

Euclid: Early Release Observations – Interplay between dwarf galaxies and their globular clusters in the Perseus galaxy cluster[★]

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ABSTRACT

We present an analysis of globular clusters (GCs) of dwarf galaxies in the Perseus galaxy cluster that explores the relationship between dwarf galaxy properties and their GCs. Our focus is on GC numbers (N_{GC}) and GC half-number radii (R_{GC}) around dwarf galaxies, and their relations with host galaxy stellar masses (M_*), central surface brightnesses (μ_0), and effective radii (R_e). This work is unique due to its large sample size and the absence of pre-selection based on μ_0 and R_{GC} for dwarf galaxies. Interestingly, we find that at a given stellar mass, R_{GC} is almost independent of the host galaxy μ_0 and R_e , while R_{GC}/R_e depends on μ_0 and R_e . Lower surface brightness and diffuse dwarf galaxies show $R_{GC}/R_e \approx 1$, while higher surface brightness and compact dwarf galaxies show $R_{GC}/R_e \approx 1.5-2$. This means that for dwarf galaxies of similar stellar mass, the GCs have a similar median extent; however, their distribution is different from the field stars of their host. Additionally, low surface brightness and diffuse dwarf galaxies on average have a higher N_{GC} than high surface brightness and compact dwarf galaxies at any given stellar mass. We also find that ultra-diffuse galaxies (UDGs) and non-UDGs in the sample have a similar R_{GC} , while UDGs have a smaller R_{GC}/R_e (typically less than one) and a three to four times higher N_{GC} than non-UDGs. Furthermore, when examining nucleated versus non-nucleated dwarf galaxies, we found that for $M_* > 10^8 M_\odot$, nucleated dwarf galaxies seem to have a smaller R_{GC} and R_{GC}/R_e , with no significant differences seen between their N_{GC} except at $M_* < 10^8 M_\odot$, where the nucleated dwarf galaxies tend to have a higher N_{GC} . Lastly, we explored the stellar-to-halo mass ratio (SHMR) of dwarf galaxies (halo mass based on N_{GC}) and conclude that the Perseus cluster dwarf galaxies follow the expected SHMR at $z = 0$ extrapolated down to $M_* = 10^6 M_\odot$.

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1. Introduction

Current galaxy formation and evolution models do not fully reproduce the properties of dwarf galaxies as observed today (e.g. Sales et al. 2020). Dwarf galaxies are low-mass dark matter-dominated objects (McConnachie 2012; Eftekhari et al. 2022), and their evolution and morphological appearance are driven by various internal and external (environmental) processes as well as the physics of dark matter, all of which act simultaneously on these systems. Depending on their initial conditions at the time of formation, dwarf galaxies can follow very diverse evolutionary paths, resulting in the diverse population of dwarf galaxies observed today. This diversity is particularly pronounced in high-density environments, such as those in galaxy clusters. Such environments appear to play a significant role in the formation of extreme dwarf galaxies, namely ultra-compact dwarf galaxies (UCDs; Hilker et al. 1999; Drinkwater et al. 2000; Mieske et al. 2013; Liu et al. 2020) and ultra-diffuse dwarf galaxies (UDGs; van Dokkum et al. 2015; Marleau et al. 2021; Iodice et al. 2023), with effective radii varying by a factor of 100 and spanning the range of a few tens of parsecs to several kiloparsecs. UDGs are dwarf galaxies with a low central surface (μ_0) brightness and large effective radius (R_e), defined as galaxies with $\mu_{0,g} > 24$ mag arcsec⁻² and $R_e > 1.5$ kpc (van Dokkum et al. 2015). They are not a stand-alone category but a subset of low surface brightness dwarf (elliptical and spheroidal) galaxies (Conselice 2018; Canossa-Gosteinski et al. 2024).

In addition to the observed diversity in galaxy size and light distribution, the globular cluster (GC) populations of dwarf galaxies also exhibit considerable diversity. Previous studies of GCs in dwarf galaxies have not revealed significant abnormalities in GC properties (Georgiev et al. 2009; Liu et al. 2019; Prole et al. 2019; Jones et al. 2023). However, some studies indicate that a small fraction of the UDG population hosts more GCs than dwarf galaxies of comparable luminosity and stellar mass (Lim et al. 2018, 2020; Gannon et al. 2022). The significance of this observed excess of GCs remains uncertain due to inconsistencies in GC number counts (N_{GC}) reported in the literature, particularly for DF44 and DFX1 in the Coma cluster (van Dokkum et al. 2017; Saifollahi et al. 2022) and MATLAS-2019, also known as NGC5648-UDG1 (Müller et al. 2021; Danieli et al. 2022). In addition to the debate surrounding N_{GC} , some studies have aimed to quantify the radial profiles of GCs around dwarf galaxies, including UDGs, by estimating the ratio between the half-number radius of the GC population (R_{GC} , the distance from the host galaxy within which half of the total number of GCs are found) and the host galaxy’s effective radius (R_e), expressed as R_{GC}/R_e . Compared with those focused on GC number counts, fewer studies have examined the constraints on the GC radial distribution. Overall, R_{GC}/R_e appears to span a range of 0.7–1.5 (Lim et al. 2018; Carleton et al. 2021; Saifollahi et al. 2022; Román et al. 2023; Marleau et al. 2024; Janssens et al. 2024; Li et al. 2025)¹. These ranges of N_{GC} and R_{GC}/R_e suggest diversity among the GC properties of dwarf galaxies. However, since most recent studies have concentrated on UDGs, a comprehensive examination of GC properties as a function of host

galaxy characteristics and environmental factors remains warranted. Given that dwarf galaxies typically host fewer GCs than massive ones, substantial samples of such objects are essential to improving the statistics for an adequate analysis and thus highlight the need for extensive observational efforts.

At distances closer than 5 Mpc, GCs are bright and large enough in angular size to be spatially resolved from the ground using medium class telescopes (2–4 m); otherwise, most are barely or not resolved and tend to appear similar to the foreground stars and background galaxies. Combining deep optical data in the u band with near-infrared data in the K band helps distinguish GCs from non-GCs (foreground stars and background galaxies) due to their distinct spectral energy distribution (SED) in this wide colour range. This method has proven to be effective in detecting GCs out to a 20 Mpc distance, such as in the Virgo and Fornax galaxy clusters, which are respectively at 16 Mpc and 20 Mpc (Muñoz et al. 2014; Powalka et al. 2016; Cantiello et al. 2018; Saifollahi et al. 2021a). Observations at high spatial resolution using the *Hubble* Space Telescope (HST) have been an invaluable alternative for studying GCs, both in nearby systems within 10 Mpc (e.g. Sharina et al. 2005; Georgiev et al. 2009) and in the more remote Virgo and Fornax clusters (Jordán et al. 2004, 2015). At the distances of the latter cluster, GCs with a typical half-light radii of 2–3 pc can appear slightly resolved (larger than point sources), making them marginally distinguishable from unresolved foreground stars and compact background galaxies. HST has also played a crucial role in investigating GCs in more distant galaxy clusters, such as the Perseus cluster at 72 Mpc (Harris et al. 2020) and the Coma cluster at 100 Mpc (Amorisco et al. 2018). Although GCs appear as point sources in HST images at these distances, the high spatial resolution of HST enables the rejection of background galaxies. The Coma cluster currently represents the distance limit of our observational capabilities for studying GCs in dwarf galaxies. Although it is possible to extend the studies further with longer exposures, this approach is inefficient in terms of telescope time. Another caveat is that the small field of view of HST cameras makes it impractical to collect data on large samples of dwarf galaxies. These limitations also exist for the *James Webb* Space Telescope (JWST) even though the deep JWST/NIRCAM images would detect GCs at larger distances, up to $z = 0.3$ (Harris & Reina-Campos 2023; Berkheimer et al. 2024).

To address these observational limitations, a wide-field survey of a significant portion of the extragalactic sky at a space-based spatial resolution would be beneficial. The recently launched ESA *Euclid* mission offers such capabilities. *Euclid* is particularly powerful for studying dwarf galaxies (Euclid Collaboration: Mellier et al. 2025) and is expected to identify over 10 000 nearby dwarf galaxies during its six-year survey. The deep, high-resolution images provided by *Euclid* in optical (VIS; Euclid Collaboration: Cropper et al. 2025) combined with near-infrared (NIR/NISP; Euclid Collaboration: Jahnke et al. 2025) and auxiliary ground-based data in u , g , r , i , and z will enable the identification of GCs in dwarf galaxies within the Local Universe. According to Euclid Collaboration: Voggel et al. (2025), for galaxies at distances smaller than 50 Mpc, Euclid Wide Survey (EWS) I_e images will detect GCs brighter than the turn-over magnitude (TOM) of the GC luminosity function (GCLF). The GCLF TOM is the peak magnitude of the (nearly) symmetric Gaussian distribution of magnitudes of

¹ Lim et al. (2018) reports $R_{GC}/R_e = 1.5$ for dwarf galaxies based on an analysis of samples from ACS Virgo Cluster Surveys (ACSVCS, Côté et al. 2004), though the paper does not provide further details on this analysis.

GCs around galaxies (Secker & Harris 1993; Rejkuba 2012), and therefore, detecting GCs brighter than this magnitude corresponds to detecting about half of the total GCs around galaxies. At 20 Mpc, GC detection in I_E with the Euclid Wide Survey is nearly complete. This is also shown in the *Euclid* Early Release Observations (ERO) of the Fornax galaxy cluster (Saifollahi et al. 2025). The completeness of GC detection drops when including detections in the NISP bands, and therefore, the overall completeness of GC selection depends on the adopted methodology for a colour-based GC selection.

Given the forthcoming data from the *Euclid* mission, it is timely to start investigating the GC properties of dwarf galaxies in detail as a function of their host galaxies and the environment on a larger scale. The uniform character of *Euclid* imaging data with a stable point spread function (PSF) facilitates a homogeneous study of GCs across the observed sky. This is particularly important because the current literature often adopts different methodologies and assumptions when studying GCs. Moreover, previous studies have largely focused on UDGs, resulting in a lack of comprehensive analyses of unbiased samples of dwarf galaxies. This uncertainty in GC properties should be addressed in the coming years using data from *Euclid* and other wide-field deep extragalactic surveys such as the Legacy Survey of Space and Time (LSST) with the Rubin Observatory, the *Nancy Grace Roman* Space Telescope’s High Latitude Wide Area Survey (Usher et al. 2023; Dage et al. 2023), and the Chinese Space Station Telescope (Qu et al. 2023).

The *Euclid* ERO programme (Cuillandre et al. 2025a) included a pointing towards the Perseus galaxy cluster (Cuillandre et al. 2025b; Kluge et al. 2025). This galaxy cluster, with a mass of $10^{15} M_\odot$ (Simionescu et al. 2012; Aguerri et al. 2020), is similar in mass but closer than the Coma cluster (72 Mpc versus 100 Mpc) and represents one of the densest environments in the local cosmic web. The ERO data for the Perseus cluster are deeper than the standard EWS image stacks and reach the GCLF TOM. With more than 1000 dwarf galaxies identified within this dataset (Marleau et al. 2025), this field is a unique and valuable resource for studies of the GCs of dwarf galaxies. Marleau et al. (2025) conducted a first analysis of the GC counts in that dataset.

In this paper, using a different technique and source catalogue, we present a detailed analysis and study of GCs in dwarf galaxies within the Perseus cluster, and we examine the differences between the GC number and the GC radial distribution in different categories of dwarf galaxies in terms of surface brightness and effective radius. Section 2 describes the *Euclid* data, while Sect. 3 outlines the methodology used to analyse the data and search for GCs. Section 4 examines the properties of the GCs and explores trends between dwarf galaxies and their GCs. Our findings are discussed in Sect. 5, and Sect. 6 presents a summary of our results.

2. Data

2.1. Imaging data

The *Euclid* ERO images (Euclid Early Release Observations 2024) of the Perseus galaxy cluster (Cuillandre et al. 2025b) cover 0.6 deg^2 of the Perseus cluster in I_E , Y_E , J_E , and H_E , with pixel sizes $0''.1$ for I_E and $0''.3$ otherwise. The data and their processing are described in detail in Cuillandre et al. (2025a). The ERO dataset for Perseus consists of images from 16 exposures in each filter, which amounts to four times the reference observation sequence that will be executed on each field of the EWS.

As a result, the stacked VIS images allowed us to reach slightly deeper than the TOM of the GCLF at the 72 Mpc distance of the Perseus cluster. The canonical TOM is expected at $I_E = 26.3$ (Marleau et al. 2025). In Sect. 3.3 we assess the depth and completeness of this dataset in I_E . The VIS data were already used to study the GC numbers of dwarf galaxies (Marleau et al. 2025) as well as intracluster GCs (Kluge et al. 2025). Here we take advantage of the near-infrared (NIR) images, and adopt a slightly different approach in the selection and analysis of GCs.

The Perseus cluster has been observed in various HST programmes, including the PIPER survey (Harris et al. 2020) which covers about 40 low surface brightness dwarf galaxies (of which some are also categorised as UDGs by definition). Figure 1 compares the I_E , Y_E , J_E , and H_E images of a dwarf galaxy with HST/ACS observations of that survey in filter F814W, as well as with ground-based CFHT/MegaCAM data in the r band. Figure 2 shows a zoomed view of the central part of the dwarf galaxy taken by *Euclid*/VIS in I_E and HST/ACS in F814W. *Euclid* ERO and HST images are comparable in depth, while HST images have a tighter PSF and better PSF sampling. The *Euclid* ERO data have such quality for more than 1000 dwarf galaxies in the Perseus cluster.

2.2. Dwarf galaxy catalogue

Using the ERO data for the Perseus cluster, Marleau et al. (2025) have identified more than 1000 dwarf galaxies (with a majority to be early-type) in a stellar mass range of 10^5 – $10^{9.5} M_\odot$, with surface brightnesses between 18 and 26 mag arcsec $^{-2}$ and effective radii between 0.5 and 6 kpc in I_E . Among many structural parameters, the Perseus cluster dwarf catalogue (Marleau et al. 2025) contains the effective radius (R_e), the central surface brightness (μ_0), and the stellar mass (M_*) of Perseus cluster dwarf galaxies. These parameters are used frequently for the analysis in this paper. We examined all the dwarf galaxies and their parameters, and excluded from the analysis the dwarf galaxies whose parameters seem unreliable (e.g. affected by a bright star). The final sample contains about 870 dwarf galaxies. All are dwarf early types, and 86 are also categorised as UDGs (defined as dwarf galaxies with a central surface brightness of $\mu_{0, I_E} > 23 \text{ mag arcsec}^{-2}$ and an effective radius of $R_e > 1.5 \text{ kpc}$).

Overall, this dataset contains a diverse population of early-type dwarf galaxies in terms of stellar mass, surface brightness, and effective radius. The absence of strong bias in size or surface brightness, when compared for instance to samples of UDGs, makes it ideal for studying the GC properties of early-type dwarf galaxies in a global way. The results presented in this work represent GCs of dwarf galaxies in dense cluster-like environments considering that all the dwarf galaxies are highly likely members of the Perseus cluster.

3. Methods

This section outlines the methodology for source detection, photometry, and identification of GC candidates around dwarf galaxies. Our approach is similar to that of Saifollahi et al. (2025), with some differences in GC selection criteria and enhancements to address flux-dependent uncertainties in colour and the source’s ellipticity for GC selection. The pipeline used for this analysis, GCEx, is publicly available². GCEx uses several astronomical packages and Python libraries and is developed specifically to analyse images and identify GCs in wide-field

² <https://github.com/teymursaif/GCEx>

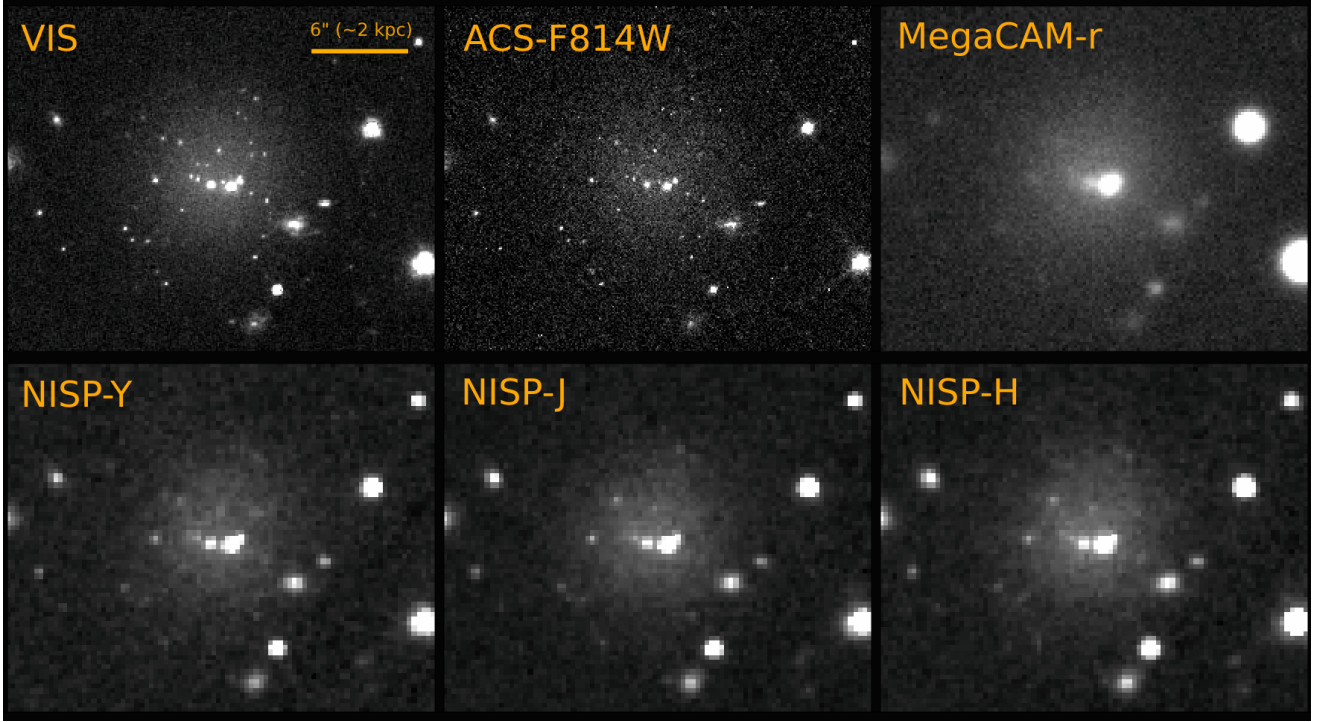


Fig. 1. View of a Perseus cluster dwarf galaxy (EDwC-0120 in [Marleau et al. 2025](#), and R84 in [Janssens et al. 2024](#)) as seen by the VIS and NISP instruments on board *Euclid* (in I_E , Y_E , J_E , and H_E) by HST/ACS (in F814W) and by CFHT/MegaCAM (in r). The galaxy seems to host several GCs, which appear as point sources on top of the galaxy, as well as a NSC in the centre of the diffuse component. The brightest object projected near the centre of the galaxy could be a foreground star or a massive GC destined to merge with the already existing NSC.

multi-wavelength data for large samples of galaxies. The methodology described below refers to the methodology adopted in GCEx.

3.1. Point spread function

Empirical PSF models are necessary in several instances throughout this analysis. Therefore, in the first step, we produced PSF models in each filter. The *Euclid* PSF is known to vary with the location of the image in the focal plane and the colour of the sources; however, such variations are minimal ([Cuillandre et al. 2025a](#); [Euclid Collaboration: McCracken et al. 2025](#)). For this analysis, we ignore these variations and use a single PSF model per filter. Here, PSF models are created by stacking 1000 cutouts of bright, non-saturated stars with magnitudes between 19 and 21 for I_E and 17 and 19 for Y_E , J_E , and H_E . Cutout sizes are 10 arcsec in all the bands. The cutouts are oversampled by a factor of ten using SWarp ([Bertin et al. 2002](#)) with the LANCZOS3 interpolation kernel, resulting in a PSF model with a pixel size ten times smaller than the instrumental pixel size. Once the cutouts are created, the centroid is estimated, and the cutouts are aligned and stacked using SWarp. The final PSF models have a full-width half-maximum (FWHM) of $0''.18$, $0''.52$, $0''.54$, and $0''.58$ in I_E , Y_E , J_E , and H_E , respectively. These FWHM values are larger than the FWHM of single exposures before stacking ($0''.16$ for I_E) mainly due to interpolation effects during resampling to create the final stacked frames ([Cuillandre et al. 2025a](#)). Our analysis of the PSF indicates that resampling primarily affects the very central region of the PSF model (within 0.1 arcsec).

3.2. Source detection and photometry

Source detection was carried out on $240'' \times 240''$ (85×85 kpc) I_E cutouts of dwarf galaxies. For this, initially we per-

formed unsharp-masking of the cutouts using a small background mesh size (BACKSIZE=8) to subtract the extended light of galaxies, thus improving the detection of fainter sources and GC candidates around galaxies. Then, we used SExtractor for source detection. Once sources were detected, we applied an initial filtering based on an estimate of their magnitudes (SExtractor output parameter MAG_AUTO), their FWHM (FWHM_IMAGE), and their ellipticity (ELLIPTICITY) to remove bright and saturated sources (brighter than 19 in I_E), very faint sources (fainter than 28.5 in I_E), extended sources (FWHM_IMAGE > 20 pixels), very compact sources (FWHM_IMAGE < 1 pixel), and highly elongated sources (ELLIPTICITY > 0.9). This filtering effectively cleans the catalogue of unwanted sources, such as large objects and artefacts, and it reduces the size of the source catalogue for the next steps of the analysis.

After the initial filtering, close to 500 000 sources remain in the detection catalogue. For these detections, we then perform aperture photometry using the photutils library in Python, employing an aperture with a diameter of 1.5 times the average FWHM of the data in the corresponding band (single value across the data). This aperture diameter corresponds to $0''.27$, $0''.78$, $0''.81$, and $0''.87$ in I_E , Y_E , J_E , and H_E , respectively. The aperture photometry values are corrected for the fraction of the missing flux beyond the aperture using the previously created PSF models. This fraction is about 20% and 10%, respectively, in the VIS and NISP bands, estimated from the PSF models out to a radius of $5''$. The fraction of the light beyond $5''$ is less than 1% ([Cuillandre et al. 2025a](#)). Furthermore, the PSF FWHM varies between 1.5 and 1.9 pixels in the field-of-view. Our analysis of the colours of bright point sources shows that the effect of this variation on the photometry and the measured colours is negligible.

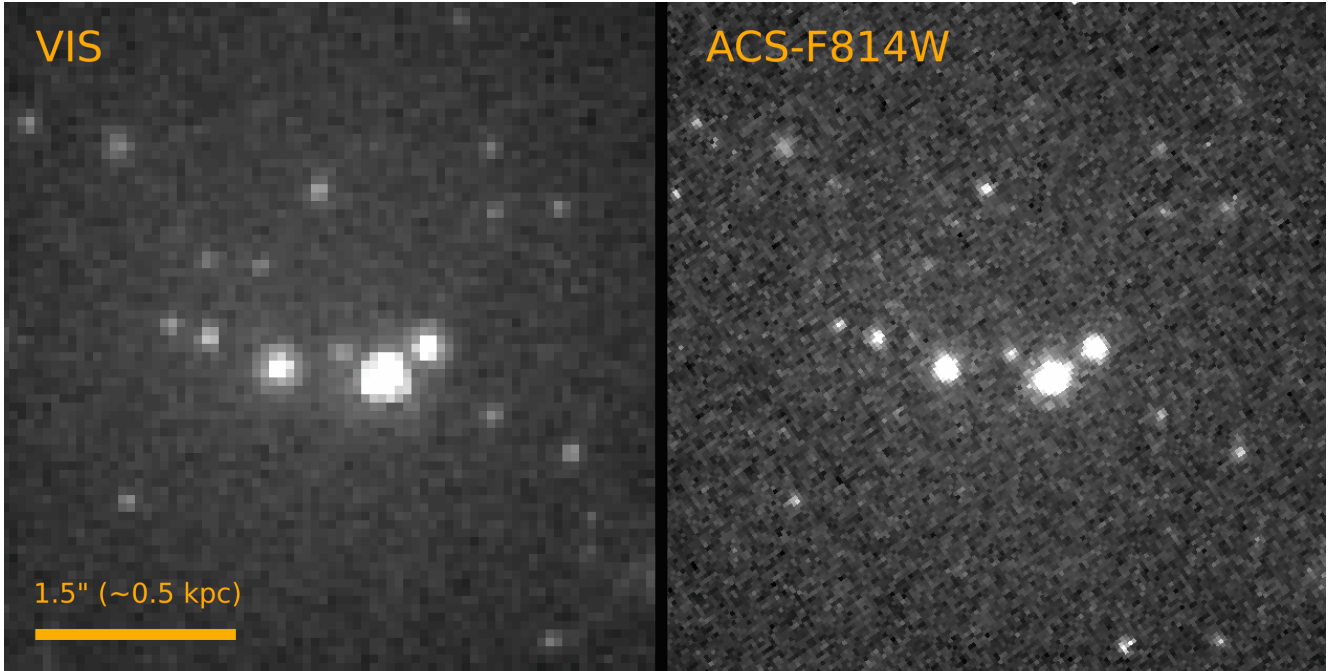


Fig. 2. Zoomed view of the dwarf galaxy in Fig. 1 as seen by *Euclid*/VIS in I_E and HST/ACS in F814W (PIPER survey, Harris et al. 2020). The total integration times for these images are 8960 s and 2100 s respectively. The depths of both datasets are comparable, while HST/ACS provides a higher spatial resolution with a better sampled PSF than *Euclid*/VIS.

At the 72 Mpc distance of the Perseus cluster, and at the *Euclid* spatial resolution, GCs appear as point sources. In this case, aperture photometry as described here properly measures the total flux of GCs, while it underestimates the total flux of more extended sources (more extended than the PSF). However, this is not a concern because these extended sources will be removed during the GC identification step. Furthermore, we measure the compactness index of the sources, C_{2-4} , defined as the difference between the magnitudes measured within aperture diameters of 2 and 4 pixels. This compactness index is used in subsequent steps to identify point sources in the data, which include the GCs as well as some faint foreground (halo) stars (see Fig. C.2 in Cuillandre et al. 2025a).

3.3. Artificial GCs and detection completeness

We perform artificial GC injection into the data to assess the completeness of source detection and then GC selection. We generate 500 artificial GCs per dwarf galaxy, randomly (uniformly) distributed within their cutout in all the bands. The artificial GCs are created using a King profile (King 1966) with a half-light radius between 1 and 5 pc and an absolute I_E between -12 and -5 , covering the full magnitude range of GCs, convolved with the PSF model of a given filter (over-sampled by a factor of 10). Before injection into the data, we re-bin the artificial GCs to the pixel scale of the images with a random sub-pixel shift in the position of their centre using SWarp. We assume an average colour of 0.45 for $I_E - Y_E$, and 0 for $Y_E - J_E$, and $J_E - H_E$. These average values are chosen considering the average GC colour in the ERO observations of the Fornax cluster (Saifollahi et al. 2025) as well as synthetic photometry of single stellar population (SSP) models for old and metal-poor GCs (see Appendix A in Saifollahi et al. 2025), considering that the GCs of dwarf galaxies are mostly metal-poor (Forbes et al. 2024). Finally, we add noise to the artificial GCs before inject-

ing them into the data. The added noise is three times larger than the pixel Poisson noise to take into account the correlated noise in the data, which we have estimated to be about 2–3 times the Poisson noise.

On the images thus modified, we carried out source detection and photometry as described above. Then, we examine the completeness of the source detection in I_E . The result is presented in Fig. 3. GC detection is almost complete to $I_E = 26$, and starts to drop after this magnitude, with a completeness of 80% and 50% achieved respectively at $I_E = 26.7$ and $I_E = 27.2$. These values were estimated by fitting a modified version of the interpolation function of Fleming et al. (1995). Before accounting for the Galactic foreground extinction along the line of sight, the canonical TOM of the GCLF is expected at 26.3, while for dwarf galaxies it will be 0.3 mag fainter, at 26.6 (Liu et al. 2019). Considering that the extinction, A_{I_E} , across the field-of-view ranges from 0.1 to 0.3 mag in I_E (Marleau et al. 2025; Kluge et al. 2025) with an average value of about 0.2 mag, we expect $I_E \approx 26.8$ for the GCLF TOM of dwarf galaxies at the distance of the Perseus cluster (72 Mpc), about the magnitude corresponding to 80% detection completeness. This assessment is based on the average extinction correction, and in practice, this correction is different for every dwarf galaxy. We took these differences into account when analysing the GC properties of dwarf galaxies.

This assessment shows that overall, we detect most GCs brighter than the GCLF TOM. The fraction of missing GCs depends on the exact shape of the GCLF of dwarf galaxies. Assuming a Gaussian GCLF with TOM $I_E = 26.6$ and width $\sigma = 1$ mag (Villegas et al. 2010), and an average extinction of 0.2 mag, we estimate that 90% of the GCs are detected. This fraction changes to 85% if we assume a narrower GCLF for dwarf galaxies with $\sigma = 0.5$.

Given the outcome of the completeness assessment, the dataset allowed us to study GC numbers (albeit with some assumptions on the shape of the GCLF, described later in

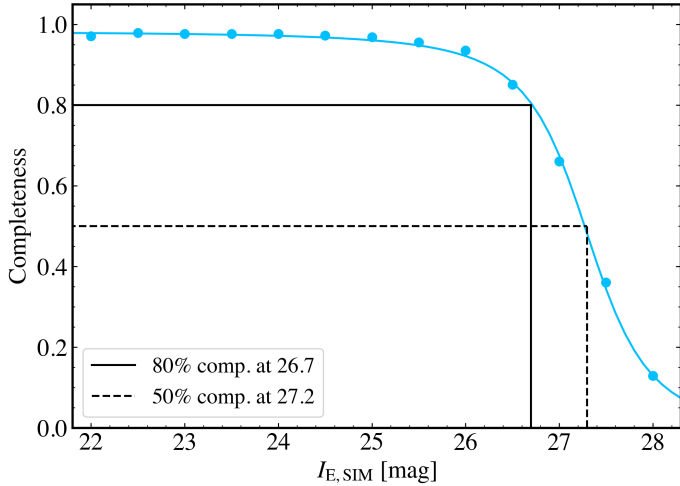


Fig. 3. Completeness of the performed source detection in I_E based on artificial GC tests. Our assessment shows that the GC detection is 80% complete at $I_E = 26.7$ (without considering the extinction). This magnitude is roughly the expected location of GCLF TOM for dwarf galaxies at the distance of the Perseus cluster (72 Mpc) (when taking into account the average Galactic foreground dust extinction $A_{I_E} = 0.2$ mag).

Sect. 4.3) as well as radial profiles of GC populations. However, deeper data are required to study the shape of the GCLF of dwarf galaxies in detail, for example its behaviour as a function of the stellar mass of galaxies. Therefore, in this paper, we do not make an assessment of the GCLF. Such studies can be done with *Euclid* with dwarf galaxies at distances below 20 Mpc, where we measure the entire GCLF in I_E .

Considering the shallower depth (about 2 mag) and the lower detection completeness of the near-infrared images (Y_E , J_E , and H_E) compared to I_E (Fig. B.1), in this work we do not limit our detection catalogue by requesting a measurement in any of those bands but instead consider every object detected in I_E . However, if an object is detected in one of the three near-infrared filters, we use the flux and estimated magnitude to measure its colour and apply flux-dependent colour cuts. This procedure is described in Sect. 3.4.

3.4. GC (candidate) selection around dwarf galaxies

We used the artificial GCs injected into the data (all filters) to examine the expected range of compactness, colour, and apparent ellipticity of real GCs as a function of their magnitude and to define the GC selection criteria. Those sources that are selected through this procedure are considered as GC candidates and are used in this work to study the GC properties of dwarf galaxies. The GC (candidate) selection procedure is demonstrated in Fig. 4, where the top, middle, and bottom panels show the selection based on the compactness index (C_{2-4}), ellipticity (ε), and colour (for example, $I_E - Y_E$). Sources with I_E magnitude between 22 and 27 that meet all these criteria are selected as GC candidates. The cuts in I_E correspond to the brightest magnitude at which we expect to find typical GCs ($I_E = 22$) and the magnitude at which the completeness of GC detection is about 50% ($I_E = 27.0$). For studying GC numbers and radial distributions, we only consider GCs brighter than the GCLF TOM, which, considering extinction, is always brighter than $I_E = 27$. Therefore, our analysis does not rely on GCs fainter than this magnitude.

The boundaries of GC selection in each parameter space are shown with red curves in Fig. 4. These boundaries are chosen

to include 98% of the artificial GCs at a given magnitude. We initially find an offset of about 0.1 mag between point sources and artificial GCs in C_{2-4} . This is due to the effects of resampling while using undersampled data to model the PSF. Therefore, we adjust the C_{2-4} values of the artificial GCs by 0.1 mag to the lower C_{2-4} values. The point of the experiment with artificial GCs is to model the range of parameters within which measurement noise, and systematic offsets, such as the one for C_{2-4} , remain unaffected by a constant shift of the compactness index values. Additionally, to include other (unexpected and random) effects in C_{2-4} of real GCs, we extend the selection boundaries (red curves) by 0.2 mag on both sides. We also extend the boundaries of colour criteria for GC selection (red curves in the bottom row of Fig. 4) to cover the full range of expected GC colours. This is necessary considering that the artificial GCs are produced using one value (average colour) of GCs. We extend the range of colour selection by 0.3 for $I_E - Y_E$ and 0.15 for $Y_E - J_E$ and $J_E - H_E$. For more details on the selection of these values, we refer the reader to Appendix A in Saifollahi et al. (2025).

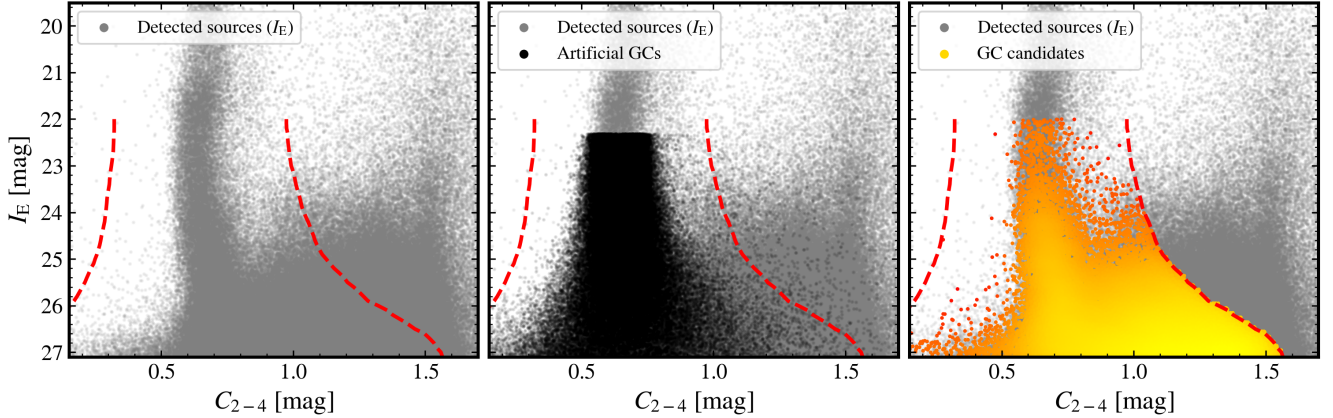
Figure 5 shows a few examples of dwarf galaxies and their identified GC candidates brighter than $I_E = 27$ (in red), as well as all other detected sources brighter than $I_E = 27$ that did not meet the GC selection criteria (in green). Figure 5 is divided into two parts: dwarf galaxies with high and low central surface brightness are shown at the top and bottom panels, respectively. Upon inspection, we see that source detection performs equally well across dwarf galaxies, GC identification can miss some GCs in central regions of high central surface brightness dwarf galaxies. This is due to residuals from unsharp masking (used to subtract the galaxy) in the centre, which can affect photometry and measurements regarding the most central GCs. As a consequence, these objects do not meet one or more GC selection criteria. This mostly limits GC detection in the central 500 pc of the brightest (and most massive) dwarf galaxies. We consider this point when studying the radial distribution of GCs and show that it has a negligible overall effect. However, this could lead to lower GC numbers for the dwarf galaxies with highest central surface brightness. Considering that these objects are the ones with most GCs, we expect that this bias in GC numbers is not significant, but must be taken into account when interpreting the results. In contrast, GC identification in low central surface brightness dwarf galaxies does not have any limitations in the centre of their host galaxies.

Before finalising the GC catalogues, for cutouts of dwarf galaxies, we produced a mask frame where nearby galaxies (both dwarf and massive) and bright stars are masked so that detections corresponding to these objects are ignored. Within some of the cutouts, edges of massive galaxies have also been masked. The motivation for such masking is to clean the background from GC over-densities around other galaxies and any false detections around bright stars. Such objects, if not masked, will artificially increase the source count around galaxies and in the background (areas far from dwarf galaxies). Additionally, given extinction corrections (A_{I_E}) in Marleau et al. (2025), we correct all the magnitudes in four filters for extinction. In Sect. 4 we use the masked and extinction-corrected GC catalogues to study the GC properties of dwarf galaxies.

4. Results

The *Euclid* ERO data of the Perseus cluster allowed us to reach the GCLF TOM expected for dwarf galaxies (at $I_E = 26.6$, without applying any extinction correction) with an overall

GC selection criteria for compactness index



GC selection criteria for ellipticity

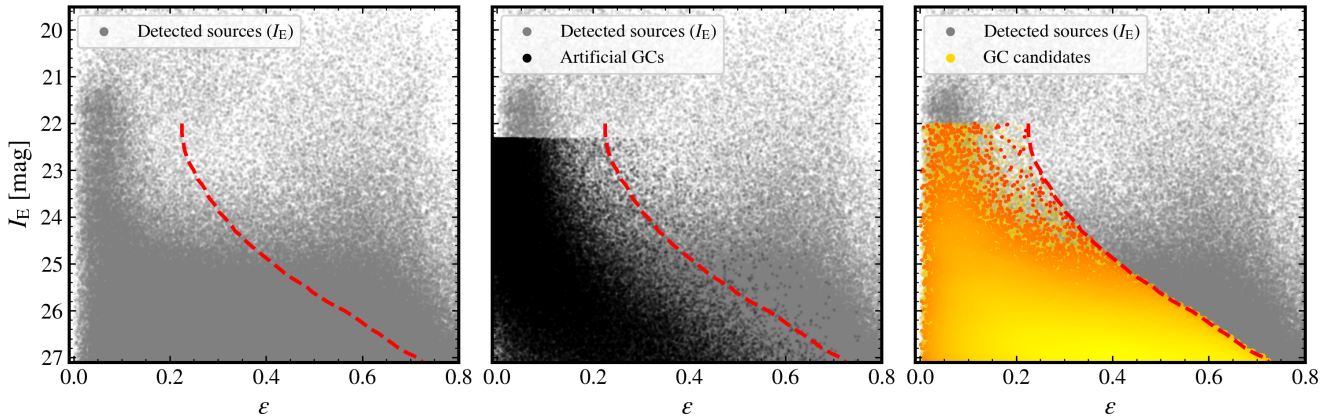
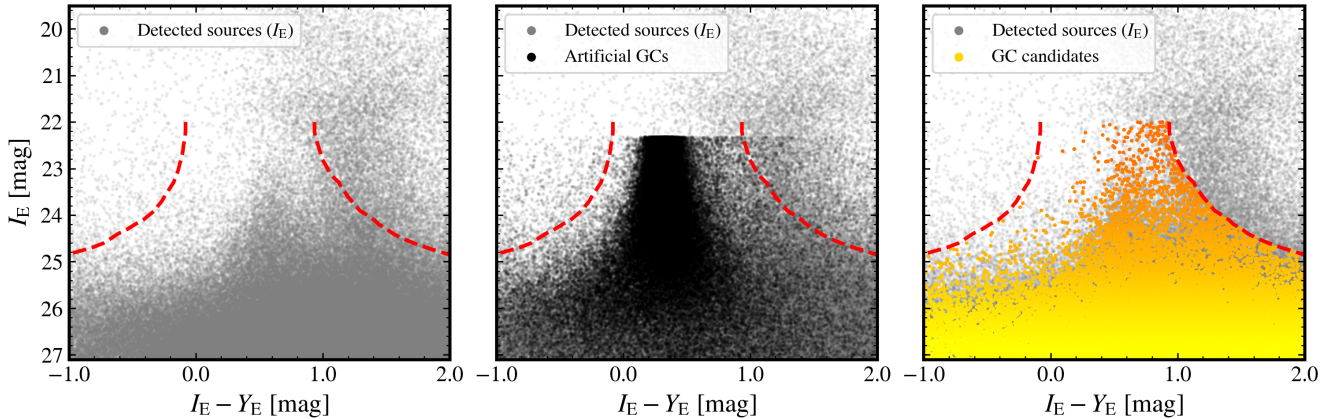
GC selection criteria for $I_E - Y_E$ colour

Fig. 4. Photometric criteria applied for GC identification when aiming for high completeness. The three rows demonstrate these criteria for compactness index (C_{2-4}), ellipticity (ϵ), and $I_E - Y_E$ colour (as an example). In each row, we show the real sky sources detected in *Euclid* I_E images (grey points), the detected artificial GCs (black points), and the real sky sources selected as GC candidates (yellow points). The dashed red curves represent the selection boundaries adopted based on the distribution of artificial GCs in each parameter space. In the top row, we see that the artificial GCs follow the vertical sequence of point sources (with an average $C_{2-4} = 0.7$), and the scatter becomes larger with fainter magnitude due to an increase in photometric uncertainties. In the bottom row, we see that the selection boundaries reject the majority of objects with $I_E < 24$ and $I_E - Y_E > 1$, which are foreground stars. However, a fraction of those stars end up in the final catalogue. This contamination is unavoidable considering the expected colour range of GCs. Tightening the selection criteria removes these contaminant stars, but leads to lower completeness.

estimated completeness of 85% down to the TOM. With such a dataset, we study the GC total numbers and their radial distribution after selecting GC candidates around dwarf galaxies. We begin by examining the GC radial distribution and the R_{GC} values around the Perseus sample. Using our results, we subsequently examine the total number of GCs, taking into account

the differences in R_{GC} between galaxies. These results³ are presented in Sects. 4.2 and 4.3. We study the average GC properties of the dwarf galaxies. Later, for the dwarf galaxies with

³ The output of the analysis and associated catalogues are available on the Strasbourg astronomical Data Center (CDS) database.

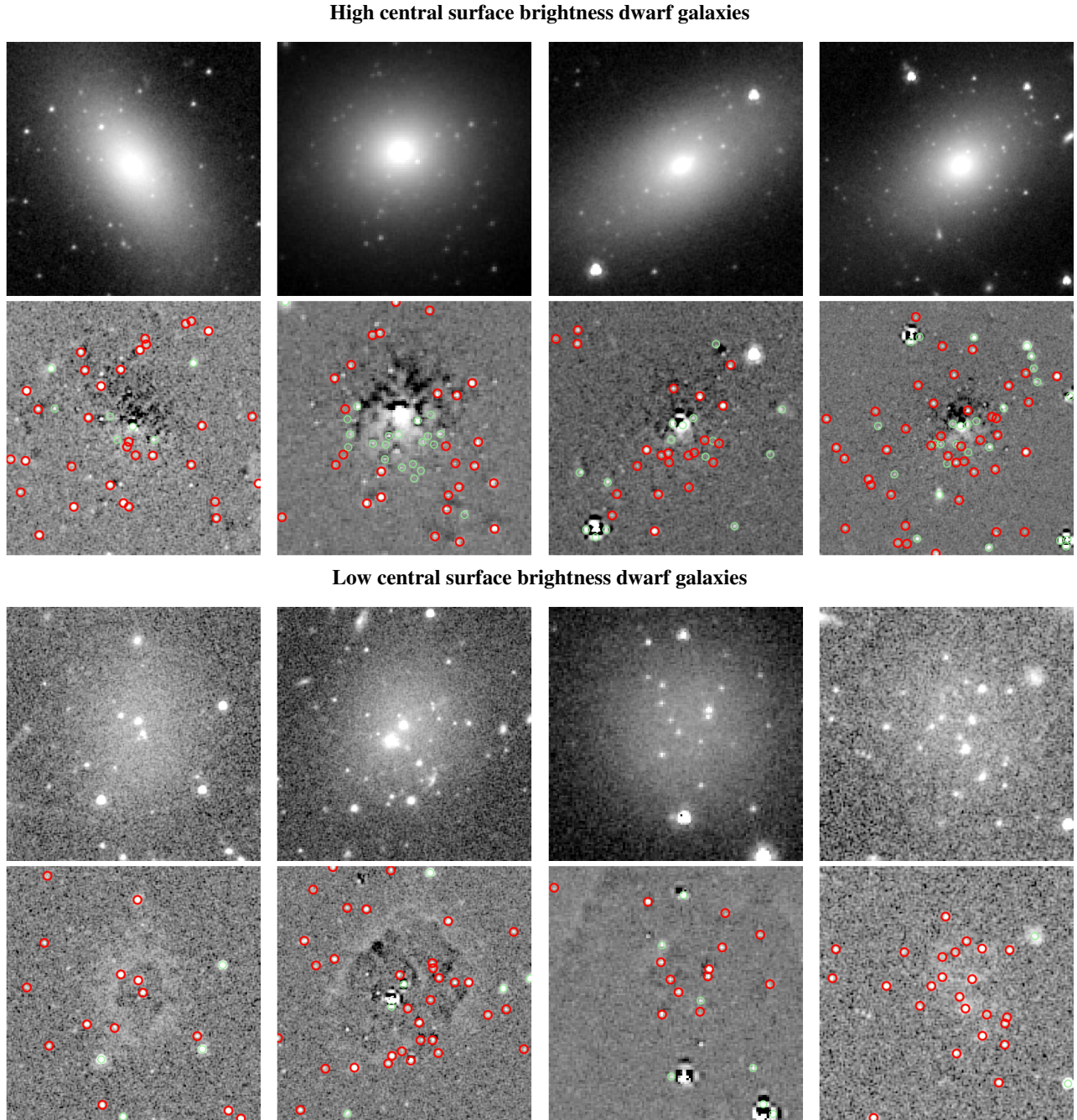


Fig. 5. Source detection and GC candidate selection for eight of the dwarf galaxies in this work. The figure shows galaxies with a high and low central brightness respectively in the top and bottom half. For each dwarf galaxy, it includes a small I_E cutout with sides of $4R_e$ and the galaxy-subtracted version of this cutout. Galaxy-subtraction has been done by unsharp-masking of the cutouts using a small background mesh size (8×8 pixels), using SEXTRACTOR. Among the detected sources, GC candidates are highlighted in red; other sources (in green) did not match the selection criteria.

more than 10 GCs, we examine the R_{GC} values for individual cases. As mentioned before, the dataset is not deep enough to conduct a comprehensive analysis of the GCLF of dwarf galaxies itself, for example, to study the GCLF TOM and width as a function of the stellar mass of the host galaxies. Additionally, the estimated colours in this work, while suitable for GC selection using flux-dependent colour cuts, are not ideal for studying the colours of individual GCs given their large uncertainties. Therefore, in this paper, we omit the study of the GCLF and of GC colours.

4.1. Dividing the sample of dwarf galaxies

Dwarf galaxies at a given stellar mass display a diversity of surface brightnesses and effective radii. Here, we further study trends between the GC populations and the properties of their host galaxy. We do this by dividing our sample of dwarf galaxies into two groups in each of two parameter spaces: the planes of central surface brightness (μ_0) vs. stellar mass (M_*), and of effective radius (R_e) vs. stellar mass (M_*). Figure 6 shows this division.

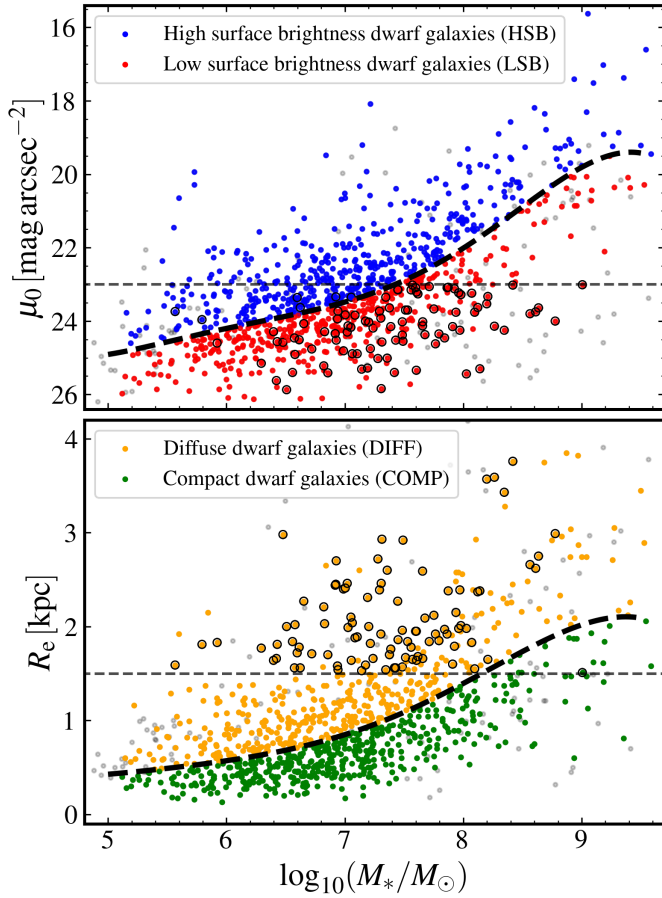


Fig. 6. Division of the dwarf galaxy sample into HSBs (blue points), LSBs (red points), COMPs (green points), and DIFFs (yellow points) based on their central surface brightness (μ_0) and effective radius (R_e). A distance of 72 Mpc was assumed to convert angular to physical scale. This division is stellar mass dependent and is indicated with the black dashed curve. Coloured points represent the dwarf galaxies that have been used for the analysis in this paper. Those dwarf galaxies that also satisfy the UDG criteria in μ_{0,I_E} and R_e ($\mu_{0,I_E} < 23 \text{ mag arcsec}^{-2}$ and $R_e > 1.5 \text{ kpc}$, shown with dashed horizontal lines) are indicated with the black circles around the coloured circles. Grey points show those dwarf galaxies that have been removed from the sample after initial assessment (as was mentioned in Sect. 2.2).

Firstly, as shown in the top panel of Fig. 6, presenting μ_0 - M_* , we divided the dwarf galaxies into low and high surface brightness sources. This division is made by estimating the median in overlapping bins for a sliding bin with a width of 1 dex along the mass axis. We note that the dividing line obtained this way is stellar mass dependent. This threshold coincides with that for UDGs (I_E central surface brightness of about $23 \text{ mag arcsec}^{-2}$) at a stellar mass of $10^8 M_\odot$. Secondly, in the lower panel of Fig. 6, showing R_e - M_* , we make a division between diffuse and compact dwarf galaxies. The terms diffuse and compact correspond respectively to larger and smaller R_e in a given stellar mass bin, which relates to the concentration of the light distribution. These divisions are made to explore how R_{GC} and N_{GC} vary with the host galaxies' surface brightness and effective radius at any given stellar mass. From now on in this article, we use the same colour code in all the figures to represent the four categories of objects defined in Fig. 6. We also assign a short label to each category. These four categories (and their labels) are:

- Blue: high surface brightness dwarf galaxies (HSBs),

- Red: low surface brightness dwarf galaxies (LSBs),
- Green: compact dwarf galaxies (COMPs),
- Yellow: Diffuse dwarf galaxies (DIFFs).

Figure. 6 also shows dwarf galaxies that satisfy the UDG criteria in central surface brightness ($\mu_{0,I_E} < 23 \text{ mag arcsec}^{-2}$) and effective-radius ($R_e > 1.5 \text{ kpc}$) with black circles. These thresholds in μ_0 and R_e are indicated with a horizontal black dashed line in each panel. From now on, along with the four categories of dwarf galaxies (LSB/HSB and COMP/DIFF), we examine GC properties of UDGs and non-UDGs, and use a consistent colour-code in the figures throughout the rest of the paper:

- Purple: UDGs,
- Grey: non-UDGs.

4.2. Radial distributions of GCs

Considering the small numbers of GCs typically associated with dwarf galaxies, particularly at the lowest galaxy masses and luminosities, studying the GC radial distribution is not in general feasible without stacking samples. With such stacked samples, we are able to study the average GC radial distribution around dwarf galaxies of various categories.

Here, we investigate these stacked radial profiles of GCs around all the dwarf galaxies in the sample across stellar masses. We define eight overlapping galaxy-mass bins, each with a width of 1 dex in stellar mass, covering the range from $10^{5.0} M_\odot$ to $10^{9.5} M_\odot$. For dwarf galaxies within such a mass bin, we estimate the galactocentric distances of the GC candidates brighter than the GCLF TOM (considering extinction) and produce a combined (stacked) radial profile. Then we estimate the Sérsic effective radius of this GC distribution by fitting a Sérsic function, fixing the Sérsic index to $n = 1$ and allowing for a constant additive background component. The value $n = 1$ is motivated by previous works for GCs of dwarf galaxies and UDGs in various environments (Saifollahi et al. 2022; Janssens et al. 2024; Tang et al. 2025), which found values between 0.5 and 2.

The fitting procedure here is carried out in three steps and in each step, we perform Sérsic fitting (with a constant background component) with a different aim. The aim of the first fit is to constrain the background component, representing the average GC density in the background, across the frame, at the position of the dwarf galaxies in the sample and in the stellar mass bin. For this step, we consider all the GC candidates within the initial cutout (out to $60''$), using a constant bin width, about $3''$ (leading to 20 radial bins in total) for the radial distribution of GC. This approach leads to more data points at larger distances to properly fit and constrain the background. The threshold of $60''$ is imposed because of the initial size of cutouts ($120'' \times 120''$). For the most massive dwarf galaxies in the sample with $M_* \sim 10^{9.5} M_\odot$ that are expected to have the most spatially extended GC populations in the sample, based on Lim et al. (2020), we expect $R_{GC} \sim 3 \text{ kpc}$. This implies that GCs typically can extend up to about 9 kpc from their host galaxies, corresponding to about three times R_{GC} . The area within this radius is about a fifth of the total area of each cutout.

Once the GC density of the background is estimated, we perform the second fit and with this fit, we aim to estimate an initial guess for the R_{GC} values, which will be used for the third step. For the second fit, we repeat the Sérsic fitting procedure before, this time for the bins within 10 kpc (about $30''$) of the galaxy (10 bins in total). This is the radius beyond which we do not expect many GCs associated with dwarf galaxies. This fit is

done with the fixed background component calculated from the first fit.

In the third step, based on the initial R_{GC} of the second fit, we only consider GC candidates within three times the initial R_{GC} (which theoretically encompasses 96% of the GCs around the host galaxy, Trujillo et al. 2001). To establish the median R_{GC} and its associated uncertainty, for the third fit, we perform the fitting 1000 times using 95% of the initial number of the GC candidates, randomly selected each time (with replacement), and we adopt the median R_{GC} and the 68% confidence interval as the final R_{GC} and its uncertainties. During this third (final) fit, we randomly chose the number of radial bins (from 6 to 10) to take into account biases that might arise from our choice for number of radial bins. We note that not all the 1000 iterations would provide a fit and the number of failed fits increases with decreasing the number of the GC candidates used for fitting. For the analysis of the stacked GC distribution, we considered fits with fewer than 100 failed iterations. In most cases, the number of failed iterations was zero.

Furthermore, we explored the initial choice of Sérsic index $n = 1$ for GCs of dwarf galaxies in this sample. Repeating the fitting procedure described above with a free Sérsic index, we find that the most likely n are in the range $n = 0.7$ – 1.3 for dwarf galaxies in different stellar mass bins, and on average about 1. This result is consistent with the Janssens et al. (2024) where authors find an average $n = 1.1$ for six Perseus cluster dwarf galaxies with $N_{GC} > 20$. Therefore, $n = 1$ seems to be a reasonable assumption for estimating R_{GC} for our sample, consistent with the previous literature.

The estimated R_{GC} for galaxies stacked in a mass bin and the associated uncertainties, as well as the ratio between R_{GC} and the host galaxy R_e , are shown in Fig. 7 for all the adopted stellar mass bins. The ratio presented in the lower panel is $R_{GC}/\langle R_e \rangle$, where $\langle R_e \rangle$ is the average effective radius of the dwarf sample for the given stellar mass bin. The value of $R_{GC}/\langle R_e \rangle$ differs from the average of the individual per galaxy ratios R_{GC}/R_e , which we write as $\langle R_{GC}/R_e \rangle$. The latter, cannot be measured because the number of GCs per dwarf galaxy is too small. Appendix C gives our assessment of the difference of these two quantities, suggesting that $R_{GC}/\langle R_e \rangle$ is up to 28% smaller than $\langle R_{GC}/R_e \rangle$. The correction factor depends little on galaxy mass (10% change across our mass bins), and drops from a 36% to a 23% correction in different subsamples because of the smaller dispersion in R_e within each of these subsets. On average, the correction factor is 28%/33% for UDGs/non-UDGs, 30%/35% for LSB/HSB, and 32%/36% for DIFF/COMP. For dwarf galaxies more massive than $M_* > 10^8 M_\odot$, the contrast between correction factors is higher; 23%/34% for LSB/HSB, and 23%/35% for DIFF/COMP. This difference between the correction factors implies that the difference between $\langle R_{GC}/R_e \rangle$ values for different dwarf categories is stronger than the one seen for $R_{GC}/\langle R_e \rangle$ in Fig. 8.

In Fig. 7, we observe an overall trend that R_{GC} and R_{GC}/R_e decrease with decreasing the stellar mass of dwarf galaxies. Such positive correlations are also seen in more massive galaxies (Forbes 2017; Hudson & Robison 2018; Lim et al. 2024). In this figure, we overplotted the analytical trends from Lim et al. (2024), extrapolated into the stellar mass range of dwarf galaxies. These equations are derived for all GCs, blue GCs, and red GCs of massive galaxies typically more massive than $10^9 M_\odot$. In such massive objects, the blue GCs are usually interpreted as originating from accreted dwarfs, while the red GCs may have formed in situ or in larger accreted objects. In the dwarf galaxies of the Perseus cluster, the observed GCs have blue colours: when

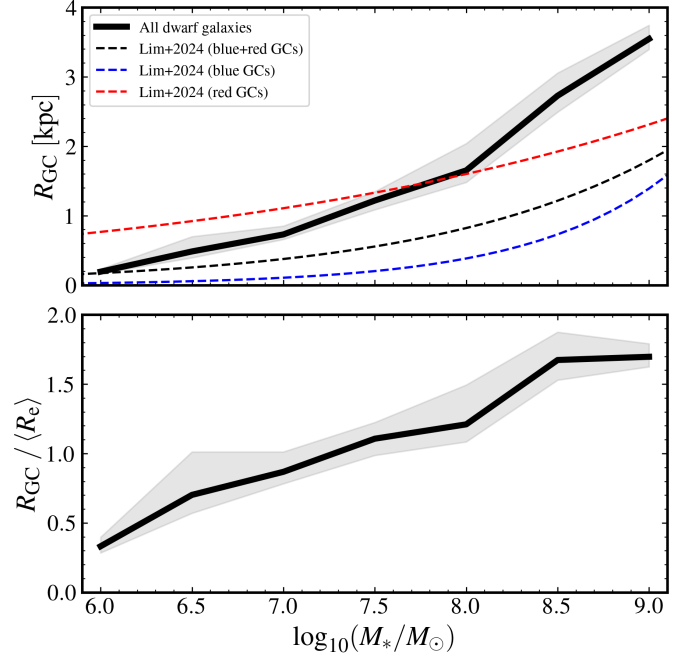


Fig. 7. Globular cluster half-number radius (R_{GC} , top panel) and ratio between the GC half-number radius and the average of the effective radius of the host galaxies ($R_{GC}/\langle R_e \rangle$, bottom panel) as functions of the stellar mass of the host galaxies (M_*) for all dwarf galaxies in the sample. The values are estimated from stacked GC radial profiles that we fit with a Sérsic function, with a Sérsic index fixed to $n = 1$ (see text). The shaded region around each line represent the 68% percentiles of the estimated average value. The dashed lines show the trends one would expect from a simple extension of the equations in Lim et al. (2024), that were derived for GCs in massive galaxies, to our lower mass regime (an extrapolation beyond the range examined by these authors).

we use single stellar population models to convert their typical Euclid colours to optical colours, these lie in the domain of the blue GCs of Lim et al. (2024). However, in dwarf galaxies, blue colours do not imply the GCs were accreted; more likely they were formed in situ, in the low-metallicity environment that also produced the dwarfs. In this sense, with all the due precautions associated with the exploitation of extrapolations, it may not be too surprising that the GCs of the Perseus dwarfs have system-radii closer to those extrapolated from the red GCs of massive galaxies than from the blue ones (as seen in Fig. 7).

We repeated the fitting for each subsample of dwarf galaxies (LSB/HSB, COMP/DIFF, UDG/non-UDGs) separately and the results are shown in Fig. 8. As seen in the middle panels, LSB dwarf galaxies (shown in red) tend to have (marginally) smaller R_{GC} , implying a more compact GC distribution, while HSB dwarf galaxies (in blue) have larger R_{GC} , implying a more extended GC distribution. There is no statistically significant difference between R_{GC} in COMP and DIFF and between UDGs/non-UDGs, except for the most massive bin with $M_* = 10^9 M_\odot$. We note that the estimated R_{GC} for this bin is based on only 4 UDGs (as presented in Table A.1). Therefore, overall it seems that R_{GC} is independent of the host galaxy's stellar distribution, with a marginal (or no) dependency on central surface brightness. Considering the $R_{GC}/\langle R_e \rangle$, as seen in Fig. 8, the difference between different categories is stronger. This clearly means that the differences are driven by R_e and not by R_{GC} . On average, for LSB dwarf galaxies, $R_{GC}/\langle R_e \rangle < 1$, while for HSB dwarf galaxies, $R_{GC}/\langle R_e \rangle = 1.5$. Furthermore, UDGs and DIFF

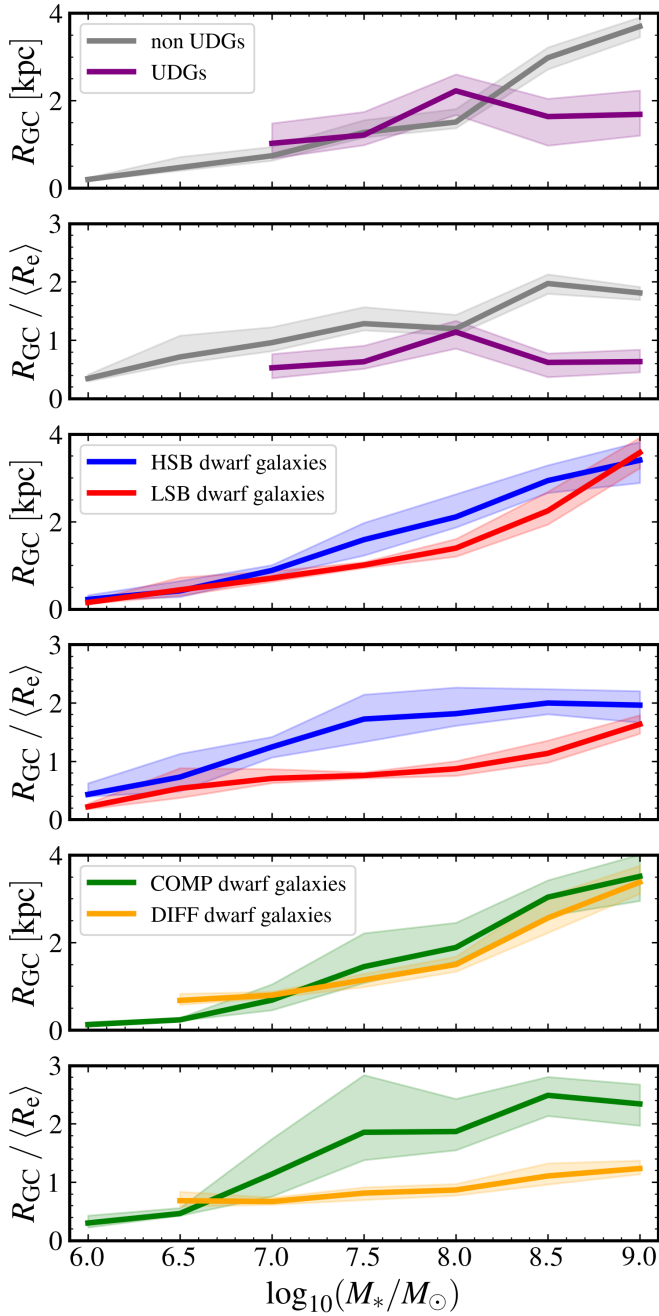


Fig. 8. Similar to Fig. 7 but for different dwarf galaxy categories. The two top panels represent the estimations for UDGs and non-UDGs (in purple and grey), the two middle panels are for those of LSBs and HSBs (in red and blue), and the bottom two panels are for DIFFs and COMPs (in yellow and green).

have $R_{GC}/\langle R_e \rangle < 1$, while non-UDGs and COMP dwarf galaxies have $R_{GC}/\langle R_e \rangle = 1.5$ – 2 .

The main indication of the analysis of GC distributions for different types of dwarf galaxies is that the GC distribution is not dependent or has weak dependencies on the host galaxy μ_0 and R_e , while $R_{GC}/\langle R_e \rangle$ would differ, leading to a smaller $R_{GC}/\langle R_e \rangle$ for LSB, DIFF, and UDGs compared to HSB, COMP, and non-UDGs. We take this into account in the next section for estimating the total GC numbers (N_{GC}) of dwarf galaxies. As shown in Appendix C, the correction factor for converting $R_{GC}/\langle R_e \rangle$ to $\langle R_{GC}/R_e \rangle$ is different between different

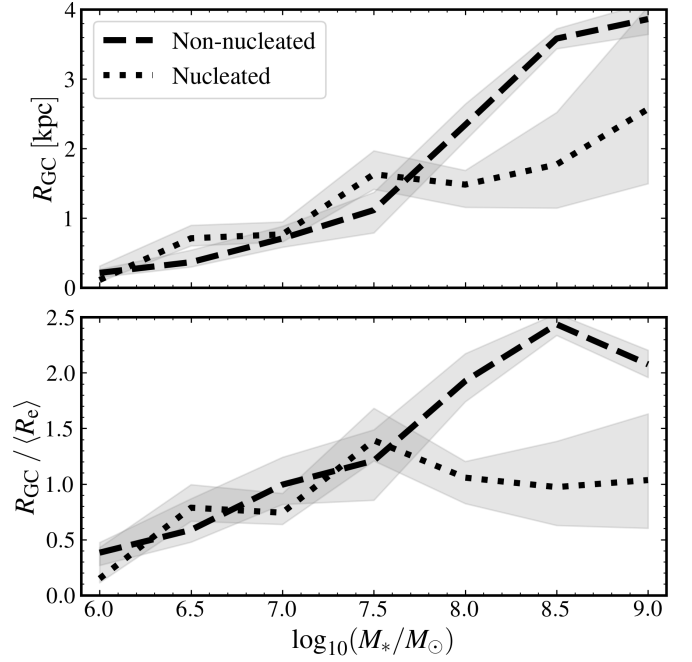


Fig. 9. Same as Fig. 7 and Fig. 8 but for nucleated and non-nucleated dwarf galaxies in our sample.

dwarf categories; on average, the correction factor on average is 1.28/1.33 for UDGs/non-UDGs, 1.30/1.35 for LSB/HSB, and 1.32/1.36 for DIFF/COMP. For dwarf galaxies more massive than $M_* > 10^8 M_\odot$, the contrast between correction factors is higher; 1.23/1.34 for LSB/HSB, and 1.23/1.35 for DIFF/COMP. This difference between the correction factors implies that the difference between the $\langle R_{GC}/R_e \rangle$ values for different dwarf categories is stronger than the one seen for $R_{GC}/\langle R_e \rangle$ in Fig. 8.

Lastly, we repeated our analysis of GC distributions for nucleated and non-nucleated dwarf galaxies. The separation between these two is based on the dwarf catalogue provided by Marleau et al. (2025). The results are shown in Fig. 9, where it is seen that there is no meaningful difference between R_{GC} and $R_{GC}/\langle R_e \rangle$ for nucleated and non-nucleated dwarf galaxies below $M_* = 10^8 M_\odot$. Above this stellar mass, nucleated dwarf galaxies show smaller R_{GC} and $R_{GC}/\langle R_e \rangle$ compared to non-nucleated dwarf galaxies. The smaller R_{GC} in nucleated dwarf galaxies could be an indication of the stronger dynamical friction of GCs within these dwarf galaxies which has also led to the formation of a nuclear star cluster (NSC; see Neumayer et al. 2020) in these dwarf galaxies (Fahrion et al. 2022; Román et al. 2023) for more than 50% of the dwarf population (den Brok et al. 2014; Sánchez-Janssen et al. 2019). The results presented in figures in this section (R_{GC} and $R_{GC}/\langle R_e \rangle$), and some details such as number of dwarf galaxies and GC candidates in each stellar mass bin are provided in Table A.1.

4.3. GC number counts

The total GC number counts of dwarf galaxies (N_{GC}) have been studied and debated extensively in recent years, with particular attention given to UDGs. The methodology for estimating N_{GC} of dwarf galaxies based on imaging data consists of several steps to take into account: (i) incompleteness in the GC candidate catalogue, (ii) the contamination from GCs that are not associated with the dwarf galaxy (including intracluster GCs,

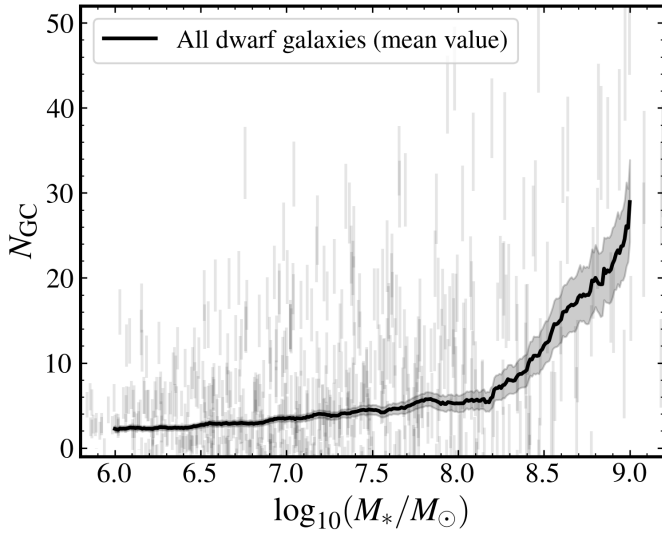


Fig. 10. Total GC numbers (N_{GC}) for all the dwarf galaxies in this work. The vertical (grey) lines represent the N_{GC} and their uncertainties for individual dwarf galaxies. The solid curve and the shaded regions around it represent the average (mean) N_{GC} and 68% confidence intervals for dwarf galaxies within a moving bin along stellar mass (x -axis) with a width of 0.5 dex. The confidence intervals are based on 1000 bootstrap-repetitions of the estimates using bootstrap sample sizes of 95% of the initial sample sizes.

foreground stars and background galaxies). These corrections would ideally require a good understanding of the properties of the GCLF (shape, width, peak magnitude) and the distribution of GCs around their host galaxies. In the case of dwarf galaxies, unless they are very rich in GCs, the GCLF and GC distribution cannot be well constrained for individual galaxies, and therefore N_{GC} estimates require some assumptions on the GCLF and the GC distribution.

As the initial step, one needs to count the GC candidates within a radius that encloses almost all the GCs. For a Sérsic distribution, this radius can be estimated using the half-number radius of the GCs (R_{GC}) for the Sérsic index (n) of the distribution. For $n = 1$, more than 99% are within $5R_{GC}$, and therefore, we do not expect any GCs beyond this radius. However, such a large radius also includes many contaminants, increasing the uncertainties in estimating N_{GC} . An alternative approach is to count GCs within a smaller radius which encircles a large fraction of GCs, and later to correct for the (small) missing fraction. For example, choosing a radius of $2R_{GC}$ (as in this work), while including 85% of the GCs, we reduce the number of contaminants by a factor of 6.25, which reduces the uncertainties (Poisson) by a factor of 2.5. In the end, the estimated N_{GC} will be corrected for the missing 15% of the GCs.

Because R_{GC} is typically not well constrained for individual dwarf galaxies, in many works it is calculated indirectly from R_{GC}/R_e by measuring the host galaxies R_e . Typically a single value of R_{GC}/R_e is used with a value between 0.7 and 1.5. However, in the previous section we showed that this ratio is dependent on the central surface brightness (μ_0) and effective radius (R_e) of the host galaxy, and therefore using a single value for a wide range of dwarf galaxies could bias the N_{GC} estimate.

Our analysis also shows that R_{GC} , in contrast, is almost independent of properties other than stellar mass, except for a possible weak dependence on the central surface brightness. Here, we investigate the N_{GC} of dwarf galaxies with this new under-

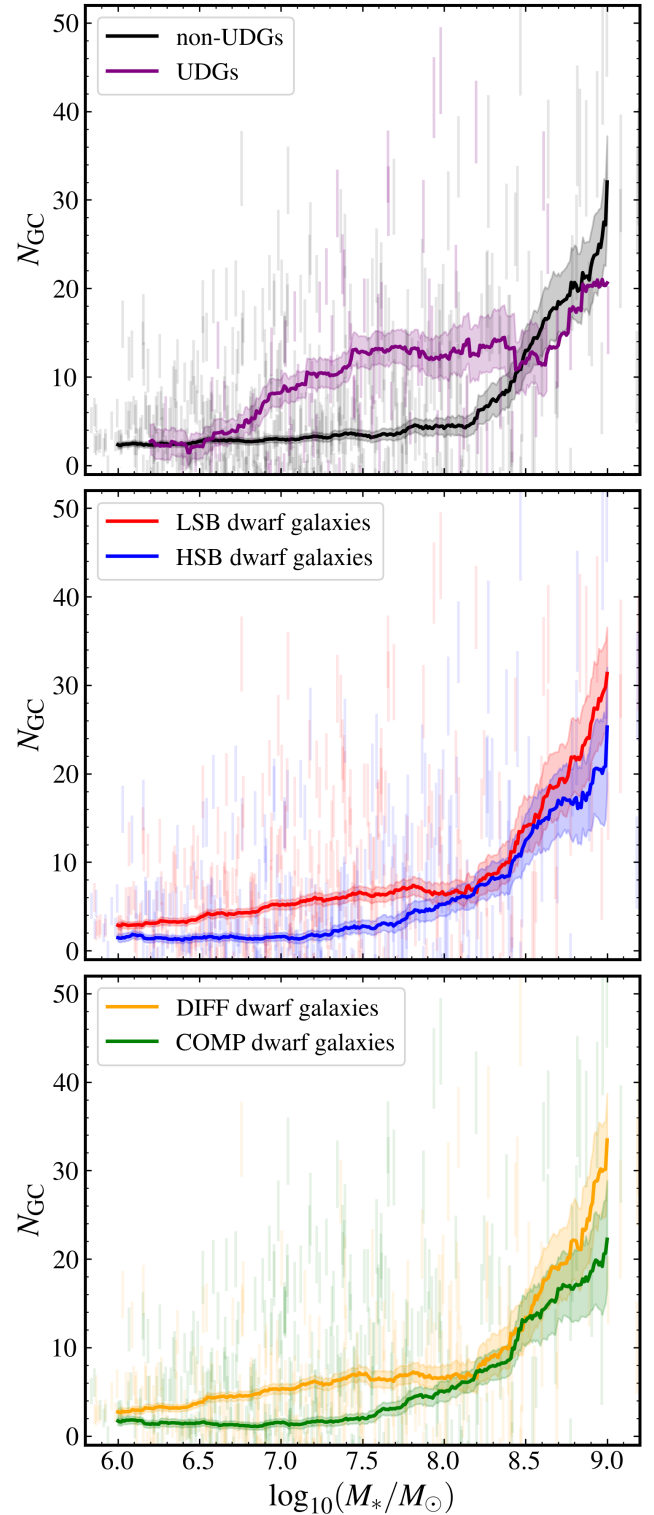


Fig. 11. Similar to Fig. 10 for UDGs/non-UDGs (top), LSB/HSB dwarf galaxies (middle), and COMP/DIFF dwarf galaxies (bottom). The vertical lines in each panel represent the N_{GC} and its uncertainties for individual galaxies. The solid line and the shaded region around it represent the average (mean) N_{GC} and 68% confidence intervals for dwarf galaxies within a moving bin along stellar mass (x -axis) with a width of 0.5 dex.

standing of R_{GC} , R_{GC}/R_e , and their possible dependence (only marginally, based on our data) on μ_0 . For each dwarf galaxy, characterised by its stellar mass and central surface brightness,

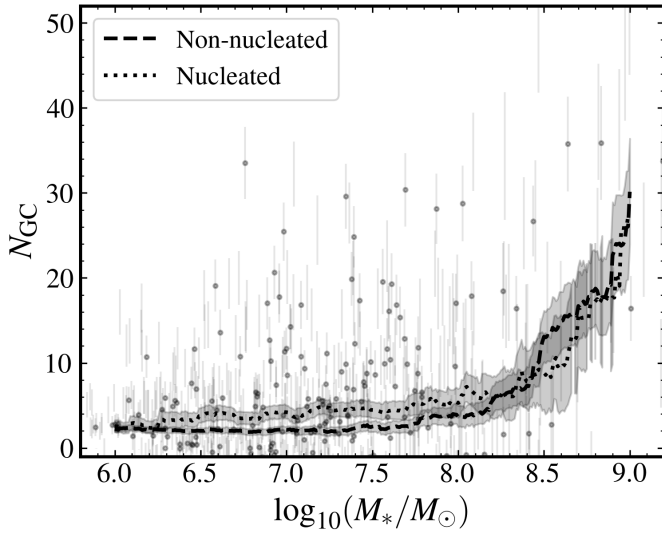


Fig. 12. Similar to Fig. 10 and Fig. 11 but for the nucleated and non-nucleated dwarf galaxies in the sample. The vertical lines with and without a central small circle indicate nucleated and non-nucleated dwarf galaxies, respectively.

we construct the stacked radial profile of the GC systems of all the dwarf galaxies within 0.5 dex in stellar mass and 1 mag in central surface brightness, and estimate R_{GC} . We then use R_{GC} (rather than a single R_{GC}/R_e value for all dwarf galaxies) to quantify the spatial extent of the GC system of the initial galaxy and to evaluate its N_{GC} . This approach should be less biased than the alternatives based on R_{GC}/R_e .

In practice, we first count the number of GC candidates brighter than the GCLF TOM within $2R_{GC}$. Here we limit our GC sample to GCs brighter than the GCLF TOM because our detection completeness drops quickly below this magnitude. The value of the GCLF TOM used here is $I_E = 26.6$ plus the I_E foreground extinction correction for the line of sight of a given dwarf galaxy. Then, we correct for this number of GC candidates for background contamination by counting GC candidates within the cutout that are farther than $5R_{GC}$ (where we do not expect any GC for $n = 1$), normalised to the ratio of areas within $2R_{GC}$ and beyond $5R_{GC}$. Additionally, using the estimated completeness function (Sect. 3.3), we correct the GC numbers for incompleteness in detection up to TOM and GCs beyond $2R_{GC}$ (about 15% for a Sérsic profile with $n = 1$), and double the resulting number for the GCs fainter than TOM to obtain the total GC number. We also estimate uncertainties in N_{GC} , taking into account the Poisson errors of GC candidate numbers within $2R_{GC}$ and farther than $5R_{GC}$ in the background. These GC numbers are also compared with the numbers in previous works (Janssens et al. 2024; Marleau et al. 2025) in Appendix D.

We examined the behaviour of N_{GC} for all the dwarf galaxies in Fig. 10, and present the counts separately for UDGs/non-UDGs, LSB/HSB and COMP/DIFF dwarf galaxies in Fig. 11. In each panel, we also show the average (mean) N_{GC} at a given stellar mass (taking into account dwarfs with zero and negative N_{GC}). As expected, N_{GC} increases with stellar mass in all cases. In general, we do not observe any hint of the presence of a distinct GC-rich population of dwarf galaxies in the N_{GC} – M_* space. Inspecting Fig. 11, we find a higher N_{GC} for LSB, DIFF, and UDGs compared to HSB, COMP, and non-UDGs. The difference between UDGs and non-UDGs is already established (Lim et al. 2018, 2020). Here we also find a higher N_{GC} when we

subdivide the samples into LSB and HSB galaxies or DIFF and COMP galaxies. The previously observed GC excess in UDGs arises from the combination of surface brightness and effective radius; neither characteristic alone is sufficient to produce this effect. Assuming that N_{GC} is a proxy for the total (dynamical) mass of galaxies (Spitler & Forbes 2009; Burkert & Forbes 2020; Zaritsky 2022), this indicates that at a given stellar mass, LSB, DIFF, and UDG dwarf galaxies are more massive (in total mass) than HSB, COMP, and non-UDG dwarf galaxies. These trends were previously discussed for UDGs; however, our results show that such trends are not unique to UDGs and can be extended to all cluster dwarf galaxies, over a range of stellar masses.

Lastly, we examined the N_{GC} values of nucleated and non-nucleated dwarf galaxies and the result is shown in Fig. 12. At a given stellar mass and for $M_* > 10^8 M_\odot$, there is no meaningful difference between the average N_{GC} in nucleated and non-nucleated dwarf galaxies. However, for $M_* < 10^8 M_\odot$, nucleated dwarf galaxies show a higher N_{GC} than non-nucleated dwarf galaxies. This observation, while it might look counter-intuitive, is consistent with the previous findings (Carlsten et al. 2022). This could be the outcome of galaxy evolution in a rich environment, which leads to a more efficient NSC formation as well as a higher N_{GC} for dwarf galaxies at a given stellar mass.

4.4. GCs of individual dwarf galaxies

In previous sections, we examined the stacked R_{GC} of dwarf galaxies, mainly because dwarf galaxies do not host many GCs. As a result, the trends discussed correspond to the general population of dwarf galaxies. However, it is likely that some dwarf galaxies, although small in number, do not behave similarly. Here, we examine R_{GC} for dwarf galaxies with more than 10 GCs (164 dwarf galaxies), where we might gain insights into their individual GC properties. Images of some of these dwarf galaxies are shown in Fig. 13. For these dwarf galaxies, we relaxed the number of failed iterations (for the third fit, as described in Sect. 4.2) by adjusting the threshold of the failed iterations to 900 (out of 1000). Additionally, we only consider dwarf galaxies with R_{GC} errors less than 2 kpc.

The results of our analysis of individual galaxies are presented in Fig. 14 for LSB/HSB, COMP/DIFF, and UDG/non-UDG dwarf galaxies. Out of 164 dwarf galaxies with more than 10 GCs, we could measure R_{GC} for 78 of them, with an overall uncertainty of less than 2 kpc. The uncertainties are estimated using bootstrapping, similar to what was done earlier for the analysis of stacked GC profiles in Sect. 4.2. We note that dwarf galaxies with $N_{GC} > 10$ are among the most massive dwarf galaxies in our Perseus sample (stellar mass mostly between $10^8 M_\odot$ and $10^9 M_\odot$), but are only a subset of those most massive dwarfs. Therefore, some caution should be exercised when comparing these figures with Fig. 7 and Fig. 8.

All panels of Fig. 14 display considerable dispersion. Simulations of GC systems drawn randomly from a Sérsic distribution with $n = 1$ lead to distributions of recovered R_{GC} of which the 2.5% and 97.5% quantiles are found at 0.55 and 1.67 times the true R_{GC} for $N_{GC} = 10$, and at 0.68 and 1.42 times the true R_{GC} for $N_{GC} = 25$. The bootstrap estimates shown by the shaded regions in the figure are representative of this inevitable uncertainty. We also compare our results for R_{GC} of individual dwarf galaxies with the values provided by Janssens et al. (2024) in Fig. D.1. Of the 40 dwarf galaxies common to both studies, five dwarf galaxies have R_{GC} values in both works. Two dwarf galaxies out of five, namely

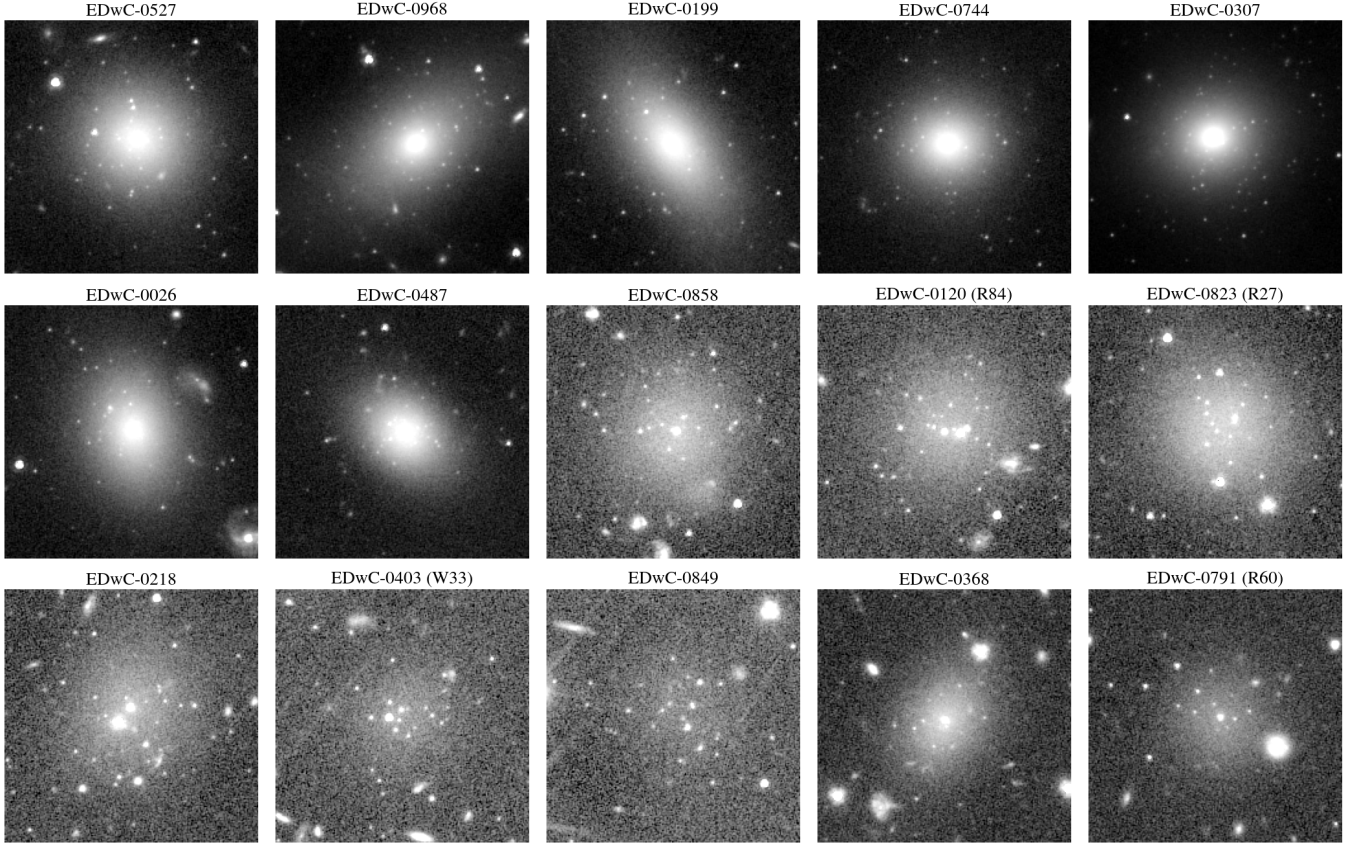


Fig. 13. Examples of Perseus cluster dwarf galaxies with more than ten GCs ($N_{GC} > 10$) in I_E data. The majority of such dwarf galaxies have a stellar mass of $M_* > 10^8 M_\odot$, while they cover a range of central surface brightness (μ_0) and effective radius (R_e). The *Euclid* ID of dwarf galaxies from Marleau et al. (2025) are written on top of each cutouts, as well as their name in Janssens et al. (2022) when available (written in parenthesis). The cutouts have dimensions of $40'' \times 40''$ (about 14×14 kpc).

EDwC-0120 and EDwC-0823 (see Fig. 13), are among the most GC-rich dwarf galaxies in our sample and within the uncertainties, the estimated R_{GC} values are consistent between the two analyses, with $R_{GC} = 2.7^{+1.0}_{-1.0}$ kpc and $R_{GC} = 1.7^{+0.5}_{-0.7}$ kpc (and $R_{GC} = 2.4^{+0.4}_{-0.3}$ kpc and $R_{GC} = 2.2^{+0.4}_{-0.3}$ kpc in Janssens et al. 2024). For the other three dwarf galaxies, namely EDwC-1011, EDwC-0403, and EDwC-0791, our values of R_{GC} are, on average, 1.5 kpc smaller. Examining these galaxies (Fig. 13), we can see a clear over-density of point sources (most likely GC candidates within the apparent half-light radius of dwarf galaxies EDwC-0403 and EDwC-0791). We take this observation as a sign that $R_{GC}/R_e < 1.0$ for these dwarf galaxies, as estimated in this work, is justified. To assess this a bit further, we compared the R_{GC} values with Li et al. (2025), which uses the same *HST* data as Janssens et al. (2024), and applies a novel approach to derive the GC properties of dwarf galaxies. We find a better consistency between our estimates of R_{GC} and Li et al. (2025) – see Fig. D.1.

We focus on R_{GC} in the left panels of Fig. 14. In the subset of dwarf galaxies with $N_{GC} > 1$, we again find a positive correlation between R_{GC} and host galaxy stellar mass, with no major systematic difference between UDG/LSB/DIFF objects on one hand, and non-UDG/HSB/COMP ones on the other. This is similar to the trends found earlier using the stacked GC radial profiles (shown with solid coloured curves). These curves, are derived and presented earlier in Sect. 4.2 and Fig. 8, and they are not a fit to the data-points in Fig. 14. Overall, be it with respect to host galaxy stellar mass or to N_{GC} , the dwarf galaxies of various types seem to follow the behaviour one obtains by extrapolat-

ing the trend seen for the red GC subpopulations of more massive galaxies (dashed red curve). This may imply that GCs of dwarf galaxies, even though they are blue and metal-poor, have formed in situ similarly to red and metal-rich GCs in massive galaxies.

The R_{GC}/R_e values of individual cases, shown in the right-hand panels of Fig. 14, in some cases seem compatible with the trends we found earlier with stacked GC radial profiles (solid lines). In particular, UDGs have smaller R_{GC}/R_e on average, with values close to or smaller than 1, while non-UDGs have a higher R_{GC}/R_e , about 1.5. Similarly, a tendency for having smaller R_{GC}/R_e values is seen for DIFF dwarf galaxies compared to the COMP dwarf galaxies. We do not see any significant difference between the R_{GC}/R_e ratios of LSB and HSB galaxies. Lastly, we want to emphasise that regardless of the average behaviour, the GC distribution covers a range of values, with R_{GC}/R_e from 0.5 to 2.

5. Discussion

In this work, we inspect several trends that relate properties of the GC systems of dwarf galaxies to properties of these host galaxies (at a given M_*), mainly central surface brightness (μ_0) and effective radius (R_e). The interpretation of such trends regarding galaxy formation models requires models that take into account galaxy evolution and GC formation and evolution simultaneously; otherwise, interpreting the results is not

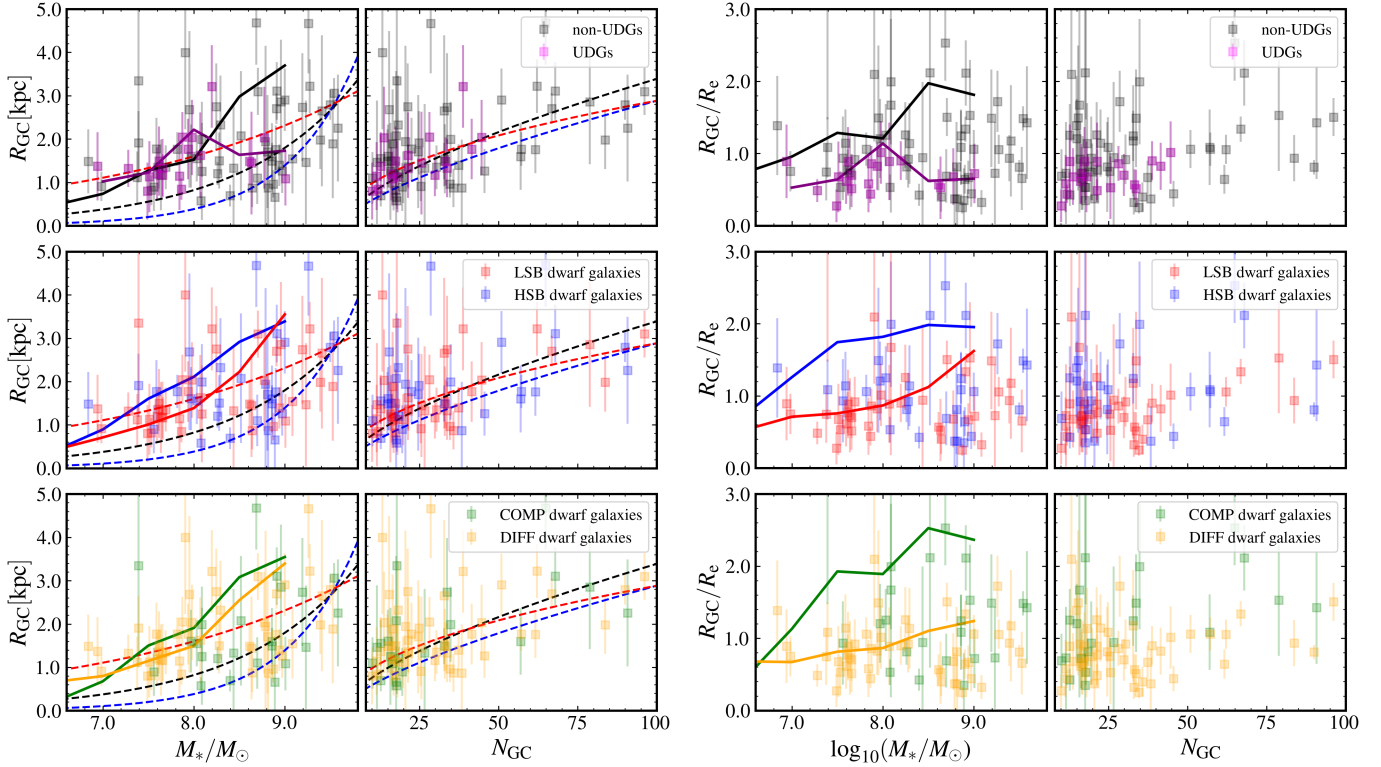


Fig. 14. Properties of the radial distributions of the GC systems of dwarf galaxies with $N_{GC} > 10$. Panels on the left and right show R_{GC} and R_{GC}/R_e values for different categories of dwarf galaxies versus their M_* and N_{GC} . The panels on the left also show the extrapolated equations from Lim et al. (2024) for all GCs (black dashed line), blue GCs (blue dashed line) and red GCs (red dashed line). We recall that these curves are our extrapolations of analytical expressions derived at higher galaxy masses into the low-mass regime. Solid lines show the results from the stacked GC distributions in Sect. 4.2 for each dwarf galaxy category. Therefore, these lines are not a fit on the presented data-points, but the curves presented previously in Fig. 8.

straightforward. Here, we briefly discuss the immediate implications of the results.

5.1. Observed trends for GCs of dwarf galaxies

One of the main findings of this work is that at a given M_* for dwarf galaxies, the distribution of GCs around dwarf galaxies, as measured through R_{GC} , is independent of R_e and almost independent of μ_0 (some hints of dependency is found). This is based on the analysis of the stacked radial distribution, which represents the average behaviour of dwarf galaxies. We see the same behaviour for individual dwarf galaxies with enough GCs that we could estimate their R_{GC} . Overall, R_{GC} increases with M_* . The increase seems to be more consistent with the extrapolation of the best-fit curve on the red GCs of massive galaxies in Lim et al. (2024). This indicates that GCs of these dwarf galaxies are mostly formed in situ. This is also the case for R_{GC} versus N_{GC} , and dwarf galaxies in this study follow the curve for red-GCs of massive galaxies (relatively, and more than the curve for blue GCs) which supports the in situ nature of the GCs within dwarf galaxies.

The trends are different for R_{GC}/R_e , both in the stacked GC profile and in individual cases. We mostly found smaller R_{GC}/R_e for diffuse galaxies and lower surface brightness galaxies. This is also the case for UDGs, which are in the overlap zone between low surface brightness and diffuse dwarf galaxies (LSB and DIFF). Our analysis shows that the excess of GC systems of LSB and DIFF dwarf galaxies is mirrored in UDGs, but with a larger amplitude. This suggests

that the excess of GC systems of UDGs results from the combination of diffuse and low surface-brightness properties; neither characteristic alone (e.g. DIFF or LSB) is sufficient to produce the quantitative trend seen for UDGs. Thus, these unique aspects of UDGs conspire to regulate their GC systems.

5.2. Stellar-to-halo mass ratios based on N_{GC}

The total number of GCs of a galaxy has been shown to correlate tightly with galaxy total mass, mostly dark matter, in the high-mass regime and to remain correlated, though with more uncertainties and dispersion, down to total masses of a few $10^9 M_\odot$ (Spitler & Forbes 2009; Harris et al. 2013; Burkert & Forbes 2020; Doppel et al. 2024; Le & Cooper 2025). Assuming that this relation between N_{GC} and total mass holds for our observed dwarf galaxies, we may convert the empirical N_{GC} to total mass (about the dark matter halo mass, M_h), using

$$M_{\text{total}}/M_\odot = 5.1 \times 10^9 N_{GC}. \quad (1)$$

This conversion is calculated based on the ratio between the total mass of GCs around galaxies and the galaxies total mass, $\eta = M_{GC}/M_{\text{total}} = 3.9 \times 10^{-5}$ (estimated value for galaxies in galaxy clusters, Harris et al. 2017), and assuming an average mass of 2×10^5 for a GC (Jordán et al. 2007). With M_{total} , we can then obtain the stellar-to-halo mass ratios (SHMRs) for the sample. Different assumptions for values of η and the average mass of GCs would lead to a slightly different result than Eq. (1). Our choice for these parameters has led to Eq. (1),

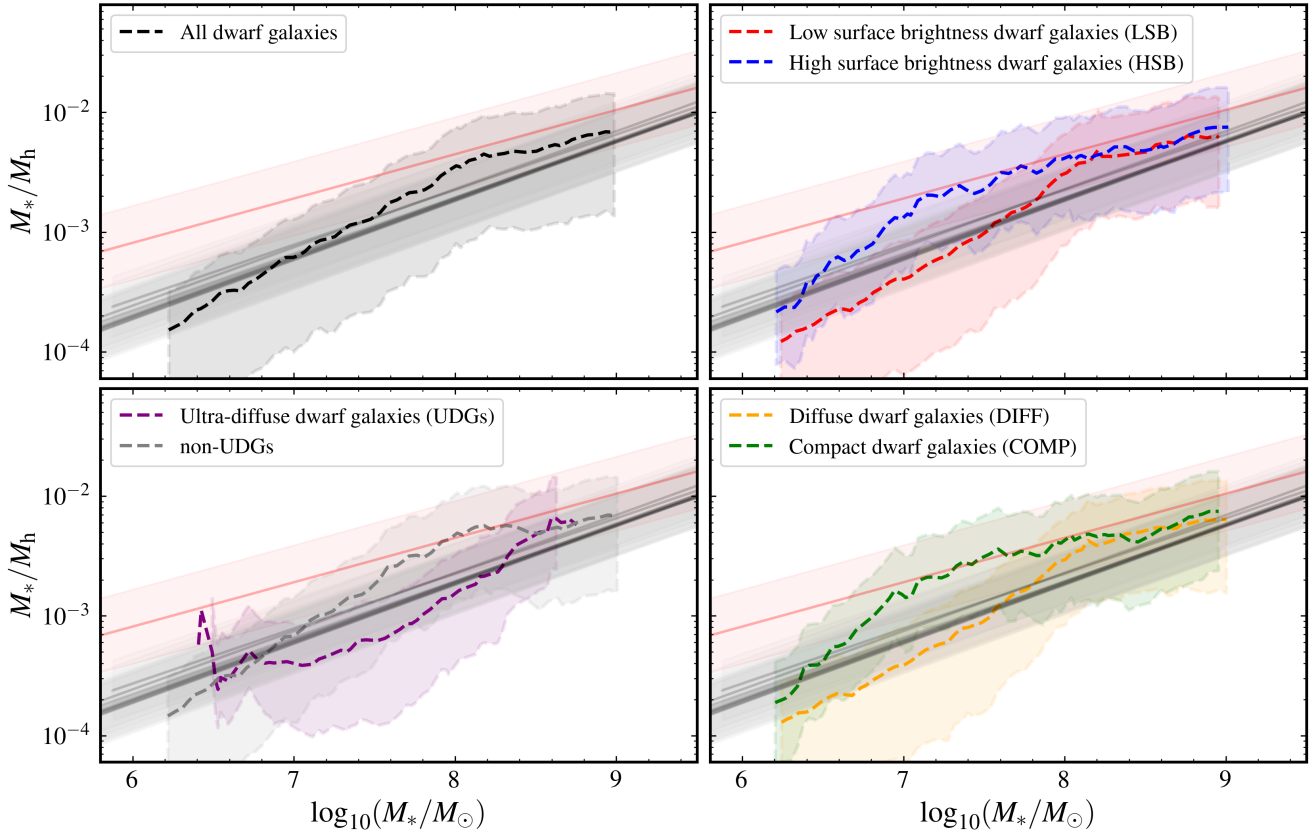


Fig. 15. Stellar-to-halo mass ratios (grey and red solid curves) of dwarf galaxies of various categories at $z = 0$ based on their GC numbers versus models (various techniques). Grey curves represent a compilation of the SHMRs as presented in Behroozi et al. (2019). The shaded area around these SHMRs correspond to 0.25 dex uncertainty for the SHMR relations. The shown SHMR at these mass regimes is an extrapolation of the SHMR obtained for the more massive galaxies. The red solid line and the shaded area around it (in panels on the left) show the SHMR relation for low-mass galaxies (at $z = 0$) in Zaritsky & Behroozi (2023) and its uncertainties (0.31 dex). The shaded area around dashed curves represents the uncertainties in M_h , converted from uncertainties in N_{GC} . These uncertainties are converted from uncertainties in N_{GC} as estimated in Sect. 4.3 (using bootstrapping) and shown in Figures 10 and 11. The shown uncertainties also include to overall uncertainties regarding the $N_{GC} - M_h$ conversion, which is $\pm(2 \times 10^9 N_{GC})$.

which is consistent with the equation $M_{total}/M_\odot = 5.0 \times 10^9 N_{GC}$ provided in Burkert & Forbes (2020), and slightly different, but by less than a factor of two, from the Zaritsky (2022) result $M_{total}/M_\odot = 2.9 \times 10^9 N_{GC}$. The conversation between N_{GC} and M_{total} provided in these works is calibrated directly on the number of GCs, and without any assumption on η and on the average mass of GCs. In general, it is important to consider the various uncertainties mentioned above when using a conversation between N_{GC} and M_{total} that results in uncertainties about $\pm(2 \times 10^9 N_{GC})$ in M_{total} .

The resulting SHMRs of dwarf galaxies are shown in Fig. 15. In each panel, we also show the SHMR fits of various works presented in Behroozi et al. (2019)⁴ at redshift $z = 0$, extrapolated to lower masses. Most of these relations are valid for galaxies with $M_h > 10^{10} M_\odot$ (or equally $M_* > 10^{6.5-7} M_\odot$), and therefore, interpreting the results for dwarf galaxies that fall below this mass threshold requires caution. These curves are valid for field galaxies, and the halo mass corresponds to the galaxies peak-halo mass, whereas the sample dwarf galaxies of this study represents galaxies in a galaxy cluster environment. In such environments, dwarf galaxies very likely have lost a fraction of their total mass due to tidal interactions with the cluster potential or other galaxies. Here we consider that tidal stripping does not

remove any GCs (Smith et al. 2013, 2015), and the current GC population represents the GCs at the peak-halo mass (before infall into the cluster). This is a plausible assumption considering that star clusters and GCs are expected to form mostly in the central regions of galaxies, where the conditions for star cluster formation are met and where they are protected by the gravitational potential of the galaxy. Figure 15 also shows the SHMR presented in Zaritsky & Behroozi (2023), shown by the red solid curve. Zaritsky & Behroozi (2023) have obtained the SHMR of low-mass galaxies down to $M_* = 10^5 M_\odot$ for four nearby ($z \sim 0$) galaxy clusters (Virgo, Fornax, Hydra, and Coma), using a mass estimator developed based on the photometric properties and scaling relations of galaxies (Zaritsky et al. 2008). Since this SHMR is based on cluster dwarf galaxies, it is not free from the effect of the environment (e.g. tidal stripping of the halo). This may explain the slightly lower dark matter fraction (higher M_*/M_h) for this relation.

As is seen in Fig. 15, within the uncertainties, the SHMR obtained for all categories of dwarf galaxies in this work are consistent with the SHMR in the literature. In particular, although UDGs are relatively more dark matter-dominated (lower M_*/M_h), they are within the scatter of the SHMR and are not exceptionally over-massive. Examining the SHMR of the different dwarf categories in comparison with Behroozi et al. (2019) SHMRs at $z = 0$ (grey solid curves), LSB/DIFF/UDG

⁴ See the references therein.

dwarf galaxies have lower M_*/M_h at a given M_* in the direction to be more compatible with the SHMR at higher redshifts. It is expected that at higher redshifts, the SHMR moves downward, towards lower M_*/M_h values. Even though the uncertainties of SHMRs are large, this may suggest that LSB/DIFF and UDG dwarf galaxies stopped forming stars already at a higher z due to a (potentially environmental) quenching mechanism; however, the exact mechanism is not clear (Forbes et al. 2023). The early quenching of these dwarf galaxies makes them failed dwarf galaxies (Buzzo et al. 2024; Ferré-Mateu et al. 2025) that have not been able to convert their mass content into stars as much as other dwarf galaxies of the same mass. Considering the average number of GCs (below 30) and corresponding halo mass (about $10^{11} M_\odot$), these objects are not failed Milky Way galaxies (van Dokkum et al. 2015). In contrast to LSB/DIFF/UDG dwarf galaxies, HSB/COMP dwarf galaxies have higher M_*/M_h at a given M_* than the average. Although most of the HSB/COMP dwarf galaxies in the Perseus cluster are red and currently quenched, they could be able to maintain their star formation long enough to convert their mass content into stars until very recently.

The scenario above is just one possible scenario that might explain the M_*/M_h of dwarf galaxies shown in Fig. 15. In any case, the failure or success of dwarf galaxies in forming stars seems to be independent of their GCs, as R_{GC} behaves similarly in all categories of dwarf galaxies. These scenarios are based on M_h from N_{GC} , which is uncertain for low-mass and dwarf galaxies. The uncertainty arises from (i) the uncertainties in the estimated N_{GC} , and (ii) the uncertainties in the estimated halo mass in dwarf galaxies. The calibration of the $N_{GC}-M_h$ relation remains one of the main tasks to be done in the coming years. In addition, having N_{GC} for dwarf galaxies in the field (rather than in galaxy clusters), where environmental effects are negligible, would provide further insight into the SHMR of dwarf galaxies.

5.3. Potential biases when estimating N_{GC} for dwarf galaxies and UDGs

The dependency of R_{GC}/R_e on central surface brightness and effective radius shows that a similar assumption on R_{GC}/R_e for all the dwarf galaxies would lead to a biased estimation of N_{GC} , resulting in even higher N_{GC} in LSB, DIFF, and UDGs. This is potentially part of the explanation behind the finding by Lim et al. (2018) of the very high N_{GC} of UDGs compared to non-UDGs in the Coma cluster, more than any other galaxy cluster/group (Prole et al. 2019; Lim et al. 2020). In some previous work, $R_{GC}/R_e = 1.5$ is adopted to calculate R_{GC} and, subsequently, to count half of the GCs (GC candidates within R_{GC}) for UDGs and non-UDGs. The total GC number N_{GC} is then estimated by multiplying this GC half-number count by two. By not taking into account such differences in R_{GC}/R_e in UDGs and non-UDGs, such an approach can overestimate N_{GC} in UDGs by about a factor of two.

Reducing the N_{GC} of UDGs in the Coma cluster does not mean that they have similar N_{GC} as non-UDGs, but it reduces the amplitude of the difference between N_{GC} of UDGs and non-UDGs. The Perseus cluster and Coma cluster have quite similar environments in richness and mass. The effect of the choice of R_{GC}/R_e for estimating N_{GC} is also addressed by Saifollahi et al. (2021b, 2022) for large UDGs (in terms of effective radius) of the Coma cluster, including DF44 and DFX1. Our results can potentially be extended to the majority of UDGs in the Coma cluster. However, we note that the values of R_{GC} and R_{GC}/R_e

cover a wide range of values, and each dwarf galaxy is unique. An average value of $R_{GC}/R_e = 1$ seems reasonable for UDGs on average (Marleau et al. 2024); however, biases in assuming an average value for all galaxies should be taken into account when estimating the total GC numbers.

6. Summary and conclusions

We used the *Euclid* ERO data of the Perseus cluster and studied the GC number counts and GC radial distribution for a uniquely large and less biased sample (regarding selection bias) of dwarf galaxies in a galaxy cluster environment. In such environments, the majority of dwarf galaxies are early types, and therefore, our sample represents early-type dwarf galaxies. These dwarf galaxies have a stellar mass ranging between $10^5 M_\odot$ and $10^{9.5} M_\odot$. The results represent the trends between dwarf galaxies and their GCs in high-density environments such as the Perseus galaxy cluster. We identified trends between GCs and their host dwarf galaxies. To do so, we divided dwarf galaxies at a given stellar mass into different categories and studied their properties: low and high surface brightness dwarf galaxies (LSBs and HSBs), and diffuse and compact dwarf galaxies (DIFFs and COMPs). We also studied UDGs and non-UDGs as well as nucleated and non-nucleated dwarf galaxies. Our findings are as follows:

1. At a given stellar mass and within the uncertainties, the average R_{GC} of LSBs and HSBs, DIFFs and COMPs, and UDGs and non-UDGs have similar values. This suggests that the R_{GC} of dwarf galaxies is independent of the effective radius of the dwarf galaxies (R_e) and central surface brightness (μ_0). However, our analysis hints at a dependence (though not significant based on our data) between R_{GC} and μ_0 .
2. At a given stellar mass and within the uncertainties, LSBs, DIFFs, and UDGs have a smaller R_{GC}/R_e , about equal to or less than one, compared to HSBs, COMPs, and non-UDGs with $R_{GC}/R_e > 1.5$. Taking into account the previous conclusion, this means that, on average, GCs are distributed over a similar radial extent for all classes of dwarf galaxies examined; however, their distributions are different from the field stars in the host galaxies.
3. For dwarf galaxies at a given stellar mass, LSBs, DIFFs, and UDGs have a higher N_{GC} compared to HSBs, COMPs, and non-UDGs. Although such behaviour for UDGs has already been established in a previous work, here we find a similar though reduced trend for the LSB and DIFF galaxies. The excess of GCs in UDGs can be attributed to the unique characteristics of lower surface brightness and a larger effective radius, as shown in the LSB and DIFF subsamples. UDGs encapsulate both of these properties, distinguishing them from other dwarf galaxy populations.
4. At a given stellar mass above $M_* = 10^8 M_\odot$, nucleated dwarf galaxies have a lower R_{GC} and R_{GC}/R_e than non-nucleated dwarf galaxies. Additionally, at a given stellar mass below this threshold, nucleated dwarf galaxies have a higher N_{GC} than non-nucleated dwarf galaxies. For more massive dwarf galaxies with $M_* > 10^8 M_\odot$, no significant difference between the N_{GC} in nucleated and non-nucleated dwarf galaxies was seen.
5. Our estimates of dark matter halo mass based on GC number counts are entirely consistent with extrapolated trends of the $z = 0$ SHMR relation at low masses. This is the case for all the different categories of dwarf galaxies. This implies that, unlike some previous claims, UDGs do not show any exceptional excess in dark matter (i.e. being over-massive)

and that their SMHR and dark matter content are within the expected ranges.

6. Our results suggest that R_{GC}/R_e depends on the central surface brightness and effective radius. This has implications for estimations of N_{GC} in dwarf galaxies of different types as well as in UDGs and non-UDGs, and it potentially points to existing biases in the estimated N_{GC} of UDGs in some previous works in the literature.

The observed trends provide a benchmark to test current as well as future cosmological simulations that simulate dwarf galaxies and their GCs in galaxy clusters (Doppel et al. 2023; Pfeffer et al. 2024). More observations of similar environments and other environments are required to further characterise these trends. *Euclid* will be fundamental for such studies, as it will observe a large number of dwarf galaxies in various environments (Euclid Collaboration: Marleau et al. 2025), including systems within the Fornax-Eridanus supercluster (Raj et al. 2020) as well as the Coma cluster and the northern part of the Virgo cluster and hundreds of nearby dwarf galaxies in isolation. A uniform analysis of such a homogeneous dataset promises a better understanding of dwarf galaxies and their GC properties across different environments. The future data releases of the *Euclid* mission also include deep ground-based data (*Euclid*-EXT) in the blue optical bands, in the u and g bands, from various surveys (UNIONS/DES/LSST). Additional colour cuts and constraints from these blue bands would improve the selection of the GCs (candidates) and further remove any remaining contamination in the GC (candidate) samples.

Data availability

Full Table A.2 is available at the CDS via <https://cdsarc.cds.unistra.fr/viz-bin/cat/J/A+A/703/A184>

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Appendix A: Tables

Here, Tables A2 and A.1 provide the results of the analysis in this paper.

Appendix B: Completeness in the near-infrared bands (Y_E , J_E , H_E)

Here we provide the results of the completeness assessment for source detection and photometry in the near-infrared bands of the ERO data, in Y_E , J_E , H_E . The results are shown in Fig. B.1.

Appendix C: Ratio between $\langle R_{GC}/R_e \rangle$ and $R_{GC}/\langle R_e \rangle$

We assessed the ratio between $\langle R_{GC}/R_e \rangle$ and $R_{GC}/\langle R_e \rangle$, referred to as α , with simulations constructed for each galaxy-mass bin under the following assumptions:

- Within a bin, the one-dimensional radial distribution of the GC-population of any galaxy is a sample originating from a single parent distribution that has a median radius R_{GC}^{th} , a theoretical quantity for which we adopt the empirical value R_{GC} shown in Fig. 7 (i.e. the median radius of the stacked GC-populations of the mass bin).
- Within a bin, the two-dimensional density distributions of the GC-populations around their host galaxy centre all follow a Sérsic law $\rho_{2D}(r)$ with the same index n (which we may vary in our experiments). This law can be written as

$$\rho_{2D}(r) \propto \exp \left\{ -b_n \left[\left(\frac{r}{r_e} \right)^{\frac{1}{n}} - 1 \right] \right\},$$

where r is the distance to the galaxy centre, and b_n is the constant that ensures that r_e indeed is the effective radius, which in our context means $r_e = R_{GC}$, since this law is used to describe the radial distribution of GCs.

For each galaxy of a given bin, the effective radius R_e of the light profile is available from measurements (Marleau et al. 2025). If the effective radius of the GC-population is the single value $R_{GC}^{th} \equiv R_{GC}$, then $\langle R_{GC}/R_e \rangle = R_{GC} \langle 1/R_e \rangle$, and on average $\alpha = \langle 1/R_e \rangle \langle R_e \rangle$ (a reference value written α_0 hereafter). The values of α_0 are between 1.28 and 1.34 for all the mass bins.

In practice, the empirical R_{GC} in Fig. 7 are only estimates of R_{GC}^{th} from finite samples. Before we divided the bins into subsets and repeated this experiment (with final numbers that are reported in the main text), we made sure that all the galaxy-mass bins that we considered contain more than 100 GCs after stacking. For Sérsic distributions with $n \leq 2$, 100 GCs provide an estimate of the effective radius (i.e. R_{GC}) with a standard deviation smaller than $\pm 15\%$ (checked numerically, and also in agreement with the width of the shaded areas in the top panel of Fig. 7). Therefore, the values of α_0 are representative of the values of α that apply in our bins. We can simulate the effect of finite samples by assigning to the galaxies of one bin somewhat dispersed effective radii for their GC distributions.

Appendix D: Comparison with the literature

Here we provide a comparison between the GC number counts (N_{GC}) and GC distributions (R_{GC} and R_{GC}/R_e) of dwarf galaxies in this work versus the values published in Marleau et al. (2025) and Janssens et al. (2024). Generally we find a good

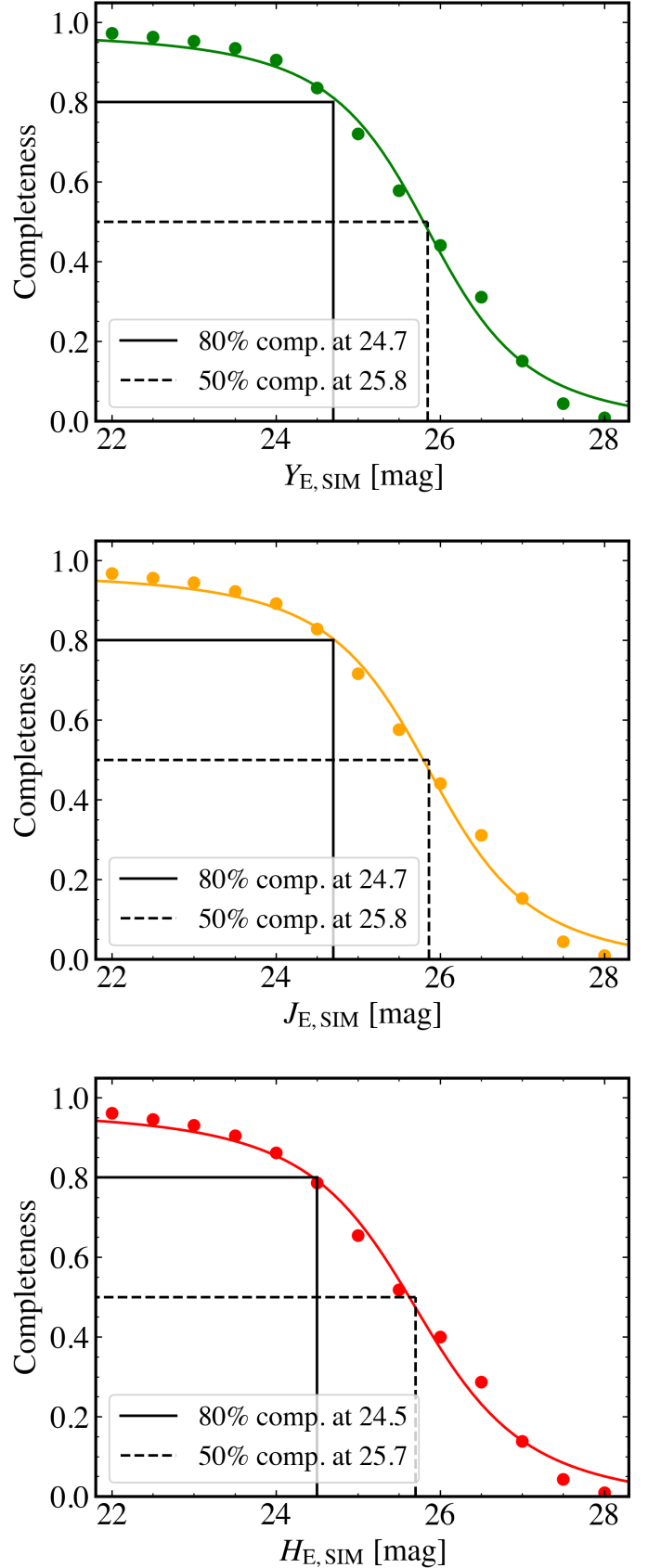


Fig. B.1. Completeness of the performed source detection in Y_E , J_E , and H_E based on artificial GC tests. Our assessment shows that the GC detection is 80% complete at magnitude 24.7, 24.7, and 24.5 in Y_E , J_E , and H_E , respectively.

Table A.1. Details and the output of the analysis of the stacked GC distributions.

Dwarf sample	Parameter	M_* bin						
		$10^6 M_\odot$	$10^{6.5} M_\odot$	$10^7 M_\odot$	$10^{7.5} M_\odot$	$10^8 M_\odot$	$10^{8.5} M_\odot$	$10^9 M_\odot$
All dwarfs	N_{dwarfs}	243	377	426	339	211	112	54
	$N_{\text{GC,cand,gal}}$	104	777	1252	1911	2229	2675	2422
	$N_{\text{GC,cand,back}}$	37.1	550.6	890.1	1422.7	1739.3	2090.2	1776.4
	R_{GC}	0.20	0.48	0.73	1.22	1.67	2.72	3.53
	$R_{\text{GC}}/\langle R_e \rangle$	0.33	0.70	0.87	1.10	1.20	1.67	1.70
HSB	N_{dwarfs}	112	187	218	165	108	59	25
	$N_{\text{GC,cand,gal}}$	131	235	572	1805	1758	1575	1351
	$N_{\text{GC,cand,back}}$	90.3	158.3	499.1	1625.3	1450.6	1214.3	1016.3
	R_{GC}	0.21	0.42	0.89	1.60	2.10	2.91	3.38
	$R_{\text{GC}}/\langle R_e \rangle$	0.43	0.75	1.25	1.74	1.81	1.98	1.95
LSB	N_{dwarfs}	130	188	206	173	102	52	28
	$N_{\text{GC,cand,gal}}$	54	421	691	894	699	899	1159
	$N_{\text{GC,cand,back}}$	18.2	273.9	404.8	582.7	478.1	686.7	845.3
	R_{GC}	0.14	0.43	0.71	1.00	1.38	2.22	3.55
	$R_{\text{GC}}/\langle R_e \rangle$	0.21	0.53	0.71	0.75	0.86	1.12	1.62
COMP	N_{dwarfs}	119	197	220	158	101	56	24
	$N_{\text{GC,cand,gal}}$	44	140	721	1000	1036	1680	1227
	$N_{\text{GC,cand,back}}$	8.6	72.3	620.8	850.5	819.8	1359.5	982.5
	R_{GC}	0.12	0.23	0.67	1.50	1.90	3.08	3.54
	$R_{\text{GC}}/\langle R_e \rangle$	0.29	0.46	1.12	1.92	1.88	2.52	2.36
DIFF	N_{dwarfs}	–	178	204	180	109	55	29
	$N_{\text{GC,cand,gal}}$	–	455	655	1053	984	1297	1309
	$N_{\text{GC,cand,back}}$	–	307.8	378.1	712.8	713.6	1002.1	905.2
	R_{GC}	–	0.67	0.79	1.14	1.49	2.54	3.39
	$R_{\text{GC}}/\langle R_e \rangle$	–	0.67	0.67	0.81	0.86	1.10	1.23
UDG	N_{dwarfs}	–	47	53	33	13	4	–
	$N_{\text{GC,cand,gal}}$	–	111	439	591	208	67	–
	$N_{\text{GC,cand,back}}$	–	36.7	267.9	427.7	156.6	37.6	–
	R_{GC}	–	1.02	1.22	2.21	1.63	1.72	–
	$R_{\text{GC}}/\langle R_e \rangle$	–	0.52	0.63	1.13	0.61	0.64	–
non-UDG	N_{dwarfs}	235	353	376	284	176	97	49
	$N_{\text{GC,cand,gal}}$	114	797	1168	1846	1652	2915	2517
	$N_{\text{GC,cand,back}}$	47.7	586.1	883.3	1499.4	1339.8	2382.7	1882.1
	R_{GC}	0.19	0.48	0.73	1.27	1.52	2.97	3.69
	$R_{\text{GC}}/\langle R_e \rangle$	0.34	0.73	0.95	1.28	1.20	1.97	1.81
Nucleated	N_{dwarfs}	200	291	291	190	115	71	35
	$N_{\text{GC,cand,gal}}$	79	442	708	898	1925	2529	1767
	$N_{\text{GC,cand,back}}$	30.1	316.2	533.8	736.3	1662.3	2037.6	1325.5
	R_{GC}	0.20	0.36	0.69	1.09	2.32	3.56	3.87
	$R_{\text{GC}}/\langle R_e \rangle$	0.38	0.58	0.98	1.19	1.91	2.42	2.08
Non-nucleated	N_{dwarfs}	41	76	112	108	56	18	6
	$N_{\text{GC,cand,gal}}$	31	207	355	949	478	233	260
	$N_{\text{GC,cand,back}}$	12.1	121.8	225.1	749.8	386.5	208.8	199.8
	R_{GC}	0.10	0.69	0.76	1.61	1.48	1.74	2.60
	$R_{\text{GC}}/\langle R_e \rangle$	0.14	0.76	0.74	1.37	1.06	0.95	1.05

Notes. The table is divided into nine parts for different categories of dwarf galaxies (column 1). Each part provides details/results for a given category, specified in column 2. These parameters are: number of dwarf galaxies in a given bin (N_{dwarf}), total number of GC candidates within $3 R_{\text{GC}}$ of the dwarf galaxies in the sample ($N_{\text{GC,cand,gal}}$), number of background sources beyond $3 R_{\text{GC}}$ and normalised to the area within $3 R_{\text{GC}}$ of the dwarf galaxies in the sample ($N_{\text{GC,cand,back}}$), the estimated GC half-number radius (R_{GC} , in kpc) and the ratio between GCs half-number radius and the average half-light radius of the dwarf sample ($R_{\text{GC}}/\langle R_e \rangle$). Columns 3 to 9 present the values corresponding to seven stellar mass bins used for the analysis of the stacked GC distributions, ranging from 10^5 to $10^9 M_\odot$. The value associated with each column corresponds to the centre of the bin. Stellar mass bins have a width of 1 dex.

agreement with their results. Comparing our results with N_{GC} from Marleau et al. (2025), as is seen in Fig. D.1 on the top, on average, we find a larger N_{GC} . In this work we used the same *Euclid* ERO data as Marleau et al. (2025) and with a different

methodology (e.g. by including colours in the GC selection procedure) for GC selection that leads to a different GC catalogue. Additionally, we adopt different approach for estimating N_{GC} , mainly, a fainter GCLF TOM and varying R_{GC} as a function

Table A2. Table of the GC properties of the dwarfs galaxies in this work.

ID	$M_*/M_\odot$	N_{GC}	R_{GC}	R_{GC}/R_e
(1)	(2)	(3)	(4)	(5)
EDwC-0017	7.95	15.5 ± 3.5	$2.68^{+1.78}_{-1.40}$	$1.49^{+0.99}_{-0.77}$
EDwC-0023	8.64	35.8 ± 5.6	$0.87^{+0.43}_{-0.33}$	$0.40^{+0.20}_{-0.15}$
EDwC-0026	8.706	96.2 ± 8.1	$3.11^{+0.53}_{-0.52}$	$1.51^{+0.26}_{-0.25}$
EDwC-0046	8.26	18.5 ± 5.6	$1.21^{+0.26}_{-0.27}$	$0.82^{+0.17}_{-0.18}$
EDwC-0082	7.76	8.8 ± 2.2	$1.04^{+0.62}_{-0.60}$	$0.54^{+0.32}_{-0.31}$
EDwC-0086	7.39	24.8 ± 3.6	$1.91^{+0.73}_{-0.72}$	$1.08^{+0.41}_{-0.41}$
EDwC-0089	9.00	50.9 ± 7.0	$2.93^{+0.63}_{-0.73}$	$1.07^{+0.23}_{-0.27}$
EDwC-0112	7.84	11.5 ± 2.6	$1.14^{+0.79}_{-0.73}$	$0.58^{+0.40}_{-0.37}$
EDwC-0120	7.94	41.6 ± 4.6	$1.77^{+0.59}_{-0.72}$	$0.82^{+0.27}_{-0.33}$
EDwC-0156	8.82	38.4 ± 6.8	$0.86^{+0.98}_{-0.46}$	$0.37^{+0.42}_{-0.20}$
...	...	...	...	...

Notes. Column 1 present the *Euclid* ID (name) of the dwarf galaxies from [Marleau et al. \(2025\)](#). Column 2 presents the stellar mass of dwarf galaxies from [Cuillandre et al. \(2025b\)](#). Columns 3, 4, and 5 present the GC properties of the dwarf galaxies, as estimated in this work. The full table is available at the Strasbourg astronomical Data Center (CDS) as supplementary material.

of stellar mass and central surface brightness of dwarf galaxies. Regarding the values R_{GC} and R_{GC}/R_e values for five dwarf galaxies, as is seen in Fig. D.1 on the bottom, our analysis leads to, on average, lower values compared to [Janssens et al. \(2024\)](#); however, for the two most GC-rich dwarf galaxies in the sample, measurements from two works are consistent with each other. For other three dwarf galaxies, the difference mainly arises from low number statistics when dealing with GCs in these dwarf galaxies.

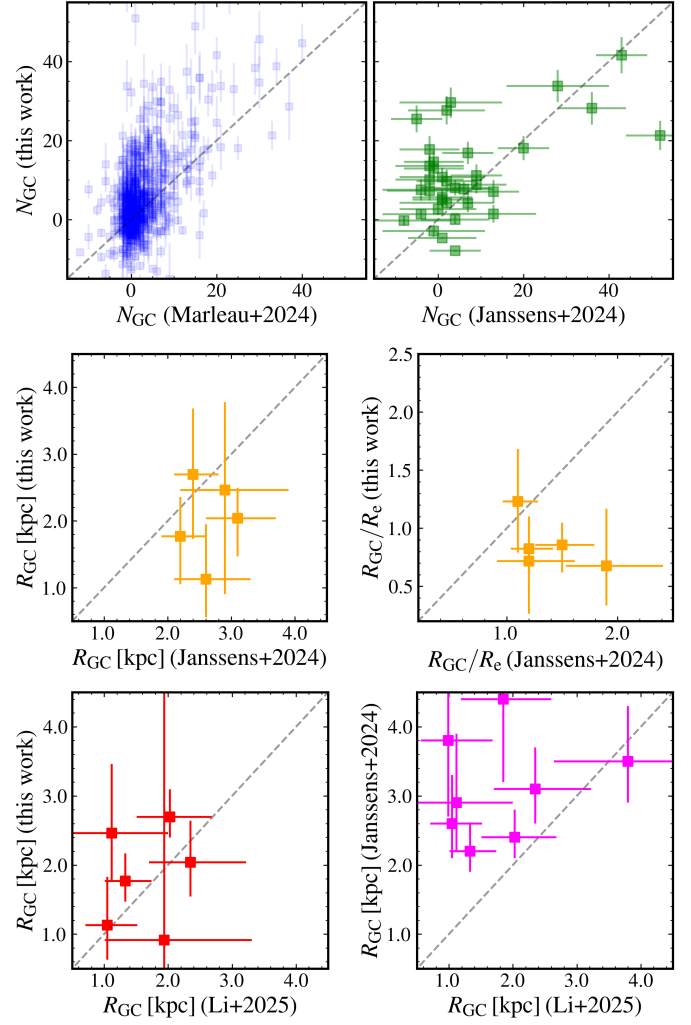


Fig. D.1. Comparison between GC properties in this work and the literature. Top: Number counts estimated in this work versus [Marleau et al. \(2025\)](#) and [Janssens et al. \(2024\)](#). Middle: R_{GC} and R_{GC}/R_e estimated in this work and [Janssens et al. \(2024\)](#). Bottom: R_{GC} values estimated in [Li et al. \(2025\)](#) versus this work, and [Janssens et al. \(2024\)](#).