

Neutral hydrogen gas within and around NGC 1316

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ABSTRACT

We present MeerKAT observations of neutral hydrogen gas (H I) in the nearby merger remnant NGC 1316 (Fornax A), the brightest member of a galaxy group which is falling into the Fornax cluster. We find H I on a variety of scales, from the galaxy centre to its large-scale environment. For the first time we detect H I at large radii (70–150 kpc in projection), mostly distributed on two long tails associated with the galaxy. Gas in the tails dominates the H I mass of NGC 1316: $7 \times 10^8 M_{\odot}$ – 14 times more than in previous observations. The total H I mass is comparable to the amount of neutral gas found inside the stellar body, mostly in molecular form. The H I tails are associated with faint optical tidal features thought to be the remnant of a galaxy merger occurred a few billion years ago. They demonstrate that the merger was gas-rich. During the merger, tidal forces pulled some gas and stars out to large radii, where we now detect them in the form of optical tails and, thanks to our new data, H I tails; while torques caused the remaining gas to flow towards the centre of the remnant, where it was converted into molecular gas and fuelled the starburst revealed by the galaxy's stellar populations. Several of the observed properties of NGC 1316 can be reproduced by a $\sim 10:1$ merger between a dominant, gas-poor early-type galaxy and a smaller, gas-rich spiral occurred 1–3 Gyr ago, likely followed by subsequent accretion of satellite galaxies.

Key words. galaxies: interactions – galaxies: ISM – galaxies: individual: NGC 1316

1. Introduction

NGC 1316 (Fornax A) is a nearby, giant early-type galaxy and a well known merger remnant¹. It is the brightest member of a galaxy group at the outskirts of the Fornax cluster and, in fact, the brightest galaxy in the entire cluster volume (Schweizer 1980; Iodice et al. 2017). The stellar body of NGC 1316 is characterised by numerous tails and loops (Schweizer 1980; Iodice et al. 2017) whose brightness and size imply at least one relatively major, recent merger with a mass ratio of approximately 10:1 or below (Schweizer 1980; Mackie & Fabbiano 1998; Goudfrooij et al. 2001). The merger(s) must have happened between ~ 1 and ~ 3 Gyr ago based on the expansion time of the outer stellar loops (Schweizer 1980) and the spread in the

ages of the globular clusters (Goudfrooij et al. 2001; Sesto et al. 2016, 2018), after which NGC 1316 has likely continued accreting smaller satellite galaxies (Iodice et al. 2017).

Although NGC 1316's merger origin is unquestionable, not all pieces of the puzzle have fallen into place yet. In particular, the composition of the galaxy's interstellar medium remains difficult to understand. Given its stellar mass of $(6 \pm 2) \times 10^{11} M_{\odot}$ (Iodice et al. 2017) NGC 1316 hosts an anomalous amount of dust: $\sim 2 \times 10^7 M_{\odot}$ (Draine et al. 2007; Lanz et al. 2010; Galametz et al. 2012) compared to the $\lesssim 10^6 M_{\odot}$ expected based on known scaling relations for early-type galaxies (Temi et al. 2009; Smith et al. 2012). This could be explained if NGC 1316 formed and acquired its dust via a $\sim 10:1$ merger between a dominant, dust-poor early-type galaxy and a smaller, dust-rich spiral. Indeed, Lanz et al. (2010) estimated that a merging spiral with $M_{\star} = 1 - 6 \times 10^{10} M_{\odot}$ would bring a sufficient mass of dust to explain the observations. The problem is that the spiral would bring into the remnant a substantial amount of cold atomic and molecular gas, too: $2-4 \times 10^9 M_{\odot}$ in total, which is not observed.

¹ Throughout this paper we assume a luminosity distance of 20 Mpc for NGC 1316, consistent with the recent estimates ($20.8 \pm 0.5_{\text{stat}} \pm 1.5_{\text{syst}}$) Mpc and $(18.8 \pm 0.3_{\text{stat}} \pm 0.5_{\text{syst}})$ Mpc by Cantello et al. (2013) and Hatt et al. (2018), respectively. We assume the same distance for all other neighbouring galaxies. At this distance, 1' corresponds to 5.8 kpc. We scale all quantities from the literature to our adopted distance.

The most complete study of the cold gas content of NGC 1316 is [Horellou et al. \(2001\)](#). These authors detected $\sim 6 \times 10^8 M_\odot$ of molecular hydrogen (H_2) in the central region of the galaxy, and a further $\sim 5 \times 10^7 M_\odot$ of atomic hydrogen ($H I$) scattered in a few clouds at the outskirts of the stellar body, at a projected radius of ~ 30 kpc. They did not detect any $H I$ in the centre of the galaxy down to $\sim 10^8 M_\odot$ and they did not detect any $H I$ at larger radii. The latter is puzzling because the long ($\gtrsim 100$ kpc) stellar tidal tails of NGC 1316 have most likely formed from the kinematically-cold stellar disc of the spiral progenitor involved in the merger. The spiral disc is gas-rich and, therefore, it would be reasonable to find gas at about the same radius (if not exactly at the same location) as the optical tails.

This is indeed the case in several late-stage mergers or merger remnants involving at least one gas-rich progenitor. Once the central stellar body has relaxed into an early-type morphology, little $H I$ is usually found in the inner regions while $H I$ is found at large radii, where stellar tails and loops are still visible. Such systems include, for example, NGC 3921 ([Hibbard & van Gorkom 1996](#)), NGC 5128 ([Schiminovich et al. 1994](#)), NGC 0680 and NGC 5557 ([Duc et al. 2011](#)), NGC 7252 ([Hibbard et al. 1994](#)), Arp 220 ([Hibbard et al. 2000](#)) and Mrk 315 ([Simkin & MacKenty 2001](#)). The lack of $H I$ at large radii in NGC 1316 is therefore a problem within the context of its supposed $\sim 10:1$ early-type+spiral merger formation.

In this paper we present MeerKAT 21 cm $H I$ line observations that shed new light on the gas content of NGC 1316 and solve the above puzzle. We show that $H I$ is actually present at large radii around NGC 1316; that it coincides at least partially with some of the known optical tidal features; and that the distribution and total mass of cold gas within and around NGC 1316 is in agreement with those expected for a $\sim 10:1$ merger between a dominant early-type galaxy and a smaller, gas-rich spiral.

2. MeerKAT observations and data reduction

We observed NGC 1316 with MeerKAT ([Camilo et al. 2018](#)) in June 2018 as part of the telescope commissioning (Table 1). The observation was carried out with 36 antennas (actually 40, but 4 did not deliver useful data). The baseline length varies between 29 m and 7.5 km. Of the 630 baselines, 15 are shorter than 100 m, 170 shorter than 500 m and 274 shorter than 1 km. We used the SKARAB correlator in the 4k mode, where 4096 channels sample the frequency interval 856–1712 MHz in full polarisation. The channel width is 209 kHz, corresponding to 44.5 km s^{-1} for $H I$ at redshift $z = 0$. The visibilities were recorded every 8 s. The total integration time on NGC 1316 was 8 h. We observed the bandpass and flux calibrator PKS 1934-638 for 10 min every 2 h, and the gain calibrator PKS 0032-403 for 2 min every 12 min.

For the purpose of this work we only processed parallel hand visibilities in the frequency interval 1381–1500 MHz. We reduced the data with a pipeline² currently under development at SARAO and INAF. The pipeline is based on *Stimela*³, a Python-based scripting framework which allows users to run several open-source radio interferometry software packages in a same script using containers. The pipeline is flexible and allows users to define their own preferred data reduction strategy. For the purpose of this work we: (i) flagged the calibrators data with *AOFlogger* ([Offringa et al. 2012](#)) based on the Stokes Q visibilities; (ii) derived a time-independent, antenna-based, complex bandpass and flux scale with *CASA/bandpass* and

Table 1. MeerKAT observations and data products.

Date	2 June 2018
Observation ID	1527936442
Pointing centre (J2000)	RA 03 ^h 22 ^m 41 ^s 7 Dec -37^{d} 12 ^m 30 ^s
Total frequency range	856–1712 MHz
Processed frequency range	1381–1500 MHz
Channel width	209 kHz
	44.5 km s^{-1} for $H I$ at $z = 0$
Number of antennas	36
Bandpass/flux calibrator	PKS 1934-638
Gain calibrator	PKS 0332-403
Time on target	8 h
$H I$ cube weighting	Briggs <i>robust</i> = 0.5 20'' tapering
$H I$ cube r.m.s. noise	$0.19 \text{ mJy beam}^{-1}$ (cube centre)
$H I$ restoring PSF	$36.7'' \times 28.1''$ FWHM PA = 109 deg
3σ $H I$ column density	$2.7 \times 10^{19} \text{ cm}^{-2}$ (cube centre, single channel)
Primary beam (FWHM)	0.9 deg at 1.4 GHz

CASA/gaincal; (iii) determined frequency-independent, time-dependent, antenna-based complex gains with *CASA/gaincal*; (iv) scaled the gain amplitudes to bootstrap the flux scale with *CASA/gaincal*; (v) applied bandpass and gains to the target visibilities with *CASA/applycal*; (vi) flagged the target data with *AOFlogger* based on the Stokes Q visibilities; (vii) iteratively imaged and self-calibrated the target radio continuum with *WSclean* ([Offringa et al. 2014](#); [Offringa & Smirnov 2017](#)) and *MeqTrees* ([Noordam & Smirnov 2010](#)), respectively, solving for frequency-independent gain phase only with a solution interval of 2 min; (viii) subtracted the best continuum model from the visibilities; (ix) removed any residual continuum emission by fitting and subtracting a 3rd order polynomial to real and imaginary visibility spectra with *CASA/mstransform*; (x) Doppler corrected the visibilities to the barycentric reference frame with *CASA/mstransform*; (xi) iteratively imaged the target $H I$ emission with *WSclean* and created 3D clean masks with *SoFIA* ([Serra et al. 2015](#); the $H I$ cube was made using Briggs *robust* = 0.5 and 20'' tapering in order to obtain a more Gaussian PSF main lobe compared to natural weighting while reaching a similar surface brightness sensitivity); (xii) created a 3D $H I$ detection mask, a total $H I$ image and an $H I$ velocity field with *SoFIA*, which we then validated with visual inspection; and (xiii) used the transformations in [Meyer et al. \(2017\)](#) to convert the $H I$ image into $H I$ column density units, and the integrated $H I$ flux of each detection into an $H I$ mass assuming a distance of 20 Mpc. Throughout the paper we adopt the optical definition of radial velocity.

The size of the Gaussian restoring PSF of the $H I$ cube is $36.7'' \times 28.1''$ at FWHM and the position angle is 109 deg (north through east). The cube noise is $0.19 \text{ mJy beam}^{-1}$ per 44.5 km s^{-1} -wide channel. This corresponds to a 3σ $H I$ column density sensitivity of $2.7 \times 10^{19} \text{ cm}^{-2}$ in a single channel. Assuming a 100 km s^{-1} line-width, the 5σ point-source M_{HI} sensitivity is $5.9 \times 10^6 M_\odot$ at the assumed distance of 20 Mpc.

Compared with the Very Large Array $H I$ cube of [Horellou et al. \(2001\)](#), ours has almost exactly the same velocity resolution (44.5 vs. 41.7 km s^{-1}), $2\times$ smaller restoring PSF area (0.33 vs. 0.67 arcmin^2), $3\times$ lower noise at the centre of the field

² <https://meerkathi.readthedocs.io>

³ <https://github.com/SpheMakh/Stimela>

(0.19 vs. 0.6 mJy beam⁻¹), and ~2× larger field of view diameter at FWHM of the primary beam (0.9 vs. 0.5 deg).

As part of our data reduction we produced a radio continuum image of the field. The image includes the bright radio lobes of NGC 1316 (Ekers et al. 1983; Fomalont et al. 1989), whose flux density is ~100 Jy in our data (mostly in the short baselines). The quality of the H I cube is good despite this strong, diffuse radio emission. We present the radio continuum image in a separate paper.

3. Results: H I within and around NGC 1316

Figure 1 shows the H I cube channels in the velocity range where we detect emission: 1370–2040 km s⁻¹ (for reference, the barycentric systemic velocity of NGC 1316 is 1760 km s⁻¹; Longhetti et al. 1998). Several H I detections already discussed by Horellou et al. (2001) are visible. These include the resolved galaxies NGC 1310 and NGC 1317 (red “x” markers in Fig. 1) and the following H I clouds within 5′ of NGC 1316 (~30 kpc in projection; we mark this galaxy with a red “+” in the image):

- the extended emission-line region (EELR) south-west of NGC 1316 originally discovered by Mackie & Fabbiano (1998), which we detect at velocities 1730–1820 km s⁻¹;
- the star cluster complex SH2 south of NGC 1316 (Schweizer 1980; Richtler et al. 2012, 2017) at 1680 km s⁻¹;
- and an H I cloud just north of NGC 1316 at 1770–1860 km s⁻¹ (hereafter, C_{N,1}).

The following new H I detections are also visible in Fig. 1:

- H I in the central arcmin (~5 kpc) of NGC 1316 at velocities 1500–1860 km s⁻¹;
- a few H I clouds forming a tail between 16′ and 26′ south of NGC 1316 (90 and 150 kpc in projection, respectively) at velocities 1590–1640 km s⁻¹ (hereafter T_S);
- a few H I clouds forming a tail between 12′ and 26′ north-west of NGC 1316 (70 and 150 kpc in projection, respectively) at velocities increasing monotonically from 1730 to 1860 km s⁻¹ when moving from north to south (hereafter T_N);
- an H I cloud 12′ north of NGC 1316 at 1770 km s⁻¹ (70 kpc in projection; hereafter C_{N,2});
- and, in the south-east corner, the galaxy PGC 12706 at 1370–1460 km s⁻¹.

Most new H I clouds are unresolved in velocity. Future data with better frequency resolution will allow us to better characterise their morphology and kinematics.

Figure 2 shows primary-beam-corrected, constant-column-density H I contours overlaid on an *r*-band optical image from the Fornax Deep Survey (Iodice et al. 2016, 2017; Venhola et al. 2018). Figure 3 shows the H I velocity field. We label individual detections or systems of H I clouds in both figures. We give their primary beam-corrected H I flux and mass, with statistical errors, in Table 2. The table also lists the velocity range of each detection.

Our H I flux values for NGC 1310, NGC 1317, EELR, SH2 and C_{N,1} agree with those published by Horellou et al. (2001) within the errors. We do not confirm the detection of the H I cloud ~5′ west of NGC 1316 shown in their Fig. 2. If real, that ~10⁷ M_⊙ unresolved cloud would have been detected at a ~10σ level in a single channel of our H I cube. Conversely, H I in the centre of NGC 1316 is too faint to be detected in the Horellou et al. (2001) data (≲1.5σ per channel); and the new H I clouds at a large distance from NGC 1316 are relatively bright in our cube (especially those in T_S) but could not be detected in their data because of the higher noise and smaller primary beam.

Table 2. H I properties of the detected sources.

Source	F_{HI} (Jy km s ⁻¹)	M_{HI} (10 ⁷ M _⊙)	Vel. range (km s ⁻¹)
NGC 1316	0.45 ± 0.04	4.3 ± 0.4	1500–1860
EELR	0.33 ± 0.03	3.1 ± 0.3	1780–1820
SH2	0.30 ± 0.02	2.8 ± 0.2	1680
C _{N,1}	0.39 ± 0.04	3.7 ± 0.3	1770–1860
C _{N,2}	0.29 ± 0.03	2.7 ± 0.3	1770
T _N	1.56 ± 0.08	14.7 ± 0.8	1780–1860
T _S	4.29 ± 0.10	40.4 ± 1.0	1590–1640
NGC 1310	5.09 ± 0.09	48.0 ± 0.9	1780–1910
NGC 1317	2.80 ± 0.06	26.4 ± 0.6	1860–2040
PGC 12706	1.46 ± 0.06	13.7 ± 0.6	1370–1460

Notes. The H I flux F_{HI} is derived from the primary-beam-corrected image. The H I mass M_{HI} is calculated from F_{HI} using the equations in Meyer et al. (2017) assuming a distance of 20 Mpc. The statistical errors are obtained propagating the noise of the primary-beam-corrected H I cube to the 3D regions within which we integrate the flux. Velocities are in the barycentric reference frame (optical definition).

This demonstrates the importance of the large primary beam of MeerKAT – combined with the excellent sensitivity to diffuse emission – for the study of nearby galaxies.

Altogether, we find $4.3 \times 10^7 M_{\odot}$ of H I in the centre of NGC 1316, another $9.6 \times 10^7 M_{\odot}$ in the three already known clouds within ~30 kpc from the galaxy (EELR, SH2 and C_{N,1}) and, finally, a total of $5.8 \times 10^8 M_{\odot}$ at larger distance, 70–150 kpc in projection (T_N, T_S and C_{N,2}). In total this amounts to $7.2 \times 10^8 M_{\odot}$ of H I associated with NGC 1316 on various scales, 14 times more than previously detected. At all distances from the centre of NGC 1316, gas to the north is redshifted relative to the systemic velocity, gas to the south blueshifted. In Sects. 3.1–3.3 we discuss the properties of these H I components and their relation to the distribution of stellar light. In Sect. 3.4 we discuss the other galaxies detected in H I in this field. In Sect. 4 we revisit the assembly history of NGC 1316 in view of the new detections.

3.1. H I in the centre of NGC 1316

Our detection of H I in the central ~5 kpc of NGC 1316 adds one more component to the complex interstellar medium in this region: dust is clearly seen both in absorption (Carlqvist 2010) and emission (Lanz et al. 2010); molecular gas is also present and tracks the distribution of dust (Horellou et al. 2001; Morokuma-Matsui et al. 2019); ionised gas appears to be distributed in a disc or ring with a polar orientation relative to the stellar rotation (Schweizer 1980; Morokuma-Matsui et al. 2019); and we have now finally detected H I.

The H I non-detection by Horellou et al. (2001) demonstrated that atomic gas is not present in large amounts in this central region, resulting in a limit on the molecular-to-atomic hydrogen mass fraction $M_{\text{H}_2}/M_{\text{HI}} \geq 6$. We now detect $4.3 \times 10^7 M_{\odot}$ of H I, which sets $M_{\text{H}_2}/M_{\text{HI}} = 14$. Values of this order are not unusual in the centre of early-type galaxies (Oosterloo et al. 2010; Serra et al. 2012; Young et al. 2014).

To first order, the newly detected H I tracks the distribution of dust and molecular gas. Its kinematics is also broadly consistent with that of molecular gas (Horellou et al. 2001; Morokuma-Matsui et al. 2019) and ionised gas (Schweizer 1980; Morokuma-Matsui et al. 2019): the southern gas is blueshifted relative to the systemic velocity; the northern gas

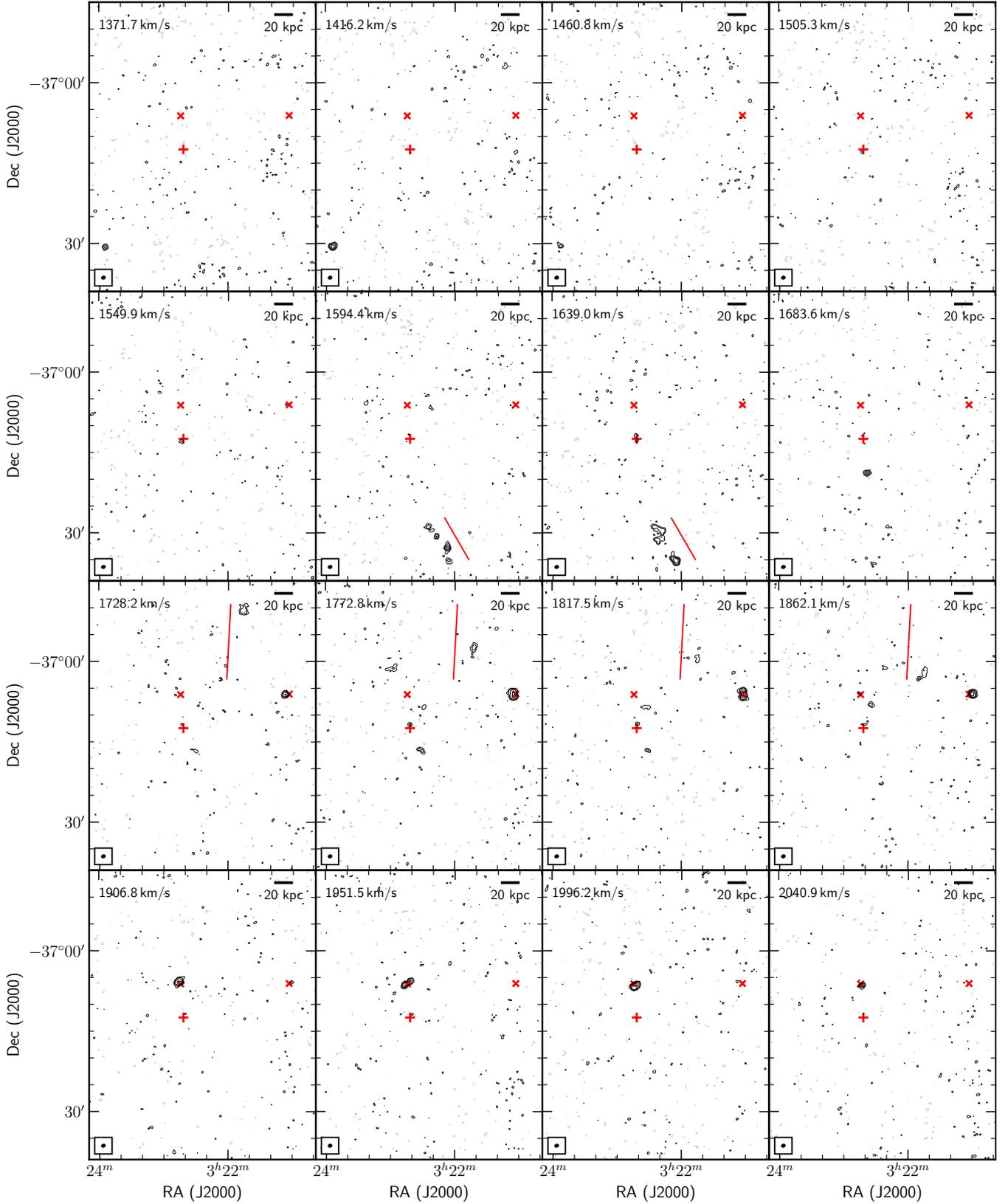


Fig. 1. Selected consecutive channel maps extracted from the H I data cube of the NGC 1316 region obtained from the MeerKAT data. The figure shows channels in the recessional velocity range where we detect H I. We list the velocity of each channel in the top-left corner. Velocities are in the barycentric reference frame (optical definition). The contour levels are $-0.5 \text{ mJy beam}^{-1}$ (grey dotted) and 0.5, 1, 2, 4, 8 mJy beam^{-1} (black solid; for reference, the noise is $0.19 \text{ mJy beam}^{-1}$). The Gaussian restoring PSF in the bottom-left corner has a FWHM of $36.7'' \times 28.1''$ and position angle 109 deg . The scale bar in the top-right corner represents 20 kpc ($\sim 3.5'$ at a distance of 20 Mpc). The red “+” marker represents the centre of NGC 1316. The two red “x” markers represent NGC 1317 (east) and NGC 1310 (west), respectively. The two red lines represent (and are drawn slightly offset from) the new, large-scale H I tails T_S (south) and T_N (north) in the channels where we detect them (see Sect. 3).

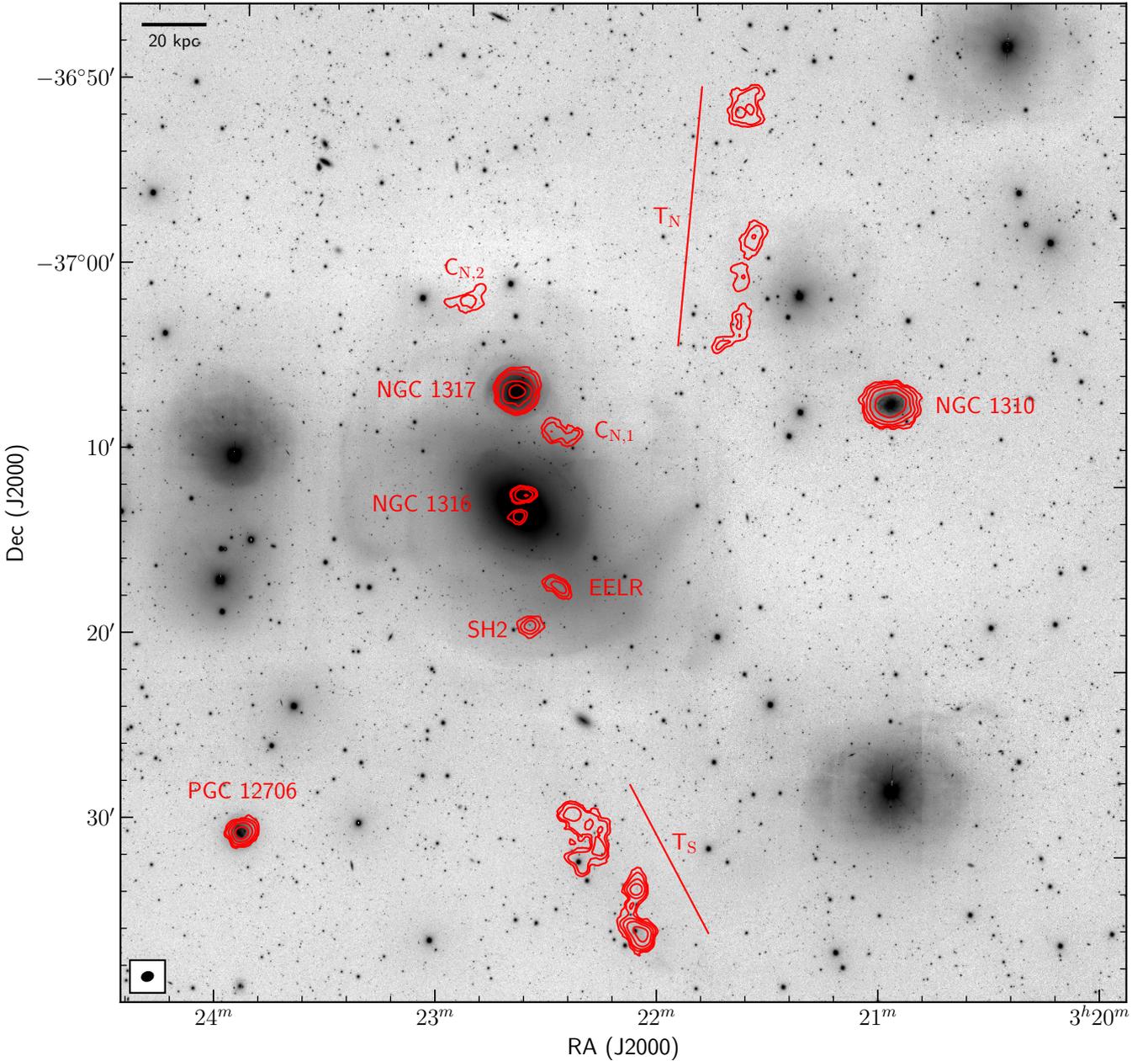


Fig. 2. Primary-beam corrected, constant column density H I contours overlaid on a *r*-band image from the Fornax Deep Survey (Iodice et al. 2017). The contour levels are $2.7 \times 10^{19} \times 2^n \text{ cm}^{-2}$ ($n = 0, 1, 2, \dots$). The lowest contour corresponds to a 3σ signal in a single channel at the centre of the H I cube (coincident with NGC 1316) but has lower significance further out. The Gaussian restoring PSF in the bottom-left corner has a FWHM of $36.7'' \times 28.1''$ and position angle 109 deg. The scale bar in the top-left corner represents 20 kpc ($\sim 3.5'$ at a distance of 20 Mpc). The H I properties of the labelled sources are listed in Table 2.

is distributed on an arc whose velocity is redshifted relative to systemic velocity on the west side, and approaches systemic velocity when moving east (Fig. 3). We will present a detailed analysis of the distribution and kinematics of the multi-phase interstellar medium in the centre of NGC 1316 in a future paper.

3.2. H I at the outskirts of NGC 1316: EELR, SH2 and $C_{N,1}$

The better sensitivity of our data allows us to obtain new details on the H I properties of EELR, SH2 and $C_{N,1}$ compared to Horellou et al. (2001). We can now measure their H I mass more precisely, and resolve the gas distribution of EELR and $C_{N,1}$.

In the EELR we detect $3.1 \times 10^7 M_{\odot}$ of H I. The emission is resolved and elongated with position angle ~ 45 deg. This is a good match to the distribution of ionised gas shown by Mackie & Fabbiano (1998). The H I column density is below 10^{20} cm^{-2} across the entire cloud at the resolution of our data. We detect H I in 3 velocity channels of our cube (1730–1820 km s^{-1} ; Fig. 1) but find no clear velocity gradient (Fig. 3). We will present a multi-wavelength analysis of the EELR in a future paper.

The H I-bright region SH2, which consists of a ring-shaped collection of 100 Myr-old star clusters ($M_{\star} \sim 10^6 M_{\odot}$; Richtler et al. 2012, 2017), remains unresolved in our H I data both in velocity (at $\sim 1680 \text{ km s}^{-1}$) and on the sky. The angular

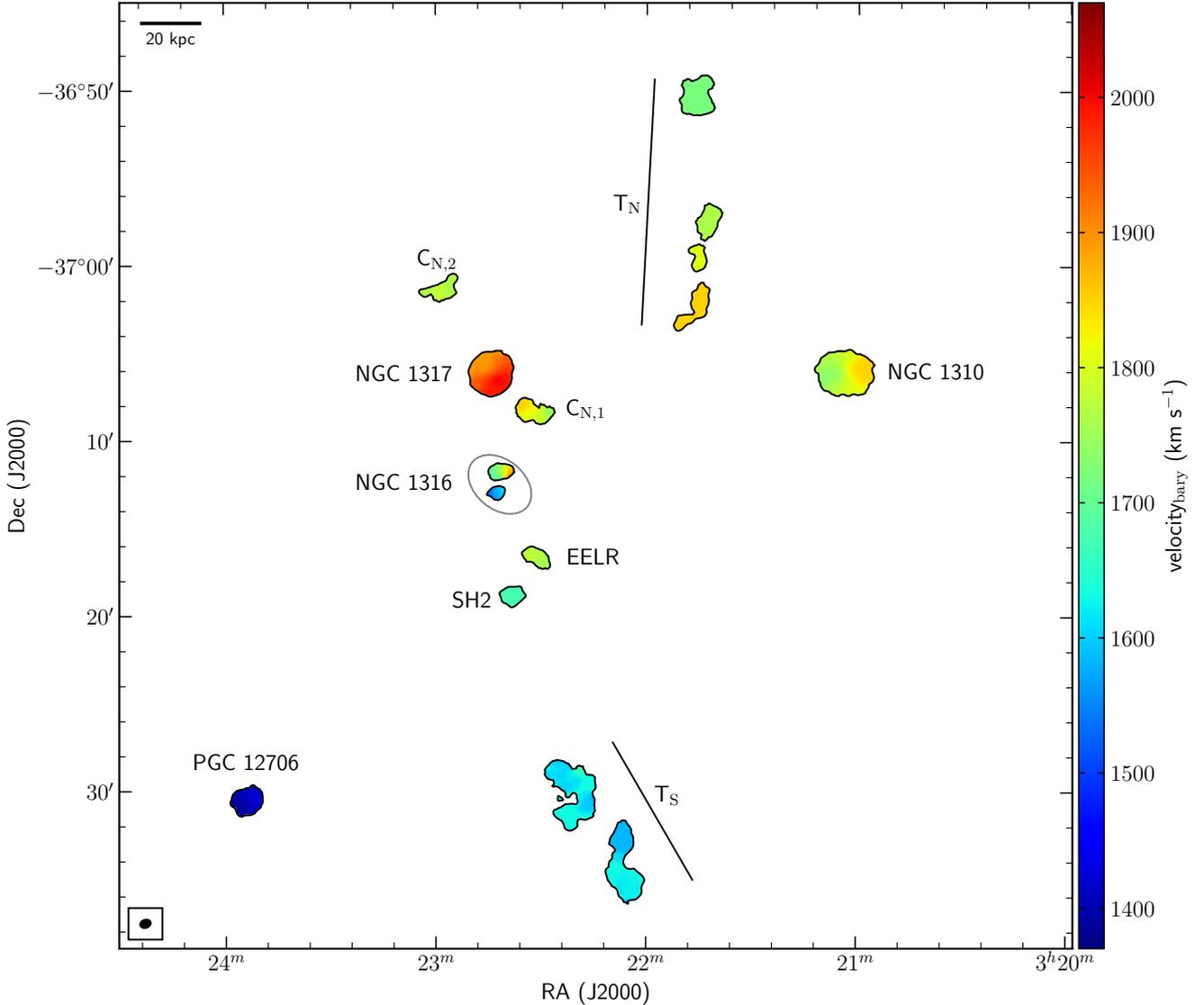


Fig. 3. H I velocity field. Velocities are in the barycentric reference frame (optical definition). The velocity colour bar is centred at the systemic velocity of NGC 1316, 1760 km s^{-1} . The black H I contour is the same as the first contour in Fig. 2 and has a column density of $2.7 \times 10^{19} \text{ cm}^{-2}$. The grey, open ellipse represents the stellar body of NGC 1316. The Gaussian restoring PSF in the bottom-left corner has a FWHM of $36.7'' \times 28.1''$ and position angle 109 deg. The scale bar in the top-left corner represents 20 kpc ($\sim 3.5'$ at a distance of 20 Mpc). The H I properties of the labelled sources are listed in Table 2.

and velocity resolution of our data is insufficient to search for any low-column-density H I extension in the direction of an associated, intermediate-age globular cluster ($\sim 15''$ to the south-west and $\sim 30 \text{ km s}^{-1}$ offset in velocity; Richtler et al. 2017). The H I mass of SH2 is $2.8 \times 10^7 M_{\odot}$, which dominates over the stellar mass by approximately an order of magnitude.

Finally, we find $3.7 \times 10^7 M_{\odot}$ of H I in $C_{N,1}$. This cloud is located at the northern edge of the stellar body of NGC 1316 and is elongated in a direction approximately parallel to the optical isophotes in that region. The H I column density is below 10^{20} cm^{-2} at the resolution of our data. Gas in this cloud is visible in 3 velocity channels of our cube (1770–1860 km s^{-1} ; Fig 1), and we observe a gradient with velocity increasing from west towards east (Fig. 3).

3.3. H I at a large distance from NGC 1316: T_N , T_S and $C_{N,2}$

The vast majority of the H I associated with NGC 1316 is distributed at large radii in the two tail-like complexes T_N and T_S .

At the resolution of our data T_N includes 4 or 5 H I clumps (Fig. 2), which form a coherent 3D structure in the H I cube from 70 to 150 kpc (projected) north-west of NGC 1316 (Figs. 1 and 3). The southernmost clump is redshifted by $\sim 100 \text{ km s}^{-1}$ relative to systemic velocity. The H I velocity then decreases progressively as we move north, and the northernmost clump is blueshifted by $\sim 30 \text{ km s}^{-1}$ relative to systemic velocity. The clumps could simply be the densest parts of an underlying, diffuse distribution of gas stretching approximately north to south. The H I distribution appears long ($\sim 80 \text{ kpc}$ in projection) and narrow ($\leq 6 \text{ kpc}$ in projection; it is mostly unresolved by the H I PSF along its transverse direction). The H I column density remains low, $\leq 10^{20} \text{ cm}^{-2}$, across the entire extent of T_N at the resolution of our data.

The structure of the more massive T_S might not be as simple as that of T_N . The H I is distributed in two main components, the southernmost of which includes in turn two individual, bright H I clumps (Fig. 2). All of the gas in T_S is detected in just two channels of our H I cube, both blueshifted relative to systemic. The

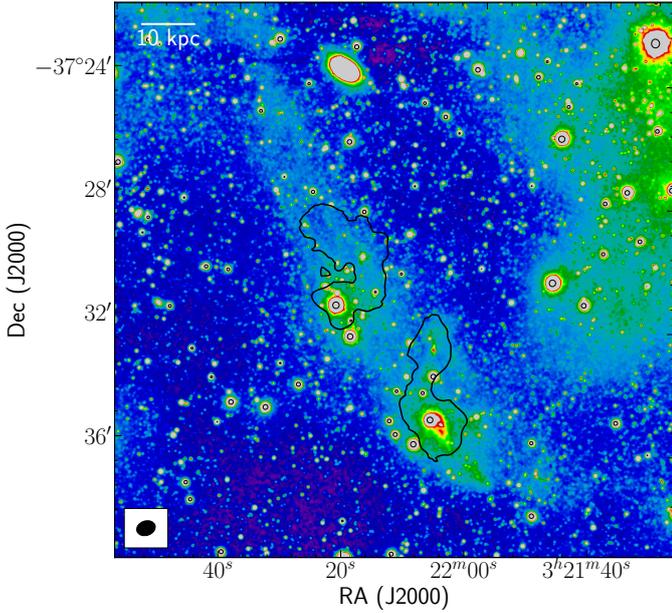


Fig. 4. Relation between the H I tail T_S and the optical tidal feature L5. The image shows a single H I contour at a column density of $2.7 \times 10^{19} \text{ cm}^{-2}$ (black) overlaid on a stack of the g -, r - and i -band images from the Fornax Deep Survey (Iodice et al. 2017) shown in false colour with a linear stretch. The optical image was convolved with a $4''$ Gaussian (FWHM). The black circles indicate bright foreground stars. The H I Gaussian restoring PSF in the bottom-left corner has a FWHM of $36.7'' \times 28.1''$ and position angle 109 deg . The scale bar in the top-left corner represents 10 kpc ($\sim 100''$ at a distance of 20 Mpc).

emission exhibits no clear ordered 3D structure in the H I cube at the velocity resolution of our data (Figs. 1 and 3). Gas in T_S reaches a relatively high column density, nearly 10^{21} cm^{-2} in the two southernmost clumps. Higher angular resolution data would likely bring the peak column density to even higher values.

Finally, $C_{N,2}$ is the only isolated H I cloud detected at a large distance from NGC 1316. It is faint and nearly unresolved. It is located $\sim 70 \text{ kpc}$ north of the galaxy in projection and, like T_N , it is redshifted relative to systemic. This cloud makes up just 5 percent of all the H I detected at a large distance from NGC 1316.

These new H I detections at large radii are associated with the fine structure identified through deep optical imaging by Schweizer (1980) and Iodice et al. (2017)⁴. The prime example is T_S , which is associated with the stellar feature L5. In particular, the overlay of H I contours on top of a gri stacked image (Fig. 4) shows that the northern, more diffuse part of T_S is associated with a faint, diffuse portion of L5; while the two southernmost H I clumps in T_S correspond to a section of L5 which is clumpier and brighter than the rest of the feature. The r -band surface brightness of this bright, southern region reaches $25 \text{ mag arcsec}^{-2}$ vs. the $>26 \text{ mag arcsec}^{-2}$ more typical of L5 as a whole. In the same region we find a high H I column density and, based on a preliminary data reduction of narrow-band imaging obtained recently by our group with the ESO/VST telescope, clumps of $H\alpha + [\text{N II}]$ emission. It is therefore likely that some star formation is on going – or has occurred recently – within them.

Figure 4 also shows that the association between H I and optical light does not hold across the entire L5. First, a bright clump of stellar light is visible $1' - 2'$ south-west of T_S , in a region

where we do not detect any H I at the sensitivity and resolution of our MeerKAT data (see Table 1). This region is not detected in $H\alpha + [\text{N II}]$ either. Furthermore, fainter, diffuse light belonging to L5 is clearly visible for several arcmin north-east of T_S . Here, too, we find no H I.

The other H I detections T_N and $C_{N,2}$ are also associated with known optical features, although more ambiguously so than T_S with L5. First, the southern end of T_N coincides with the south-west end of the L9 feature discovered by Iodice et al. (2017). However, while T_N continues northwards to a region with no optical emission, L9 bends towards east. Similarly, $C_{N,2}$ is located within another optical feature identified by Iodice et al. (2017), L7. However, the optical surface brightness of this feature does not change significantly at the location of the H I cloud.

3.4. H I in galaxies near NGC 1316

Besides NGC 1316, the observed field includes three relatively large, H I-rich galaxies: NGC 1317, NGC 1310 and PGC 12706 (in order of increasing projected distance from NGC 1316; see Fig. 2). These have a $+200$, ~ 0 and -350 km s^{-1} offset relative to the recessional velocity of NGC 1316, respectively. The first two are likely part of the NGC 1316 group. NGC 1310 is slightly in the foreground (Fomalont et al. 1989; Anderson et al. 2018), while the exact 3D location of NGC 1317 is more uncertain. The case of PGC 12706 is less clear as the velocity difference from NGC 1316 is relatively large. This galaxy could be a few Mpc in front of the group.

Our data do not reveal any striking new H I features in NGC 1317 compared with the data of Horellou et al. (2001). Despite our slightly more sensitive data, we confirm their conclusion that the H I disc is confined to within the stellar disc and that the galaxy is H I-deficient, indicating that some gas has likely been removed from its outskirts possibly following the interaction with the intra-group medium. We also confirm the lack of any asymmetries or tidal features in the H I disc. It is therefore likely that the 3D distance between NGC 1317 and the ~ 10 times more massive NGC 1316 is too large for NGC 1317 to be tidally perturbed.

Our data for NGC 1310 are significantly more sensitive than those of Horellou et al. (2001) because of our larger primary beam (the galaxy is $\sim 20'$ from the pointing centre). Unlike in that previous study, we find the H I column density to be symmetric across the disc. Therefore, this galaxy, too, shows no signs of interaction with NGC 1316. At the resolution of our data there appears to be an H I depression in the galaxy centre. Furthermore, the H I disc is more extended than the stellar disc, which is common for galaxies of this late morphological type.

Finally, we detect for the first time H I in PGC 12706. This galaxy has a peculiar optical morphology characterised by a bright, distorted ring of star formation, likely originating from a collision (Madore et al. 2009). The H I is unresolved in our data and, therefore, gives us no clue about the origin of the ring. The H I is very bright, though, and future observations with higher angular and velocity resolution will be useful to understand the nature of this object.

4. The assembly history of NGC 1316

Our detection of $7 \times 10^8 M_\odot$ of H I within and around NGC 1316 – mostly at large radii and partly associated with optical tidal features – provides new constraints on the assembly history of this galaxy. In particular, it gives renewed support to the hypothesis

⁴ See also <https://images.datacentral.org.au/malin/DEEP/060>.

that NGC 1316 formed through a merger between a dominant, gas-poor early-type galaxy and a smaller, gas-rich spiral.

As discussed in Schweizer (1980) and Iodice et al. (2017), NGC 1316 has likely experienced multiple mergers with mass ratio $\leq 10:1$ in the last few billion years. The picture emerging from those studies is that the oldest merger is likely traced by the longest and faintest tidal features – L5 to the south, shown in Fig. 4, and L9 to the north – and has occurred at least 1 Gyr ago. Subsequent mergers might be responsible for the smaller and brighter tidal features visible in Fig. 2, such as those to the east (known as L2) and west (L1). Our detection of H I exactly coincident with L5 and in the proximity of L9 establishes an important new fact: the oldest merger was gas-rich. The lack of H I at the location of the brighter, more recent tidal features suggests that, in comparison, subsequent mergers were gas-poor.

This result suggests a possible causal relation between that first, gas-rich merger and other evidence that cold gas has flowed recently into the centre of NGC 1316, where it has been partly consumed by star formation. The galaxy hosts a multi-phase interstellar medium dominated by molecular gas and kinematically decoupled from the stars in the central few kiloparsecs of the stellar body (Schweizer 1980; Horellou et al. 2001; Morokuma-Matsui et al. 2019; this work). Furthermore, optical and infrared spectroscopy reveal intermediate-age stellar populations, which is best interpreted in terms of a ~ 10 percent (in mass) of ~ 1 Gyr-old stars on top of an older population (Kuntschner 2000; Silva et al. 2008). And, finally, the timescale of this star formation event is comparable to that for the formation of a relatively young (1–3 Gyr old) population of globular clusters, which indeed Goudfrooij et al. (2001) and Sesto et al. (2018) used to date the initial merger.

We therefore propose that the same gas-rich merger that produced stellar and gaseous tidal features at large radii is also responsible for the accretion of gas into the central regions of NGC 1316. During mergers, non-axisymmetries generated by tidal forces can drive some of the gas inward, where it triggers a starburst, while the same tidal forces can pull stars and the remaining gas out to large radii, where they can remain visible for billions of years (e.g., Di Matteo et al. 2007, 2008). In other known merger remnants, such as those discussed in Sect. 1, cold gas at large radius is mostly atomic while, in the central region, most of it is in molecular form and associated with recent star formation. The same is true in NGC 1316, where we find identical masses of neutral hydrogen in the centre (mostly molecular) and at large radii (mostly atomic; Horellou et al. 2001; Morokuma-Matsui et al. 2019).

The association between stellar light and H I in merger remnants is not always tight. Gas-poor optical features and “dark” H I features exist in many systems like those mentioned in Sect. 1. In this sense, the lack of stellar light across most of the H I tail T_N and the lack of H I across, e.g., most of L9 are not too surprising (see Sect. 3.3). The definite association between T_S and L5 is more consequential in the discussion about the origin of NGC 1316 than the lack of association between stellar light and H I at other locations.

The above, qualitative discussion shows that, based on the new H I data and on the large body of previous observational results, many features of NGC 1316 can be traced back to a gas-rich merger occurred at most a few billion years ago – likely followed by subsequent mergers involving smaller amounts of cold gas. We supplement that discussion with some more quantitative arguments on the nature of that initial merger. As a first step we revisit Lanz et al. (2010)’s calculation of the expected H I mass of NGC 1316 in the context of its merger origin. As

discussed in Sect. 1, NGC 1316 hosts >10 times more dust than expected given the typical dust-to-stellar mass ratio of galaxies with its morphology and stellar mass (Smith et al. 2012). We therefore assume that all of the dust currently present within NGC 1316 ($\sim 2 \times 10^7 M_\odot$; Draine et al. 2007; Lanz et al. 2010; Galametz et al. 2012) has an external origin and comes from an Sa progenitor that merged with an early-type galaxy⁵. Galaxies of Sa type have a typical dust-to-stellar mass ratio of $\sim 5 \times 10^{-4}$ (Smith et al. 2012). Therefore, the Sa’s stellar mass was likely to be $\sim 4 \times 10^{10} M_\odot$. The merger that formed NGC 1316, whose stellar mass is now $\sim 6 \times 10^{11} M_\odot$ (Iodice et al. 2017), appears then to have had a mass ratio of 14:1. This would be sufficient to form the long tidal features visible in its optical images such as L5 and L9.

How much H I would the Sa progenitor bring into NGC 1316? Galaxies with this Hubble type and stellar mass have a stellar surface mass density $\mu_\star \sim 10^9 M_\odot \text{ kpc}^{-2}$ (based on the sample of Nair & Abraham 2010). The scaling relation between M_{HI}/M_\star , M_\star and μ_\star (Brown et al. 2015) tells us that the galaxy likely had $M_{\text{HI}}/M_\star \sim 0.05$. We thus expect the Sa progenitor of NGC 1316 to have had $M_{\text{HI}} \sim 2 \times 10^9 M_\odot$. This result is in good agreement with the earlier calculation by Lanz et al. (2010).

Our detection of $7 \times 10^8 M_\odot$ of H I around NGC 1316 increases the H I mass of this galaxy by a factor of 14 compared to previous work, bringing it within a factor of just ~ 3 of the H I mass expected for the progenitor Sa galaxy. This is now an acceptable difference considering: (i) the large scatter around many of the scaling relations used for the above calculation (e.g., scaling relations involving M_{HI}/M_\star have a typical scatter between 0.5 and 1 dex; see Catinella et al. 2018); (ii) the uncertainty on some of the quantities that go into those scaling relations (in this case, dust mass, stellar mass, and stellar mass surface density); (iii) the possibility that the Sa progenitor was H I-deficient following the removal of some of its cold gas through an interaction with the intra-group medium (indeed, as discussed in Sect. 3.4, the nearby galaxy NGC 1310 is H I-deficient); (iv) the possibility that some H I in the field is below our detection limit; and (v) the fact that during a merger some H I will inevitably be lost to star formation in the remnant (as discussed above) and to ionisation once tidal forces diffuse it into the inter-galactic medium.

The above calculation should not be over-interpreted. For example, it is possible that the spiral progenitor was a slightly more massive galaxy with a slightly lower M_{HI}/M_\star ratio. However, the same calculation suggests that a dramatically different merger is unlikely. A merger between two equal-mass galaxies – of which at least one gas-rich – is ruled out as it would involve a much larger gas mass than observed. An additional consistency argument is that galaxies similar to the hypothetical Sa progenitor discussed above are present in the immediate surroundings of NGC 1316, namely NGC 1310 and NGC 1317. It is possible that one such galaxy has merged with a larger early-type galaxy to form NGC 1316.

We further investigate the consistency between our observational results and the expectation for a $\sim 10:1$ early-type + spiral merger remnant using the GalMer database of simulated galaxy mergers (Chilingarian et al. 2010). GalMer includes

⁵ Our choice of progenitors aims at keeping the merger mass ratio below $\sim 10:1$. Given the low cold gas-to-stellar mass ratio of early-type spirals, choosing an Sa galaxy as the smaller progenitor maximises its stellar mass and, therefore, minimises the merger mass ratio at fixed cold-gas mass. Choosing a later-type spiral progenitor would give a mass ratio too large to be consistent with previous results on the merger that generated the extended optical tidal features like L5 (Schweizer 1980; Mackie & Fabbiano 1998; Goudfrooij et al. 2001).

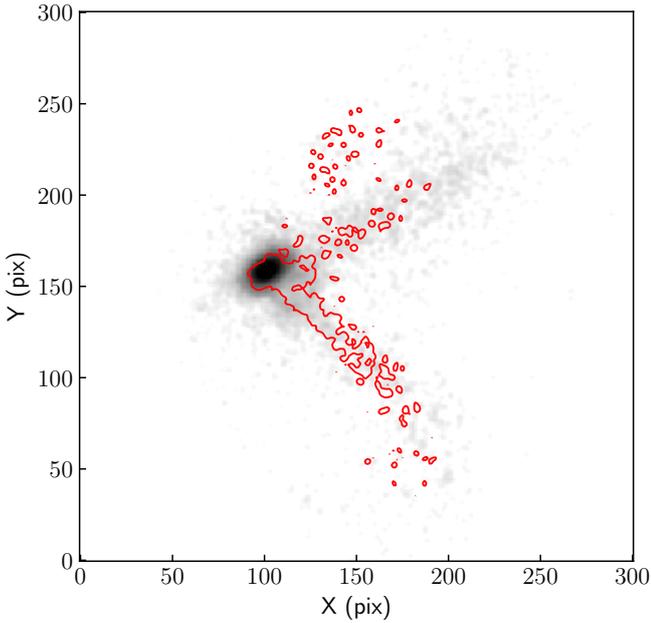


Fig. 5. Snapshot of GalMer merger # 971 (Chilingarian et al. 2010) taken 2.5 Gyr after the start of the simulation and ~ 2 Gyr after the first pericentral passage. The background greyscale image shows the distribution of stellar mass on a logarithmic scale. The red contours show the distribution of gas. Both stellar and gas images have been smoothed for clarity. The image is $750 \times 750 \text{ kpc}^2$, but note that we ignore the physical scale of the system in our comparison to NGC 1316 for reasons explained in the text.

mergers over a range of galaxy morphologies, mass ratios and orbital parameters. We have inspected several S0+Sa 10:1 mergers and found that those on prograde orbits typically result in the presence of low-surface-brightness tidal material (both stars and gas) at large radii due the large total angular momentum of the system – like in NGC 1316.

Given the limited size of the GalMer database it would be unreasonable to expect an exact match to the properties of NGC 1316. For example, the total stellar mass of the S0+Sa 10:1 mergers is ~ 7 times smaller than that of NGC 1316, and the range of orbital parameters explored in the database is limited. Furthermore, the resolution of the GalMer simulations limits a thorough treatment of the complex gas physics, and no feedback from an active galactic nucleus is included. With these caveats in mind, we note however the following important consistencies between NGC 1316 and the GalMer merger # 971 observed ~ 2 Gyr after the first pericentral passage (Fig. 5):

- the merger exhibits multiple large-scale stellar tidal tails which include a gas counterpart for at least part of their extent; this agrees with our detection of H I along part of L5 and at the start of L9; the tails are still expanding at the time of the snapshot in Fig. 5; the exact shape and size of the tails varies with viewing angle (e.g., nearly radial in Fig. 5, more curved/broad on other projections), but overall they are always clearly visible;
- on some projections, the velocities of the two gas tails lay on opposite sides relative to the systemic velocity of the remnant, as in NGC 1316; and the velocity of the gas approaches the systemic velocity when moving towards larger radii along a tail, as in T_N (Fig. 3);
- the remnant hosts approximately equal amounts of gas in the central region and in the outer tails; this is consistent with NGC 1316, where equal amounts of neutral hydrogen

are found in the centre (mostly molecular) and in the tails (mostly atomic; GalMer does not distinguish between these two phases);

- the stellar tidal tails are only ~ 0.1 mag bluer than the stellar body in $g - r$ colour, consistent with the idea that they formed several billion years ago, and with the small colour difference between most tidal features and the main stellar body in NGC 1316 (Iodice et al. 2017);
- the stellar body rotates about its short axis as in NGC 1316 (Bosma et al. 1985; Arnaboldi et al. 1998).

A comparison between NGC 1316 and a set of GalMer mergers was presented also by Iodice et al. (2017). They found that the same 10:1 merger analysed here (#971 in the GalMer database, which they also observed ~ 2 Gyr after the first pericentral passage) provides a good match to the structure of the stellar body of NGC 1316 and to the surface brightness of its outer tidal features L5 and L9 – while a 1:1 merger between two disc galaxies does not. In that paper, the 10:1 merger is favoured also because it has no gaseous tails, consistent with the H I results available at that time (Horellou et al. 2001). We now detect extended H I tails in NGC 1316 but, as argued above, this does not rule out a 10:1 merger. In fact, Fig. 5 shows that the 10:1 merger remnant, too, hosts extended gas tails. Compared to the analysis in Iodice et al. (2017) the gas tails are now visible because we smooth the simulations, enhancing low-surface-brightness features, and we choose a projection which makes them stand out. This projection does not match the orientation of NGC 1316’s stellar body (unlike in Iodice et al. 2017) but this is not of real concern since there are no reasons to expect an exact match between the galaxy and the simulation.

There are also significant differences between GalMer merger # 971 and NGC 1316. For example, gas in its centre is kinematically aligned with the stars (unlike in NGC 1316; Schweizer 1980; Morokuma-Matsui et al. 2019); and the length of its tidal tails exceeds the observed one. However, considering the limitations mentioned above and the complex recent history of NGC 1316, which includes episodes of nuclear activity (e.g., Ekers et al. 1983; Fomalont et al. 1989) and further minor mergers (Schweizer 1980; Iodice et al. 2017), this comparison should only be viewed as a consistency check. The main message is that some of the first-order features of NGC 1316 can be reproduced by an S0+Sa $\sim 10:1$ prograde merger, which, as argued by Lanz et al. (2010) and confirmed above in this section, is required to explain the overall composition of the galaxy’s interstellar medium.

To conclude, it would be unreasonable to expect that a single merger can explain all the properties of NGC 1316. As thoroughly discussed by Schweizer (1980) and Iodice et al. (2017), NGC 1316 is likely continuing to accrete satellites. These might explain, for example, some of the minor H I clouds in Fig. 2. However, our results of a similar amount of neutral gas in the inner regions of NGC 1316 (mostly molecular) as in the tails at large radii (atomic) associated with optical tidal features, together with the evidence of central star formation in the last few Gyr, are consistent with the merger hypothesis. A quantitative comparison between NGC 1316 and simulations is challenging, but the picture emerging from this diverse set of observations is becoming increasingly clearer.

5. Summary

We have observed the nearby merger remnant NGC 1316 at the outskirts of the Fornax galaxy cluster with MeerKAT in search of H I associated with its prominent stellar tidal features. We detect

a total of $7 \times 10^8 M_{\odot}$ of H I within and around this galaxy, 14 times more than in previous observations. Most of the H I is distributed in two tail-like complexes at large radii, 70–150 kpc in projection. The southern tail is associated with the optical tidal feature known as L5. It hosts clumps of dense H I ($\sim 10^{21} \text{ cm}^{-2}$) and exhibits signs of recent star formation. The northern tail is only partly associated with a known optical feature and is characterised by a low H I column density ($\lesssim 10^{20} \text{ cm}^{-2}$). In addition to several scattered H I clouds at various distance from NGC 1316, we also detect for the first time H I in the centre of the galaxy: $4 \times 10^7 M_{\odot}$ closely associated with dust and molecular gas in the central ~ 5 kpc.

These results shed new light on the assembly history of NGC 1316. The total H I mass and the location of most of the H I at large radii, near known stellar tidal features, together with the presence of a similar mass of gas in a molecular form in the galaxy centre, where a starburst has recently occurred, strengthen the hypothesis that NGC 1316 formed through a gas-rich merger a few billion years ago, followed by subsequent, gas-poorer accretion of satellite galaxies. To first order, the properties of NGC 1316 are consistent with this initial merger having occurred between a dominant, gas-poor early-type galaxy and a ~ 10 times smaller gas-rich spiral.

Future MeerKAT observations of this field conducted as part of the MeerKAT Fornax Survey (Serra et al. 2016) will achieve significantly better sensitivity as well as angular and velocity resolution than those presented here. We will reach a similar H I column density sensitivity ($\sim 3 \times 10^{19} \text{ cm}^{-2}$) at ~ 3 times better angular resolution ($10''$) and with >40 times narrower velocity channels ($\sim 1 \text{ km s}^{-1}$); while at the $30''$ angular resolution of the current data we will detect H I down to a column density of a few times 10^{18} cm^{-2} . The future data will thus unveil additional details of the distribution of H I within and around NGC 1316.

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