

Probing fragmentation and velocity sub-structure in the massive NGC 6334 filament with ALMA[★]

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ABSTRACT

Context. *Herschel* imaging surveys of galactic interstellar clouds support a paradigm for low-mass star formation in which dense molecular filaments play a crucial role. The detailed fragmentation properties of star-forming filaments remain poorly understood, however, and the validity of the filament paradigm in the intermediate- to high-mass regime is still unclear.

Aims. Here, following up on an earlier 350 μm dust continuum study with the ArTéMiS camera on the APEX telescope, we investigate the detailed density and velocity structure of the main filament in the high-mass star-forming region NGC 6334.

Methods. We conducted ALMA Band 3 observations in the 3.1 mm continuum and of the $\text{N}_2\text{H}^+(1-0)$, $\text{HC}_3\text{N}(36-35)$, $\text{HNC}(1-0)$, $\text{HC}_3\text{N}(10-9)$, $\text{CH}_3\text{CCH}(6-5)$, and $\text{H}_2\text{CS}(3-2)$ lines at an angular resolution of $\sim 3''$, corresponding to 0.025 pc at a distance of 1.7 kpc.

Results. The NGC 6334 filament was detected in both the 3.1 mm continuum and the N_2H^+ , HC_3N , HC_5N , CH_3CCH , and H_2CS lines with ALMA. We identified twenty-six compact (< 0.03 pc) dense cores at 3.1 mm and five velocity-coherent fiber-like features in N_2H^+ within the main filament. The typical length (~ 0.5 pc) of, and velocity difference (~ 0.8 km s⁻¹) between, the fiber-like features of the NGC 6334 filament are reminiscent of the properties for the fibers of the low-mass star-forming filament B211/B213 in the Taurus cloud. Only two or three of the five velocity-coherent features are well aligned with the NGC 6334 filament and may represent genuine, fiber sub-structures; the other two features may trace accretion flows onto the main filament. The mass distribution of the ALMA 3.1 mm continuum cores has a peak at $\sim 10 M_\odot$, which is an order of magnitude higher than the peak of the prestellar core mass function in nearby, low-mass star-forming clouds. The cores can be divided into seven groups, closely associated with dense clumps seen in the ArTéMiS 350 μm data. The projected separation between ALMA dense cores (0.03–0.1 pc) and the projected spacing between ArTéMiS clumps (0.2–0.3 pc) are roughly consistent with the effective Jeans length (0.08 ± 0.03 pc) in the filament and a physical scale of about four times the filament width, respectively, if the inclination angle of the filament to line of sight is $\sim 30^\circ$. These two distinct separation scales are suggestive of a bimodal fragmentation process in the filament.

Conclusions. Despite being one order of magnitude denser and more massive than the Taurus B211/B213 filament, the NGC 6334 filament has a density and velocity structure that is qualitatively very similar. The main difference is that the dense cores embedded in the NGC 6334 filament appear to be an order of magnitude denser and more massive than the cores in the Taurus filament. This suggests that dense molecular filaments may evolve and fragment in a similar manner in low- and high-mass star-forming regions, and that the filament paradigm may hold in the intermediate-mass (if not high-mass) star formation regime.

Key words. ISM: clouds – ISM: structure – stars: formation

1. Introduction

Herschel imaging observations of galactic molecular clouds reveal an omnipresence of filamentary structures and suggest that filaments dominate the mass budget of the dense molecular gas where stars form (André et al. 2010; Molinari et al. 2010; Hill et al. 2011; Schisano et al. 2014; Könyves et al. 2015). At least in the nearby clouds of the Gould Belt, detailed studies of the radial column density profiles have found a common inner filament width of ~ 0.1 pc, with a dispersion of a factor $\lesssim 2$, when averaged over the filament crests (Arzoumanian et al. 2011, 2019; Palmeirim et al. 2013; Koch & Rosolowsky 2015). Furthermore, most of the prestellar cores identified with *Herschel* are found to be embedded within such filamentary structures, showing that dense molecular filaments are the main sites of solar-type star formation (e.g., Könyves et al. 2015, 2019; Marsh et al. 2016). Overall, the *Herschel* findings in nearby clouds support a filament paradigm for solar-type star formation in two main steps

(cf. André et al. 2014, 2017; Inutsuka et al. 2015): first, multiple large-scale compressions of interstellar material in supersonic flows generate a pervasive web of ~ 0.1 -pc wide filaments in the cold interstellar medium (ISM); second, the densest filaments fragment into prestellar cores by gravitational instability near or above the critical mass per unit length of nearly isothermal gas cylinders, $M_{\text{line,crit}} = 2c_s^2/G \sim 16 M_\odot \text{pc}^{-1}$, where $c_s \sim 0.2$ km s⁻¹ is the isothermal sound speed for molecular gas at $T \sim 10$ K.

Since molecular filaments are also known to be present in other galaxies (cf. Fukui et al. 2015), this paradigm may potentially have implications for star formation on galaxy-wide scales. The star formation efficiency in dense molecular gas is indeed observed to be remarkably uniform over a wide range of scales from pc-scale filaments and clumps to entire galactic disks (Gao & Solomon 2004; Lada et al. 2010, 2012; Shimajiri et al. 2017; Zhang et al. 2019), with possible deviations in extreme environments, such as the central molecular zone (Longmore et al. 2013). Assuming that all star-forming filaments have similar inner widths, as seems to be the case in nearby clouds (Arzoumanian et al. 2019), it is argued that the microphysics of filament fragmentation into prestellar cores may ultimately

[★] The final data used in the paper (FITS) are only available at the CDS via anonymous ftp to cdsarc.u-strasbg.fr (130.79.128.5) or via <http://cdsarc.u-strasbg.fr/viz-bin/cat/J/A+A/632/A83>

be responsible for this quasi-universal star formation efficiency (André et al. 2014; Shimajiri et al. 2017).

The validity and details of the two-step filament paradigm remain controversial, however, especially beyond the Gould Belt and in high-mass star-forming clouds. In particular, an alternative scenario is proposed based on the notion of global hierarchical cloud collapse (e.g. Vázquez-Semadeni et al. 2019), which is especially attractive in the case of high-mass star formation to account for the structure of strongly self-gravitating hub-filament systems, where a massive cluster-forming hub is observed at the center of a converging network of filaments (Myers 2009; Peretto et al. 2013). In this alternative scenario, most if not all filaments would be generated by gravitational effects, as opposed to large-scale compression, and would represent accretion flows onto dense hubs.

The evolution and detailed fragmentation manner of star-forming filaments are also poorly understood. In particular, the mere existence of massive, ~ 0.1 -pc wide, filaments with masses per unit length M_{line} that exceed the critical line mass¹ of an isothermal filament $M_{\text{line,crit}}$ by one to two orders of magnitude is a paradox. Indeed, such filaments may be expected to undergo rapid radial contraction into spindles before any fragmentation into prestellar cores (cf. Inutsuka & Miyama 1997). A possible solution for this paradox is that massive filaments accrete background cloud material while evolving (cf. Palmeirim et al. 2013; Shimajiri et al. 2019), and that this accretion process drives magneto-hydrodynamic (MHD) waves, generating sub-structure within dense filaments, and leading to a dynamical equilibrium (Hennebelle & André 2013), where M_{line} approaches the virial mass per unit length $M_{\text{line,vir}} = 2\sigma_{\text{1D}}^2/G$ (Fiege & Pudritz 2000, where σ_{1D} is the one-dimensional velocity dispersion). The detection through molecular line observations of velocity-coherent fiber-like sub-structure within several nearby supercritical filaments (Hacar et al. 2013, 2018) may possibly be the manifestation of such a process.

The physical origin of observed fibers is nevertheless still under debate (Tafalla & Hacar 2015; Smith et al. 2014; Clarke et al. 2017, 2018), and it is not yet clear whether dense molecular filaments in massive star-forming regions have similar characteristics to those observed in nearby clouds. Our recent APEX/ArTéMiS 350 μm study of the massive star-forming complex NGC 6334 showed that the main filament of the complex has an observed line mass of $\sim 1000 M_{\odot} \text{pc}^{-1}$, consistent within uncertainties with the estimated virial mass per unit length $\sim 500 M_{\odot} \text{pc}^{-1}$, and an inner width $\sim 0.15 \pm 0.05$ pc all along its length (André et al. 2016), within $\sim 50\%$ of the typical inner filament width observed with *Herschel* in nearby clouds (Arzoumanian et al. 2019). NGC 6334 is a very active star-forming region at a distance $d \sim 1.7$ kpc, which contains 150 OB stars and more than 2000 young stellar objects (Persi & Tapia 2008; Russeil et al. 2013; Willis et al. 2013; Tigé et al. 2017), for a total gas mass of $\sim 2.2 \times 10^5 M_{\odot}$. Here, thanks to the high angular resolution and sensitivity of Atacama Large Millimeter Array (ALMA) data at 3 mm, we investigate the density and velocity sub-structure of the massive NGC 6334 filament and compare the results with those obtained for nearby, lower-mass star-forming filaments.

The paper is organized as follows. In Sect. 2, we describe our ALMA observations of NGC 6334 in both the $\text{N}_2\text{H}^+(1-0)$, $\text{HC}_3\text{N}(10-9)$, $\text{HC}_3\text{N}(36-35)$, $\text{CH}_3\text{CCH}(6_0-5_0)$, $\text{H}_2\text{CS}(3_{1,2}-2_{1,1})$

lines and the 3.1 mm continuum. In Sect. 3, we show the spatial distributions of the detected molecular lines and continuum emission. We also analyze the velocity features observed in $\text{N}_2\text{H}^+(1-0)$ emission and extract dense structures such as compact cores and fiber-like components. In Sect. 4, we discuss the evidence of a bimodal fragmentation pattern in the NGC 6334 filament, emphasize the presence of both unusually massive cores and fiber-like velocity-coherent components in the filament, and comment on the possible origin of these multiple velocity components. Our conclusions are summarized in Sect. 5.

2. ALMA observations of the NGC 6334 filament

We carried out ALMA Cycle 3 observations in Band 3 toward NGC 6334 with both the 12 m antennas (C36-2 configuration) and the Atacama Compact Array (ACA) 7 m antennas, as part of project 2015.1.01404.S (PI: Ph. André). We imaged the main filament in the NGC 6334 region using a 17-pointing mosaic with the 12 m array and a 8-pointing mosaic with the ACA array. The $\text{N}_2\text{H}^+(1-0)$ line was observed in narrow-band mode at a frequency resolution of 61.035 kHz, corresponding to ~ 0.2 km s^{-1} . The HC_3N ($J = 36-35$) line was included in the same narrow-band, high spectral resolution window. The 3.1 mm continuum emission was observed using three wide bands, each covering a bandwidth of 1875.0 MHz. The HNC (1-0), HC_3N (10-9), CH_3CCH (6_0-5_0), and H_2CS ($3_{1,2}-2_{1,1}$) lines were included in these wide bands and observed at a frequency resolution of 31.250 MHz, corresponding to ~ 100 km s^{-1} . The 12 m-array observations were carried out on 23, 24, 26 January 2016 with 42, 46, and 37 antennas, respectively, and projected baseline lengths ranging from 11.7 to 326.8 k λ . The ACA observations were carried out between 16 March 2016 and 8 July 2016 with 7-9 antennas and projected baseline lengths ranging from 4.5 to 48.2 k λ . In the calibration process, we made additional flagging of a few antennas that had too low gain or showed large amplitude dispersion in time. The bandpass calibration was achieved by observing the quasar J1617-5848 with the 12 m array and the three quasars J1924-2914, J1517-2422, and J1427-4206 with the ACA array. The complex gain calibration was carried out using the quasar J1713-3418 with the 12 m array and the quasar J1717-3342 with ACA. Absolute flux calibration was achieved by observing two solar system objects (Callisto and Titan) with the 12 m array and planets (Mars and Neptune) with ACA. Calibration and data reduction were performed with the Common Astronomy Software Application (CASA) package (version 4.5.3 for 12 m data calibration and imaging, version 4.6.0 for 7 m data calibration). For imaging, we adopted the Briggs weighting scheme with a robust parameter of 0.5, as a good compromise between maximizing angular resolution and sensitivity to extended structures. The resulting synthesized beam sizes ($\sim 3''$, corresponding to 0.025 pc) and rms noise levels for each line and the continuum are summarized in Table 1.

With a minimum projected baseline length of 4.5 k λ , our ALMA+ACA observations are estimated to be sensitive to angular scales up to $\sim 37''$ (corresponding to ~ 0.3 pc) at the 10% level (Wilner & Welch 1994). For comparison, we expect the transverse size of any filament sub-structures to be smaller than the ~ 0.15 pc inner width of the NGC 6334 filament (André et al. 2016). Likewise with a minimum projected baseline of 11.7 k λ , our 12 m-only observations are sensitive to angular scales up to $\sim 14''$ (corresponding to ~ 0.1 pc) at the 10% level. For comparison, the scales of individual dense cores

¹ The critical line mass is $M_{\text{line,crit}} = 2c_s^2/G$, where c_s is the isothermal sound speed, corresponding to $M_{\text{line,crit}} \approx 33 M_{\odot}/\text{pc}$ for a molecular gas temperature $T = 20$ K.

Table 1. Parameters for the ALMA observations.

	Rest freq. GHz	θ_{beam} (PA) " × " (°)	rms mJy beam ⁻¹	dV km s ⁻¹
3.1 mm continuum (12 m + 7 m)	–	2.99 × 1.97 (82.6)	0.2	–
3.1 mm continuum (12 m)	–	2.80 × 1.84 (83.4)	0.14	–
HNC (1–0, $E_u = 4.35$ K)	90.663572	3.82 × 2.48 (85.9)	0.3	60
HC ₃ N (10–9, $E_u = 24.015$ K)	90.978989	3.81 × 2.47 (85.9)	0.27	60
N ₂ H ⁺ (1–0, $E_u = 4.472$ K)	93.176265	3.78 × 2.40 (84.1)	4.4	0.2
HC ₅ N (36–35, $E_u = 80.504$ K)	93.188126	3.77 × 2.39 (84.1)	5.4	0.2
CH ₃ CCH (6 ₀ –5 ₀ , $E_u = 17.225$ K)	102.54798	3.75 × 2.36 (79.1)	0.36	60
H ₂ CS (3 _{1,2} –2 _{1,1} , $E_u = 23.213$ K)	104.61699	3.40 × 2.17 (84.3)	0.22	60

are at most ~ 0.1 pc (Könyves et al. 2015). Moreover, simulations of ALMA observations performed with CASA, using the same set of baselines as the real ALMA+ACA data, show that interferometric-filtering effects do not generate spurious sub-structures within the filament (see Fig. A.1).

3. Results and analysis

In this section, we show the results of our ALMA 3.1 mm continuum and N₂H⁺ (1–0) line observations, and then extract compact 3.1 mm continuum sources and fiber-like velocity-coherent structures from the ALMA data.

3.1. 3.1 mm continuum emission

Our ALMA 3.1 mm continuum map of the NGC 6334 filament region is shown in panel a of Fig. 1 (color scale and contours). It is also overlaid as contours on a *Spitzer* 8 μm emission map in panel d of Fig. 1. Hereafter, we call the map resulting from the combination of APEX/ArTéMiS 350 μm (André et al. 2016) and *Herschell*/SPIRE 350 μm (Russeil et al. 2013) data the ArTéMiS 350 μm map. The counterpart of the dense filament seen in the ArTéMiS 350 μm map can be recognized in the northern part of the ALMA 3.1 mm continuum map. In the southern part of the field, the filament is not clearly detected in the ALMA 3.1 mm continuum map, but a shell-like structure can be seen. Conversely, the shell-like structure is not seen in the ArTéMiS 350 μm map. A compact HII region associated with a 4.9 GHz radio continuum source labeled source D in Rodriguez et al. (1982) lies close to the contours of this shell-like structure. The shell-like structure in the ALMA 3.1 mm continuum map also coincides with bright, extended mid-infrared emission detected at 8 μm with *Spitzer* as shown in Fig. 1d. We conclude that the 3.1 mm continuum emission in the shell-like structure is most likely dominated by free–free emission from the compact HII region around the luminous embedded star associated with source D.

3.2. Molecular lines

3.2.1. Spatial distribution of the N₂H⁺ (1–0) emission

Figure 1b shows the integrated intensity map of the N₂H⁺ (1–0) line emission around the NGC 6334 main filament. It can be seen that the prominent dusty filament in the ArTéMiS 350 μm map (Fig. 1c) is very well traced by the ALMA N₂H⁺ (1–0) data. In addition, a few N₂H⁺ blobs can be recognized outside the main filament.

3.2.2. Spatial distributions of other molecular line tracers

The ALMA maps obtained in all other molecular line tracers, except HNC(1–0), also show the NGC 6334 filament (see Fig. A.2a–e). The map obtained in HNC(1–0) differs from the other maps in that it shows a rather scattered distribution of discrete blobs. Due to a lower effective excitation density (cf. Shirley 2015), the emission from the HNC(1–0) transition may be more extended than the emission in the other dense gas tracers observed here and may be resolved out due to interferometric filtering effects, even with ACA. The maps obtained in CH₃CCH(6₀–5₀), H₂CS(3_{1,2}–2_{1,1}), HC₃N(10–9), and HC₅N(36–35) emission appear to trace the same filamentary structure as detected in N₂H⁺ (1–0). The HC₅N(36–35) line has a high upper state energy of 80.504 K, implying that some of the gas in the main filament has a high temperature and/or density. Indeed, André et al. (2016) estimated the average column density and average volume density of the entire NGC 6334 filament to be $1\text{--}2 \times 10^{23}$ and 2.2×10^5 cm⁻³, respectively. This exceeds the (column) density values observed by Palmeirim et al. (2013) for the low-mass star-forming filament B211/B213 in Taurus ($\sim 1.4 \times 10^{22}$ cm⁻² and 4.5×10^4 cm⁻³) by an order of magnitude.

3.2.3. Observed line profiles

The N₂H⁺ (1–0) and HC₅N(36–35) line observations were obtained in narrow-band mode with a velocity resolution of 0.2 km s⁻¹, which allows us to investigate the velocity structure of the NGC 6334 filament.

Here, we discuss the velocity profiles obtained in N₂H⁺ (1–0) and HC₅N(36–35).

Figure 2 shows the N₂H⁺ (1–0) and HC₅N(36–35) spectra averaged over the pixels where significant emission was detected above the 5 σ level. Two peaks at $V_{\text{sys}} = -2.6$ and $V_{\text{sys}} = -1.0$ km s⁻¹ can be recognized in the HC₅N(36–35) spectrum. We set the velocity scale of the N₂H⁺ (1–0) spectrum using the rest frequency of the isolated component of the N₂H⁺ ($v = 0$, $J = 1\text{--}0$, $F_1 = 0\text{--}1$, $F = 1\text{--}2$, 93.176265 GHz) hyperfine structure (HFS) multiplet as a reference. Hereafter, we call this isolated N₂H⁺ component “HFS 1” and the other components “HFS 2–7” (see Fig. 2). The peak velocity of HFS 1 (approximately -2.6 km s⁻¹) is consistent with the peak velocity of the HC₅N(36–35) emission. In addition, the N₂H⁺ HFS 1 line profile exhibits an emission wing at highly blueshifted velocities (up to -12 km s⁻¹, see Fig. 2). This blueshifted emission does not appear to be associated with the NGC 6334 filament itself since it is mainly detected outside the main filamentary structure (see

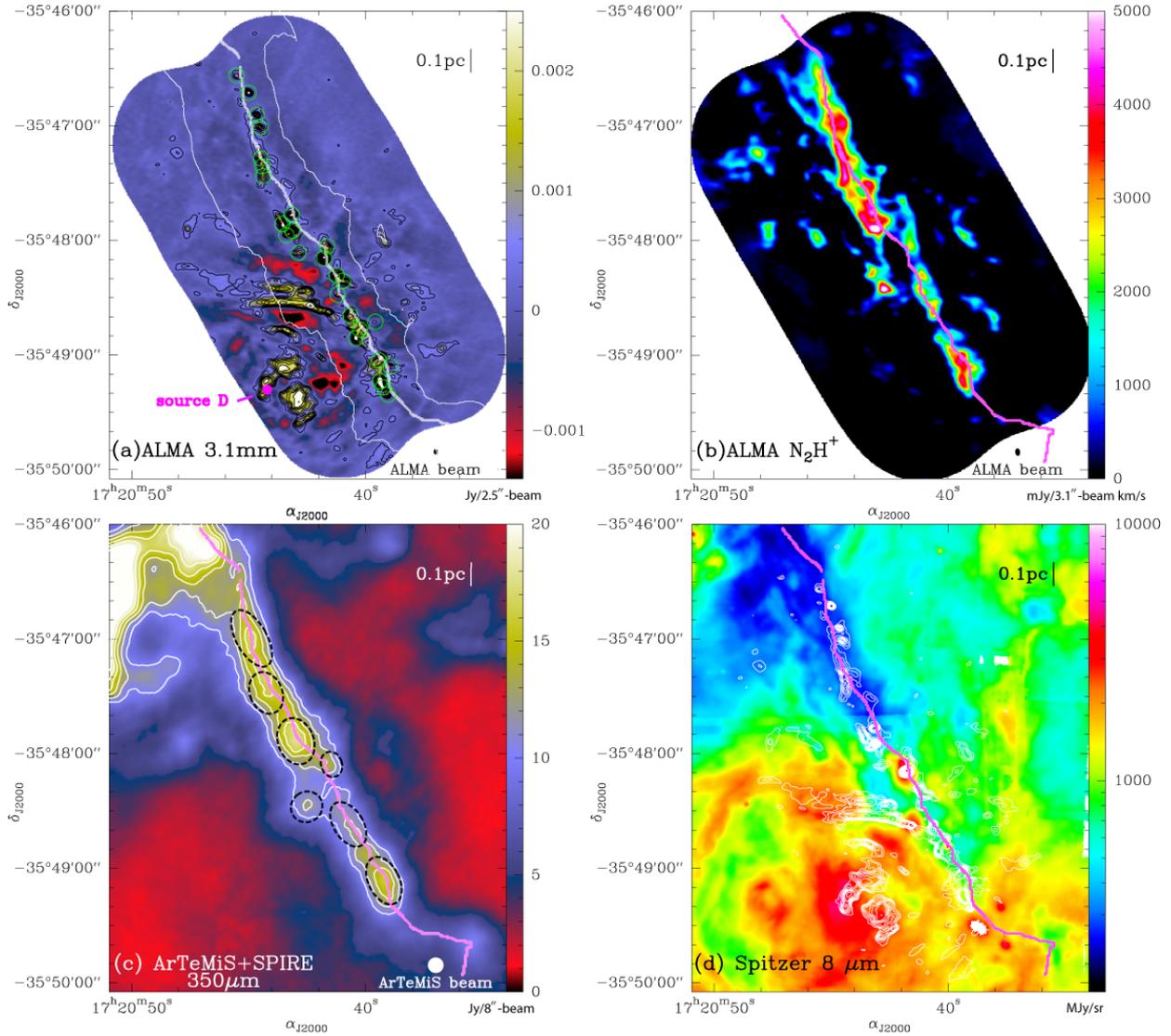


Fig. 1. Maps of the NGC 6334 main filament region in 3.1 mm continuum emission with ALMA 12 m (*panel a*), $N_2H^+(1-0)$ integrated intensity with ALMA 12 m+7 m (*panel b*), $350\ \mu\text{m}$ dust continuum emission with ArTéMiS (*panel c*, from André et al. 2016), and $8\ \mu\text{m}$ dust continuum emission with Spitzer (*panel d*). In *panel a*, the magenta filled circle indicates the position of a radio continuum compact HII region (source D in Rodríguez et al. 1982), and green open circles indicate the positions of the compact 3.1 mm continuum sources identified with getsources (see Sect. 3.3.1). In *panel a*, the solid white contour marks the footprint of the main filament, defined as the intersection of the area within $\pm 30''$ from the filament crest and the interior of the $5\ \text{Jy}\ \text{beam}^{-1}$ contour in the ArTéMiS $350\ \mu\text{m}$ map. In *panels a and d*, contours of 3.1 mm continuum emission are overlaid with levels of 2, 4, 6, 8, 10, 15, 20, 25, 30, 35, and 40σ , where $1\sigma = 0.2\ \text{mJy/ALMA-beam}$. In *panel b*, the N_2H^+ intensity was integrated over the velocity range from $-11.4\ \text{km s}^{-1}$ to $+16.8\ \text{km s}^{-1}$, including all hyperfine components. In *panel c*, contours of $350\ \mu\text{m}$ dust continuum emission are shown with levels of 10, 12, 14, 16, 18, and $20\ \text{Jy/8''-beam}$, and dashed ellipses indicate the positions of the ArTéMiS dense clumps identified with GAUSSCLUMPS (see Sect. 4.1). In all panels, the crest of the NGC 6334 main filament as traced with the DisPerSE algorithm in the ArTéMiS $350\ \mu\text{m}$ map (André et al. 2016) is shown as a solid curve.

Fig. E.1). At positions where the blueshifted emission between -12 to $-6\ \text{km s}^{-1}$ overlaps with the main velocity component at $\sim -2.6\ \text{km s}^{-1}$ in the plane of the sky, the HFS 2–7 counterparts of the blueshifted emission can be confused with HFS 1 emission from the $-2.6\ \text{km s}^{-1}$ component since the velocity offsets of the HFS 2–7 components range from 7.3 to $14.9\ \text{km s}^{-1}$ with respect to the velocity of HFS 1 (see Appendix B).

3.3. Extraction of compact continuum sources and “fiber-like” velocity structure from the data

To extract compact continuum sources from the ALMA/ACA 3.1 mm continuum map and fiber-like velocity-coherent features

from the ALMA/ACA $N_2H^+(1-0)$ data cube, we made use of the getsources, getfilaments, and getimages algorithms (Men’shchikov et al. 2012; Men’shchikov 2013, 2017, and in prep.).

3.3.1. Compact source extraction from the ALMA 3.1 mm continuum data with getsources

To identify compact sources in the ALMA 3.1 mm continuum map, we applied the getsources algorithm (e.g., Men’shchikov et al. 2012). getsources is a multi-scale source extraction algorithm primarily developed for the exploitation of multi-wavelength far-infrared and submillimeter continuum data resulting

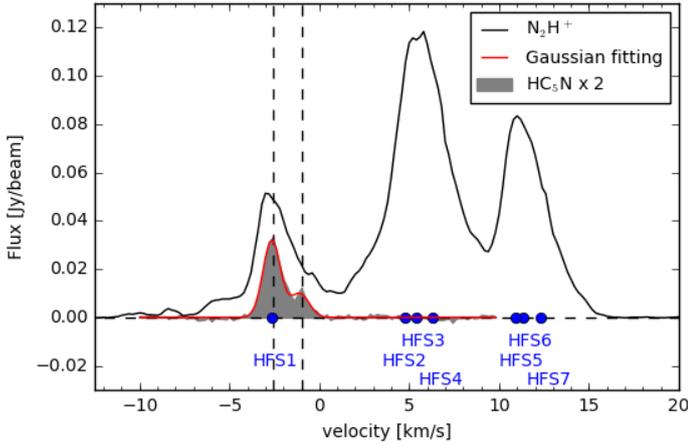


Fig. 2. $\text{N}_2\text{H}^+(1-0)$ (black) and $\text{HC}_5\text{N}(36-35)$ (gray) spectra averaged over pixels where the emission was detected above 5σ . The red curve shows a two-component Gaussian fit to the HC_5N line profile. The blue circles indicate the peak velocity positions of each $\text{N}_2\text{H}^+(1-0)$ HFS component expected from the $\text{HC}_5\text{N}(36-35)$ systemic velocity of -2.6 km s^{-1} . The vertical dashed lines mark the peaks of the two HC_5N velocity components.

from *Herschel* surveys of Galactic star-forming regions (see Könyves et al. 2015), but can also be used with single-band continuum data and spectral line data. Source extraction with this algorithm has only one free parameter, namely the maximum size of the structures to be extracted from the images. Here, we adopted a maximum size of $15''$ (or $\sim 0.12 \text{ pc}$), which is comparable to the transverse full width at half maximum (FWHM) size of the filament as measured with ArTéMiS (André et al. 2016).

After running *getsources* on the 12 m-only data² and applying post-extraction selection criteria (see Appendix C), we identified a total of 40 candidate compact 3.1 mm continuum sources. As shown in Fig. 1a, 28 of these 40 sources are located within the main filament defined as the intersection of the area within $\pm 30''$ (or $\pm 0.25 \text{ pc}$) from the filament crest and the area enclosed within the 5 Jy beam^{-1} contour in the ArTéMiS $350 \mu\text{m}$ map (see also Fig. 3 and Table 2). Two of these 28 sources (IDs 15, 18 in Table 2) are probably affected by contamination from free-free emission as mentioned in Sect. 3.1. The positions and sizes of all of these 40 compact continuum sources are summarized in Table 2, along with their basic properties.

To investigate whether the compact continuum sources identified above are associated with N_2H^+ emission, we also applied *getsources* to each N_2H^+ velocity channel map (see Appendix C for details). We found that 23 of the 40 compact continuum sources are associated with N_2H^+ emission in at least two consecutive channels. In addition, five of the 40 compact sources (IDs 3, 7, 13, 33, 35 in Table 2) are associated with $\text{N}_2\text{H}^+(1-0)$ line emission in only one channel, but these sources did not pass our post-extraction selection criteria (see Appendix C). Although contamination of the 3.1 mm continuum data by free-free emission is an issue, compact sources robustly detected in both 3.1 mm continuum and N_2H^+ emission are unlikely to be affected.

We estimated the mass (M_{tot}) of each compact continuum source, under the assumption that all of the 3.1 mm continuum emission arises from dust and that the emission is optically thin.

² We used the 12 m-only instead of the 12 m+7 m image as the sensitivity of the data is better for the detection of compact ($< 0.1 \text{ pc}$) sources (cf. Table 2).

The mass was obtained from the integrated 3.1 mm flux density $S_{3.1 \text{ mm}}^{\text{tot}}$ derived from two-dimensional Gaussian fitting in the 12 m+7 m continuum image³ using the formula:

$$M_{\text{tot}} = \frac{S_{3.1 \text{ mm}}^{\text{tot}} d^2}{\kappa_{3.1 \text{ mm}} B_{3.1 \text{ mm}}(T_d)}, \quad (1)$$

where d , $\kappa_{3.1 \text{ mm}}$, and $B_{3.1 \text{ mm}}(T_d)$ are the distance to the target, the dust opacity (per unit mass of gas + dust), and the Planck function at dust temperature T_d , respectively. We adopted the same dust opacity law as the *Herschel* Gould Belt survey team, namely $\kappa_\lambda = 0.1 (\lambda/300 \mu\text{m})^{-\beta} \text{ cm}^2 \text{ g}^{-1}$ (cf. Hildebrand 1983; Roy et al. 2014) with $\beta = 2$ and here $\lambda = 3.1 \text{ mm}$. If $\beta = 1.5$ instead of $\beta = 2$, the core masses would be a factor of ~ 3 lower than the values listed in Table 2. For most sources, we adopted $T_{\text{dust}} = 20 \text{ K}$ which corresponds to the median dust temperature derived from *Herschel* data along the crest of the filament (André et al. 2016; Tigé et al. 2017). For source ID 1, we adopted a higher temperature value ($T_{\text{dust}} = 50 \text{ K}$) as this object coincides with a bright *Spitzer* $8 \mu\text{m}$ source and is most likely an internally-heated, relatively massive protostellar core. The 5σ mass sensitivity of the ALMA 12 m continuum data for compact sources is $\sim 2.0 M_\odot$, corresponding to $S_{3.1 \text{ mm}}^{\text{tot}} = 0.7 \text{ mJy}$, assuming $\beta = 2$ and $T_{\text{dust}} = 20 \text{ K}$. The average gas density ($\equiv \bar{n}_{\text{H}_2}$) of each source was then derived assuming spheroidal cores as follows:

$$\bar{n}_{\text{H}_2} = \frac{M_{\text{tot}}}{\frac{4}{3}\pi [\sqrt{R_{\text{major}} R_{\text{minor}}}]^3}, \quad (2)$$

where R_{major} and R_{minor} are the deconvolved FWHM sizes along the major and minor axis of the source, respectively (cf. Könyves et al. 2015). The median mass of the 26 compact continuum sources embedded in the main filament⁴ is $9.6^{+3.0}_{-1.9} M_\odot$ (lower quartile: $7.7 M_\odot$, upper quartile: $12.6 M_\odot$) and their median volume-averaged density $1.6 \times 10^7 \text{ cm}^{-3}$ (lower quartile: $1.0 \times 10^7 \text{ cm}^{-3}$, upper quartile: $2.2 \times 10^7 \text{ cm}^{-3}$). The 26 continuum sources embedded in the main filament are compact (with an estimated typical outer radius $\sim 5000 \text{ au}$) and have a very high volume-averaged density ($\sim 10^7 \text{ cm}^{-3}$), suggesting that they are on their way to form stars. Moreover, their spatial distribution closely follows that of the ArTéMiS $350 \mu\text{m}$ emission clumps (see Fig. 3 and Sect. 4.1). In the following, we therefore regard these 26 compact 3.1 mm continuum sources as dense cores.

3.3.2. Extraction of fiber-like structures from the N_2H^+ data cube with *getfilaments*

Removal of unrelated velocity components. In the NGC 6334 region, we found evidence of the presence of several velocity components along the line of sight (see Sect. 3.2.3). To avoid contamination of the $\text{N}_2\text{H}^+(1-0)$ HFS 1 line emission from the main velocity component at $V_{\text{sys}} = -2.6 \text{ km s}^{-1}$ by the HFS 2–7 emission from other velocity components, we

³ We used the 12 m+7 m data to estimate the integrated flux densities and masses of the sources detected in the 12 m-only continuum map, in order to avoid missing-flux problems due to interferometric filtering of large scales as much as possible. For four weak unresolved sources, undetected above the 3σ level in the 12 m+7 m map, the 12 m-only data were used instead (cf. Table 2).

⁴ The median mass of all 40 compact continuum sources is $9.4^{+3.4}_{-1.9} M_\odot$ (lower quartile: $7.5 M_\odot$, upper quartile: $12.8 M_\odot$) and the median volume-averaged density $1.4 \times 10^7 \text{ cm}^{-3}$ (lower quartile: $0.8 \times 10^7 \text{ cm}^{-3}$, upper quartile: $2.0 \times 10^7 \text{ cm}^{-3}$).

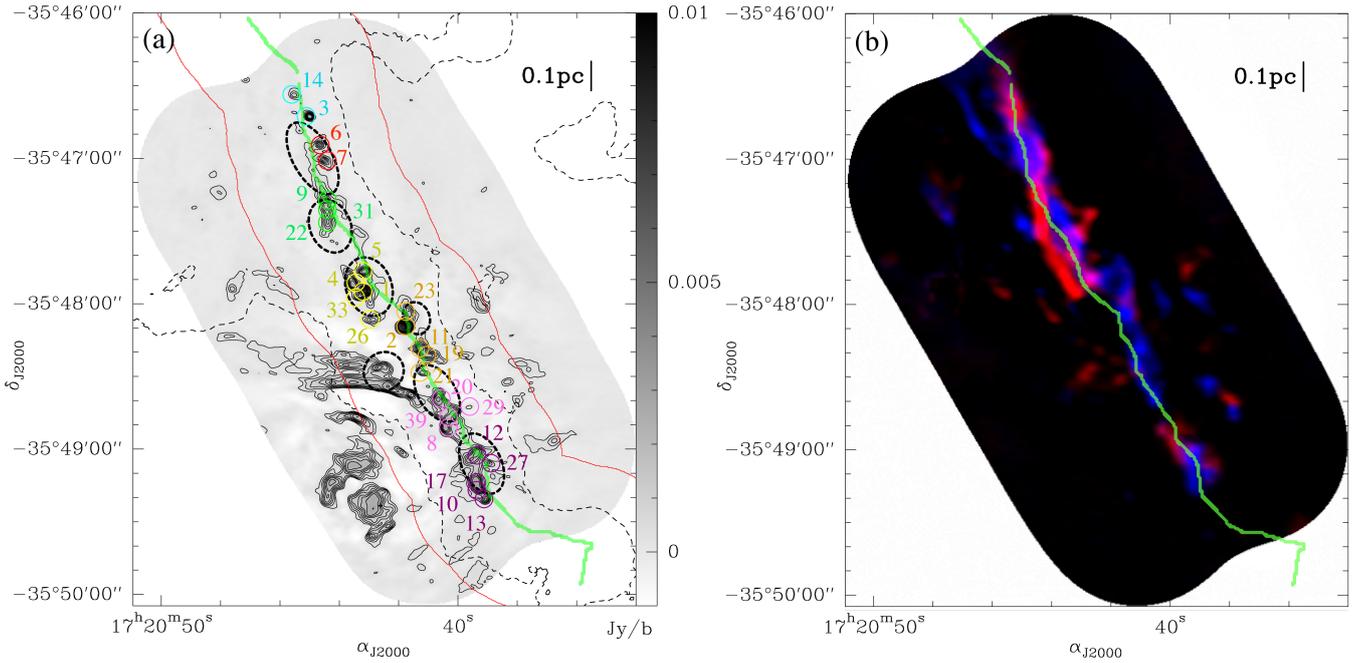


Fig. 3. *Panel a:* spatial distribution of the 26 ALMA compact cores (magenta, pink, orange, yellow, green, red, green, and light blue circles) and 7 ArTéMiS clumps (larger black dashed ellipses – see Sect. 4.1) embedded in the NGC 6334 main filament, overlaid on the ALMA 3.1 mm continuum image (grayscale and same solid contours as in Fig. 1). The black dashed contour shows an ArTéMiS 350 μm continuum emission level of 5 Jy/8''-beam. The green curve marks the crest of the dusty filament traced with DisPerSE in the ArTéMiS 350 μm continuum map (André et al. 2016, see Fig. 1c). The red curves mark the area within $\pm 30''$ (or ± 0.25 pc) from this filament crest. *Panel b:* spatial distribution of the $\text{N}_2\text{H}^+(1-0)$ velocity-coherent fiber-like sub-structures F-1 (blue, integrated over the velocity range -3 to -2.4 km s^{-1}) and F-2 (red, integrated over the velocity range -2.0 to -1.2 km s^{-1}) in the NGC 6334 filament. In *panel a*, the seven groups of ALMA 3.1 mm compact sources found by the nearest neighbor separation algorithm (see Sect. 4.1 and Fig. 8) are indicated by small circles of different colors. The green curve is the same as in *panel a*.

fit the N_2H^+ emission observed at each pixel with a multiple-velocity-component HFS model and subtracted the corresponding HFS 2–7 emission from the data cube. As shown in Fig. B.1, if the velocity width of each component is ~ 1.0 km s^{-1} , then the N_2H^+ HFS 1 emission of the main component can be contaminated by the HFS 2–7 emission of a more blue-shifted velocity component when the velocity difference between the two components along the line of sight is ~ 5.0 – 7.0 km s^{-1} . Figure B.2 shows four examples of $\text{N}_2\text{H}^+(1-0)$ spectra obtained after subtraction of contaminating components. Hereafter, we use the subtracted data cube in our scientific analysis.

Extraction of velocity-coherent fiber-like features with getfilaments. We extracted filamentary structures from the ALMA N_2H^+ data cube after subtraction of compact N_2H^+ emission using the *getfilaments* algorithm (e.g., Men'shchikov 2013). Each N_2H^+ velocity channel map was decomposed into a set of spatially-filtered single-scale images from small ($3''$) to large ($16''$) scales. At each scale, the algorithm separated filamentary structures and compact sources, and extracted significant filamentary structures in the source-subtracted image component. All extracted filamentary structures were merged into a single image per velocity channel map. We selected only filamentary structures whose $>90\%$ pixels along the crest were detected above the $>3\sigma$ level in each channel. As the velocity resolution of our ALMA N_2H^+ data (0.2 km s^{-1} , that is, less than the isothermal sound speed in the cloud) should be sufficient to resolve the velocity width of the filament, we also imposed that a selected filamentary structure should be detected in at least two consecutive velocity channels. By comparing velocity-channel maps, we therefore associated

filamentary structures with matching spatial distributions among channels (see Appendix D). Our procedure is similar to that used by Hacar et al. (2013) to trace velocity-coherent structures in data cubes, in the sense that both methods work in position–position–velocity (PPV) space⁵. In this way, we identified five distinct fiber-like structures, labeled F-1 to F-5, whose crests are displayed in Figs. 4 and 5. Examples of individual N_2H^+ (HFS 1) spectra across the two main fiber-like structures F-1 and F-2 are shown in Fig. 6. The spatial distribution of the structure labeled F-2 partly overlaps with the distribution of F-1 at -2.4 km s^{-1} . But the northern portion of F-2 lies slightly to the west of F-1, while the southern part of F-2 lies to the east of F-1 (see, e.g., Fig. 4b). Moreover, at positions where F-1 and F-2 overlap in the plane of the sky, the N_2H^+ spectra clearly exhibit distinct velocity components (see, e.g., positions (c) and (d) in Fig. 6). Therefore, F-1 and F-2 seem to be distinct velocity-coherent features. The five fiber-like structures are likely associated with the dust continuum filament seen in the ArTéMiS 350 μm map. Similar sub-filamentary structures have been found in the low-mass star-forming filament B211/B213 in Taurus (Hacar et al. 2013) and have been called fibers in the literature (cf. Tafalla & Hacar 2015). At this point, we refer to the five features F-1–F-5 as fiber-like structures. (In Sect 4.3 below, we argue that only two of them, F-1 and F-2, may be genuine fibers, that is velocity-coherent sub-structures of the main filament itself.)

⁵ As only a small (≥ 5) number of velocity-coherent features appeared to be present in our N_2H^+ data cube (see Table 3) and we had no experience with the FIVE method developed by Hacar et al. (2013), we preferred to follow a procedure based on channel-by-channel *getfilaments* extractions and careful visual inspection of the data cube (see Appendix D).

Table 2. Results of compact source extractions on the ALMA 12 m-only 3.1 mm continuum map.

ID ⁽¹⁾	RA	Dec	$R_{\text{major}} \times R_{\text{minor}}$	PA	$S_{3.1\text{mm}}^{\text{peak}}$	S/N	$S_{3.1\text{mm}}^{\text{tot}}$	M_{tot}	\bar{n}_{H_2}	Velocity channel at which source is detected in N_2H^+	Association with N_2H^+ features
	[hms]	[dms]	[$'' \times ''$] ^{(2),(3)}	[deg] ⁽²⁾	[mJy/beam] ⁽²⁾	[σ]	[mJy] ⁽²⁾	[M_{\odot}] ⁽⁴⁾	[$\times 10^6 \text{ cm}^{-3}$] ⁽⁵⁾	[km s^{-1}]	
1	17:20:43.13	-35:47:54.9	2.9×2.3 (0.02 pc \times 0.02 pc)	1	5.1 ± 0.2	36.4	11.4	31.5	11.1	[-2.2, -1.8]	F-2
2	17:20:41.78	-35:48:9.2	2.8×1.7 (0.02 pc \times 0.01 pc)	10	6.9 ± 0.2	49.2	13.5	13.8	8.1		No N_2H^+
3	17:20:45.00	-35:46:42.7	1.7×0.8 (0.01 pc \times 0.01 pc)	-33	2.9 ± 0.2	21.0	4.1	11.2	42.5	-1.2 ⁽⁶⁾	
4	17:20:43.56	-35:47:51.4	2.2×1.1 (0.02 pc \times 0.01 pc)	36	3.0 ± 0.2	21.4	4.7	12.8	22.6	[-1.4, -1.2]	F-2
5	17:20:43.16	-35:47:46.6	1.2×0.8 (0.01 pc \times 0.01 pc)	-8	2.5 ± 0.2	18.2	3.2	8.9	54.3	[-2.6, -2.4], [-2.0, -1.8]	F-1
6	17:20:44.66	-35:46:54.3	1.7×1.3 (0.01 pc \times 0.01 pc)	6	2.4 ± 0.2	17.0	3.5	9.6	19.1	[-3.0, -2.8], [-1.8], [-1.4, -1.2]	F-2
7	17:20:44.51	-35:47:0.8	2.0×1.1 (0.02 pc \times 0.01 pc)	-32	2.4 ± 0.2	17.3	3.7	10.1	19.1	-3.8, -2.8 ⁽⁶⁾	F-1(marginal)/F-3(marginal)
8	17:20:40.42	-35:48:51.1	1.9×1.4 (0.02 pc \times 0.01 pc)	-22	3.0 ± 0.2	21.5	4.7	12.9	18.5	[-2.6, -1.4]	
9	17:20:44.48	-35:47:17.1	1.9×1.5 (0.02 pc \times 0.01 pc)	40	2.7 ± 0.2	19.4	4.3	11.9	14.7	[-2.8, -2.6], [-1.8, -1.2]	F-2
10	17:20:39.27	-35:49:16.9	2.1×0.9 (0.02 pc \times 0.01 pc)	35	3.3 ± 0.2	23.8	5.0	13.7	35.1	[-3.8, -3.0]	F-1
11	17:20:41.25	-35:48:18.3	1.5×1.0 (0.01 pc \times 0.01 pc)	-39	2.3 ± 0.2	16.6	3.2	8.7	28.4	[-3.6, -2.6]	
12	17:20:39.44	-35:49:2.8	2.4×1.4 (0.02 pc \times 0.01 pc)	-42	3.5 ± 0.2	24.8	6.0	16.6	15.8	[-3.8, -2.4]	
13	17:20:39.05	-35:49:20.7	2.4×0.9 (0.02 pc \times 0.01 pc)	19	1.8 ± 0.2	13.1	2.9	8.1	15.3	-3.8 ⁽⁶⁾	
14	17:20:45.52	-35:46:33.6	1.7×1.5 (0.01 pc \times 0.01 pc)	-28	1.7 ± 0.2	12.1	2.6	7.1	10.7	[-3.2, -2.8], [-2.2, -2.0], -1.2	F-2
15 ⁽⁷⁾	17:20:41.58	-35:48:37.5	1.5×1.0 (0.01 pc \times 0.01 pc)	-23	2.6 ± 0.2	18.3	2.1	7.0	24.4		No N_2H^+ /free-free contamination
16	17:20:43.38	-35:48:47.0	2.3×0.9 (0.02 pc \times 0.01 pc)	-32	4.5 ± 0.2	32.3	7.1	19.6	37.8		No N_2H^+ /Out of main filament
17	17:20:39.39	-35:49:13.7	2.1×1.7 (0.02 pc \times 0.01 pc)	6	3.9 ± 0.2	28.1	6.6	18.2	17.7	[-3.8, -2.8]	
18	17:20:42.35	-35:48:34.2	2.0×1.7 (0.02 pc \times 0.01 pc)	-39	4.8 ± 0.2	34.5	8.1	22.3	21.9	-3.4, [-1.6, -1.4]	free-free contamination
19 ⁽⁷⁾	17:20:40.97	-35:48:21.2	1.5×1.0 (0.01 pc \times 0.01 pc)	-19	2.0 ± 0.2	13.9	1.8	5.4	18.6		No N_2H^+
20	17:20:40.66	-35:48:39.4	1.3×1.0 (0.01 pc \times 0.01 pc)	44	2.1 ± 0.2	15.0	2.8	7.6	31.0	[-3.4, -2.6]	F-1
21 ^{(7),(2)}	17:20:41.20	-35:48:28.5	1.4×0.9 (0.01 pc \times 0.01 pc)	-33	1.1 ± 0.1	8.1	1.4	3.1	13.4	[-2.4, -1.2]	F-1
22	17:20:44.39	-35:47:26.4	2.1×1.6 (0.02 pc \times 0.01 pc)	-2	2.2 ± 0.2	16.0	3.7	10.2	10.7	[-3.8, -3.2], [-1.8, -1.2]	F-2
23	17:20:41.65	-35:48:3.1	2.0×1.3 (0.02 pc \times 0.01 pc)	44	1.7 ± 0.2	12.2	3.2	8.9	7.0	[-3.4, -2.4]	F-1
24	17:20:47.61	-35:47:53.7	1.6×1.5 (0.01 pc \times 0.01 pc)	31	1.2 ± 0.2	8.3	1.7	4.8	8.1	[-3.4, -3.2], -2.0, -1.2	Out of main filament
25	17:20:42.87	-35:49:36.0	1.9×0.8 (0.02 pc \times 0.01 pc)	34	1.9 ± 0.2	13.7	2.8	7.6	23.9		No N_2H^+ /Out of main filament
26	17:20:42.87	-35:48:6.6	2.0×0.2 (0.02 pc \times 0.00 pc)	28	1.5 ± 0.2	10.5	2.0	5.6	141.9	[-1.4, -1.2]	F-2
27	17:20:38.94	-35:49:5.9	2.1×1.0 (0.02 pc \times 0.01 pc)	-19	1.9 ± 0.2	13.6	2.9	7.9	17.0	[-2.2, -1.6]	
28	17:20:41.88	-35:49:1.6	3.4×0.9 (0.03 pc \times 0.01 pc)	-41	2.2 ± 0.2	16.0	4.3	11.9	14.4		No N_2H^+ /Out of main filament
29 ^{(7),(2)}	17:20:39.59	-35:48:42.6	1.4×0.9 (0.01 pc \times 0.01 pc)	15	0.6 ± 0.1	4.5	0.3	1.7	7.4		No N_2H^+
30	17:20:42.35	-35:49:31.8	1.7×1.7 (0.01 pc \times 0.01 pc)	1	2.7 ± 0.2	19.1	4.2	11.6	14.6		No N_2H^+ /Out of main filament
31	17:20:44.42	-35:47:21.2	2.2×1.5 (0.02 pc \times 0.01 pc)	22	2.1 ± 0.2	15.2	3.5	9.7	10.7	-3.8, [-2.0, -1.8]	F-2
32 ^{(7),(2)}	17:20:44.26	-35:49:14.1	1.4×0.9 (0.01 pc \times 0.01 pc)	39	3.3 ± 0.1	23.4	1.1	9.0	38.5		No N_2H^+ /Out of main filament
33 ⁽⁷⁾	17:20:43.39	-35:47:57.3	1.5×1.0 (0.01 pc \times 0.01 pc)	42	2.3 ± 0.2	16.4	0.7	6.3	21.9	-1.8 ⁽⁶⁾	F-2(marginal)
34	17:20:39.27	-35:47:59.3	2.3×1.3 (0.02 pc \times 0.01 pc)	34	2.0 ± 0.2	14.4	3.4	9.2	10.7	[-3.8, -3.4], [-2.8, -2.6]	Out of main filament
35 ^{(7),(2)}	17:20:38.34	-35:48:32.4	1.4×0.9 (0.01 pc \times 0.01 pc)	15	0.9 ± 0.1	6.7	1.2	2.6	11.0	-2.2 ⁽⁶⁾	Out of main filament
36	17:20:42.92	-35:49:24.7	2.1×1.9 (0.02 pc \times 0.02 pc)	-19	3.8 ± 0.2	27.0	6.7	18.5	14.1		No N_2H^+ /Out of main filament
37	17:20:36.73	-35:48:56.6	1.9×1.4 (0.02 pc \times 0.01 pc)	38	1.4 ± 0.2	9.9	2.2	6.0	8.4		No N_2H^+ /Out of main filament
38	17:20:43.44	-35:49:6.4	2.2×1.8 (0.02 pc \times 0.01 pc)	-4	1.8 ± 0.2	12.9	3.1	8.7	7.2		No N_2H^+ /Out of main filament
39	17:20:40.25	-35:48:46.7	2.5×1.5 (0.02 pc \times 0.01 pc)	-36	2.3 ± 0.2	16.4	4.0	11.1	10.0	[-2.6, -2.2]	
40	17:20:42.82	-35:49:20.4	2.0×1.9 (0.02 pc \times 0.02 pc)	0	3.4 ± 0.2	24.4	5.9	16.3	13.8		No N_2H^+ /Out of main filament

Notes. ⁽¹⁾ID numbers of 3.1 mm compact sources associated with N_2H^+ emission in two or more contiguous channels are labeled in bold. ⁽²⁾Results of the 2D Gaussian fitting on the 12 m+7 m map. For IDs 21, 29, 32, 35, the 12 m map was used since the 3.1 mm continuum emission is not detected with a $>3\sigma$ on the 12m+7m map. ⁽³⁾Deconvolved FWHM source diameters along the major and minor axes. ⁽⁴⁾We adopted $T_{\text{dust}} = 20$ K for all sources, except for source ID 2 for which we assumed $T_{\text{dust}} = 50$ K (see Sect. 3.3.1). ⁽⁵⁾Volume-averaged gas density derived assuming spheroidal cores as in Eq. (2). ⁽⁶⁾Marginal N_2H^+ detection since it is detected in only one channel for this source. ⁽⁷⁾Deconvolved source size are not obtained since the structure is not resolved. The peak flux was used to estimate the mass.

Their typical length is 0.5 ± 0.4 pc (see Table 3, which also gives their velocity ranges).

A portion of the dusty filament seen in the ArTéMiS 350 μm continuum map was not traced by `getfilaments` in the N_2H^+ data, especially in the southern part of the field in Fig. 1 although the filament can be recognized by eye in the N_2H^+ (1-0) integrated intensity map (Fig. 1b). The reason why this southern portion was not traced in the N_2H^+ velocity channel maps with `getfilaments` is that the N_2H^+ emission becomes very clumpy in this area and was subtracted out as a collection of point-like sources by `getsources` before identification of filamentary structures.

F-1 is the longest and F-2 the second longest of the five extracted velocity features. Both of these fiber-like structures are roughly parallel to the dust continuum filament seen in the ArTéMiS 350 μm map. In contrast, the crest of F-5 is roughly perpendicular to the dust filament. In the northern part of the

N_2H^+ map (north of $-35^\circ 47' 10''$ in Figs. 4 and 7a), the spatial distributions of F-1, F-2, F-4 partly overlap in the plane of sky. The three features F-1, F-2, F-4 are nevertheless separated from each other in velocity space (see Table 3). The crests of F-1, F-2, F-4 are also slightly shifted from one another by typically one ALMA beam ($\sim 3''$, see Fig. 7a). In the northern part, they are distributed in the sequence F-1, F-2, F-4 from east to west (see Fig. 7a). In contrast, the projected LSR velocities are in the order F-4 ($\sim -3.3 \text{ km s}^{-1}$), F-1 ($\sim -2.8 \text{ km s}^{-1}$), F-2 ($\sim -1.7 \text{ km s}^{-1}$). Assuming F-1, F-2, F-4 are part of the same filament, this different ordering between the spatial and velocity distributions cannot be explained by a simple transverse velocity gradient across the parent filament due to either accretion onto and/or rotation of the filament. In the southern part of the field (south of $-35^\circ 47' 10''$ in Figs. 4 and 7b), F-4 is not detected and F-1, F-2 are distributed in the order F-2, F-1 from east to west (see Fig. 7b). These opposite spatial configurations for F-1,

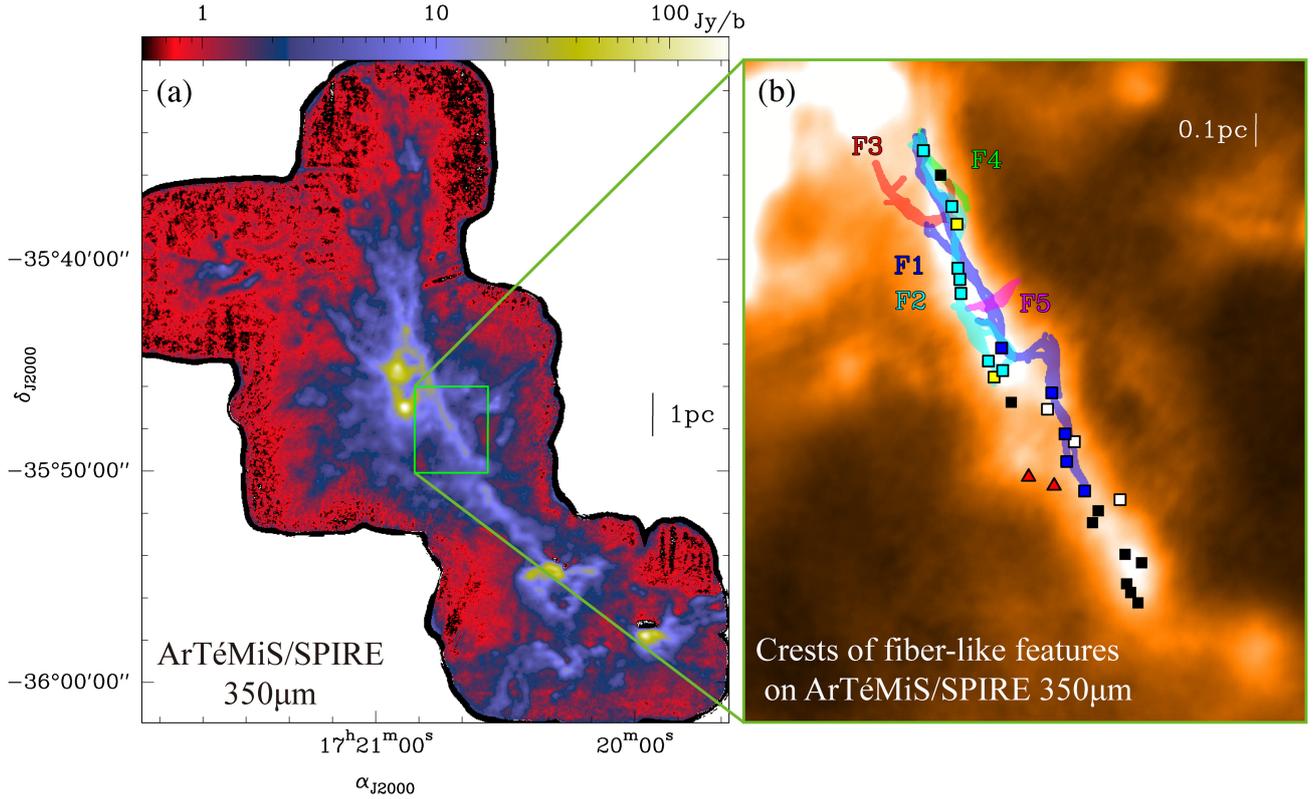


Fig. 4. *Panel a:* ArTéMiS 350 μm dust continuum map of the NGC 6334 region (André et al. 2016) and *panel b:* blow-up of the field imaged here with ALMA, with the skeleton of the velocity-coherent fiber-like structures identified in the N_2H^+ data cube overlaid as colored curves (see Sect. 3.3.2 and Table 3). In *panel b*, dark blue squares mark 3.1 mm compact sources associated with N_2H^+ emission and embedded in F-1, light blue squares 3.1 mm sources associated with N_2H^+ emission and embedded in F-2, yellow squares 3.1 mm sources embedded in F-1, F-2, or F-3 and marginally detected in only one N_2H^+ channel, black squares 3.1 mm sources associated with N_2H^+ emission in the main filament but not embedded in any N_2H^+ velocity-coherent structure, and white squares 3.1 mm sources in the main filament not associated with any N_2H^+ emission. The red triangles indicate 3.1 mm compact sources possibly contaminated by free-free emission.

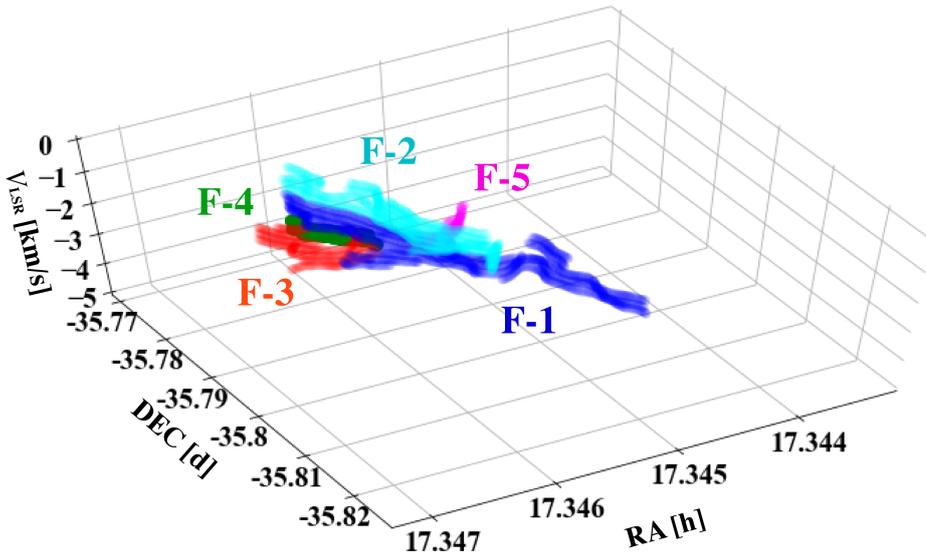


Fig. 5. $\text{N}_2\text{H}^+(1-0)$ position–position–velocity plot of the NGC 6334 filament system. Colored curves indicate the crests of the N_2H^+ velocity-coherent structures identified with *getfilaments*, after applying the post-extraction selection criteria described in Sect. 3.3.2, and using the same color coding as in Fig. 4b.

F-2 between the northern and the southern part of the dust filament are suggestive of an intertwined, double helix-like pattern (DNA-like) (see Fig. 3b).

We also extracted filamentary structures from the ALMA HC_3N , HC_5N , CH_3CCH , and H_2CS integrated intensity maps using *getfilaments*. We detected one filamentary structure in HC_3N , one filamentary structure in HC_5N , one filamentary

structure in CH_3CCH , and two filamentary structures in H_2CS . These filamentary structures roughly correspond to the F-1 and/or F-2 “fiber-like” features identified in N_2H^+ . Due to the lower effective sensitivity of the $\text{H}_2\text{CS}(3_{1,2}-2_{1,1})$ data, the NGC 6334 filament is broken up into two segments in H_2CS . The H_2CS filamentary structures coincide with the northern and southern parts of the filamentary structures traced in other lines.

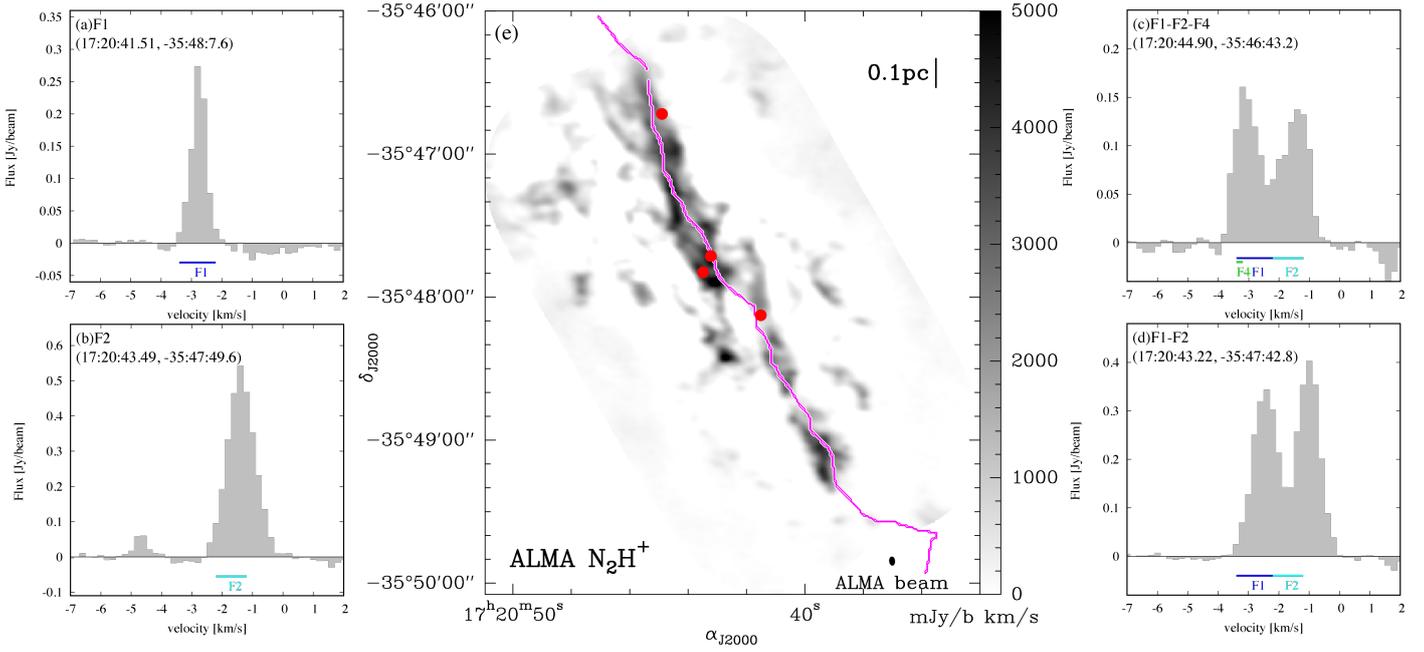


Fig. 6. Four examples of spectra showing the isolated component of the $\text{N}_2\text{H}^+(1-0)$ multiplet (HFS 1) in the original data cube (i.e. before subtracting unrelated velocity components). The positions of these spectra are indicated at the top-left corner of each panel. The central panel is the $\text{N}_2\text{H}^+(1-0)$ intensity map (integrated from -11.4 km s^{-1} to $+16.8 \text{ km s}^{-1}$ – cf. Fig. 1b). The red filled circles in the central panel also mark the positions of the four spectra. Colored horizontal segments indicate the velocity ranges of the various fiber-like structures (see Table 3).

Table 3. N_2H^+ “fiber-like” structures identified with *getfilaments*.

ID	Length	Velocity range
F-1	1.21 pc	$[-3.4, -2.2 \text{ km s}^{-1}]$
F-2	0.72 pc	$[-2.2, -1.2 \text{ km s}^{-1}]$
F-3	0.33 pc	$[-3.6, -2.4 \text{ km s}^{-1}]$
F-4	0.26 pc	$[-3.4, -3.2 \text{ km s}^{-1}]$
F-5	0.17 pc	$[-2.2, -1.6 \text{ km s}^{-1}]$

Owing to the lower line intensities and/or lower spectral resolution of our data in HC_3N , HC_5N , CH_3CCH , and H_2CS , we could not carry out as detailed a PPV analysis in these lines as in N_2H^+ .

4. Discussion

Our ALMA study has revealed, for the first time, the presence of density and velocity sub-structure in the NGC 6334 filament. We discuss the two types of sub-structure in turn in the following.

4.1. Bimodal fragmentation in the NGC 6334 filament

While it was not clear whether dense cores were embedded in the NGC 6334 filament based only on the ArTéMiS $350 \mu\text{m}$ data (André et al. 2016), our ALMA 3.1 mm continuum map with an angular resolution of $\sim 2.3''$ has allowed us to identify 26 candidate dense cores within the filament (cf. Sect. 3.3.1). These ALMA cores appear to be clustered in several groups closely associated with ArTéMiS clumps (see Fig. 3). Figure 8 shows a dendrogram tree obtained when applying a nearest neighbor separation (NNS) analysis to the population of ALMA dense cores (with the *scipy* function *cluster.hierarchy.linkage* and the *single* method). Here, we adopted two values of the NNS

grouping threshold, 0.1 and 0.15 pc, consistent with both the typical inner width of the filament and the typical FWHM size of the ArTéMiS clumps (see Table 4). Based on the NNS analysis with a grouping threshold of 0.1 pc, the 26 dense cores can be divided into seven groups of cores (groups 1–7). Using a grouping threshold of 0.15 pc instead of 0.1 pc, groups 2, 3 join group 1 and groups 4, 5, 6, 7 merge into the same group (see Fig. 8).

Remarkably, the seven groups of ALMA compact cores roughly correspond to dense clumps visible as closed contours in the ArTéMiS $350 \mu\text{m}$ map of the NGC 6334 filament at $8''$ resolution (see Figs. 1c and E.1). For the purpose of this paper, we used GAUSSCLUMPS (Stutzki & Guesten 1990) to characterize the properties of the ArTéMiS clumps. We applied GAUSSCLUMPS with a detection threshold of $1 \text{ Jy}/8''\text{-beam}$ corresponding to 5σ (where $1\sigma = 0.2 \text{ Jy}/8''\text{-beam}$). A total of seven clumps were identified in this way within the footprint of the main filament, defined as the intersection of the area within $\pm 30''$ (or $\pm 0.25 \text{ pc}$)⁶ from the filament crest and the interior of the $5 \text{ Jy}/8''\text{-beam}$ ArTéMiS contour (see solid white contour in Fig. 8), whose positions, sizes, and estimated masses are given in Table 4. These seven ArTéMiS clumps are also identified with the *getsources* (Men’shchikov et al. 2012; Men’shchikov 2013) and REINHLOD (Berry et al. 2007) source extraction algorithms, as described in Appendix F. They correspond to six groups of ALMA cores (groups 1, 2, 3, 5, 6, 7). More precisely, if we consider a group of ALMA dense cores to be associated with an ArTéMiS clump when all of the cores in the group lie within the closed contours of the clump (Table 4), then we find that each ArTéMiS clump consists of at least two ALMA cores (Fig. 3a).

⁶ The area within $\pm 0.25 \text{ pc}$ of the filament crest extends well into the power-law wing of the filament radial profile (see Fig. 3a of André et al. 2016) and includes the bulk of the ArTéMiS $350 \mu\text{m}$ emission from the filament (see Fig. 8).

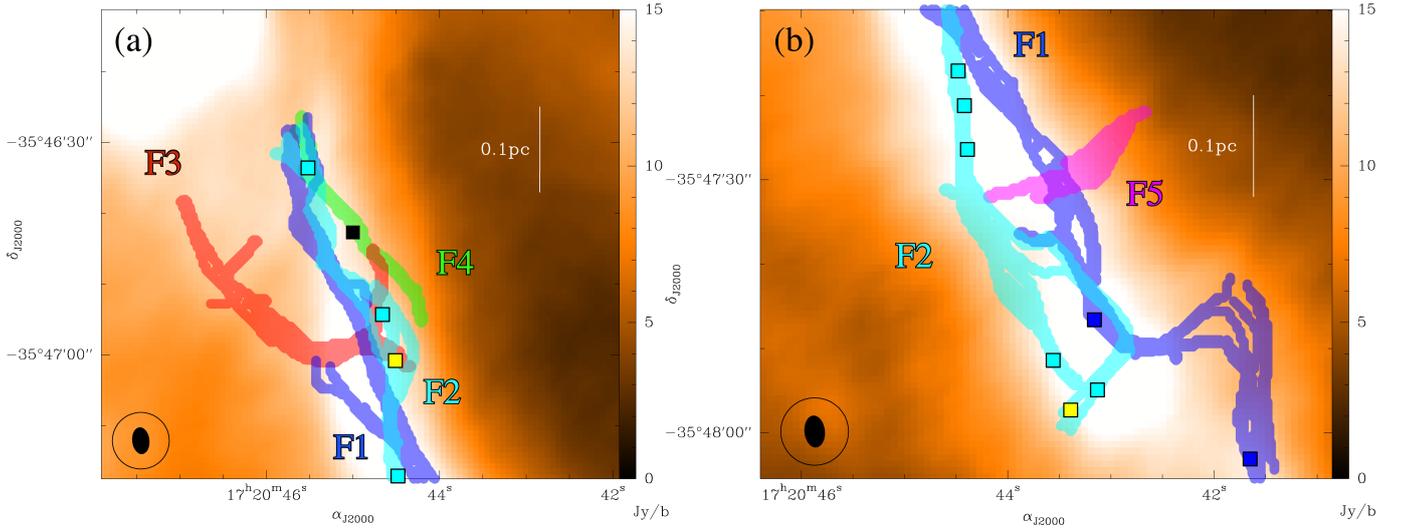


Fig. 7. Close-up views of the northern (a) and southern (b) parts of the dusty filament seen in the ArTéMiS $350\ \mu\text{m}$ map. Symbols and curves are the same as in Fig. 4b. The ALMA 3.1 mm continuum and ArTéMiS $350\ \mu\text{m}$ beams are shown as a filled ellipse and an open circle at the bottom left of each panel.

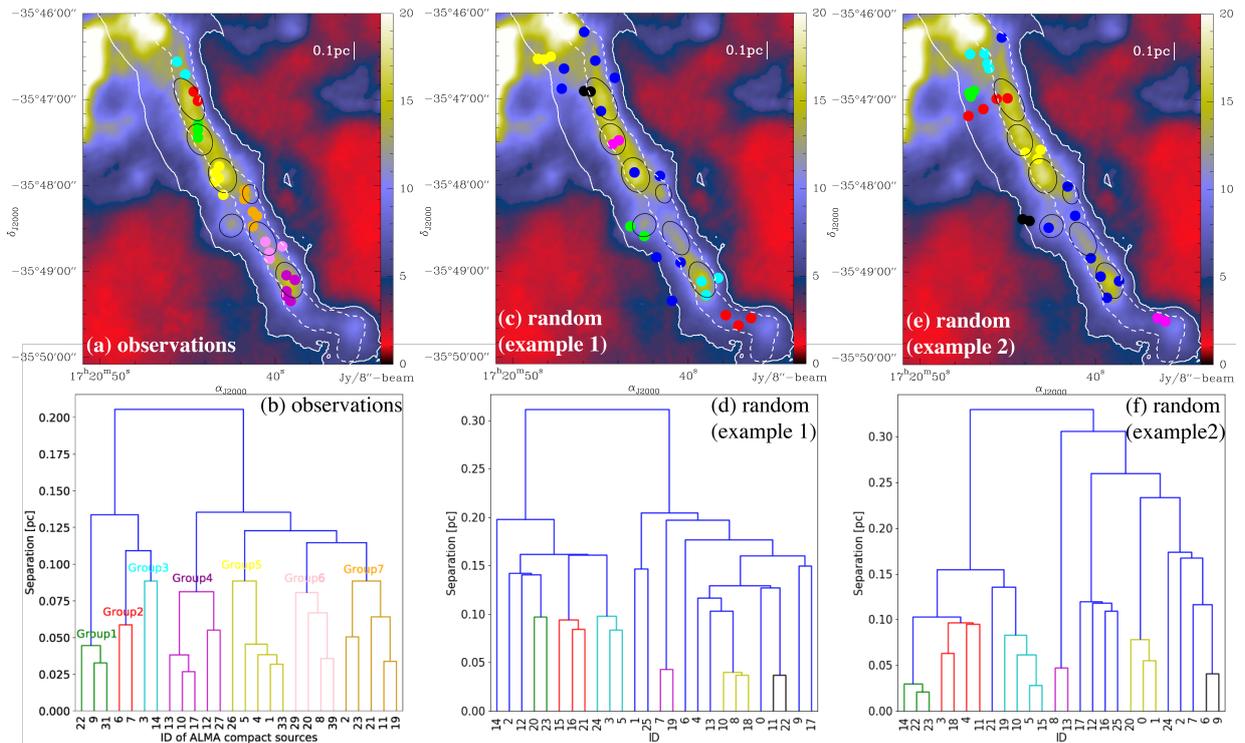


Fig. 8. Comparison of the spatial distributions and NNS analysis results obtained for the observed ALMA cores (*left column*) with those obtained for two examples of randomly-distributed source populations (*middle and right columns of panels*). *Top row*: distributions of observed cores (*panel a*) and randomly-placed sources (*panels c and e*) overlaid on the ArTéMiS $350\ \mu\text{m}$ continuum map (color scale). In each panel, the solid white contour marks the footprint of the main filament, defined as the intersection of the area within $\pm 30''$ from the filament crest and the interior of the $5\ \text{Jy}\ \text{beam}^{-1}$ contour in the ArTéMiS $350\ \mu\text{m}$ map, within which randomly-placed sources were distributed. The dashed white contour marks the area within $\pm 6''$ (or $0.05\ \text{pc}$) of the filament crest, defining the inner filament footprint also discussed in the text. *Bottom row*: NNS analysis results for the observed cores (*panel b*) and the randomly-placed sources (*panels d and f*). Colored circles mark core positions in *panels a, c, e*, with colors corresponding to the groups defined by the NNS dendrograms shown in *panels b, d, f*.

To test whether the observed clustering of ALMA cores and close association with ArTéMiS clumps may also be present in the case of randomly distributed cores, we inserted 26 sources at random positions within the footprint of the main filament (see solid white contour in Fig. 8) using the python module *random*

and then applied the same NNS analysis. A total of 100 realizations of such random source distributions were constructed. Two examples of resulting NNS dendrogram trees are displayed in Fig. 8, where they are compared to the NNS tree obtained for the real ALMA cores. It can be seen that randomly distributed

Table 4. Properties of ArTéMiS dense clumps.

ArTéMiS clump	RA [h:m:s]	Dec [d:m:s]	$R_{\text{major}} \times R_{\text{minor}}^{(a)}$	PA [deg]	$S_{350\mu\text{m}}^{\text{peak}}$ [Jy/beam]	$S_{350\mu\text{m}}^{\text{tot}}$ [Jy]	$M_{\text{clump}}^{(b)}$ [M_{\odot}]	\bar{n}_{H_2} [$\times 10^5 \text{ cm}^{-3}$]	$\sigma_{\text{N}_2\text{H}^+}$ [km s $^{-1}$]	IDs of associated ALMA cores
1	17:20:44.9	-35:46:60	31''9 \times 13''7 (0.26 pc \times 0.11 pc)	15	17	227	320	2.1	0.5	6, 7 (Group 2)
2	17:20:44.3	-35:47:28	21''2 \times 14''8 (0.17 pc \times 0.12 pc)	22	11	106	149	1.7	0.8	9, 19, 24 (Group 1)
3	17:20:43.0	-35:47:53	23''6 \times 16''9 (0.19 pc \times 0.14 pc)	24	18	216	303	2.3	1.0	1, 4, 5, 21, 25 (Group 5)
4	17:20:41.4	-35:48:06	10''9 \times 7''5 (0.09 pc \times 0.06 pc)	26	9	36	51	4.2	0.3	2, 11, 16, 18, 20 (Group 7)
5	17:20:42.5	-35:48:28	14''7 \times 13''7 (0.12 pc \times 0.11 pc)	-80	12	82	116	2.5	0.8	
6	17:20:40.7	-35:48:37	24''6 \times 13''0 (0.20 pc \times 0.11 pc)	13	11	117	164	1.8	0.4	8, 17, 23, 26 (Group 6)
7	17:20:39.2	-35:49:06	25''3 \times 14''0 (0.21 pc \times 0.12 pc)	19	16	177	249	2.3	0.5	10, 12, 13, 15, 22 (Group 4)

Notes. ^(a)Deconvolved FWHM source diameters along the major and minor axes. ^(b)Estimated using an equation similar to Eq. (1) at $\lambda = 350 \mu\text{m}$ assuming $T_{\text{dust}} = 20 \text{ K}$.

sources tend to be less clustered than the observed cores, that is, a higher number of randomly-placed sources are isolated (not grouped by the NNS analysis) compared to the observations. In contrast to the observed cores, there is also no clear association between the groups of randomly-placed sources and the ArTéMiS clumps. The clustering of the observed ALMA cores within the ArTéMiS clumps is highly significant, at the $>5\sigma$ level according to binomial statistics: 16 out of 26 cores ($>60\%$) lie within the FWHM ellipses of the clumps, while only 3 ± 2 ($\sim 13\%$) of 26 randomly-placed objects would be expected within the clumps. Furthermore, the mean separation between observed cores ($0.04 \pm 0.01 \text{ pc}$) is smaller than the mean separation between randomly-placed sources found over 100 realizations ($0.06 \pm 0.01 \text{ pc}$).

It is also apparent in the top row of Fig. 8 that the observed cores lie significantly closer to the filament crest than randomly-placed objects within the filament footprint (solid white contour). Quantitatively, the distribution of offsets between the observed cores and the filament crest has a median value of only 0.02 pc (lower quartile: 0.01 pc , upper quartile: 0.03 pc) while the median offset between randomly-placed sources and the crest is 0.07 pc (lower quartile: 0.03 pc , upper quartile: 0.13 pc). In other words, the spatial distribution of observed cores is almost one-dimensional (1D) along the filament crest. To further test whether significant 1D grouping of the cores exists along the crest, we repeated the same experiment with randomly-placed objects as described above but using a narrower footprint, defined as the area within $\pm 6''$ (or 0.05 pc) of the filament crest (dashed white contour in Fig. 8). The transverse width of this inner filament footprint is $\sim 0.1 \text{ pc}$, which is comparable to the dispersion of observed core positions about the crest. In this case, the mean separation between randomly-placed sources becomes identical to the mean separation between observed cores ($0.04 \pm 0.01 \text{ pc}$). The observed ALMA cores nevertheless remain more closely associated with the ArTéMiS clumps than randomly-distributed objects along the filament crest: only 8 ± 3 of 26 randomly-placed sources would be expected within the clumps, while 16 are observed, a difference which is significant at the $\sim 3\sigma$ level according to binomial statistics. We conclude that, even in 1D, there is marginal evidence of non-random core grouping along the filament axis.

The projected nearest-neighbor separation between the ArTéMiS clumps ranges from 0.2 to 0.3 pc , and the typical projected separation between ALMA dense cores embedded within a given ArTéMiS clump is $0.04 \pm 0.01 \text{ pc}$. These observational findings are interesting to compare with theoretical expectations. When the line mass of a cylindrical gas filament is close to the critical mass per unit length (thermal case) or the virial mass per

unit length (nonthermal case, more appropriate here – see end of Sect. 1), the filament is expected to fragment into overdensities with a typical separation of about four times the filament width according to self-similar solutions which describe the collapse of an isothermal filament under the effect of self-gravity (e.g. Inutsuka & Miyama 1992, 1997). With a typical inner width of $0.15 \pm 0.05 \text{ pc}$ from the ArTéMiS results (André et al. 2016), the NGC 6334 filament can therefore be expected to fragment with a characteristic separation of $\sim 0.6 \pm 0.2 \text{ pc}$. This separation scale is roughly consistent with the projected separation of 0.2 – 0.3 pc observed between ArTéMiS clumps, assuming a plausible inclination angle of $\sim 30^\circ$ between the filament axis and the line of sight.

On the other hand, the effective Jeans length $\lambda_{\text{J,eff}}$ or Bonnor-Ebert diameter $D_{\text{BE,eff}}$ within the filament and its clumps may be estimated as (cf. Bonnor 1956):

$$\lambda_{\text{J,eff}} \equiv D_{\text{BE,eff}} \approx 0.98 \sqrt{\frac{C_{\text{s,eff}}^2}{G \rho_{\text{clump}}}}, \quad (3)$$

where

$$C_{\text{s,eff}} = \sqrt{\sigma_{\text{N}_2\text{H}^+}^2 + kT_{\text{k}} \left(\frac{1}{m} - \frac{1}{m_{\text{N}_2\text{H}^+}} \right)}, \quad (4)$$

and G , ρ_{clump} , $\sigma_{\text{N}_2\text{H}^+}$, T_{k} , m , and $m_{\text{N}_2\text{H}^+}$ are the gravitational constant, the density of each ArTéMiS clump (see Table 4), the velocity dispersion measured in N_2H^+ , the gas kinetic temperature, the mean molecular mass, and the mass of the N_2H^+ molecule, respectively. Assuming a gas kinetic temperature of 20 K , which corresponds to the median dust temperature derived from *Herschel* data along the crest of the filament (André et al. 2016; Tigé et al. 2017) and using the velocity dispersion measured in N_2H^+ for each ArTéMiS clump ($\sigma_{\text{N}_2\text{H}^+} = \delta V_{\text{FWHM}} / \sqrt{8 \ln(2)}$, where δV_{FWHM} is the FWHM linewidth), the effective Jeans length in the clumps of the filament is estimated to range from ~ 0.04 to $\sim 0.3 \text{ pc}$ (median: 0.08 pc). This is roughly consistent with the typical projected separation between ALMA dense cores ($0.04 \pm 0.01 \text{ pc}$) assuming random projection effects within the ArTéMiS clumps.

These two characteristic separation scales are suggestive of two distinct fragmentation modes within the NGC 6334 filament: (i) a cylindrical fragmentation mode into clumps or groups of cores with a separation of ~ 4 times the filament width, and (ii) a spherical Jeans-like fragmentation mode into compact cores with a separation on the order of the effective Jeans length. Similar bimodal fragmentation patterns were first reported by

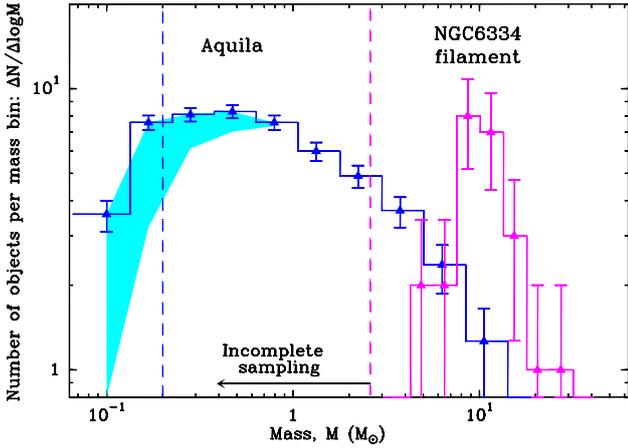


Fig. 9. Mass distribution of the 26 ALMA cores identified in the NGC 6334 main filament (magenta) compared to a scaled version of the prestellar CMF observed in the Aquila cloud (blue, from Könyves et al. 2015). In both NGC 6334 and Aquila, cores were extracted from the data using the same algorithm `getsources`. For easier comparison, the Aquila CMF was re-normalized to have a peak value comparable to that of the NGC 6334 CMF. All error bars correspond to \sqrt{N} counting statistics. The magenta vertical dashed line indicates the estimated 90% completeness level ($2.6 M_{\odot}$) of our ALMA 3.2 mm census of dense cores in the NGC 6334 filament. The blue vertical dashed line marks the 90% completeness level of the Aquila CMF ($0.2 M_{\odot}$). (The blue shaded area reflects uncertainties associated with the uncertain classification of starless cores as bound or unbound objects.)

Takahashi et al. (2013) in Orion OMC3 and Kainulainen et al. (2013) in the infrared dark cloud (IRDC) G11.11-0.12 (see also Teixeira et al. 2016 and Kainulainen et al. 2017). Recent theoretical works on filament fragmentation have tried to account for these two fragmentation modes based on perturbations to standard cylinder fragmentation models (cf. Clarke et al. 2016; Gritschneider et al. 2017; Lee et al. 2017).

4.2. Unusually massive cores in the NGC 6334 filament

Remarkably, the median mass of the ALMA cores ($9.6^{+3.0}_{-1.9} M_{\odot}$) is an order of magnitude higher than the peak of the prestellar core mass function (CMF) at $\sim 0.6 M_{\odot}$ as measured with *Herschel* in nearby clouds (e.g. Könyves et al. 2015). This is illustrated in Fig. 9 which compares a rough estimate of the CMF in the NGC 6334 filament based on the present ALMA study to the CMF derived from *Herschel* data in the Aquila cloud. Although the two data sets differ somewhat in nature and wavelength (e.g., ALMA 3.1 mm vs. *Herschel* 160–500 μm), cores were extracted from dust continuum maps using the same algorithm (`getsources`) in both cases. Moreover, the ALMA 3.1 mm continuum data used here are sensitive to the typical size scales of dense cores identified with *Herschel* in Gould Belt clouds (~ 0.02 to ~ 0.1 pc – see Sect. 2 and Fig. 7 of Könyves et al. 2015). The two CMFs shown in Fig. 9 are thus directly comparable. The difference in typical core mass between NGC 6334 and Aquila largely exceeds possible uncertainties in the dust emissivity at 3.1 mm. For instance, if the dust emissivity index $\beta = 1.5$ (instead of $\beta = 2$ as assumed in Table 2 and Fig. 9), the median core mass in NGC 6334 would still be a factor ~ 5 higher than the CMF peak in Aquila. The median mass of observed ALMA cores essentially corresponds to the peak of the NGC 6334 CMF, and lies more than a factor of three above the estimated 90% completeness level ($2.6 M_{\odot}$) of the present core

census⁷. While the NGC 6334 CMF shown in Fig. 9 is admittedly quite uncertain due to, for instance, low-number statistics and uncertainties in the 3.1 mm dust emissivity, we stress that it represents one of the first estimates of the CMF generated by a single, massive filament⁸ (see Takahashi et al. 2013; Zhang et al. 2015; Ohashi et al. 2016 for examples of CMFs in somewhat less massive filaments or IRDCs with $M_{\text{line}} < 500 M_{\odot} \text{pc}^{-1}$).

The median core mass derived from the ALMA data is roughly consistent with the effective critical Bonnor-Ebert mass $M_{\text{BE,eff}}$ in the clumps of the filament, which can be expressed as (Bonnor 1956):

$$M_{\text{BE,eff}} = 1.18 \frac{C_{\text{s,eff}}^3}{\sqrt{G^3 \rho_{\text{clump}}}}, \quad (5)$$

or about $7.2 M_{\odot}$ (lower quartile: $5.0 M_{\odot}$, upper quartile: $23.9 M_{\odot}$). Together with the results of Sect. 4.1 on core grouping, this supports the view that the ALMA dense cores arise from effective Jeans-like fragmentation of the ArTéMiS clumps. Our results also suggest that more massive cores may form in denser/more massive filaments and are consistent with a picture in which the global prestellar CMF, and possibly the stellar initial mass function (IMF) itself, originate from the superposition of the CMFs generated by individual filaments with a whole spectrum of masses per unit length (André et al. 2019).

4.3. Likelihood of the NGC 6334 filament being a system of two velocity-coherent fibers

As described in Sect. 3.3, our ALMA observations show that the ~ 0.15 -pc wide filament detected in the ArTéMiS 350 μm dust continuum map is sub-structured in five fiber-like $\text{N}_2\text{H}^+(1-0)$ components with different velocities. The typical length of these fiber-like sub-structures is 0.5 ± 0.4 pc and the typical velocity difference between them is 0.8 km s^{-1} (standard deviation of the five N_2H^+ velocity components). The N_2H^+ velocity-coherent sub-structures may be broadly categorized into two groups. The first group consists of the N_2H^+ fiber-like sub-structures F-1 and F-2, which are also detected in other molecular line tracers of dense gas such as CH_3CCH , H_2CS , HC_3N , and HC_5N and harbor compact dense cores detected in 3.1 mm continuum emission. The second group (F-3, F-4, and F-5) is made up of velocity-coherent sub-structures detected only in N_2H^+ and which do not seem to contain dense cores. There are at least two possible explanations for these two groups. First, the physical excitation conditions in the two groups of sub-structures may differ. Second, the low velocity resolution of the current CH_3CCH , H_2CS , HC_3N line data (see Table 1) may prevent the identification of fiber-like features in these tracers. To test the former explanation, observations in several transitions of each species would be required to estimate the excitation temperature, column density,

⁷ The 90% completeness level was estimated to be $2.6 M_{\odot}$ assuming Gaussian statistics and a 5σ peak-flux detection threshold of $0.7 \text{ mJy beam}^{-1}$. Under these assumptions, $\sim 10\%$ of sources with intrinsic peak fluxes at the 6.3σ level of $\sim 0.9 \text{ mJy beam}^{-1}$ will have measured peak fluxes below 5σ . Thus, one can expect the 90% completeness level to roughly correspond to the 6.3σ level (cf. Belloche et al. 2011).

⁸ Despite the presence of fiber-like sub-structures, we regard the prominent filamentary structure seen in the ArTéMiS 350 μm continuum map and the ALMA $\text{N}_2\text{H}^+(1-0)$ integrated intensity map (cf. Fig. 1) as a single filament, comparable to the many filaments detected with *Herschel* in nearby clouds (Arzoumanian et al. 2011, 2019) and to the B211/B213 filament in Taurus (Palmeirim et al. 2013; Hacar et al. 2013, and Sect. 4.3).

and chemical abundance of the various molecules. To investigate the latter effect, ALMA observations of the same tracers at higher spectral resolution would be required.

The presence of velocity-coherent fiber-like sub-structures in molecular filaments was first reported by Hacar et al. (2013) in the case of the low-mass star-forming filament B211–B213 in Taurus ($d \sim 140$ pc). In this filament, Hacar et al. (2013) used their friends in velocity (FIVE) algorithm to identify at least 20 velocity-coherent components in N_2H^+ and C^{18}O (see Table 3 in Hacar et al. 2013), which were subsequently called fibers. Since then, similar velocity-coherent components have also been detected in N_2H^+ in the IRDC G035.39-00.33 (Henshaw et al. 2014), the NGC 1333 protocluster (Hacar et al. 2017), IRDC G034.43+00.24 (Barnes et al. 2018), and the Orion A integral-shaped filament (Hacar et al. 2018). The fiber-like sub-structures identified in NGC 1333 and Orion A may, however, differ in nature from those observed in the Taurus B211/B213 filament and in the present target NGC 6334 (see also Clarke et al. 2017). The velocity-coherent sub-structures observed in the NGC 1333 and Orion A cases are indeed well separated in the plane of the sky, while those in the Taurus and NGC 6334 filaments overlap in the sky and can mostly be distinguished in PPV space.

The typical length (0.6 ± 0.5 pc) and velocity difference between components (0.7 km s^{-1}) reported by Hacar et al. (2013) for the fiber-like sub-structures of the (low-mass) B211/B213 filament in Taurus are remarkably similar to the properties estimated here for the velocity-coherent sub-structures of the (high-mass) NGC 6334 filament. We further note that Hacar et al. (2013) divided up their B211/B213 fibers into two groups, fertile and sterile, depending on whether they contained dense cores or not. Most of the 35 velocity-coherent sub-structures identified in B211/B213 were sterile and detected mostly in $\text{C}^{18}\text{O}(1-0)$, while only 7 fiber-like sub-structures were fertile and also detected in $\text{N}_2\text{H}^+(1-0)$. This is reminiscent of the situation found here for the NGC 6334 filament, where the two sub-structures detected in multiple dense gas tracers (F-1 and F-2) are the only two fertile fibers harboring ALMA dense cores. We argue below that these two categories of fiber-like structures differ in physical nature and possibly origin.

4.4. Possible origin(s) of the fiber-like sub-structures

Three scenarios for the formation of velocity-coherent fiber-like sub-structures have been proposed in the literature. One is the “fray and fragment” scenario proposed by Tafalla & Hacar (2015) and supported by Clarke et al. (2017). In this scenario, a main filament forms first by collision of two supersonic turbulent gas flows. Then, the main filament fragments into an intertwined system of velocity-coherent sub-structures, due to a combination of residual turbulent motions and self-gravity. In this picture, the sub-structures are formed by fragmentation of a single filament, and the velocity-coherent sub-structures are expected to be roughly aligned with the main filament (cf. Smith et al. 2014). The second scenario is the “fray and gather” scenario, in which turbulent compression first generates short, velocity-coherent filamentary structures within the parent cloud, which are then gathered by large-scale collapse of the cloud, as proposed by Smith et al. (2014). In a third, alternative picture, a dense star-forming filament forms within a shell-like molecular gas layer as a result of large-scale anisotropic compression associated with, for instance, expanding bubble(s) or cloud–cloud collision (Chen & Ostriker 2014; Inutsuka et al. 2015; Inoue et al. 2018), and subsequently grows in mass by accreting ambient gas from the surrounding shell-like structure due to its own gravitational

potential (cf. Palmeirim et al. 2013; Shimajiri et al. 2019). This accretion process supplies gravitational energy to the dense filament, which is then converted into turbulent kinetic energy in the form of MHD waves (Hennebelle & André 2013), explaining the increase in velocity dispersion with column density observed for thermally supercritical filaments (Arzoumanian et al. 2013). A quasi-stationary state is reached as a result of a dynamical equilibrium between accretion-driven MHD turbulence and the dissipation of this MHD turbulence owing to ion-neutral friction, possibly accounting for the roughly constant filament width of ~ 0.1 pc (Hennebelle & André 2013). In this “compress, accrete, and fragment” scenario, sterile fiber-like structures would correspond to portions of the accretion flow onto the central filament (see also Clarke et al. 2018), while fertile fiber-like structures would be the direct imprint of accretion-driven MHD waves within the main filament system.

While further observational constraints will be needed to fully discriminate between the above three pictures, the available constraints tend to support a scenario intermediate between the fray and fragment and the fray and gather picture, perhaps more similar to the compress, accrete, and fragment picture that we propose here. Indeed, the NGC 6334 region may be affected by cloud–cloud collision (Fukui et al. 2018, see also Appendix E), and some of the N_2H^+ velocity-coherent sub-structures identified here (e.g., F-5) are not aligned with, but roughly perpendicular to, the main filament traced at $350 \mu\text{m}$ by ArTéMiS and in N_2H^+ by the two main sub-structures (F-1, F-2 in Table 3). Furthermore, Shimajiri et al. (2019) recently reported kinematic evidence that the B211/B213 filament in Taurus may have formed inside a shell-structure resulting from large-scale compression.

5. Conclusions

To study the detailed density and velocity structure of the massive filament in the NGC 6334 complex, we carried out ALMA observations at $\sim 3''$ resolution in the 3.1 mm continuum and the $\text{N}_2\text{H}^+(1-0)$, $\text{HC}_5\text{N}(36-35)$, $\text{HNC}(1-0)$, $\text{HC}_3\text{N}(10-9)$, $\text{CH}_3\text{CCH}(6-5)$, $\text{H}_2\text{CS}(3-2)$ lines. Our main results may be summarized as follows:

- Both the 3.1 mm continuum emission and the $\text{N}_2\text{H}^+(1-0)$, $\text{HC}_5\text{N}(36-35)$, $\text{HC}_3\text{N}(10-9)$, $\text{CH}_3\text{CCH}(6-5)$, $\text{H}_2\text{CS}(3-2)$ lines detected with ALMA trace the dusty filament imaged earlier at $350 \mu\text{m}$ at lower ($8''$) resolution with the ArTéMiS camera on APEX.
- We identified a total of 40 compact dense cores in the ALMA 3.1 mm continuum map, 26 of them being embedded in the NGC 6334 filament. The majority (21/26 or 80%) of these dense cores are also detected in $\text{N}_2\text{H}^+(1-0)$ emission. The median core mass is $10^{+3}_{-2} M_\odot$ (lower quartile: $8 M_\odot$, upper quartile: $13 M_\odot$), compared to a 5σ mass sensitivity of $2 M_\odot$.
- The CMF derived from the sample of ALMA cores in the NGC 6334 filament (Fig. 9) presents a peak at the median core mass of $\sim 10 M_\odot$, which lies an order of magnitude higher than the peak of the prestellar CMF measured with *Herschel* in nearby clouds.
- The projected separation between ALMA dense cores is 0.03–0.1 pc, which is roughly consistent with the effective Jeans length within the filament. The ALMA cores can be grouped into seven groups, approximately corresponding to dense clumps seen in the ArTéMiS $350 \mu\text{m}$ continuum map. The projected separation between these groups is 0.2–0.3 pc, which roughly agrees with the characteristic spacing of four times the filament width expected from the linear fragmentation theory of nearly isothermal gas cylinders. These two

distinct fragmentation scales are suggestive of two fragmentation modes: a cylindrical mode corresponding to groups of cores, and a spherical, Jeans-like mode corresponding to cores within groups.

- We also identified five fiber-like, velocity-coherent sub-structures within the filament by applying the `getfilaments` algorithm to the ALMA $N_2H^+(1-0)$ data cube. The typical length of these fiber-like structures is 0.5 pc and the projected velocity difference between them is $\sim 0.8 \text{ km s}^{-1}$. Only two or three of these five velocity-coherent features are well aligned with the NGC 6334 filament and may represent genuine, intertwined fiber sub-structures. The other two detected velocity-coherent features may rather trace accretion flows onto the main filament.
- With the important exception of the typical core mass (which is here an order of magnitude higher), the fragmentation properties and velocity structure of the massive ($\geq 500 M_\odot \text{ pc}^{-1}$) filament in NGC 6334 are remarkably similar to the properties observed by Hacar et al. (2013) and Tafalla & Hacar (2015) for the low-mass ($\sim 50 M_\odot \text{ pc}^{-1}$) B211/B213 filament in the Taurus cloud.
- As both regions appear to be affected by large-scale compressive flows, we suggest that the density and velocity sub-structure observed in the NGC 6334 and the Taurus filament may have originated through a similar mechanism, which we dub “compress, accrete, and fragment”.

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Appendix A: Complementary figures

Figures A.1 and A.2 are complementary figures.

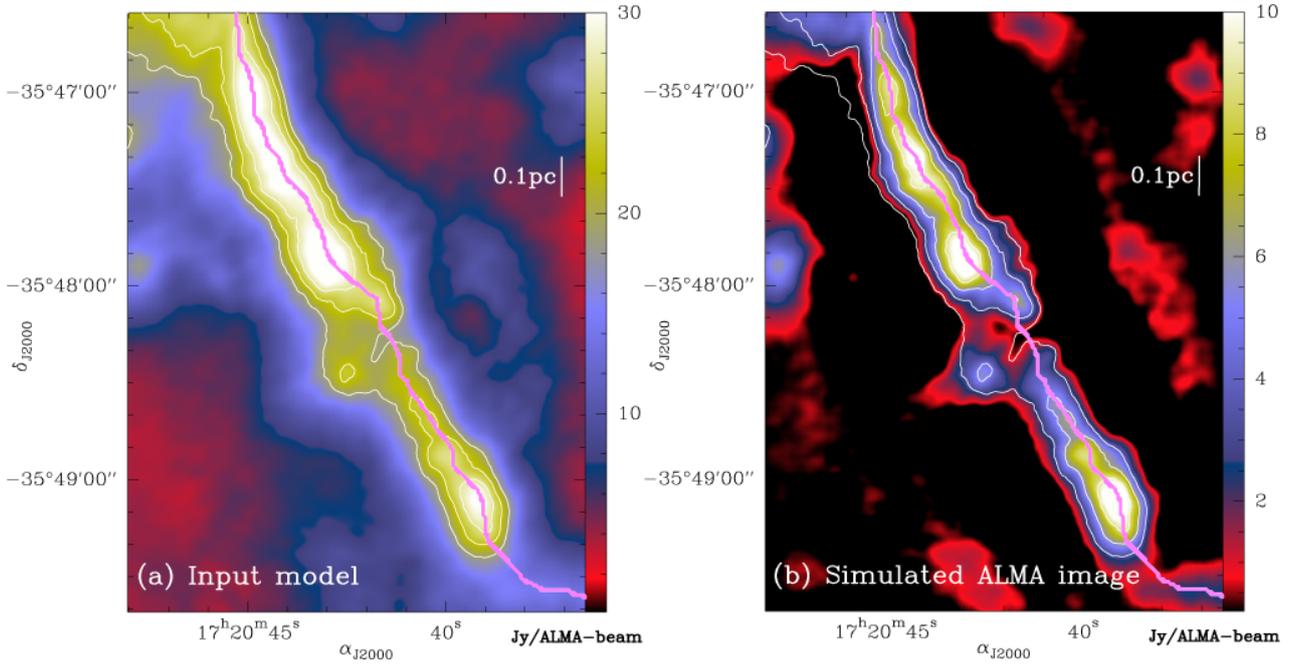


Fig. A.1. Simulations of ALMA+ACA observations performed with CASA and using the ArTéMiS+SPIRE $350\ \mu\text{m}$ image (André et al. 2016) as input model. *Panel a:* input model: ArTéMiS $350\ \mu\text{m}$ map in units of Jy/ALMA-beam. *Panel b:* simulated ALMA+ACA image using the same set of baselines (i.e., $u-v$ coverage) as the real data. While extended ($>37''$) emission is filtered out, we note that interferometric-filtering effects do not generate spurious sub-structures within the filament.

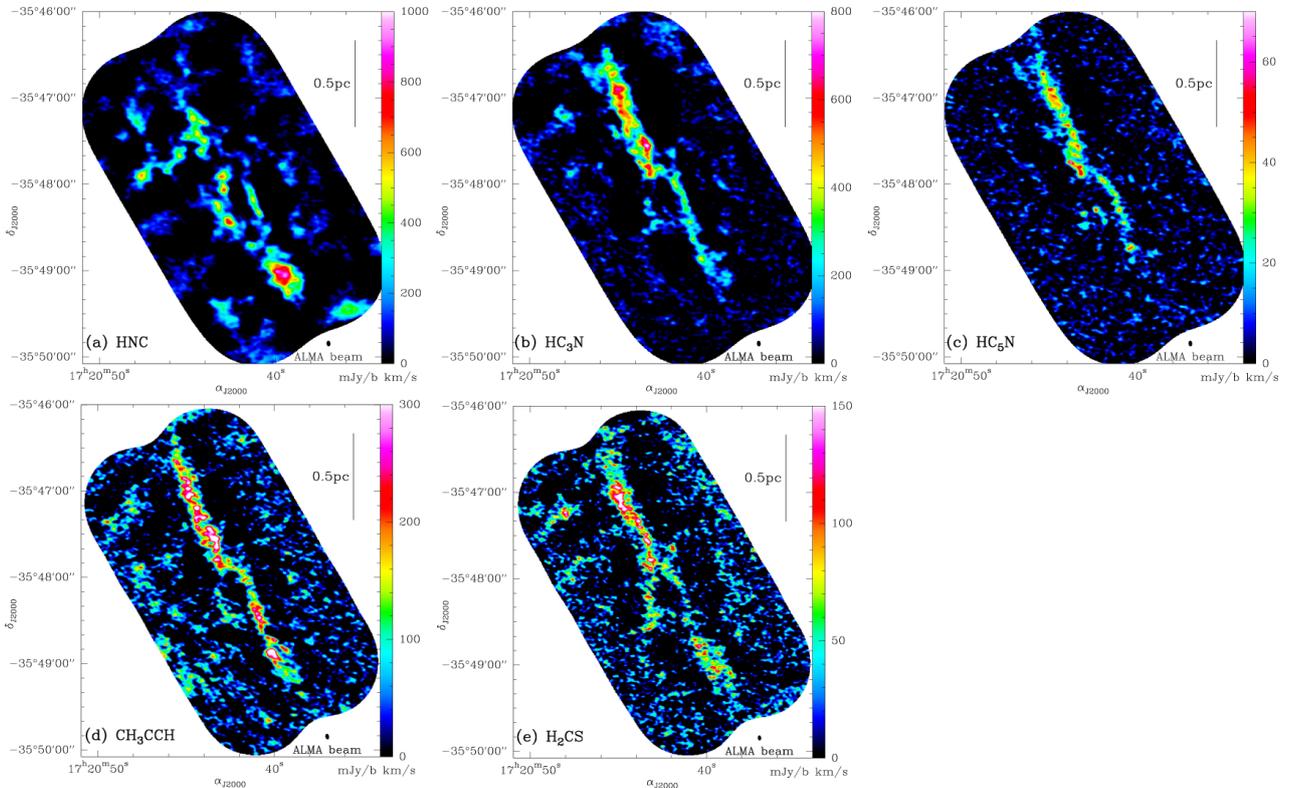


Fig. A.2. Integrated intensity maps of the NGC 6334 main filament region in HNC(1–0) (a), HC₃N(10–9) (b), HC₅N(36–35) (c), CH₃CCH(6₀–5₀) (d), and H₂CS(3_{1,2}–2_{1,1}) (e) with ALMA 12m+7m.

Appendix B: Treatment of multiple velocity components in the N₂H⁺(1–0) spectra

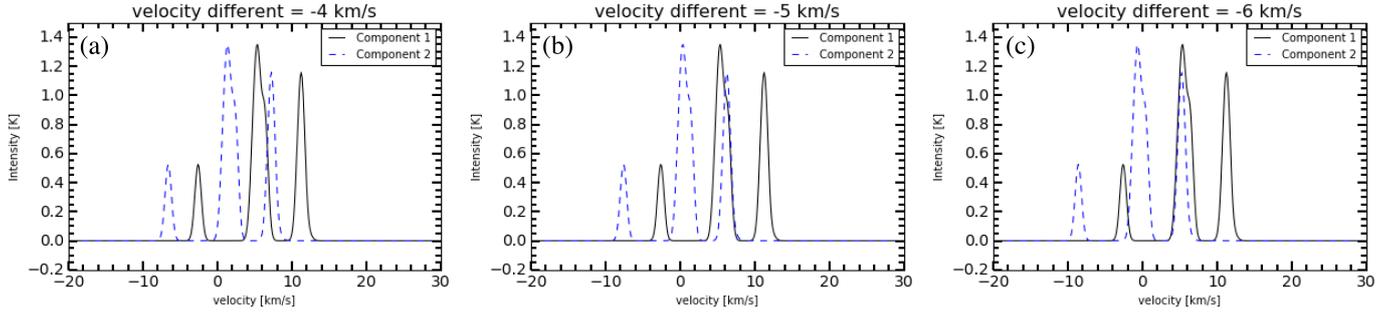


Fig. B.1. Schematic pictures of the contamination of the main velocity component the N₂H⁺(1–0) emission (1, solid line) by the HFS 2–7 emission from other velocity components. The velocity differences between component 1 and component 2 are 4, 5, and 6 km s^{−1} in panels *a*, *b*, and *c*, respectively.

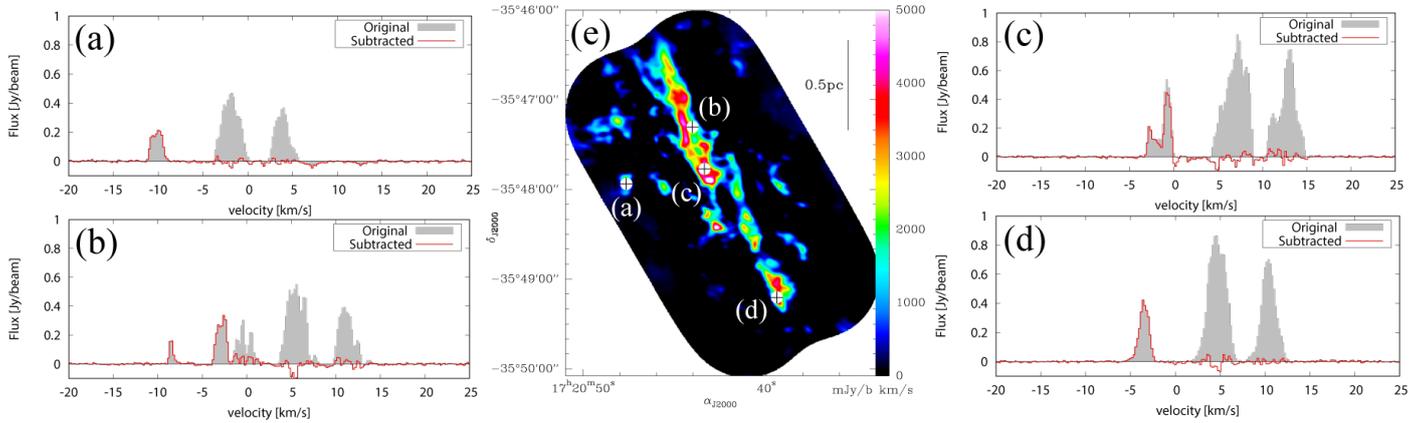


Fig. B.2. Examples of N₂H⁺(1–0) spectra before (gray) and after (red) subtraction of contaminating emission from hyperfine components HFS 2–7 (see Sect. 3.3.2). Panel *e*: black plus symbols within white circles mark the positions of the spectra shown in the other panels.

In order to disentangle between the N₂H⁺(1–0) HFS 1 emission from the main velocity component at $V_{\text{LSR}} = -6$ km s^{−1} and the HFS 2–7 emission from other velocity components, we compared the observations with synthetic N₂H⁺ emission models including two velocity components along the line of the sight.

Assuming the same excitation temperature and the same linewidth for all hyperfine transitions (HFS 1–7), and Gaussian line opacity profiles as a function of velocity, the N₂H⁺ temperature and the N₂H⁺ opacity at velocity V are given by:

$$T(V) = \frac{p_1}{p_4} (1 - e^{-\tau(V)}), \quad (\text{B.1})$$

where

$$\tau(V) = p_4 \times \sum_{i=1}^7 r_i \exp \left[-4 \ln 2 \left(\frac{V - V_i - p_2}{p_3} \right)^2 \right], \quad (\text{B.2})$$

where the parameters p_1 , p_2 , p_3 , p_4 represent $T \times \tau$, V_{LSR} , linewidth, and opacity τ for the main HFS component (HFS 3)⁹.

Figure B.1 shows model spectra in the case of two velocity components (components 1 and 2) along the line of the sight. To make it simple, we assumed $p_1 = 1.0$, $p_3 = 1.0$ km s^{−1}, and $p_4 = 0.2$ for both velocity components. We adopted a LSR

velocity $p_2 = -2.6$ km s^{−1} for component 1. For component 2, we added a velocity offset of -4 , -5 , and -6 km s^{−1} to $p_2 = -2.6$ km s^{−1}. It can be seen that, for an offset of -4 km s^{−1}, the HFS 1 emission of component 1 is not contaminated by the HFS 2–7 emission of component 2. However, for offsets between -5 and -6 km s^{−1}, the HFS 1 emission of component 1 can be significantly contaminated. Thus, to avoid contamination by the HFS 2–7 emission from other velocity components, we performed a multi-component HFS fitting of the observed spectra for all pixels showing evidence of more than one velocity component along the line of sight with blueshifted velocity offsets >4 km s^{−1} in magnitude. This HFS fitting with N HFS components was applied to each pixel, where $N = 1, 2$, or 3 . An initial HFS fitting step was first performed with $N = 1$. Whenever the peak signal to noise ratio (S/N) of the residual intensity was lower than 10, the fit was deemed to be acceptable and the corresponding spectrum was assumed to consist of $N = 1$ Gaussian components. Whenever the peak S/N of the residuals was larger than 10, an additional HFS fitting step was performed with $N+1$ components. In practice, all pixels with detected N₂H⁺ emission were fit with $N \leq 3$ components. The fit HFS 2–7 emission was then subtracted from the observed spectra. Accordingly, the resulting data cube only has significant emission from HFS 1 as shown in Fig. B.2.

⁹ See <https://www.iram.fr/IRAMFR/GILDAS/doc/pdf/class.pdf> for details.

Appendix C: Post-extraction selection of `getsources` results

In Sect. 3.3.1, we applied the `getsources` algorithm to extract compact sources from the ALMA 3.1 mm continuum map and N_2H^+ data cube. Here, we describe the additional criteria used to select robust compact sources from the `getsources` extraction results.

The `getsources` code identified a total of 49 compact sources in the 3.1 mm continuum map. First, we removed all sources lying closer than 3 ALMA beams from the edge of the map. With this criterion, 3 of the 49 initial sources were removed. We also removed 6 sources identified by `getsources` with a peak intensity below $<5\sigma$. As a result of these two selection steps, only 40 robust continuum sources remained.

Then, in order to investigate the potential association of these compact 3.1 mm continuum sources with compact N_2H^+ emission/sources, we also applied the `getsources` algorithm to each velocity channel map in the N_2H^+ data cube. In this case, we removed extracted sources with a peak intensity below $<3\sigma$ in individual channel maps. We selected only compact N_2H^+ sources detected in at least two contiguous velocity channels (where the channel width corresponds to 0.2 km s^{-1}) and lying within one ALMA beam ($\sim 3''$ or $\sim 0.02 \text{ pc}$) of a 3.1 mm continuum source. As a result, we found that 21 (or 81%) of the 26 continuum cores are associated with compact N_2H^+ emission.

Appendix D: Post-extraction selection of `getfilaments` results

In Sect. 3.3.2, we summarized the additional criteria used to select robust filamentary (sub-)structures from the `getfilaments` results. Here, we provide more details.

Our first criterion was to impose that $>90\%$ of the pixels on the crest of a robust filamentary structure are detected above the 3σ level. In Fig. D.1a, for instance, $>90\%$ of the pixels on the crest of the filamentary structure are detected above 3σ . Thus, this filamentary structure would pass the first criterion. In Fig. D.1b, $<90\%$ of the pixels on the crest of the filamentary structure are detected above 3σ . Thus, this structure would be regarded as a fake filamentary structure and would be removed from the detection list.

Then, we connected filamentary structures detected in distinct velocity channels by comparing their spatial distributions

in adjacent velocity-channel maps: whenever two structures detected in adjacent velocity-channel maps overlapped over $>60\%$ of the pixels, they were connected together.

In practice, our step-by-step procedure to associate velocity-coherent features seen in adjacent velocity channels, from the most blueshifted to the most redshifted channel, can be described as follows. Figures D.2 and D.3 show the velocity channel maps of the compact-source-subtracted N_2H^+ emission cube that were used for the identification of filamentary structures with `getfilaments` and the crests of these structures in each velocity channel (after post-extraction selection).

In the channel map at $V_{\text{LSR}} = -3.6 \text{ km s}^{-1}$, one filamentary structure is detected. This filamentary structure is labeled F-3¹⁰. In the channel maps at $V_{\text{LSR}} = -3.4$ and -3.2 km s^{-1} , three filamentary structures are detected in both channels. The spatial distributions of the detected filamentary structures are consistent in both channels. The filamentary structure whose distribution coincides with that of F-3 at -3.6 km s^{-1} is identified as F-3. The other two filamentary structures are labeled F-4 and F-1.

In the channel map at $V_{\text{LSR}} = -3.0 \text{ km s}^{-1}$, three discrete filaments are detected. The spatial distribution of one part of the northernmost filamentary structure coincides with the distribution of F-3 at -3.6 to -3.2 km s^{-1} . Thus, this filamentary structure is identified as being F-3. The spatial distribution of another one overlaps that of F-1 at -3.4 km s^{-1} . The spatial distribution of the third structure coincides with the distribution of one of the filamentary structures detected in the next velocity channel (-2.8 km s^{-1}). Thus, these features are identified as being the same filamentary structure (F-3).

In the channel maps for $V_{\text{LSR}} = -2.8$ and -2.6 km s^{-1} , two filamentary structures corresponding to F-3 and F-1 are detected. At -2.6 km s^{-1} , the filamentary structure shown as a black curve was removed as it is only detected in this channel.

In the channel map at $V_{\text{LSR}} = -2.4 \text{ km s}^{-1}$, three filamentary structures are detected. One of them was removed since it is only detected in this channel. The other two discrete structures are identified as F-1 since their distributions overlap with that of F-1 at -2.6 km s^{-1} .

In the channel map at $V_{\text{LSR}} = -2.2 \text{ km s}^{-1}$, four components can be identified. Two discrete structures correspond to F-1 since their distributions overlap with that of F-1 at -2.4 km s^{-1} .

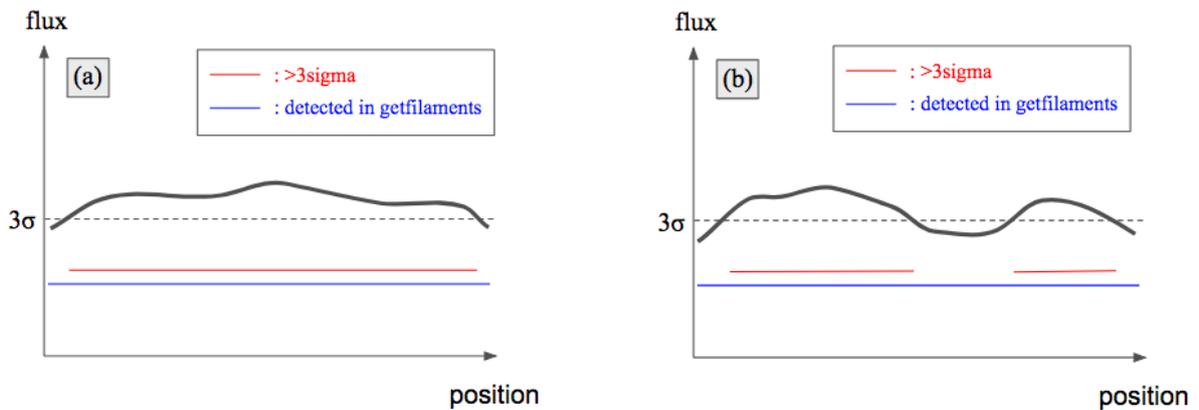


Fig. D.1. Schematic illustration of the first criterion used in the post-selection of `getfilaments` results. *Panel a:* $>90\%$ of the pixels on the crest of the filamentary structure are detected above 3σ . *Panel b:* $<90\%$ of the pixels on the crest of the filamentary structure identified with `getfilaments` are detected above 3σ .

¹⁰ In Table 3, the identified fiber-like features are listed by decreasing order of length.

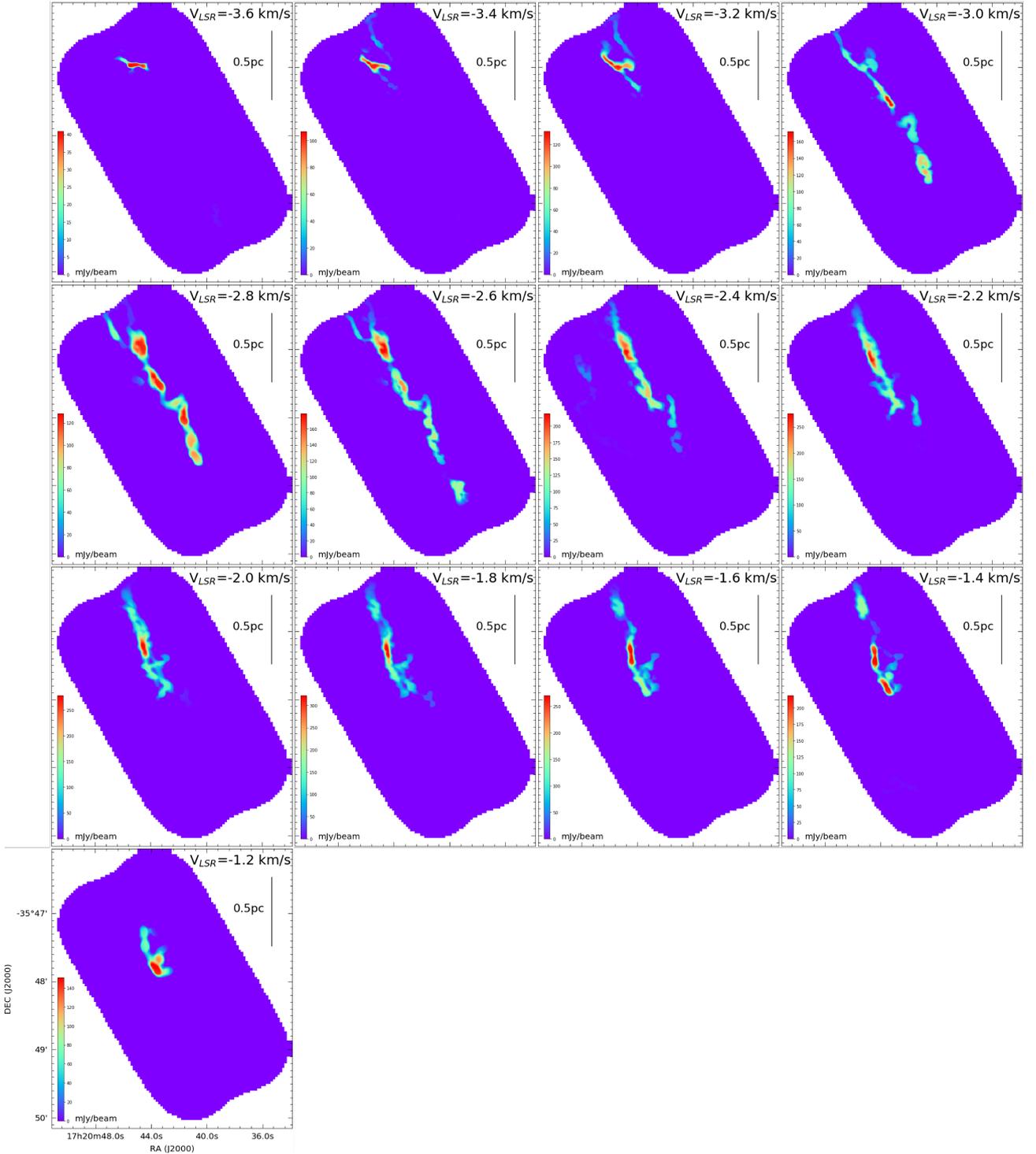


Fig. D.2. N_2H^+ velocity channel maps after subtraction of compact N_2H^+ emission using the `getfilaments` algorithm.

The filamentary structure labeled as F-3 is roughly perpendicular to the other filamentary structures. The spatial distribution of the structure labeled F-2 partly overlaps with the distribution of F-1 at -2.4 km s^{-1} . Thus, there is a possibility that F-1 and F-2 are actually part of the same physical structure. But the northern portion of F-2 lies slightly to the west of F-1, while the southern part of F-2 lies to the east of F-1. Moreover, at positions where F-1 and F-2 overlap in the plane of the sky, the N_2H^+ spectra clearly exhibit distinct velocity components (see, e.g.,

positions (c) and (d) in Fig. 6). Therefore, F-1 and F-2 seem to be distinct velocity-coherent features.

In the channel maps at $V_{\text{LSR}} = -2.0$ to -1.6 km s^{-1} , two filamentary structures (F-5 and F-2) can be identified. But F-2 lies slightly to the west of F-1 in the northern part ($\text{Dec} > -35:47:30$), while the southern part of F-2 lies to the east of F-1 in the southern part ($\text{Dec} < -35:47:30$). At -1.8 km s^{-1} , one structure was removed since it is detected in only one channel.

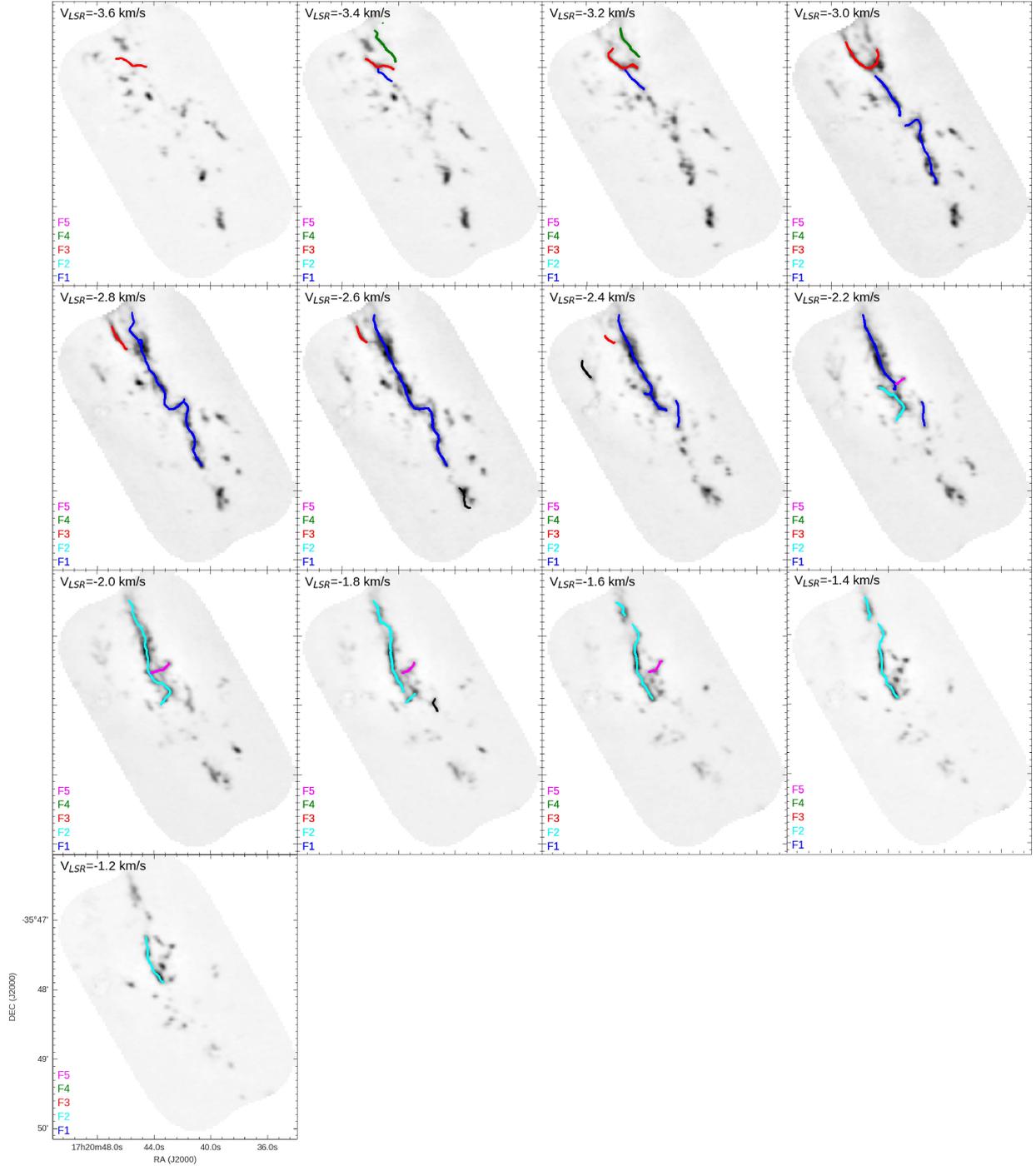


Fig. D.3. Crests of the filamentary structures identified in the N_2H^+ velocity channel maps. In each panel, the grayscale shows the distribution of N_2H^+ emission. Dark blue, light blue, red, green, and magenta curves mark the crests of F-1, F-2, F-3, F-4, F-5, respectively. Black curves indicate removed filamentary structures which are detected in only one velocity channel.

Appendix E: Origin of the highly blueshifted N_2H^+ blobs

We detected highly blue-shifted N_2H^+ blobs, with LSR velocities up to -12 km s^{-1} compared to the filament systemic velocity of -2.6 km s^{-1} (see Sect. 3.2.3). There are at least two possibilities for their origin.

The first possibility is that the highly blue-shifted N_2H^+ blobs are associated with cloud–cloud collision. Fukui et al. (2018) found the presence of highly blue-shifted emission up

to -20 km s^{-1} on ~ 10 -pc scales in their wide-field NANTEN2 ^{12}CO (1–0, 2–1) data around Galactic longitude of $l = 351.1$ – 351.5 deg which includes our ALMA observed area. They suggested that a large-scale cloud–cloud collision occurred a few Myr ago and produced this highly blue-shifted ^{12}CO emission in the NGC 6334 complex.

The second possibility is that the highly blue-shifted N_2H^+ emission blobs are associated with the expanding shell produced by source D. As described in Sect. 3.1, the distribution of the 3.1 mm continuum emission is shaped like a shell around source

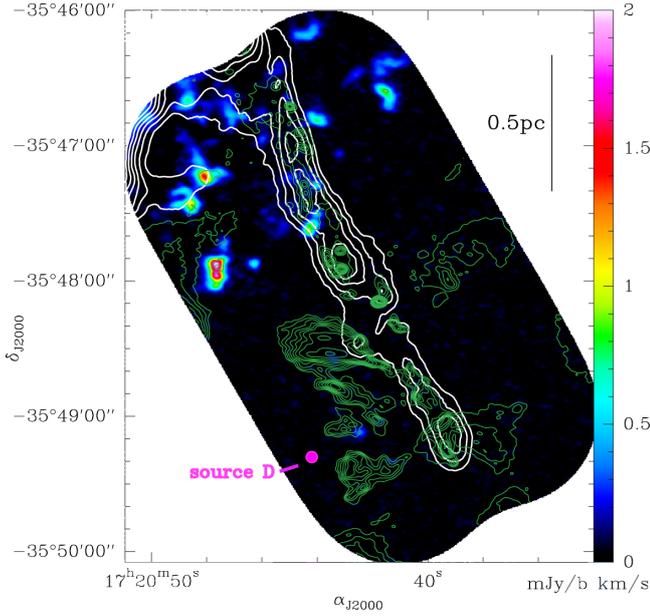


Fig. E.1. ALMA 12 m+7 m map of the $\text{N}_2\text{H}^+(1-0)$ highly blue shifted emission integrated from -12 to -6 km s^{-1} (color scale). The superimposed white contours correspond to ArTéMiS $350\ \mu\text{m}$ continuum levels of 10, 12, 14, 16, 18, and 20 $\text{Jy}/8''$ -beam. Green contours of ALMA 3.1 mm continuum emission are also overlaid with levels of 2, 4, 6, 8, 10, 15, 20, 25, 30, 35, and 40σ , where $1\sigma = 0.2$ $\text{mJy}/\text{ALMA-beam}$. The magenta filled circle marks the position of radio continuum source D (Rodríguez et al. 1982) associated with a compact HII region.

D. Most of the highly blue-shifted N_2H^+ blobs are distributed on the outskirts of this shell-like structure. While the northern part of the NGC 6334 filament is seen in absorption in the *Spitzer* $8\ \mu\text{m}$ map, such absorption is not seen in the southern part (around the shell-like structure). This suggests that the southern part of the NGC 6334 filament is in the back of the exciting star (source D). If the N_2H^+ blobs were associated with the shell-like structure, their velocity would be expected to be redshifted compared to the surrounding gas material. This is not consistent with the fact that the N_2H^+ blobs are blueshifted.

We conclude that the highly blueshifted N_2H^+ blobs are more likely related to the cloud–cloud collision scenario proposed by Fukui et al. (2018).

Appendix F: Comparison between GAUSSCLUMPS, getsources, and REINHOLD extractions of clumps in the ArTéMiS map

As described in Sect. 4.1, we applied GAUSSCLUMPS to identify seven clumps in the ArTéMiS $350\ \mu\text{m}$ map. To investigate

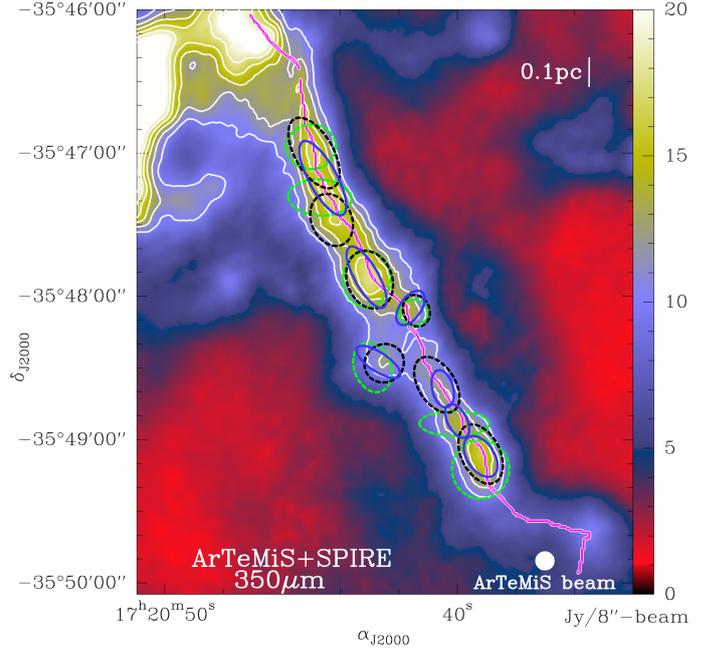


Fig. F.1. Comparison of the clumps identified in the ArTéMiS $350\ \mu\text{m}$ map (color scale) with GAUSSCLUMPS (black dashed ellipses), getsources (blue solid ellipses), and REINHOLD (green dashed ellipses).

the robustness of this identification, we also applied the getsources algorithm, already used to identify compact sources in the ALMA 3.1 mm continuum map (see Sect. 3.3.1), as well as the REINHOLD algorithm, available as a python package (*pycupid*; Berry et al. 2007). REINHOLD marks the edges of emission clumps that have shell or ring shapes. Then, all pixels within each ring or shell are assigned to be a single clump. Here, we adopted the following REINHOLD parameters: $\text{RMS} = 0.08$ Jy beam^{-1} , $\text{FLATSLOPE} = 1\sigma$, $\text{FWHMBEAM} = \text{beam size}$, $\text{MAXBAD} = 0.05$, $\text{MINLEN} = \text{beam size}/\text{pixel size}$, $\text{THRESH} = 10$, and $\text{MINPIX} = \text{beam size}$. Seven clumps were identified in the main filament. Figure F.1 compares the distributions of the clumps identified with GAUSSCLUMPS, getsources, and REINHOLD. The getsources algorithm identified seven sources within the main filament. Five of these seven getsources objects coincide with a clump found with GAUSSCLUMPS. REINHOLD also identified seven clumps within the main filament. One clump identified with GAUSSCLUMPS is not identified with REINHOLD, but is identified with getsources. All clumps identified with GAUSSCLUMPS are also identified with getsources and/or REINHOLD. Thus, we conclude that the seven clumps identified with GAUSSCLUMPS are robust.