

XMMSL2 J144605.0+685735: a slow tidal disruption event

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ABSTRACT

Aims. We investigate the evolution of X-ray selected tidal disruption events.

Methods. New events are found in near real-time data from *XMM-Newton* slews, and are monitored by multi-wavelength facilities.

Results. In August 2016, X-ray emission was detected from the galaxy XMMSL2 J144605.0+685735 (also known as 2MASX 14460522+6857311), that was 20 times higher than an upper limit from 25 years earlier. The X-ray flux was flat for ~100 days and then fell by a factor of 100 over the following 500 days. The UV flux was stable for the first 400 days before fading by a magnitude, while the optical (*U,B,V*) bands were roughly constant for 850 days. Optically, the galaxy appears to be quiescent, at a distance of 127 ± 4 Mpc ($z = 0.029 \pm 0.001$) with a spectrum consisting of a young stellar population of 1–5 Gyr in age, an older population, and a total stellar mass of $\sim 6 \times 10^9 M_{\odot}$. The bolometric luminosity peaked at $L_{\text{bol}} \sim 10^{43}$ ergs s⁻¹ with an X-ray spectrum that may be modelled by a power law of $\Gamma \sim 2.6$ or Comptonisation of a low-temperature thermal component by thermal electrons. We consider a tidal disruption event to be the most likely cause of the flare. Radio emission was absent in this event down to $< 10 \mu\text{Jy}$, which limits the total energy of a hypothetical off-axis jet to $E < 5 \times 10^{50}$ ergs. The independent behaviour of the optical, UV, and X-ray light curves challenges models where the UV emission is produced by reprocessing of thermal nuclear emission or by stream-stream collisions. We suggest that the observed UV emission may have been produced from a truncated accretion disc and the X-rays from Compton upscattering of these disc photons.

Key words. galaxies: individual: XMMSL2 J144605.0+685735 – galaxies: individual: 2MASX 14460522+6857311 – X-rays: galaxies

1. Introduction

The close approach of a star to a supermassive black hole (SMBH) can lead to the destruction of the stellar body in a process known as a tidal disruption event (TDE; Hills 1975). Gravitationally bound material returns to the black hole and is accreted, giving rise to a flare whose electromagnetic signature peaks in the extreme ultraviolet (EUV) band (Rees 1988; Ulmer 1999). These flares were first detected in the soft X-ray band by ROSAT (Komossa et al. 2004; Bade et al. 1996; Komossa & Greiner 1999; Komossa & Bade 1999), later by *XMM-Newton* and *Chandra* (Esquej et al. 2007; Saxton et al. 2012a, 2017; Maksym et al. 2010; Lin et al. 2015) (see review by Komossa 2017), and also in the UV band by GALEX (Gezari et al. 2006, 2008, 2009). In recent years, large-area optical surveys have detected candidate TDEs emitting at temperatures of a few $\times 10^4$ K (van Velzen et al. 2011; Cenko et al. 2012a; Gezari et al. 2012; Arcavi et al. 2014; Holoien et al. 2016), ostensibly too cool to be coming from an accretion disc (e.g. Bonning et al. 2007). This optical radiation has been interpreted as being due to reprocessing of the accretion radiation by an optically thick screen (Metzger & Stone

2016; Roth & Kasen 2018; Dai et al. 2018) or to emission from shocks (Piran et al. 2015). Super-Eddington accretion in the initial phase of the disruption causes a large-scale, radiation-driven outflow of material from the central engine (Strubbe & Quataert 2009), which Metzger & Stone (2016) showed would initially completely absorb the radiation from the central engine and convert it into optical/UV photons with an effective temperature similar to that observed. In this model, the screen density is expected to drop after a few months to the point where the inner thermal radiation would become visible, with the delay time and ratio of X-ray to optical/UV flux depending on the line of sight (Metzger & Stone 2016; Dai et al. 2018). Observationally, the evidence for differences in the X-ray and UV/optical timescales is mixed. The X-rays may have lagged the UV by ~32 days in ASASSN-14li (Pasham et al. 2017), but broadly fell on the same timescale (Brown et al. 2017), as they did in 2MASX 0740-85 (Saxton et al. 2017). In SDSS J1201+30 the UV flux did not change, while the X-rays dropped by a factor of 100 (Saxton et al. 2012a), whereas in ASASSN-15oi the X-ray luminosity was quite low ($L_X \sim 10^{41}$ ergs s⁻¹) at the peak of the optical flare, but 200–300 days later had increased to $L_X \sim 10^{42}$ ergs s⁻¹ (Holoien et al. 2018; Gezari et al. 2017).

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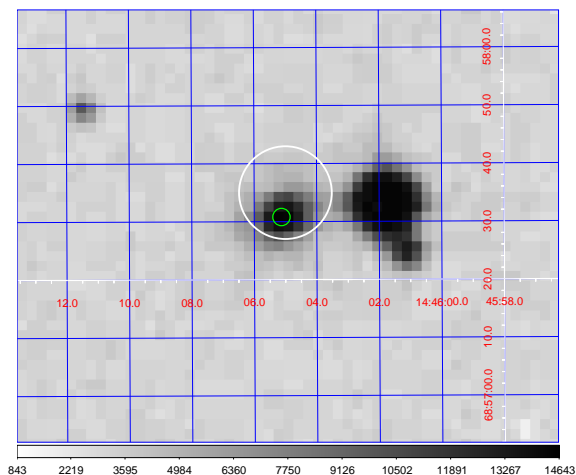


Fig. 1. Digital sky survey image. Shown is the *XMM-Newton* slew error circle ($8''$ radius; white) and UVOT-enhanced *Swift*, 90% confidence, error circle ($1.5''$ radius; green) centred on the detections and coincident with the nucleus of the galaxy 2MASX 14460522+6857311. The bright object immediately to the west of the galaxy is the 11th magnitude star HD 23235465.

Clear evidence of reprocessing is offered by broad low-ionisation optical lines, which indicate a large covering angle of material orbiting the nucleus. Komossa et al. (2008) identified broad hydrogen and helium lines and highly ionised narrow iron lines in SDSS J095209.56+214313.3, which were seen to fade over time. Similar bright, broad lines have been seen in other archival SDSS data (Wang et al. 2011) and in more recent optically discovered TDE such as ASSASN-14li (Holoien et al. 2016). Reprocessing features are, however, not generally seen in the post-disruption optical spectra of X-ray selected TDEs on timescales of days to months (e.g. Saxton et al. 2012a, 2017) or years (e.g. Komossa & Greiner 1999; Lin et al. 2017). Hence, the relationship between the emission mechanisms in the X-ray and UV bands remains an outstanding problem.

In August 2016, the *XMM-Newton* slew survey (Saxton et al. 2008) detected a flare from the nucleus of the quiescent galaxy 2MASX 14460522+6857311. Further monitoring has shown enhanced emission in the UV band. In Sect. 2 we discuss the discovery of this flare and the source identification. In Sects. 3–5 we present X-ray, UV, optical, and radio follow-up observations. In Sect. 6 we perform a temporal and spectral analysis of the source, and in Sect. 7 we discuss the flare characteristics within the TDE model. The paper is summarised in Sect. 8.

A Λ cold dark matter cosmology with $(\Omega_M, \Omega_\Lambda) = (0.27, 0.73)$ and $H_0 = 70 \text{ km}^{-1} \text{ s}^{-1} \text{ Mpc}^{-1}$ has been assumed throughout.

2. X-ray flare identification

During the slew 9305900002, performed on August 22, 2016, *XMM-Newton* (Jansen 2001) detected a source, XMMSL2 J144605.0+685735 (hereafter XMMSL2 J1446+68), with an EPIC-pn, medium filter, 0.2–2 keV count rate of $1.2 \pm 0.4 \text{ count s}^{-1}$.

The source position, with a 1σ error of $8''$, lies $6''$ from the galaxy 2MASX 14460522+6857311 (hereafter 2MASX 1446+6857) and $18''$ from the 11th magnitude A star, HD 23235465 (Fig. 1). On September 12, 2016, *XMM-Newton* made a pointed observation of the source and found an X-ray source consistent with the coordinates of the nucleus of 2MASX 1446+6857 (also known as LEDA 2725953) and no emission from the star.

It was possible to perform a very crude analysis on the 11 photons in the slew spectrum to investigate the gross spectral properties of the detection. Detector matrices were calculated, taking into account the transit of the source across the detector, using a technique outlined in Read et al. (2008). A simple absorbed power-law fit gives a slope $\Gamma = 2.7 \pm 1.0$ assuming no intrinsic absorption above the Galactic value of $1.9 \times 10^{20} \text{ cm}^{-2}$ (Willingale et al. 2013). This corresponds to an absorbed flux of $F_{0.2-2} \sim 2 \pm 0.7 \times 10^{-12} \text{ ergs s}^{-1} \text{ cm}^{-2}$ using the above model. We calculate a 2σ upper limit from the ROSAT All-Sky Survey (RASS) at this position of $0.010 \text{ count s}^{-1}$ (see Esquej et al. 2007, for a description of the upper limit calculation), a factor of 20 lower 0.2–2 keV flux using the same spectral model.

3. X-ray and UV observations

An X-ray monitoring programme was initiated with the *Neil Gehrels Swift* Observatory (hereafter *Swift*; Gehrels et al. 2004) to follow the evolution of the source flux and spectrum. Snapshot 3ks observations were made, initially once a week and then less frequently, with the X-ray telescope (XRT; Burrows et al. 2005) in photon counting mode and the UV optical telescope (UVOT; Roming et al. 2005). The *Swift*-XRT observations were analysed following the procedure outlined in Evans et al. (2009) and the UVOT data were reduced as described in Poole et al. (2008). An accurate position for the source in the *Swift*-XRT field can be determined by matching the UVOT field of view with the USNO-B1 catalogue and registering the XRT field accordingly (Goad et al. 2007). The resulting source position, $\alpha_{J2000} = 14 : 46 : 05.13$, $\delta_{J2000} = 68 : 57 : 30.8 (\pm 1.5'')$; 90% confidence) is coincident with the 2MASS position of the galactic nucleus (see Fig. 1) to within the 90% confidence uncertainty.

In parallel, five *XMM-Newton* pointed observations were triggered between September 8, 2016 and April 25 2018 (observation ID=0763640201, 0763640401, 0763640501, 0763640601, 0823330101). A summary of observations and exposure times is given in Table 1. In each observation, the EPIC-pn and MOS-1 cameras were operated in full frame mode with the *thin1* filter in place, while the MOS-2 camera was used in small window mode with the medium filter. The source was too faint for statistically significant data to be collected from the reflection grating spectrometers.

The XMM data were analysed with the *XMM-Newton* Science Analysis System (SAS v16.0.1; Gabriel et al. 2003). Light curves were extracted from the observations and searched for periods of high background flaring revealing that the first four observations were considerably affected by background. Observations 1, 3, and 4 could be cleaned by filtering out periods of high background, but observation 2 had moderate background for the whole exposure and has been used in its entirety. For observations 1 and 4 we used a single contiguous section of data from mission reference time 589743669–589748442 and 596642336–596651995, respectively. For observation 3 we created a light curve over the whole field of view from single-pixel events in the energy range 10–12 keV. We then produced a series of good time intervals where the count rate in this light curve was $\leq 1.4 \text{ c/s}$ and used these to extract the source spectrum.

4. Optical observations

A set of optical observations were initiated shortly after the discovery of XMMSL2 J1446+68. Photometry was obtained to monitor the flux evolution in the optical bands and complement the UVOT observations. Spectroscopic observations were made

Table 1. X-ray observation log of XMMSL2 J1446+68.

Mission ^(a)	Date	Exp time ^(b) (s)	Flux ^(c)
ROSAT-Survey	1990-07-11	958.2	<0.07
XMM-Newton slew	2004-11-19	10.5	<0.78
XMM-Newton slew	2013-11-03	4.9	<1.05
XMM-Newton slew	2014-04-24	8.2	<0.99
XMM-Newton slew	2014-11-07	9.5	<0.55
XMM-Newton slew	2016-08-22	8.0	1.68 ± 0.59
XMM-Newton pointed	2016-09-08	3918	0.83 ± 0.03
Swift	2016-09-12	2899	1.06 ± 0.13
Swift	2016-09-15	1071	0.50 ± 0.12
XMM-Newton pointed	2016-09-16	16140	0.84 ± 0.02
Swift	2016-09-20	2085	1.01 ± 0.20
Swift	2016-09-27	2073	1.01 ± 0.18
Swift	2016-10-04	1731	1.31 ± 0.18
Swift	2016-10-10	1064	0.82 ± 0.16
Swift	2016-10-13	1211	0.98 ± 0.18
Swift	2016-10-17	1963	1.13 ± 0.16
Swift	2016-10-22	879	1.05 ± 0.25
XMM-Newton pointed	2016-10-30	5990	0.95 ± 0.03
Swift	2016-11-02	2160	0.96 ± 0.13
Swift	2016-11-23	1408	0.81 ± 0.17
XMM-Newton pointed	2016-11-27	7982	0.69 ± 0.02
Swift	2016-12-13	2138	0.49 ± 0.10
Swift	2017-01-13	682	0.21 ± 0.13
Swift	2017-01-26	1885	0.37 ± 0.23
Swift	2017-02-17	1923	0.15 ± 0.06
Swift	2017-03-28	2210	0.126 ± 0.048
Swift	2017-04-20	1975	0.186 ± 0.057
Swift	2017-08-14	4947	0.118 ± 0.028
XMM-Newton pointed	2018-04-25	17820	0.016 ± 0.003
Swift	2018-12-20	2934	0.025 ± 0.015

Notes. ^(a)XMM-Newton, EPIC-pn camera: slew observations performed in *full frame* mode with the *medium* filter; pointed observations performed in *full frame* mode with the *thin1* filter. Swift-XRT observations performed in *pc* mode. ^(b)Useful exposure time after removing times of high background flares. ^(c)Absorbed flux in the 0.2–2 keV band, units of 10^{-12} ergs s^{-1} cm^{-2} . For simplicity this was calculated using a power-law model of slope 2.5 and Galactic absorption of 1.9×10^{20} cm^{-2} in all cases.

to check for the presence of a persistent AGN and to look for broad or narrow line features which may have been produced by material reprocessing the nuclear flare emission.

4.1. Photometry

XMMSL2 J1446+68 was observed with the ESA Test-Bed Telescope Cebreros (MPC code Z58; Ocaña et al. 2016) with the prime focus 4Kx4K CCD camera using B,V,R Johnson filters. Short observations of 30–90 seconds were taken on September 6, 2016, September 20, 2016, and October 31, 2016. Differential photometry was provided by using up to 2000 nearby stars from the catalogue UCAC-4. Photometric residuals are better than 0.1 magnitude.

4.2. Spectroscopy

A low-resolution ($R \sim 350$, $\lambda = 4000\text{--}8000$ Å) optical spectrum of XMMSL2 J1446+68 was obtained using the SPRAT

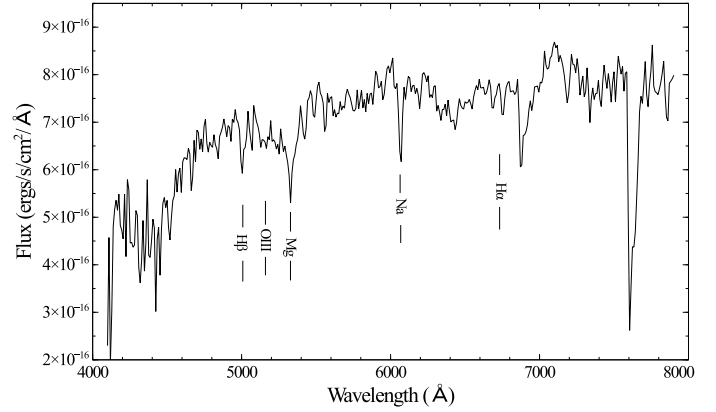


Fig. 2. Optical spectrum (observed frame) of XMMSL2 J1446+68 taken with the SPRAT spectrograph on the 2 m Liverpool Telescope on September 5, 2016.

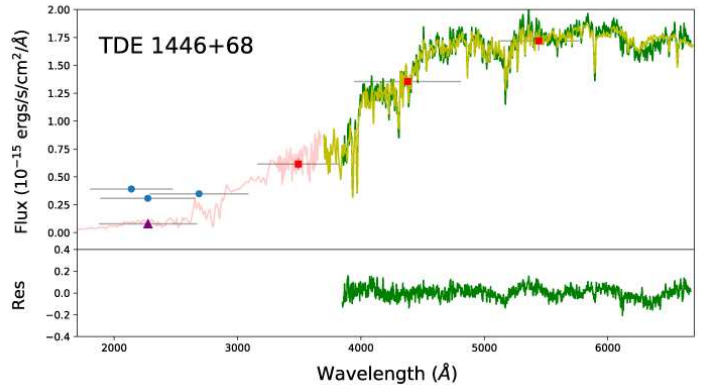


Fig. 3. Optical spectrum of XMMSL2 J1446+68 taken on September 4, 2016 with the FAST instrument on the FLWO telescope (light green) fitted with a stellar population model shown in dark green and pink (see text for details of the fitting procedure). The mean Swift-UVOT filter fluxes from the period September 12, 2016 to August 14, 2017 are overplotted as red squares (U,B,V) and blue circles (UVW1,UVW2). The GALEX-NUV filter point from 2007 is shown as a purple triangle.

spectrograph (Piascik et al. 2014) of the 2.0 m Liverpool Telescope (Steele et al. 2004) on September 5, 2016. The exposure time used was 300 s. The spectrum was reduced and calibrated using the standard pipeline (Smith et al. 2016), and showed a non-active galaxy (Fig. 2) at a redshift of $z = 0.029 \pm 0.001$.

XMMSL2 J1446+68 was also observed with the FAST spectrograph on the 60-inch telescope at Fred Lawrence Whipple Observatory (FLWO) on September 4, 2016 (Fig. 3). The two 900 s spectra were extracted and wavelength-calibrated using an automated pipeline, and flux-calibrated using a standard star taken on the same night. These observations confirmed the redshift of $z = 0.029 \pm 0.001$.

The stellar population of XMMSL2 J1446+68 was analysed using the STARLIGHT spectral population synthesis code (Cid Fernandes et al. 2005, 2011). The code uses the Bruzual & Charlot (2003) single stellar population models and requires the data to be corrected for foreground Galactic extinction and taken to the rest frame. We used the Cardelli et al. (1989) extinction curve to correct for foreground reddening and adopted $A_V = 0.064$ and shifted the spectrum to the rest wavelength using $z = 0.029$. Figure 3 presents the results from the spectral population synthesis analysis showing the observed and total synthetic

spectra and its residuals. The mean *Swift*-UVOT filter fluxes from the flat part of the UV light curve (September 12, 2016 to August 14, 2017) are overplotted on the synthetic spectrum, as is a GALEX-NUV pre-flare flux from 2007 (see Sect. 6.1). The fit needs a relatively young stellar population of 1–5 Gyr to model the blue flux in conjunction with an older population. The total initial mass of the galaxy is $1.1 \times 10^{10} M_{\odot}$ and the current mass $5.7 \times 10^9 M_{\odot}$ after correcting for the gas restored to the ISM. The residuals clearly show that no emission lines are present in XMMSL2 J1446+68.

5. Radio observations

We observed XMMSL2 J1446+68 at 6.0 GHz and 21.7 GHz with the National Science Foundation *Karl G. Jansky* Very Large Array (VLA) on 2016 September 15.16 UT and 2017 February 22.52 UT. The first observation was taken while the VLA was moving from B configuration into its most extended A configuration and used 3C48 as the flux calibrator and J1436+6336 as the gain calibrator at both frequencies. The second was taken in the most compact D configuration of the VLA and used J1438+6211 as the gain calibrator at 6 GHz, J1436+6336 as the gain calibrator at 21.7 GHz, and 3C286 as the flux calibrator at both frequencies. We reduced and imaged the data using standard Common Astronomy Software Applications (CASA) routines (McMullin et al. 2007).

We did not detect any radio emission at the enhanced *Swift* position of XMMSL2 J1446+68 or at the position of the nearby star HD 23235465. We determined the image root mean square (rms) noise at the source position using the *imtool* package (part of *pwkit*; Williams et al. 2017). In the first observation, we find rms values of $9 \mu\text{Jy}$ at 6.0 GHz and $28 \mu\text{Jy}$ at 21.7 GHz; in the second the rms values are $6 \mu\text{Jy}$ at 6.0 GHz and $28 \mu\text{Jy}$ at 21.7 GHz. The more constraining 5σ upper limits on the 6 GHz radio luminosity of XMMSL2 J1446+68 are shown in Fig. 4 in comparison to other TDEs. The radio limits for XMMSL2 J1446+68 are the deepest radio limits for any X-ray TDE obtained to date, and the deepest limits for any TDE on timescales of a few months. Only one optical TDE, iPTF16fnl, was found to have a deeper radio limit just days after discovery (Blagorodnova et al. 2017). However, iPTF16fnl remains the faintest and fastest-fading optical TDE discovered to date, and therefore may not be representative of the typical population. XMMSL2 J1446+68 is $\geq 10^5$ times fainter than the radio-loud jetted TDE Swift J164449.3+573451 and ≥ 20 times fainter than the canonical radio-weak TDE ASASSN-14li at comparable epochs (Berger et al. 2012; Alexander et al. 2016), making it highly unlikely that XMMSL2 J1446+68 launched either a relativistic jet or a sub-relativistic outflow.

6. X-ray light curve

In Fig. 5 we show the historical light curve of XMMSL2 J1446+68. The soft X-ray (0.2–2 keV) flux is effectively constant for the first ~ 90 days, with a temporary drop of a factor 2 after 24 days, which we attribute to small variations in the local accretion rate. It then drops by a factor of 10 over the following ~ 100 days and a further factor of 10 by day 600. Modelling the observed X-ray emission as a constant, from the discovery date $t_0 = \text{August 22, 2016}$ to a later date t_{drop} when the flux follows a power-law decay, we find that at time t

$$f(x) = \begin{cases} 1.06 \pm 0.03 \times 10^{-12} & \text{if } t_0 < t < t_{\text{drop}}, \\ 9.3^{+10.6}_{-5.0} \times 10^{-9} (t - t_0)^{-2.02 \pm 0.16} & \text{if } t \geq t_{\text{drop}}, \end{cases} \quad (1)$$

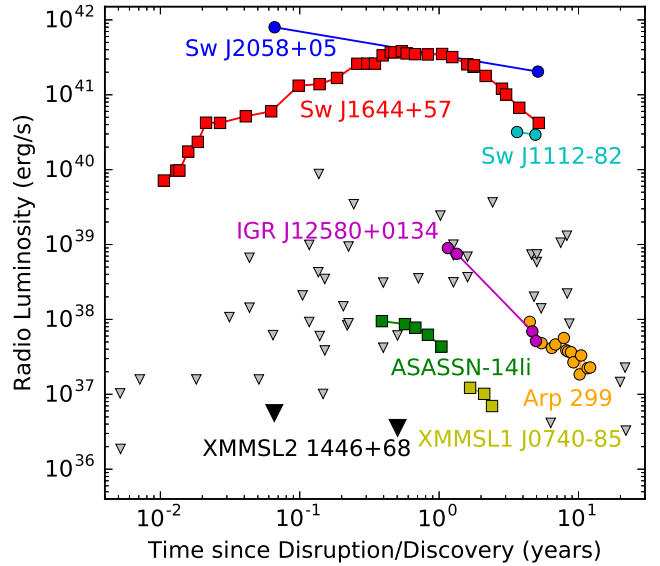


Fig. 4. Upper limits on the radio luminosity of XMMSL2 J1446+68 (black triangles) compared with other TDEs and plotted against observed time since disruption or discovery. The coloured points are radio detections of TDEs from the literature (Zauderer et al. 2011, 2013; Berger et al. 2012; Eftekhari et al. 2018; Cenko et al. 2012b; Pasham et al. 2015; Brown et al. 2015; Irwin et al. 2015; Perlman et al. 2017; Alexander et al. 2016, 2017; Mattila et al. 2018), while the grey triangles are upper limits from non-detections (Bower et al. 2013; van Velzen et al. 2013, 2019; Arcavi et al. 2014; Chornock et al. 2014; Blagorodnova et al. 2017; Blanchard et al. 2017; Alexander et al. 2019 and references therein). All upper limits are 5σ .

where $(t - t_0)$ is in units of days and $t_{\text{drop}} = \text{November 19, 2017} \pm 2$ days.

In Fig. 6 we show the exposure-corrected, background-subtracted, 0.2–2.0 keV short-term light curve from the *XMM-Newton* pointed observation of September 16, 2016. The X-ray flux shows small variations on short timescales, which can be used to obtain an estimate of the black hole mass. We used the method of Ponti et al. (2012) on a 10 ks segment of the September 9, 2016 *XMM-Newton* pointed observation, finding an insignificant excess variance of 0.007 ± 0.02 , from which we extracted a 90% confidence lower limit on the mass of $M_{\text{BH}} \geq 2 \times 10^6 M_{\odot}$. We note, however, that this technique was calibrated on AGNs, which may not be strictly applicable to the variability seen in a TDE.

Quasi-periodic oscillations (QPO) have been found in the X-ray light curves of some TDEs (Reis et al. 2012; Pasham et al. 2019; Lin et al. 2013). We barycentred the event file of the *XMM-Newton* observation of September 9, 2016, created source and background time series, and searched for a QPO in the power spectrum. No significant peaks were found.

6.1. UV light curve

During the five pointed *XMM-Newton* observations, the optical monitor (OM) cycled between the *B*, *U*, *UVW1*, *UVM2*, and *UVW2* filters. *Swift*-UVOT observations were performed with the *UVW1*, *UVM2*, and *UVW2* filters, except for the first six observations and the final observation, which also used the optical *U*, *B*, *V* filters. The galaxy was detected in all filters in all the observations except for the final *XMM-Newton* observation, which did not detect the galaxy in the *UVW2* filter using the source search software. A faint object is visible by eye, however,

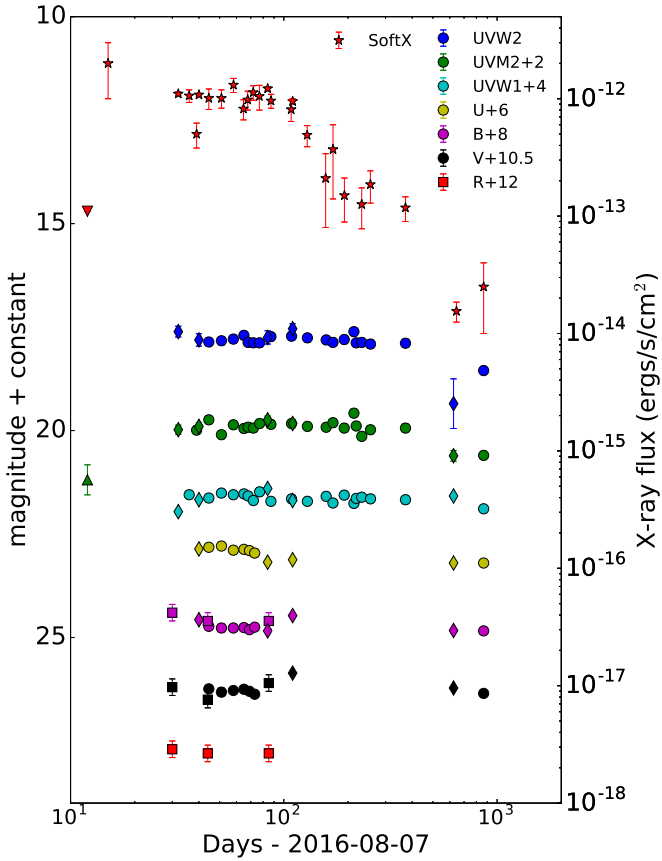


Fig. 5. Long-term X-ray, UV, and optical light curve of XMMSL2 J1446+68 plotted against time. From top to bottom: 0.2–2 keV, UVW2 (1928 Å), UVM2 (2246 Å), UVW1 (2600 Å), U (3465 Å), B (4500 Å), V (5430 Å), and R (6580 Å). X-ray points are shown as red stars, *Swift*-UVOT filters as circles, *XMM-Newton*-OM filter points as diamonds, and the Cebalros telescope points as squares. A GALEX-NUV measurement, taken in 2007 and calibrated to the UVM2 filter values, is shown as a green triangle and the RASS soft X-ray upper limit from 1990–1991 as a downward red triangle. All points show the observed fluxes and magnitudes without host galaxy subtraction.

and so the magnitude of this source was extracted by analysing the image directly. Relative filter fluxes were determined using several nearby sources of comparable brightness as references. The absolute flux scale was taken from the *Swift*-UVOT filters, with the *XMM-Newton*-OM points scaled to these values.

GALEX observed the position of the galaxy between 2005-03-05 and 2007-02-18. There is no detection of the galaxy present in the catalogue (Bianchi et al. 2017); however, a faint enhancement is visible by eye in the near-UV (NUV filter) image (2267 Å) lying within the radial profile of the bright A star, HD 23235465. By using the catalogue magnitudes of the surrounding stars and galaxies, and an analysis of the count rates in this image we calculate that the NUV magnitude of the galaxy was 21.0 ± 0.15 (AB magnitude) during the GALEX observations, which is 1.2 mag below the OM-UVM2 magnitude from September 8, 2016. The NUV filter has a similar bandpass to the UVM2 filter and gives consistent magnitudes to within $\Delta m = 0.026$, with a conservative standard deviation due to colour-dependence and other systematic effects of 0.33 mag (Page et al. 2012). To illustrate the variation in the UVM2/GALEX-NUV flux, in Fig. 7 we show images of the field around XMMSL2 J1446+68 in the GALEX observation and the

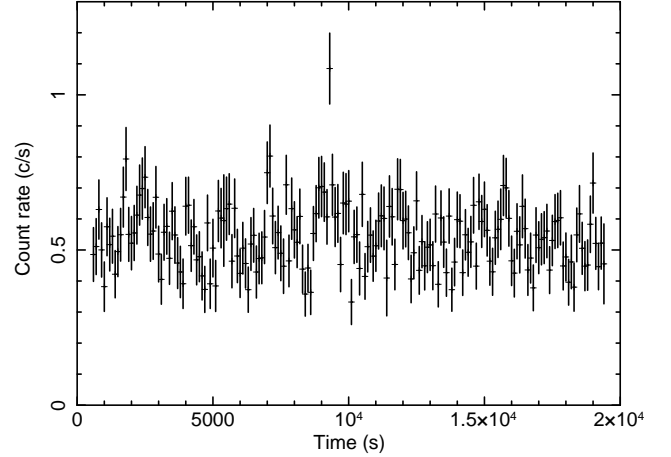


Fig. 6. Background subtracted, exposure corrected, EPIC-pn, 0.2–2 keV light curve for the September 16, 2016 *XMM-Newton* pointed observation, displayed in 100 s bins.

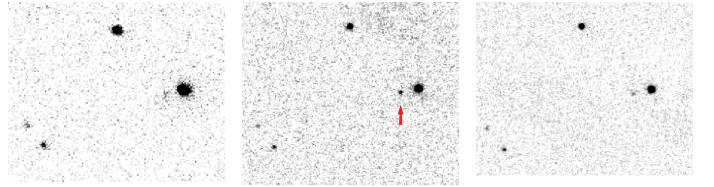


Fig. 7. GALEX-NUV image taken in 2007 (left), XMM-OM UVM2 image from September 30, 2016 (middle; XMMSL2 J1446+68 indicated by red arrow), and XMM-OM UVM2 image from April 25, 2018 (right).

OM-UVM2 observations from September 30, 2016 and April 25, 2018.

In Fig. 5 we see that the UVW2 and UVM2 fluxes remained roughly constant for the first ~ 400 days, but had fallen by ~ 1 magnitude by day 600¹. Using the GALEX-NUV observation as a measurement of the quiescent galaxy flux we find a galaxy-subtracted UVM2 flux of 1.2×10^{-13} ergs s⁻¹ cm⁻². The optical filters do not vary significantly over the 850 days of monitoring, indicating that the optical flux is dominated by the host galaxy.

6.2. X-ray spectral analysis

We extracted spectra from the second *XMM-Newton* pointed observation (hereafter XMM2), which was taken 25 days after discovery and has the highest exposure time, from circles with optimum radii determined by the task *eregionanalyse* for the EPIC-pn, MOS-1, and MOS-2 cameras. Background spectra were created from nearby source-free regions. The source spectra were grouped to have a minimum of 20 counts per bin and were fitted simultaneously within the XSPEC package (v12.8.2). Fits were performed using the chi-squared statistic over the energy range 0.3–10 keV. Quoted errors are at the 90% confidence level unless otherwise stated.

As a first step we fit the XMM2 observation with a simple power-law model and galactic absorption of 1.9×10^{20} cm⁻², modelled by TBABS with elemental abundances set to those in Wilms et al. (2000). The fit is good ($\chi^2_r = 354/347$) with a slope of $\Gamma = 2.59 \pm 0.03$. If we fit only over the 2–10 keV

¹ We confirmed that none of the *Swift*-UVOT observations fell on the regions of low sensitivity reported in Edelson et al. (2015).

Table 2. Spectral fits to *XMM-Newton* pointed observations of XMMSL2 J1446+68.

Date	Γ ^(a)	Norm ^(b) keV ⁻¹ cm ⁻² s ⁻¹	Flux (0.2–2 keV) 10 ⁻¹³ ergs s ⁻¹ cm ⁻²	Flux (2–10 keV) 10 ⁻¹³ ergs s ⁻¹ cm ⁻²	χ^2 /d.o.f.
2016-09-08	2.56 ± 0.06	2.27 ± (0.10) × 10 ⁻⁴	8.4 ± 0.3	2.6 ± 0.3	116/99
2016-09-16	2.59 ± 0.03	2.29 ± (0.07) × 10 ⁻⁴	8.6 ± 0.2	2.5 ± 0.2	354/347
2016-10-30	2.58 ± 0.04	2.52 ± (0.08) × 10 ⁻⁴	9.7 ± 0.2	2.8 ± 0.2	241/237
2016-11-27	2.55 ± 0.05	1.85 ± (0.06) × 10 ⁻⁴	7.0 ± 0.2	2.2 ± 0.2	228/220
2018-04-25	2.68 ^{+0.48} _{-0.40}	4.0 ± (1.0) × 10 ⁻⁶	0.16 ± 0.03	0.04 ^{+0.04} _{-0.02}	15/13
Combined ^(c)	2.58 ± 0.03	2.20 ± (0.03) × 10 ⁻⁴	8.5 ± 0.2	2.5 ± 0.2	425/410

Notes. Fit to the broad-band *XMM-Newton* spectra, from 0.3–10 keV, of a power law absorbed by the Galactic column (model TBABS, $N_H = 1.9 \times 10^{20}$ cm⁻²). Errors are at the 90% confidence level. Observations of 2016-09-08, 2016-09-16, 2016-10-30, and 2016-11-27 use the EPIC-pn, MOS1, and MOS2 data; the 2018-04-25 observation only uses the EPIC-pn data. ^(a)Power-law slope, ^(b)power-law normalisation, ^(c)Combination of observations 1-4 (2016-09-08 – 2016-11-27), EPIC-pn data only, without filtering for periods of higher background.

range, the slope flattens to 2.48 ± 0.26 and small residuals are present at lower energies. We attempted to model the low-energy excess as an extra emission component, with a black body of $kT = 132_{-14}^{+20}$ eV ($\chi_r^2 = 344/345$). The extra component is not statistically significant. We also tried modelling the soft component with a Bremsstrahlung ($\chi_r^2 = 347/345$; $kT = 428_{-75}^{+63}$ eV) and power law ($\chi_r^2 = 354/345$). We repeated this for the five *XMM-Newton* pointed observations², finding that an absorbed power law provided a satisfactory fit for each. In Table 2 we show the results of these fits, all of which are consistent with a power-law slope of 2.5–2.6.

The same fitting process was applied to the *Swift* XRT observations from 0.3–10 keV. Figure 8 shows the fitted slope of each observation, which is effectively constant, within the errors, between discovery and 600 days later.

To improve the statistics we combined the counts from the clean data in observations 1, 3, and 4 from the *XMM-Newton* EPIC-pn camera. This yielded a power-law slope of 2.58 ± 0.04 and $\chi^2 = 165/183$. Using the full EPIC-pn dataset from observations 1-4, without applying any background filtering, we get a consistent slope of $\Gamma = 2.58 \pm 0.03$ and $\chi^2 = 425/410$ suggesting that the periods of higher background do not cause significant spectral deviations (Fig. 9). We then used this “dirty” combined spectrum of 45.9 ks to look for spectral features beyond the simple power law. A small improvement can be found by adding a black-body component ($\chi^2 = 414/408$) with $kT = 113_{-13}^{+12}$ eV when the power-law slope reduces to 2.43 ± 0.08 . In this case $13 \pm 6\%$ of the total 0.2–2 keV flux is provided by the soft component. A Comptonised black-body model (COMPBB in XSPEC; Nishimura et al. 1986) gives a similar result ($\chi^2 = 416/408$) with a very low black-body temperature ($kT = 10_{-8}^{+10}$ eV), electron temperature ($kT_e = 35_{-7}^{+4}$ keV), and optical depth ($\tau = 0.31_{-0.17}^{+0.24}$). In summary, the X-ray spectrum is dominated by a somewhat steep power law of slope ~ 2.6 with weak evidence for a faint excess at low energies. The fifth *XMM-Newton* observation has low statistics, but gives a consistent spectral slope of $\Gamma = 2.68_{-0.40}^{+0.48}$, to within the errors.

The unabsorbed 0.2–10 keV luminosity from the *XMM-Newton* slew observation of August 22, 2016 fit with a power law of $\Gamma = 2.58$ is $L_X = 6 \times 10^{42}$ ergs s⁻¹. The bolometric correction is quite uncertain for this spectrum and we need to model the full dataset carefully to try and estimate the flux that emerges in

the EUV band. To find the UV flux from the flare we subtract the contribution of the stellar population of the host galaxy using the GALEX flux for the UVM2 filter (see Sect. 6.1), and use the fifth *XMM-Newton* observation for the UVW2 filter. The contribution of the galaxy to the total filter flux, from the first *XMM-Newton* observation, is then 20% (UVW2) and 33% (UVM2). Under the assumption that the UV and X-ray emission both come from a coherent accretion structure, we fit the galaxy-subtracted UVM2 and UVW2 data and the X-ray spectrum from XMM observation 1, with the optxagnf model (Done et al. 2012). This model consists of a multi-colour disc and a corona of optically thick electrons which Compton scatter disc photons to X-ray energies³. Given the number of free parameters in the model we assume a non-spinning black hole ($a = 0$) and use a prior for the black hole mass.

From the relationship of black hole mass to bulge K-band luminosity (Marconi & Hunt 2003) we get $M_{BH} \sim 5 \times 10^6 M_\odot$ from the point source 2MASS magnitude of $m_K = 13.06 \pm 0.05$ (or 12.2 extended), with a systematic error of 0.3 dex ($2 \times 10^6 M_\odot$). Using the current galaxy stellar mass of $5.7 \times 10^9 M_\odot$, we find a mass of $M_{BH} = 2 \times 10^7 M_\odot$ (rms error of 0.6 dex; Reines & Volonteri 2015). The black hole mass in 2MASX 1446+6857 has also been calculated from the $M_{BH} - \sigma$ relation (Wevers et al. 2019) as $M_{BH} = 7_{-5}^{+17} \times 10^7 M_\odot$. We use the lower and higher of these estimated masses in the optxagnf fits (see Table 3). The bolometric luminosity from the XMM1 observation can be obtained from these fits using $L_{bol} = 1.3 \times 10^{38} M_{BH} \dot{m}$, where \dot{m} is the accretion rate in units of the Eddington accretion rate, giving $L_{bol} = 1.0(1.3) \times 10^{43}$ ergs s⁻¹ for a black hole mass of 5×10^6 (7×10^7) M_\odot . The radius where emission transitions from a thermal black-body spectrum to a Comptonised spectrum is 46.3 (100) R_g for the fits. At this radius, very little of the thermal emission falls within the soft X-ray band, which is therefore dominated by Compton upscattering of the UV photons. Using the same model, the *Swift* observation of August 14, 2017 (which has a similar UV flux to XMM1, but an X-ray flux lower by a factor of 10) has $L_{bol} \sim 30\%$ that of XMM1, with a lower optical depth of the Comptonising electrons. The absorption-corrected luminosity in the UVM2 and UVW2 filters, after subtraction of the quiescent galaxy flux, was $L_{uv} = 2 - 3 \times 10^{41}$ ergs s⁻¹ in both observations.

As a further test of the stability of the X-ray spectrum we calculated the ratio between a soft (0.3–1.5 keV) and hard (1.5–8 keV) energy band for all of the X-ray observations (Fig. 8). In

² Only observations 1, 2, 3, and 4 had sufficient statistics to warrant the inclusion of the MOS camera data. Observation 5 was fit with the EPIC-pn data alone.

³ It also includes a power-law component, which we set to zero in the fits.

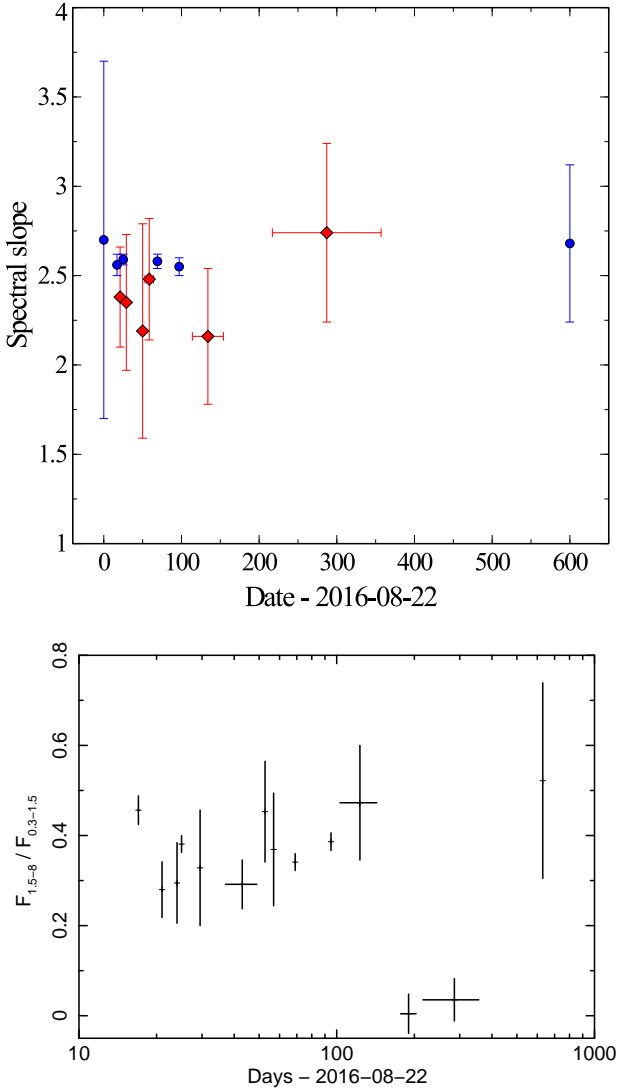


Fig. 8. *Upper:* 0.3–10.0 keV spectral slope from absorbed power-law fits to the *XMM-Newton* (blue circles) and *Swift* (red diamonds) observations of XMMSL2 J1446+68. Some of the *Swift* observations have been combined to reduce the error bar size. *Lower:* source spectral hardness expressed as the ratio of the 1.5–8 keV to 0.3–1.5 keV fluxes.

this plot we see an apparent strong softening of two spectra produced from the combination of *Swift* observations 00034707017, 00034707018, 00034707019 (February 15, 2017 to March 13, 2017) and 00034707020, 00034707021, 00034707022 (March 26, 2017 to August 14, 2017). The data from *Swift* observations 00034707017–00034707022 were combined into a single spectrum to maximise the signal-to-noise ratio, and compared with the last two *XMM-Newton* observations and the first *Swift* observation, all fit with an absorbed power law of $\Gamma = 2.58$ (Fig. 9). The combined *Swift* spectrum does not appear to be softer than the other spectra, but rather to be lacking flux above 1.5 keV. It may be fit with a power law ($\Gamma = 2.69 \pm 0.44$; $\chi^2 = 46/39$) or alternatively by a black body ($kT = 225^{+45}_{-33}$ eV; $\chi^2 = 43/39$), both absorbed by the Galactic column. Adding an empirical edge of energy ($E = 1.62^{+0.37}_{-0.10}$ keV) and optical depth ($\tau = 10.0$) to a $\Gamma = 2.58$ power-law model improves the fit (to $\chi^2 = 37/38$). We note that this trend is *not* what is expected from cold absorption, which would have exactly the opposite effect on an X-ray spectrum. Detailed photo-ionisation modelling of the TDE

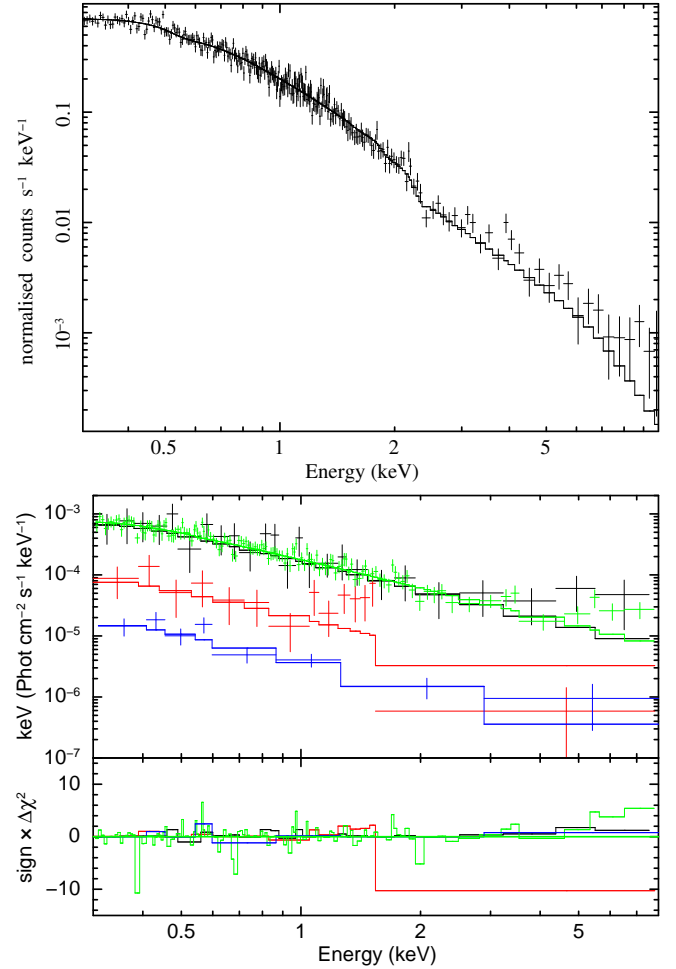


Fig. 9. *Upper:* fit of the EPIC-pn spectrum combined from the first four *XMM-Newton* observations to a power law absorbed by the Galactic column of 1.9×10^{20} cm $^{-2}$. *Lower:* spectra from the first *Swift* observation (September 12, 2016; black), the *XMM-Newton* observations of November 27, 2016 (green) and April 25, 2018 (blue), and the combined data from *Swift* between February 15, 2017 and August 14, 2017 (red), all fitted with a power law of $\Gamma = 2.58$ absorbed by the Galactic column. Residuals to the fits are shown below.

in NGC 5905 has shown that individual deep absorption edges of high column density, ionised material can potentially imprint sharp spectral features (e.g. Komossa & Bade 1999). However, the deepest edge is that expected from oxygen, around 0.7 keV rest frame (see also Kara et al. 2018). While we formally can fit the spectra with a single absorption edge at 1.62 keV, this does not correspond to any known transition, and we consider this solution highly artificial, and do not discuss it further. Alternative interpretations could involve changes in the property of the accretion disc corona, which would primarily affect the hard X-ray photons. However, at present we do not have enough data to constrain this scenario further.

7. Discussion

The lack of emission lines in optical spectra makes it unlikely that XMMSL2 J1446+68 was a steadily emitting AGN before the 2016 flare. The Wide-field Infrared Survey Explorer (WISE) colours, $W1-W2 = 0.0$, are also suggestive of a non-active galaxy (Stern et al. 2012), and with a peak bolometric luminosity of $\sim 10^{43}$ ergs s $^{-1}$ it seems more likely that this event was caused

Table 3. Spectral fits with the *optxagnf* model to UV and X-ray data of XMMSL2 J1446+68.

Date	M_{BH} (M_{\odot})	$\log L/L_{\text{edd}}$ ^(a)	r_{cor} ^(b) (R_g)	τ ^(c)	L_{bol} ergs s ⁻¹	$\chi^2/\text{d.o.f.}$
2016-09-08	5×10^6	$-1.80^{+0.03}_{-0.06}$	46.3 ± 8.0	$1.96^{+0.03}_{-0.04}$	1.0×10^{43}	191/89
2017-08-14	5×10^6	$-2.28^{+0.21}_{-0.17}$	58^{+42}_{-49}	$1.44^{+0.23}_{-0.48}$	3.4×10^{42}	8/15
2016-09-08	7×10^7	-2.83 ± 0.02	100^{+0}_{-9}	2.29 ± 0.04	1.3×10^{43}	215/89
2017-08-14	7×10^7	$-3.35^{+0.05}_{-0.20}$	100^{+0}_{-65}	$1.66^{+0.18}_{-0.23}$	4.1×10^{42}	10/15

Notes. Fit to the galaxy-subtracted UVM2 and UVW2 filter count rates and the X-ray spectrum from the *XMM-Newton* observation of 2016-09-08 and the *Swift* observation of 2017-08-14 with the *optxagnf* (Done et al. 2012) model. The hard power-law fraction is set to zero, the electron temperature (kT_e) is fixed at 10 keV, the disc outer radius is fixed at $10^3 R_g$, and a Schwarzschild non-spinning black hole is assumed. ^(a)Log of the accretion rate in units of the Eddington limit. ^(b)The disc radius at which the transition from colour-corrected black-body emission to a Comptonised spectrum occurs in units of the gravitational radius (GM/c^2). The model applies a hard cut-off at $100 R_g$. ^(c)The optical depth of the 10 keV electrons.

by a TDE. From the density of WISE objects we calculate a 0.6% probability that there is a foreground object lying within $1.5''$ of the galaxy nucleus. The possibility of galactic sources mimicking a TDE spectrum and light curve was discussed in the early TDE literature (e.g. Komossa & Bade 1999; Komossa & Greiner 1999) and the only remaining credible galactic TDE impostor is a nova (e.g. Mainetti et al. 2016). The rate of galactic novae is estimated to be 50 per year (Shafter 2017), and as the space density of galaxies is 0.0177 Mpc^{-3} (Driver 2005) then the probability of one of these novae lying within $1.5''$ of the nucleus of a galaxy out to a distance of 127 Mpc is 3×10^{-5} per year. As the slew survey covers $\sim 10\%$ of the sky each year and has been running for ten years, then the probability that the slew survey has seen one such event is 3×10^{-5} . Therefore, we conclude that the flare most likely originates within the galaxy 2MASX 1446+6857.

If we integrate the luminosity over the full light curve we obtain a total emitted bolometric luminosity of 4×10^{50} ergs. As we have no knowledge of the flux prior to discovery, this is necessarily a lower limit. This total energy release is equivalent to a consumed mass of $2 \times 10^{-3} M_{\odot}$ assuming a conversion efficiency of $\eta = 0.1$. While we may well have discovered the event post-peak and so be underestimating the total emitted radiation, this result does agree with the trend seen in many other TDEs where the radiated mass is just a fraction of a percent of a solar mass (e.g. Komossa et al. 2004; Holoien et al. 2018; Saxton et al. 2017; Maksym et al. 2010). In some earlier work this low integrated luminosity has led to very low estimates of the mass of the disrupted object, and suggestions that a giant planet or the outer layers of an evolved star had been accreted rather than debris from a main-sequence star (Li et al. 2002; Nikolajuk & Walter 2013). The apparent ubiquity of the low integrated luminosity (only 3XMM, J150052.0+015452 has so far radiated the equivalent of a large fraction of a solar mass; Lin et al. 2017) argues instead for a mechanism where the majority of the disrupted mass is not promptly accreted or where the multi-wavelength radiation is strongly suppressed. Some evidence for this is given by late-time IR emission produced by the heating of material lying at ~ 1 pc from the nucleus, which implies a true emitted luminosity ten times higher than that seen directly (e.g. Komossa et al. 2009; van Velzen et al. 2016).

This particular TDE is unusual in two key ways: a) the X-ray flux begins to fade long before the UV flux; b) the X-ray spectrum is characterised by a single steep ($\Gamma \sim 2.6$) power law rather than a soft pseudo-thermal component or a hard power law of $\Gamma \sim 1.3$ – 2.3 , as seen in the radio-loud TDE (Nikolajuk & Walter 2013; Burrows et al. 2011; Cenko et al. 2012b).

Piran et al. (2015) describe a scenario, backed up by simulations (Shiokawa et al. 2015; Bonnerot et al. 2017), where the UV and optical emission is produced by shocks from inter-stream collisions, while the X-ray radiation is produced by the later accretion of this same material. This model predicts a delay in the X-rays with respect to the UV. Support was given to this model by the X-ray flux apparently lagging behind the UV flux by 32 days in ASASSN-14li (Pasham et al. 2017) and by a delay in the onset of observed X-rays in ASASSN-15oi (Gezari et al. 2017).

Another attempt to describe the mutual exclusivity of optical and X-ray TDEs is given by the unified TDE model (Dai et al. 2018), which conjectures that X-ray TDEs are seen when the viewing angle looks down a funnel directly onto the accretion structure, while optical TDEs are produced by the reprocessing of accretion-driven X-ray photons (see also Metzger & Stone 2016; Roth et al. 2016). In this case the UV should be simultaneous with the X-rays, or a little delayed due to light-travel time, as the reprocessing is more or less instantaneous.

Neither of these models naturally explains the dichotomy between the X-ray and UV light curves in this TDE. We need a mechanism that maintains the near-UV luminosity from the TDE at an approximately constant value of $\sim 3 \times 10^{41}$ ergs s⁻¹ for 400 days while the X-ray luminosity drops by an order of magnitude. As the X-ray spectrum is consistent between the 5 *XMM-Newton* observations and the *Swift* observations before February 15, 2017, the reduction in flux in these observations is likely to be intrinsic rather than due to intervening absorption. We note that the combined *Swift* spectrum of February 15, 2017 to August 14, 2017 (days 177–356), has a relative deficit of photons at high energies, but we are unable to find a physically meaningful interpretation for this spectrum.

If the UV were purely caused by X-ray reprocessing then it should scale with the X-rays. On the other hand, if the UV is generated from inter-stream shocks then this would indicate a constant rate of return of material over the first 400 days.

A possible solution could be that the matter around the black hole formed a disc that spread outwards and accreted viscously (Cannizzo et al. 1990). van Velzen et al. (2019) saw evidence for this in late-time UV measurements of ten TDEs, consistent with a strong flattening of the far-UV and near-UV emission after 1–2 years and an average black-body temperature of 10^{4-5} K. They showed that this could result in a UV light curve that decayed very slowly ($t^{-0.1}$) with a roughly constant temperature. They predicted a low X-ray flux for their TDE sample based purely on the expected emission from disc thermal processes. We see in XMMSL2 J1446+68 and other events (Saxton et al. 2012b;

Lin et al. 2017) that Compton upscattering of thermal photons by a warm electron population provides an important contribution to the X-ray emission at all energies that cannot be neglected. In Sect. 6.2 we saw that the fading X-rays could be reproduced by a reduction in the optical depth of a population of Comptonising electrons, while the seed UV photons from the disc remain fairly constant.

The variety of X-ray spectra that are seen in TDEs is striking. Many events, for example NGC 5905, RX J1242.6-1119, RX J1420.4+5334, NGC 3599, SDSS 1323+48, ASASSN-14li, and 3XMM J150052.0+015452, were dominated by a very soft component near peak, reasonably well fit by thermal (blackbody-like) emission (e.g. Komossa & Bade 1999; Komossa et al. 2004; Esquej et al. 2008; Miller et al. 2015; Lin et al. 2017). Others, such as SDSS J095209.56+214313.3, Swift 1644+57, 2MASX 0740-85, 2MASS 0619-65, and IGR J12580+0134, displayed a hard power law (Komossa et al. 2008; Burrows et al. 2011; Saxton et al. 2017, 2014; Nikolajuk & Walter 2013) of slope $\Gamma \sim 1.3-2.3$. XMMSL2 J1446+68, however, has a steep power law ($\Gamma \sim 2.6$) that does not change, while the flux drops by a factor of 100^4 . We constrained the radio emission in this TDE to $<10\mu\text{Jy}$, the deepest upper limit for a TDE to date, which makes it unlikely that the X-rays were produced by a jet. The total energy of any off-axis jet can be constrained to $E < 5 \times 10^{50}$ ergs for circumnuclear densities $>0.1\text{ cm}^{-3}$ (Generozov et al. 2017).

The event may be self-consistently explained if the debris formed a viscously evolving, truncated accretion disc that emitted thermally in the UV and EUV bands, and a population of warm electrons that upscattered the thermal photons into a pseudo-power law of X-ray photons.

8. Summary

A flare was detected from the galaxy 2MASX 1446+6857 on August 22, 2016, reaching a bolometric luminosity, $L_{\text{bol}} \sim 10^{43}$ ergs s^{-1} . The source flux subsequently decayed by a factor of 100 in the 0.2–2 keV X-ray band over 600 days, but barely changed in the UV filters for the first 400 days and then dropped by about one magnitude in the UVM2 and UVW2 bands. The galaxy shows no signs of previous AGN activity, and we attribute the flare to the accretion of debris from a star that was tidally destroyed during a close approach to the nuclear black hole. The apparent independence of the X-ray and UV light curves may point to the creation of a slowly evolving, long-lived accretion disc structure in this event. The X-ray spectrum can be modelled with a power law of $\Gamma \sim 2.6$ throughout the evolution of the event (although a deficit of high-energy photons was noticed during the middle of the monitoring) and may be solely due to Compton upscattering of thermal photons from the disc. This event is radio-quiet to a level $L_R < 4 \times 10^{36}$ ergs s^{-1} , which constrains the total energy emitted by an off-axis jet to $<5 \times 10^{50}$ ergs, the deepest limit yet achieved for an X-ray selected TDE.

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⁴ There is evidence for a deficit of high-energy photons in the later *Swift* observations. This had disappeared by the time of the last XMM-Newton observation.

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