

LETTER TO THE EDITOR

C_{60}^+ diffuse interstellar band correlations and environmental variations

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ABSTRACT

Context. The diffuse interstellar bands (DIBs) are absorption features seen in the spectra of astronomical objects that arise in the interstellar medium. Today, more than 500 DIBs have been observed, mostly in the optical and near-infrared wavelengths. The origin of the DIBs is unclear; only ionised buckminsterfullerene, C_{60}^+ , has been identified as a viable candidate for two strong and three weaker DIBs.

Aims. We investigate the correlations between the strengths of the two strongest C_{60}^+ DIBs as well as their environmental behaviour.

Methods. We analysed measurements of the strengths of the two C_{60}^+ DIBs at 9577 and 9633 Å for 26 lines of sight. We used two different methods, including Monte Carlo simulations, to study their correlations and the influence of measurement errors on the correlation coefficients. We examined how the strength of the C_{60}^+ DIBs changes as a result of different environmental conditions, as measured by the concentration of H/H₂ and the strength of the ambient UV radiation.

Results. In contrast to results recently reported by Galazutdinov et al. (2021, AJ, 161, 127), we find a high correlation between the strengths of the C_{60}^+ DIBs. We also discovered that the behaviour of the correlated C_{60}^+ bands is quite distinct from other DIBs at 5780, 5797, and 6203 Å in different environments.

Key words. ISM: lines and bands – ISM: molecules – ISM: clouds – dust, extinction

1. Introduction

In the last century, more than 500 diffuse interstellar bands (DIBs) have been found arising from the interstellar medium (ISM; Fan et al. 2019). The first two DIBs were found by Mary Lea Heger in 1919, who observed two broad features around 5797 and 5780 Å that originate from the ISM (Heger 1922; see also Herbig 1995 for a review). The carrier molecules that give rise to these spectral features are unknown for most of the DIBs, with one exception: C_{60}^+ .

Foing & Ehrenfreund (1994) first discovered two relatively strong and broad interstellar features near 9577 and 9633 Å. They proposed that these features were carried by ionised buckminsterfullerene, C_{60}^+ , based on two strong absorption bands recorded for C_{60}^+ embedded in rare gas matrices (d’Hendecourt et al. in Joblin 1992; Fulara et al. 1993). Recently, five near-infrared DIBs were attributed to C_{60}^+ by Campbell et al. (2015), who used ion trap technology to record laboratory spectra of C_{60}^+ -He complexes and C_{60}^+ embedded in He droplets. These measurements confirmed the proposition that the 9633 and 9577 DIBs are due to C_{60}^+ . Additionally, they reported three weaker bands at 9428, 9365, and 9348 Å that showed up in the laboratory data.

Today, the identification of C_{60}^+ is supported by an impressive array of observational and experimental studies (see Linnartz et al. 2020, for a review). However, there has been some debate in the literature about this assignment due, at least partly, to several factors that complicate the comparison between the laboratory data and astronomical observations. First, all C_{60}^+ features lie in spectral regions that are heavily contaminated by telluric water vapour lines, making their detection and proper characterisation difficult. This is especially hard for the weaker bands, but even the two stronger DIBs at 9633 and 9577 Å suffer from telluric contamination, which results in uncertainties on the measurements (Foing & Ehrenfreund 1994, 1997; Galazutdinov et al. 2000, 2021). Cordiner et al. (2017, 2019), however, used the *Hubble* Space Telescope Imaging Spectrograph (HST/STIS) to observe several targets that exhibit the 9577 and 9632 DIBs. The use of space telescopes eliminates the need for error-prone telluric correction methods, at the cost of some spectral resolution. This allowed these authors to confirm the presence of the weaker C_{60}^+ DIBs in several targets; however, their observations did not include the wavelength range covering the 9632 Å DIB. The second complicating factor is that for some spectral types the 9633 DIB overlaps with a stellar Mg II line, which can affect equivalent width (EW) measurements. When

studying the C_{60}^+ DIBs, it is thus necessary to study objects for which contamination is not an issue (either because the stellar line is not present or because the stellar radial velocity moves it well away from the DIB) or, alternatively, provide a good estimate for the strength of the stellar line, which is not a straightforward task.

A key issue that has been used to challenge the C_{60}^+ identification is based on correlation studies between the C_{60}^+ DIBs. Correlation studies have a long history in DIB research, dating back to Snow (1973). Researchers have used correlation studies to investigate the nature of the DIB carriers by comparing the DIB EWs to different interstellar quantities, such as reddening E_{B-V} and other extinction parameters (e.g. Snow 1973; Megier et al. 2005; Friedman et al. 2011; Raimond et al. 2012; Puspitarini et al. 2013), the column densities of interstellar atoms and small molecules (e.g. Krelowski et al. 1992; Galazutdinov et al. 2004; Friedman et al. 2011), and the strengths of other DIBs (e.g. Snow 1973; Cami et al. 1997; Moutou et al. 1999; McCall et al. 2010; Xiang et al. 2012; Fan et al. 2017; Bondar 2020; Smith et al. 2021), often with the aim to group DIBs in families that could have similar (or even the same) carriers (Krelowski & Walker 1987; Westerlund & Krelowski 1989; Cami et al. 1997; Wszolek & Godłowski 2003).

At the base of many of these correlation studies is the idea that if two DIBs arise from the same carrier, their EWs should in principle exhibit a perfect correlation since they both arise from the electronic ground state. Based on this assumption, Galazutdinov et al. (2021) state that the two strong bands at 9577 Å and 9632 Å should show a perfect correlation (to within uncertainties). According to their measurements and their analysis, however, the two bands in fact show a remarkably poor correlation, even accounting for measurement uncertainties, which led them to the conclusion that the two DIBs cannot be due to the same carrier and hence cannot be due to C_{60}^+ . This is a surprising result in view of the good spectral matches with the laboratory experiments and begs a careful and critical examination.

In this Letter we compose a dataset from different literature sources to study the correlations of the two C_{60}^+ bands with each other as well as with several environmental parameters.

2. Data

The data we use in this Letter include EW measurements of the 9577 and 9632 Å DIBs and environmental parameters for 26 lines of sight that have been collected from different sources in the literature. We summarise them in Table 1. The environmental data, in particular the colour excess E_{B-V} , the column densities of neutral hydrogen and molecular hydrogen, $N(H)$ and $N(H_2)$, and the molecular fraction of hydrogen, f_{H_2} , were taken from Fan et al. (2017), who studied the behaviour of eight relatively strong DIBs in different interstellar environments. In the Fan et al. (2017) dataset, most of the spectra were taken with the ARC echelle spectrograph (ARCES) on the 3.5 m telescope at the Apache Point Observatory (APO). Fan et al. (2017), however, did not include the 9577 and 9633 Å DIBs. Therefore, most of the EWs for the 9633 and 9577 DIBs were taken from Galazutdinov et al. (2021), who collected spectra of these bands using several high resolving power echelle spectrographs. Furthermore, for the line of sight toward HD 43384, the EWs were taken from Jenniskens et al. (1997), who used the 1.4 m ESO Coudé Auxiliary Telescope (CAT) and the Coudé Echelle Spectrometer (CES); for BD +63 1964, we used the values from Cordiner et al. (2017) obtained with HST/STIS.

For HD 183143, the EWs for the C_{60}^+ bands were examined in different studies, with a rather large difference between the reported values. Foing & Ehrenfreund (1997) report 266 mÅ and 311 mÅ for the EWs of the 9577 and 9633 DIBs, respectively, while Cox et al. (2014) found 260 ± 5 mÅ and 263 ± 3 mÅ and Galazutdinov et al. (2021) list 105 ± 20 mÅ and 300 ± 20 mÅ. The disagreement in the values for the 9633 DIB could be (partly) due to the different methods used for the Mg II line correction. For the 9577 Å DIB, any discrepancies are most likely due to telluric correction methods. The standard technique for the removal of telluric water vapour lines is the division of the reddened star spectrum by one from an un-reddened star of a similar spectral type observed through the same air mass. For the purposes of this Letter, we chose to use the EWs determined by Cox et al. (2014) for HD 183143.

3. Methods

For our correlation studies, we used the Pearson linear correlation coefficient, r . As one can see in Eq. (1),

$$r = \frac{\sum_{i=1}^n (x_i - \bar{x})(y_i - \bar{y})}{\sqrt{\sum_{i=1}^n (x_i - \bar{x})^2} \sqrt{\sum_{i=1}^n (y_i - \bar{y})^2}}, \quad (1)$$

this coefficient is described by the ratio between the covariance of two variables, x_i and y_i , and the product of their standard deviation. The result of this equation always gives a value between -1 and 1 , where $r = 1$ (or $r = -1$) indicates a strong positive (negative) relationship between the two variables and $r = 0$ indicates no relationship at all.

Pair-wise comparisons between different quantities can help in identifying general relationships. To include the measurement error in the correlation calculations, we also calculated the weighted correlation coefficient $r_w(x, y; w)$ using the weighted covariance $\text{Cov}(x, y; w)$ at a given weight, w :

$$\text{Cov}(x, y; w) = \frac{\sum_i w_i \left(x_i - \frac{\sum_i w_i x_i}{\sum_i w_i}\right) \left(y_i - \frac{\sum_i w_i y_i}{\sum_i w_i}\right)}{\sum_i w_i}. \quad (2)$$

The weighted correlation coefficient $r_w(x, y; w)$ can then be calculated with Eq. (3):

$$r_w(x, y; w) = \frac{\text{Cov}(x, y; w)}{\sqrt{\text{Cov}(x, x; w)\text{Cov}(y, y; w)}}. \quad (3)$$

For the purpose of this study, the measurement error σ was taken as a weight in the form of $1/\sigma^2$. We also adopted an independent approach to assess the influence of measurement errors by performing Monte Carlo simulations in which we first constructed a few thousand new datasets. Each of these datasets was constructed from the original measurements, but we added a random value to each of the measurements drawn from a normal distribution with a standard deviation equal to the measurement error. These thousand datasets thus represent a thousand realisations of our original measurements. For each dataset, we calculated the Pearson correlation coefficient, r . We then determined the mean and the standard deviation of the correlation coefficients corresponding to each of the thousands of datasets.

Table 1. Data of the 9632 and 9577 Å DIBs and environmental data for 26 lines of sight.

Name	SpT	E_{B-V} [mag]	$\log_{10}N(\text{H})$	$\log_{10}N(\text{H}_2)$	f_{H_2}	W(9633) [mÅ]	W(9577) [mÅ]
HD 22951	B0.5V	0.27 ± 0.03 ^(a)	21.04 ± 0.13 ^(a)	20.46 ± 0.14 ^(a)	0.34 ± 0.14 ^(a)	72 ± 22 ^(b)	
HD 23180	B1III	0.31 ± 0.03 ^(a)	20.82 ± 0.09 ^(a)	20.60 ± 0.12 ^(a)	0.55 ± 0.12 ^(a)	141 ± 46 ^(b)	77 ± 34 ^(b)
HD 27778	B3V	0.37 ± 0.03 ^(a)	20.95 ± 0.15 ^(a)	20.79 ± 0.06 ^(a)	0.58 ± 0.12 ^(a)	64 ± 17 ^(b)	50 ± 30 ^(b)
HD 36861	O8e	0.15 ± 0.03 ^(a)	20.81 ± 0.12 ^(a)	19.12 ± 0.14 ^(a)	0.04 ± 0.02 ^(a)	71 ± 22 ^(b)	70 ± 40 ^(b)
HD 37022	O6	0.34 ± 0.03 ^(a)	21.54 ± 0.11 ^(a)	17.55 ^(a)	0.00021 ^(a)	72 ± 15 ^(b)	71 ± 20 ^(b)
HD 40111	B0.5II	0.20 ± 0.03 ^(a)	21.03 ± 0.09 ^(a)	19.73 ± 0.14 ^(a)	0.09 ± 0.04 ^(a)	76 ± 23 ^(b)	67 ± 35 ^(b)
HD 43384	B3Ib	0.58 ± 0.03 ^(a)	21.27 ± 0.30 ^(a)	20.87 ± 0.14 ^(a)	0.44 ± 0.25 ^(a)	40 ± 18 ^(d)	41 ± 18 ^(d)
HD 54662	O7III	0.35 ± 0.05 ^(a)	21.23 ± 0.10 ^(a)	20.00 ± 0.14 ^(a)	0.11 ± 0.05 ^(a)	60 ± 20 ^(b)	60 ± 20 ^(b)
HD 57061	O9III	0.16 ± 0.03 ^(a)	20.80 ± 0.08 ^(a)	15.48 ± 0.14 ^(a)	0 ± 0 ^(a)	65 ± 17 ^(b)	94 ± 47 ^(b)
HD 143275	B0.3IV	0.17 ± 0.03 ^(a)	21.01 ± 0.08 ^(a)	19.41 ± 0.14 ^(a)	0.05 ± 0.02 ^(a)	94 ± 26 ^(b)	
HD 144470	B1V	0.22 ± 0.03 ^(a)	21.18 ± 0.09 ^(a)	20.05 ± 0.08 ^(a)	0.13 ± 0.04 ^(a)	103 ± 21 ^(b)	58 ± 25 ^(b)
HD 145502	B3V	0.24 ± 0.03 ^(a)	21.20 ± 0.12 ^(a)	19.89 ± 0.12 ^(a)	0.09 ± 0.04 ^(a)	158 ± 20 ^(b)	120 ± 30 ^(b)
HD 147165	B2III+O9V	0.41 ± 0.04 ^(a)	21.38 ± 0.08 ^(a)	19.79 ± 0.12 ^(a)	0.05 ± 0.02 ^(a)	230 ± 33 ^(b)	190 ± 36 ^(b)
HD 147889	B2V	1.07 ± 0.04 ^(a)	21.80 ± 0.05 ^(a)		0.48 ± 0.19 ^(a)		26 ± 18 ^(b)
HD 148605	B2V	0.13 ± 0.03 ^(a)	20.66 ± 0.08 ^(a)	18.74 ± 0.14 ^(a)	0.02 ± 0.01 ^(a)	119 ± 35 ^(b)	80 ± 20 ^(b)
HD 149757	O9.5Ve	0.32 ± 0.03 ^(a)	20.69 ± 0.10 ^(a)	20.64 ± 0.06 ^(a)	0.64 ± 0.08 ^(a)	96 ± 30 ^(b)	72 ± 50 ^(b)
HD 167971	O8e	1.08 ± 0.06 ^(a)	21.60 ± 0.30 ^(a)	20.85 ± 0.12 ^(a)	0.26 ± 0.19 ^(a)	178 ± 13 ^(b)	190 ± 20 ^(b)
HD 168607	B9Ia	1.57 ± 0.04 ^(a)					335 ± 120 ^(b)
HD 169454	B1.5Ia	1.12 ± 0.04 ^(a)	21.51 ± 0.10 ^(a)	20.98 ± 0.13 ^(a)	0.37 ± 0.12 ^(a)	130 ± 20 ^(b)	82 ± 10 ^(b)
HD 170740	B2V	0.48 ± 0.03 ^(a)	21.04 ± 0.15 ^(a)	20.86 ± 0.08 ^(a)	0.57 ± 0.13 ^(a)	150 ± 20 ^(b)	93 ± 20 ^(b)
HD 183143	B7Iae	1.27 ± 0.06 ^(a)				263 ± 5 ^(c)	260 ± 5 ^(c)
HD 190603	B1.5IaE	0.72 ± 0.06 ^(a)	21.41 ± 0.10 ^(a)	20.54 ± 0.13 ^(a)	0.21 ± 0.09 ^(a)	155 ± 111 ^(b)	150 ± 20 ^(b)
Cyg OB2 12	B5Ie	3.31 ± 0.10 ^(a)					390 ± 70 ^(b)
HD 204827	B0V	1.11 ± 0.04 ^(a)	21.35 ± 0.05 ^(a)		0.65 ± 0.10 ^(a)		10 ± 10 ^(b)
HD 208501	B8Ib	0.75 ± 0.04 ^(a)					97 ± 54 ^(b)
BD +63 1964	B0II	1.00 ± 0.03 ^(a)	21.70 ± 0.10 ^(a)	21.10 ± 0.18 ^(a)	0.33 ± 0.14 ^(a)		74 ± 3 ^(e)

References. ^(a)Fan et al. (2017), ^(b)Galazutdinov et al. (2021), ^(c)Cox et al. (2014), ^(d)Jenniskens et al. (1997), ^(e)Cordiner et al. (2017).

4. Results and discussion

4.1. Correlation between the 9577 and 9633 Å DIBs

We first investigate the correlation between the two strong C₆₀⁺ DIBs. It should be noted that in our sample of stars we have measurements for the 9577 DIB for 24 objects and for the 9633 DIB for 20 objects.

As one can see in Fig. 1, we find a relatively high correlation coefficient of $r = 0.90$ for these two DIBs. For this dataset, the weighted correlation coefficient, $r_w = 0.93$, is even slightly larger than the regular Pearson correlation coefficient, r . Using Monte Carlo simulations, where r was calculated 4000 times, the calculated correlation coefficient is slightly smaller but still fairly high, at $r_{\text{MCS}} = 0.82 \pm 0.06$.

This result is inconsistent with Galazutdinov et al. (2021), who report a poor mutual correlation of only $r = 0.37$ between the EWs of the strong C₆₀⁺ DIBs. However, it should be noted that in their study they used observations of 62 lines of sight obtained using several high resolving power echelle spectrographs, with measurements of the EWs for both C₆₀⁺ DIBs for 43 lines of sight. They state that they calculated the correlation coefficients with the measurement error σ taken as a weight in the form of $1/\sigma^2$. Because of the different results of the correlation calculations of this dataset and that of Galazutdinov et al. (2021), we recalculated the correlation coefficient with their published data (the black data points in Fig. 1) using different methods. As one can see in Fig. 1, we find instead a correlation coefficient of $r = 0.67$, a weighted correlation coefficient $r_w = 0.53$ (where we used $1/\sigma^2$ as the weight), and a Monte Carlo result of $r_{\text{MCS}} = 0.60 \pm 0.05$. While these values are smaller than ours, none of them come close to their reported value of $r = 0.37$.

It is important in this context to realise that if two variables are intrinsically highly correlated, measurement uncertainties

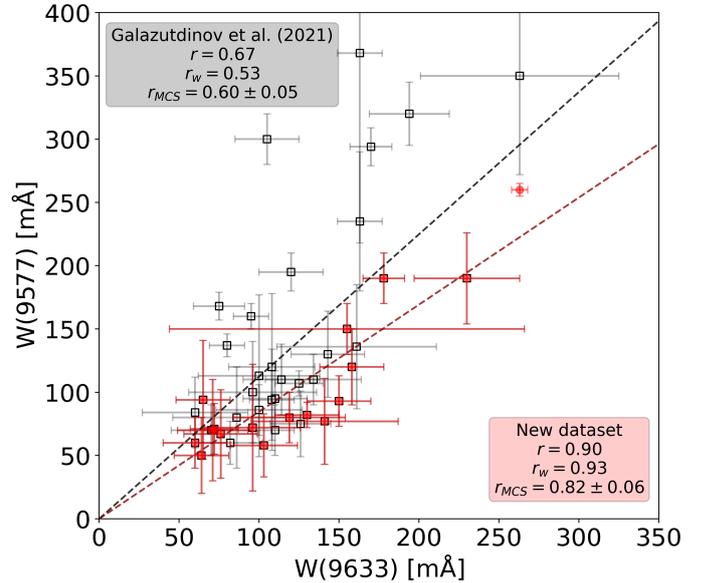


Fig. 1. Equivalent width measurements of the 9577 DIB versus those of the 9633 DIB and the various resulting correlation coefficients. The red data points represent the measurements from Table 1. The open black squares show the EWs of 43 lines of sight from Galazutdinov et al. (2021), with measurements for both C₆₀⁺ DIBs. We note the relatively high correlation coefficients for our dataset.

will always reduce the correlation between them. Thus, in the presence of noise, even auto-correlations can yield values significantly lower than unity. To get a sense of the magnitude of this effect on our calculations, we calculated the ‘self-correlation’ of both the 9633 and 9577 DIBs for our dataset as well as that

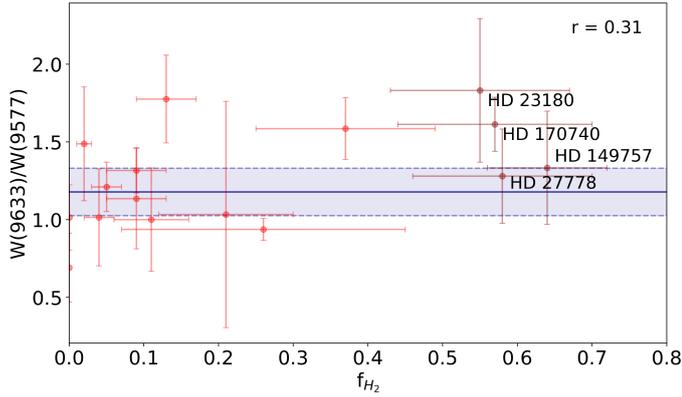


Fig. 2. $W(9633)/W(9577)$ ratio as a function of f_{H_2} .

from Galazutdinov et al. (2021). Not surprisingly, the correlation coefficient, r , and the weighted correlation coefficient, r_w , have a value of 1 for both datasets. However, when performing such a correlation with a Monte Carlo simulation, the achieved correlation is 0.87–0.93 – of the same magnitude as the values for the mutual correlations. This shows that even if two DIBs were perfectly correlated intrinsically, they would yield an effective correlation of about 0.9 due to measurement errors. Cami et al. (1997) found that this systematic ‘decorrelation bias’ of 1.8σ in the auto-correlation has to be considered to check the reliability of the calculated correlation. Given the large uncertainties on the EW measurements, the correlation we find between the two C_{60}^+ DIBs is thus an indication of an even tighter intrinsic correlation, and possibly even a perfect correlation.

It is important in this context to point out that the idea that absorption bands originating from the same electronic ground state should show a perfect correlation has indeed long been a key assumption that has driven many correlation studies. However, this assumption does not always hold, and C_{60}^+ is one such exception. Indeed, Lykhin et al. (2019) carried out detailed computations to elucidate the electronic structure of C_{60}^+ relevant to the DIBs. They conclude that the bands at 9577 and 9632 Å arise from excitation in the non-Franck-Condon region of the ground electronic state. They further suggest that ro-vibrational coupling can affect the Franck-Condon factors and therefore the intensities of the bands. If true, this would then imply that the ratio of the two DIB strengths could be used as a measure of the interstellar temperature. In any event, it illustrates that the band ratio does not need to be constant, even for transitions from the same species and the same electronic ground state.

4.2. Behaviour of C_{60}^+ DIBs with H and H_2 column densities

The column densities of H and H_2 were measured in a number of our targets, as reported by Fan et al. (2017). For a homogeneous line of sight, f_H and f_{H_2} represent the average atomic and molecular fraction. However, interstellar lines of sight are heterogeneous, including the conditions that represent the outside, inside, and edge of interstellar clouds (and sometimes of multiple clouds). Ionised buckminsterfullerene is mostly related to conditions in the outer edge of clouds, where there are sufficient photons with energy ≥ 7 eV to ionise C_{60} , in balance with the local electron density, n_e . We note that in H II regions with an abundance of photons ≥ 13.6 eV, C_{60}^+ would be further ionised to C_{60}^{++} (the ionisation potential being 11.35 eV) but nonetheless would not break its cage structure.

The molecular hydrogen fraction, f_{H_2} , can be interpreted as an indicator of the typical local hydrogen density in the main interstellar clouds in each line of sight (Fan et al. 2017). As H_2 dissociates under the influence of UV radiation, f_{H_2} can also be used as an approximation of the strength of the local UV field (Cami et al. 1997). In Fig. 2, the ratio $W(9633)/W(9577)$ is indicated versus this average f_{H_2} indicator, showing values in the range 1.0–1.5 at $f_{H_2} < 0.3$ and 1.3–1.8 for $f_{H_2} > 0.35$. At first approximation, there is no significant variation in the ratio across our sightlines, taking into account the noise in data and mixing of heterogeneous interstellar conditions.

In Fig. 3a the normalised EWs of the 9633 and 9577 DIBs are plotted against the fraction of neutral hydrogen, f_H . Using Eq. (4) the neutral hydrogen fraction, f_H , can be calculated from the column densities of neutral hydrogen and molecular hydrogen, $N(H)$ and $N(H_2)$:

$$f_H = \frac{N(H)}{N(H) + 2N(H_2)}. \quad (4)$$

The normalised EWs of the C_{60}^+ DIBs experience a ‘u-shaped’ behaviour, seen in Fig. 3a. While $W(\text{DIB})/E(B-V)$ increases strongly with f_H at high f_H , the lines of sight HD 23180, HD 170740, HD 149757, and HD 27778 are relatively strong at around $f_H = 0.4$. This would not be expected for ionised molecules such as C_{60}^+ . The reason could be the presence of multiple velocity components in the line of sight, indicating the presence of clouds with different compositions. Therefore, the molecular fraction of H_2 could arise from a different region where C_{60}^+ is not that abundant. For example, HD 170740 has two evident components in the profiles of He I and C II (Galazutdinov et al. 2017). Furthermore, a stellar Mg II line could influence the accuracy of some of the measured EWs of the 9633 DIB. Walker et al. (2017) found close spectral standards to be superior to corrections based on line profiles from model atmosphere calculations introduced by Galazutdinov et al. (2017) in correcting for contamination by this stellar Mg II line.

We expect C_{60}^+ to be present in the outer edges of clouds, where more UV-processed material can be found. In Table 2 the normalised EWs of the DIBs at 6203 Å, 5780 Å, and 5797 Å from Fan et al. (2017) are shown for the 26 investigated lines of sight. These DIBs have been subject to many past studies, and the measurement errors on their EWs are very small. In Fig. 3b the normalised EWs of these DIBs are plotted against f_H once again. In contrast to the C_{60}^+ DIBs, here the DIB strengths increase with f_H until a maximum is reached, after which the DIB strengths decrease with f_H . This rise and fall of the normalised DIB strength in Fig. 3b is quite different from the physical behaviour of the correlated C_{60}^+ bands in Fig. 3a.

4.3. Correlations with E_{B-V} and hydrogen fraction

The colour excess E_{B-V} is calculated from the difference between the observed colour index of an object $(B-V)_{\text{observed}}$ and the intrinsic colour index $(B-V)_{\text{intrinsic}}$ predicted from its spectral type. Thus, the optical depth of interstellar dust can be related to E_{B-V} . In Fig. 3c the relation of the two normalised C_{60}^+ DIBs to E_{B-V} is shown. The dataset of Table 1 contains two outliers, Cyg OB2 12 and HD 168607, where E_{B-V} is larger than 1.5. To study the behaviour and ionisation properties of the DIBs at low E_{B-V} , we focused on the lines of sight with $E_{B-V} < 1.6$. We note that the ionisation potential (IP) of C_{60}^+ is 11.35 eV (Baral et al. 2016). Sonnentrucker et al. (1997) developed a simple photoionisation model for the DIB carriers based on the DIB variations at

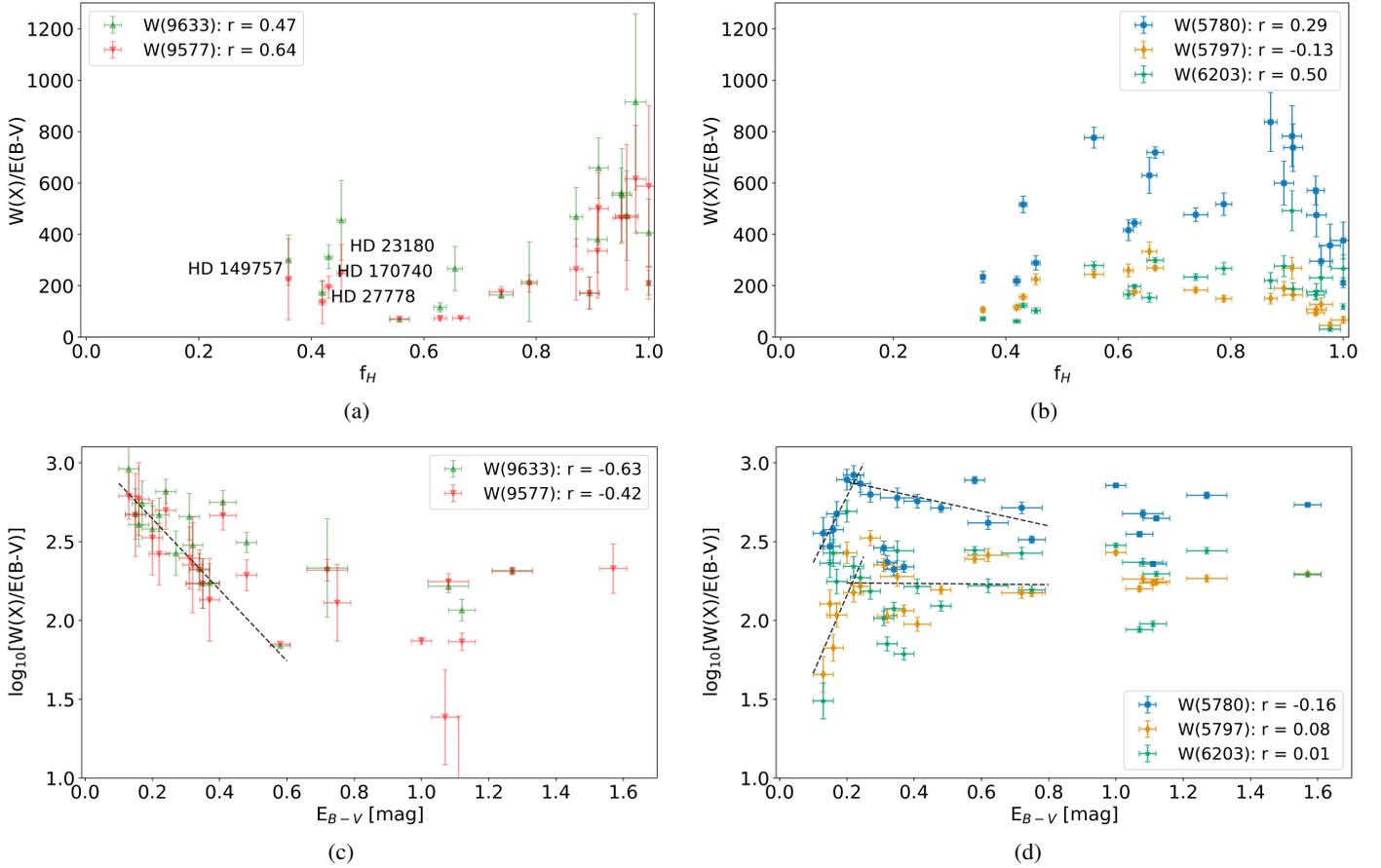


Fig. 3. Correlation studies for the C_{60}^+ bands. The normalised EWs of the two C_{60}^+ DIBs and the DIBs at 6203, 5780, and 5797 Å are plotted against the fraction of neutral hydrogen f_H (top) and the colour excess E_{B-V} (bottom). We note the very specific behaviour of C_{60}^+ compared to other DIB carriers. (a) Normalised EWs of the C_{60}^+ DIBs as functions of f_H . (b) Normalised EWs of the other DIBs as functions of f_H . (c) Normalised EWs of the C_{60}^+ DIBs as functions of E_{B-V} . (d) Normalised EWs of the other DIBs as functions of E_{B-V} .

low reddening. Their model enables us to estimate the DIB IPs from the slope of the variation in the normalised DIB strength as a function of reddening. The prediction for C_{60}^+ would be a slope of 3.1 for $\log(W/E_{B-V})$ versus E_{B-V} . For ionised carriers, such as C_{60}^+ , they are predicted to become neutral in higher density regions. Thus, the normalised EWs of the corresponding DIBs would decrease with E_{B-V} . This coincides with our results in Fig. 3c, where the negative slope for E_{B-V} values between 0.1 and 0.6 indicates that more C_{60}^+ is transformed into C_{60} . The regime where C_{60}^+ might be ionised to the doubly charged fullerene cation C_{60}^{++} cannot be seen in Fig. 3c, and yet this might happen at $E_{B-V} < 0.1$ for single clouds. This result suggests that ionised buckminsterfullerenes reside mostly outside the clouds.

In Fig. 3d the normalised EWs of the 5780, 5797, and 6203 DIBs are plotted against E_{B-V} . Here a rise and fall can be seen for low values of E_{B-V} . This lambda-shaped behaviour was also reported by Fan et al. (2017) for the 5780 and 5797 DIBs. We note that the negative slope of -1 for $\log(W_{5780}/E)$ for $E_{B-V} = 0.3$ to 0.7 can be explained by a simple ionisation model where the DIB carrier column density is carried in the edge of a cloud (Sonnentrucker et al. 1997). However, this slope might be affected if other factors such as hydrogenation (vs. H/UV) or electron density balance are affected within the clouds (Vuong & Foing 2000). For C_{60}^+ the negative slope of -2 for $\log(W/E_{B-V})$ versus E_{B-V} for $0.2 < E_{B-V} < 0.4$ is sharper than the -1 slope expected for a pure electron recombination effect

(and as observed for W(5780)). This could indicate that C_{60}^+ is reduced in the inner edge of clouds due not only to the formation of C_{60} but also possibly to other compounds, for instance via hydrogenation.

5. Conclusions

We have analysed the C_{60}^+ DIBs at 9633 and 9577 Å, performing correlation studies with different methods. We find a high correlation of $r = 0.90$ between these two DIBs for a dataset of 26 lines of sight, in contrast to results by Galazutdinov et al. (2021) who found a poor mutual correlation of the C_{60}^+ DIBs and questioned the common origin of these DIBs. We have not been able to reproduce their results even when using their listed measurement values. A recent study by Nie et al. (2021) also supports a close 9577 and 9633 Å correlation.

Correlation studies of the C_{60}^+ DIBs with different environmental parameters show that the ratio of the two C_{60}^+ DIBs does not change significantly within errors for a very large range of conditions. A distribution of C_{60}^+ band ratios of 1–1.4 and 1.3–1.8 for an extreme range of conditions may be investigated with more accurate measurements and better modelling of intrinsic environment conditions. This could lead to a better understanding of possible changes in the band ratio and ratios that are slightly different from those measured in the laboratory.

Table 2. Equivalent widths of the DIBs at 6203, 5780, and 5797 Å for the 26 lines of sight from Fan et al. (2017).

Name	W(6203) [mÅ]	W(5780) [mÅ]	W(5797) [mÅ]	Name	W(6203) [mÅ]	W(5780) [mÅ]	W(5797) [mÅ]
HD 22951	41.42 ± 1.68	169.9 ± 1.9	89.9 ± 1.1	HD 147889	93.39 ± 1.70	377.4 ± 3.3	169.9 ± 0.7
HD 23180	31.97 ± 1.57	89.5 ± 1.4	69.7 ± 0.4	HD 148605	4.00 ± 0.50	46.4 ± 1.0	5.9 ± 0.7
HD 27778	22.66 ± 0.86	81.0 ± 1.1	42.7 ± 0.7	HD 149757	22.73 ± 0.92	74.8 ± 1.4	34.0 ± 0.7
HD 36861	34.60 ± 2.10	44.3 ± 1.6	19.1 ± 0.6	HD 167971	252.48 ± 2.86	514.3 ± 5.9	197.7 ± 1.3
HD 37022	40.37 ± 1.23	71.7 ± 1.4		HD 168607	306.31 ± 5.94	850.3 ± 9.0	310.1 ± 1.6
HD 40111	98.40 ± 4.80	156.4 ± 2.8	53.7 ± 2.5	HD 169454	220.95 ± 2.64	497.7 ± 2.0	196.7 ± 1.6
HD 43384	161.78 ± 2.36	450.2 ± 3.3	141.8 ± 1.1	HD 170740	59.22 ± 2.29	247.8 ± 1.5	75.0 ± 0.2
HD 54662	96.73 ± 3.31	209.7 ± 2.1	66.5 ± 0.5	HD 183143	351.05 ± 1.31	789.5 ± 8.4	234.0 ± 0.6
HD 57061	42.80 ± 2.50	60.2 ± 2.0	10.7 ± 0.5	HD 190603	192.31 ± 3.53	372.7 ± 2.7	107.7 ± 0.9
HD 143275	30.00 ± 1.00	80.7 ± 0.8	18.4 ± 0.5	Cyg OB2 12	399.81 ± 4.74	843.0 ± 8.5	355.9 ± 3.1
HD 144470	48.44 ± 2.08	184.2 ± 1.8	33.0 ± 1.0	HD 204827	105.66 ± 1.68	253.2 ± 1.6	191.7 ± 0.5
HD 145502	44.77 ± 1.85	177.0 ± 2.1	39.5 ± 1.0	HD 208501	116.94 ± 1.54	244.1 ± 2.3	112.4 ± 1.3
HD 147165	67.28 ± 2.82	234.0 ± 3.0	38.8 ± 1.3	BD +63 1964	299.47 ± 4.59	718.4 ± 4.6	269.4 ± 0.9

In lines of sight dominated by atomic H, we find a correlation between the normalised EWs of the C_{60}^+ DIBs and f_H . We find also a negative correlation of relative band strength at E_{B-V} from 0.15–0.5, indicating the decrease in C_{60}^+ entering the edge of clouds due to electron and H recombination. We find that C_{60}^+ survives outside and in the outer edges of clouds, as opposed to many other DIB carriers that are depleted for single clouds thinner than $E_{B-V} = 0.2$. Especially in the near-infrared wavelength range of the C_{60}^+ bands, it can be challenging to correct for telluric and stellar spectral lines in the lines of sight (Cox et al. 2017). We look forward to extending the analysis with ESO Diffuse Interstellar Bands Large Exploration Survey (EDIBLES) spectra (Lallement et al. 2018) in future studies.

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References

Baral, K. K., Aryal, N. B., Esteves-Macaluso, D. A., et al. 2016, *Phys. Rev. A*, **93**, 033401
 Bondar, A. 2020, *MNRAS*, **496**, 2231
 Cami, J., Sonnentrucker, P., Ehrenfreund, P., & Foing, B. H. 1997, *A&A*, **326**, 822
 Campbell, E. K., Holz, M., Gerlich, D., & Maier, J. P. 2015, *Nature*, **523**, 322
 Cordiner, M. A., Cox, N. L. J., Lallement, R., et al. 2017, *ApJ*, **843**, L2
 Cordiner, M. A., Linnartz, H., Cox, N. L. J., et al. 2019, *ApJ*, **875**, L28
 Cox, N. L. J., Cami, J., Kaper, L., et al. 2014, *A&A*, **569**, A117
 Cox, N. L., Cami, J., Farhang, A., et al. 2017, *A&A*, **606**, A76
 Fan, H., Welty, D. E., York, D. G., et al. 2017, *ApJ*, **850**, 194
 Fan, H., Hobbs, L. M., Dahlstrom, J. A., et al. 2019, *ApJ*, **878**, 151
 Foing, B. H., & Ehrenfreund, P. 1994, *Nature*, **369**, 296

Foing, B. H., & Ehrenfreund, P. 1997, *A&A*, **317**, L59
 Friedman, S. D., York, D. G., McCall, B. J., et al. 2011, *ApJ*, **727**, 33
 Fulara, J., Jakobi, M., & Maier, J. P. 1993, *Chem. Phys. Lett.*, **211**, 227
 Galazutdinov, G. A., Krelowski, J., Musaeu, F. A., Ehrenfreund, P., & Foing, B. H. 2000, *MNRAS*, **317**, 750
 Galazutdinov, G. A., Manicò, G., Pirronello, V., & Krelowski, J. 2004, *MNRAS*, **355**, 169
 Galazutdinov, G. A., Shimansky, V. V., Bondar, A., Valyavin, G., & Krelowski, J. 2017, *MNRAS*, **465**, 3956
 Galazutdinov, G. A., Valyavin, G., Ikhsanov, N. R., & Krelowski, J. 2021, *AJ*, **161**, 127
 Heger, M. L. 1922, *Lick Obs. Bull.*, **10**, 146
 Herbig, G. H. 1995, *ARA&A*, **33**, 19
 Jenniskens, P., Mulas, G., Porceddu, I., & Benvenuti, P. 1997, *A&A*, **327**, 337
 Joblin, C. 1992, PhD Thesis, Université Paris VII, France
 Krelowski, J., & Walker, G. A. H. 1987, *ApJ*, **312**, 860
 Krelowski, J., Snow, T. P., Seab, C. G., & Papaj, J. 1992, *MNRAS*, **258**, 693
 Lallement, R., Cox, N. L. J., Cami, J., et al. 2018, *A&A*, **614**, A28
 Linnartz, H., Cami, J., Cordiner, M., et al. 2020, *J. Mol. Spectrosc.*, **367**, 111243
 Lykhin, A., Ahmadvand, S., & Varganov, S. 2019, *J. Phys. Chem. Lett.*, **10**, 115
 McCall, B. J., Drosback, M. M., Thorburn, J. A., et al. 2010, *ApJ*, **708**, 1628
 Megier, A., Krelowski, J., & Weselak, T. 2005, *MNRAS*, **358**, 563
 Moutou, C., Krelowski, J., D’Hendecourt, L., & Jamrozczak, J. 1999, *A&A*, **351**, 680
 Nie, T. P., Xiang, F. Y., & Li, A. 2021, *MNRAS*, submitted [arXiv:2111.05769]
 Puspitarini, L., Lallement, R., & Chen, H. C. 2013, *A&A*, **555**, A25
 Raimond, S., Lallement, R., Vergely, J. L., Babusiaux, C., & Eyer, L. 2012, *A&A*, **544**, A136
 Smith, F. M., Harriott, T. A., Majaess, D., Massa, L., & Matta, C. F. 2021, *MNRAS*, **507**, 5236
 Snow, T. P., Jr. 1973, *AJ*, **78**, 913
 Sonnentrucker, P., Cami, J., Ehrenfreund, P., & Foing, B. H. 1997, *A&A*, **327**, 1215
 Vuong, M. H., & Foing, B. H. 2000, *A&A*, **363**, L5
 Walker, G. A. H., Campbell, E. K., Maier, J. P., & Bohlender, D. 2017, *ApJ*, **843**, 56
 Westerlund, B. E., & Krelowski, J. 1989, *A&A*, **218**, 216
 Wszolek, B., & Godłowski, W. 2003, *MNRAS*, **338**, 990
 Xiang, F., Liu, Z., & Yang, X. 2012, *PASJ*, **64**, 31