

# Neutral gas properties of Lyman continuum emitting galaxies: Column densities and covering fractions from UV absorption lines

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## ABSTRACT

*Context.* The processes allowing the escape of ionizing photons from galaxies into the intergalactic medium are poorly known.

*Aims.* To understand how Lyman continuum (LyC) photons escape galaxies, we constrain the H I covering fractions and column densities using ultraviolet (UV) H I and metal absorption lines of 18 star-forming galaxies that have Lyman series observations. Nine of these galaxies are confirmed LyC emitters.

*Methods.* We fit the stellar continuum, dust attenuation, metal, and H I properties to consistently determine the UV attenuation, as well as the column densities and covering factors of neutral hydrogen and metals. We used synthetic interstellar absorption lines to explore the systematics of our measurements. Then we applied our method to the observed UV spectra of low-redshift and  $z \sim 3$  galaxies.

*Results.* The observed H I lines are found to be saturated in all galaxies. An indirect approach using O I column densities and the observed O/H abundances yields H I column densities of  $\log(N_{\text{H I}}) \sim 18.6\text{--}20 \text{ cm}^{-2}$ . These columns are too high to allow the escape of ionizing photons. We find that the known LyC leakers have H I covering fractions less than unity. Ionizing photons escape through optically thin channels in a clumpy interstellar medium. Our simulations confirm that the H I covering fractions are accurately recovered. The Si II and H I covering fractions scale linearly, in agreement with observations from stacked Lyman break galaxy spectra at  $z \sim 3$ . Thus, with an empirical correction, the Si II absorption lines can also be used to determine the H I coverage. Finally, we show that a consistent fitting of dust attenuation, continuum, and absorption lines is required to properly infer the covering fraction of neutral gas and subsequently to infer the escape fraction of ionizing radiation.

*Conclusions.* These measurements can estimate the LyC escape fraction, as we demonstrate in a companion paper.

**Key words.** galaxies: ISM – ISM: abundances – ISM: lines and bands – ultraviolet: ISM – dust, extinction – dark ages, reionization, first stars

## 1. Introduction

Star-forming galaxies are ideal laboratories to understand how the early universe became reionized. Galaxies likely reionized the universe because quasars are too rare at high redshifts (Fontanot et al. 2012, 2014). Compact galaxies with intense star formation rates produce large amounts of ionizing photons which, under certain circumstances, escape the interstellar medium (ISM) and ionize the intergalactic medium (IGM). To reionize the universe, studies suggest that 10–20% of the ionizing photons produced by star-forming galaxies must escape galaxies (Ouchi et al. 2009; Robertson et al. 2013; Dressler et al. 2015). However, it has been challenging to detect Lyman continuum (LyC) radiation from individual galaxies.

Three unambiguous observations of ionizing photons have been reported at  $z \sim 3$  (Vanzella et al. 2015; de Barros et al. 2016; Shapley et al. 2016; Bian et al. 2017). The observed escape fractions (near 50 %) are greater than those required to reionize the universe. Additionally, there are nine low-redshift ( $z < 0.3$ )

galaxies with  $1\% \leq f_{\text{esc}}^{\text{LyC}} \leq 13\%$  (Bergvall et al. 2006; Leitert et al. 2013; Borthakur et al. 2014; Leitherer et al. 2016; Izotov et al. 2016a,b; Puschnig et al. 2017), and one recent detection at  $z = 0.37$  with  $f_{\text{esc}}^{\text{LyC}} \approx 46\%$  (Izotov et al. 2018). The low number of detections emphasizes the difficulty of detecting Lyman continuum emitters (LCEs).

Zackrisson et al. (2013) proposed two theoretical models to explain how ionizing photons escape galaxies. In the first scenario, low H I column densities allow Lyman continuum photons to pass through without being completely absorbed; this is called the density-bounded scenario (Jaskot & Oey 2013; Nakajima & Ouchi 2014). This scenario manifests itself as low H I column densities ( $< 10^{18} \text{ cm}^{-2}$ ). In the second scenario, ionizing photons leak into the IGM through holes in the neutral gas (Heckman et al. 2011). This scenario is called the picket-fence model and relies on a patchy neutral gas. A patchy neutral ISM manifests as H I absorption lines with a covering fraction less than one. It is unclear which of these scenarios describes

how ionizing photons leak from galaxies. Constraining neutral gas properties, especially the H I covering fraction and column density, is an effective way to disentangle how ionizing photons escape galaxies.

H I absorption lines in the rest-frame far ultraviolet (the Lyman series: 912–1026 Å) directly probe the H I covering fraction and column density. However, the Lyman series is challenging to observe for several reasons. First, this requires deep rest-frame far-ultraviolet observations blueward of Ly $\alpha$ , which is notoriously difficult to observe at low redshifts. Second, the Lyman series is unavailable at redshift  $z > 3$  because the Ly $\alpha$  forest completely absorbs this region. Third, several observational obstacles need to be accounted for to measure the H I properties. In particular, foreground contamination (Vanzella et al. 2010), intervening absorbers, stellar continuum from the galaxy itself, ISM absorption lines, and, at low redshifts, Milky Way and geocoronal emission need to be identified.

Interstellar medium metal absorption lines (i.e., Si II 1260 Å, C II 1334 Å) are easier to observe than the Lyman series (Heckman et al. 2011; Alexandroff et al. 2015). However, using metal absorption lines to trace the H I assumes that metals directly probe the neutral gas. Recent studies indicate that ISM metal lines may have a factor of 2 times smaller covering fractions than H I absorption lines (Reddy et al. 2016b). As a result, metal absorption lines may not trace the H I.

In this article, we directly observe the Lyman series of individual high and low-redshift star-forming galaxies to determine their neutral gas properties, and to compare H I measurements to ISM metal properties. For the first time, we measure the H I properties of spectroscopically confirmed LyC emitters to determine which physical process enables ionizing photons to escape galaxies. A companion paper uses these observed H I properties to predict the escape fractions of ionizing photons (Chisholm et al. 2018, hereafter Paper II).

This paper is organized as follows: Sect. 2 describes the observational data. Section 3 defines the various models and equations used to fit the stellar continua and UV absorption lines. In Sect. 4 we use synthetic spectra to illustrate how accurately we recover H I column densities and covering fractions from observations. Section 5 discusses the measured H I covering fractions of the LyC emitters, the relation between the H I and Si II covering fractions, the effects of the assumed dust geometry, and comparisons to previous studies. We summarize our results in Sect. 6.

## 2. Observed data

We studied the neutral gas properties of a sample of 18 star-forming galaxies listed in Table 1. Our selection was driven by the need to observe the Lyman series, i.e., available rest-frame UV spectroscopy between Lyman- $\beta$  and the Lyman limit. We selected the low-redshift galaxies observed with the Cosmic Origins Spectrograph (COS) on the *Hubble* Space Telescope (HST; Green et al. 2012) with this wavelength coverage. Given the sensitivity and wavelength range of the G130M grating on COS, the Lyman series is observable with a spectral resolution  $R \gtrsim 1500$  for galaxies at  $z > 0.18$ . Therefore, we selected the 15 galaxies at low redshifts ( $z < 0.36$ ) with Lyman series observations from the COS archive. Each galaxy is at a redshift such that at least the Ly $\beta$  line is observable with the COS G140L or G130M gratings.

Many of our low-redshift galaxies were originally targeted to observe possible LyC emission, but only nine of them are

**Table 1.** Sample of galaxies with Lyman series observations.

Galaxy name (1)	$z$ (2)	$12 + \log(\text{O}/\text{H})$ (3)	$f_{\text{esc}}^{\text{LyC}}$ (4)	$R$ (5)
J0921+4509	0.23499	8.67 <sup>a</sup>	0.010 <sup>g</sup>	15 000
J1503+3644	0.3537	7.95 <sup>b</sup>	0.058 <sup>b</sup>	1500
J0925+1409	0.3013	7.91 <sup>c</sup>	0.072 <sup>c</sup>	1500
J1152+3400	0.3419	8.00 <sup>b</sup>	0.132 <sup>b</sup>	1500
J1333+6246	0.3181	7.76 <sup>b</sup>	0.056 <sup>b</sup>	1500
J1442–0209	0.2937	7.93 <sup>b</sup>	0.074 <sup>b</sup>	1500
Tol1247–232	0.0482	8.10 <sup>d</sup>	0.004 <sup>h</sup>	1500
Tol0440–381	0.0410	8.20 <sup>d</sup>	0.019 <sup>h</sup>	1500
Mrk54	0.0451	8.60 <sup>d</sup>	<0.002 <sup>h</sup>	1500
J0926+4427	0.18069	8.01 <sup>e</sup>	–	15 000
J1429+0643	0.1736	8.20 <sup>e</sup>	–	15 000
GP0303–0759	0.16488	7.86 <sup>e</sup>	–	15 000
GP1244+0216	0.23942	8.17 <sup>e</sup>	–	15 000
GP1054+5238	0.25264	8.10 <sup>e</sup>	–	15 000
GP0911+1831	0.26223	8.00 <sup>e</sup>	–	15 000
SGAS J1226	2.92525	–	–	4000
SGAS J1527	2.76228	<8.5 <sup>i</sup>	–	2700
Cosmic Eye	3.07483	8.60 <sup>f</sup>	–	2500

**Notes.** Column 1: Galaxy name; Col. 2: redshift; Col. 3: metallicities derived from optical emission lines; Col. 4: Lyman continuum escape fraction; and Col. 5: spectral resolution of the observations. Dashes indicate that the quantities have not been measured.

**Reference.** <sup>(a)</sup>Pettini & Pagel (2004); <sup>(b)</sup>Izotov et al. (2016b); <sup>(c)</sup>Izotov et al. (2016a); <sup>(d)</sup>Leitherer et al. (2016); <sup>(e)</sup>Izotov et al. (2011); <sup>(f)</sup>Stark et al. (2008); <sup>(g)</sup>Borthakur et al. (2014); <sup>(h)</sup>Chisholm et al. (2017a); <sup>(i)</sup>Wuyts et al. (2012).

confirmed LyC emitting galaxies. Four galaxies are from the Green Pea sample of Henry et al. (2015) and two are Lyman break analogs from Heckman et al. (2011). The other nine galaxies are known LyC emitters (J1503+3644, J0925+1409, J1152+3400, J1333+6246, J1442–0209, J0921+4509, Tol1247–232, Tol0440–381, and Mrk 54 from Izotov et al. 2016a,b; Borthakur et al. 2014; Leitherer et al. 2016). The nine leakers with COS/HST data were reduced using CALCOS v2.21 and a custom method for faint COS spectra (Worseck et al. 2016). The other COS/HST data were reduced with CALCOS v2.20.1 and the methods from Wakker et al. (2015). The five Izotov et al. (2016a,b) spectra were smoothed with a 5 pixel boxcar.

Finally, we also included three gravitationally lensed galaxies at  $z \approx 3$  (Stark et al. 2008; Koester et al. 2010). These galaxies are part of the *Magellan* Evolution of Galaxies Spectroscopic and Ultraviolet Reference Atlas (MEGASaURA; Rigby et al. 2018) and were selected because they are the only galaxies in the MEGASaURA sample with a signal-to-noise ratio (S/N) greater than 2 near the Lyman series. These are moderate resolution ( $R \sim 3000$ ) spectra observed with the MagE spectrograph on the *Magellan* Telescopes (Marshall et al. 2008). Instead of the full galaxy names, we used the following short names for the two sources: SGAS J1226 = SGAS J122651.3+215220 and SGAS J1527 = SGAS J152745.1+065219.

The complete sample is summarized in Table 1, where we list a few host properties and (nominal) spectral resolutions of the observations. An upper limit on the [N II]/H $\alpha$  ratio constrains the metallicity for SGAS J1527 ( $12 + \log(\text{O}/\text{H}) < 8.5$ ; Wuyts et al. 2012); for SGAS J1226 these lines are not accessible from the

ground. Meanwhile, Stark et al. (2008) measured the metallicity of the Cosmic Eye using the  $R_{23}$  index.

### 3. Ultraviolet spectral fitting methods and results

We now describe the theory and methods we used to fit the UV spectra including stellar continua, ISM absorption lines, and Milky Way absorption lines. The method is then applied to simulated data (in particular to determine systematic errors) and to the observed spectra.

#### 3.1. Ultraviolet continuum and interstellar absorption line modeling

##### 3.1.1. Adopted geometries and basic formulas

To describe the radiation transfer in the host galaxy, we adopted the classical picket-fence model with different assumptions on the geometric distribution of gas and dust. In practice we considered two cases: (a) the picket-fence model with a uniform foreground dust screen, and (b) a clumpy picket-fence model where dust is only in the neutral gas clumps. Two parameters describe the two models: the dust attenuation (here parametrized by  $E_{B-V}$ ) and the geometric covering fraction of neutral gas ( $C_f$ ), which is defined as the fraction of the total lines of sight of the emitted UV radiation that are intercepted by neutral gas in the direction toward the distant observer.

In (a) both the radiation emerging from the gas clumps (with geometric coverage  $C_f$ ) and radiation directly escaping ( $1 - C_f$ ) is attenuated by a uniform foreground dust screen. In (b), only a fraction,  $C_f$ , of radiation is processed through the gas clumps, imprinting interstellar absorption lines and attenuating the stellar continuum. The rest escapes unaltered. Gas within these clumps is assumed to be homogeneous and the interclump medium has a negligible column density of neutral gas, i.e., is assumed to be completely transparent.

These simple models have already been examined and assumed by other authors (e.g., Zackrisson et al. 2013; Borthakur et al. 2014; Vasei et al. 2016; Reddy et al. 2016b). We, therefore, only briefly list the main equations used in our spectral modeling. For a picket-fence model with a uniform foreground dust screen (a) the emergent flux  $F_\lambda$  is

$$F_\lambda = F_\lambda^* \times 10^{-0.4k_\lambda E_{B-V}} \times (C_f \exp(-\tau_\lambda) + (1 - C_f)), \quad (1)$$

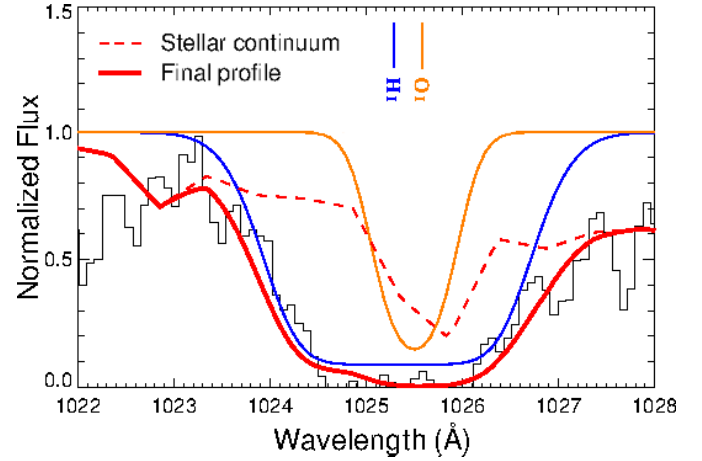
where  $F_\lambda^*$  is the intrinsic stellar emission prior to alteration by the ISM,  $k_\lambda$  describes the attenuation law, and  $\tau_\lambda$  is the optical depth of the interstellar absorption lines. For a picket-fence model, with a clumpy gas distribution (b), the emergent flux becomes

$$F_\lambda = F_\lambda^* \times 10^{-0.4k_\lambda E_{B-V}} \times C_f \exp(-\tau_\lambda) + F_\lambda^* \times (1 - C_f), \quad (2)$$

where the second term describes the unattenuated, directly escaping radiation. This light is unattenuated because the holes are assumed to be free of gas and dust. For high covering fractions ( $C_f \rightarrow 1$ ) or low attenuations ( $E_{B-V} \rightarrow 0$ ), Eq. (1) and Eq. (2) are identical.

We define the residual flux,  $R$ , as the ratio of the flux density at the observed wavelength of the line to the continuum flux density. The  $R$  gives the fraction of light unabsorbed by the neutral gas. For saturated lines ( $\tau_\lambda \gg 1$ ) the residual flux becomes

$$R = 1 - C_f, \quad (3)$$



**Fig. 1.** Total fit (red solid line) of the Ly $\beta$  absorption line for the galaxy GP 1244+0216 (black data). The contributions from the stellar continuum (dashed red line), Ly $\beta$  absorption line (blue solid line), and O I absorption line (orange solid line) blend together. O I absorption lines at 989 and 1039 Å are unblended and robustly constrain the O I profile. All three components need to be accounted for when fitting the Lyman series.

for a uniform dust screen (a), and

$$R = \frac{(1 - C_f)}{10^{-0.4k_\lambda E_{B-V}} C_f + (1 - C_f)}, \quad (4)$$

for case (b) (Vasei et al. 2016). We note that  $R$  of the uniform screen model is always greater than, or equal to, the  $R$  of the clumpy model (see the discussion in Sect 5.3).

Finally, the absolute escape fraction of radiation for monochromatic radiation,  $f_{\text{esc}} = F_\lambda / F_\lambda^*$ , becomes

$$f_{\text{esc}} = 10^{-0.4k_\lambda E_{B-V}} \times (C_f \exp(-\tau_\lambda) + (1 - C_f)), \quad (5)$$

for the uniform dust screen (a), and

$$f_{\text{esc}} = 10^{-0.4k_\lambda E_{B-V}} \times C_f \exp(-\tau_\lambda) + (1 - C_f), \quad (6)$$

in clumpy geometry (b). It is important to note that the values of  $C_f$  and  $E_{B-V}$  differ a priori between the two model geometries, except for the case of high covering fractions ( $C_f \rightarrow 1$ ) or low attenuation ( $E_{B-V} \rightarrow 0$ ). They must be determined consistently from fits to the observed data, adopting the equations corresponding to the assumed geometry.

##### 3.1.2. Fitting method

The main spectral region modeled here is the Lyman series, H I absorption lines from Ly $\beta$  to the Lyman break ( $\approx 912$ – $1050$  Å). In practice, owing to the lower S/N close to the Lyman break, the bluest Lyman line that we include is Ly6 (930 Å). Figure 1 emphasizes the complicated nature of the reddest Lyman series line, Ly $\beta$ : strong stellar continuum features (red dashed line) blend with the broad H I interstellar absorption (blue line) and the weaker O I interstellar absorption line (orange line). Consequently, bluer Lyman series lines (especially Ly $\gamma$ ) have fewer complications because of their simpler stellar continua.

We modeled ISM metal absorption lines from O I, O VI, Si II, C II, and C III (listed in Table 2). O I has a similar ionization

**Table 2.** Fitted absorption lines.

Ion	$\lambda_{\text{rest}} [\text{\AA}]$
H I	920.947 <sup>a</sup>
	923.148 <sup>a</sup>
	926.249 <sup>a</sup>
	930.751
	937.814
	949.742
	972.517
	1025.728
O I	924.950 <sup>a</sup>
	929.517
	930.257
	936.629
	948.686
	950.885
	971.738
	976.448
	988.578
	988.655
O VI	988.773
	1025.762
	1039.230
C II	1302.168 <sup>b</sup>
	1031.912
C III	1037.613
	1036.337
Si II	977.030
	989.870
	1020.70
	1190.42 <sup>c</sup>
	1193.28 <sup>c</sup>
	1260.42 <sup>c</sup>

**Notes.** Wavelengths are in vacuum. <sup>(a)</sup>Used only for generating synthetic spectra, see Sect. A. <sup>(b)</sup>Used for O I measurements in Tol0440–381 and Mrk54 spectra, see Sect. 3.3.2. <sup>(c)</sup>Used only to measure Si II covering fraction, see Sect. 3.3.3.

structure as H I, such that O I absorption lines blend with all H I lines, except for O I 989 and 1039 Å. These two lines constrain the O I profile.

First, we fit the stellar continuum. We started with an initial stellar model using a linear combination of ten single-age stellar continuum models with ages of 1, 2, 3, 4, 5, 8, 10, 15, 20, and 40 Myr. We also used stellar continuum metallicities of 0.05, 0.2, 0.4, 1, or 2  $Z_{\odot}$ . These spectra were drawn from the fully theoretical STARBURST99 library (S99; Leitherer et al. 1999), computed with the WM-Basic method (Leitherer et al. 2010), and have a spectral resolution  $R(\text{S99}) \approx 2500$ . We chose the stellar continuum metallicity closest to the ISM metallicity (Table 1). The STARBURST99 models use a Kroupa initial mass function with a high (low) mass exponent of 2.3 (1.3), a high-mass cutoff at 100  $M_{\odot}$ , and the stellar evolution tracks with high mass loss from

Meynet et al. (1994). We fit for a linear combination of the stellar continuum flux,  $F_{\text{S99}}$ , as

$$F^* = \sum_{i=1}^{10} X_i F_i^{99}, \quad (7)$$

where  $X_i$  are the linear coefficients for a given age ( $i$ ) and  $F_i^{99}$  are the STARBURST99 theoretical stellar continuum models for a given age.

Absorption lines of different ions are added using Voigt profiles defined by the velocity shift  $v$ ,  $b$ -parameter, column density  $N$ , and  $C_f$ . The metal covering fraction is initially fixed to 1 to reduce the number of free parameters. We included the O I absorption lines that are directly blended with the Lyman series. Each ion is considered independent, as are its parameters. For each galaxy, we tested whether including the remaining metal lines listed in Table 2 improves the Lyman series fits. If, by eye, they did not improve the fits, we did not include the lines. Finally, the STARBURST99 models are convolved with the nominal spectral resolution of the observations.

We accounted for dust attenuation using the attenuation law from Reddy et al. (2016a), a uniform dust screen model, and fitting for the dust attenuation parameter ( $E_{\text{B-V}}$ ; similar to Chisholm et al. 2015). The linear combination of stellar continuum models, interstellar absorption lines, and dust attenuation produces the final fitted spectrum.

We fit the data using an IDL routine based on the nonlinear least squares method, MPFIT (Markwardt 2009). The MPFIT method returns the best fit and errors for  $E_{\text{B-V}}$ ,  $b$ ,  $v$ ,  $N$ , and  $C_f$  of each ion. For observed spectra, the first step consists of masking the ISM absorption lines and the contaminating Milky Way absorption lines, geocoronal emission, and intervening absorbers. We applied these masks on the data and fit for the linear combination of dust-attenuated STARBURST99 models. In a second step, we fixed the stellar continuum and fit for the ISM absorption lines and Milky Way absorption lines adjacent to the ISM lines. We fit for all of the observed Lyman series lines up to Ly6, provided that they are not near geocoronal emission or intervening absorbers, and do not have a S/N below one. Since the simulated data do not contain the extra complications of adjacent Milky Way and geocoronal lines, all parameters are simultaneously fit in one step.

### 3.2. Fitting simulated data

We tested our fitting method with both noise-free and noisy synthetic data to determine how well the estimated parameter errors represent the actual errors. Additionally, we tested how these errors depend on the S/N and spectral resolution. In Appendix A we fully describe the generation of synthetic spectra, but here we summarize the steps. The synthetic spectra were produced for two scenarios: one assuming the picket-fence model ( $\log(N_{\text{H I}}[\text{cm}^{-2}]) = 20$ ,  $C_f = 0.9$ ) and one describing a uniform ISM in the density-bounded regime ( $\log(N_{\text{H I}}[\text{cm}^{-2}]) = 17.57$ ,  $C_f = 1$ ). Both scenarios correspond to an escape fraction  $f_{\text{esc}} = 0.1$ , regardless of the chosen dust distribution, since we set  $E_{\text{B-V}} = 0$  (Eqs. (5) and (6)).

We created synthetic spectra using the parameters of the picket-fence and density-bounded regime (see Table A.1) for 7 different spectral resolutions between  $R = 600$ –15 000. For each of these 14 set-ups (7 spectral resolutions and two scenarios) we created 50 different realizations by adding various sets of random Gaussian noise. The level of Gaussian noise was chosen to produce a final S/N per pixel between 2–50, in 7 total S/N steps. In total we created 98 different configurations (2 different

**Table 3.** Derived H I properties from the Lyman series absorption lines.

Galaxy name	$E_{B-V}$	$\log(N_{\text{H I}})$ [ $\log(\text{cm}^{-2})$ ]	$b$ [ $\text{km s}^{-1}$ ]	$v$ [ $\text{km s}^{-1}$ ]	$C_f$ fits	$C_f$ depth	$C_f$ final
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)
J0921+4509	$0.222 \pm 0.015$	$17.94 \pm 3.22$	$95 \pm 49$	$-56 \pm 13$	$0.754 \pm 0.111$	$0.769 \pm 0.116$	$0.761 \pm 0.080$
J1503+3644	$0.274 \pm 0.014$	$20.50 \pm 2.33$	$91 \pm 24$	$-127 \pm 12$	$0.634 \pm 0.109$	$0.754 \pm 0.062$	$0.725 \pm 0.054$
J0925+1409	$0.164 \pm 0.015$	$17.81 \pm 3.26$	$81 \pm 73$	$-214 \pm 151$	$0.652 \pm 0.218$	$0.635 \pm 0.094$	$0.638 \pm 0.086$
J1152+3400	$0.134 \pm 0.022$	$20.94 \pm 2.70$	$187 \pm 48$	$-468 \pm 39$	$0.506 \pm 0.067$	$0.619 \pm 0.088$	$0.548 \pm 0.053$
J1333+6246	$0.151 \pm 0.043$	$21.14 \pm 1.46$	$102 \pm 19$	$-126 \pm 48$	$0.731 \pm 0.122$	$0.826 \pm 0.066$	$0.804 \pm 0.058$
J1442-0209	$0.140 \pm 0.015$	$19.60 \pm 3.53$	$123 \pm 183$	$-261 \pm 34$	$0.621 \pm 0.120$	$0.549 \pm 0.040$	$0.556 \pm 0.038$
Tol1247-232	$0.156 \pm 0.010$	$21.89 \pm 3.09$	$66 \pm 87$	$260 \pm 31$	$0.587 \pm 0.061$	$0.690 \pm 0.084$	$0.623 \pm 0.049$
Tol0440-381	$0.271 \pm 0.028$	–	–	–	–	$0.570 \pm 0.084$	$0.570 \pm 0.084$
Mrk54	$0.359 \pm 0.013$	–	–	–	–	$0.504 \pm 0.077$	$0.504 \pm 0.077$
J0926+4427	$0.114 \pm 0.010$	$16.39 \pm 0.40$	$213 \pm 23$	$-199 \pm 12$	$0.723 \pm 0.048$	$0.814 \pm 0.048$	$0.768 \pm 0.034$
J1429+0643	$0.108 \pm 0.015$	$16.17 \pm 0.50$	$236 \pm 56$	$-241 \pm 36$	$0.897 \pm 0.071$	$0.960 \pm 0.061$	$0.930 \pm 0.046$
GP0303-0759	$0.121 \pm 0.045$	$16.07 \pm 1.50$	$192 \pm 99$	$-266 \pm 92$	$0.908 \pm 0.207$	–	$0.908 \pm 0.207$
GP1244+0216	$0.290 \pm 0.043$	$16.34 \pm 0.80$	$220 \pm 79$	$-78 \pm 48$	$0.909 \pm 0.357$	$0.950 \pm 0.131$	$0.946 \pm 0.123$
GP1054+5238	$0.204 \pm 0.044$	$19.60 \pm 3.81$	$164 \pm 78$	$-166 \pm 29$	$0.702 \pm 0.131$	$0.889 \pm 0.158$	$0.778 \pm 0.101$
GP0911+1831	$0.352 \pm 0.038$	$16.76 \pm 1.19$	$188 \pm 43$	$-273 \pm 40$	$0.731 \pm 0.150$	$0.765 \pm 0.116$	$0.752 \pm 0.092$
SGAS J1226	$0.201 \pm 0.001$	$17.16 \pm 1.06$	$380 \pm 24$	$-264 \pm 21$	$0.932 \pm 0.038$	$0.998 \pm 0.009$	$0.994 \pm 0.009$
SGAS J1527	$0.370 \pm 0.002$	$17.15 \pm 1.22$	$269 \pm 47$	$-208 \pm 25$	$1.000 \pm 0.075$	$0.990 \pm 0.038$	$0.992 \pm 0.034$
Cosmic Eye	$0.405 \pm 0.001$	$22.59 \pm 1.28$	$199 \pm 5$	$56 \pm 16$	$0.918 \pm 0.072$	$0.998 \pm 0.024$	$0.990 \pm 0.023$

**Notes.** Column 1: Galaxy name; Col. 2: dust attenuation parameter ( $E_{B-V}$ ); Col. 3: logarithm of the H I column density; Col. 4: H I Doppler  $b$ -parameter; Col. 5: H I velocity shift; and Col. 6: H I covering fraction from the fits. Uncertainties of the H I column density and covering fraction from the fits include the systematic errors detailed in Sect 4.2. Column 7: H I covering fraction measurements derived from the residual flux of the individual H I absorption lines (Sect. 3.3.1; Table C.3). GP0303-0759 does not have a reliable depth measurement because Milky Way absorption lines overlap the Lyman series; Col. 8: weighted mean of Cols. 6 and 7.

scenarios, 7 spectral resolutions, and 7 S/Ns), and each configuration has 50 individual spectra, for a total of 4900 synthetic spectra. Figures B.1 and B.2 show synthetic spectra and the best-fit models for high and low spectral resolution ( $R = 15\,000$  and  $R = 600$ ) for three S/N configurations (noise-free, 10 and 2).

We fit the synthetic spectra with the methods outlined in Sect. 3.1.2. From this fitting we define two types of error:

*The statistical error ( $Err_{\text{stat}}$ ).* This is the individual errors returned by MPFIT.

*The systematic error ( $Err_{\text{syst}}$ ).* This is the deviation of the 50 parameter estimates of each scenario from the actual parameter value that created the original line profile. The  $Err_{\text{syst}}$  is therefore a function of S/N and resolution.  $Err_{\text{syst}}$  is calculated as

$$Err_{\text{syst}} = \sqrt{\frac{1}{50} \sum_{i=1}^{50} (x_i - x_T)^2}, \quad (8)$$

where  $x_i$  is the estimated parameter and  $x_T$  is the true parameter value. The systematic error accounts for differences between the estimated parameters and the actual parameters that are unaccounted for by MPFIT. We express  $Err_{\text{syst}}$  as a percent error, defined as the deviation of the measured parameter from the true parameter divided by the true parameter value. This fractional error is more broadly applicable to a variety of measurement values.

By definition,  $Err_{\text{syst}}$  is not included in the statistical uncertainties reported by MPFIT, but it critically describes the ability of MPFIT to recover the actual parameter values. We account for  $Err_{\text{syst}}$  by defining the total error of each parameter ( $Err_{\text{tot}}$ ) as the quadratic sum of  $Err_{\text{syst}}$  and  $Err_{\text{stat}}$  as follows:

$$Err_{\text{tot}} = \sqrt{Err_{\text{syst}}^2 + Err_{\text{stat}}^2}. \quad (9)$$

The systematic errors derived here are included in the errors of the H I covering fraction and column density (Cols. 3 and 6 of Table 3).

### 3.3. Fitting observed data

Using the fitting method described in Sect. 3.1.2, we fit the observed spectra of the 18 galaxies in our sample. The resulting fit parameters  $E_{B-V}$ ,  $C_f$ ,  $v$ ,  $b$ , and  $N$  are listed in Table 3, and the corresponding best-fit spectra are shown in Appendix B. The reported uncertainties on the H I column densities and covering fractions account for the systematic errors derived from simulations.

We note that Tol0440-381 and Mrk54 have much lower redshifts than the rest of the sample. Consequently there are more Milky Way absorption lines near the Lyman series. This high density of Milky Way Lyman series absorption lines means that we cannot accurately fit the Lyman series with Voigt profiles.

Rather, we used nonparametric fits to describe the H I properties (see Sect. 3.3.1).

The H I column densities are affected by large uncertainties because the low S/N and insufficient resolution of the spectra do not allow a reliable estimate of  $N_{\text{HI}}$  given the saturation of the H I absorption lines (see Sect. 4.1). Therefore direct  $N_{\text{HI}}$  measurements from the Lyman series are not reliable and an indirect method is used (Sect. 3.3.2). To examine whether saturation and resulting degeneracies affect the  $C_f$  or metal column densities that we derived, we refit the spectra fixing the H I column density to  $N_{\text{HI}} = 10^{18} \text{ cm}^{-2}$ . This  $N_{\text{HI}}$  corresponds to an optically thick portion of the curve of growth for the observed Lyman series lines (from Ly $\beta$  to Ly6). We find that all the fit parameters (except for  $b(\text{H I})$ ) are consistent within  $1\sigma$  with the results listed in Table 3. The fact that  $b$  changes is consistent with  $b$  being degenerate with  $N_{\text{HI}}$  at these column densities. This indicates that saturated H I absorption does not affect the measured velocities and covering fractions.

### 3.3.1. Nonparametric measurements of the H I covering fraction

Fitting the Lyman series using the radiative transfer equation (Eq. (1)) is one way to measure the H I covering fraction. However, this assumes that the absorption profiles follow a single Lorentzian velocity distribution. Observed absorption profiles arising from galactic outflows are highly non-Lorentzian (Heckman et al. 2000; Pettini et al. 2002; Shapley et al. 2003; Weiner et al. 2009; Chisholm et al. 2017b). Consequently, we also measured  $C_f$  from the residual flux of the Lyman series lines after removing the stellar continuum (see Eq. (3)). This assumes a uniform dust geometry and that the Lyman series is fully saturated. The Lyman series is saturated from  $N_{\text{HI}} \gtrsim 10^{16} \text{ cm}^{-2}$  (for Ly $\beta$  to Ly6). Importantly, this nonparametric approach does not assume a velocity distribution of the H I gas and accounts for arbitrary line profiles. In other words, the nonparametric approach does not assume that the gas has the same covering fraction at every velocity (cf. discussion in Vasei et al. 2016).

The covering fraction is derived as the maximum of  $(1 - F_{\text{Gau}})$  in a velocity range chosen by eye near the deepest part of each Lyman series line. The  $F_{\text{Gau}}$  is the stellar continuum removed flux, modified by a Gaussian kernel centered on zero with standard deviation corresponding to the error array. We measured  $(1 - F_{\text{Gau}})$  1000 times, where each time the flux value of each pixel is determined from a different noise distribution. We then took the median and standard deviation of this distribution as the  $C_f$  value and uncertainty for each Lyman series transition. Table C.3 lists the  $C_f$  derived from the residual flux of each Lyman series transition in each galaxy. We then defined the  $C_f(\text{H I})$  from the residual flux as the error weighted mean ( $M_W$ ) of the  $i$  observed Lyman series transitions and the  $C_f$  error as the error on  $M_W$  as

$$M_W = \frac{\sum_{i=1}^n C_{f_i} \times \omega_i}{\sum_{i=1}^n \omega_i} \text{ with } \omega_i = \frac{1}{\sigma_i^2},$$

$$\sigma_{M_W} = \sqrt{\frac{1}{\sum_{i=1}^n \omega_i}}.$$
(10)

The corresponding values, denoted as  $C_f$  depth, are reported in Table 3. The  $C_f$  depth values are consistent, within the errors, with the  $C_f$  values derived from the fits in Sect. 3.3. Except for Tol0440–381 and Mrk54, where we did not fit the absorption lines with the method in Sect. 3.1.2, we computed the final

$C_f(\text{H I})$  as the weighted mean of the  $C_f$  depth and the fitted  $C_f$  values (column 8 in Table 3).

### 3.3.2. Indirect measurements of the H I column density

Since the Lyman series lines are saturated, but not damped, and the spectra have insufficient spectral resolution and S/N direct H I column density measurements are largely unconstrained, as discussed further in Sect. 4.1. Therefore, indirect methods of measuring  $N_{\text{HI}}$  are needed. We used the O I absorption lines, constrained when possible by the unsaturated O I 1039 Å absorption line, to derive the O I column density. Using the known metallicity ( $12 + \log(\text{O}/\text{H})$ ) of the galaxy, we then indirectly inferred the hydrogen column density. This approach assumes that the emission-line based oxygen abundance, tracing the chemical composition of the ionized gas, is identical to that of the neutral gas. If the metallicity was lower in the neutral gas, for example, because of the presence of some pristine gas (see, e.g., Leboutteiller et al. 2013, and references therein) or incomplete mixing of the metals (see Sect. 5.2), the resulting  $N_{\text{HI}}$  would be higher (and more saturated) than inferred here.

The O I fit parameters  $b$ ,  $v$ ,  $N_{\text{OI}}$ , and the H I column density derived using the metallicities from Table 1, are listed in Table 4. A curve of growth analysis indicates that O I 1039 Å saturates at  $N_{\text{OI}} > 10^{16.5} \text{ cm}^{-2}$ , whereas all but one of our  $N_{\text{OI}}$  are less than  $10^{16.2} \text{ cm}^{-2}$ . If the O I 1039 line is saturated, then the  $N_{\text{HI}}$  values in Table 4 would underestimate the actual  $N_{\text{HI}}$  (i.e.,  $N_{\text{HI}}$  is not lower than the quoted values). Figure B.3 demonstrates that O I 1039 Å is resolved from the C II 1036 Å and the O VI 1038 Å absorption lines. However, O I is undetected for J0925+1409 (see Fig. B.5), and we used O I 1302 Å for Tol0440–381 and Mrk54. We could not determine  $N_{\text{HI}}$  for the three MEGaSAURA sources, since the O I line was either saturated, or there was not a literature metallicity. The resulting H I column densities are found to be  $\log(N_{\text{HI}}) \sim 18.6\text{--}20 \text{ cm}^{-2}$ .

We fixed the O I covering fractions to the H I values (i.e.,  $C_f(\text{O I}) = C_f(\text{H I})$ ). This is plausible because O I and H I have similar ionization potentials and their ionization fractions are locked together by charge exchange. We examined how tying the covering fractions impact the derived O I column densities. As an extreme case, we fixed the O I covering fraction to 1 for J1152+3400, which has one of the lowest  $C_f(\text{H I})$ , and J0921+4509, which has the lowest  $N_{\text{HI}}$  derived from O I. We found that  $\log(N_{\text{OI}})$  is reduced by 0.30 and 0.15 dex, respectively, which is comparable to the  $N_{\text{OI}}$  errors. Therefore, we conclude that tying  $C_f(\text{O I})$  to  $C_f(\text{H I})$  does not drastically change the measured O I column density.

### 3.3.3. The Si II covering fraction

We measured the Si II covering fraction ( $C_f(\text{Si II})$ ) using the Si II 1190 Å doublet and the Si II 1260 Å singlet. For the Si II 1260 Å singlet,  $C_f(\text{Si II})$  was derived using the same procedure as for the residual flux measurements of the Lyman series (Sect. 3.3.1). A different approach, detailed in Chisholm et al. (2017b), was adapted for the doublet. As the two transitions share the same  $C_f$  and the ratio of the two optical depths is given by the ratio of their oscillator strengths,  $f$ , the velocity-resolved covering fraction was measured from a system of equations (Hamann et al. 1997) as

$$C_f(v) = \frac{F_W(v)^2 - 2F_W(v) + 1}{F_S(v) - 2F_W(v) + 1},$$
(11)

**Table 4.** Fitted O I properties.

Galaxy name	$\log(N_{\text{O I}})$ [ $\log(\text{cm}^{-2})$ ]	$b$ [ $\text{km s}^{-1}$ ]	$v$ [ $\text{km s}^{-1}$ ]	$\log(N_{\text{H I}})$ [ $\log(\text{cm}^{-2})$ ]
(1)	(2)	(3)	(4)	(5)
J0921+4509	$15.30 \pm 0.13$	$45 \pm 15$	$62 \pm 11$	$18.63 \pm 0.19$
J1503+3644	$15.55 \pm 0.16$	$302 \pm 98$	$102 \pm 77$	$19.60 \pm 0.17$
J0925+1409	–	–	–	–
J1152+3400	$15.43 \pm 0.17$	$227 \pm 129$	$-102 \pm 83$	$19.43 \pm 0.18$
J1333+6246	$15.54 \pm 0.35$	$287 \pm 277$	$-213 \pm 152$	$19.78 \pm 0.37$
J1442–0209	$15.62 \pm 0.58$	$178 \pm 145$	$82 \pm 101$	$19.69 \pm 0.58$
Tol1247–232	$15.29 \pm 0.43$	$278 \pm 168$	$58 \pm 143$	$19.19 \pm 0.44$
Tol0440–381 <sup>a</sup>	$15.47 \pm 0.02$	$623 \pm 28$	$12 \pm 19$	$19.27 \pm 0.10$
Mrk54 <sup>a</sup>	$16.07 \pm 0.01$	$619 \pm 8$	$15 \pm 6$	$19.37 \pm 0.10$
J0926+4427	$15.77 \pm 0.02$	$118 \pm 5$	$-141 \pm 4$	$19.76 \pm 0.05$
J1429+0643	$15.55 \pm 0.24$	$218 \pm 74$	$-43 \pm 98$	$19.35 \pm 0.25$
GP0303–0759	$15.41 \pm 0.18$	$209 \pm 84$	$31 \pm 70$	$19.55 \pm 0.19$
GP1244+0216	$16.12 \pm 0.14$	$157 \pm 47$	$-67 \pm 36$	$19.95 \pm 0.15$
GP1054+5238	$15.73 \pm 0.14$	$256 \pm 72$	$39 \pm 55$	$19.63 \pm 0.15$
GP0911+1831	$15.73 \pm 0.15$	$221 \pm 79$	$-102 \pm 67$	$19.73 \pm 0.16$
SGAS J1226	$16.11 \pm 0.03$	$200 \pm 7$	$-290 \pm 11$	–
SGAS J1527	$15.57 \pm 0.13$	$133 \pm 44$	$-89 \pm 33$	–
Cosmic Eye	$20.64^b \pm 0.66$	$24 \pm 77$	$208 \pm 9$	$24.04^b \pm 0.67$

**Notes.** Column 1: Galaxy name; Col. 2: logarithm of O I column density; Col. 3: O I Doppler b-parameter; Col. 4: velocity shift of the O I line. The O I covering fraction is fixed to the H I final value from Table 3. We did not detect the O I absorption lines in the J0925+1409 spectra. Column 5: H I column density derived from the product of (3) and  $12 + \log(\text{O}/\text{H})$ ; Table 1). We cannot estimate  $N_{\text{H I}}$  for SGAS J1226 and SGAS J1527 because they do not have a measured  $12 + \log(\text{O}/\text{H})$ . <sup>(a)</sup>Used O I 1302 Å <sup>(b)</sup>saturated O I line, hence unreliable  $N_{\text{H I}}$  determination. From the damped Ly $\alpha$  profile, [Quider et al. \(2010\)](#) obtained  $N_{\text{H I}} = (3.0 \pm 0.8) \times 10^{21} \text{ cm}^{-2}$ .

where  $F_W$  is the continuum subtracted flux of the weaker doublet line (Si II 1190 Å) and  $F_S$  is the continuum subtracted flux of the stronger doublet line (Si II 1193 Å). This method accounts for the possibility that the lines are not saturated. We measured the covering fraction 1000 times by varying the flux in a similar way to the residual flux measurements of the Lyman series. The median and standard deviation of this distribution is taken as the  $C_f(\text{Si II})$  value and error (Table 5). Using the Si II 1260 Å singlet assumes that the line is saturated, whereas the doublet method accounts for unsaturated absorption. We derived statistically consistent values using both methods (all within  $1\sigma$ ) for all the galaxies for which we can measure  $C_f(\text{Si II})$  with both methods. This indicates that Si II 1260 Å is indeed saturated.

### 3.4. Effect of ultraviolet attenuation law on covering fraction

A priori the measurements made here also depend on the dust attenuation law used. In this study, we used the attenuation law from [Reddy et al. \(2016a\)](#) because it is defined blueward of Ly $\alpha$ . We refit J1503+3644 and GP0911+1831, which are two galaxies with high  $E_{\text{B-V}}$  and low  $C_f$ , using a Small Magellanic Cloud (SMC) attenuation law<sup>1</sup> still assuming a uniform dust foreground. The SMC law is significantly steeper than the [Reddy et al. \(2016a\)](#) law. With the SMC dust law, we measured  $C_f(\text{H I})$

<sup>1</sup> Values have been taken from the IDL routine from J. Xavier Prochaska: <https://github.com/profxj/xidl/tree/master/Dust>

of  $0.653 \pm 0.109$  and  $0.744 \pm 0.156$ , respectively. These are consistent, within  $1\sigma$ , with the  $C_f(\text{H I})$  estimated using the [Reddy et al. \(2016a\)](#) attenuation law ( $0.634 \pm 0.109$  and  $0.731 \pm 0.150$ ). The  $E_{\text{B-V}}$  values change based upon the attenuation law used to match the observed continuum, but these changes do not affect the measured  $C_f$ . We therefore conclude that the adopted attenuation law does not significantly impact the measured covering fractions.

## 4. Recovering H I properties from simulated spectra

We simulated synthetic spectra and fit these mock H I lines with the method in Sect. 3.2. Comparing the fitted results with the parameters that created the spectra characterizes how accurately the method returns the H I parameters. We discuss these results in the context of the H I column density (Sect. 4.1) and the covering fraction (Sect. 4.2). This discussion illustrates that  $C_f$  is accurately measured for most resolutions and S/Ns, while the H I column density has large uncertainties. These simulations are especially helpful for planning future observations by determining the S/Ns and resolutions required to accurately measure the covering fractions of LyC emitters.

### 4.1. H I column densities

The synthetic spectra allow us to quantify how accurately our method reproduces the H I properties. For S/Ns (per pixel) less than 10 and resolutions less than 3000, the simulations have

**Table 5.** Si II covering fractions.

Galaxy name (1)	Si II 1190 Å (2)	Si II 1260 Å (3)	Mean (4)
J0921+4509	$0.66 \pm 0.32$	$0.59 \pm 0.15$	$0.60 \pm 0.14$
J1503+3644	$0.38 \pm 0.35$	$0.56 \pm 0.45$	$0.45 \pm 0.28$
J0925+1409	$0.40 \pm 0.23$	$0.43 \pm 0.34$	$0.41 \pm 0.19$
J1152+3400	$0.27 \pm 0.26$	$0.23 \pm 0.55$	$0.27 \pm 0.24$
J1333+6246	$0.29 \pm 0.26$	$0.56 \pm 0.35$	$0.39 \pm 0.21$
J1442–0209	$0.46 \pm 0.23$	$0.48 \pm 0.34$	$0.47 \pm 0.19$
Tol1247–232	–	$0.26 \pm 0.01$	$0.26 \pm 0.01$
Tol0440–381	–	$0.37 \pm 0.04$	$0.37 \pm 0.04$
Mrk54	–	$0.32 \pm 0.01$	$0.32 \pm 0.01$
J0926+4427	$0.37 \pm 0.19$	$0.36 \pm 0.06$	$0.36 \pm 0.06$
J1429+0643	$0.73 \pm 0.26$	$0.78 \pm 0.11$	$0.77 \pm 0.10$
GP0303–0759	$0.54 \pm 0.29$	$0.47 \pm 0.12$	$0.48 \pm 0.11$
GP1244+0216	$0.49 \pm 0.39$	$0.51 \pm 0.18$	$0.50 \pm 0.16$
GP1054+5238	$0.41 \pm 0.34$	$0.49 \pm 0.22$	$0.46 \pm 0.19$
GP0911+1831	$0.39 \pm 0.36$	$0.45 \pm 0.33$	$0.43 \pm 0.24$
SGAS J1226	$0.97 \pm 0.06$	$0.91 \pm 0.04$	$0.93 \pm 0.03$
SGAS J1527	$0.83 \pm 0.16$	$0.79 \pm 0.10$	$0.80 \pm 0.08$
Cosmic Eye	$0.98 \pm 0.06$	$0.94 \pm 0.03$	$0.95 \pm 0.02$

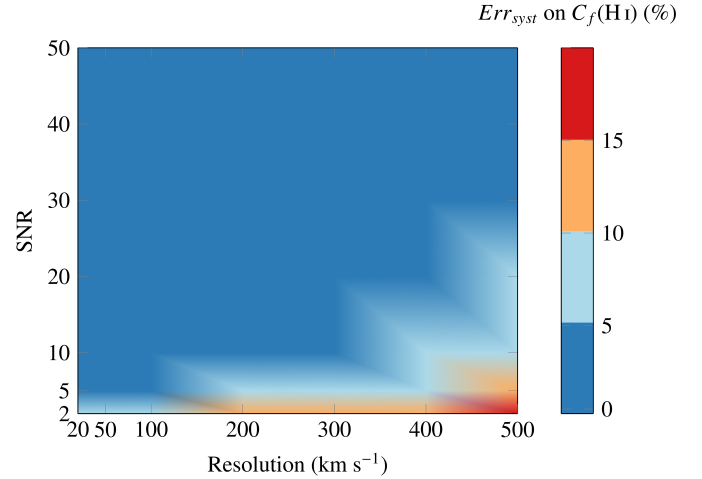
**Notes.** Column 1: Galaxy name; Col. 2:  $C_f(\text{Si II})$  derived from the Si II 1190 Å doublet using Eq. (11); Col. 3:  $C_f(\text{Si II})$  derived from the Si II 1260 Å absorption line; Col. 4: weighted mean between Cols. 2 and 3. We do not observe the Si II 1190 Å doublet for Tol1247–232, Tol0440–381, and Mrk54 because these lines fall in the COS detector gap.

$N_{\text{HI}}$  percent errors greater than 300% (Table C.1). Even at higher resolutions ( $R = 15\,000$ ), the percent error is greater than 200%, unless the S/N is greater than 5. For the lower S/Ns typical of our observations, we measured order of magnitude systematic uncertainties on  $N_{\text{HI}}$ . These large uncertainties are inherent because the Lyman series transitions saturate for HI column densities between  $N = 10^{16} \text{ cm}^{-2}$  to  $N = 10^{22} \text{ cm}^{-2}$ . For these so-called Lyman limit systems, high quality and very high-resolution spectra ( $R \sim 30\,000$ ) are needed to constrain  $N_{\text{HI}}$  with Voigt fitting methods (see, e.g., O’Meara et al. 2007). Therefore, we conclude that  $N_{\text{HI}}$  cannot be directly fitted from the Lyman series absorption lines. However, the O I absorption lines included in our fits remain unsaturated for  $N_{\text{OI}} < 10^{16.5} \text{ cm}^{-2}$  and do not suffer from these large uncertainties. Consequently, the neutral column density is most accurately inferred by converting  $N_{\text{OI}}$  into  $N_{\text{HI}}$  using the gas-phase metallicity, as in Sect. 3.3.2.

#### 4.2. HI covering fractions

Conversely, the simulations show that  $C_f(\text{HI})$  has a low percent error, under typical observing conditions. At  $R > 1500$  and  $S/N > 5$ , the  $C_f$  systematic percent errors are less than 6% of the measured value (Fig. 2, Table C.2). Therefore, the neutral gas covering fractions are accurately recovered from our observational conditions.

The *James Webb* Space Telescope (JWST) will accurately measure  $C_f$  of metal absorption lines from high-redshift



**Fig. 2.** Color map of the systematic percent error of the covering fraction as a function of the resolution and S/N. The synthetic spectra are created with  $R < 120 \text{ km s}^{-1}$  combine a theoretical stellar continuum spectra with  $R(S99) = 120 \text{ km s}^{-1}$  and absorption lines of spectral resolution  $R$ . The covering fraction is recovered to within 5% of the estimated parameter for all observations within the dark blue region.

galaxies. The JWST’s Near-Infrared Spectrograph (NIRSpec) is expected to have  $R \sim 3000$  (1000) in the high- (medium-) resolution configurations. This means that the systematic errors will be 3% (6%) of the measured  $C_f$  for  $S/N = 5$  observations, illustrating the feasibility of measuring  $C_f$  from high-redshift galaxies with JWST.

## 5. Covering fraction of LyC emitters and comparison sources

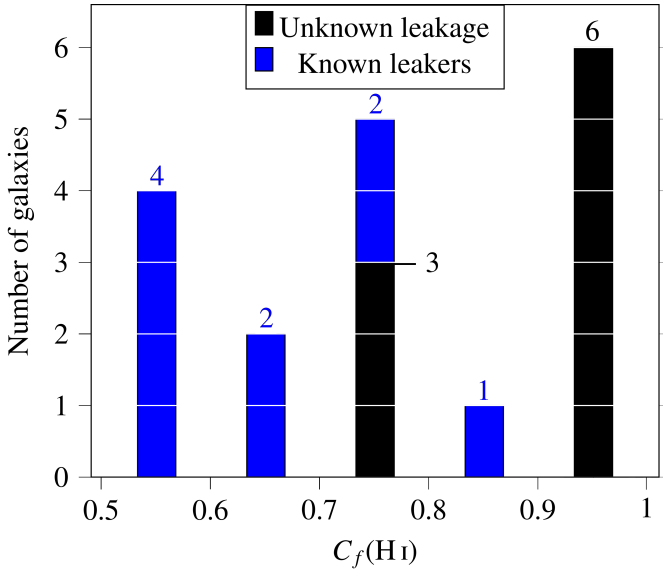
We now examine the derived covering fractions from the Lyman series and the Si II absorption lines. Then we discuss different geometrical model assumptions and compare our results to earlier work.

### 5.1. Leakers have low neutral gas covering fractions

The HI covering fraction describes the porosity of the neutral gas and demonstrates whether the neutral gas is clumpy. A smaller  $C_f(\text{HI})$  means that there are more low-density channels for ionizing photons to escape through.

Our sample has HI covering fractions ranging from 0.50 to unity (Fig. 3). Only 5 of the 18 galaxies (28%) have an HI covering fraction consistent with unity at  $1\sigma$ . The low  $C_f(\text{HI})$  values are likely because the sample is biased: 15 of these 18 galaxies were targeted as potential LyC leaker or for being particularly strong line emitters. Since a nonunity  $C_f$  is a possible LyC escape mechanism, it is not too surprising that many of these galaxies have low  $C_f(\text{HI})$ . The three galaxies that were not targeted as LyC leakers (the MEGA SaURA galaxies) have  $C_f(\text{HI})$  consistent with unity. In contrast, the galaxies with the highest confirmed escape fractions of ionizing photons have the lowest  $C_f(\text{HI})$  values (Fig. 3). The leakers have a median  $C_f(\text{HI}) = 0.62 \pm 0.10$ , while the unknown leakers have a median  $C_f(\text{HI}) = 0.95 \pm 0.10$ . Lyman continuum emitters have low HI covering fractions, which allows LyC photons to escape through low-density channels.

The high  $N_{\text{HI}}$  values, estimated from the O I column densities (Table 4), further emphasize that ionizing photons must escape



**Fig. 3.** Histogram of the H I covering fraction ( $C_f(\text{H I})$ ) from the 18 galaxies in our sample. We split the sample into galaxies known to leak LyC photons (blue) and galaxies without measured LyC emission (black). The leakers have the lowest  $C_f(\text{H I})$  and the unknown emitters have the highest  $C_f(\text{H I})$ .

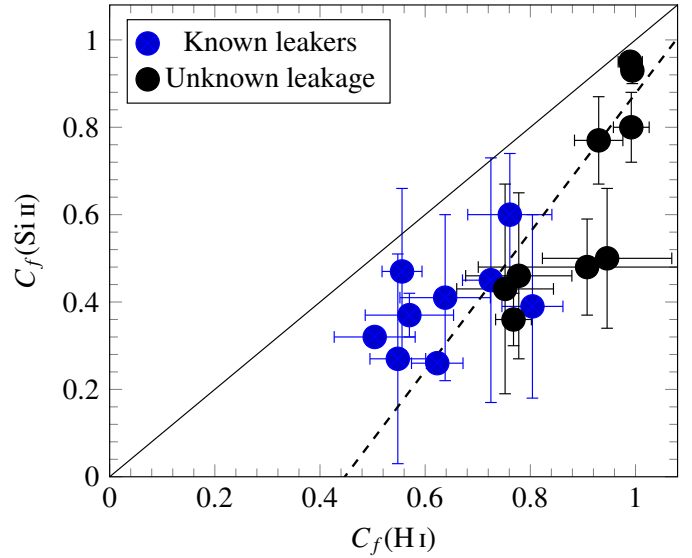
through holes in the H I.  $N_{\text{H I}}$ , calculated from  $N_{\text{O I}}$ , is greater than  $10^{18.6} \text{ cm}^{-2}$  for the entire sample. Even without converting into  $N_{\text{H I}}$ , the high  $N_{\text{O I}}$  values require unphysically high metallicities ( $12 + \log(\text{O}/\text{H}) > 10$ ) for ionizing photons to escape through low-density regions. At these column densities the Lyman series and Lyman limit are saturated. Ionizing photons cannot pass through the neutral gas unabsorbed. In other words, even LyC galaxies have optically thick H I; ionizing photons must escape through low-density channels.

Low  $C_f(\text{H I})$  values indicate that the escape of LyC photons is dominated by a patchy ISM or the picket-fence model. However, we find that a low covering fraction is not the only parameter leading to a high  $f_{\text{esc}}$ . For example, Mrk 54 has a  $C_f(\text{H I}) \sim 0.5$  and a  $f_{\text{esc}} < 1\%$ . Dust crucially impacts the LyC escape fraction by removing ionizing photons (see Eq. (5)), which assumes a uniform dust screen). Consequently, the escape fractions cannot simply be inferred from the measured covering fractions, but requires a joint determination of  $C_f$  and  $E_{\text{B-V}}$  for a given set of geometrical assumptions (e.g., cases (a) or (b) from in Sect. 3.1.1). This is discussed in depth in Sect. 5.3 and in Paper II, in which we show that the LyC escape fractions derived using the Lyman series absorption features are consistent with the directly observed escape fractions.

Using Eq. (5), we can predict which of the nine galaxies in the sample without measured LyC emission should emit ionizing photons. To emit LyC photons, both a low  $E_{\text{B-V}}$  and a low  $C_f$  are required. J0926+4427, J1429+0643, and GP1054+5238 are the best candidates in our sample to leak ionizing photons (in order of most likely to emit ionizing photons). Follow-up observations should expect to find  $f_{\text{esc}}$  values between 0.02–0.06.

## 5.2. H I and Si II covering fractions

Low-ionization metal absorption lines are often used as proxies for the neutral gas covering fraction (Shapley et al. 2003; Jones et al. 2013; Alexandroff et al. 2015; Trainor et al. 2015). Metal lines are in redder portions of the spectra. Consequently, instruments can more efficiently observe metals lines at low



**Fig. 4.** Comparison of the Si II and H I covering fractions. The black solid line shows a one-to-one relation and the dotted line is the fitted relation (Eq. (12)). The average offset between  $C_f(\text{H I})$  and  $C_f(\text{Si II})$  is 0.25.

redshifts and the Lyman forest does not obscure metal lines at high redshifts. Therefore, metal absorption lines are ideal probes of neutral gas properties when the Lyman series is not observed. However, recent observations of  $z \sim 3$  stacked spectra have suggested that metal absorption lines have covering fractions a factor of two smaller than H I absorption lines (Reddy et al. 2016b). Here, we test whether  $C_f(\text{Si II})$  traces  $C_f(\text{H I})$ .

The  $C_f(\text{Si II})$  is systematically lower than  $C_f(\text{H I})$  and has a mean offset of 0.25 (Fig. 4).  $C_f(\text{H I})$  and  $C_f(\text{Si II})$  are linearly related at the  $3\sigma$  significance level (p-value  $< 0.001$ ; Pearson's correlation coefficient of 0.79) as

$$C_f(\text{H I}) = (0.63 \pm 0.12) \times C_f(\text{Si II}) + (0.44 \pm 0.07). \quad (12)$$

While  $C_f(\text{Si II})$  is not equal to  $C_f(\text{H I})$ , Eq. (12) estimates  $C_f(\text{H I})$  from  $C_f(\text{Si II})$ . This empirically derived relation estimates  $C_f(\text{H I})$  from measurements of the metal line covering fraction (here from Si II lines), even if the Lyman series lines are not accessible. This emphasizes a clear practical advantage of the uniform dust screen geometry: the stellar attenuation can be estimated without directly measuring  $C_f(\text{H I})$ . The stellar attenuation then defines the continuum level from which  $C_f(\text{metal})$  can be estimated and, in turn,  $C_f(\text{H I})$  estimated from Eq. (12). Consequently,  $C_f(\text{H I})$  can be estimated using  $C_f(\text{Si II})$  and stellar continuum fits of high-redshift galaxies, where the Lyman series is not directly observable. This is not possible in the clumpy geometry because it requires  $C_f(\text{H I})$  to fit the stellar continuum. In Paper II we used the Lyman series absorption lines and dust attenuation to reproduce the observed Lyman continuum escape fraction.

There are numerous reasons why  $C_f(\text{Si II})$  could be related to, but not equal to,  $C_f(\text{H I})$  (see Reddy et al. 2016b, for an in-depth discussion). First, narrow, high covering fraction, metal lines may be unresolved in low-resolution spectra ( $\sim 1500$ ). While this is possible for low-resolution data, it is less likely for the high-resolution COS data. The Si II lines could possibly be unsaturated, but the doublet method accounts for this possibility and does not remove the systematic offset. Alternatively, the Si II ionization potential overlaps with, but is not equal to,

H I. Therefore, the Si II and H I gas may trace similar, but not equal, gas. However, Reddy et al. (2016b) have found that the Si II and H I line profiles and central velocities are similar (also see Chisholm et al. 2016). This indicates that the two transitions are comoving and trace similar gas. Finally, the Si II lines are also potentially affected by scattering and fluorescent emission in-filling, although it is difficult to predict whether the amount the emission increases the covering fraction (cf. Prochaska et al. 2011; Scarlata & Panagia 2015).

Another explanation is that the neutral gas and metal-enriched gas are not fully mixed. In this situation, Si II only probes metal-rich regions, while the H I gas probes both high and low metallicity gas. Therefore, the neutral gas covers more of the background source than the metals alone. At lower metallicities this is exaggerated because there is less metal-rich gas. This leads to fewer high-density metal regions to absorb the background continuum, and systematically reduces the metal covering fraction. To test the effect of metallicity on the relation between the neutral and metal covering fractions, we fit a multiple linear relationship between  $C_f(\text{H I})$ ,  $C_f(\text{Si II})$ , and  $12 + \log(\text{O}/\text{H})$ . We find a significant ( $3\sigma$ , p-value < 0.001, Pearson's correlation coefficient of 0.79) trend between these three parameters that scales as

$$C_f(\text{H I}) = (1.8 \pm 0.8) - (0.18 \pm 0.10) \times [12 + \log(\text{O}/\text{H})] + (0.75 \pm 0.16) \times C_f(\text{Si II}). \quad (13)$$

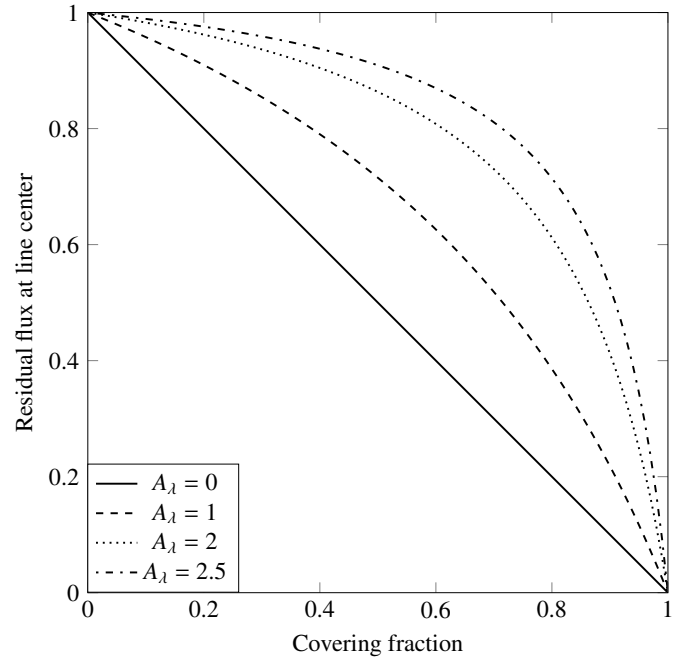
While Eq. (13) is only marginally more significant than the simpler relation between  $C_f(\text{Si II})$  and  $C_f(\text{H I})$ , it physically explains the relation between the neutral and metal covering fractions.

The difference between Eqs. (12) and (13) is the metallicity dependence of the correction added to  $C_f(\text{Si II})$ . As the metallicity increases, the required correction to convert  $C_f(\text{Si II})$  into  $C_f(\text{H I})$  decreases. While at lower metallicities, the factor added to  $C_f(\text{Si II})$  must be larger. Physically, this means that at lower metallicities the Si II traces a smaller fraction of the total H I. Therefore, the linear relationship between  $C_f(\text{H I})$  and  $C_f(\text{Si II})$  may arise because metal-rich clumps cover the background source differently than the H I gas.

The empirical relationship, Eq. (12), is recovered from Eq. (13) when the median metallicity of the sample (8.06) is used. However, Eq. (12) would change if the median  $12 + \log(\text{O}/\text{H})$  of the sample changes. This suggests that the empirical relationship between  $C_f(\text{H I})$  and  $C_f(\text{Si II})$  is only constrained near the metallicities of the sample. For divergent  $12 + \log(\text{O}/\text{H})$  values, the metallicity should be accounted for when estimating  $C_f(\text{H I})$ .

### 5.3. Is the dust geometry clumpy or uniform?

In Sect. 5.1, we emphasized that dust heavily contributes to the escape of ionizing photons. However, since the impact of dust depends on the assumed geometry, we explore the two cases described above (Sect. 3.1.1), namely (a) a uniform dust screen and (b) a clumpy geometry dust model. To test the impact of these assumed geometries on  $C_f$  and  $E_{\text{B-V}}$ , we modify our fitting routine from (a) a uniform dust screen model to (b) a clumpy model. We refit two galaxies: J0921+4509 and J1152+3400, which have an average  $C_f(\text{H I})$  and low  $C_f(\text{H I})$ , respectively. Adopting a clumpy geometry (b), we derive  $E_{\text{B-V}} = 0.236$  and 0.239 (versus  $E_{\text{B-V}} = 0.224$  and 0.134 in model (a)) and  $C_f(\text{H I}) = 0.976$  and 0.912 (versus  $C_f = 0.754$  and 0.506 in model (a)), respectively. These values and Eq. (6) predict a LyC escape



**Fig. 5.** Residual flux  $R$  of a saturated absorption line from Eq. (4) for various values of  $A_\lambda = k_\lambda E_{\text{B-V}}$  in a clumpy ISM geometry. When dust is present, the residual flux is higher than  $(1 - C_f)$ . However, in a uniform dust screen geometry, the residual flux is always  $R = (1 - C_f)$ . The covering fraction derived from the observed residual flux and a clumpy dust geometry is underestimated when dust is present.

fraction of  $f_{\text{esc}} = 0.024$  and 0.088 at  $\lambda = 912 \text{ \AA}$ , which is consistent, within  $1\sigma$ , with the values derived using the uniform dust screen model ( $f_{\text{esc}} = 0.017 \pm 0.007\%$  and  $0.092 \pm 0.027\%$ ).

As just shown, using the uniform dust screen leads to a lower  $E_{\text{B-V}}$  and  $C_f(\text{H I})$  compared to the clumpy model. In the clumpy model, all of the light that escapes through holes in the neutral gas is unattenuated. Therefore, the holes in the neutral gas must be smaller, while any continuum passing through the neutral gas must be more heavily attenuated to match the observed flux.  $C_f(\text{H I})$  and  $E_{\text{B-V}}$  must take values that fit both the geometry and the observed flux (extending to Lyman continuum). Both quantities are therefore model dependent. However, disentangling the actual geometric model is not mandatory to predict  $f_{\text{esc}}$  with  $C_f(\text{H I})$  and  $E_{\text{B-V}}$ , as long as the geometry is consistently modeled (see Paper II).

On the other hand, the observed residual flux of the H I absorption lines provide clues to the geometric model. Equation (4) shows that the residual flux of the H I absorption lines are wavelength dependent in the clumpy geometry. Therefore, if the residual flux of the individual Lyman series absorption lines changes with wavelength (for example from Ly $\beta$  to Ly6), then the data would suggest a clumpy dust geometry. Figure 5 shows the relationship between the residual flux of a saturated absorption line ( $R$ ) and the neutral covering fraction for various values of  $A_\lambda = k_\lambda E_{\text{B-V}}$  in a clumpy geometry. As  $A_\lambda$  increases (larger  $E_{\text{B-V}}$  or a bluer transition), the residual flux decreases as dust removes more continuum photons. In clumpy dust geometry,  $R$  is less than  $1 - C_f$  if there is dust. For a covering fraction of 0.8, the measured  $R$  is 0.2 when  $A_\lambda$  is 0, 0.4 when  $A_\lambda$  is 0.1, and 0.7 when  $A_\lambda$  is 0.2 (Fig. 5). The  $C_f$  estimated from  $1 - R$  is underestimated in the clumpy dust geometry case.

Since the residual intensity increases with  $A_\lambda$  and  $k_\lambda$ , this effect, if present, should be detectable in the Lyman series

**Table 6.** Theoretical residual flux for the Lyman series (Ly $\beta$  to Ly6) of a galaxy with a clumpy dust geometry.

H I lines	$E_{B-V}$			
	0.0	0.1	0.2	0.3
Ly $\beta$	0.100	0.245	0.489	0.737
Ly $\gamma$	0.100	0.254	0.512	0.764
Ly $\delta$	0.100	0.259	0.523	0.775
Ly5	0.100	0.261	0.530	0.782
Ly6	0.100	0.262	0.533	0.785

**Notes.** Theoretical measurements of the residual flux of a saturated absorption line for five Lyman series lines. These values are calculated assuming  $f_{\text{esc}} = 0.10$ , a clumpy dust geometry, saturated Lyman lines, and the Reddy et al. (2016a) dust attenuation law. The differences found between various Lyman lines for a given  $E_{B-V}$  are small and accurate residual flux measurements are required to distinguish the geometries.

absorption lines. Nevertheless, the differences between the residual flux of the Lyman series should be small and most easily detected in galaxies with large  $E_{B-V}$ . Table 6 shows the expected residual flux of saturated Lyman series lines, assuming an  $f_{\text{esc}} = 0.10$  ( $C_f = 0.9$  in the clumpy geometry). When  $E_{B-V}$  is 0.1, the difference between the Ly $\beta$  and Ly6 residual flux is less than 0.02. For  $E_{B-V} = 0.30$ , the offset rises to 0.05. While we do not measure a trend in our observations (Table C.3), the measured residual intensities have large uncertainties. Our simulations imply that  $S/N = 30$  and  $R = 15\,000$  are required to distinguish the two geometries. We cannot determine the dust geometry with the current data and both a uniform dust screen or clumpy geometry are allowed by the observations.

#### 5.4. Comparison with other studies

Reddy et al. (2016b) recently determined the H I covering fraction of  $z \sim 3$  Lyman break galaxies from Lyman series fits. Analyzing stacked spectra, they found high  $C_f(\text{H I}) \sim 0.92\text{--}0.97$ . They also examined the relation between  $C_f(\text{H I})$  and the  $C_f$  of metal lines (Si II, C II, and Al III), finding a much lower covering fraction for the metals than for H I. The differences between their  $C_f(\text{metal})$  and their  $C_f(\text{H I})$  are nearly twice as large as we find for  $C_f(\text{Si II})$  in Fig. 4 (their  $C_f(\text{metal})$  is 0.4–0.6 smaller than  $C_f(\text{H I})$ , versus our average of 0.25).

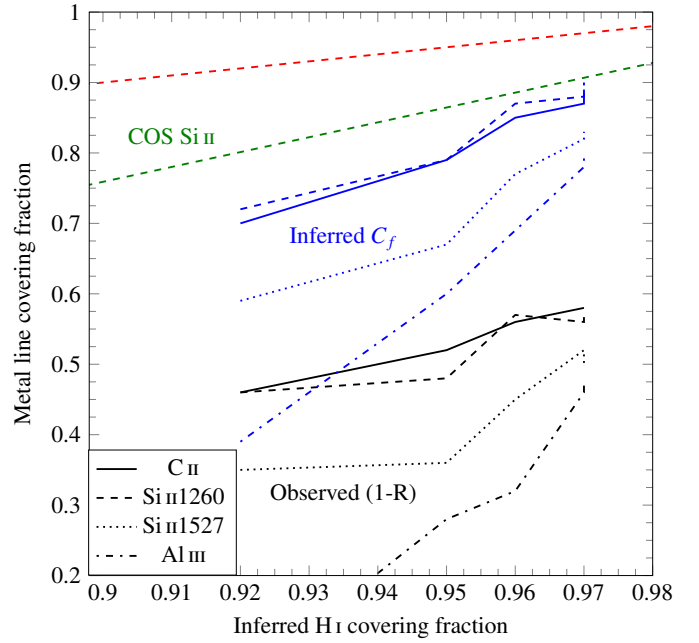
This discrepancy arises from an inconsistency in determining  $C_f(\text{metal})$ . To fit their  $C_f(\text{H I})$ , Reddy et al. (2016b) adopted a clumpy geometry with dust only in the H I. The continuum level used to obtain the normalized composite spectra in their study is determined as

$$F_{\lambda}^{\text{cont}} = F_{\lambda}^* 10^{-0.4E_{B-V}k_{\lambda}} \times C_f(\text{H I}) + F_{\lambda}^*[1 - C_f(\text{H I})]. \quad (14)$$

In this geometry, the residual flux of a saturated metal absorption line is

$$R(\text{metal}) = \frac{F_{\lambda}}{F_{\lambda}^{\text{cont}}} = \frac{1 - C_f(\text{metal})}{10^{-0.4E_{B-V}k_{\lambda}} \times C_f(\text{H I}) + 1 - C_f(\text{H I})}. \quad (15)$$

We note that Eq. (15) reduces to Eq. (4) when  $C_f(\text{metal}) = C_f(\text{H I})$ . We stress that the continuum normalization impacts the determination of  $C_f$  from R because R depends on the assumed dust geometry. In the uniform dust screen geometry (a; the geometry that we predominately use, i.e., in Fig 3),



**Fig. 6.** Metal covering fractions from Si II, C II, and Al III (as given by the line types in the legend) as a function of the H I covering fraction ( $C_f(\text{H I})$ ) from Reddy et al. (2016b).  $C_f(\text{H I})$  is always measured in a clumpy geometry. The black lines show 1–R measured from the dereddened stellar continuum normalized stacks. This assumes a clumpy geometry with a uniform dust screen. The blue lines show when the  $C_f(\text{metal})$  is derived in the same geometry as  $C_f(\text{H I})$ . When calculated with consistent geometries,  $C_f(\text{Si II } 1260 \text{ \AA})$  (blue dashed line) is offset from  $C_f(\text{H I})$  by a similar amount as our observations in Fig. 4. The green dashed line shows our fitted relationship between  $C_f(\text{Si II})$  and  $C_f(\text{H I})$  from Fig. 4, after converting it to clumpy geometry. The Reddy et al. (2016b)  $C_f(\text{H I})$ – $C_f(\text{Si II})$  relationship is consistent, within the errors, with ours. The red dashed line shows a one-to-one relation.

the covering fraction is simply 1–R. Conversely, in the clumpy model (b), the covering fraction of a saturated absorption lines is *not* simply 1–R. The covering fraction cannot be read off from the profile; a measured  $E_{B-V}$  is required to determine  $C_f$  (see the large differences in Fig. 5; Vasei et al. 2016). However, Reddy et al. (2016b) measured  $C_f(\text{metal})$  as 1–R from the continuum normalized spectra. Meanwhile, they determined  $C_f(\text{H I})$  from the full fits of their spectra, which properly accounts for their assumed geometry. This means that  $C_f(\text{metal})$  and  $C_f(\text{H I})$  are calculated with different geometric assumptions.

The inconsistency in the derivation of  $C_f(\text{metal})$  and  $C_f(\text{H I})$  produces the large mismatch between our observations and Reddy et al. (2016b). We correct the observed residual flux from Reddy et al. (2016b) to recompute consistently the  $C_f(\text{metal})$  values using Eq. (15) and their measured  $E_{B-V}$  (using their SMC extinction law) and  $C_f(\text{H I})$  values (Fig. 6). The large offset between the Si II 1260 Å and H I is halved (compare the dashed blue line to the dashed black line), although  $C_f(\text{metal})$  is still systematically smaller by  $\sim 0.2$ . This offset is similar to the offset that we measure in Fig. 4 (green dashed line in Fig. 6). The discrepancy between  $C_f(\text{metal})$  and  $C_f(\text{H I})$  emphasizes that  $C_f$  can only be read off from the residual flux of absorption lines in the uniform dust screen geometry because this continuum normalization accounts for the dust attenuation. In other geometries, the dust attenuation must be accounted for when measuring  $C_f$ .

The small remaining offset ( $\sim 0.2$ ) between the  $C_f(\text{H I})$  and  $C_f(\text{metal})$  is comparable to the  $C_f(\text{Si II})$ – $C_f(\text{H I})$  relation found

in Eq. (12). To demonstrate this, we convert our  $C_f(\text{H I})$  and  $C_f(\text{Si II})$  values, measured in a uniform geometry, to a clumpy geometry, and recalculate how  $C_f(\text{Si II})$  scales with  $C_f(\text{H I})$  (green line in Fig. 6). This recalculated relation is within the errors of the Reddy et al. (2016b) Si II 1260 Å relation (blue dashed line). The consistent scaling between  $C_f(\text{H I})$  and  $C_f(\text{Si II})$  from our study and Reddy et al. (2016b) demonstrates first that  $C_f(\text{Si II}) \neq C_f(\text{H I})$ , and second that empirical relations can convert  $C_f(\text{Si II})$  into  $C_f(\text{H I})$  for a variety of galaxy types. Both findings are crucial for indirect measurements of the Lyman continuum escape fraction, as discussed in Paper II.

There is still a  $C_f$  spread among the different transitions: strong Si II 1260 Å and C II lines have similar  $C_f$ , but weaker Si II 1527 and Al III lines have smaller  $C_f$ . Puzzlingly, the Si II 1260 Å and Si II 1527 Å lines arise from the same transition, and should have the same  $C_f$ . Different  $C_f(\text{Si II } 1260 \text{ Å})$  and  $C_f(\text{Si II } 1527 \text{ Å})$  would be expected if Si II 1527 Å was not saturated, but Reddy et al. (2016b) showed that the Si II 1527 Å/Si II 1260 Å equivalent width ratio is unity. Consequently, both transitions are saturated. Alternatively, since Si II 1260 Å is a stronger line than Si II 1527 Å ( $f$  value of 1.22 versus 0.133), there could be more optically thick Si II 1260 Å regions along the line of sight than optically thick Si II 1527 Å regions. Intersecting more regions along the line of sight covers more of the background continuum, such that  $C_f(\text{Si II } 1260 \text{ Å})$  is larger than  $C_f(\text{Si II } 1527 \text{ Å})$ . This differential covering fraction is similar to the proposed physical mechanism creating the difference between  $C_f(\text{Si II})$  and  $C_f(\text{H I})$  in Sect. 5.3. In the Si II 1260 Å and Si II 1527 Å case, the number of absorbers is different owing to differences in line strength, not metallicity.

## 6. Conclusions

We have analyzed the H I, O I, and Si II low ionization interstellar absorption lines from a sample of 18 star-forming galaxies with rest-frame ultraviolet spectroscopy of the Lyman series. The majority of the sources have COS spectra taken with the HST and are at low redshift ( $z < 0.4$ ). Our sample includes nine Lyman continuum leaking galaxies. We fit the stellar continuum, dust attenuation, H I Lyman series absorption lines, and several low ionization absorption lines (Si II  $\lambda\lambda$ 1190, 1193 and  $\lambda$ 1260 Å, O I  $\lambda$  1039 Å in particular). These fits determine the UV attenuation, as well as column densities and covering fractions of neutral hydrogen and metals (Sect. 3). Additionally, we applied our fitting method to synthetic Lyman series ISM absorption lines to investigate the systematic errors of the covering fractions and column densities. Then, we studied the observed Lyman series lines to constrain the H I properties. The direct H I properties were compared to indirect estimates of the neutral gas properties using O I and Si II absorption lines.

Our main results are summarized as follows:

- The H I covering fraction is accurately recovered from the Lyman series. Synthetic spectra recover the covering fractions with low systematics for a wide range of  $S/N$  and resolutions ( $S/N \geq 2$  and  $R > 3000$ , or  $S/N \geq 5$  and  $R \geq 600$ ; Fig. 2). Future observatories, such as JWST, will accurately measure covering fractions of high-redshift galaxies.
- The observed H I lines are found to be saturated in all galaxies. Assuming the same O/H abundance in the neutral and ionized gas, we derive H I column densities of  $\log(N_{\text{H I}}) \sim 18.6\text{--}20 \text{ cm}^{-2}$  from the O I absorption lines (Sect. 3.3.2).
- The H I column densities derived for the known LyC leakers are too high to allow ionizing photons to escape. Rather, we find that the LyC emitting galaxies have H I covering fractions below unity. Ionizing photons escape through optically thin holes in a clumpy ISM. The median H I covering fraction of confirmed LyC emitting galaxies is 0.62, as compared to 0.95 for galaxies that do not have LyC detections.
- The Si II covering fraction is systematically lower than the H I covering fraction (Fig. 4). However, the Si II covering fraction is found to scale linearly with the H I covering fraction (Eq. (12)). We show that this relation is compatible with the relationship from Reddy et al. (2016b) of stacked  $z \sim 3$  spectra. Thus, with an empirical correction, the Si II absorption lines can be used to determine the H I covering fraction when the Lyman series is not observable. This is especially powerful at high redshift when the Lyman series cannot be observed.
- The assumed dust geometry (here a uniform screen) impacts the measured covering fractions and dust attenuations (Fig. 5), but it does not impact the inferred escape fractions of ionizing photons (Sect. 5.3). Crucially, the geometric covering fraction cannot be read off from the residual line flux in all geometries. A consistent fitting of the dust attenuation, continuum, and absorption lines is required to determine the covering fraction properly.

By relating the H I covering fraction to metal absorption lines, we have provided the framework to measure the neutral gas covering fractions in star-forming galaxies. This will be of particular interest at high redshift, where the Lyman series and Lyman continuum are not observable and indirect measurements are required to measure the escape fraction of ionizing photons. In a companion paper (Chisholm et al. 2018), we use the absorption lines to accurately predict the escape fraction of the known LyC leakers. These methods also yield consistent results for a sample of  $z \sim 2 - 2.4$  galaxies. Our analysis emphasizes that UV spectra of sufficient quality ( $S/N > 5$ ) from JWST and extremely large telescopes may constrain how the Universe became reionized.

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## References

- Alexandroff, R. M., Heckman, T. M., Borthakur, S., Overzier, R., & Leitherer, C. 2015, *ApJ*, **810**, 104
- Bergvall, N., Zackrisson, E., Andersson, B.-G., et al. 2006, *A&A*, **448**, 513
- Bian, F., Fan, X., McGreer, I., Cai, Z., & Jiang, L. 2017, *ApJ*, **837**, L12
- Borthakur, S., Heckman, T. M., Leitherer, C., & Overzier, R. A. 2014, *Science*, **346**, 216
- Chisholm, J., Tremonti, C. A., Leitherer, C., et al. 2015, *ApJ*, **811**, 149
- Chisholm, J., Tremonti, C. A., Leitherer, C., Chen, Y., & Wofford, A. 2016, *MNRAS*, **457**, 3133
- Chisholm, J., Orlitová, I., Schaerer, D., et al. 2017a, *A&A*, **605**, A67
- Chisholm, J., Tremonti, C. A., Leitherer, C., & Chen, Y. 2017b, *MNRAS*, **469**, 4831
- Chisholm, J., Gazagnes, S., Schaerer, D., et al. 2018, *A&A*, **616**, A30 (Paper II)
- de Barros, S., Vanzella, E., Amorín, R., et al. 2016, *A&A*, **585**, A51
- Dressler, A., Henry, A., Martin, C. L., et al. 2015, *ApJ*, **806**, 19
- Fontanot, F., Cristiani, S., & Vanzella, E. 2012, *MNRAS*, **425**, 1413
- Fontanot, F., Cristiani, S., Pfrommer, C., Cupani, G., & Vanzella, E. 2014, *MNRAS*, **438**, 2097
- Green, J. C., Froning, C. S., Osterman, S., et al. 2012, *ApJ*, **744**, 60
- Hamann, F., Barlow, T. A., & Junkkarinen, V. 1997, *ApJ*, **478**, 87
- Heckman, T. M., Lehnert, M. D., Strickland, D. K., & Armus, L. 2000, *ApJS*, **129**, 493
- Heckman, T. M., Borthakur, S., Overzier, R., et al. 2011, *ApJ*, **730**, 5

- Henry, A., Scarlata, C., Martin, C. L., & Erb, D. 2015, *ApJ*, 809, 19
- Izotov, Y. I., Guseva, N. G., & Thuan, T. X. 2011, *ApJ*, 728, 161
- Izotov, Y. I., Orlitová, I., Schaerer, D., et al. 2016a, *Nature*, 529, 178
- Izotov, Y. I., Schaerer, D., Thuan, T. X., et al. 2016b, *MNRAS*, 461, 3683
- Izotov, Y. I., Schaerer, D., Worseck, G., et al. 2018, *MNRAS*, 474, 4514
- Jaskot, A. E., & Oey, M. S. 2013, *ApJ*, 766, 91
- Jones, T. A., Ellis, R. S., Schenker, M. A., & Stark, D. P. 2013, *ApJ*, 779, 52
- Koester, B. P., Gladders, M. D., Hennawi, J. F., et al. 2010, *ApJ*, 723, L73
- Lebouteiller, V., Heap, S., Hubeny, I., & Kunth, D. 2013, *A&A*, 553, A16
- Leitet, E., Bergvall, N., Hayes, M., Linné, S., & Zackrisson, E. 2013, *A&A*, 553, A106
- Leitherer, C., Schaerer, D., Goldader, J. D., et al. 1999, *ApJS*, 123, 3
- Leitherer, C., Ortiz Otálvaro, P. A., Bresolin, F., et al. 2010, *ApJS*, 189, 309
- Leitherer, C., Hernandez, S., Lee, J. C., & Oey, M. S. 2016, *ApJ*, 823, 64
- Markwardt, C. B. 2009, in *Astronomical Data Analysis Software and Systems XVIII*, ed. D. A. Bohlender, D. Durand, & P. Dowler, *ASP Conf. Ser.*, 411, 251
- Marshall, J. L., Burles, S., Thompson, I. B., et al. 2008, in *Ground-based and Airborne Instrumentation for Astronomy II*, *Proc. SPIE*, 7014, 701454
- Meynet, G., Maeder, A., Schaller, G., Schaerer, D., & Charbonnel, C. 1994, *A&AS*, 103, 97
- Nakajima, K., & Ouchi, M. 2014, *MNRAS*, 442, 900
- O'Meara, J. M., Prochaska, J. X., Burles, S., et al. 2007, *ApJ*, 656, 666
- Ouchi, M., Mobasher, B., Shimasaku, K., et al. 2009, *ApJ*, 706, 1136
- Pettini, M., & Pagel, B. E. J. 2004, *MNRAS*, 348, L59
- Pettini, M., Rix, S. A., Steidel, C. C., et al. 2002, *ApJ*, 569, 742
- Prochaska, J. X., Kasen, D., & Rubin, K. 2011, *ApJ*, 734, 24
- Puschnig, J., Hayes, M., Östlin, G., et al. 2017, *MNRAS*, 469, 3252
- Quider, A. M., Shapley, A. E., Pettini, M., Steidel, C. C., & Stark, D. P. 2010, *MNRAS*, 402, 1467
- Reddy, N. A., Steidel, C. C., Pettini, M., & Bogosavljević, M. 2016a, *ApJ*, 828, 107
- Reddy, N. A., Steidel, C. C., Pettini, M., Bogosavljević, M., & Shapley, A. E. 2016b, *ApJ*, 828, 108
- Rigby, J. R., Bayliss, M. B., Sharon, K., et al. 2018, *AJ*, 155, 104
- Robertson, B. E., Furlanetto, S. R., Schneider, E., et al. 2013, *ApJ*, 768, 71
- Scarlata, C., & Panagia, N. 2015, *ApJ*, 801, 43
- Shapley, A. E., Steidel, C. C., Pettini, M., & Adelberger, K. L. 2003, *ApJ*, 588, 65
- Shapley, A. E., Steidel, C. C., Strom, A. L., et al. 2016, *ApJ*, 826, L24
- Stark, D. P., Swinbank, A. M., Ellis, R. S., et al. 2008, *Nature*, 455, 775
- Trainor, R. F., Steidel, C. C., Strom, A. L., & Rudie, G. C. 2015, *ApJ*, 809, 89
- Vanzella, E., Siana, B., Cristiani, S., & Nonino, M. 2010, *MNRAS*, 404, 1672
- Vasei, K., Siana, B., Shapley, A. E., et al. 2016, *ApJ*, 831, 38
- Wakker, B. P., Hernandez, A. K., French, D. M., et al. 2015, *ApJ*, 814, 40
- Weiner, B. J., Coil, A. L., Prochaska, J. X., et al. 2009, *ApJ*, 692, 187
- Worseck, G., Prochaska, J. X., Hennawi, J. F., & McQuinn, M. 2016, *AJ*, 825, 144
- Wuyts, E., Rigby, J. R., Gladders, M. D., et al. 2012, *ApJ*, 745, 86
- Zackrisson, E., Inoue, A. K., & Jensen, H. 2013, *ApJ*, 777, 39

## Appendix A: Simulated spectra

**Table A.1.** Absorption lines, wavelengths, and parameters used to create the synthetic absorption lines.

Ion	$v$ [km s <sup>-1</sup> ]	$b$ [km s <sup>-1</sup> ]	$\log(N)$ [cm <sup>-2</sup> ]	$C_f$
(1)	(2)	(3)	(4)	(5)
H I	0	50	17.57 <sup>a</sup> 20.00 <sup>b</sup>	1 <sup>a</sup> 0.9 <sup>b</sup>
O I	0	75	16	1
O VI	0	100	15	1
C II	0	80	15.5	1

**Notes.** Column 1: Ion; Col. 2: velocity shift of the lines; Col. 3: Doppler parameter; Col. 4: logarithm of the column density; Col. 5: covering fraction. <sup>(a)</sup>Denotes values used for the simulated density-bounded spectra. <sup>(b)</sup>denotes values used for the simulated picket-fence spectra.

Here we detail the procedure we used to create the simulated spectra in Sect. 3.2. We simulated synthetic spectra to mimic the observed Lyman series. Therefore, we created synthetic spectra using a single STARBURST99 stellar continuum model of  $R(S99) \approx 2500$  (Leitherer et al. 1999, 2010) with a stellar age of 3 Myr, a solar metallicity, and without dust attenuation. This stellar age has a strong O IV P-Cygni profile (broad blueshifted absorption and redshifted emission) that blends with the Ly $\beta$  absorption line (see Fig. 1). The synthetic spectra include the wavelength regime between 910–1050 Å with a pixel separation of 0.02 Å (6 km s<sup>-1</sup> at 1035 Å). These wavelengths include the Lyman series, from Ly $\beta$  to the Lyman break ( $\approx 911.8$  Å). We imprinted on top of the stellar continuum ISM absorption lines from H I, O I, O VI, and C II (see Table 2) with Voigt profiles. The Doppler broadening parameters ( $b$ ), column densities ( $N$ ), and covering fractions ( $C_f$ ) are created with independent values for each element (see Table A.1). We fixed the velocities of each line at 0 km s<sup>-1</sup>. The Einstein-A coefficients and the oscillator strengths of each transition were taken from the NIST Atomic Database<sup>2</sup>.

We defined two categories of synthetic spectra corresponding to an ionizing escape fraction of 10%. This value is drawn from the largest known publicly available low-redshift LyC

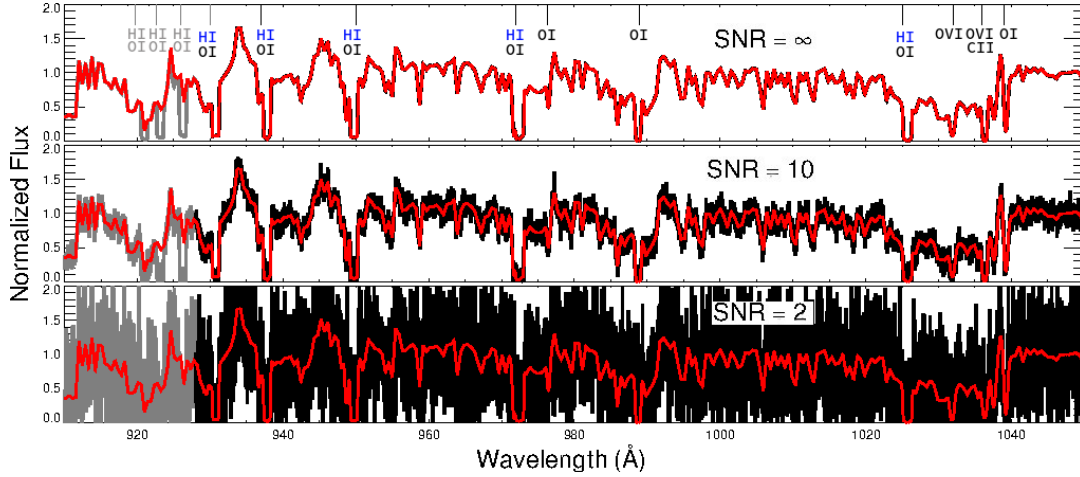
escape fractions (Izotov et al. 2016a,b). For the first category, we assumed holes in an otherwise optically thick ISM allow ionizing photons to escape. We called this the picket-fence model. For the picket-fence model, the H I parameters are  $N_{\text{HI}} = 10^{20}$  cm<sup>-2</sup> ( $e^{-\tau_{\text{H}}} \rightarrow 0$ ) and  $C_f(\text{H I}) = 0.90$ . For the second scenario, we assumed a low H I column density allows ionizing photons to escape (the density-bounded regime). We fixed  $N_{\text{HI}} = 10^{17.57}$  cm<sup>-2</sup> ( $e^{-\tau_{\text{H}}} \approx 0.10$ ) and  $C_f(\text{H I}) = 1.0$ . Since we fixed  $E_{\text{B-V}} = 0$  in all our simulations, the geometry does not change our simulations (Sect. 3.1.1).

To reproduce the observations, we created synthetic spectra at various instrumental resolutions and S/Ns. We simulated seven resolutions,  $R = \lambda/\delta\lambda$  (where  $\delta\lambda$  is the FWHM of the spectra), of 15 000, 6000, 3000, 1500, 1000, 750, and 600. For the synthetic spectra having  $R > R(S99)$ , this is performed by combining the theoretical stellar continuum with  $R(S99) = 2500$  with absorption lines of spectral resolution  $R$ . The range of resolutions chosen correspond to spectrographs such as the COS on HST, the Low and High Resolution Imaging Spectrometer (LRIS/HRIS) on Keck, MagE on *Magellan*, MUSE, or the upcoming NIRSpec on JWST. For each  $R$ , we generated seven S/Ns (2, 5, 10, 20, 30, 40, and 50) in a sufficiently large range to study typically delivered noise levels. At each S/N, we generated 50 different sets of random Gaussian noise and added the noise to the same synthetic spectra. This produces a sample of 50 synthetic spectra where only the random noise varies between the spectra. Figures B.1 and B.2 show simulated spectra with resolutions of 15 000 and 600 for three S/Ns (noise-free, 10 and 2). These results are discussed in Sect. 4.

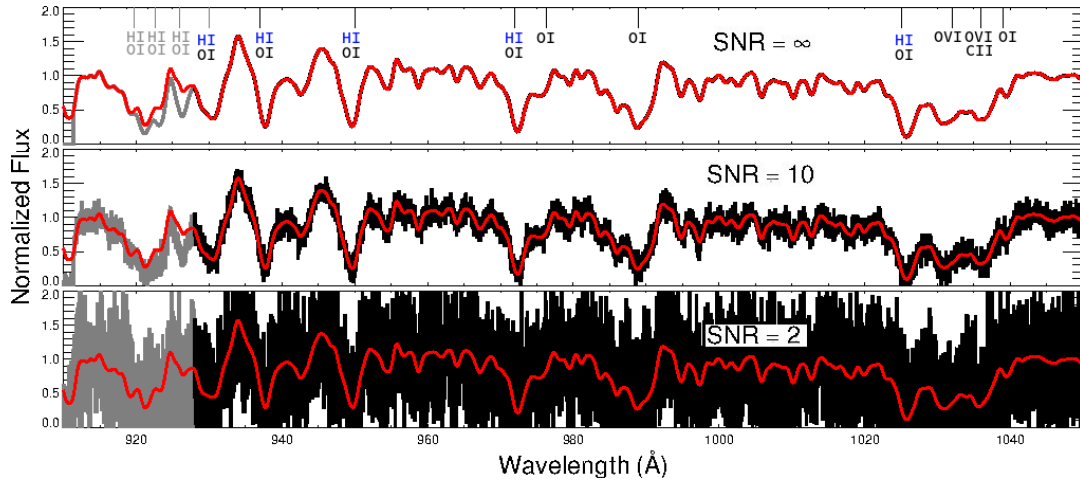
## Appendix B: Fits

We provide fits of the synthetic and observed spectra detailed in this article. Fig. B.1 and Fig. B.2 show the synthetic spectra for three S/Ns ( $\infty$ , 10 and 2) with  $R = 15\,000$  and  $R = 600$ , respectively. In Figs. B.3–B.20 we present the fits for the 18 galaxies in our sample. The top panels of each figure show the full wavelength coverage; the observed flux is indicated in black and the total fit in red. Gray flux indicates that these regions were masked during the fit. The green lines show the error on the observed flux. The lower panels zoom in on individual Lyman series lines. The upper portions of each plot indicate the fitted Lyman series lines (blue), ISM metal absorption lines (solid black line), and Milky Way absorption lines (black dashed lines). The line labels are gray if they are not fit. Geocoronal regions are denoted by shaded gray regions.

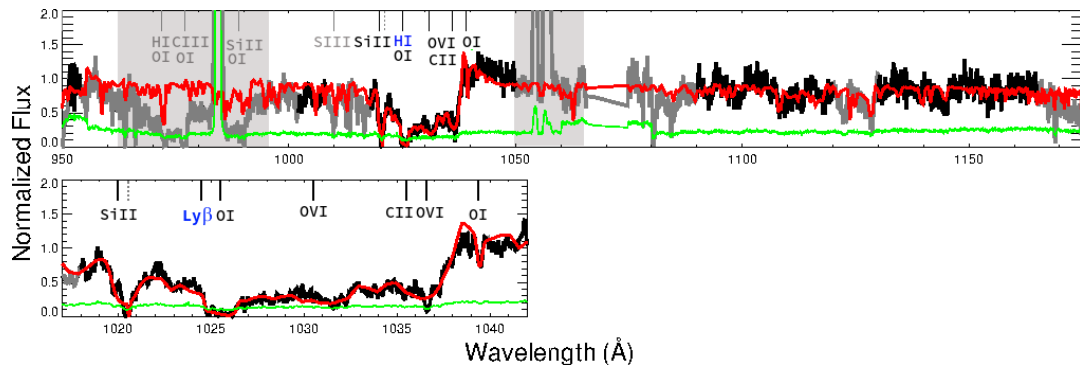
<sup>2</sup> [https://physics.nist.gov/PhysRefData/ASD/lines\\_form.html](https://physics.nist.gov/PhysRefData/ASD/lines_form.html)



**Fig. B.1.** Fits (in red) of the synthetic spectra at high resolution ( $R = 15\,000$ ) for the picket-fence model ( $C_f(\text{H I}) = 0.9$ ,  $N_{\text{H I}} = 10^{20}\text{cm}^{-2}$ ) with S/N equal to  $\infty$  (top panel), 10 (middle panel), and 2 (bottom panel). We exclude the gray portion when fitting. In the upper portion of the top panel, we indicate and label the fitted absorption lines.



**Fig. B.2.** Same as Fig. B.1 but for a spectral resolution of  $R = 600$ .



**Fig. B.3.** Fit (in red solid line) of the COS G130M spectrum of the galaxy J0921+4509. Black is the observed flux included in the routine either to fit the stellar continuum or the ISM absorption lines. Gray portions are masked out for both steps; gray shaded regions indicate those masked because of geocoronal emission. The flux error array appears in green. We display the ISM and Milky Way absorption lines as solid and dotted lines in the upper portion of each panel. Black or blue labels indicate that the lines are fit, whereas gray labels indicate that they are not fit. When present, red labels indicate lines that are not detected. Reference for this observation: [Borthakur et al. \(2014\)](#).

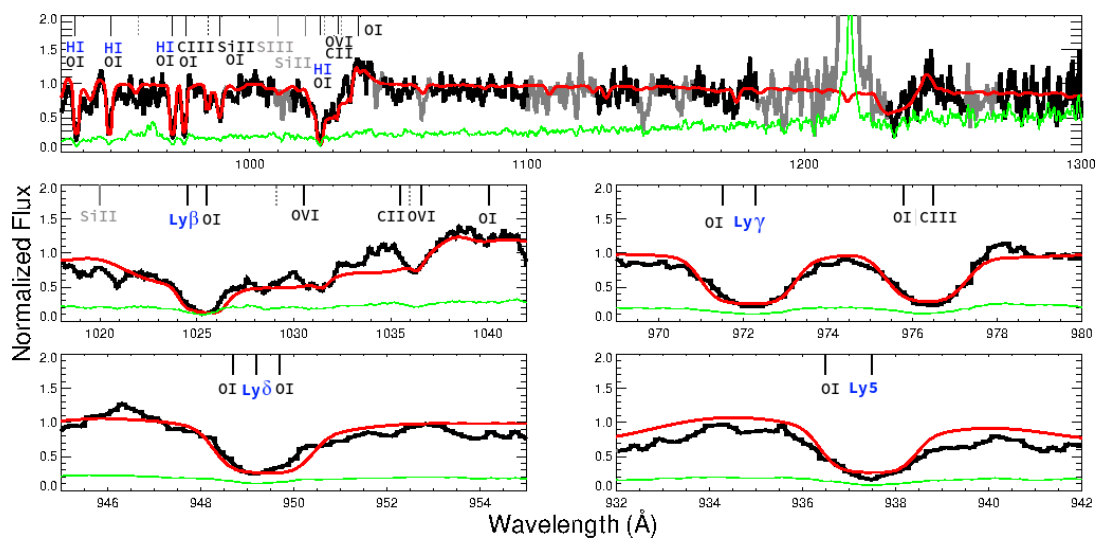


Fig. B.4. Same as Fig. B.3 but for the COS G140L spectrum of J1503+3644. Reference for this observation: Izotov et al. (2016b).

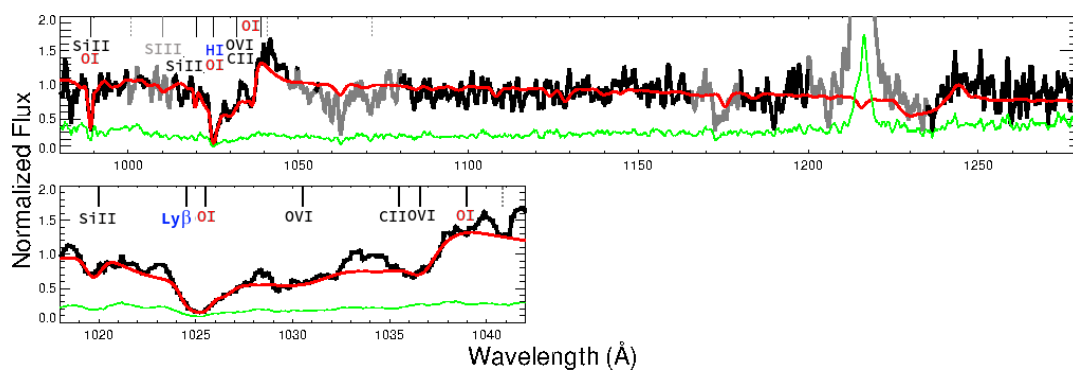


Fig. B.5. Same as Fig. B.3 but for the COS G140L spectrum of J0925+1409. Reference for this observation: Izotov et al. (2016a).

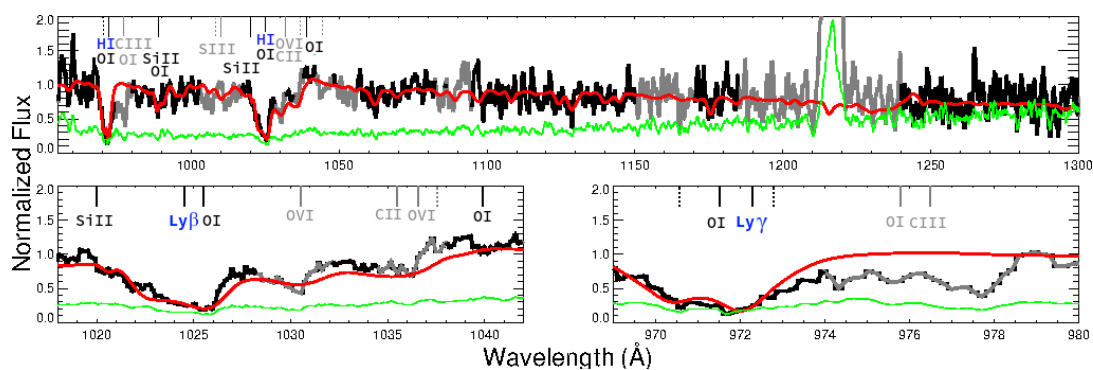
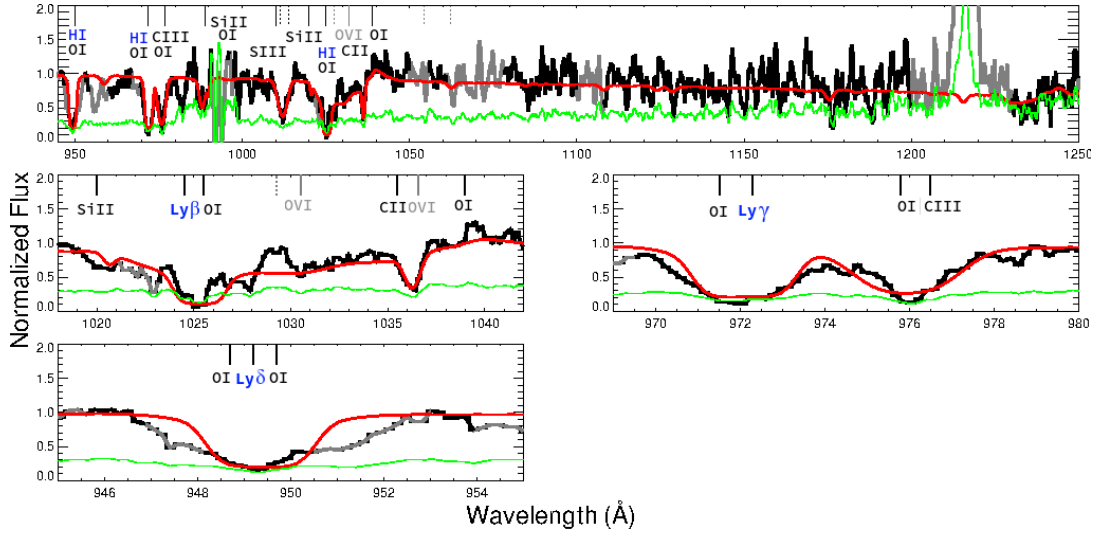
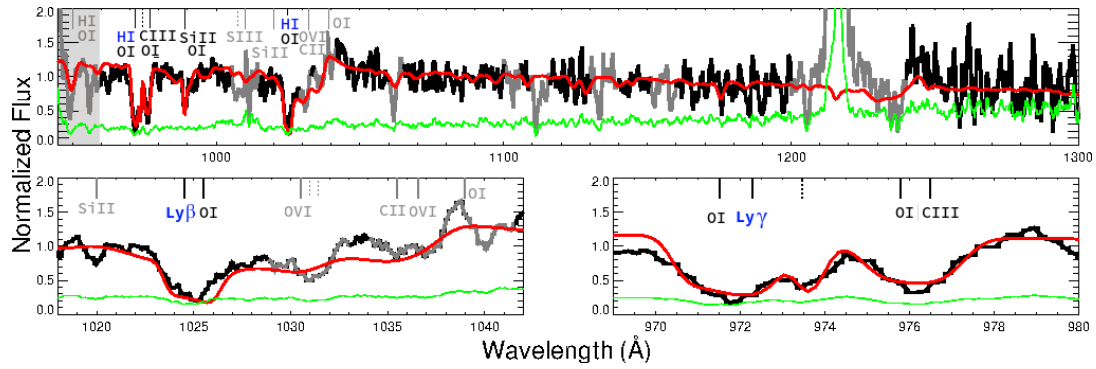


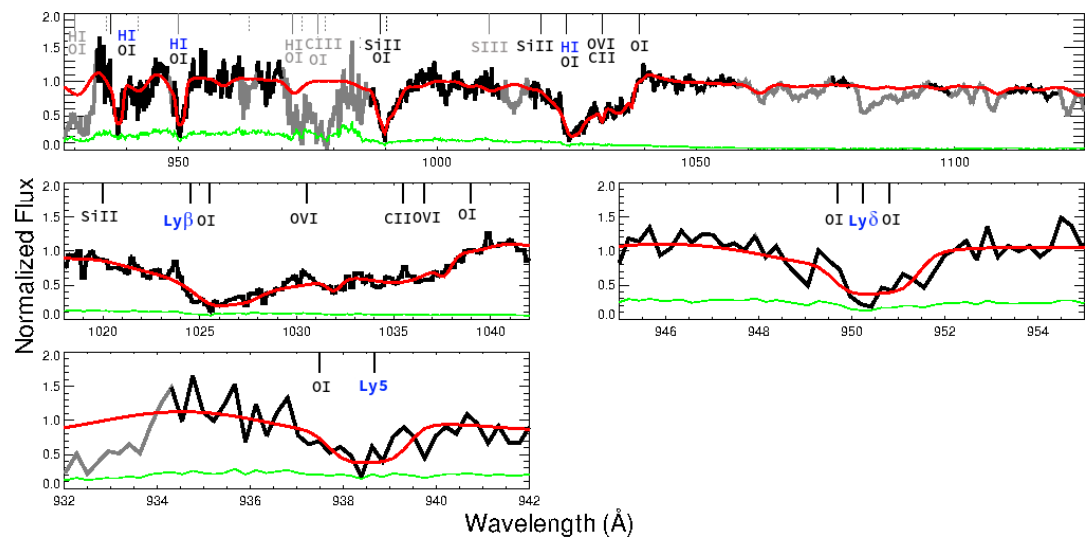
Fig. B.6. Same as Fig. B.3 but for the COS G140L spectrum of J1152+3400. Reference for this observation: Izotov et al. (2016b).



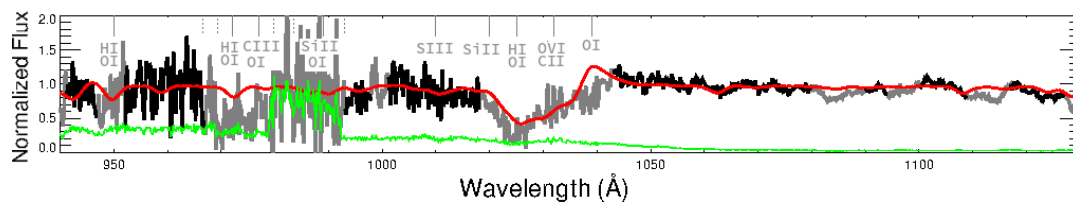
**Fig. B.7.** Same as Fig. B.3 but for the COS G140L spectrum of J1333+6246. Reference for this observation: Izotov et al. (2016b).



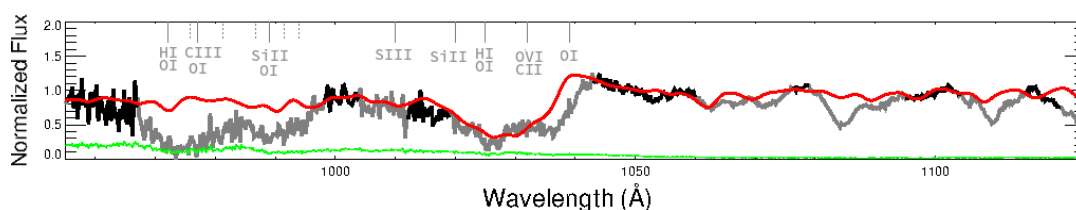
**Fig. B.8.** Same as Fig. B.3 but for the COS G140L spectrum of J1442-0209. Reference for this observation: Izotov et al. (2016b).



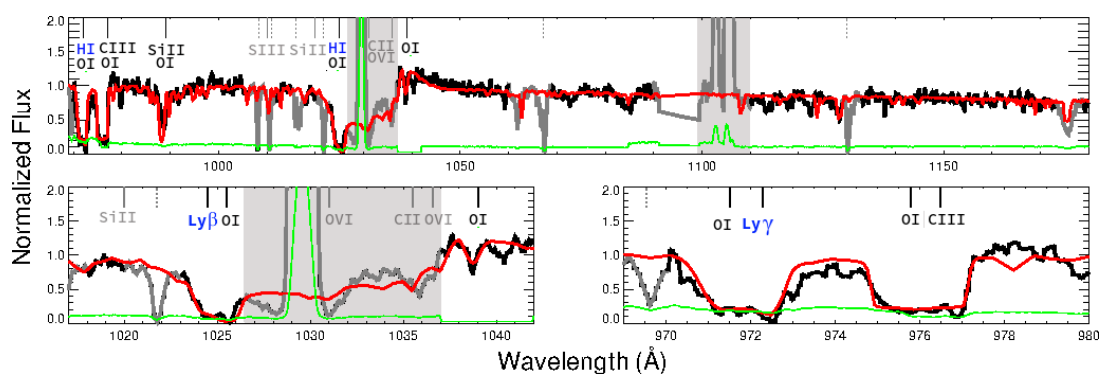
**Fig. B.9.** Same as Fig. B.3 but for the COS G140L spectrum of Tol1247-232. Reference for this observation: Leitherer et al. (2016).



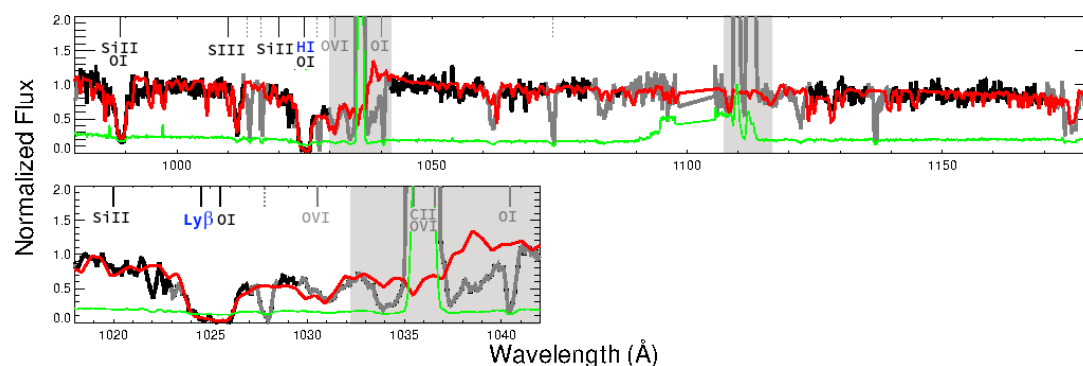
**Fig. B.10.** Best fit for the stellar continuum for the COS G140L spectrum of Tol0440–381. This galaxy has a low redshift and the Milky Way absorption lines contaminate the Lyman series. Consequently, we do not fit for the ISM absorption lines. Reference for this observation: [Leitherer et al. \(2016\)](#).



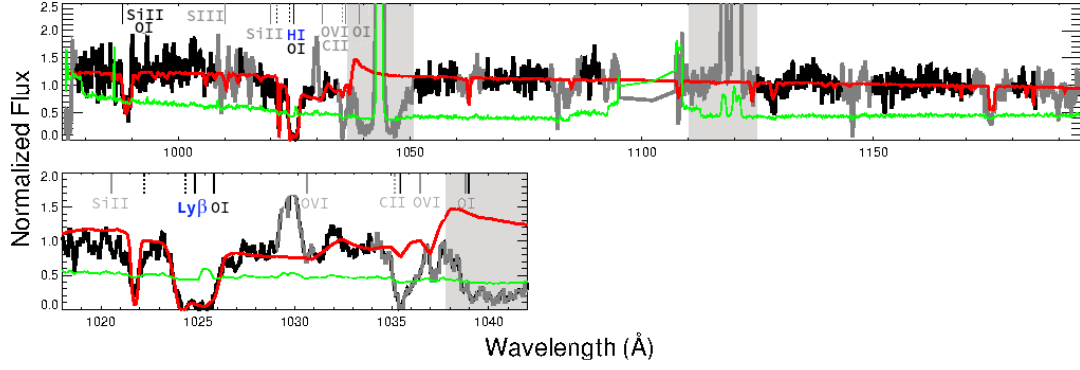
**Fig. B.11.** Same as Fig. B.10 but for the COS G140L spectrum of Mrk54. Reference for this observation: [Leitherer et al. \(2016\)](#).



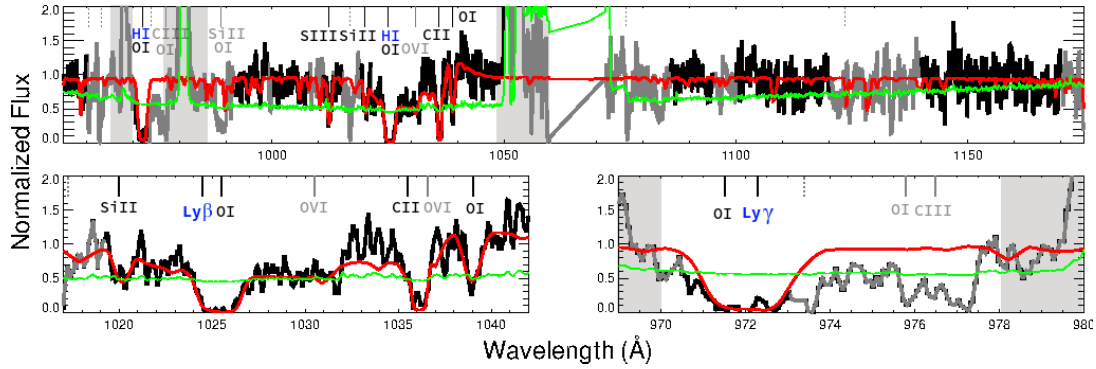
**Fig. B.12.** Same as Fig. B.3 but for the COS G130M spectrum of J0926+4427. Reference for this observation: [Heckman et al. \(2011\)](#).



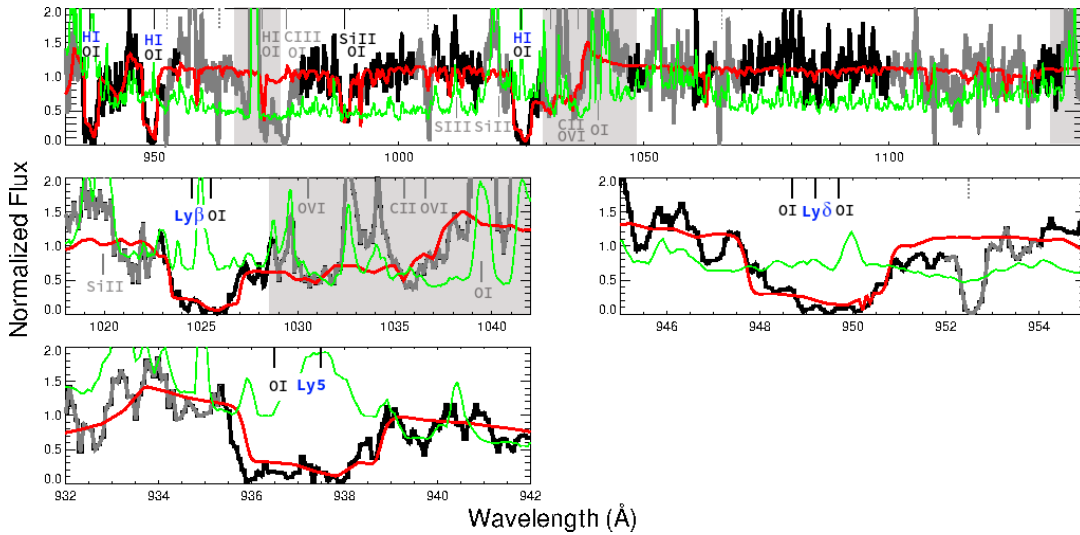
**Fig. B.13.** Same as Fig. B.3 but for the COS G130M spectrum of J1429+0643. Reference for this observation: [Heckman et al. \(2011\)](#).



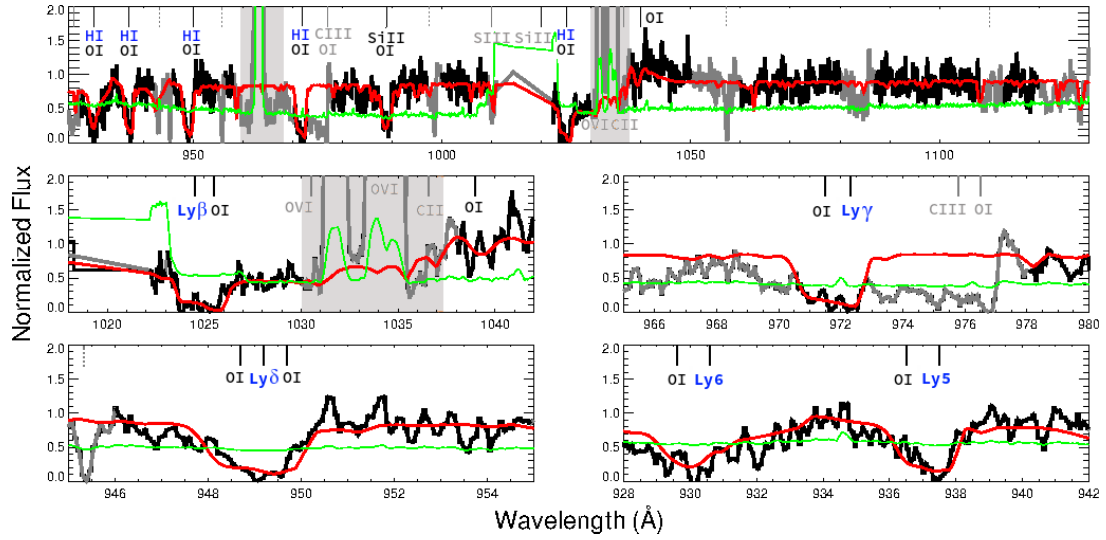
**Fig. B.14.** Same as Fig. B.3 but for the COS G130M spectrum of GP0303–0759. Reference for this observation: [Henry et al. \(2015\)](#).



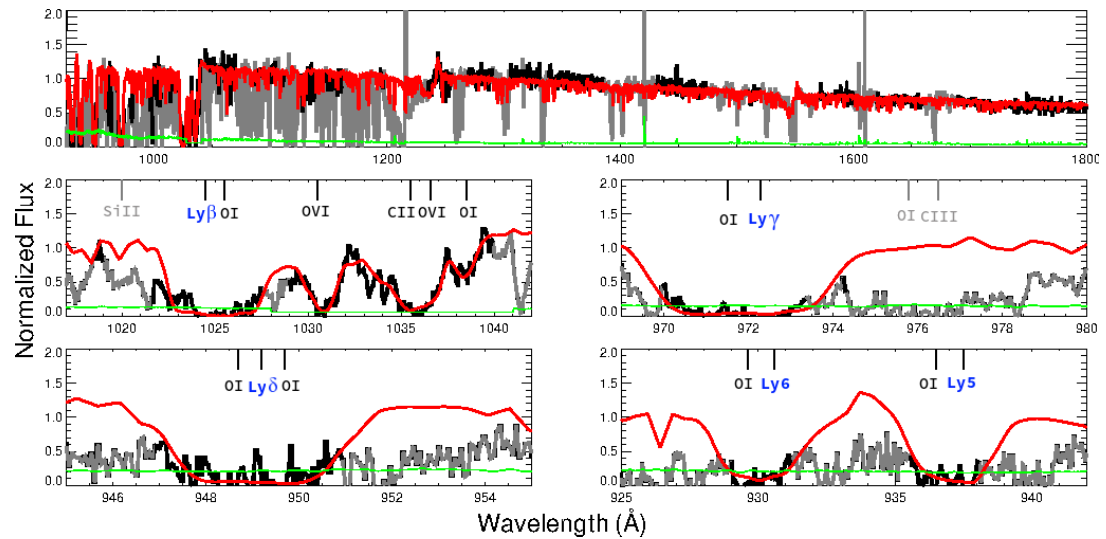
**Fig. B.15.** Same as Fig. B.3 but for the COS G130M spectrum of GP1244+0216. Reference for this observation: [Henry et al. \(2015\)](#).



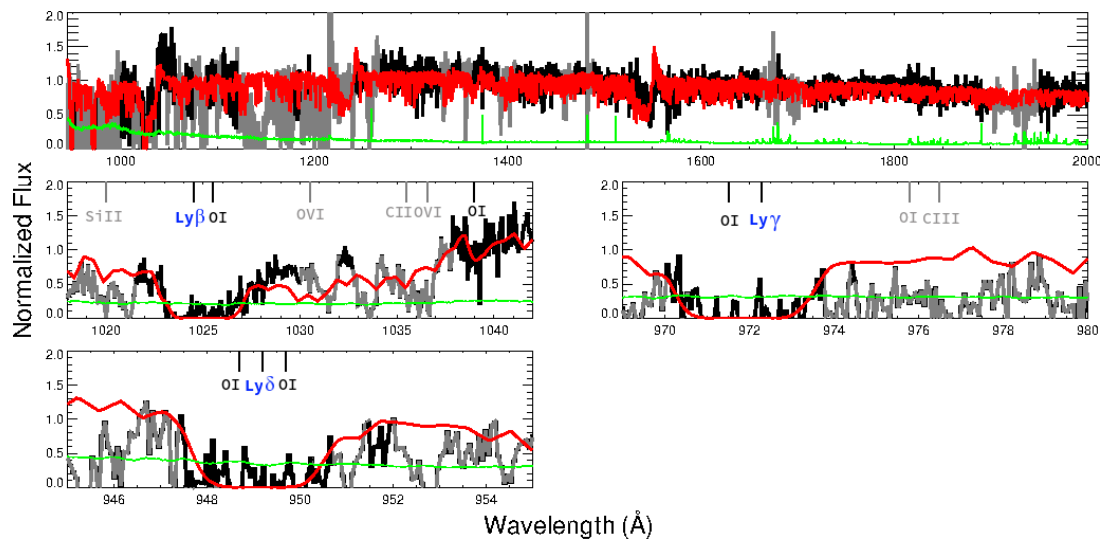
**Fig. B.16.** Same as Fig. B.3 but for the COS G130M spectrum of GP1054+5238. Reference for this observation: [Henry et al. \(2015\)](#).



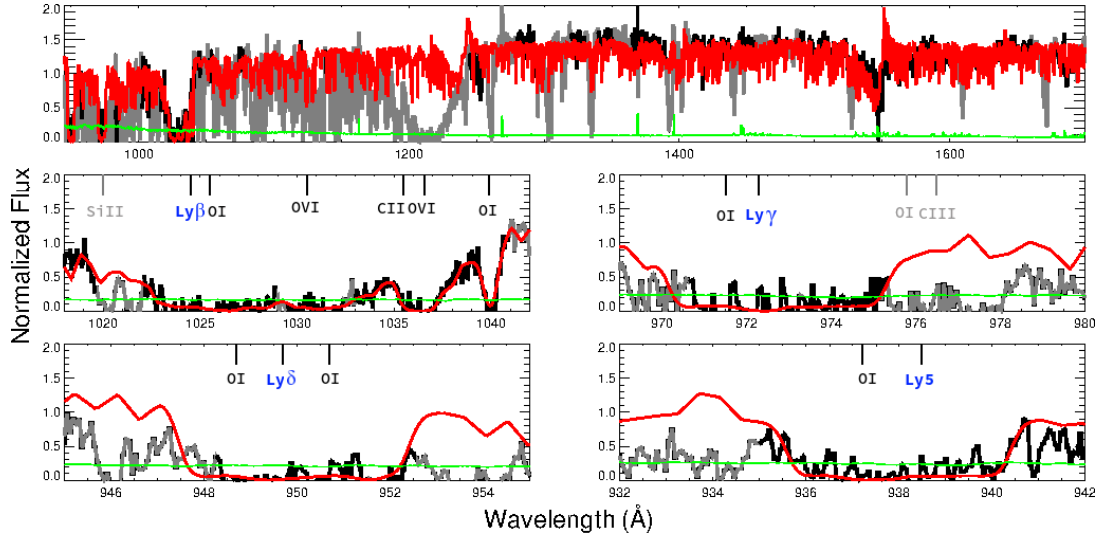
**Fig. B.17.** Same as Fig. B.3 but for the COS G130M spectrum of GP0911+1831. Reference for this observation: [Henry et al. \(2015\)](#).



**Fig. B.18.** Same as Fig. B.3 but for the MagE spectrum of SGAS J122651.3+215220. Reference for this observation: [Rigby et al. \(2018\)](#).



**Fig. B.19.** Same as Fig. B.3 but for the MagE spectrum of SGAS J152745.1+065219. Reference for this observation: [Rigby et al. \(2018\)](#).



**Fig. B.20.** Same as Fig. B.3 but for the MagE spectrum of the Cosmic Eye. Reference for this observation: [Rigby et al. \(2018\)](#).

## Appendix C: Additional tables

### C.1. Measured systematic errors from single synthetic spectra

We tabulate the systematic errors measured from the synthetic spectra (Sect. 4). Table C.1 gives the  $N_{\text{HI}}$  systematic errors, and Table C.2 gives the  $C_f$  systematic errors. The percent errors plotted in Fig. 2 are computed by subtracting 0.9 from the errors and dividing the difference by 0.9 (the  $C_f$  used to generate the spectra).

**Table C.1.** Systematic errors on the logarithm of the H I column density estimated from synthetic spectra.

$R$	$S/N$						
	50	40	30	20	10	5	2
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)
15 000	0.04	0.04	0.05	0.07	0.17	0.33	1.13
6000	0.05	0.07	0.10	0.14	0.31	0.78	1.27
3000	0.09	0.09	0.12	0.18	0.57	1.05	1.13
1500	0.10	0.17	0.20	0.28	0.65	1.04	1.28
1000	0.12	0.14	0.17	0.30	0.87	1.21	1.18
750	0.14	0.17	0.24	0.33	0.86	1.20	1.53
600	0.13	0.19	0.25	0.30	0.81	1.21	1.54

**Notes.** Columns 2–8 give the systematic  $\log(N_{\text{HI}})$  error at various spectral resolutions ( $R$ ; see Col. 1) and  $S/N$ s (see Cols. 2–7 for  $S/N$  50–2) for simulated spectra with  $N_{\text{HI}} = 10^{17.57} \text{ cm}^{-2}$ . See Sect. 4.1.

**Table C.2.** Systematic errors on the covering fraction estimated from synthetic spectra.

$R$	$S/N$						
	50	40	30	20	10	5	2
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)
15 000	<0.01	<0.01	<0.01	0.01	0.01	0.02	0.06
6000	<0.01	<0.01	<0.01	0.01	0.01	0.03	0.07
3000	<0.01	<0.01	<0.01	0.01	0.01	0.03	0.07
1500	<0.01	0.01	0.01	0.01	0.03	0.05	0.10
1000	0.01	0.011	0.02	0.02	0.04	0.07	0.12
750	0.02	0.021	0.03	0.04	0.06	0.09	0.14
600	0.03	0.032	0.04	0.05	0.09	0.10	0.17

**Notes.** Columns 2–8 give the systematic  $C_f$  errors at different spectral resolutions ( $R$ ; Col. 1) and  $S/N$ s (see Cols. 2–7 for  $S/N$  50–2) for simulated spectra with  $C_f = 0.9$ . See Sect. 4.2.

### C.2. Lyman series residual flux

Table C.3 lists the residual flux of the individual Lyman series lines (Sect. 3.3.1).

**Table C.3.** Measurement of H I covering fraction derived from the residual flux of the individual Lyman series. See Sect. 3.3.1.

Galaxy name	Ly $\beta$	Ly $\gamma$	Ly $\delta$	Ly5	Weighted mean
J0921+4509	0.769 $\pm$ 0.116	–	–	–	0.769 $\pm$ 0.116
J1503+3644	0.847 $\pm$ 0.157	0.723 $\pm$ 0.129	0.741 $\pm$ 0.112	0.744 $\pm$ 0.113	0.754 $\pm$ 0.062
J0925+1409	0.635 $\pm$ 0.094	–	–	–	0.635 $\pm$ 0.094
J1152+3400	0.619 $\pm$ 0.088	–	–	–	0.619 $\pm$ 0.088
J1333+6246	0.773 $\pm$ 0.152	0.870 $\pm$ 0.103	0.804 $\pm$ 0.106	–	0.826 $\pm$ 0.066
J1442–0209	0.589 $\pm$ 0.049	0.471 $\pm$ 0.068	–	–	0.549 $\pm$ 0.040
Tol1247–232	0.543 $\pm$ 0.135	–	0.775 $\pm$ 0.134	0.791 $\pm$ 0.175	0.690 $\pm$ 0.084
Tol0440–381	0.507 $\pm$ 0.139	–	0.615 $\pm$ 0.167	0.602 $\pm$ 0.137	0.570 $\pm$ 0.084
Mrk54	0.397 $\pm$ 0.136	–	0.652 $\pm$ 0.120	0.406 $\pm$ 0.148	0.504 $\pm$ 0.077
J0926+4427	0.817 $\pm$ 0.057	0.807 $\pm$ 0.087	–	–	0.814 $\pm$ 0.048
J1429+0643	0.955 $\pm$ 0.061	–	–	–	0.955 $\pm$ 0.061
GP0303–0759	–	–	–	–	–
GP1244+0216	0.985 $\pm$ 0.211	0.894 $\pm$ 0.204	1.000 $\pm$ 0.292	–	0.950 $\pm$ 0.131
GP1054+5238	0.936 $\pm$ 0.318	–	0.891 $\pm$ 0.203	0.798 $\pm$ 0.420	0.889 $\pm$ 0.158
GP0911+1831	0.718 $\pm$ 0.361	0.781 $\pm$ 0.189	0.825 $\pm$ 0.198	0.635 $\pm$ 0.282	0.765 $\pm$ 0.116
SGAS J1226	1.000 $\pm$ 0.010	1.000 $\pm$ 0.030	0.981 $\pm$ 0.057	0.940 $\pm$ 0.060	0.998 $\pm$ 0.009
SGAS J1527	0.858 $\pm$ 0.141	1.000 $\pm$ 0.045	1.000 $\pm$ 0.087	1.000 $\pm$ 0.201	0.990 $\pm$ 0.038
Cosmic Eye	1.000 $\pm$ 0.043	1.000 $\pm$ 0.063	1.000 $\pm$ 0.033	0.899 $\pm$ 0.166	0.998 $\pm$ 0.024

**Notes.** Dashes indicate that these transitions were not observed owing to Milky Way absorption, geocoronal emission, or low S/N.