R. T., Müller, H. S. P., et al. 2019, *A&A*, **628**, A10
- Belloche, A., Maury, A. J., Maret, S., et al. 2020, *A&A*, **635**, A198
- Belloche, A., Garrod, R. T., Müller, H. S. P., et al. 2025, *A&A*, **698**, A143
- Bernal, J. J., Sephus, C. D., & Ziurys, L. M. 2021, *ApJ*, **922**, 106
- Beuther, H., Schilke, P., Sridharan, T. K., et al. 2002, *A&A*, **383**, 892
- Blair, S. K., Magnani, L., Brand, J., & Wouterloot, J. G. A. 2008, *Astrobiology*, **8**, 59
- Bonfand, M., Belloche, A., Menten, K. M., Garrod, R. T., & Müller, H. S. P. 2017, *A&A*, **604**, A60
- Bouscasse, L., Csengeri, T., Wyrowski, F., Menten, K. M., & Bontemps, S. 2024, *A&A*, **686**, A252
- Bouvier, M., Viti, S., Mangum, J. G., et al. 2025, *A&A*, **698**, A261
- Brand, J., & Wouterloot, J. G. A. 2007, *A&A*, **464**, 909
- Brand, J., Wouterloot, J. G. A., Rudolph, A. L., & de Geus, E. J. 2001, *A&A*, **377**, 644
- Busch, L. A., Belloche, A., Garrod, R. T., Müller, H. S. P., & Menten, K. M. 2022, *A&A*, **665**, A96
- Busch, L. A., Belloche, A., Garrod, R. T., Müller, H. S. P., & Menten, K. M. 2024, *A&A*, **681**, A104
- Calcutt, H., Jørgensen, J. K., Müller, H. S. P., et al. 2018, *A&A*, **616**, A90
- Ceccarelli, C. 2004, in *Astronomical Society of the Pacific Conference Series*, **323**, Star Formation in the Interstellar Medium: In Honor of David Hollenbach, eds. D. Johnstone, F. C. Adams, D. N. C. Lin, D. A. Neufeld, & E. C. Ostriker, 195
- Cesaroni, R. 2005, in *IAU Symposium*, **227**, Massive Star Birth: A Crossroads of Astrophysics, eds. R. Cesaroni, M. Felli, E. Churchwell, & M. Walmsley, 59
- Chen, H.-R. V., Keto, E., Zhang, Q., et al. 2016, *ApJ*, **823**, 125

- Chen, Y., van Gelder, M. L., Nazari, P., et al. 2023, *A&A*, **678**, A137
- Chen, Y., Garrod, R. T., Rachid, M., et al. 2025, *A&A*, **696**, A198
- Choudhury, S., Subramaniam, A., & Cole, A. A. 2016, *MNRAS*, **455**, 1855
- Choudhury, S., Subramaniam, A., Cole, A. A., & Sohn, Y. J. 2018, *MNRAS*, **475**, 4279
- Choudhury, S., de Grijs, R., Bekki, K., et al. 2021, *MNRAS*, **507**, 4752
- Colzi, L., Rivilla, V. M., Beltrán, M. T., et al. 2021, *A&A*, **653**, A129
- Cortés, P. C., Carpenter, J., Kamenó, S., et al. 2025, *ALMA Technical Handbook*, ALMA Doc. 12.3, ver. 1.5
- Coudert, L. H., Zemouli, M., Motiyenko, R. A., Margulès, L., & Klee, S. 2014, *J. Chem. Phys.*, **140**, 064307
- De Simone, M., Podio, L., Chahine, L., et al. 2024, *A&A*, **686**, L13
- Di Francesco, J., Myers, P. C., Wilner, D. J., Ohashi, N., & Mardones, D. 2001, *ApJ*, **562**, 770
- Ellingsen, S. P. 2006, *ApJ*, **638**, 241
- Endres, C. P., Schlemmer, S., Schilke, P., Stutzki, J., & Müller, H. S. P. 2016, *J. Mol. Spectrosc.*, **327**, 95
- Evans, II, N. J., Di Francesco, J., Lee, J.-E., et al. 2015, *ApJ*, **814**, 22
- Fadul, A. M. A., Schwarz, K. R., van'T Hoff, M. L. R., et al. 2025, *AJ*, **169**, 307
- Fayolle, E. C., Öberg, K. I., Garrod, R. T., van Dishoeck, E. F., & Bisschop, S. E. 2015, *A&A*, **576**, A45
- Flower, D. R., Pineau des Forêts, G., & Rabli, D. 2010, *MNRAS*, **409**, 29
- Fontani, F., Colzi, L., Bizzocchi, L., et al. 2022a, *A&A*, **660**, A76
- Fontani, F., Schmiedeke, A., Sánchez-Monge, A., et al. 2022b, *A&A*, **664**, A154
- Fontani, F., Vermariën, G., Viti, S., et al. 2024, *A&A*, **691**, A180
- García de la Concepción, J., Colzi, L., Jiménez-Serra, I., et al. 2022, *A&A*, **658**, A150
- Garrod, R. T., & Herbst, E. 2006, *A&A*, **457**, 927
- Garrod, R. T., & Pauly, T. 2011, *ApJ*, **735**, 15
- Garrod, R. T., Widicus Weaver, S. L., & Herbst, E. 2008, *ApJ*, **682**, 283
- Garrod, R. T., Jin, M., Matis, K. A., et al. 2022, *ApJS*, **259**, 1
- Giannetti, A., Leurini, S., König, C., et al. 2017, *A&A*, **606**, L12
- Gigli, D., Fontani, F., Colzi, L., et al. 2025, *A&A*, **704**, A171
- Goldsmith, P. F., & Langer, W. D. 1999, *ApJ*, **517**, 209
- Green, J. A., Breen, S. L., Fuller, G. A., et al. 2017, *MNRAS*, **469**, 1383
- Hachisuka, K., Brunthaler, A., Menten, K. M., et al. 2009, *ApJ*, **696**, 1981
- Hamedani Golshan, R., Sánchez-Monge, Á., Schilke, P., et al. 2024, *A&A*, **688**, A3
- Hasegawa, T. I., Herbst, E., & Leung, C. M. 1992, *ApJS*, **82**, 167
- Herbst, E., & van Dishoeck, E. F. 2009, *ARA&A*, **47**, 427
- Hocuk, S., Szűcs, L., Caselli, P., et al. 2017, *A&A*, **604**, A58
- Hsu, S.-Y., Lee, C.-F., Johnstone, D., et al. 2025, *ApJ*, **989**, 56
- Huang, K.-Y., Behrens, E., Bouvier, M., et al. 2025, *A&A*, **699**, A70
- Ikeda, T., Shimonishi, T., Izumi, N., et al. 2025, *ApJ*, **988**, 111
- Ikeda, T., Shimonishi, T., Kaneko, H., et al. 2026, *ApJ*, **999**, 194
- Immer, K., Li, J., Quiroga-Núñez, L. H., et al. 2019, *A&A*, **632**, A123
- Jiménez-Serra, I., Caselli, P., Martín-Pintado, J., & Hartquist, T. W. 2008, *A&A*, **482**, 549
- Jin, M., & Garrod, R. T. 2020, *ApJS*, **249**, 26
- Jin, M., Lam, K. H., McClure, M. K., et al. 2022, *ApJ*, **935**, 133
- Jørgensen, J. K., van der Wiel, M. H. D., Coutens, A., et al. 2016, *A&A*, **595**, A117
- Jørgensen, J. K., Müller, H. S. P., Calcutt, H., et al. 2018, *A&A*, **620**, A170
- Jørgensen, J. K., Belloche, A., & Garrod, R. T. 2020, *ARA&A*, **58**, 727
- Kauffmann, J., Bertoldi, F., Bourke, T. L., Evans, II, N. J., & Lee, C. W. 2008, *A&A*, **487**, 993
- Koelemay, L. A., & Ziurys, L. M. 2025, *ApJ*, **986**, 122
- Kurtz, S., Cesaroni, R., Churchwell, E., Hofner, P., & Walmsley, C. M. 2000, in *Protostars and Planets IV*, eds. V. Mannings, A. P. Boss, & S. S. Russell, 299
- Lee, J.-E., Lee, S., Baek, G., et al. 2019, *Nat. Astron.*, **3**, 314
- Liu, C.-F., Shang, H., Johnstone, D., et al. 2025, *ApJ*, **979**, 17
- Long, F., Herczeg, G. J., Harsono, D., et al. 2019, *ApJ*, **882**, 49
- López-Gallifa, Á., Rivilla, V. M., Beltrán, M. T., et al. 2024, *MNRAS*, **529**, 3244
- López-Gallifa, Á., Rivilla, V. M., Beltrán, M. T., et al. 2025, *A&A*, **704**, A288
- Luck, R. E., & Lambert, D. L. 2011, *AJ*, **142**, 136
- Maret, S., Hily-Blant, P., Pety, J., Bardeau, S., & Reynier, E. 2011, *A&A*, **526**, A47
- Martín, S., Mangum, J. G., Harada, N., et al. 2021, *A&A*, **656**, A46
- Maud, L. T., Moore, T. J. T., Lumsden, S. L., et al. 2015, *MNRAS*, **453**, 645
- Maury, A. J., André, P., Testi, L., et al. 2019, *A&A*, **621**, A76
- McClure, M. K., Rocha, W. R. M., Pontoppidan, K. M., et al. 2023, *Nat. Astron.*, **7**, 431
- Méndez-Delgado, J. E., Amayo, A., Arellano-Córdova, K. Z., et al. 2022, *MNRAS*, **510**, 4436
- Minier, V., Ellingsen, S. P., Norris, R. P., & Booth, R. S. 2003, *A&A*, **403**, 1095
- Mininni, C., Beltrán, M. T., Rivilla, V. M., et al. 2020, *A&A*, **644**, A84
- Mininni, C., Beltrán, M. T., Colzi, L., et al. 2023, *A&A*, **677**, A15
- Motiyenko, R. A., & Margulès, L. 2025, *A&A*, **699**, A348
- Mottram, J. C., van Dishoeck, E. F., Schmalzl, M., et al. 2013, *A&A*, **558**, A126
- Müller, H. S. P., Thorwirth, S., Roth, D. A., & Winnewisser, G. 2001, *A&A*, **370**, L49
- Müller, H. S. P., Schlöder, F., Stutzki, J., & Winnewisser, G. 2005, *J. Mol. Struct.*, **742**, 215
- Ossenkopf, V., & Henning, T. 1994, *A&A*, **291**, 943
- Panagia, N. 1973, *AJ*, **78**, 929
- Pestalozzi, M. R., Minier, V., & Booth, R. S. 2005, *A&A*, **432**, 737
- Pickett, H. M., Poynter, R. L., Cohen, E. A., et al. 1998, *J. Quant. Spec. Radiat. Transf.*, **60**, 883
- Qin, S.-L., Liu, T., Liu, X., et al. 2022, *MNRAS*, **511**, 3463
- Rachid, M. G., Rocha, W. R. M., & Linnartz, H. 2022, *A&A*, **665**, A89
- Reid, M. J., Menten, K. M., Brunthaler, A., et al. 2019, *ApJ*, **885**, 131
- Requena-Torres, M. A., Martín-Pintado, J., Martín, S., & Morris, M. R. 2008, *ApJ*, **672**, 352
- Rocha, W. R. M., van Dishoeck, E. F., Ressler, M. E., et al. 2024, *A&A*, **683**, A124
- Rudolph, A. L., Brand, J., de Geus, E. J., & Wouterloot, J. G. A. 1996, *ApJ*, **458**, 653
- Sewilo, M., Indebetouw, R., Charnley, S. B., et al. 2018, *ApJ*, **853**, L19
- Sewilo, M., Cordiner, M., Charnley, S. B., et al. 2022, *ApJ*, **931**, 102
- Shimonishi, T. 2024, arXiv e-prints [arXiv:2411.04451]
- Shimonishi, T., Onaka, T., Kawamura, A., & Aikawa, Y. 2016, *ApJ*, **827**, 72
- Shimonishi, T., Das, A., Sakai, N., et al. 2020, *ApJ*, **891**, 164
- Shimonishi, T., Izumi, N., Furuya, K., & Yasui, C. 2021, *ApJ*, **922**, 206
- Shimonishi, T., Tanaka, K. E. I., Zhang, Y., & Furuya, K. 2023, *ApJ*, **946**, L41
- Shope, B. M., El-Abd, S. J., Brogan, C. L., et al. 2024, *ApJ*, **972**, 146
- Simons, M. A. J., Lamberts, T., & Cuppen, H. M. 2020, *A&A*, **634**, A52
- Suutarinen, A. N., Kristensen, L. E., Mottram, J. C., Fraser, H. J., & van Dishoeck, E. F. 2014, *MNRAS*, **440**, 1844
- Tafalla, M., Santiago-García, J., Hacar, A., & Bachiller, R. 2010, *A&A*, **522**, A91
- Tobin, J. J., Hartmann, L., Bergin, E., et al. 2012, *ApJ*, **748**, 16
- Tychoniec, Ł., Hull, C. L. H., Kristensen, L. E., et al. 2019, *A&A*, **632**, A101
- Tychoniec, Ł., van Dishoeck, E. F., van't Hoff, M. L. R., et al. 2021, *A&A*, **655**, A65
- van Dishoeck, E. F., & Blake, G. A. 1998, *ARA&A*, **36**, 317
- Vastel, C., Ceccarelli, C., Lefloch, B., & Bachiller, R. 2014, *ApJ*, **795**, L2
- Walmsley, M. 1995, in *Revista Mexicana de Astronomía y Astrofísica Conference Series*, 1, Revista Mexicana de Astronomía y Astrofísica Conference Series, eds. S. Lizano, & J. M. Torrelles, 137
- Walsh, C., Millar, T., Nomura, H., et al. 2013, in *Protostars and Planets VI Posters*
- Watanabe, N., & Kouchi, A. 2002, *ApJ*, **571**, L173
- Wilson, T. L., & Rood, R. 1994, *ARA&A*, **32**, 191
- Wouterloot, J. G. A., & Brand, J. 1989, *A&AS*, **80**, 149
- Wouterloot, J. G. A., & Brand, J. 1996, *A&AS*, **119**, 439
- Wouterloot, J. G. A., Brand, J., & Henkel, C. 1988, *A&A*, **191**, 323
- Wouterloot, J. G. A., Brand, J., & Fiegle, K. 1993, *A&AS*, **98**, 589
- Yan, Y. T., Henkel, C., Kobayashi, C., et al. 2023, *A&A*, **670**, A98
- Yang, Y.-L., Green, J. D., Pontoppidan, K. M., et al. 2022, *ApJ*, **941**, L13
- Zeng, S., Jiménez-Serra, I., Rivilla, V. M., et al. 2018, *MNRAS*, **478**, 2962

Appendix A: Observational parameters

Tables A.1 and A.2 summarize the observational setups, along with the synthesized beam sizes and rms noise levels in the final data cubes.

Table A.1: Observational setups of the spectral line survey.

Setup	Frequency range		Date of observation yyyy.mm.dd	Baseline range m	$t_{\text{int}}^{(a)}$ h	Calibrators ^(b)			
	LSB GHz	USB GHz				P	A	B	F
S1	215.1–223.2	230.5–238.7	2023.12.03	23.1–328.0	2.6	1,2	1,2	3	4
			2023.12.05	19.5–313.6	3.4	1,2	1,2	3	4
S2	222.8–230.9	238.3–246.4	2023.10.05	17.7–368.0	5.2	1,2	1,2	3	4,5

Notes. ^(a) On-source integration time. ^(b) Phase (P), amplitude (A), bandpass (B), and absolute flux (F) calibrators. The calibrators are: 1: J0244+6228, 2: 0224+671, 3: 3C84, 4: LKHA101, 5: MWC349.

Table A.2: Beam sizes and noise levels.

Setup	SPW ^(a)	mode ^(b)	Frequency range MHz	Synthesized beam		Linear size au	rms	
				HPBW "×"	PA °		line mJy beam ^{−1}	continuum mJy beam ^{−1}
S1	LO	SURVEY	215065–219123	0.89×0.78	137	5340×4680	6.27	0.12
	LI	SURVEY	219130–223188	0.86×0.76	136	5160×4560	5.56	0.14
	UI+CO	SPLIT	230549–230772	0.82×0.72	135	4920×4320	6.76	0.39
	UI−CO	SURVEY	230772–234605	0.82×0.72	135	4920×4320	5.95	0.13
	UO	SURVEY	234614–238673	0.81×0.70	137	4860×4200	6.91	0.15
S2	LO	SURVEY	222807–226863	0.90×0.70	170	5400×4200	5.77	0.13
	LI−CO	SURVEY	226872–230368	0.87×0.70	0	5220×4200	5.54	0.12
	LI+CO	SPLIT	230368–230930	0.87×0.70	0	5220×4200	6.07	0.25
	UI	SURVEY	238290–242346	0.84×0.66	169	5040×3960	6.16	0.19
	UO	SURVEY	242355–246413	0.83×0.65	0	4980×3900	6.77	0.17

Notes. ^(a) Spectral window. S1-UI and S2-LI are each divided into two subwindows: +CO indicates the subwindow that contains the CO 2–1 line and −CO the subwindow that does not. ^(b) Pipeline mode used in IMAGER for imaging.

Appendix B: LTE model parameters

Tables B.1-B.3 list the input parameters of our best-fit LTE models of MM1, MM2, and MM3.

Table B.1: Parameters of our best-fit LTE model of MM1.

Molecule	Tag ^a	N_{det}^b	size ^c "	T_{rot}^d (K)	N^e (cm ^{−2})	$T_{\text{rot,pd}}^f$ (K)	N_{pd}^g (cm ^{−2})	C_{vib}^h	C_{conf}^i	V_{off}^j (km s ^{−1})	ΔV^k
HDO	19002	2	0.2	260	2.5(17)	260±0	3.0±0.1(17)	1.00	1.00	0.4	5.0
			0.2	260	2.0(16)			1.00	1.00	−1.4	2.0
CN	26504	4	3.9	10	2.5(14)	...	...	1.00	1.00	0.0	4.0
CO	28503	1	3	75	1.0(18)	...	...	1.00	1.00	−7.4	51.0
			3	75	1.0(18)			1.00	1.00	2.0	13.0
			3	75	6.0(16)			1.00	1.00	−82.6	18.5
¹³ CO	29802	1	2	75	4.0(17)	...	...	1.00	1.00	0.0	4.5
C ¹⁸ O	30502	1	2.2	75	5.6(16)	...	...	1.00	1.00	0.0	3.0
C ¹⁷ O	29603	1	2.2	75	1.2(16)	...	...	1.00	1.00	0.0	3.0
DCN	28802	1	1.9	75	5.6(13)	...	...	1.00	1.00	0.1	4.3
H ₂ CO	30501	5	0.3	175	7.5(16)	103±3	7.6±0.3(16)	1.00	1.00	0.0	4.9
			0.9	50	1.9(15)			1.00	1.00	0.0	2.8
H ₂ ¹³ CO	31503	1	0.3	175	1.5(15)	...	...	1.00	1.00	0.0	4.9
CH ₃ OH, $v = 0$	32701	52	0.3	175	9.5(17)	175±2	2.1±0.1(18)	1.00	1.00	1.5	3.5
			0.3	175	6.5(17)			1.00	1.00	−1.3	3.0
CH ₃ OH, $v_t = 1$	32702	12	0.3	175	9.5(17)	175±2	2.1±0.1(18)	1.00	1.00	1.5	3.5
			0.3	175	6.5(17)			1.00	1.00	−1.3	3.0
CH ₃ OH, $v_t = 2$	32703	4	0.3	175	9.5(17)	175±2	2.1±0.1(18)	1.00	1.00	1.5	3.5

Table B.1: Continued.

Molecule	Tag ^a	N_{det}^b	size ^c "	T_{rot}^d (K)	N^e (cm ⁻²)	$T_{\text{rot,pd}}^f$ (K)	N_{pd}^g (cm ⁻²)	C_{vib}^h	C_{conf}^i	V_{off}^j (km s ⁻¹)	ΔV^k
			0.3	175	6.5(17)			1.00	1.00	-1.3	3.0
CH ₂ DOH	33902	4	0.3	80	9.0(15)	76±19	1.1±0.5(16)	1.00	1.00	-1.0	2.9
¹³ CH ₃ OH, $v = 0$	33401	15	0.3	115	3.0(16)	117±9	5.9±0.6(16)	1.00	1.00	1.5	3.5
			0.3	115	2.5(16)			1.00	1.00	-1.3	3.0
¹³ CH ₃ OH, $v_t = 1^{(l)}$	33402	0	0.3	115	3.0(16)	...	...	1.00	1.00	1.5	3.5
			0.3	115	2.5(16)			1.00	1.00	-1.3	3.0
H ₂ S	34502	1	1.0	75	4.2(15)	...	...	1.00	1.00	0.3	4.0
CH ₃ CCH, $v = 0$	40502	6	1.9	75	6.4(14)	78±17	6.4±2.4(14)	1.00	1.00	0.0	2.7
CH ₃ CN, $v = 0$	41601	17	0.3	190	6.0(15)	193±6	1.2±0.1(16)	1.00	1.00	1.5	3.5
			0.3	190	4.0(15)			1.00	1.00	-1.3	3.0
CH ₃ CN, $v_8 = 1$	41603	17	0.3	190	6.0(15)	193±6	1.2±0.1(16)	1.00	1.00	0.5	6.1
			0.3	190	4.0(15)			1.00	1.00	-1.3	3.0
CH ₂ CO, $v = 0$	42501	4	0.3	110	2.0(15)	103±19	3.6±1.0(15)	1.00	1.00	1.2	3.4
			0.3	110	1.5(15)			1.00	1.00	-1.8	2.6
HNCO, $v = 0$	43511	7	0.3	180	2.5(16)	177±4	2.9±0.1(16)	1.02	1.00	0.9	5.0
			0.3	180	2.5(15)			1.02	1.00	-1.2	2.0
CH ₃ CHO, $v = 0$	44301	9	0.8	65	3.5(14)	63±13	5.4±1.9(14)	1.00	1.00	1.5	3.5
			0.8	65	1.5(14)			1.00	1.00	-1.3	3.0
CS	44501	1	1.7	75	5.2(15)	...	...	1.00	1.00	0.0	4.0
¹³ CS	45403	1	0.8	75	3.6(14)	...	...	1.00	1.00	0.0	4.0
C ³⁴ S	46501	1	1.1	75	8.7(14)	...	...	1.00	1.00	0.0	4.0
C ³³ S	45404	1	0.8	75	4.1(14)	...	...	1.00	1.00	0.0	4.0
¹³ C ³⁴ S ^(l)	47501	1	0.8	75	1.8(13)	...	...	1.00	1.00	0.0	4.0
SiO	44505	1	1.2	75	8.5(13)	...	...	1.00	1.00	0.5	4.0
NH ₂ CHO, $v = 0$	45512	11	0.3	110	1.0(15)	110±26	2.0±0.6(15)	1.02	1.00	1.5	3.5
			0.3	110	6.0(14)			1.02	1.00	-1.3	3.0
t-HCOOH	46506	14	0.3	190	3.7(16)	191±22	3.9±0.5(16)	1.02	1.00	0.7	5.0
H ₂ CS	46509	7	1.	105	2.6(15)	103±7	2.8±0.3(15)	1.00	1.00	0.2	4.9
CH ₃ OCH ₃ , $v = 0$	46514	11	0.3	120	5.1(16)	117±5	9.0±0.8(16)	1.00	1.00	-1.2	3.2
			0.3	120	3.5(16)			1.00	1.00	1.7	2.4
C ₂ H ₅ OH, $v = 0$	46524	23	0.3	85	2.0(16)	88±10	3.2±0.7(16)	1.05	1.00	1.0	3.0
			0.3	85	1.1(16)			1.05	1.00	-2.1	1.8
SO	48501	4	1.2	75	1.1(16)	...	...	1.00	1.00	0.0	4.0
³³ SO	49501	1	0.4	75	1.5(15)	...	...	1.00	1.00	0.0	4.0
³⁴ SO	50501	1	0.5	75	4.5(15)	...	...	1.00	1.00	0.0	4.0
S ¹⁸ O	50502	2	0.5	75	3.2(14)	...	...	1.00	1.00	0.0	4.0
HC ₃ N, $v = 0$	51501	4	0.5	135	5.1(14)	136±14	5.1±1.0(14)	1.41	1.00	0.9	5.0
HC ₃ N, $v_7 = 1$	51502	2	0.5	135	5.1(14)	...	...	1.41	1.00	0.9	5.0
C ₂ H ₅ CN, $v = 0$	55913	20	0.3	190	1.7(15)	...	...	1.73	1.00	1.5	3.5
CH ₃ C(O)CH ₃ , $v = 0$	58912	13	0.3	82	5.0(15)	82±8	9.1±1.7(15)	1.00	1.00	1.5	3.5
			0.3	82	4.0(15)			1.00	1.00	-1.3	3.0
CH ₂ (OH)CHO, $v = 0^{(l)}$	60501	3	0.3	135	1.2(15)	...	...	1.21	1.00	0.5	2.7
CH ₃ COOH, $v_t = 0$	60523	13	0.3	135	1.0(16)	...	...	1.01	1.00	0.6	2.7
CH ₃ OCHO, $v = 0$	60103	71	0.3	135	5.7(16)	136±4	1.2±0.1(17)	1.14	1.00	1.7	2.7
			0.3	135	5.7(16)			1.14	1.00	-1.7	3.0
CH ₃ OCHO, $v_t = 1$	60104	44	0.3	135	5.7(16)	136±4	1.2±0.1(17)	1.14	1.00	1.7	2.7
			0.3	135	5.7(16)			1.14	1.00	-1.7	3.0
OCS, $v=0$	60503	3	0.4	115	5.0(16)	111±7	5.7±0.5(16)	1.00	1.00	0.0	5.0
O ¹³ CS	61502	1	0.3	115	2.0(15)	...	...	1.00	1.00	0.0	5.0
OC ³⁴ S	62505	2	0.3	115	6.8(15)	...	...	1.00	1.00	0.0	5.0
OC ³³ S ^(l)	61503	3	0.4	115	8.0(14)	...	...	1.00	1.00	0.0	5.0
a-(CH ₂ OH) ₂	62503	23	0.3	175	1.5(16)	...	...	1.91	1.19	0.5	3.9
SO ₂ , $v = 0$	64502	19	0.3	142	3.0(17)	149±4	3.6±0.3(17)	1.00	1.00	0.0	4.5
SO ₂ , $v_2 = 1^{(l)}$	64503	0	0.3	142	3.0(17)	...	...	1.00	1.00	0.0	4.5
³³ SO ₂ , $v = 0$	65501	3	0.3	142	3.8(15)	...	...	1.00	1.00	0.5	4.5
³⁴ SO ₂ , $v = 0$	66501	13	0.3	142	1.8(16)	...	...	1.01	1.00	0.0	4.5

Notes. ^(a) Entry number in our Weeds local database; see Sect. 3.4.1. ^(b) Number of detected lines. One line of a given species may mean a group of transitions of this species that are blended together. ^(c) Source diameter (FWHM). ^(d) Rotational temperature used for the Weeds model. ^(e) Column density used for the Weeds model. $X(Y)$ means $X \times 10^Y$. ^(f) Rotational temperature derived from the population diagram. ^(g) Column density derived from the population diagram. In cases where molecules have two velocity components, the value represents the total column density summed over both components. ^(h) Correction factor that was applied to the column density to account for the contribution of vibrationally or torsionally excited states, in the cases where this contribution was not included in the partition function of the spectroscopic predictions. In most cases, this factor was estimated in the harmonic approximation (Eq. (1) in Belloche et al. 2025). ⁽ⁱ⁾ Correction factor that was applied to the column density to account for the contribution of conformers, in the cases where this contribution was not included in the partition function of the spectroscopic predictions. ^(j) Velocity offset with respect to the systemic velocity of MM1 (-71.7 km s^{-1}). ^(k) Linewidth (FWHM). ^(l) Tentative detection.

Table B.2: Parameters of our best-fit LTE model of MM2.

Molecule	Tag ^a	N_{det}^b	size ^c "	T_{rot}^d (K)	N^e (cm^{-2})	$T_{\text{rot,pd}}^f$ (K)	N_{pd}^g (cm^{-2})	C_{vib}^h	C_{conf}^i	V_{off}^j (km s^{-1})	ΔV^k
CN	26504	1	4.5	7	3.0(14)	...	...	1.00	1.00	-0.7	1.5
CO	28503	1	4.5	12	1.0(17)	...	...	1.00	1.00	2.0	13
¹³ CO	29802	1	4.5	12	2.0(15)	...	...	1.00	1.00	0.0	4.5
H ₂ CO	30501	4	4.5	41	1.3(14)	41±5	1.3±0.2(14)	1.00	1.00	-1.0	1.5
CS	44501	1	4.5	12	6.2(14)	...	...	1.00	1.00	-1.5	1.5
C ³⁴ S	46501	1	4.5	12	3.1(13)	...	...	1.00	1.00	-1.0	1.5
CH ₃ OH, $v = 0$	32701	3	4.5	60	2.0(14)	49±3	5.7±0.6(14)	1.00	1.00	-0.5	2
H ₂ CS	46509	4	4.5	12	1.3(14)	12±5	1.4±2.8(14)	1.00	1.00	-1.0	1.5

Notes. Same as Table B.1, but a fixed source size of 4.5" was adopted, for a beam filling factor of about unity, as the source size could not be fitted and the molecular lines trace extended emission.

Table B.3: Parameters of our best-fit LTE model of MM3.

Molecule	Tag ^a	N_{det}^b	size ^c "	T_{rot}^d (K)	N^e (cm^{-2})	$T_{\text{rot,pd}}^f$ (K)	N_{pd}^g (cm^{-2})	C_{vib}^h	C_{conf}^i	V_{off}^j (km s^{-1})	ΔV^k
CN	26504	1	4.5	10	4.5(14)	...	...	1.00	1.00	1.0	2.0
CO	28503	1	4.5	43	6.0(17)	...	...	1.00	1.00	5.0	4.0
			4.5	43	4.0(16)			1.00	1.00	0.0	2.5
			4.5	43	3.0(16)			1.00	1.00	-3.5	2.0
			4.5	43	6.0(16)			1.00	1.00	-18.0	12.0
¹³ CO	29802	1	4.5	43	2.5(16)	...	...	1.00	1.00	1.0	2.0
			4.5	43	2.5(16)			1.00	1.00	4.0	3.0
C ¹⁷ O	29603	1	4.5	43	1.6(15)	...	...	1.00	1.00	1.0	2.0
C ¹⁸ O	30502	1	4.5	43	5.5(15)	...	...	1.00	1.00	1.0	2.0
			4.5	43	3.0(15)			1.00	1.00	4.0	3.0
DNC	28508	1	4.5	43	2.0(12)	...	...			1.0	1.5
DCN	28802	1	4.5	43	1.1(13)	...	...	1.00	1.00	1.0	1.5
H ₂ CO	30501	4	4.5	87	2.8(14)	89±8	3.5±0.4(14)	1.00	1.00	1.0	1.5
CH ₃ OH, $v = 0$	32701	10	4.5	37	4.0(14)	36±8	4.4±1.7(14)	1.00	1.00	1.0	1.5
H ₂ S	34502	1	4.5	43	2.0(14)	...	...	1.00	1.00	1.5	1.5
CH ₃ CCH, $v = 0$	40502	7	4.5	55	1.8(14)	54±12	1.8±8.7(13)	1.00	1.00	1.0	1.5
CS	44501	1	4.5	43	2.1(14)	...	...	1.00	1.00	1.0	2.0
			4.5	43	3.5(13)			1.00	1.00	4.0	3.0
¹³ CS	45403	1	4.5	43	1.0(13)	...	...	1.00	1.00	1.0	2.0
C ³³ S	45404	1	4.5	43	1.5(13)	...	...	1.00	1.00	1.0	2.0
C ³⁴ S	46501	1	4.5	43	4.5(13)	...	...	1.00	1.00	1.0	2.0
SiO	44505	1	4.5	43	1.5(12)	...	...	1.00	1.00	1.0	1.5
			4.5	43	1.6(12)			1.00	1.00	4.5	2.0
H ₂ CS	46509	6	4.5	43	1.3(14)	42±2	1.4±0.2(14)	1.00	1.00	1.0	1.5
SO	48501	1	4.5	43	1.3(14)	...	...	1.00	1.00	1.0	2.0
			4.5	43	6.1(13)			1.00	1.00	4.0	3.0
HC ₃ N, $v = 0$	51501	4	4.5	27	8.0(13)	27±6	8.8±11.3(13)	1.41	1.00	1.0	1.5
SO ₂ , $v = 0$	64502	1	4.5	87	1.0(14)	...	...	1.00	1.00	1.0	2.0
			4.5	87	8.0(13)			1.00	1.00	4.0	3.0

Notes. Same as Table B.1, but a fixed source size of 4.5" was adopted, for a beam filling factor of about unity, as the source size could not be fitted and the molecular lines trace extended emission.

Appendix C: Emission size

Figure C.1 presents the deconvolved size of CH₃OH plotted against upper-level energy and line optical depth. Figure C.2 displays the deconvolved emission sizes for all detected molecules.

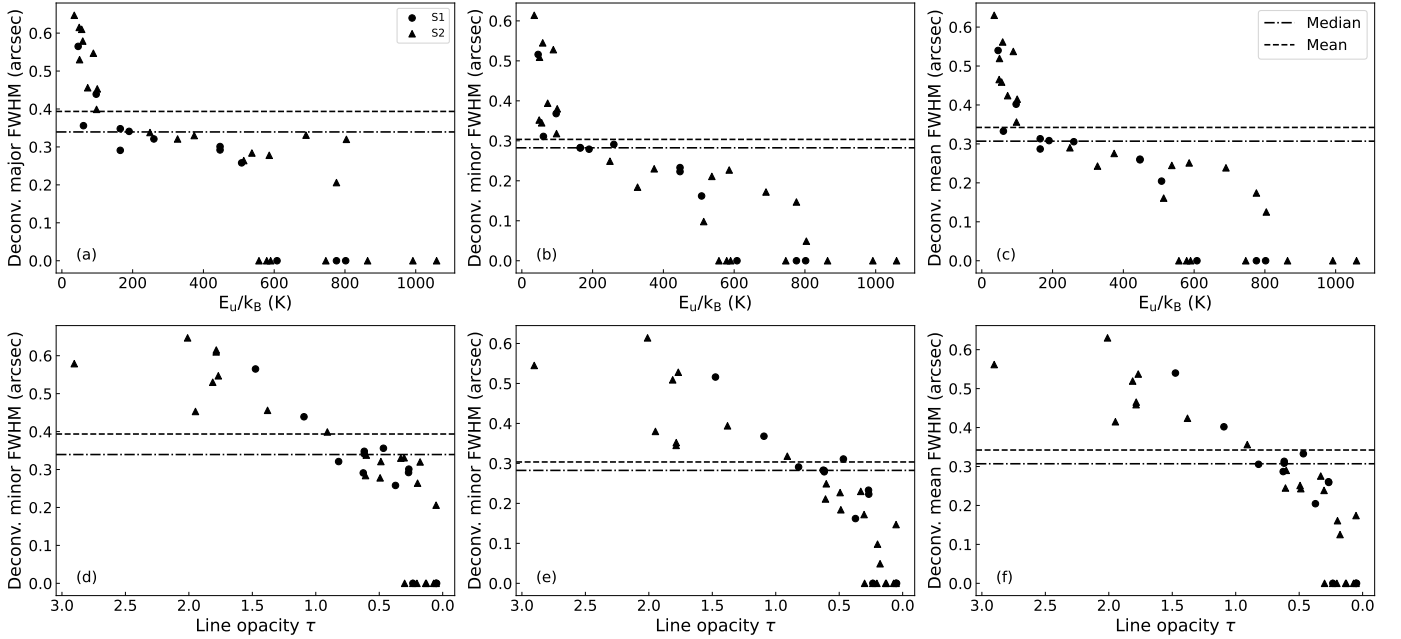


Fig. C.1: Deconvolved emission size of methanol transitions as a function of upper-level energy (top) and line opacity (bottom). The left (a and d), middle (b and e), and right (c and f) panels show the major, minor, and mean FWHM, respectively. The filled circle and triangle symbols indicate the spectral setups S1 and S2, respectively. The dashed line represents the mean value, while the dot-dashed line represents the median value.

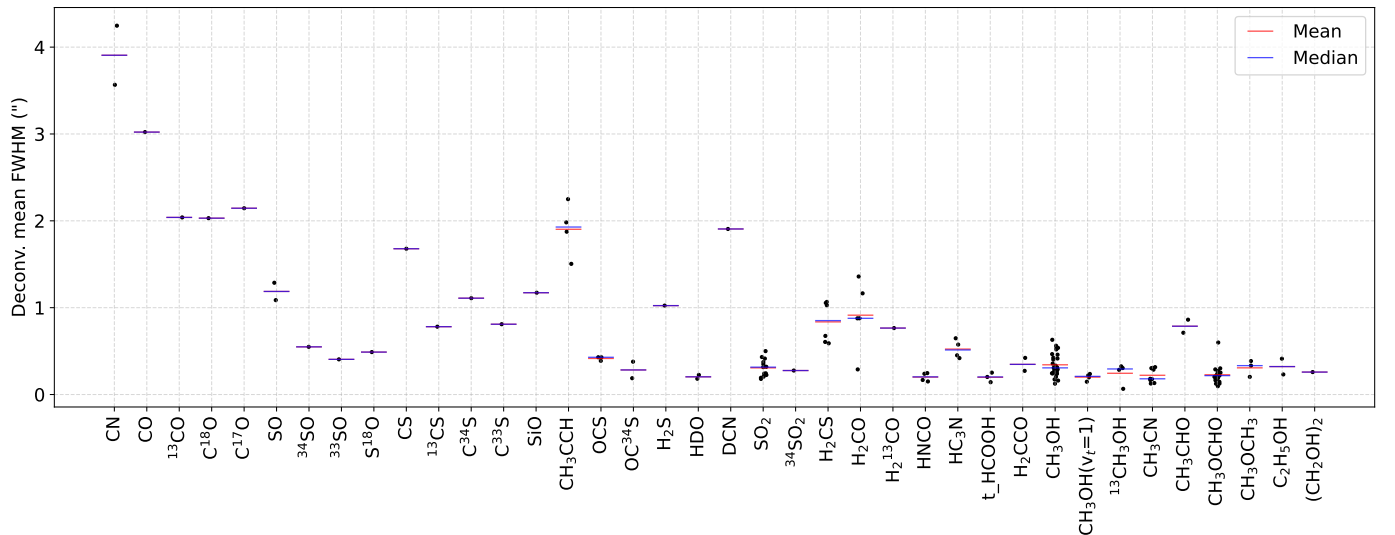


Fig. C.2: Deconvolved mean source size determined for all identified molecules through 2D Gaussian fitting. For each molecule, the deconvolved emission sizes of different transitions are shown as black dots, while the mean and median values are indicated by red and blue lines, respectively.

Appendix D: Observed spectrum of MM1

Figures D.1 and D.2 present the full spectral survey of G135.27 MM1. In Fig. D.1, the spectra are shown with a zoomed y-axis to highlight the weak COM lines, while Fig. D.2 shows the spectra un-zoomed. The LTE model parameters of the synthetic spectra overlaid on the observed spectra are listed in Table B.1.

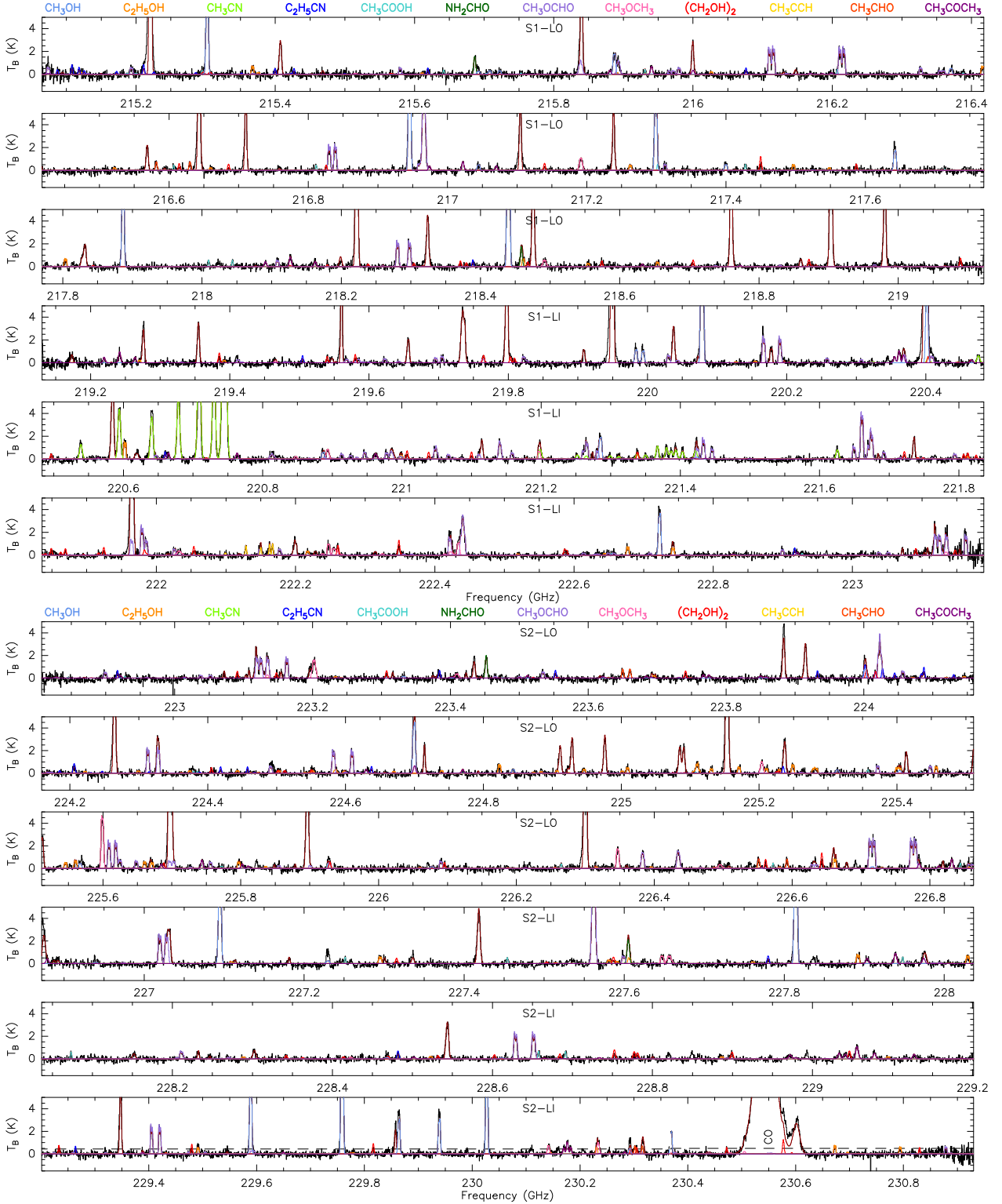


Fig. D.1: NOEMA spectrum of the outer Galaxy hot core G135.27 MM1. The observed spectrum and the synthetic spectrum that contains the contribution of all identified molecules are shown in black and dark red, respectively. The synthetic spectra of selected COMs are overlaid in various colors, as labeled at the top. The horizontal dashed line marks the 3σ noise level. The spectra have been truncated to emphasize the faint lines. The synthetic spectrum of S1-UI does not contain the CO contribution because the CO line profile is incomplete in this spectral window.

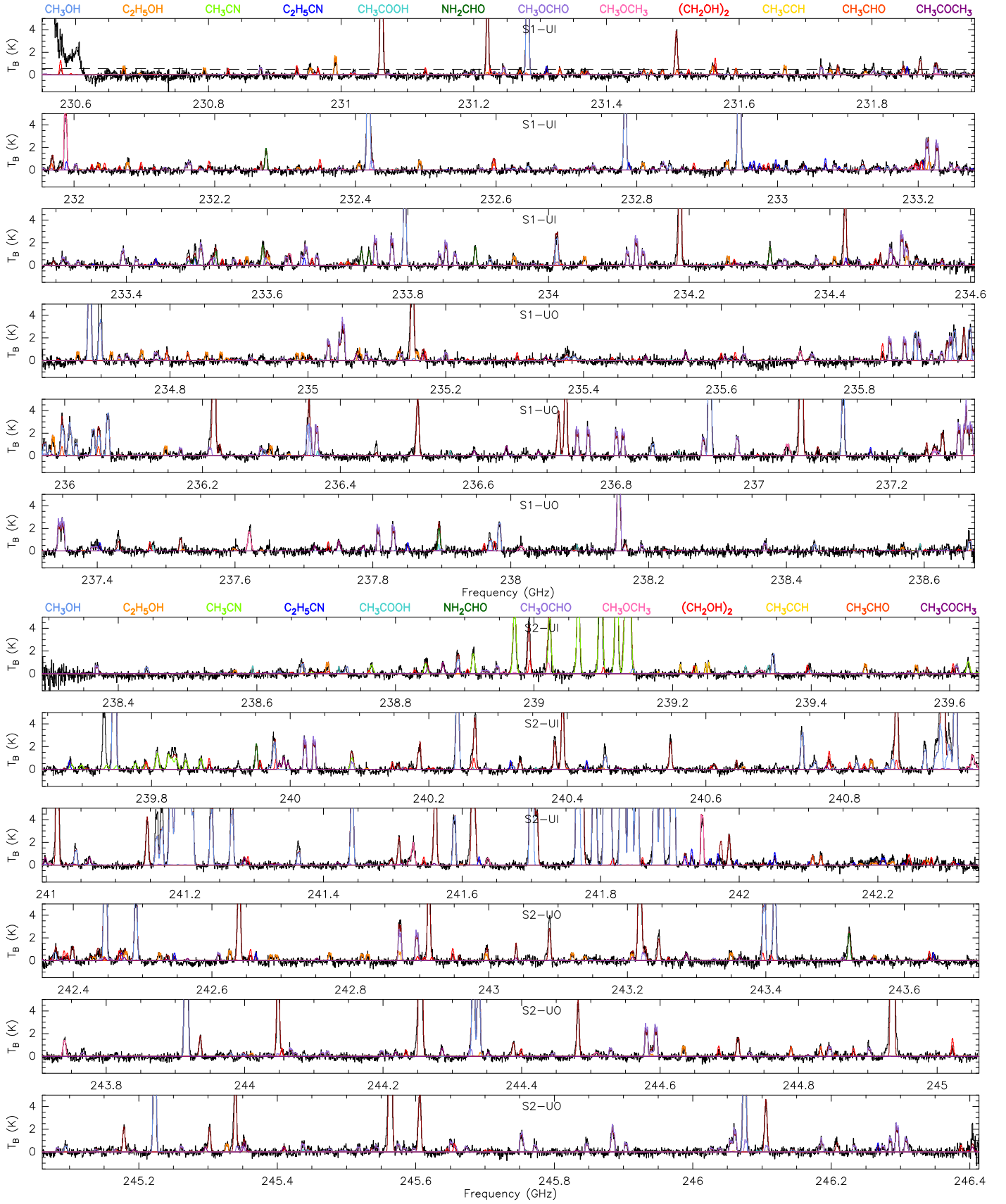


Fig. D.1: Continued.

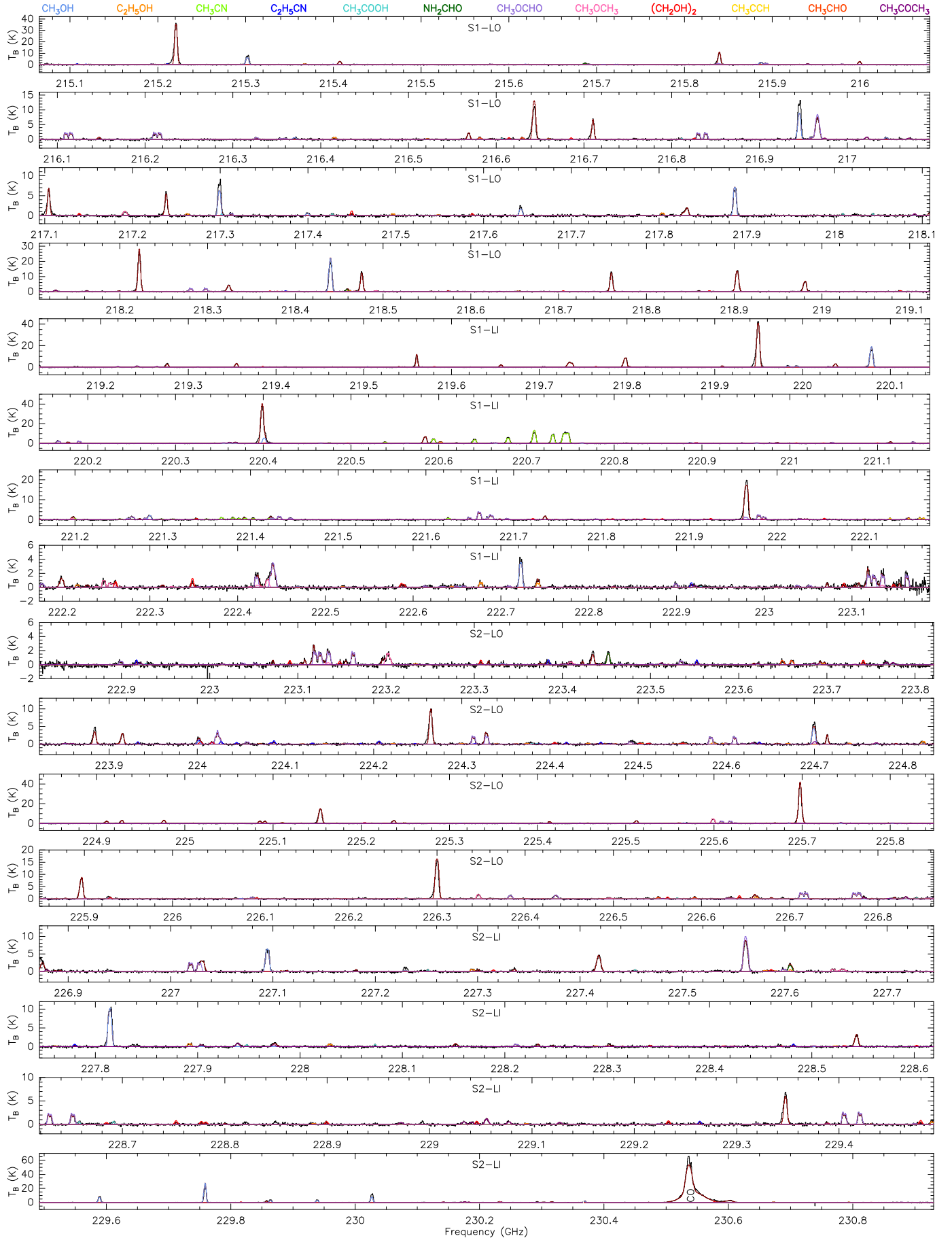


Fig. D.2: Same as Fig. D.1 but showing the un-zoomed spectrum.

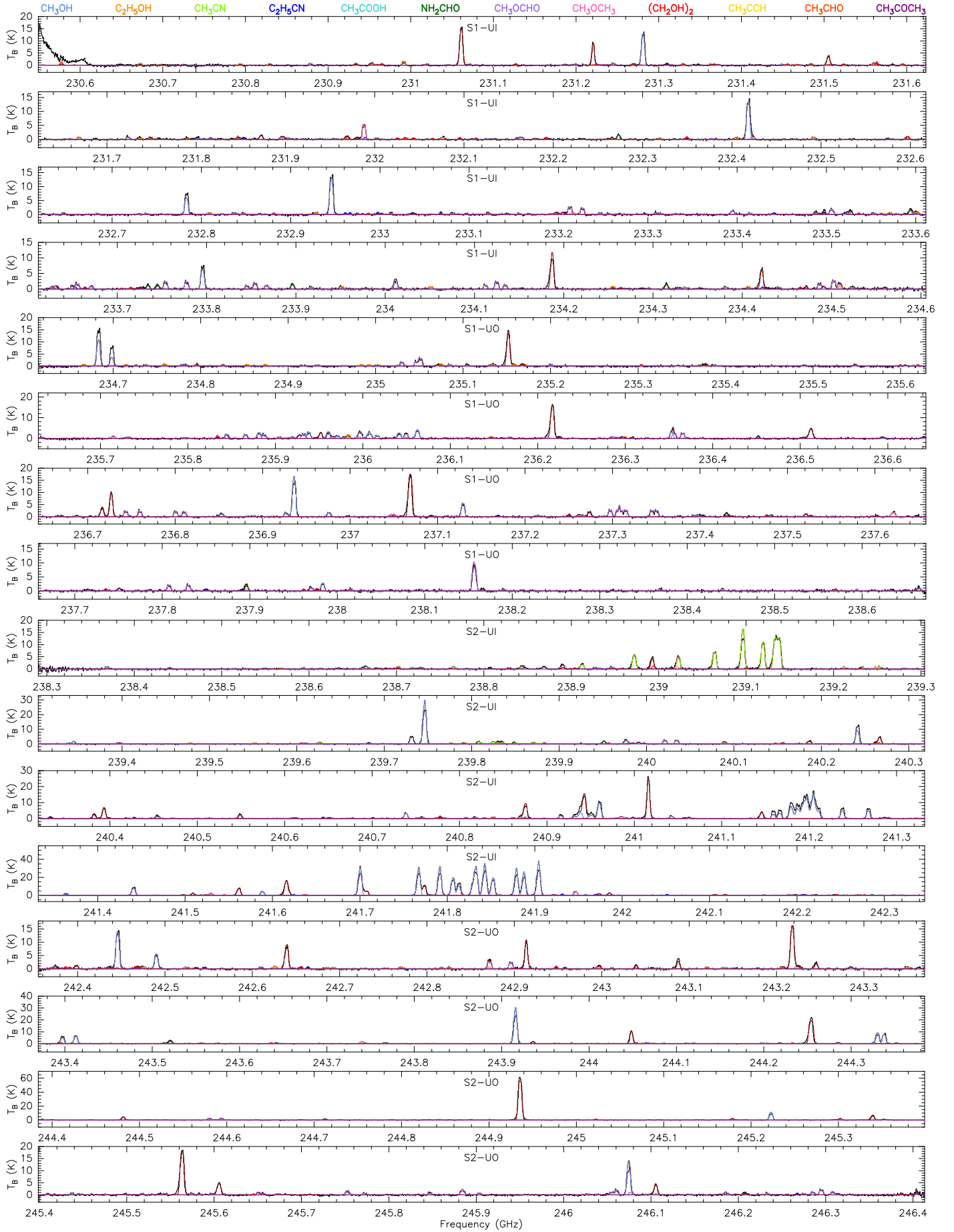


Fig. D.2: Continued.
A303, page 28 of 28