

Toward a realistic macroscopic parametrization of space plasmas with regularized κ -distributions

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ABSTRACT

So-called κ -distributions are widely invoked in the analysis of nonequilibrium plasmas from space, although a general macroscopic parametrization as known for Maxwellian plasmas near thermal equilibrium is prevented by the diverging moments of order $l \geq 2\kappa - 1$. To overcome this critical limitation, recently novel regularized κ -distributions (*RK*) have been introduced, including various anisotropic models with well-defined moments for any value of $\kappa > 0$. In this paper, we present an evaluation of the pressure and heat flux of electron populations, as provided by moments of isotropic and anisotropic *RK*s for conditions typically encountered in the solar wind. We obtained finite values even for low values of $\kappa < 3/2$, for which the pressure and heat flux moments of standard κ -distributions are not defined. These results were also contrasted with the macroscopic parameters obtained for Maxwellian populations, which show a significant underestimation especially if an important suprathermal population is present (e.g., for $\kappa < 2$), but ignored. Despite the collisionless nature of solar wind plasma, a realistic characterization as a fluid becomes thus possible, retaining all nonthermal features of plasma particles.

Key words. solar wind – plasmas – methods: statistical

1. Introduction

Collision-poor plasmas from astrophysical and experimental setups cannot easily reach thermal equilibrium. Plasma particles exhibit non-Maxwellian velocity distributions that are well represented by the κ -power laws (Vasyliunas 1968; Collier et al. 1996; Maksimovic et al. 2005; Štverák et al. 2008), including isotropic and anisotropic κ -distribution functions which greatly facilitate descriptions of plasma effects at small scales; see Pierrard & Lazar (2010) for a review. These departures from thermal equilibrium are expected as a consequence of long-range (collective) interactions of plasma particles mediated by electromagnetic field fluctuations (Treumann et al. 2004; Leubner & Vörös 2005). Scenarios based on the existence of κ -distributions are widely invoked, especially in the expansion of solar atmosphere, to explain coronal plasma heating (Scudder 1992), differential acceleration of the solar wind (Maksimovic et al. 1997), and its interaction with planetary environments (Meyer-Vernet 2001). These scenarios may also offer indirect insights from the physics of nonthermal emissions, as observed from flaring sources in the solar corona (Dudík et al. 2017) or from more distant astrophysical objects (Nicholls et al. 2012).

Similar to a standard Maxwellian for plasmas at (or near) thermal equilibrium, the main moments of the velocity distribution are expected to provide a macroscopic characterization of nonequilibrium plasmas, but moments of a κ -distribution are however not well defined for any value of κ (Scherer et al. 2017, 2019a). Thus, the moment of order l can be defined only for $\kappa > (l+1)/2$, and the pressure (or kinetic temperature) parameter

is notorious, being directly derived from the moment of order $l = 2$, which is convergent only for $\kappa > 3/2$. If the temperature is taken as a fitting parameter, the observations not satisfying this condition are simply discarded and lost from the analysis (Štverák et al. 2008), leading to questionable or even unrealistic interpretations. In order to overcome these limitations, standard κ -distributions have recently been adjusted by introducing the regularized κ -distribution (*RK*) with well-defined moments for any value of $\kappa > 0$ (Scherer et al. 2017). General expressions of the principal moments ($l \leq 3$) were also provided for various anisotropic *RK*s, able to describe temperature anisotropies or drifting populations, see Scherer et al. (2019a).

These recent results are applied in the present paper to evaluate the main properties of electron populations, for example, pressure P (or temperature T) given by the second-order moment and the heat flux derived from the third-order moment. We chose conditions typically observed in the solar wind, where plasma particles are predominantly nonrelativistic ($k_B T \ll mc^2$) and the main heat-flux is carried by the beaming populations, i.e., by the electron strahl. The solar wind is not highly magnetized, for instance, at 1 AU the electron plasma frequency is much higher than cyclotron frequency, but the interplanetary magnetic field ensures a gyrotropy of the velocity distributions, i.e., the distributions are almost isotropic in the plane perpendicular (subscript \perp) to the magnetic field (Maksimovic et al. 2005; Štverák et al. 2008). Thus, deviations from isotropy in general take the form of electron beams (or strahls) propagating parallel (subscript \parallel) to the magnetic field and anisotropic pressure P (or temperature T) of electron populations with different components $P_\perp \neq P_\parallel$ (or $T_\perp \neq T_\parallel$).

In Sect. 2 we consider non-drifting electrons associated with central populations in the observed distributions (Štverák et al. 2008), i.e., the core and halo electrons, which are together well reproduced by a κ -distribution function (Lazar et al. 2017) without a major drift with respect to the massive ions. In the solar wind (ions) referential the net heat flux of these electron populations is zero, and we may restrict the analysis to the pressure (or temperature) components. Additional drifting or beaming electrons responsible for the electron heat flux are enhanced in fast winds or during more energetic events like coronal mass ejections (CMEs). These electron beams or strahls are nonthermal and are well described by the same κ -distribution functions with field-aligned drifts (Nieves-Chinchilla & Viñas 2008; Štverák et al. 2009; Viñas et al. 2010; Lazar et al. 2017). In Sect. 3 we evaluate the pressure components and the (parallel) heat flux transported by the electron beams with isotropic or anisotropic pressures. Macroscopic properties predicted for RK s are compared with the corresponding moments of standard κ -distributions. These properties are also contrasted with Maxwellian limits reproducing the core of the observed distributions, which are often invoked to motivate a fluid description and ignore suprathermal populations. In the last section we summarize the present results and discuss their importance from a perspective of a realistic macroscopic characterization of space plasmas as fluids.

2. Non-drifting distributions

We assume a plasma of electrons (subscript e) and protons (subscript p) specific to the solar wind conditions at 1 AU. The following basic set of plasma parameters characterizes the low-energy core or Maxwellian populations: number densities $n_e = n_p = n_0 = 10 \text{ cm}^{-3}$; temperatures $T_e \simeq T_p = 10^5 \text{ K}$, implying a mean energy per electron $k_B T_e \simeq 8.6 \text{ eV}$ and for thermal velocity of electrons $\theta = (2k_B T_e / m_e)^{1/2} \simeq 0.005c$, where c is the speed of light in vacuum; and magnetic field $B_0 = 10^{-4} \text{ G} = 10 \text{ nT}$, implying for the ratio of plasma and cyclotron frequencies $\omega_{pe} / \omega_{ce} \simeq 100$. These values are used in our evaluations for each type of κ -distribution function, standard or regularized, intended to account for suprathermal electrons and their contribution to the macroscopic properties of solar wind plasma. Heavier ions may also be present in the solar wind but usually with a very low number density $n_i / n_{\text{total}} < 4\%$, and these are therefore ignored in the present analysis.

2.1. Isotropic temperatures

For isotropic temperatures $T_{\parallel} = T_{\perp} = T$ the observed velocity distributions are described by a standard κ -distribution function, denoted with subscript K (Appendix A summarizes the abbreviations used for the distribution functions), i.e.,

$$f_K(v) = \frac{n_0 \Gamma[\kappa]}{(\pi \theta^2 \kappa)^{3/2} \Gamma[\kappa - 1/2]} \left(1 + \frac{v^2}{\kappa \theta^2}\right)^{-\kappa-1}, \quad (1)$$

where θ is given by the second order moment $k_B T = (2/3)(m < v^2 > / 2) = (m/3) \int d^3 v v^2 f_K = (m \theta^2 / 2) \kappa / (\kappa - 3/2)$. The same θ may be associated with the Maxwellian (subscript M) limit (Lazar et al. 2015, 2016)

$$f_M(v) = \frac{n_0}{(\pi \theta^2)^{3/2}} \exp\left(-\frac{v^2}{\theta^2}\right), \quad (2)$$

describing the core (low-energy) population in the observed distribution (assuming $n_{\text{core}} \simeq n_0$). Thermal velocity of the

core electrons is directly related to their temperature by $\theta = (2k_B T_M / m)^{1/2}$.

The regularized κ -distribution (subscript RK ; Scherer et al. 2017)

$$f_{RK}(v) = \frac{n_0 \mathcal{U}_{[0]}}{(\pi \theta^2 \kappa)^{3/2}} \left(1 + \frac{v^2}{\kappa \theta^2}\right)^{-\kappa-1} \exp\left(-\frac{\alpha^2 v^2}{\theta^2}\right) \quad (3)$$

combines standard κ -distribution with an exponential cutoff controlled by the positive cutoff parameter $\alpha < 1$, which must be small enough such that the RK retains the relevance of extended suprathermal tails in the observed distributions. In the normalization constant

$${}_{[n]} \mathcal{U}_{[1]} = 1/{}_{[1]} \mathcal{U}_{[n]} = U\left(\frac{3+n}{2}, \frac{3+n}{2} - \kappa, \alpha^2 \kappa\right), \quad (4)$$

where $U(a, b, z)$ is the Kummer U (or Tricomi) function (Scherer et al. 2019a). For $\alpha = 0$ the RK reduces to a standard κ -distribution, but a finite value $\alpha > \theta/c$ is required for the cutoff parameter in order to restrain for speeds lower than the speed of light c in vacuum and to minimize the unrealistic effects of superluminal particles with $v > c$ (Scherer et al. 2019b). For the solar wind plasma conditions assumed in this work, we find $\alpha = 0.012$ large enough to admit only negligible effects of superluminal particles (see condition $R_R < 1\%$ for the relative pressure contribution of superluminal particles in Scherer et al. (2019b)) satisfied in this case for any value of $\kappa > 0$. With this value for α we use the results in Scherer et al. (2019a) to evaluate the relative pressures P_i / P_M , and obtain

$$P_K / P_M = \frac{\kappa}{\kappa - 3/2}, \quad (5)$$

for a standard κ -distribution, and

$$P_{RK} / P_M = \kappa {}_{[2]} \mathcal{U}_{[0]}, \quad (6)$$

for RK , where $P_M = (1/2)n_0 m_e \theta^2 \simeq 1.2 \cdot 10^{-11} \text{ N m}^{-2}$ is the pressure given by Maxwellian limit in Eq. (2) (with the parameters from above)¹. Normalized pressures are plotted in Fig. 1 as a function of κ . Variations with κ of the normalized temperatures T_{RK} / T_M and T_K / T_M have the same profiles (not shown here) as P_{RK} / P_M and P_K / P_M , respectively. These profiles tend to coincide for large values of $\kappa > 10$, stating that this happens only in the absence of suprathermal populations. Most probably, very strong suprathermal tails corresponding to very low values of $\kappa < 1$ and to P_{RK} (or T_{RK}) values several orders of magnitude higher than P_M (or T_M) are not encountered in the solar wind. However, distributions corresponding to moderately low values of $\kappa < 10$ are more expected; these distributions include those slightly below the limit imposed by a standard κ -distribution (subscript K), i.e., $1 \leq \kappa \leq 3/2$, which are usually discarded if fitted with such a power-law model (Štverák et al. 2008). The bottom panel of Fig. 1 shows only this interval of κ values, where the pressure (or the temperature) is highly dependent on the model function for the observed velocity distribution. For instance, we can evaluate the relative pressure for $\kappa = 1$, and find $P_{RK} / P_M = 62.4$, implying for the absolute pressure $P_{RK} \simeq 0.7 \cdot 10^{-9} \text{ N m}^{-2} > P_M$, that means a boost of at least one order of magnitude under the effects of suprathermals. For $\kappa > 3/2$ these parameters can be evaluated for a standard K , but leading to overestimations. For instance, for $\kappa = 1.6$ we find $P_K / P_M = 16$, roughly two times higher than $P_{RK} / P_M = 7.65$, which leads to a realistic value $P_{RK} \simeq 0.9 \cdot 10^{-10} \text{ N m}^{-2}$.

¹ Comparison with the corresponding pressures in Scherer et al. (2019b) suggests a typo in their Eqs. (5)–(7), where the factor $3/2$ must be replaced by $1/2$.

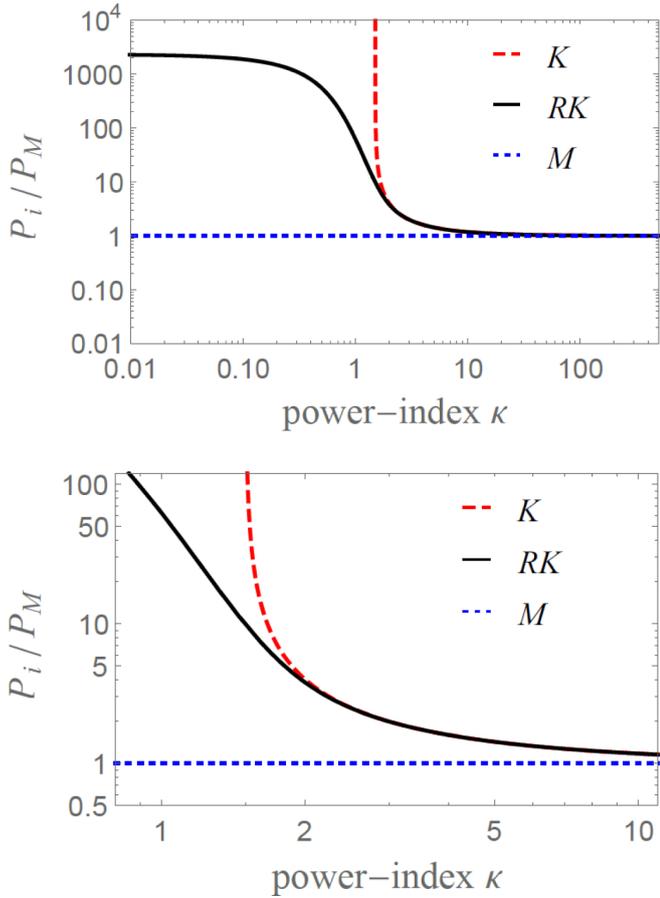


Fig. 1. Normalized pressure P_i/P_M for a standard (isotropic) κ -distribution ($i = K$, red dashed) and for the regularized κ -distribution ($i = RK$, black solid) in *bottom panel* with a zoom-in only for a restraint range of conditions relevant for the solar wind.

2.2. Anisotropic temperatures

For anisotropic temperatures $T_{\parallel} \neq T_{\perp}$, where \parallel and \perp denote directions relative to the uniform magnetic field lines, the observed velocity distributions can be described by a bi- κ -distribution (subscript BK ; Maksimovic et al. 2005; Štverák et al. 2008)

$$f_{BK}(v_{\parallel}, v_{\perp}) = \frac{n_0 \Gamma[\kappa]}{(\pi\kappa)^{3/2} \theta_{\parallel} \theta_{\perp}^2 \Gamma[\kappa - 1/2]} \left(1 + \frac{v_{\parallel}^2}{\kappa \theta_{\parallel}^2} + \frac{v_{\perp}^2}{\kappa \theta_{\perp}^2} \right)^{-\kappa-1}, \quad (7)$$

where $\theta_{\parallel, \perp}$ are given by the second-order moment $2k_B T_{\parallel, \perp}/m = \int d^3v v_{\parallel, \perp}^2 f_K = \kappa \theta_{\parallel, \perp}^2 / (\kappa - 3/2)$. The same $\theta_{\parallel, \perp} = (2k_B T_{M, \parallel, \perp}/m)^{1/2}$ describe the anisotropic (low-energy) core within a bi-Maxwellian (subscript BM)

$$f_{BM}(v_{\parallel}, v_{\perp}) = \frac{n_0}{\pi^{3/2} \theta_{\parallel} \theta_{\perp}^2} \exp\left(-\frac{v_{\parallel}^2}{\theta_{\parallel}^2} - \frac{v_{\perp}^2}{\theta_{\perp}^2}\right). \quad (8)$$

The regularized bi- κ -distribution (subscript RBK , see also Appendix A for a summary of these abbreviations)

$$f_{RBK}(v_{\parallel}, v_{\perp}) = \frac{n_0 {}_{[0]} \mathcal{U}_{\perp}}{(\pi\kappa)^{3/2} \theta_{\parallel} \theta_{\perp}^2} \left(1 + \frac{v_{\parallel}^2}{\kappa \theta_{\parallel}^2} + \frac{v_{\perp}^2}{\kappa \theta_{\perp}^2} \right)^{-\kappa-1} \times \exp\left(-\frac{\alpha^2 v_{\parallel}^2}{\theta_{\parallel}^2} - \frac{\alpha^2 v_{\perp}^2}{\theta_{\perp}^2}\right) \quad (9)$$

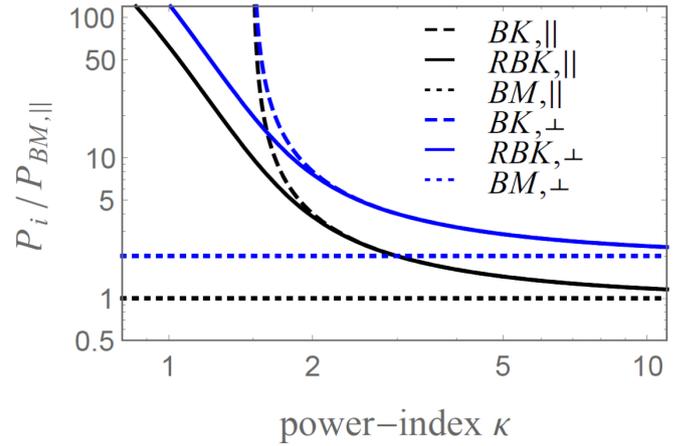


Fig. 2. Normalized pressure components, parallel $P_{i, \parallel}/P_{BM, \parallel}$ (black) and perpendicular $P_{i, \perp}/P_{BM, \parallel}$ (blue), and for BK ($i = BK$, dashed), RBK ($i = RBK$, solid), and BM ($i = BM$, dotted).

combines a BK function with an exponential cutoff factor controlled by a small $\alpha < 1$ (Scherer et al. 2019a). This cutoff parameter can be directly related to the maximum energy channel of the instrument measuring the particle distributions independently, which is the same for all pitch angles. Different cutoffs corresponding to parallel and perpendicular directions $\alpha_{\parallel} \neq \alpha_{\perp}$ may introduce additional anisotropies (Scherer et al. 2019a), which are not of interest in the present study. For $\alpha = 0$ the RBK reduces to a standard BK , but again a finite value $\alpha > \theta/c$ is required for the cutoff parameter to restrain to realistic speeds $v < c$ (Scherer et al. 2019b), such that for the solar wind conditions at 1 AU we find $\alpha = 0.012$. With this value for α and a moderate anisotropy $A = \theta_{\perp}^2/\theta_{\parallel}^2 = 2$ we evaluate the relative pressures $P_i/P_{BM, \parallel, \perp}$ for BK ($i = BK$) as follows:

$$\frac{P_{BK, \parallel}}{P_{BM, \parallel}} = \frac{\kappa}{\kappa - 3/2}, \quad \frac{P_{BK, \perp}}{P_{BM, \perp}} = \frac{\kappa}{\kappa - 3/2} A, \quad (10)$$

and for RBK ($i = RBK$)

$$\frac{P_{BRK, \parallel}}{P_{BM, \parallel}} = \kappa {}_{[2]} \mathcal{U}_{[0]}, \quad \frac{P_{BRK, \perp}}{P_{BM, \perp}} = \kappa {}_{[2]} \mathcal{U}_{[0]} A, \quad (11)$$

where ${}_{[m]} \mathcal{U}_{[n]} = {}_{[m]} \mathcal{U}_{\perp} / {}_{[n]} \mathcal{U}_{\parallel}$. For Maxwellian pressures we find $P_{BM, \parallel} = (1/2) n_0 m_e \theta_{\parallel}^2 \approx 1.2 \cdot 10^{-11} \text{ N m}^{-2}$ as given by the plasma parameters above, assuming $T_{\parallel} = T_e = 10^5 \text{ K}$ and $P_{BM, \perp} = A P_{BM, \parallel} = (1/2) n_0 m_e \theta_{\perp}^2 \approx 2.4 \cdot 10^{-11} \text{ N m}^{-2}$. These normalized pressures are plotted in Fig. 2 as a function of κ . Values obtained for the isotropic pressure in the previous section, i.e., for $\kappa = 1$ and $\kappa = 1.6$, are also relevant in this case for parallel pressure, while perpendicular pressure would be two times higher. Variations with κ of the normalized temperatures T_{RBK}/T_{BM} and T_{BK}/T_{BM} must show the same profiles (not shown) as P_{RBK}/P_{BM} and P_{BK}/P_{BM} , respectively. For opposite anisotropies $A < 1$ variations with κ are similar (not shown here); the only difference is that perpendicular components are lower than parallel components, $P_{\perp} < P_{\parallel}$ (or $T_{\perp} < T_{\parallel}$).

3. Beaming distributions

Firstly, we consider the general situation when a beaming population, for example, electron strahl, exhibits an intrinsic pressure (or temperature) anisotropy, and later we particularize for isotropic temperatures. The anisotropic components of pressure

(or temperature) show the same dependence on the model of velocity distribution, as already described above for non-drifting distributions, although in the solar wind frame (fixed to protons) the effective pressure (or temperature) is increased in parallel direction by the antisuward beaming speed ($k_B T_{\text{eff}} = k_B T + mV^2$). On the other hand, in this case we shall evaluate the higher order parameter giving the electron heat flux. The electron strahl, or beaming population, carries a significant amount of heat, which increases during (or after) the energetic events such as fast winds and CMEs, when the strahl relative density and beaming velocity may reach extreme values. However, the strahl also suffers a significant erosion with distance caused by a self-induced scattering triggered by the self-generated beam-plasma instabilities (Vocks et al. 2005; Maksimovic et al. 2005; Viñas et al. 2010; Verscharen et al. 2019; Tong et al. 2019). Values assumed in the present study for the relative densities and beaming velocities are in general relevant for various specific conditions in the solar wind, for example, slow or fast winds at various heliocentric distances.

3.1. Anisotropic temperatures

The electron strahl with anisotropic temperatures $T_{\parallel} \neq T_{\perp}$ can be reproduced by a drifting bi- κ -distribution (*DBK*) function (Nieves-Chinchilla & Viñas 2008; Viñas et al. 2010) given by

$$f_{\text{DBK}}(v_{\parallel}, v_{\perp}) = \frac{n_0 \Gamma[\kappa]}{(\pi\kappa)^{3/2} \theta_{\parallel} \theta_{\perp}^2 \Gamma[\kappa - 1/2]} \left(1 + \frac{(v_{\parallel} - V)^2}{\kappa \theta_{\parallel}^2} + \frac{v_{\perp}^2}{\kappa \theta_{\perp}^2} \right)^{-\kappa-1}, \quad (12)$$

where the beaming (drifting) velocity V is in parallel direction (along the magnetic field), and $\theta_{\parallel, \perp}$ are given by the second order moments $2k_B T_{\parallel, \perp}/m = \int d^3v v v_{\perp}^2 f_{\kappa} = \kappa \theta_{\parallel, \perp}^2 / (\kappa - 3/2)$. The same $\theta_{\parallel, \perp} = (2k_B T_{M, \parallel, \perp}/m)^{1/2}$ describe the anisotropic beaming core within a drifting bi-Maxwellian (*DBM*)

$$f_{\text{DBM}}(v_{\parallel}, v_{\perp}) = \frac{n_0}{\pi^{3/2} \theta_{\parallel} \theta_{\perp}^2} \exp\left(-\frac{(v_{\parallel} - V)^2}{\theta_{\parallel}^2} - \frac{v_{\perp}^2}{\theta_{\perp}^2}\right). \quad (13)$$

The regularized *DBK* (*RDBK*) given by

$$f_{\text{RDBK}}(v_{\parallel}, v_{\perp}) = \frac{n_0 [0] \mathcal{U}_{[1]}}{(\pi\kappa)^{3/2} \theta_{\parallel} \theta_{\perp}^2} \left(1 + \frac{(v_{\parallel} - V)^2}{\kappa \theta_{\parallel}^2} + \frac{v_{\perp}^2}{\kappa \theta_{\perp}^2} \right)^{-\kappa-1} \times \exp\left(-\frac{\alpha^2 (v_{\parallel} - V)^2}{\theta_{\parallel}^2} - \frac{\alpha^2 v_{\perp}^2}{\theta_{\perp}^2}\right) \quad (14)$$

combines a standard *DBK* with an exponential cutoff factor controlled by $\alpha < 1$ (Scherer et al. 2019a). For $\alpha = 0$ the *RDBK* reduces to a *DBK*, but, again, a finite value $\alpha > \theta/c$ is required for the cutoff parameter in order to restrain to realistic speeds $v < c$ (Scherer et al. 2019b). We use the general expression of the vector heat flux, for example, Eq. (6.5) in Paschmann et al. (1998)², and derive the heat flux $q = q_{\parallel}$ given by the strahl population along the magnetic field direction. It results from the general definition in statistical plasma physics (Paschmann et al. 1998) for the solar wind conditions, where the relative (counter-)drift of the central (core) population is in general very small, such that the suprathermal electron beaming

populations remain the main contributor to the antisuward electron heat flux (Gary et al. 1999; Dudík et al. 2017). This is what we obtain, for a *DBM*,

$$q_{\text{DBM}} = \frac{1}{2} n m_e V \left(\frac{3}{2} \theta_{\parallel}^2 + \theta_{\perp}^2 + V^2 \right), \quad (15)$$

see also Gary et al. (1999)³, for a *DBK*,

$$q_{\text{DBK}} = \frac{1}{2} n m_e V \left[\frac{\kappa}{\kappa - 3/2} \left(\frac{3}{2} \theta_{\parallel}^2 + \theta_{\perp}^2 \right) + V^2 \right], \quad (16)$$

and for a *RDBK*,

$$q_{\text{RDBK}} = \frac{1}{2} n m_e V \left[\kappa_{[2]} \mathcal{U}_{[0]} \left(\frac{3}{2} \theta_{\parallel}^2 + \theta_{\perp}^2 \right) + V^2 \right]. \quad (17)$$

A comparison with previous results, for example, Eqs. (12) and (41) in Scherer et al. (2019a), suggests the existence of a typo affecting the contribution of perpendicular kinetic energy in their equations, but also the absence of the common factor 1/2, as resulting from their Eq. (7e) for the general definition of the heat flux². Normalizing to Maxwellian limit, we find

$$\frac{q_{\text{DBK}}}{q_{\text{DBM}}} = \frac{\left(\frac{3}{2} + A\right) \frac{\kappa}{\kappa - 3/2} + W^2}{\frac{3}{2} + A + W^2}, \quad (18)$$

and

$$\frac{q_{\text{RDBK}}}{q_{\text{DBM}}} = \frac{\left(\frac{3}{2} + A\right) \kappa_{[2]} \mathcal{U}_{[0]} + W^2}{\frac{3}{2} + A + W^2}, \quad (19)$$

in terms of anisotropy $A = \theta_{\perp}^2 / \theta_{\parallel}^2$ and normalized drift $W = V / \theta_{\parallel}$.

To evaluate the heat flux, we first consider $A = \theta_{\perp}^2 / \theta_{\parallel}^2 = 2 > 1$, and temperatures characteristic to a hotter strahl $T_{\parallel} = 3 \cdot 10^5$ K may be few times higher than temperature of central core+halo component (Viñas et al. 2010). The relative drift may vary with number density of beaming electrons. Thus, a low number density $n_1 = 0.02 n_0$ and low beaming velocity $W_1 = V_1 / \theta_{\parallel} = 0.5$ can be specific to slow winds at low distances < 1 AU, or for any kind of slow or fast winds at large distances > 1 AU, as suggested in Maksimovic et al. (2005). In this case, for Maxwellian limit of the parallel (non-drifting) pressure, we find $P_{\text{BM}, \parallel} = (1/2) n_1 m_e \theta_{1, \parallel}^2 \approx 0.07 \cdot 10^{-11} \text{ N m}^{-2}$ and $P_{\text{BM}, \parallel}^{\text{eff}} = (1 + 2W^2) P_{\text{BM}, \parallel} \approx 0.11 \cdot 10^{-11} \text{ N m}^{-2}$, while the heat flux, as given by Eq. (15), is $q_{\text{DBM}} \approx 0.45 \cdot 10^{-5} \text{ W m}^{-2}$. The relative heat flux is plotted in Fig. 3, top panel, for both the *DBK* and *RDBK* distribution functions. For lower values of $\kappa < 3/2$ the heat flux may significantly increase up to, for example, $q_{\text{RDBK}} \approx 0.25 \cdot 10^{-3} \text{ W m}^{-2}$ for $\kappa = 1$. A comparison can also be made for $\kappa > 3/2$, when, for instance, for $\kappa = 1.6$, we find $q_{\text{DBK}} \approx 0.35 \cdot 10^{-4} \text{ W m}^{-2}$ and a reduced $q_{\text{RDBK}} \approx 0.25 \cdot 10^{-4} \text{ W m}^{-2}$ (by minimizing the unrealistic contribution of superluminal particles with $v > c$).

More intense strahls with $n_2 = 0.1 n_0$ and $W_2 = V_2 / \theta_{\parallel} = 2.5$ are characteristic to fast winds at low distances < 1 AU (Maksimovic et al. 2005; Štverák et al. 2009). In the Maxwellian limit for the parallel component of the (non-drifting) pressure we find $P_{\text{BM}, \parallel} = (1/2) n_2 m_e \theta_{2, \parallel}^2 \approx 0.4 \cdot 10^{-11} \text{ N m}^{-2}$ and $P_{\text{BM}, \parallel}^{\text{eff}} = (1 + 2W^2) P_{\text{BM}, \parallel} \approx 5.0 \cdot 10^{-11} \text{ N m}^2$, while for the heat

² Equation (7e) in Scherer et al. (2019a) differs by a factor of 1/2 and is not consistent with the definition ($mv^2/2$) used in our paper for the kinetic energy of plasma particles.

³ This expression is also adopted in Gary et al. (1999), even though it is not consistent with their definition (mv^2) for the kinetic energy of plasma particles.

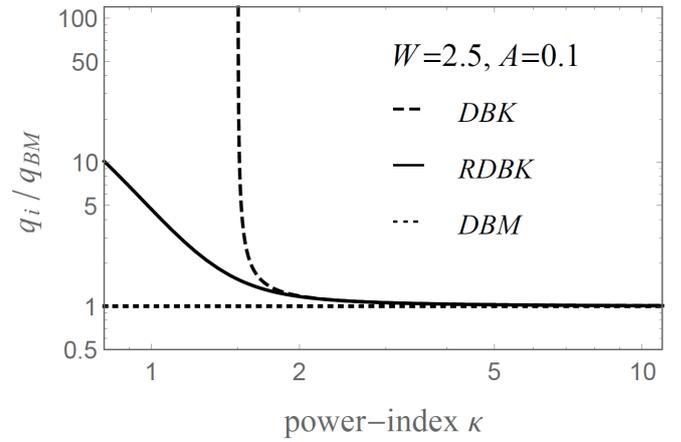
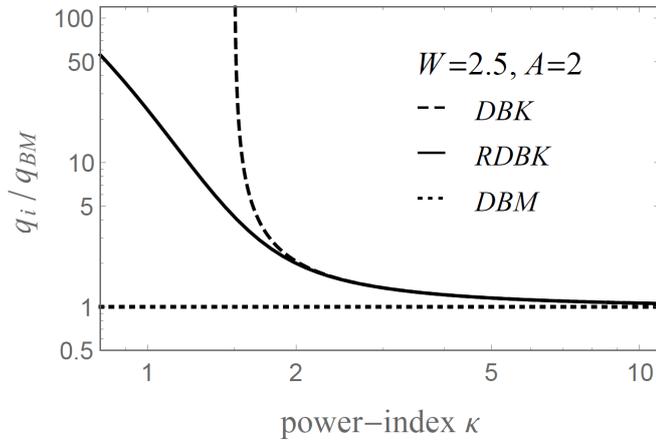
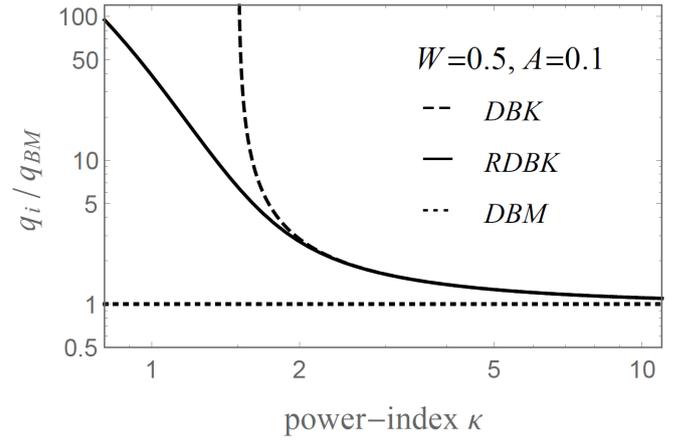
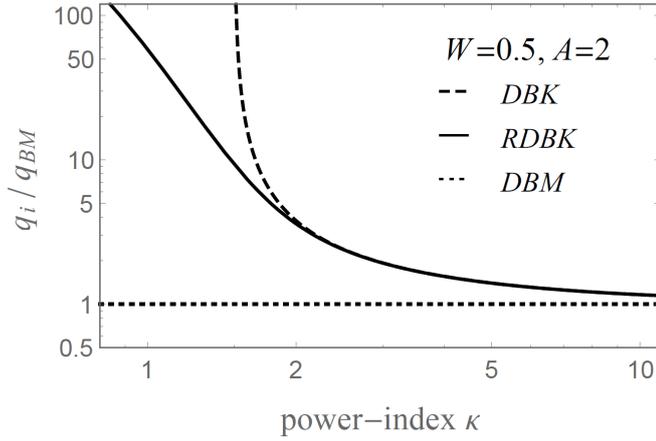


Fig. 3. Normalized heat flux for $A = 2$, $W = 0.5$ (top panel) and $W = 2.5$ (bottom panel), evaluated for DBK ($i = DBK$, dashed) and $RDBK$ ($i = RDBK$, solid).

Fig. 4. Normalized heat flux for $A = 0.1$, $W = 0.5$ (top panel) and $W = 2.5$ (bottom panel), evaluated for DBK ($i = DBK$, dashed) and $RDBK$ ($i = RDBK$, solid).

flux $q_{DBM} \approx 0.3 \cdot 10^{-3} \text{ W m}^{-2}$. The relative heat flux is plotted in Fig. 3, bottom panel, for both the DBK and $RDBK$ distribution functions, and for $\kappa = 1$ we find a higher heat flux $q_{RDBK} \approx 0.7 \cdot 10^{-2} \text{ W m}^{-2}$. For $\kappa = 1.6$ we find a lower difference between $q_{DBK} \approx 1.1 \cdot 10^{-3} \text{ W m}^{-2}$ and $q_{RDBK} \approx 0.9 \cdot 10^{-3} \text{ W m}^{-2}$. In any of these cases, the heat flux decreases with increasing κ -index, i.e., with reducing the contribution of suprathermal or non-Maxwellian electrons in the strahl population. We note that for large $\kappa \rightarrow \infty$ the strahl distribution reduces to a quasi-thermal Maxwellian representation, which should not be confused with the Maxwellian core in the central population of electrons observed in the solar wind.

In the opposite case, when the strahl exhibits a reversed anisotropy $A = 0.1 < 1$, we assume the same $T_{\parallel} = 3 \cdot 10^5 \text{ K}$ and parallel pressure remains the same, but the heat flux may significantly change: see, for instance, Fig. 4, bottom panel. For less energetic strahls ($W = 0.5$), we obtain $q_{DBM} \approx 0.23 \cdot 10^{-5} \text{ W m}^{-2}$, as given for reference by Eq. (15). The top panel of Fig. 4 shows the variation with κ of the relative heat flux, and both profiles for DBK and $RDBK$ are similar to those obtained in Fig. 3 top panel. Thus, for $\kappa < 1.5$ the heat flux values remain much higher than those predicted for a BM ; for example, for $\kappa = 1$ we find $q_{RDBK} \approx 0.24 \cdot 10^{-3} \text{ W m}^{-2}$, while for $\kappa = 1.6$ $q_{DBK} \approx 0.35 \cdot 10^{-4} \text{ W m}^{-2}$ and $q_{RDBK} \approx 0.25 \cdot 10^{-4} \text{ W m}^{-2}$ is slightly reduced. The situation is much different for more energetic beams with $W = 2.5$ (bottom panel) when $q_{DBM} \approx 0.05 \cdot 10^{-3} \text{ W m}^{-2}$ and values obtained for the heat flux in the limit of very small val-

ues of $\kappa \rightarrow 1$ may be one order of magnitude lower than that obtained in top panel; this also highly contrasts with the corresponding result in Fig. 3, bottom panel. In this case, for $\kappa = 1$ we find $q_{RDBK} \approx 0.7 \cdot 10^{-2} \text{ W m}^{-2}$, while for $\kappa = 1.6$ we get $q_{DBK} \approx 1.6 \cdot 10^{-3} \text{ W m}^{-2}$ and $q_{RDBK} \approx 0.12 \cdot 10^{-3} \text{ W m}^{-2}$.

3.2. Isotropic temperatures

If the temperature of beaming population is (almost) isotropic, i.e., $T_{\parallel} \approx T_{\perp}$, the heat flux can be calculated (for $A = 1$) using the same expressions (15)–(17) for their absolute values or (18)–(19) for their relative values. Profiles of the (relative) heat-flux dependence on κ are shown in Fig. 5 and are very similar to those obtained in Fig. 3. Values provided for the heat flux by the standard κ -distribution and by the regularized model are also comparable with those obtained for $A = 2$. The contrast should increase for higher anisotropies $A \gg 1$.

4. Discussion and conclusions

A macroscopic description of collisionless plasmas out of thermal equilibrium is not always straightforward because it is highly dependent on the model function describing the velocity distribution of plasma particles. For κ -distribution functions a macroscopic description is simply altered by the divergence of higher order moments, such as temperature, pressure, or heat flux, for plasma states with significant suprathermal populations.

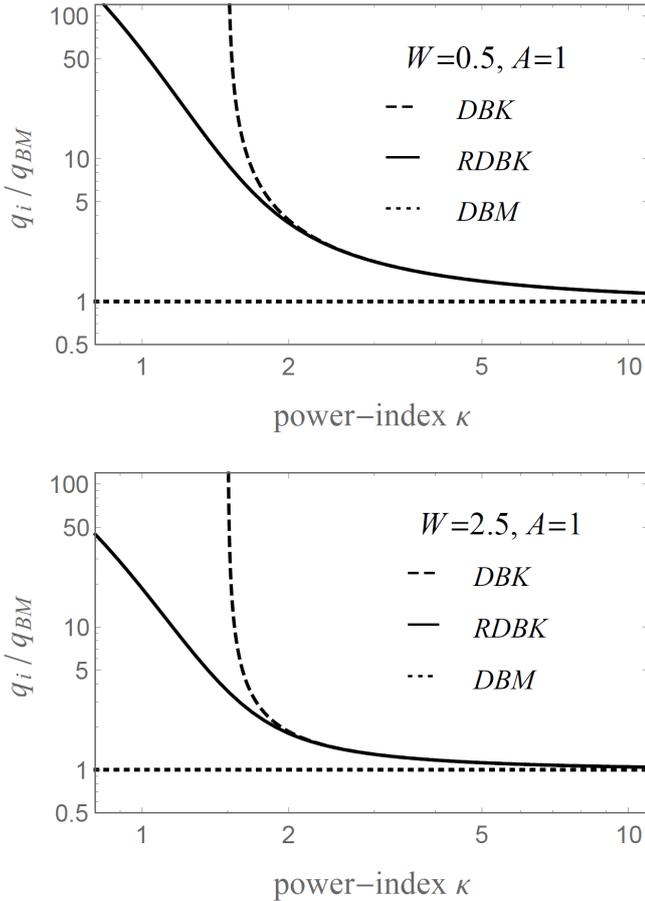


Fig. 5. Normalized heat flux for $A = 1$, $W = 0.5$ (top panel) and $W = 2.5$ (bottom panel), evaluated for DBK ($i = \text{DBK}$, dashed) and RDBK ($i = \text{RDBK}$, solid).

For that reason the distributions with hard suprathermal tails are discarded from the observational analysis. In the absence of relativistic corrections, moments of standard κ -distributions may also include unphysical contributions of superluminal particles with speeds larger than the speed of light in vacuum.

Regularized κ -distribution (RK) functions (Scherer et al. 2017) were introduced to overcome these limitations, and in this work we invoked these models to evaluate the pressure and heat flux of electron populations for solar wind conditions at 1 AU. The heat-flux values summarized in Tables 1 and 2 and discussed in the previous section are all in the range of the electron heat flux measured in the solar wind by different missions, from very early Imp 6, 7, or 8, to Ulysses and Wind spacecraft (Feldman et al. 1975; Cuperman 1980; Scime et al. 2001; Pagel et al. 2005). Of particular interest are plasma states dominated by the suprathermal populations, for example, for $\kappa < 2$, when the pressure or the heat flux transported by the beaming populations may be at least one order of magnitude higher than estimates made by ignoring suprathermal populations. In this case, the corresponding moments obtained for a standard κ -distribution are defined only for $\kappa > 3/2$ and can markedly overestimate predictions made by the RKs. We have to point out that values chosen for the cutoff parameter α may have various motivations; for example, for $R_R < 1\%$ Scherer et al. (2019b) have considered reducing the unrealistic effects of superluminal particles to a reasonably acceptable level. With the parameters adopted in our study this condition may indeed let α be closer to $\theta/c \simeq 0.005$, than our alternative $\alpha = 0.012$ that rather reduces

Table 1. Heat flux (W m^{-2}) for $A = 2$.

Model	DBM $\kappa \rightarrow \infty$	DBK $\kappa = 1.6$	RDBK $\kappa = 1.6$	RDBK $\kappa = 1$
$W = 0.5$	0.45×10^{-5}	0.35×10^{-4}	0.25×10^{-4}	0.25×10^{-3}
$W = 2.5$	0.3×10^{-3}	1.1×10^{-3}	0.9×10^{-3}	0.7×10^{-2}

Table 2. Heat flux (W m^{-2}) for $A = 0.1$.

Model	DBM $\kappa \rightarrow \infty$	DBK $\kappa = 1.6$	RDBK $\kappa = 1.6$	RDBK $\kappa = 1$
$W = 0.5$	0.23×10^{-5}	0.35×10^{-4}	0.25×10^{-4}	0.24×10^{-3}
$W = 2.5$	0.05×10^{-3}	1.6×10^{-3}	0.12×10^{-3}	0.7×10^{-2}

the implications of superluminal particles. For slightly higher values of α our results move toward a Maxwellian prediction, while lower values increase the effects of suprathermals, especially for $\kappa \leq 2$. Other constraints on the α value may be introduced by the observational spectra showing dropping shoulders, as an indication of different populations present at higher energies but with lower intensities.

We can conclude trusting that based on the RK model functions, a realistic and physically well-defined parametrization becomes possible even for collisionless plasmas with various nonthermal features such as suprathermal populations or kinetic anisotropies. Values derived in this work for macroscopic properties of RK electrons (either isotropic or anisotropic) can be confronted with those obtained from a direct numerical integration of observational data (Paschmann et al. 1998). Once accurately estimated, these parameters can be used as initial conditions as well as boundary conditions, for example, in plasma transport models for the evolution of the expanding solar wind in heliosphere (Pierrard et al. 2001, 2011) and for its interaction with planetary atmospheres.

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Appendix A: Abbreviations

The abbreviations used for various distribution functions are explained and summarized below in Tables A.1 and A.2.

Table A.1. Non-drifting distributions.

Isotropic temperature	Anisotropic temperature
<i>K</i> : κ -distribution	<i>BK</i> : Bi- κ -distribution
<i>M</i> : Maxwellian distribution	<i>BM</i> : Bi-Maxwellian dist.
<i>RK</i> : Regularized κ -dist.	<i>RBK</i> : Regularized Bi- κ dist.

Table A.2. Drifting (or beaming) distributions (apply for both cases of anisotropic or isotropic temperatures, see Sect. 3).

<i>DBK</i> : Drifting Bi- κ -distribution
<i>DBM</i> : Drifting Bi-Maxwellian distribution
<i>RDBK</i> : Regularized Drifting Bi- κ -distribution