FilDReaMS

II. Application to the analysis of the relative orientations between filaments and the magnetic field in four Herschel fields

J.-S. Carrière®, K. Ferrière®, I. Ristorcelli, and L. Montier

Abstract

Context. Both simulations and observations of the interstellar medium show that the study of the relative orientations between filamentary structures and the magnetic field can bring new insight into the role played by magnetic fields in the formation and evolution of filaments and in the process of star formation.

Aims. We provide a first application of FilDReaMS, the new method presented in the companion paper to detect and analyze filaments in a given image. The method relies on a template that has the shape of a rectangular bar with variable width. Our goal is to investigate the relative orientations between the detected filaments and the magnetic field.

Methods. We apply FilDReaMS to a small sample of four Herschel fields (G210, G300, G82, G202) characterized by different Galactic environments and different evolutionary stages. First, we look for the most prevalent bar widths, and we examine the networks formed by filaments of different bar widths as well as their hierarchical organization. Second, we compare the filament orientations to the magnetic field orientation inferred from Planck polarization data and, for the first time, we study the statistics of the relative orientation angle as functions of both spatial scale and H$_2$ column density.

Results. We find preferential relative orientations in the four Herschel fields: small filaments with low column densities tend to be slightly more parallel than perpendicular to the magnetic field; in contrast, large filaments, which all have higher column densities, are oriented nearly perpendicular (or, in the case of G202, more nearly parallel) to the magnetic field. In the two nearby fields (G210 and G300), we observe a transition from mostly parallel to mostly perpendicular relative orientations at an H$_2$ column density $\equiv 1.1 \times 10^{21}$ cm$^{-2}$ and $1.4 \times 10^{21}$ cm$^{-2}$, respectively, consistent with the results of previous studies.

Conclusions. Our results confirm the existence of a coupling between magnetic fields at cloud scales and filaments at smaller scale. They also illustrate the potential of combining Herschel and Planck observations, and they call for further statistical analyses with our dedicated method.

Key words. ISM: clouds – ISM: structures – ISM: magnetic fields – dust – techniques: image processing

1. Introduction

The physical processes involved in the early stages of star formation are still poorly understood. There is growing evidence that magnetic fields, together with turbulence and gravity, contribute significantly to the formation of stars, but their respective roles at different spatial and temporal scales are not clearly established (McKee & Ostriker 2007; Dobbs et al. 2014). Making progress requires in-depth observational studies of the formation and evolution of dense structures and star-forming regions. Filamentary structures are observed to be ubiquitous in all Galactic (neutral) environments. Moreover, these observations show that the many detected pre-stellar cores seem to be predominantly distributed in the densest filaments (Polychroni et al. 2013; Könyves et al. 2015; Montillaud et al. 2015). The question of the origin and evolution of dense cores is thus strongly connected to that of filaments (see review by Andrè et al. 2014).

Filamentary structures are also observed to be dominant features in numerical simulations of magnetohydrodynamic (MHD) turbulence, indicating the potential influence of magnetic fields in their formation (see review by Hennebelle & Inutsuka 2019). Furthermore, simulations make it clear that studying the relative orientation between magnetic fields and filaments can shed light on the role played by magnetic fields in the formation and evolution of filaments (see for instance Soler & Hennebelle 2017; Wu et al. 2017). In particular, the simulations of Soler et al. (2013) showed that in a weakly magnetized medium, the magnetic field is preferentially oriented parallel to the density structures, whereas in a strongly magnetized medium, there is a change in relative orientation from parallel to perpendicular at a critical density (see also Chen et al. 2016).

A number of observational studies reveal preferential relative orientations between magnetic fields and filaments, as initially inferred from starlight polarization (e.g., Goodman et al. 1990; Pereya & Magalhães 2004; Sugitani et al. 2011; Palmeirim et al. 2013; Li et al. 2013; Clark et al. 2014). More recently, the statistical analysis based on the Planck survey of dust polarization evidence showed that elongated structures are predominantly aligned parallel to the magnetic field in the diffuse Galactic environment.
(neutral) medium, while they are mostly perpendicular in dense molecular clouds (Planck Collaboration XXXII 2016a; Planck Collaboration XXXV 2016b), with a transition at a column density \( N_\text{HI} \approx 10^{10.7} \text{ cm}^{-2} \) (Planck Collaboration XXXV 2016b). A similar trend (low-\(N_\text{HI}\) filaments mostly parallel and high-\(N_\text{HI}\) filaments mostly perpendicular to the magnetic field) within individual clouds, with a similar transition column density, were also found by Malinen et al. (2016); Cox et al. (2016); Soler (2019), who compared the orientations of the plane-of-sky (PoS) magnetic field traced by Herschel (36′ resolution) in a sample of nearby molecular clouds. These results provide evidence that a coupling exists between magnetic fields at cloud scales and filaments at smaller scales (Soler 2019). They also demonstrate the benefit of combining Planck and Herschel data sets despite their different angular resolutions.

This benefit, in turn, calls for further statistical analyses based on Planck and Herschel data together to explore relative orientation trends towards star-forming regions in different Galactic environments. The Herschel Galactic cold core (GCC) key-program (Juvela et al. 2010, 2012), which includes 116 Galactic fields (\(\sim 40 \times 40\) size on average) hosting Planck clumps (Planck catalogue of Galactic cold clumps (PGCC) from Planck Collaboration XXVIII 2016c) is particularly well suited for such a statistical study. It covers a wide range of Galactic positions, environments, physical conditions, and evolutionary stages (Montillaud et al. 2015). Most of the fields display filaments over a broad range of column densities (Juvela et al. 2012), often including low-density strata (as in the GCC field L1642 studied by Malinen et al. 2016). Future statistical analyses of relative orientations will require an efficient and robust method that makes it possible to extract filaments over a broad range of sizes and column densities and to determine their orientations.

The purpose of the companion paper (Carrière et al. 2022, hereafter Paper I) was precisely to develop such a method, which we dubbed FilDReaMS (Filament Detection & Reconstruction at Multiple Scales). Our purpose here is to apply FilDReaMS to a first selection of four Herschel fields and to study the properties of the detected filaments, with a special attention to their relative orientations to the magnetic field.

In Sect. 2, we present the data used for our study. In Sect. 3, we review our new FilDReaMS method, giving just enough details for its application. In Sect. 4, we apply FilDReaMS to a sample of four Herschel fields selected amongst the 116 fields of the Herschel-GCC program. In Sect. 5, we provide a detailed discussion of our results. In Sect. 6, we summarize and conclude our study.

### 2. Data

We will apply our new FilDReaMS method to a sample of four molecular clouds, for which we will exploit the combination of Planck-HFI and Herschel data. In brief, Planck-HFI observations at 353 GHz provide a whole-sky map of the dust polarized emission with an angular resolution of 4.7′; from this map, one can infer (a dust-weighted line-of-sight average of) the orientation of the PoS component of the magnetic field. Meanwhile, Herschel observations provide dust emission maps at several wavelengths (with an angular resolution of 18′ at 250 μm); based on these maps, column density maps have been constructed with an angular resolution of 36′.

### 2.1. Planck-HFI

The polarized emission measured by Planck-HFI at 353 GHz is dominated by the thermal emission from Galactic dust (Planck Collaboration XIX 2015). This emission, which is polarized perpendicular to the local magnetic field, provides a good probe of the magnetic field orientation in the densest regions of the interstellar medium (ISM). The magnetic field PoS orientation angle, \(\psi_0\), can be written in terms of the Stokes parameters, \(Q\) and \(U\), as

\[
\psi_0 = \frac{1}{2} \arctan \left( \frac{5}{3} \frac{U}{Q} \right) \pm 90^\circ, \tag{1}
\]

where \(\arctan\) is the two-argument arctangent function defined from \(-180^\circ\) to \(180^\circ\). Since polarization data do not give access to the magnetic field direction, but only to its orientation, we may arbitrarily require that \(\psi_0\) must lie in the range \([-90^\circ, +90^\circ]\); the \(+\) or \(-\) sign in the last term of Eq. (1) is then chosen accordingly. Here, we adopt the IAU polarization convention in Galactic coordinates, such that \(\psi_0\) increases counterclockwise from Galactic north. Since this is opposite to the Healpix convention used in the Planck community, we have to change the sign of \(U\) taken from the Planck Legacy Archive\(^1\).

We extracted from the all-sky Planck \(Q\) and \(U\) maps at 353 GHz a set of four much smaller maps (2′ × 2′, approximately twice the size of Herschel maps), each centered on one of our four selected Herschel fields. We smoothed the original Planck \(Q\) and \(U\) maps from 4.7′ to 7′ resolution. The smoothing of these maps and their uncertainties was performed as described in Appendix of Planck Collaboration XIX (2015). Our choice of a 7′ resolution results from a compromise between the need to reduce the integration time required to sample the magnetic field in a given area and the desire to access the smallest possible magnetic field fluctuations. The typical improvement of polarization SNR when smoothing from 4.7′ to 7′ is roughly a factor of \(7′/4.7′ \approx 1.5\); in the case of our four fields, the fraction of pixels with SNR <3 in the Planck \(Q\) and \(U\) maps smoothed to 7′ is \(\lesssim 1\%\).

The uncertainty in the polarization angle was computed as described in the appendix of Planck Collaboration XIX (2015). Following Malinen et al. (2016), we only retained pixels with uncertainty \(<10^\circ\). It is difficult to directly estimate the impact of our smoothing on the inferred polarization angles, because the SNR at 4.7′ resolution is not high enough. However, we verified that smoothing from 7′ to 10′ (two resolutions leading to SNR >3 in almost all pixels) had very little impact on the derived polarization angles. We are, therefore, confident that our results will not be significantly affected by our choice of resolution for the Planck \(Q\) and \(U\) maps.

### 2.2. Herschel

#### 2.2.1. Herschel data

The Herschel satellite (Pilbratt et al. 2010) provided a fantastic probe of filaments within molecular clouds at a much better angular resolution than that of Planck. The Herschel open-time key program Galactic Cold Cores (Juvela et al. 2010), which we will refer to as the Herschel-GCC program, was designed to map out a sample of cold regions of interstellar clouds previously detected in the Planck all-sky survey. This follow-up is composed of 116 fields (with a typical size of 40′) observed

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\(^1\) Planck Legacy Archive: https://www.cosmos.esa.int/web/planck/pla
with PACS and SPIRE (Poglitsch et al. 2010; Griffin et al. 2010), which were selected using the Planck Catalogue of Galactic cold clumps (PGCC, Planck Collaboration XXVIII 2016c) to cover a diversity of clump physical properties and Galactic environments (Juvela et al. 2012). The fields are presented in detail in Montillaud et al. (2015), who built a catalogue of ≈4000 cold and compact sources, identifying their various evolutionary stages from gravitationally unbound to prestellar and protostellar cores.

In this study, we use the H$_2$ column density maps presented in Montillaud et al. (2015) and derived from the dust spectral energy distribution (SED) fit based on the 250, 350, and 500 µm SPIRE maps. The authors assumed a fixed emissivity spectral index $\beta=2.0$, with a dust opacity $k = 0.1$ cm$^2$/g($\nu/1000$ GHz)$^\beta$. The $N_{H_2}$ maps have an angular resolution of 36″.

2.2.2. Sample of four fields from the Herschel-GCC program

We focus on four Herschel fields out of the 116 fields of the Herschel-GCC program. Their characteristics are summarized in Table 1, while their H$_2$ column density maps are displayed in Fig. 1. The reason why we selected these four molecular clouds is because they present a good diversity, which allows us to explore various Galactic environments and filament morphologies. Furthermore, they are already well-known and benefit from both observations with multiple tracers and various in-depth analyses carried out in the past decade, which provided us with information concerning their evolutionary stage or their dynamics (see references below).

The G210 molecular cloud is located at high Galactic latitude ($|b| \approx 36°$) in a diffuse environment. It was studied by Malinen et al. (2014, 2016), who found an unusually high star formation rate (SFR) considering its low mass, as well as a complex interplay between the large-scale magnetic field and the cloud structure.

The G300 cloud lies at medium Galactic latitude ($|b| \approx 9°$) inside the well-known Musca filament. It was part of a statistical analysis by Arzoumanian et al. (2019), and it was studied in greater detail by Cox et al. (2016) and Kainulainen et al. (2016). It was found to be a filament at a very early stage of evolution with an SFR ≈1% (Cox et al. 2016), and it was shown to be mainly thermally super-critical and fragmenting faster than the center of the Musca filament (fragments 4 and 5 in Kainulainen et al. 2016).

The G82 cloud is located close to the Galactic plane ($|b| \approx 2°$). It was found by Saajasto et al. (2017) to be a filament at a late stage of evolution, and it was described as a “debris filament” already forming several cold cores and young stellar objects.

Finally, the G202 cloud, also close to the Galactic plane ($|b| \approx 3°$), is part of the Monoceros B molecular complex. It was studied in great detail by Montillaud et al. (2019a); Alina et al. (2020). They both found this filament to be thermally super-critical. Montillaud et al. (2019a,b) also found that this filament is actively forming stars, especially in its center, the “junction” region, which lies between two colliding filaments and represents $\approx$60% of the cloud mass with a high SFR $\approx$20 – 38%.

### Table 1. Basic parameters of the H$_2$ column density maps of the four Herschel fields displayed in Fig. 1.

<table>
<thead>
<tr>
<th>Field</th>
<th>$l$</th>
<th>$b$</th>
<th>$d$</th>
<th>$\Delta l \times \Delta b$</th>
<th>$N_{H_2}$</th>
</tr>
</thead>
<tbody>
<tr>
<td>G210</td>
<td>210.90°</td>
<td>-36.55°</td>
<td>140 ± 20 pc</td>
<td>$1.28^\circ \times 1.22^\circ \rightarrow 3.1^\circ \times 3.0^\circ$</td>
<td>$[0.2, 7.5] \times 10^{21}$ cm$^{-2}$</td>
</tr>
<tr>
<td>G300</td>
<td>300.86°</td>
<td>-9.00°</td>
<td>150 ± 30 pc</td>
<td>$0.64^\circ \times 0.68^\circ \rightarrow 1.7^\circ \times 1.8^\circ$</td>
<td>$[0.6, 8.2] \times 10^{21}$ cm$^{-2}$</td>
</tr>
<tr>
<td>G82</td>
<td>82.65°</td>
<td>-2.00°</td>
<td>620^±31 pc</td>
<td>$0.036^\circ$</td>
<td>$[2.7, 42] \times 10^{21}$ cm$^{-2}$</td>
</tr>
<tr>
<td>G202</td>
<td>202.02°</td>
<td>+2.85°</td>
<td>760 ± 100 pc</td>
<td>$1.22^\circ \times 1.17^\circ \rightarrow 16.2^\circ \times 15.6^\circ$</td>
<td>$[0.6, 58] \times 10^{21}$ cm$^{-2}$</td>
</tr>
</tbody>
</table>

**Notes.** Parameters are ordered by increasing distance: Galactic longitude, $l$, Galactic latitude, $b$, distance $d$ (taken from Saajasto et al. 2017 for G82 and Montillaud et al. 2015 for the other fields), angular and physical sizes of a pixel (equal to one-third of the beam size), angular and physical dimensions of the map, $\Delta l \times \Delta b$, and range of the H$_2$ column density, $N_{H_2}$.

3. FilDReaMS in a nutshell

3.1. Overview of the method

The purpose of FilDReaMS (described in details in Paper I) is to detect and characterize filamentary structures in an image, which in the present paper is a Herschel (intensity or column-density) map. To that end, FilDReaMS resorts to a special template, which has the shape of a rectangular bar (referred to as the model bar) of width $W_b$, length $L_b$, and aspect ratio $r_b = L_b/W_b$. In the following astrophysical applications, we adopt $r_b = 3$ (Panopoulou et al. 2014; Arzoumanian et al. 2019) and we consider values of $W_b$ spanning the range $[W_b]_{min} = 5 \times [W_b]_{max}$, with $[W_b]_{max} = 5$ px and $[W_b]_{max}$ equal to one-ninth the size of the Herschel map. The orientation angle of the model bar, $\psi_b$, follows the same convention as the magnetic field orientation angle, $\psi_b$ (Eq. (1)), such that $\psi_b$ is defined in the range $[-90°, +90°]$ and increases counterclockwise from Galactic north. The same will hold true for the orientation angle of a filament, $\psi_f$, defined in the next paragraph.

For any given value of $W_b$, FilDReaMS filters out structures broader than $W_b$ in the initial image and converts the filtered image into a binary map. At each pixel $i$ of the binary map, FilDReaMS considers a model bar centered on $i$, derives the bar orientation angle that provides the best match to the binary map, $(\psi_b)_i$, and computes the corresponding significance, $S_i$, based on a comparison with an ideal case. If $S_i > 1$, FilDReaMS concludes that a significant filament with orientation angle $(\psi_f)_i = (\psi_b)_i$ is detected at pixel $i$. The true shape of this filament is then reconstructed from the binary map and its intensity from the initial image. Iterating over all the pixels of the binary map (minus a band of width $L_b/2$ adjacent to the border) and superposing all the associated reconstructed filaments yields the
entire network of physical filaments of bar width $W_b$. Each pixel $i'$ of this network is assigned a filament orientation angle, $(\psi_f^*)_{i'}$, equal to the orientation angle of the most significant filament (filament with the highest $S_i$) whose model bar contains $i'$.

The procedure is repeated for different values of $W_b$, each of which leads to a new network of filaments. Each pixel $i'$ of the different networks is assigned a most significant bar width, $(W_b^*)_{i'}$, equal to the bar width of the most significant filament whose model bar contains $i'$. Finally, the histogram of $W_b^*$ over all pixels makes it possible to identify the most prevalent bar widths for the entire map, $W_b^{\text{peak}}$. In the following, to make it easier to compare the different Herschel fields, we will work with the normalized histogram $(N_{\text{pix}}/N_{\text{map}})$ versus $W_b^*$, where $N_{\text{pix}}$ is the number of pixels whose most significant bar width is $W_b^*$ and $N_{\text{map}}$ is the total number of pixels in the map.

For convenience, the meaning of all the parameters introduced above can be found in Table 2.

### 3.2. Applications

A first obvious application of FilDReaMS concerns the bar widths of filaments, with, in particular, the idea of uncovering the most prevalent bar widths. Related goals are the visualization of the network formed by filaments of a given bar width and the examination of the spatial connections between filaments of different bar widths. All these, in turn, may provide some insight into the process of filament formation.

A second, very important application of FilDReaMS, which we will devote most of our attention to, concerns the orientations of filaments relative to the local magnetic field. The idea is to compare the filament orientation angles, $(\psi_f^*)_{i'}$, derived with FilDReaMS to the magnetic field PoS orientation angle, $\psi_B$, inferred from the dust polarized emission observed by Planck (Eq. (1)). The natural quantity to work with in this context is the relative orientation angle, $(\psi_f^* - \psi_B)$. For convenience, since both $\psi_f^*$ and $\psi_B$ are defined in the range $[-90^\circ, +90^\circ]$ (see Sects. 3.1 and 2.1, respectively), we require that $(\psi_f^* - \psi_B)$ must also lie in the range $[-90^\circ, +90^\circ]$; if needed, this can be achieved by adding or subtracting $180^\circ$. FilDReaMS will enable us to study, for the first time, the statistics of $(\psi_f^* - \psi_B)$ as functions of spatial scale (Sect. 4.2) and as functions of both spatial scale and $H_2$ column density (Sect. 4.3).

Following previous studies (Soler et al. 2013; Planck Collaboration XXXII 2016a; Planck Collaboration XXXV 2016b; Soler 2019), we will construct histograms of relative orientation (HROs), which are just histograms of $(\psi_f^* - \psi_B)$, similar...
to histograms \((N_{\text{pix}}/N_{\text{map}})\) versus \((\psi_{\text{f}} - \psi_{\text{B}})\). The novelty is that we will do so in three separate ranges of bar widths. We will also construct 2D histograms of \((\psi_{\text{f}} - \psi_{\text{B}})\) as functions of \(H_2\) column density, \(N_{\text{H}_2}\) (i.e., 2D histograms \((N_{\text{pix}}/N_{\text{map}})\) versus \((N_{\text{H}_2},(\psi_{\text{f}} - \psi_{\text{B}}))\)) again in three separate ranges of bar widths. This will enable us to explore the idea of a bi-modal distribution of \((\psi_{\text{f}} - \psi_{\text{B}})\), with a transition at a certain \(N_{\text{H}_2}\). To be more quantitative, we will resort to the multivariate analysis methods presented by Malinen et al. (2016), namely, a non-negative matrix factorization (NMF, Lee & Seung 1999) and a principal component analysis (PCA, Timmerman 2003). NMF makes it possible to derive the first (in our case, the first two) principal components of the 2D histograms, together with their respective weights (i.e., their relative contributions) in each \(N_{\text{H}_2}\) bin. PCA estimates the capability to reconstruct the entire 2D histograms by including only the considered (in our case, the first two) principal components. This reconstruction capability quantifies the relevance of a bi-modal relative orientation.

### 4. Results

We now present and analyze the results obtained with FilDReaMS for a sample of four Herschel fields. We focus on two important characteristics of the reconstructed filaments: their bar widths, \(W_b\) (Sect. 4.1), and their orientations relative to the magnetic field, with a detailed discussion of how the relative orientation angles, \((\psi_{\text{f}} - \psi_{\text{B}})\), statistically vary with bar width (Sect. 4.2) and with column density (Sect. 4.3). In each subsection, we start with the G82 field, which exhibits a rich filamentary network and covers the broadest range of scales; we continue with the G210 field (only in Sects. 4.2 and 4.3), which shows interesting new features; and we finish with the general ammonium and disconnected filaments can form striation patterns, faint and without their uncertainties. FilDReaMS allows us to detect a wide variety of filaments, mostly visible with the naked eye in the initial Herschel G82 map (left panel of Fig. 2). Filaments are detected in all areas of the map, over a broad range of sizes and column densities, with \(N_{\text{H}_2}\) varying from \(\lesssim 10^{21} \text{ cm}^{-2}\) to \(\approx 4 \times 10^{22} \text{ cm}^{-2}\).

All the reconstructed filaments are plotted in the middle and right panels of Fig. 2. Because many pixels belong to several filaments (usually filaments of different bar widths), it is not possible to visualize all the filaments in a single plot. Therefore, we show two extreme views, in which filaments of different bar widths are superposed, either with the smallest filaments in the background and the largest in the foreground (middle panel) or vice-versa (right panel). A fraction of the smaller filaments appear to be spatially connected to larger filaments, forming either crests, internal sub-structures (such as strands), or ramifications (which appear to emerge from, or merge with, the larger filaments). The rest of the smaller filaments are disconnected from larger filaments. Both ramifications and disconnected filaments can form striation patterns, faint and periodic structures similar to those detected by Goldsmith et al. (2008) and Narayanan et al. (2008) in the diffuse 13CO emission map from the Taurus molecular cloud. Crescents are roughly parallel to their parent filaments, while strands are observed at all angles. Ramifications tend to connect with their parent filaments at large (roughly perpendicular) angles. Disconnected small filaments are found at all angles, although in the vicinity of a large filament they tend to line up roughly parallel or perpendicular to this large filament.

### 4.1. Filament bar widths

#### 4.1.1. G82

Following the guidelines given in Sect. 3.1, we restrict the values of \(W_b\) to the range \([5, 30]\) px. The corresponding range in parsecs can be derived using the relation

\[
W_b^{[\text{pc}]} = d \gamma W_b^{[\text{px}]},
\]

where \(d\) is the distance of the Herschel field (in pc) and \(\gamma\) the angular size of a pixel (in rad). With \(d = 620_{-31}^{+54}\) pc and \(\gamma = 12''\) (see Table 1), the range in parsecs is \([0.18, 1.08]\) pc. The uncertainty in \(W_b^{[\text{pc}]}, \Delta W_b^{[\text{pc}]},\) is related to the uncertainty in \(W_b^{[\text{px}]}, \Delta W_b^{[\text{px}]},\) \(1\) px, and the uncertainty in \(d, \Delta d,\) through

\[
\Delta W_b^{[\text{pc}]} = W_b^{[\text{pc}]} \left( \frac{\Delta d}{d} + \frac{\Delta W_b^{[\text{px}]}}{W_b^{[\text{px}]}} \right).
\]

Hence, with \(\Delta d/d = 6\%\) and \(\Delta W_b^{[\text{px}]}/W_b^{[\text{px}]},\) ranging from \(1\) px/30 px \(\approx 3\%\) to \(1\) px/5 px \(\approx 20\%,\) we find that \(\Delta W_b^{[\text{pc}]}/W_b^{[\text{pc}]},\) ranges from \(\lesssim 3\%\) to \(19\%.\) From now on, for simplicity, we give the values of \(W_b^{[\text{pc}]},\) without their uncertainties.

Table 2. List of all the symbols used in the paper.

<table>
<thead>
<tr>
<th>Symbol</th>
<th>Description</th>
</tr>
</thead>
<tbody>
<tr>
<td>(i)</td>
<td>Index of the considered pixel in the binary map</td>
</tr>
<tr>
<td>(i')</td>
<td>Index of the considered pixel in the map of reconstructed filaments</td>
</tr>
<tr>
<td>(\psi_{\text{B}})</td>
<td>Orientation angle of the PoS component of the magnetic field</td>
</tr>
<tr>
<td>(W_b)</td>
<td>Width of the model bar</td>
</tr>
<tr>
<td>(L_b)</td>
<td>Length of the model bar</td>
</tr>
<tr>
<td>(r_b = L_b/W_b)</td>
<td>Aspect ratio of the model bar</td>
</tr>
<tr>
<td>(\psi_{\text{f}})</td>
<td>Orientation angle of the model bar</td>
</tr>
<tr>
<td>(S)</td>
<td>Significance of a filament detection</td>
</tr>
<tr>
<td>(\psi_{\text{f}})</td>
<td>Orientation angle of a filament</td>
</tr>
<tr>
<td>(\psi_{\text{f}}^{*})</td>
<td>Orientation angle of the most significant filament</td>
</tr>
<tr>
<td>(W_b^{*})</td>
<td>Bar width of the most significant filament</td>
</tr>
<tr>
<td>(N_{\text{pix}})</td>
<td>Number of pixels whose most significant bar width is (W_b^{*})</td>
</tr>
<tr>
<td>(N_{\text{map}})</td>
<td>Total number of pixels in the map</td>
</tr>
<tr>
<td>(W_b^{\text{peak}})</td>
<td>Most prevalent bar width for the entire map</td>
</tr>
</tbody>
</table>
The histogram \( N_{\text{pix}} / N_{\text{map}} \) versus \( W_b^* \) (defined in Sect. 3.1) is displayed in the top panel of Fig. 3. Five peaks clearly emerge at \( W_b^* = 5 \) px, 9 px, 20 px, 25 px, and 30 px, corresponding to 0.18 pc, 0.32 pc, 0.72 pc, 0.90 pc, and 1.08 pc, respectively. These peaks represent the most prevalent bar widths. The peak at \( W_b^* = 5 \) px is probably partly artificial, because it arises at the lower boundary of the \( W_b^* \) range and, therefore, includes not only the contribution from filaments with \( W_b^* = 5 \) px, but also the contribution from all the enclosed smaller filaments (not considered in this study). However, the fact that \( N_{\text{pix}} / N_{\text{map}} \) rises almost continuously from \( W_b^* = 8 \) px all the way in to \( W_b^* = 5 \) px speaks in favor of a true physical peak at \( W_b^* = 5 \) px. The peak at 30 px, which arises at the upper boundary of the \( W_b^* \) range, might also be partly artificial, in the sense that it could correspond to the crest of a larger filament with higher significance \( S \). In any case, the peaks at 9 px, 20 px and 25 px are probably real.

In the bottom panel of Fig. 3, we show three sets of reconstructed filaments, each with a bar width \( W_b \) equal to one of the most prevalent bar widths. We choose \( W_b^* = 5 \) px (dark green), 9 px (middle green), and 25 px (light green). In case of overlap, we let the smaller filaments appear in the foreground. Together, the three sets of filaments provide a good picture of the entire filamentary network displayed in Fig. 2, while they bring out the different morphological trends more clearly.

### Table 3. Summary of the main results obtained for the reconstructed filaments in each of the four Herschel fields.

<table>
<thead>
<tr>
<th>Herschel field</th>
<th>Most prevalent widths (pc)</th>
<th>Preferential relative orientations</th>
</tr>
</thead>
<tbody>
<tr>
<td>Summary figure</td>
<td>( W_b^* ) peak</td>
<td>( 2R_d ) ((p=2.2))</td>
</tr>
<tr>
<td>----------------</td>
<td>--------------------------</td>
<td>-------------------------------</td>
</tr>
<tr>
<td>G210</td>
<td>0.04 ± 0.01</td>
<td>0.07 ± 0.03</td>
</tr>
<tr>
<td>Fig. A1</td>
<td>0.15 ± 0.02</td>
<td>0.23 ± 0.05</td>
</tr>
<tr>
<td></td>
<td>0.18 ± 0.03</td>
<td>0.27 ± 0.05</td>
</tr>
<tr>
<td>G300</td>
<td>0.04 ± 0.01</td>
<td>0.06 ± 0.02</td>
</tr>
<tr>
<td>Fig. A2</td>
<td>0.06 ± 0.01</td>
<td>0.09 ± 0.03</td>
</tr>
<tr>
<td></td>
<td>0.17 ± 0.03</td>
<td>0.25 ± 0.06</td>
</tr>
<tr>
<td>G82</td>
<td>0.18 ± 0.01</td>
<td>0.29 ± 0.06</td>
</tr>
<tr>
<td>Fig. A3</td>
<td>0.32 ± 0.02</td>
<td>0.49 ± 0.07</td>
</tr>
<tr>
<td></td>
<td>0.72 ± 0.04</td>
<td>1.05 ± 0.11</td>
</tr>
<tr>
<td>G202</td>
<td>0.22 ± 0.03</td>
<td>0.35 ± 0.09</td>
</tr>
<tr>
<td>Fig. A4</td>
<td>0.71 ± 0.09</td>
<td>1.05 ± 0.18</td>
</tr>
<tr>
<td></td>
<td>0.97 ± 0.13</td>
<td>1.42 ± 0.23</td>
</tr>
</tbody>
</table>

**Notes.** Column 1: name of the Herschel field and summary figure with the graphical results. Columns 2–3: most prevalent bar widths, \( W_b^* \) and corresponding Plummer widths, \( 2R_d(p=2.2) \), for a Plummer power-law index \( p = 2.2 \) (see Fig. 8 and Table 4 of Paper I). Columns 4–6: preferential relative orientations with respect to the magnetic field, for the three ranges of bar width defined in Sect. 4.2.1: small (S), medium (M), and large (L). The ranges in parsecs are given in the first row, and the dominant trends are summarized in the next rows, with the symbols ||, \( \perp \), and / meaning preferentially parallel, preferentially perpendicular, and without clear preference, respectively.
Fig. 3. Results obtained for the bar widths in the *Herschel* G82 field. *Top:* number of pixels, $N_{\text{pix}}$, whose most significant bar width is $W_\star^b$, normalized to the number of pixels in the map, $N_{\text{map}}$, as a function of $W_\star^b$. The peaks of the histogram correspond to the most prevalent bar widths, $W_{\star \text{peak}}^b$, three of which (highlighted in green) were selected for visualization. *Bottom:* reconstructed filaments with bar width equal to one of the three selected $W_{\star \text{peak}}^b$. Smaller filaments are overlaid on top of larger filaments.

Fig. 4. Results obtained for the bar widths in the four *Herschel* fields. *Top four panels:* histograms of bar width in the four individual *Herschel* fields. *Bottom panel:* combined histograms of the four *Herschel* fields together. In each panel, the solid line gives the normalized number of pixels whose most significant bar width is $W_\star^b$ as a function of $W_\star^b$, whereas the dashed line gives the normalized number of pixels belonging to a reconstructed filament of bar width $W_b$ as a function of $W_b$.

also seen in the other *Herschel* fields. The reconstructed maps of G202 (Fig. A.4) look similar to those of G82 (Fig. A.3), with this small difference that G202 contains no visible striations. In G210 (Fig. A.1), a few striations are again present, but the largest filaments do not show the same degree of internal sub-structuring. All this is even more true in G300 (Fig. A.2), where, in addition to the more numerous striations, crests are visible inside the largest filaments. More globally, G300 shows a higher degree of large-scale ordering, with most filaments being roughly aligned along two orthogonal directions.

The histograms ($N_{\text{pix}}/N_{\text{map}}$) versus $W_\star^b$ of the four individual *Herschel* fields are displayed in Fig. 4 (top four panels, solid line), along with the combined histogram of the four *Herschel* fields together (bottom panel, solid line). These histograms include only the most significant filaments, which means that for every value of $W_\star^b$, they give the normalized number of pixels whose most significant bar width is $W_\star^b$. For completeness, we also plot the histograms ($N_{\text{pix}}/N_{\text{map}}$) versus $W_b$ that include all the reconstructed filaments (dashed line), which means that for every value of $W_b$, they give the normalized number of pixels belonging to a reconstructed filament of bar width $W_b$. All the histograms are plotted along a common $W_\star^b$ (or $W_b$) axis, which encompasses the different $W_\star^b$ ranges of the individual *Herschel* fields ($\approx[0.04, 0.2]$ pc for G210 and G300 and $\approx[0.2, 1.1]$ pc for G82 and G202).

Each of the individual histograms ($N_{\text{pix}}/N_{\text{map}}$) versus $W_\star^b$ (solid line) contains between 4 and 6 peaks. As explained for G82 (Sect. 4.1.1), the first and last peaks, which arise at the lower and upper boundaries, could be partly artificial. Accordingly, the
peaks near 0.04 pc, 0.2 pc, and 1.1 pc in the combined histogram could also be partly artificial. The other peaks in the histograms of G210 and G300 are probably too weak to be truly significant. In contrast, both G82 and G202 exhibit two strong peaks near 0.7 pc and 0.9 pc, which clearly stand out in the combined histogram. No such peaks appear in the histograms \((N_{\text{pix}}/N_{\text{map}})\) versus \(W_0\) (dashed line), where the contributions from the most significant filaments are lost in the contributions from all the reconstructed filaments. Altogether, two preferential bar widths \(\simeq 0.7\,\text{pc}\) and \(\simeq 0.9\,\text{pc}\) emerge from our analysis, though it is not clear whether these are specific to G82 and G202 or whether they could have more generality. Uncovering general trends in the preferential bar widths would require a complete statistical analysis over many more \textit{Herschel} fields.

For each \textit{Herschel} field, we can integrate the curve \((N_{\text{pix}}/N_{\text{map}})\) versus \(W_0^\star\) (solid line) over \(W_0^\star\) to obtain the fraction of the map covered by reconstructed filaments with bar width in the considered range \(\simeq [0.04, 0.2]\,\text{pc}\) for G210 and G300 and \(\simeq [0.2, 1.1]\,\text{pc}\) for G82 and G202. The result is \(\simeq 33\%\) for G210, \(\simeq 25\%\) for G300, \(\simeq 51\%\) for G82, and \(\simeq 55\%\) for G202. Hence, between one quarter and one half of the \textit{Herschel} maps are found to be covered by reconstructed filaments, in agreement with the visual impression left by the reconstructed maps in Appendix A. Let us emphasize, though, that these fractions are probably highly dependent on the angular resolution of the map, in the sense that a given field mapped at higher resolution might be expected to appear more extensively covered by filaments.

It is also interesting to examine how the \(H_2\) column density varies with bar width in the four \textit{Herschel} fields (see the four panels of Fig. 5). In each \textit{Herschel} field, we first consider all the pixels whose most significant bar width is \(W_0^\star\) and calculate their average column density, \(N_{\text{HI}}\), as a function of \(W_0^\star\) (solid curve). We then consider all the pixels belonging to a reconstructed filament of bar width \(W_0\) and calculate their average column density, \(\overline{N_{\text{HI}}}\), as a function of \(W_0\) (dashed curve). When all the reconstructed filaments are included (dashed curve), the average column density clearly increases with increasing bar width, which means that larger filaments tend to have higher column densities. When only the most significant filaments are included (solid curve), this trend becomes much weaker (in G82 and G202) or even disappears entirely (in G210 and G300). What emerges instead are pronounced peaks at intermediate values of \(W_0^\star\). The associated high values of \(N_{\text{HI}}\) suggest that the corresponding intermediate filaments are actually part of larger filaments. Thus, this figure could indicate the existence of special bar widths at which crests and strands inside large filaments with high column densities are more significant than their enclosing filaments.

4.2. Filament relative orientations: variations with bar width

4.2.1. G82

We define three ranges of bar width: the small (S) range includes a single bin at the smallest bar width, \(W_0 = 0.18\,\text{pc}\), the large (L) range includes a single bin at the largest bar width, \(W_0 = 1.08\,\text{pc}\), and the Medium (M) range includes all the intermediate bar widths, \(W_0 = [0.22, 1.04]\,\text{pc}\). In Sect. 3.1, we explained how to reconstruct physical filaments with a given bar width, \(W_0\), and we defined at each pixel a filament orientation angle, \(\psi_f^\star\), for that bar width. We now apply this procedure to the S, M, and L ranges, thereby obtaining S, M, and L filaments, respectively. S and L filaments are simply the reconstructed filaments with the smallest and largest bar widths, respectively. M filaments are formed by the superposition of all the reconstructed filaments of intermediate bar widths, and the associated filament orientation angle at each pixel is the filament orientation angle derived for the most significant intermediate bar width.

The left, middle, and right panels of Fig. 6 show the absolute value of the relative orientation angle, \(|\psi_f^\star - \psi_B|\), inside the S, M, and L filaments, respectively, with the magnetic field orientation plotted with Linear Integral Convolution (LIC) in greyscale in the background. The derived values of \(|\psi_f^\star - \psi_B|\) are in good agreement with the relative orientations inferred from a direct by-eye inspection. The S and M maps exhibit a broad range of relative orientations. In contrast, the L map is dominated by one big filament, which is oriented at \(\approx 90^\circ\) from the magnetic field – except for a short portion beyond the kink in the lower part, which is nearly parallel to the magnetic field.

These general trends appear more clearly and in more detail in the three normalized HROs, \((N_{\text{pix}}/(N_{\text{pix}})_{\text{max}})\) versus \((\psi_f^\star - \psi_B)\),
plotted in polar representation over $[-90^\circ, +90^\circ]$ in Fig. 7. The HRO of S filaments (blue) covers the entire range of relative orientation, with a slight asymmetry that favors small angles. At the other extreme, the HRO of L filaments (purple) is strongly dominated by a pronounced peak in relative orientation near $90^\circ$. Between these two extremes, the HRO of M filaments (orange) exhibits both trends: it covers almost the entire range of relative orientation and has a pronounced peak near $90^\circ$; it also has a weaker and broader peak near $-25^\circ$.

4.2.2. G210

We define again S, M, and L filaments, but now with $W_b = 0.04$ pc, $W_b = [0.05, 0.21]$ pc, and $W_b = 0.22$ pc, respectively. The S, M, and L filaments, together with their $|\psi^*-\psi_B|$ are plotted in the left, middle, and right panels of Fig. 8. S filaments have again a broad range of relative orientations, with a slight asymmetry in favor of small angles, but they form a more structured spatial pattern than in G82 (Fig. 6): a whole group of filaments on the right is nearly parallel to the magnetic field, while a group on the left is nearly perpendicular. L filaments reduce to a pair of filaments threaded along a same straight line, which makes a large angle ($\approx 70^\circ - 85^\circ$) to the magnetic field – except in the uppermost portion of the upper filament, where field lines turn to become nearly parallel to the filament. M filaments for the most part are either nearly parallel ($|\psi^*-\psi_B| \approx 0^\circ - 20^\circ$) or nearly perpendicular ($|\psi^*-\psi_B| \approx 70^\circ - 90^\circ$) to the magnetic field, and they too show more spatial coherence than in G82.

These general trends are nicely confirmed by the HROs plotted in Fig. 9, which in addition provide more quantitative information on the relevant angular ranges.

4.2.3. The four Herschel fields

The S, M, and L maps as well as the four HROs of the four Herschel G82 field, for the three ranges of bar width defined in Sect. 4.2.1: small (left), medium (middle), and large (right). The magnetic field orientation is visualized with LIC in greyscale in the background.

4.2.3. The four Herschel fields

The S, M, and L maps as well as the four HROs of the four Herschel fields are plotted in their respective sheets in Appendix A.

G300 is qualitatively similar to G210. Its S, M, and L ranges are $W_b = 0.04$ pc, $W_b = [0.05, 0.16]$ pc, and $W_b = 0.17$ pc, respectively. The main difference is that the S and M maps of G300 follow a common global pattern. In the M map, the biggest filaments in the middle are nearly perpendicular to the magnetic field, while the smaller filaments on either side are more nearly parallel. This global pattern is reflected in the S map, which contains two overlapping families of filaments with nearly orthogonal orientations. A more minor difference is that the HRO of S filaments in G300 has a stronger asymmetry toward small angles.

G202 also has many similarities with the other three Herschel fields. Its S, M, and L ranges are $W_b = 0.22$ pc, $W_b = [0.27, 1.07]$ pc, and $W_b = 1.11$ pc, respectively. What clearly sets it apart is that its L filaments are at small angles ($|\psi^*-\psi_B| \leq 30^\circ$) to the magnetic field. M filaments automatically follow a similar trend. S filaments, which already show a slight preference for small angles in the other Herschel fields, only have this preference slightly enhanced in G202. Altogether,
the three types of filaments tend to be more parallel than perpendicular to the magnetic field, and this trend gradually increases from S to M to L filaments.

When the four Herschel fields are considered together, a few general trends emerge in each of the S, M, and L ranges, even though these ranges do not refer to the same linear scales for the different fields (see Fig. 5):

- S filaments are numerous and observed at all relative orientations, with a slight general trend toward alignment parallel to the magnetic field. This trend becomes increasingly noticeable along the sequence G82, G210, G300, G202;
- The L map is dominated by one or two filaments covering a restricted range of relative orientations, which is close to 90° in G82, G210, and G300, and broadly around 0° in G202;
- M filaments are observed at most relative orientations, with a strong inclination toward the range covered by L filaments.

An important conclusion of this first analysis is that filaments of different widths align differently with respect to the magnetic field.

### 4.3. Filament relative orientations: variations with column density

#### 4.3.1. G82

We now inquire into a possible correlation between the relative orientation angle, $(\psi_f^\star - \psi_B)$, and the Herschel column density, $N_{H_2}$, for the S, M, and L filaments separately. Considering only pixels that belong to at least one reconstructed filament, we divide the full range of $N_{H_2}$ into 18 bins containing the same number of pixels. For each of the S, M, and L sets of filaments, we construct the HRO in every $N_{H_2}$ bin and we combine the 18 HROs into the 2D histograms $(N_{pix}/N_{map})$ versus $(N_{H_2}, \psi_f^\star - \psi_B)$ plotted in the top row of Fig. 10. Also plotted in Fig. 10 are the 2D histograms of the S filaments that belong (S $\in$ [M, L], middle left panel) or do not belong (S $<$ [M, L], bottom left panel) to larger (M or L) filaments as well as the 2D histograms of the M filaments that belong (M $\in$ L, central panel) or do not belong (M $<$ L, bottom middle panel) to L filaments. The computation of which (S or M) filaments belong to larger filaments is performed on a pixel-by-pixel basis.

In the left column of Fig. 10, the top panel confirms that S filaments exist at all relative orientations, and shows that this is true at all column densities; the slight asymmetry in favor of small angles detected in the blue HRO of Fig. 7 is hardly noticeable. The middle and bottom panels, for their part, tell us that (1) the above conclusions directly apply to S $\in$ [M, L] filaments, which represent the vast majority of S filaments, and (2) the few S $<$ [M, L] filaments are predominantly found at low column densities.

The rightmost panel of Fig. 10 confirms that L filaments have a strong preference for relative orientations $\approx 90°$. It now clearly appears that this preference applies only at high column densities, where all the L filaments are found. The latter actually reduce to one big filament, as already mentioned in connection with the right panel of Fig. 6.

In the middle column of Fig. 10, the top panel confirms that M filaments exist at all relative orientations, and shows that this is true at all column densities; the slight asymmetry in favor of small angles detected in the blue HRO of Fig. 7 is hardly noticeable. The middle and bottom panels, for their part, tell us that (1) the above conclusions directly apply to S $\in$ [M, L] filaments, which represent the vast majority of S filaments, and (2) the few S $\in$ [M, L] filaments are predominantly found at low column densities.

The rightmost panel of Fig. 10 confirms that L filaments have a strong preference for relative orientations $\approx 90°$. It now clearly appears that this preference applies only at high column densities, where all the L filaments are found. The latter actually reduce to one big filament, as already mentioned in connection with the right panel of Fig. 6.

In the middle column of Fig. 10, the top panel confirms that M filaments exist at nearly all relative orientations, with a marked preference for $\approx 90°$. We now see that this preference applies mostly at high column densities. Furthermore, the middle and bottom panels bring to light a clear dichotomy between M $\in$ L filaments, which are almost exclusively found at high column densities, with relative orientations $\approx 90°$, and M $<$ L filaments,
Fig. 10. 2D histograms of the relative orientation angle between the reconstructed filaments and the magnetic field, $(\psi_f^* - \psi_B)$, as a function of H$_2$ column density, $N_{\text{H}_2}$, in the Herschel G82 field, for the three ranges of bar width defined in Sect. 4.2.1: small (left, blue), medium (middle, orange), and large (right, purple). Top: all of the S, M, and L filaments. Middle: S and M filaments that belong to larger filaments. Bottom: S and M filaments that do not belong to larger filaments.

Fig. 11. Same as Fig. 10, but for the Herschel G210 field.

which are mostly found at low-to-intermediate column densities, with a slight clustering around $-25^\circ$. The M$_{\text{L}}$ histogram is strikingly similar to the L histogram, which can be explained by the fact that the L filament in the right panel of Fig. 6 has an almost identical M counterpart in the middle panel.

4.3.2. G210

Following the same procedure as for G82 (Sect. 4.3.1), we obtain the 2D histograms ($N_{\text{pix}}/N_{\text{map}}$) versus $(N_{\text{H}_2}, \psi_f^* - \psi_B)$ displayed in Fig. 11.

The S filaments (left column) show nearly the same trends as in G82. Two minor differences are that (1) the slight asymmetry in favor of small relative orientation angles detected in the blue HRO of Fig. 9 is now visible at low column densities, especially in the S$_{\text{[M,L]}}$ histogram, and (2) the few S$_{\text{[M,L]}}$ filaments are almost exclusively found at low column densities and at specific relative orientations ($\approx$10$^\circ$, 50$^\circ$, and 90$^\circ$).

For L filaments (right column), the broad peak at $\approx$70$^\circ$–85$^\circ$ observed in the purple HRO of Fig. 9 is now seen to arise from a wide range of intermediate-to-high column densities. In contrast, the few pixels with relative orientations between $\approx$ –35$^\circ$ and 70$^\circ$...
are all concentrated at the highest column densities. Remember that these pixels belong to the uppermost portion of the upper filament in the right panel of Fig. 8, where field lines undergo significant bending.

For M filaments (middle column), the two preferential relative orientations $\approx 0^\circ$–$20^\circ$ and $\approx 70^\circ$–$90^\circ$ emerging from the orange HRO of Fig. 9 are now seen to apply mostly to low-to-intermediate and intermediate-to-high column densities, respectively. Roughly speaking, the former pertain to the $M_L$ filaments and the latter to the $M_{HL}$ filaments. The strong similarity between the $M_L$ and L histograms can again be explained by the morphological resemblance between L filaments and their enclosed M filaments (see right and middle panels of Fig. 8).

We now focus on the 2D histogram of M filaments and use the NMF and PCA methods introduced in Sect. 3.2 to determine the column density at which the transition in $|\psi^* - \psi|$, from $\approx 0^\circ$–$20^\circ$ to $\approx 70^\circ$–$90^\circ$, occurs. The result is displayed in Fig. 12. The 2D histogram ($N_{pix}/N_{map}$) versus $(N_{HL}, |\psi^* - \psi|)$ directly obtained from the top middle panel of Fig. 11) is plotted in the top panel. The residual of the reconstruction with the first two principal components derived with NMF (2D histogram minus the sum of the two principal components) is overplotted with contour lines. The first two principal components together have an estimated reconstruction capability of 65%. Their profiles as functions of $|\psi^* - \psi|$ are shown in the right panel of Fig. 12, where the blue and green components mostly represent filaments that are approximately parallel and perpendicular, respectively, to the magnetic field. The relative weights of the two components as functions of column density are shown in the bottom panel. The “parallel component” dominates at $N_{HL} < 1.1 \times 10^{21}$ cm$^{-2}$, while the “perpendicular component” dominates at $N_{HL} > 1.1 \times 10^{22}$ cm$^{-2}$ – except in the bin $[1.85, 2.95] \times 10^{21}$ cm$^{-2}$, which is dominated by the “parallel component”. This nearly parallel orientation at high column density comes from the uppermost portion of the M filament associated with the upper L filament in the right panel of Fig. 8. What can be retained from the application of the NMF and PCA methods is that the relative orientation of filaments undergoes a transition from mostly parallel to mostly perpendicular to the magnetic field at a column density $N_{HL} \approx 1.1 \times 10^{21}$ cm$^{-2}$.

4.3.3. The four Herschel fields

The 2D histograms ($N_{pix}/N_{map}$) versus $(N_{HL}, |\psi^* - \psi|)$ of the S, M, and L filaments in the four Herschel fields are displayed in their respective sheets in Appendix A. For compactness, we do not show the 2D histograms of the $S_{G10}$ and $S_{G210}$ filaments or those of the $M_L$ and $M_{HL}$ filaments separately, as we did for G82 in Fig. 10 and G210 in Fig. 11. However, we did construct and examine the separate 2D histograms for the four Herschel fields, and we found that they all share the same general properties.

G300 has many similarities with G210 (perhaps in part because both are nearby fields), as well as its own peculiarities. S filaments with low column densities show a clear asymmetry toward small relative orientation angles, as well as its own peculiarities. S filaments with low column densities show a clear asymmetry toward small relative orientation angles, and so do $S_{G10}$ and $S_{G210}$ filaments. L filaments reduce to a single intermediate-to-high $N_{HL}$ filament, with relative orientation $\approx 90^\circ$. M filaments can be divided into two classes: low-to-intermediate $N_{HL}$ filaments, with relative orientations between $-45^\circ$ and $+45^\circ$, and intermediate-to-high $N_{HL}$ filaments, with relative orientations $\approx 90^\circ$. These two classes can be roughly identified with the $M_L$ and $M_{HL}$ filaments, respectively. They also correspond to the two principal components derived with the NMF method applied to the M filaments, with an estimated reconstruction capability of 71% (see Fig. 13). The bottom panel of Fig. 13 indicates that the transition between the two components occurs at a column density $N_{HL} \approx 1.4 \times 10^{21}$ cm$^{-2}$, which turns out to be very close to the transition column density $\approx 1.1 \times 10^{21}$ cm$^{-2}$ obtained for G210.

G202 shows similar trends to the other three Herschel fields, but only for S, $S_{G10}$, $S_{G210}$, and $M_{HL}$ filaments. L and $M_L$ filaments, which are still found at high column densities, now cover a broad range of relative orientations around $0^\circ$.

To sum up, the general trends emerging from the 2D histograms of S, M, and L filaments in the four Herschel fields are the following:

- S filaments are found at all column densities and all relative orientations. Those with low column densities have a general tendency (hardly noticeable in G82, weak in G210 and G202, stronger in G300) to be more parallel than perpendicular to
the magnetic field. The small fraction of S filaments present outside larger (M or L) filaments mostly have low column densities, with no preference for parallel alignment; – L filaments have high or intermediate-to-high column densities. In G82, G210, and G300, they tend to be nearly perpendicular to the magnetic field, while in G202, they tend to be more nearly parallel; – M filaments span the entire ranges of column densities and relative orientations, with however a much more structured distribution than S filaments. Most filaments behave very similarly to L filaments.

If we now consider S, M, and L filaments together, we can derive the complete 2D histograms \(\frac{N_{\text{pix}}}{N_{\text{map}}}\) versus \((N_{\text{H}_2}, \psi^*_f - \psi_B)\) for each of the four individual Herschel fields (top four panels of Fig. 14), as well as the combined 2D histogram for the four Herschel fields together (bottom panel). The full range of \(N_{\text{H}_2}\), which encompasses the different \(N_{\text{H}_2}\) ranges of the individual Herschel fields, is divided into 36 bins with the same sum (over the four fields) of \(N_{\text{pix}}/N_{\text{map}}\). Clearly, for each Herschel field, the complete histogram (Fig. 14) does not closely resemble any of the S, M, or L histogram (Appendix A). This results from a combination of two different factors: (1) Each pixel retained in the complete histogram is assigned the orientation angle of the most significant filament passing through it, which can now be a S, M, or L filament. (2) The \(N_{\text{H}_2}\) grid based on the four Herschel fields is very different from the individual \(N_{\text{H}_2}\) grids.

The combined histogram (bottom panel of Fig. 14), which is simply the superposition of the four individual complete histograms, tentatively shows a bimodal distribution in relative orientation, with a tendency for low/high-\(N_{\text{H}_2}\) filaments to be more nearly parallel/perpendicular to the magnetic field. The transition from parallel to perpendicular orientation occurs over a range of column densities \(\approx [2, 3] \times 10^{21} \text{ cm}^{-2}\). However, this range is biased upward by G202, which contains mostly parallel filaments throughout its \(N_{\text{H}_2}\) range. If G202 were excluded from the combined histogram, the transition would occur at \(N_{\text{H}_2} \approx (1 - 2) \times 10^{21} \text{ cm}^{-2}\), consistent with the values obtained earlier for G210 (\(\approx 1.1 \times 10^{21} \text{ cm}^{-2}\)) and G300 (\(\approx 1.4 \times 10^{21} \text{ cm}^{-2}\)). Here,
too, a more accurate estimate of the transition column density would require a complete statistical analysis over a large number of Herschel fields.

5. Discussion

5.1. Filament bar widths

By applying FilDReaMS to the $H_2$ column density maps of the four Herschel fields of our sample, we were able to construct histograms of the most significant bar size, $W^\text{peak}$ (solid line in the top four panels of Fig. 4), and from the peaks of the histograms we were able to derive the most prevalent bar widths, $W^\text{peak}$. The histogram of each Herschel field is found to contain between 4 and 6 peaks, with the outermost peaks (located at the boundaries) being probably partly artificial. When the four fields are considered together (bottom panel of Fig. 4), no significant peak emerges in the range $\approx [0.04, 0.2]$ pc covered by the nearby fields, G210 and G300, whereas two pronounced peaks near 0.7 pc and 0.9 pc stand out in the range $\approx [0.2, 1.1]$ pc covered by the distant fields, G82 and G202. Our current small-number statistics do not allow us to ascribe any generality to these preferential bar widths.

The most prevalent bar widths, $W^\text{peak}$, derived in our study can causally be converted to Plummer widths, $2R_{\text{flat}}$, with the help of Fig. 8 and Table 4 of Paper I. We consider the case of default noise level (i.e. the typical noise level of Herschel maps) and a Plummer power-law index of $p = 2.2$ (median value obtained by Arzoumanian et al. 2019, for a sample of 599 filaments including G300). However, one has to be aware that the resulting $2R_{\text{flat}}$ are just rough estimates obtained very indirectly through a method that is not designed to derive the transverse column density profiles of filaments. Therefore, it is probably not very meaningful to compare our $2R_{\text{flat}}$ to the Plummer widths derived in previous studies, for instance, based on true Plummer-type fits to the transverse column density profiles of filaments all along their lengths (Juvela et al. 2012; Kainulainen et al. 2016; Cox et al. 2016; Arzoumanian et al. 2011, 2019).

In the case of G210 and G300, the most prevalent bar widths could potentially be found in the range $W^\text{peak} \approx [0.04, 0.2]$ pc, corresponding to Plummer widths in the range $2R_{\text{flat}} \approx [0.06, 0.3]$ pc. In the case of G82 and G202, the relevant ranges are $W^\text{peak} \approx [0.2, 1.1]$ pc, corresponding to $2R_{\text{flat}} \approx [0.3, 1.6]$ pc. The two most prevalent bar widths in G82 and G202, $W^\text{peak} \approx 0.7$ pc and 0.9 pc, translate into Plummer widths $2R_{\text{flat}} \approx 1.0$ pc and 1.3 pc, respectively. No other preferential Plummer width is uncovered in our sample.

We do not recover the characteristic width $\sim 0.1$ pc found in previous studies (Arzoumanian et al. 2011, 2019; Palmeirim et al. 2013; Cox et al. 2016). The existence of this characteristic width was called into question by a re-analysis of the Arzoumanian et al. (2011, 2019) methodology (Panopoulou et al. 2017, 2022) as well as by the conclusions of an anisotropic wavelet analysis applied to the Herschel column-density maps of the nearby Aquila Rift and Polaris Flare (Ossenkopf-Okada & Stepanov 2019). Here we do not have enough statistics to settle the debate. The angular resolution (36") of our column density maps precludes detecting filaments widths $\sim 0.1$ pc in molecular clouds more distant than $\sim 200$ pc (Kainulainen et al. 2016), which means that a filament width $\sim 0.1$ pc is undetectable in our two distant fields (G82 and G202). For the two nearby fields (G210 and G300), where a filament width $\sim 0.1$ pc falls within the detectability range, its non-detection could suggest that either these two fields are not particularly representative, or the level of noise is too high for the relation between bar width and Plummer width to be really meaningful, or the $\sim 0.1$ pc is not a universal characteristic filament width.

FilDReaMS is not designed to give information on the transverse column density profile of filaments. However, the appearance of small filaments superposed on larger filaments (see right panel of Fig. 2, bottom panel of Fig. 3, and similar figures in Appendix A) gives some rough, qualitative idea of the internal structure of large filaments. In particular, the fact that a fraction of the small (S) filaments form the crests of larger filaments may provide some indirect support for the often-used Plummer-type profile.

The ramifications and striations observed in Figs. 2, 3, and Appendix A are similar to those observed in dust-continuum and molecular-line maps (see, for instance, Sugitani et al. 2011; Arzoumanian et al. 2013; Peretto et al. 2013; Saaasto et al. 2017; Palmeirim et al. 2013; André et al. 2014). They suggest a hierarchical process in filament formation, with small filaments potentially feeding larger filaments. Such a scenario is consistent with the predictions of numerical simulations (see Balsara et al. 2001; Gómez & Vázquez-Semadeni 2014).

5.2. Filament orientations to the magnetic field

5.2.1. Comparison with previous studies

The results obtained for the relative orientations between filaments and the local magnetic field in our sample of Herschel fields confirm some of the general trends observed in previous studies (as introduced in Sect. 1), but they also display a variety of behaviors. Globally we find that filaments at high column densities are mostly perpendicular to the magnetic field, while a subset of low column-density filaments (mainly striations) are mostly parallel. In two of our fields (G210 and G300), we find a transition between parallel and perpendicular alignments at $H_2$ column densities ($\approx 1.1 \times 10^{14}$ cm$^{-2}$ and $1.4 \times 10^{23}$ cm$^{-2}$, respectively), consistent with the results from previous studies (Planck Collaboration XXVI 2016a; Planck Collaboration XXXV 2016b; Malinen et al. 2016; Alina et al. 2019). However, these general results are not systematically observed in our sample: low-column-density filaments can cover all relative orientations, and in G202, all filaments, including those at high column densities, tend to be more parallel than perpendicular to the magnetic field.

The main novelty of our methodology is that we now have control over the bar widths of filaments and we can investigate the filament relative orientations in different ranges of bar width, down to the angular resolution of the maps. The HROs of our Herschel fields show different behaviors for S, M, and L filaments, respectively. In each field, a large number of S filaments are detected at all relative orientations, although with a slight trend toward parallel alignment. The largest filaments are mostly perpendicular and are also associated with the highest column densities. Many S and M filaments are parts of larger filaments, with S filaments forming crests, internal sub-structures (such as strands), or ramifications. Our method makes it possible to separate them from isolated filaments and to study their relative orientations separately.

In Sect. 4, we presented the results obtained when applying our methodology to $N_{HI}$ maps of our sample. This enabled us to identify truly material filamentary structures. However, the angular resolution of the column density maps is limited to 36" (the resolution of the 500 μm Herschel band). In contrast, the Herschel intensity maps at 250μm offer a better angular
resolution (18”), which therefore gives access to thinner filaments. For comparison, we performed the same analysis based on the intensity maps of our sample. The results are displayed in Figs. B1–B4. The gain in angular resolution clearly leads to an increase in the number of detected filaments in the S range. Indeed, improving the angular resolution allows us to detect filaments that are either thinner or at lower column densities. These filaments are expected to be parallel to the magnetic field, and this would explain the observed trend in the 2D histograms. Some of these filaments could potentially be striations, or fibers in the sense defined by Hacar et al. (2013, 2018). Such a result was already noticed in the study by Clark et al. (2014) when using HI data at higher resolution. In the nearby fields G210 and G300, many more striations and strands are detected; the HROs of S filaments are now dominated by nearly parallel orientations, and their 2D histograms show that this trend prevails mainly at low column densities. In G82, there is an increased number of parallel S filaments at all column densities. In G202, the trends observed in both the HROs and the 2D histograms from the parallel S filaments are scale-dependent, and so is their relative importance. Chen et al. (2016) showed with three-dimensional (3D) MHD simulations that turbulence dominates at large (cloud) scales, magnetic fields at intermediate (filament) scales, and self-gravity at small (core) scales. This conclusion also emerges from a simple toy model of a homogeneous cloud of size ℓ: if interactions between the three processes are ignored, turbulent energy density varies as ℓ (Larson’s law), magnetic energy density as ℓ₀, and self-gravitational energy density as ℓ⁻¹ (C-Y. Chen’s talk at the Sofia on-line workshop ‘Magnetic Fields and the structure of the filamentary Interstellar Medium’, June 2021). This ordering has direct implications for the formation of filaments and for their orientations with respect to the ambient magnetic field.

At early stages (large scales, low densities), turbulence dominates and leads to the formation of low-density structures through compression (Padoan et al. 2001) and shear (Hennebelle 2013) of the gas and the frozen-in field lines; these structures tend to be elongated parallel to the magnetic field. Field-aligned striations can also form through the non-linear coupling of MHD waves in the presence of density inhomogeneities (Tritsis & Tassis 2016). At late stages (small scales, high densities), self-gravity becomes dominant and causes gas to contract preferentially along field lines; the resulting high-density structures tend to be elongated perpendicular to the magnetic field (e.g., Chen et al. 2016; Seifried et al. 2020; Girichidis 2021). Soler & Hennebelle (2017) showed that a system in ideal-MHD turbulence naturally tends to evolve towards a configuration where iso-density contours (and hence density structures) are either parallel or perpendicular to the magnetic field. They also found that, in the presence of a relatively strong magnetic field, compressive motions, resulting from either gravitational collapse or converging flows, can produce a transition from mostly parallel at low column densities to mostly perpendicular at higher column densities.

One can think of several reasons why our observational results do not consistently conform to the theoretical predictions and why different clouds exhibit different behaviors. There are at least two strict physical reasons. (1) The transition from preferentially parallel to preferentially perpendicular filaments occurs at a critical density where self-gravity takes over the dominant role from magnetic fields. This critical density is strongly dependent on the initial physical conditions (density, ionization fraction, magnetic field, turbulence...) in the parent molecular cloud, which in turn vary significantly from cloud to cloud. (2) The magnetic field provides a natural reference direction to measure the orientations of filaments, but other reference directions

The Herschel G202 field displays a different behaviour, with an ordered magnetic field structure at large scale over the whole field. Our study shows that filaments at all scales (S, M, L), including the densest and largest filaments, tend to be parallel to the magnetic field. This conclusion is in agreement with the results of Alina et al. (2020). Combining Planck observations of dust polarized emission and CO molecular line tracers, they studied the large-scale magnetic field and its interchange with the gas dynamics in this region. Their analysis reveals a shock region with colliding filaments, and they suggest that the magnetic field remains dragged during the evolution of the cloud, resulting in an orientation parallel to the filaments.

5.2.2. Physical interpretation

Physically, filaments arise from an interplay between turbulence, magnetic fields, and self-gravity (see, for instance, André et al. 2014; Li et al. 2014; Myers 2017). These three processes are scale-dependent, and so is their relative importance. These three processes are scale-dependent, and so is their relative importance. These three processes are scale-dependent, and so is their relative importance. These three processes are scale-dependent, and so is their relative importance.
can arise from factors such as the 3D shape of the parent molecular cloud, the large-scale stratification of the Galactic disk, shock waves driven by nearby supernovae, shearing motions in the ambient medium... These additional factors will generally tilt filaments away from orientations strictly parallel or perpendicular to the magnetic field.

We can also mention two geometrical reasons. (1) The 3D orientation of filaments relative to the magnetic field is generally modified upon projection on the sky. Filaments that are parallel to the magnetic field in 3D remain parallel in projection, but filaments that are perpendicular to the magnetic field in 3D generally do not appear perpendicular in the sky (e.g., Planck Collaboration XXXII 2016a). This projection effect leads to a bias in the histogram of relative orientation. (2) The filaments observed in 2D intensity or column density maps as well as the magnetic field orientation inferred from 2D polarization maps result from an integration along the line of sight, which generally involves a superposition of several features. Therefore, relative orientations measured in 2D sky maps are not necessarily representative of the actual situation in 3D space. The difference between relative orientations in 2D and 3D was recently quantified by Girichidis (2021) in high-resolution MHD simulations of the turbulent ISM. Here, the issue of line-of-sight integration may be a real problem for the two distant Herschel fields, G82 and G202, which lie close to the midplane: in these fields, there is no guarantee that the magnetic field orientation from Planck truly represents the local magnetic field of the clouds, especially for low $N_{\text{H}_2}$ filaments.

6. Summary and conclusion

In Paper I, we presented a new method, called Fi1DReaMS, which makes it possible to detect filaments of a given bar width in an image, to identify the most prevalent bar widths (and corresponding Plummer widths) in the image, and to derive the local orientation angles of the detected filaments.

Here, we applied Fi1DReaMS to a small sample of four Herschel fields located at different distances (two fields at 150 pc and two fields beyond 600 pc), embedded in different Galactic environments, and captured at different evolutionary stages of star formation. Our fields cover broad dynamic ranges of spatial scales and $N_{\text{H}_2}$ column densities (see Table 1), thereby making it possible to detect all kinds of filamentary structures from large and massive filaments down to striations.

In a second step, we compared the filament orientation angles derived with Fi1DReaMS to the local magnetic field orientation inferred from Planck polarization data. This enabled us to study (for the first time) the statistics of the relative orientation angle as functions of both spatial scale and $N_{\text{H}_2}$ column density. We emphasize that combining spatial scales and $H$ stages of star formation. Our fields cover broad dynamic ranges of ambient physical conditions and evolutionary stage. In this manner, our results will hopefully contribute to a better physical understanding of the process of filament and star formation.

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References

Regarding the filament relative orientations to the magnetic field, a few general trends emerge from our study:

- Many small (S) filaments are detected in each Herschel field, over broad ranges of column densities and relative orientations. At low column densities, S filaments tend to be slightly more parallel than perpendicular to the magnetic field. This trend, which increases along the sequence G82, G210, G300, G202, is more pronounced in intensity maps (which have a twice better angular resolution and can, therefore, reveal smaller filaments) than in column density maps.

- Only one or two large (L) filaments appear in each Herschel field, always at high or intermediate column densities. Large filaments clearly show a preferential orientation to the magnetic field, which is nearly perpendicular in G82, G210, and G300, but more nearly parallel in G202. These trends are very similar in column density and intensity maps.

- The two nearby fields G210 and G300 undergo a transition in relative orientation, from mostly parallel to mostly perpendicular to the magnetic field, at a column density $N_{\text{H}_2} \approx 1.1 \times 10^{21}$ cm$^{-2}$ for G210 and $N_{\text{H}_2} \approx 1.4 \times 10^{21}$ cm$^{-2}$ for G300. This transition appears more clearly in intensity maps than in column density maps. No such transition is seen in the distant fields G82 and G202.

In the future, we plan to extend our sample to the 116 fields of the Herschel-GCC program. This will enable us to probe different places in the Galaxy, with a wide range of physical parameters and across the entire sequence of star formation. This, in turn, will open the way to a complete statistical analysis of the relative orientation between filaments and the local magnetic field, not only as a function of spatial scale, but also as a function of ambient physical conditions and evolutionary stage. In this manner, our results will hopefully contribute to a better physical understanding of the process of filament and star formation.
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Appendix A: Summary of the results obtained from column-density maps

Fig. A.1. Summary of the main graphical results obtained when applying Fi1DReaM5 to the H$_2$ column density map of the Herschel G210 field. The top, second, third, bottom-left, and bottom-right panels correspond to Figs. 2, 3, 6, 7, and the top row of Fig. 10, respectively.
Fig. A.2. Same as Fig. A.1, but for the G300 field.
Fig. A.3. Same as Fig. A.1, but for the G82 field.
Fig. A.4. Same as Fig. A.1, but for the G202 field.
Appendix B: Summary of the results obtained from intensity maps

Fig. B.1. Same as Fig. A.1, starting from the Herschel 250 μm intensity map of the G210 field.
Fig. B.2. Same as Fig. B.1, but for the G300 field.
Fig. B.3. Same as Fig. B.1, but for the G82 field.
Fig. B.4. Same as Fig. B.1, but for the G202 field.