

Binary energy source of the HH 250 outflow and its circumstellar environment[★]

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ABSTRACT

Aims. Herbig-Haro flows are signposts of recent major accretion and outflow episodes. We aim to determine the nature and properties of the little-known outflow source HH 250-IRS, which is embedded in the Aquila clouds.

Methods. We have obtained adaptive optics-assisted *L*-band images with the NACO instrument on the Very Large Telescope (VLT), together with *N*- and *Q*-band imaging with VISIR also on the VLT. Using the SINFONI instrument on the VLT we carried out *H*- and *K*-band integral field spectroscopy of HH 250-IRS, complemented with spectra obtained with the SpeX instrument at the InfraRed Telescope Facility (IRTF) in the *JHKL* bands. Finally, the SubMillimeter Array (SMA) interferometer was used to study the circumstellar environment of HH 250-IRS at 225 and 351 GHz with CO (2–1) and CO (3–2) maps and 0.9 mm and 1.3 mm continuum images.

Results. The HH 250-IRS source is resolved into a binary with 0′.53 separation, corresponding to 120 AU at the adopted distance of 225 pc. The individual components show heavily veiled spectra with weak CO absorption indicative of late-type stars. Both are Class I sources, but their spectral energy distributions between 1.5 μm and 19 μm differ markedly and suggest the existence of a large cavity around one of the components. The millimeter interferometric observations indicate that the gas mainly traces a circumbinary envelope or disk, while the dust emission is dominated by one of the circumstellar envelopes.

Conclusions. HH 250-IRS is a new addition to the handful of multiple systems where the individual stellar components, the circumstellar disks and a circumbinary disk can be studied in detail, and a rare case among those systems in which a Herbig-Haro flow is present.

Key words. circumstellar matter – stars: winds, outflows – stars: individual: IRAS 19190+1048 – stars: pre-main sequence – ISM: jets and outflows – binaries: general

1. Introduction

The majority of stars are members of binary systems (e.g. Duquennoy & Mayor 1991; Raghavan et al. 2010). Over the past 25 years it has been documented that the binary fraction is enhanced among T Tauri stars compared to the field (e.g. Reipurth & Zinnecker 1993; Duchêne & Kraus 2013), and increases for Class I objects (Connelley et al. 2008a,b) and even further for Class 0 objects (Chen et al. 2013). Larson (1972) suggested that all stars may be born in small multiple systems, which decay through dynamical interactions into the mixture of single, binary, and higher-order systems observed among field stars. This idea finds significant support from both observations and numerical simulations (e.g. Reipurth et al. 2014). When a triple system decays, the usual outcome is that the lightest member is expelled into a distant orbit or into an escape, while the remaining newly bound binary shrinks. Since this process in at least

half the cases occurs during the embedded phase, then the hardened binary evolves in a viscous environment leading to orbital decay (Reipurth et al. 2010). The increasingly frequent periastron passages cause accretion events and thus outflow episodes. Reipurth (2000) suggested that the dynamical evolution of newly formed binaries can explain the structure of giant Herbig-Haro (HH) flows, and demonstrated that HH energy sources very often are binaries. We have undertaken a survey of outflow sources and high-accretion objects with the aim to find and study binaries in this population. In the present paper we present a detailed study of the driving source of the little-known HH 250 flow in Aquila, largely based on new, high angular resolution near-infrared observations that show it to be a resolved binary. We also present millimeter and submillimeter interferometric observations of the surroundings of the binary, showing the different distributions of gas and dust that we interpret in terms of a combination of circumstellar and circumbinary disks.

2. Observations and data reduction

2.1. *L*-band imaging

Adaptive optics-assisted imaging in the *L*' (3.8 μm) band were obtained in Service Mode using NACO (Lenzen et al. 2003; Rousset et al. 2003), the adaptive optics-assisted near-infrared

[★] Based on observations obtained with the VLT (Cerro Paranal, Chile) in programs 089.C-0196(A), 095.C-0488(A), and 095.C-0488(B), as well as with IRTF (Mauna Kea, Hawaii), SMA (Mauna Kea, Hawaii), and the Nordic Optical Telescope (La Palma, Canary Islands, Spain).

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Table 1. Summary of observations.

Instrument	Band	Date	Mode
SINFONI (VLT)	<i>H</i> and <i>K</i>	24/25 August 2015	Integral field spectroscopy
NACO (VLT)	<i>L'</i>	18/19 July 2012	Imaging
VISIR (VLT)	<i>N</i>	20/21 July 2015	Imaging
VISIR (VLT)	<i>Q</i>	6/7 August–3/4 September 2015	Imaging
SpeX (IRTF)	0.7 to 4.2 μm	17/18 June 2015	Spectroscopy
SMA	0.9 mm	31 March 2015	Integral field spectroscopy
SMA	1.3 mm	27 February and 6 March 2015	Integral field spectroscopy
EMMI(NTT)	$H\alpha$	18/19 August 1994	Imaging
ALFOSC (NOT)	$H\alpha$	4/5 December 2016	Imaging

imager and spectrograph at the Very Large Telescope (VLT), on the night of 18/19 July 2012. The brightness of HH 250-IRS in the infrared allowed us to use the source itself for wavefront sensing, using a dichroic that directed the light in the *IJK* bands to the wavefront sensor and transmitted light in the *L'* band to the camera. The optics used produced an image scale of $0''.027$ per pixel on the 1024×1024 pixels detector, yielding a field of view of $27'' \times 27''$.

A total of 180 images of 0.175 s exposure time each were obtained and stacked at each point of a dither pattern consisting of 15 positions with a random distribution of telescope pointings within a box $8''$ in size. Flat fields obtained as part of the observatory calibration plan were used in the data reduction. Sky frames were obtained by median combining the dithered images of the field without correcting for telescope offsets. The sky-subtracted, flat-fielded images were then combined into a single frame by coadding the individual images obtained at each telescope pointing, after correcting for the offsets from each position to the next. The final coadded frame thus has an equivalent exposure time of 473 s.

2.2. *H*- and *K*-band integral field spectroscopy

Integral field spectroscopy of HH 250-IRS in the 1.5–2.4 μm range was obtained using SINFONI (Eisenhauer et al. 2003; Bonnet et al. 2004), the adaptive optics-assisted near-infrared integral field spectrograph at the VLT, on the night of 24/25 August 2015 (Table 1). HH 250-IRS is sufficiently bright to be used for tip-tilt correction with SINFONI, and higher order turbulence correction was achieved with the use of the laser guide star facility available at the fourth unit of the VLT where SINFONI is mounted. The image quality thus obtained allowed us to clearly split both components of the binary system, and to extract uncontaminated spectra of each of them. Optics delivering a field of view of $3'' \times 3''$ to the focal plane, and a grism covering the entire *H* and *K* bands at a spectral resolution $R \approx 1500$ were used. Twenty-four exposures of 30 s each were taken by placing the target successively in each of the four quadrants of the integral field unit (IFU) field of view, thus giving a total integration time of 12 min. Since the separation between the components of HH 250-IRS is much smaller than the size of the quadrants this allowed us to confine the target to only one of the quadrants in each exposure, thus removing the need to obtain separate sky exposures. Given that each of the sections of the image slicer of SINFONI partly covers the field of view of two quadrants, the sky background of each exposure was estimated by averaging the exposures obtained two steps before and two steps after it so that there would be no superposition between the (positive) spectral trace of the object and

the (negative) spectral traces of the same object in the subtracted sky frame.

The spectral traces at each slice of the field were determined using dedicated calibration frames that record the trace of a continuum point source for each of the slices. In this way we obtained for each slice the function giving the position of the source along the spatial direction at each wavelength, and defined pseudo-apertures centered on each pixel of the slice along the spatial direction. A wavelength calibration frame was constructed by median-averaging and sigma-clipping all the science frames to remove the spectral traces of the target from the resulting frame, thus leaving the sky lines only. Sky spectra were extracted for each of the pseudo-apertures defined as described above and individually calibrated in wavelength using the OH airglow emission lines as reference (Oliva & Origlia 1992). Spectra of the sky-subtracted science frames were likewise extracted and wavelength calibrated, and the same procedure was followed for an observation of Hip 97787, a B7V star used for telluric feature removal. The extracted spectra of the telluric reference star were combined into a single 1D spectrum, and all the extracted spectra of the science frame were divided by it and multiplied by a blackbody spectrum of 13000 K temperature to provide an approximate relative flux calibration. All the sky-subtracted, wavelength-calibrated, telluric feature-corrected spectra extracted from the science frames were then combined into a single data cube, from which spectra of both components of the binary at arbitrary apertures and images at arbitrarily selected bands could be extracted in a straightforward manner.

Similar observations were carried out for a binary star previously found in the *L'*-band images and suspected to be physically related to HH 250 IRS (see Sect. 4), which was observed on the nights of 1/2 and 5/6 June 2015. The $0''.14$ separation between both components measured in the *L'*-band image, coupled with the absence of a natural guide star bright enough for tip-tilt correction sufficiently close to the position of the target, prevented us from splitting them even with laser-assisted observations in the integral field datacube. Those observations were therefore carried out without using adaptive optics, and the extracted spectrum of this object is thus the blend of both components. The optics used yielded a larger field of view $8''$ across, with a scale of $0''.25$ per pixel. The observation strategy was the same as for HH 250-IRS, now taking 32 observations of 60 s each on both observing nights, thus resulting in 64 min of exposure time.

2.3. *H*-, *K*-, and *L*-band spectroscopy

Observations were carried out with the 3.0 m NASA Infrared Telescope Facility (IRTF) on Maunakea, Hawaii on the night

of 17/18 June 2015, using the recently upgraded SpeX (Rayner et al. 2003) instrument in the short cross-dispersed (SXD) mode and the long cross-dispersed (short grating setting) mode (LXD), which together cover $0.7 \mu\text{m}$ to $4.2 \mu\text{m}$. The total exposure time in SXD was 24 min for HH 250 IRS and the same for the nearby source 2MASS 19212367+1054108. The total exposure time in LXD was 12.8 min for each of these two sources. An A0 telluric standard star was observed in both modes for telluric correction, within 0.1 air masses of the target. An argon lamp was observed for wavelength calibration and a quartz lamp for flat fielding. An arc/flat calibration set was observed for each target/standard pair. The SpeX data were flat fielded, extracted, and wavelength calibrated using *SpeXtool* (Cushing et al. 2004).

2.4. *N*- and *Q*- band imaging

N- and *Q*-band imaging of HH 250-IRS was obtained with VISIR (Lagage et al. 2004), the thermal infrared camera and spectrograph at the VLT, on the nights of 20/21 July 2015 (*N*-band) and on that same night, and also on 6/7 August and 3/4 September 2015 (*Q* band). The actual filters used were the B10.7, centered on $\lambda = 10.65 \mu\text{m}$ with a full width at half-maximum of $1.37 \mu\text{m}$; and the Q2 filter centered on $\lambda = 18.72 \mu\text{m}$ with a full width at half-maximum of $0.88 \mu\text{m}$. The images were obtained using the chopping and nodding technique to remove the effect of the rapidly changing sky emission and of the difference in the telescope thermal signature pattern due to the slight change of position of the secondary when chopping. A chop throw of $4''$ was used, which kept the target within the $38'' \times 38''$ field of view at all times. Only one chop and nod cycle was obtained, with a chopping frequency of 4 Hz. The total exposure time was 60 s in the *N* band, and 90 s for each of the observations in the *Q* band. The images produced by the imaging data reduction pipeline could be directly used for our purposes.

2.5. *H α* imaging

Images of HH 250 and its environment through *H α* and [SII] filters, as well as through a broadband *R* filter, were obtained using the EMMI imager and spectrograph at the ESO New Technology Telescope (NTT) at La Silla Observatory on 18 August 1994 and can be found at the ESO Science Archive. The years elapsed since those images were taken provide an excellent opportunity to investigate changes having taken place in the region in the intervening time. For this purpose we obtained a ten-minute exposure with the ALFOSC imager and spectrograph at the 2.5 m Nordic Optical Telescope at the Roque de los Muchachos Observatory on the night of 4/5 December 2016, also using an *H α* filter, which provides a depth similar to that obtained with EMMI at the NTT in 1994.

2.6. Submillimeter imaging

SMA observations toward HH 250-IRS were conducted with the extended configuration at 225 GHz on 27 February and 6 March 2015 and at 351 GHz on 31 March 2015. The pointing center was $\alpha(J2000) = 19^{\text{h}}21^{\text{m}}23^{\text{s}}.1$, $\delta(J2000) = 10^{\circ}54'04''$. In these observations, six antennae were used. 3C279, Titan, and 1751+096 were observed as bandpass, flux, and gain calibrators, respectively. The flux of 1751+096 at 225 GHz was 2.15 Jy on 27 February 2015 and 2.29 Jy on 6 March 2015, and that at 351 GHz was 1.97 Jy on 31 March 2015. The typical systemic

Table 2. Summary of SMA Imaging Parameters.

	Synthesized beam	Noise (mJy Beam ⁻¹)
1.3 mm	1''.4 × 0''.9(106°)	1.3
0.9 mm	0''.8 × 0''.5(100°)	3.9
CO (2–1)	1''.3 × 0''.8(105°)	150 ^a
CO (3–2)	0''.9 × 0''.5(103°)	210 ^a

Notes. ^(a) Noise per channel with a velocity width of 0.5 km s^{-1} .

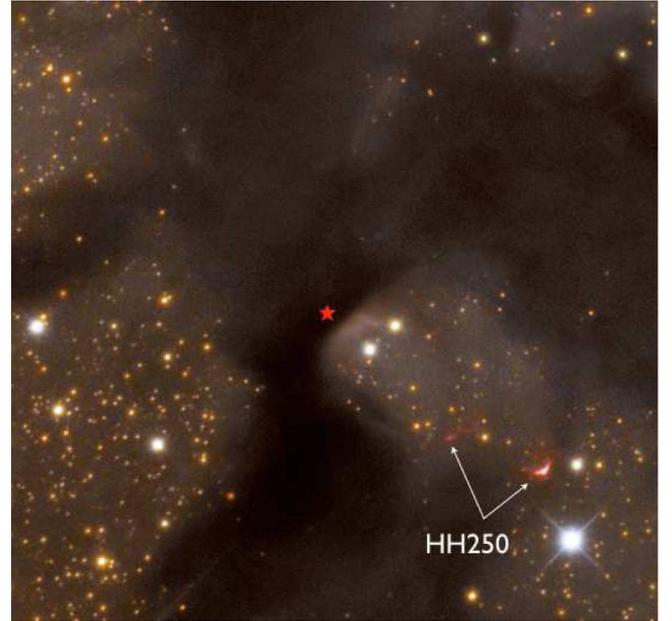


Fig. 1. Optical (*BGR*) overview of the star forming region centered on HH 250 IRS = IRAS 19190+1048. The red asterisk marks the embedded source, and two components of the HH 250 flow are marked (A is the brighter knot, B is the fainter one). The image is $5''.7$ across, with north to the top and east to the left. Image courtesy Adam Block.

temperature of the observations was 120–300 K on 27 February 2015, 110–180 K on 6 March 2015, and 210–640 K on 31 March 2015. CO (2–1; 230.538 GHz) and 1.3 mm continuum were observed simultaneously in the 225 GHz observations, and CO (3–2; 345.796 GHz) and 0.9 mm continuum in the 351 GHz observation. 512 channels were assigned to a chunk with a bandwidth of 104 MHz for the CO lines. All the data were calibrated using the MIR software package (Scoville et al. 1993). The calibrated visibility data were Fourier-transformed with natural weighting and CLEANed with MIRIAD (Sault et al. 1995) to produce images. The CO images were generated at a velocity resolution of 0.5 km s^{-1} . The angular resolutions and noise levels of our continuum and CO images are listed in Table 2.

3. Environment and Herbig-Haro flow

HH 250-IRS, also known as IRAS 19190+1048 (2MASS J19212262+1053569, WISE J192122.62+105357.0) is located in a compact cloud core in the southern part of the large highly structured cloud complex Lynds 673 (Fig. 1; for a review of star formation in the Aquila clouds see Prato et al. 2008). The galactic coordinates of the source are $l = 46.03$, $b = -1.58$, which places it in Cloud B that is located at higher galactic longitudes

than the large Aquila Rift (Dame & Thaddeus 1985). Cloud B appears to be a small extension of the Aquila Rift, because of continuity in location and velocity, which places it at a distance of about 300 pc (Dame et al. 1987; Sakamoto et al. 1997). Straizys et al. (2003) determined a distance of ~ 270 pc to the core of the Aquila Rift, with a distance to the cloud face of ~ 225 pc. Prato et al. (2003) adopted a distance of 150 pc to the nearby T Tauri star AS 353A. Recently Ortiz-León et al. (2017) found a distance to the neighboring Serpens/W40 complex of 436 ± 9 pc, suggesting that the Aquila Rift is more distant at lower longitudes. Evidently the distances to objects at this galactic longitude are poorly known, in the following we adopt the distance of 225 pc by Straizys et al. (2003) while recognizing the uncertainty in this value.

Devine et al. (1997) discovered a bright compact Herbig-Haro object, labeled HH 250A, near IRAS 19190+1048, and just 14 arcmin southeast of the famous HH 32 (Herbig & Jones 1983). HH 250A, located at $\alpha_{2000} = 19:21:13.8$, $\delta_{2000} = +10:52:30$, has a clear bow shock morphology, facing away from the IRAS source and is located in what appears to be an illuminated outflow cavity stretching away from it (see Fig. 1). A fainter knot (B) is located about $50''$ further to the ENE and near the line joining HH 250A and the IRAS source. At an angular distance of $157''$, IRAS 19190+1048 is the nearest IRAS source to HH 250A and the only embedded source in the cloud. The second nearest source, IRAS 19188+1050, is located $4:2$ away, slightly outside the dark cloud in a direction that is almost perpendicular to the axis of the concavity of HH 250A, and the third nearest source is over $9:1$ away, outside the cloud and also in a direction far removed from that axis. Given the close proximity to HH 250-IRS, its location close to the direction defined by the axis of the HH 250A bow shock, the morphology of the dark cloud in the area, and the lack of other suitable candidates in the area, we consider that IRAS 19190+1048 is the only realistic candidate driving source of HH 250A and B.

A comparison between the images obtained in 1994 and 2016 is presented in Fig. 2, where some obvious changes can be easily seen. Figure 2 shows that the diffuse (presumably reflection) nebulosity NW of HH 250-IRS seen in 1994 has become much fainter or disappeared in 2016, whereas HH 250B has become brighter in the meantime. Another area of nebulosity, close to the embedded IRAS source and near the eastern edge of the image, was barely visible in 1994 but has gained prominence 22 years later. These changes point to a dynamic environment near HH 250-IRS, with dusty clumps in its immediate surroundings causing different parts of the region being alternatively exposed to, or shadowed from, direct illumination by the central source.

Dramatic changes affect HH 250A itself, as clearly seen in the lower panels of Fig. 2. The arc shape is preserved, but a bright knot has developed near the apsis of the arc in the intervening years and now dominates the appearance of the object. Other subtler changes in the bright emission can be seen, including the appearance of a faint knot downstream along the eastern branch of the arc. The changes in the structure of HH 250A make it difficult to precisely determine its proper motion, but a bulk displacement of the whole structure $\approx 1:0$ to the southwest can be measured. This translates into $\approx 0:046/\text{yr}$, or $\approx 50 \text{ km s}^{-1}$ at the adopted distance of the region. The direction of the proper motion away from IRAS 19190+1048 further confirms the identity of the latter as the driving source of HH 250, and the angular distance between both indicates that the Herbig-Haro object is the result of an outburst having taken place about 3500 years ago.

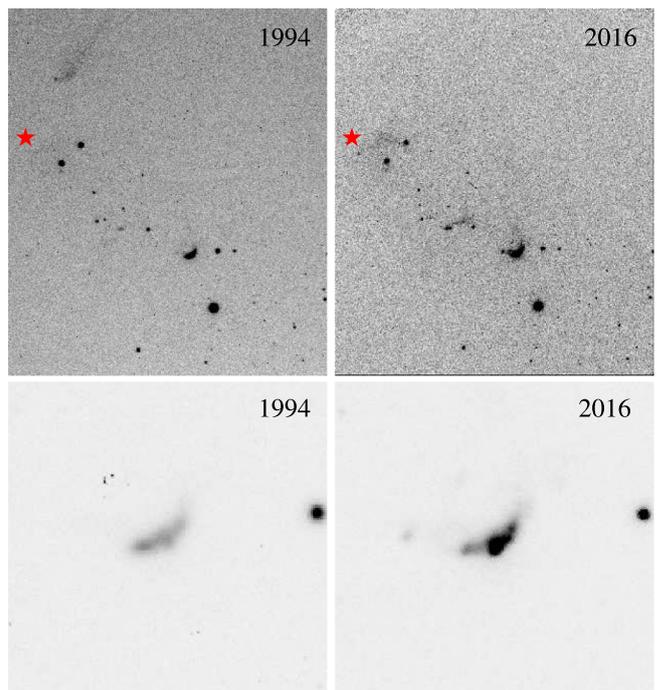


Fig. 2. *Upper panels:* HH 250 region imaged with an $H\alpha$ filter at two different epochs 22 years apart. The 1994 images were obtained with the NTT/EMMI, and the 2016 ones with the NOT/ALFOSC. The field is $4:0 \times 4:5$ across. The reflection nebula seen to the NNW of the embedded source has disappeared in the intervening time, whereas nebulosity just E of the source has brightened in the most recent image. The red triangle marks the position of the embedded source HH 250-IRS. *Lower panels:* close-up view of the bow shock reveals a dramatic brightening of the $H\alpha$ emission between both epochs, as well as other changes in its structure. The field is $41'' \times 41''$ across. The grayscale of the *lower panels* has been stretched with respect to the upper ones to better show structure in the bow shock.

4. HH 250-IRS

HH 250-IRS, the powering source of HH 250, appears resolved into a binary in all our images. The northwestern and southeastern components, hereafter referred to as HH 250-IRS NW and HH 250-IRS SE respectively, are separated by an angular distance of $0:53$ with the SE component located at a position angle of 142° with respect to the NW one. At the adopted distance the angular separation corresponds to a projected distance of 120 AU between both components.

4.1. Infrared spectroscopy and photometry

The $1.5\text{--}2.4 \mu\text{m}$ spectra of both components of HH250-IRS are similar and typical of deeply embedded Class I or flat spectrum sources. They are essentially featureless, with the exception of hints of weak CO absorption longward of $2.293 \mu\text{m}$ in the SE component, moderately intense emission lines in both spectra of H_2 at $2.122 \mu\text{m}$, and weaker but also clear H_2 emission at $2.04 \mu\text{m}$. The non-detection of H_2 at $2.248 \mu\text{m}$ despite the strength of the line at $2.122 \mu\text{m}$ indicates that the emission is produced by non-dissociative shocks rather than fluorescence. The images obtained by integrating the data-cubes in a narrow spectral range around $2.12 \mu\text{m}$ show no traces of extended structure, suggesting that the H_2 emission probably arises either from an unresolved region at the base of an outflow, or from an accretion shock. The weak or absent CO absorption is suggestive of

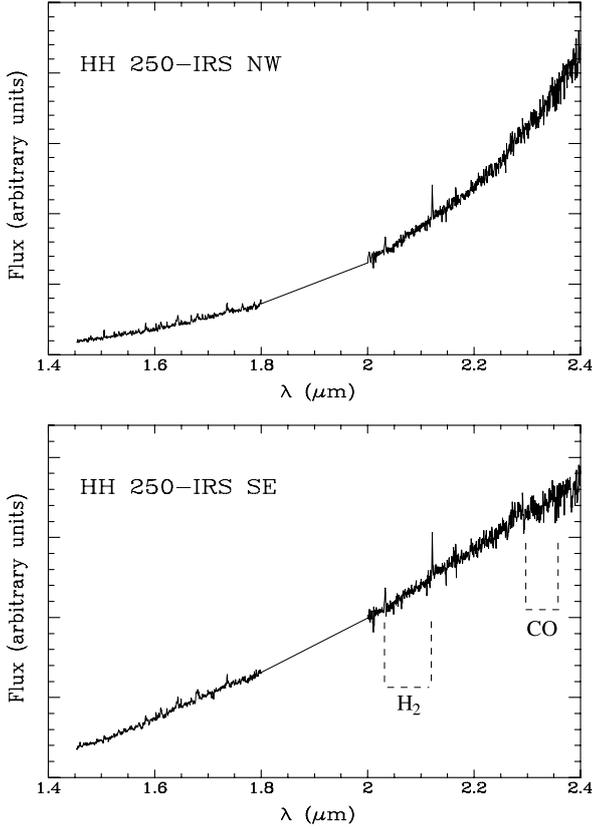


Fig. 3. *H*- and *K*-band SINFONI spectra of both components of HH 250-IRS. Both sources display a heavily reddened and veiled continuum with emission in the H_2 line at $2.122 \mu\text{m}$ as the only outstanding feature. HH 250-IRS SE, the least reddened of both components, also shows hints of CO absorption starting at $2.293 \mu\text{m}$.

a strongly veiled cool photosphere. The SpeX spectrum in the $2.8\text{--}4.2 \mu\text{m}$ range is dominated by the very strong $3 \mu\text{m}$ absorption feature which indicates a substantial amount of water ice along the line of sight (e.g. [Graham 1998](#)). The $3.45 \mu\text{m}$ depression is also clearly visible, it is ascribed to saturated aliphatic hydrocarbons in grains ([Pendleton et al. 1994](#)) and correlates closely with water ice ([Brooke et al. 1999](#)). Two emission lines are visible, the Pf γ line at $3.74 \mu\text{m}$ and the Br α line at $4.05 \mu\text{m}$.

Figure 5 shows the binary system in all the bands in which we observed it, between $1.5 \mu\text{m}$ and $19 \mu\text{m}$. The figure makes obvious that the NW component has overall redder colors than the SE, but the brightness ratio between both components does not follow a monotonical trend with wavelength. While the SE component dominates in the *H* and *K* bands and yields to the NW component at longer wavelengths, the long wavelength dominance of the NW component is greatly reduced in the *N* band.

We have measured individual magnitudes in the *H* and *K* bands by performing photometry on an image obtained by integrating the SINFONI datacube over the wavelength ranges covered by the 2MASS filters in each of those bands. The individual magnitudes of each component were obtained by adding up their fluxes and assuming the sum to be equal to the flux derived from the combined magnitude listed in the 2MASS All-Sky Point Source Catalog ([Skrutskie et al. 2006](#)). A similar comparison was made to derive the *L'* magnitudes, now using the instrumental photometry of the NACO image in that band and comparing it to the flux in the *W1* band listed in the WISE

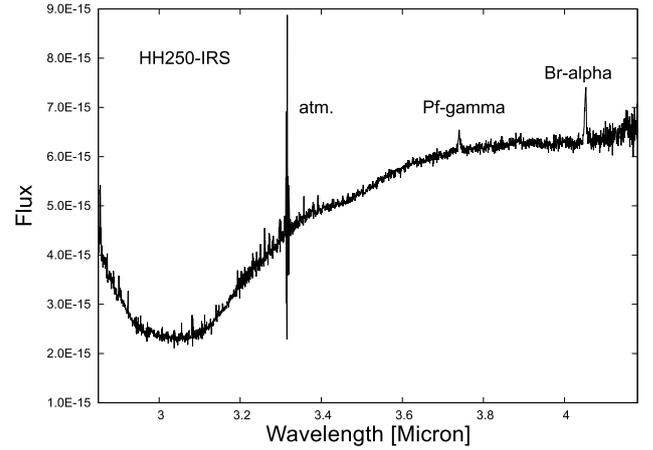


Fig. 4. Combined SpeX spectrum of both components of HH 250-IRS covering the $2.85 \mu\text{m}$ to $4.18 \mu\text{m}$ spectral region. The Pf γ line at $3.7406 \mu\text{m}$ and the Br α line at $4.0523 \mu\text{m}$ are prominently in emission. Strong water ice dominates the $3 \mu\text{m}$ region of the spectrum.

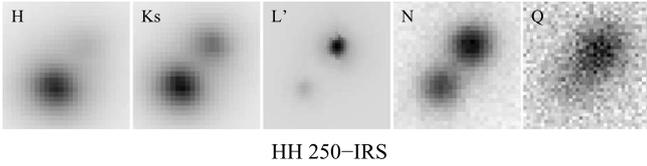
All-Sky Point Source catalog ([Wright et al. 2010](#)). For the VISIR observation in the *N* band an observation of the star β CrA carried out shortly before the observation of HH 250-IRS was taken as a reference, using the published magnitude of that star ($W3 = 1.348$) as a reference. Finally, for the *Q*-band magnitude we used the WISE catalog values, correcting for the fact that the *Q* filter used with VISIR has a central wavelength lying between those of the *W3* and *W4* bands. The correction was made assuming that the combined flux of both components of HH 250-IRS has a power-law dependence with wavelength, $f \propto (\lambda/\lambda_0)^\alpha$. We thus obtained an interpolated magnitude for the combined components $Q = W3 - 0.741(W3 - W4)$. The noisy VISIR *Q*-band image and the proximity of the diffraction limit of the VLT to the separation between both components of the binary prevented us from carrying out an accurate derivation of the respective contributions, but a rough two-Gaussian fit yielded a magnitude difference $\Delta Q \approx 1.6$ between both components. Table 3 lists the magnitudes thus obtained in each band.

With the exception of the *N*-band measurement, which was calibrated through the contiguous observation of a star of constant brightness, our derivation of fluxes in all the other bands relies on the cataloged combined fluxes of both components of HH250-IRS, and are thus in principle subjected to uncertainties caused by possible variability. However, the cataloged variability likelihood of HH250-IRS is low in all the WISE bands and consistent with a source of constant brightness, thus ruling out significant errors due to variability in the *L'*, *N* and *Q* bands. Since the near-infrared spectra of both components is likely to be dominated by the emission of the circumstellar envelope rather than by photospheric emission, the non-variability of the emission at the longer wavelengths suggests that the *H* and *K_S* fluxes are also approximately constant. This is also expected on statistical grounds, as the fraction of sources with variability amplitudes exceeding 0.5 mag among disk-bearing young stellar objects in nearby star forming regions does not exceed 3% ([Scholz 2012](#)). We thus consider that the bases of the flux calibrations presented here are reliable, and that the features of the spectral energy distributions discussed below are unaffected by the non-simultaneity of our observations.

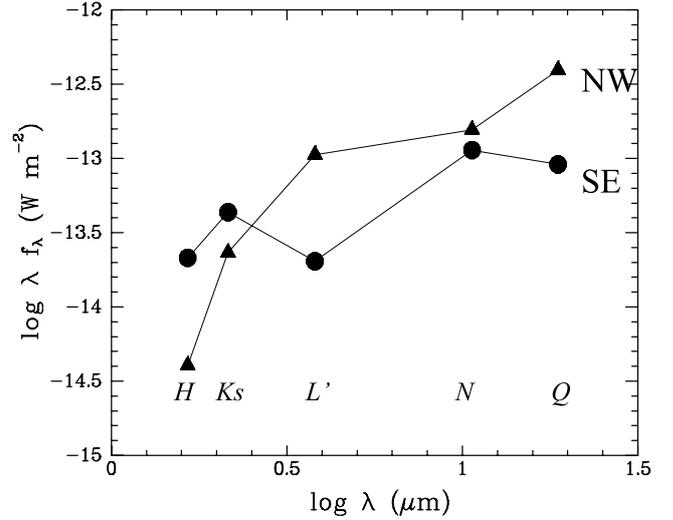
The spectral energy distributions of both components of HH 250-IRS are presented in Fig. 6. While both sources show an overall rising shape with wavelength that characterizes them

Table 3. Magnitudes of HH 250-IRS.

Band	HH 250-IRS NW	HH 250-IRS SE
<i>H</i>	14.15 ± 0.1	12.34 ± 0.05
<i>K_S</i>	11.50 ± 0.05	10.82 ± 0.05
<i>L'</i>	8.18 ± 0.05	9.98 ± 0.1
<i>N</i>	4.55 ± 0.1	4.89 ± 0.1
<i>Q</i>	1.7 ± 0.2	3.3 ± 0.3


Fig. 5. Images of the resolved HH 250-IRS binary from $1.5 \mu\text{m}$ to $19 \mu\text{m}$, showing the predominance of the SE components shortward of $3 \mu\text{m}$ and of the NE component at longer wavelengths. *H* and *K_S* images were obtained with SINFONI, *L'* with NACO, and *N* and *Q* with VISIR. The separation between both components is $0''.53$. North is to the top and east to the left.

as Class I sources, in consistency with their spectra, clear differences already inferred from the sequences of images shown in Fig. 5 appear between both. Overall HH 250-IRS NW is more deeply embedded probably due to the presence of a more massive circumstellar envelope around it. On the other hand, the *N*-band fluxes hint at the presence of strong silicate spectral features in that region appearing in absorption in the NW component and in emission in the SE one. This is consistent with the overall trend observed in the Class I/Class II sequence (Kessler-Silacci et al. 2005), where the strong and broad absorption produced by amorphous silicates near $9.7 \mu\text{m}$ gradually turns into emission as the circumstellar dust distribution evolves from an envelope to a disk configuration and the spectral energy distribution turns bluer as the central object progressively emerges. It is thus intriguing that the two components of the HH 250-IRS binary system, being essentially coeval, display differences in their spectral energy distributions that are normally associated with the evolutionary stages of their circumstellar environments. The differences between both spectral energy distributions are much above what may be accounted for by the magnitude uncertainties. Furthermore, whereas HH 250-IRS NW displays a clear excess flux in the *L'* band as compared to that expected from the extrapolation of the *H* and *K_S* fluxes if produced by a reddened photosphere, the *HK_SL'* colors of HH 250-IRS SE are consistent with a reddened photosphere with little infrared excess, which manifests itself only at longer wavelengths. This points to a lack of warm dust in the immediate vicinity of the central source of HH 250-IRS SE, hinting at the existence of a dust-free cavity around the central star of HH 250-IRS SE that has no analog in the envelope of the NW component. Large inner holes in the distribution of dust are the hallmark of transition disks, which are usually thought to be caused by circumstellar dust evolving from optically thick to optically thin due to grain growth, formation of planetesimals, and eventual clearance of dust from the inner regions of disks by newly formed planets (e.g. Cieza et al. 2010). It would be surprising that such processes had had time to operate in a system as young as HH 250-IRS, and we prefer instead to relate the existence of the gap in the inner disk of HH 250-IRS SE to the presence of the NW binary component, as we describe in Sect. 5.


Fig. 6. Spectral energy distributions of both components of the HH 250-IRS binary system.

4.2. Cold material around HH 250-IRS

4.2.1. CO emission

Figure 7a presents the spectra of the CO (2–1) and (3–2) emission lines observed with the SMA. The spectra are extracted at the continuum peak position and averaged over the region within a radius of $0''.5$. The area for averaging corresponds to the area of the synthesized beam of the CO (2–1) image. Two velocity components are detected in the CO spectra, one centered at $V_{\text{LSR}} \sim 4.6 \text{ km s}^{-1}$ (hereafter blueshifted component) and one at $V_{\text{LSR}} \sim 10.7 \text{ km s}^{-1}$ (hereafter redshifted component). Figure 8 shows the integrated intensity maps of the two velocity components detected in the CO lines. In both the CO (2–1) and (3–2) emission, the redshifted component is extended and elongated along the northwest-southeast direction with a position angle of $\sim 150^\circ$. This orientation is perpendicular to the outflow direction and matches well the line between components SE and NW of HH 250 IRS. The blueshifted component is not resolved and has a size comparable to the angular resolution. Because of the limited signal-to-noise ratios, the velocity structures of the CO (2–1) and (3–2) emission of both components are barely revealed.

To compare the brightness temperature T_b of the CO (2–1) and (3–2) emission we convolved the CO (3–2) image with the angular resolution of the CO (2–1) image after correcting for primary beam attenuation. The spectra of the CO (2–1) and (3–2) emission extracted at the same angular resolutions are presented in Fig. 7b. By fitting a Gaussian line profile to the spectra, we measured their peak brightness temperatures, the centroid velocities in the LSR frame, and the FWHM line widths. The results of the fit are listed in Table 4. For each velocity component, we assume that the CO (2–1) and (3–2) emission lines trace the same bulk of gas, although there is an offset of 0.5 km s^{-1} in the centroid velocities of their redshifted component. We then used the radiative transfer code RADEX (van der Tak et al. 2007) to simultaneously fit the observed brightness temperatures in both CO transitions with the observed line widths. We performed spectral fits with various fixed values of the volume density n_{H_2} and varied the kinetic temperature T_k and the column density N_{CO} to obtain a predicted T_b to be compared with the actually observed values, thus obtaining the best-fit excitation temperature T_{ex} , optical depth

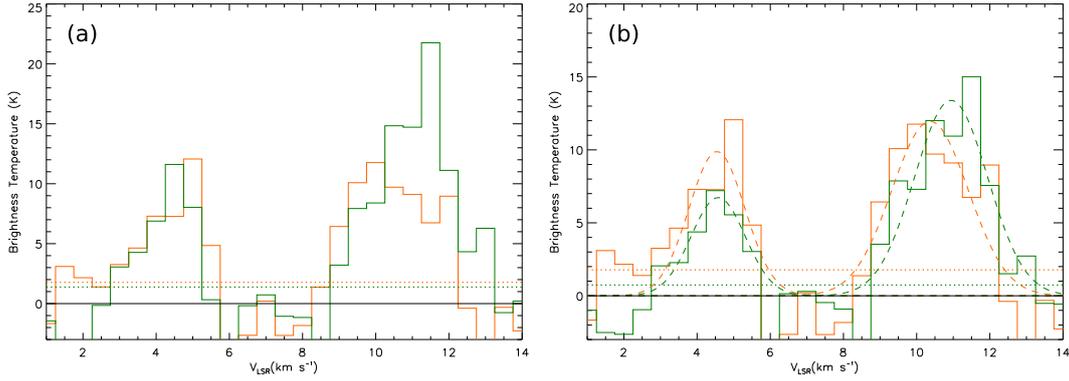


Fig. 7. CO spectra obtained with our SMA observations. Orange and green histograms present the CO (2–1) and (3–2) emission, respectively. The spectra in both *panel a* and *b* were extracted by averaging over the region with a radius of $0''.5$ centered on the continuum peak position from the images with the correction of the primary beam attenuation. Dashed lines denote 1σ noise levels of the spectra. In *panel b*, the CO (3–2) spectrum was extracted after convolving the CO (3–2) image with the same angular resolution of the CO (2–1) image. Red and blue curves show the Gaussian fitting results of the line profiles.

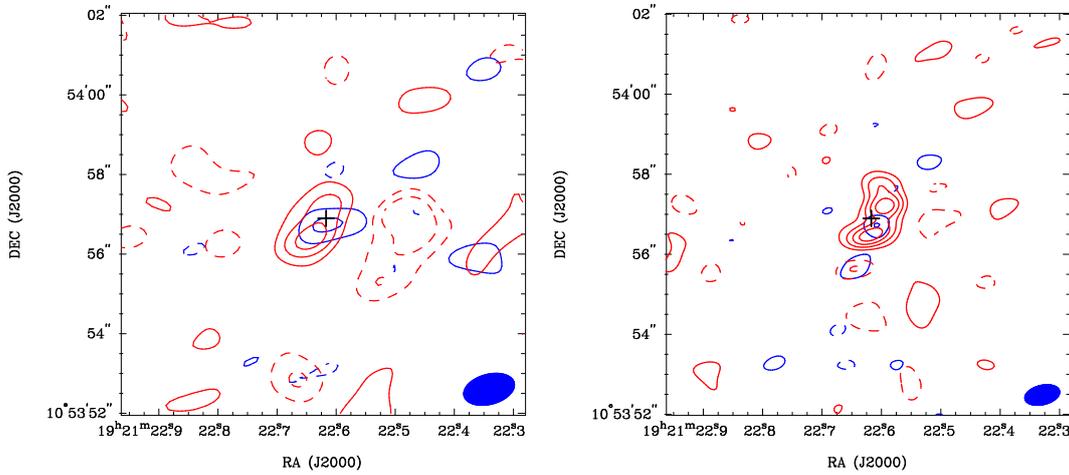


Fig. 8. Integrated intensity maps of the CO (2–1; *left*) and (3–2; *right*) emission obtained with our SMA observations. Blue and red contours present the blueshifted and redshifted velocity components centered at $V_{\text{LSR}} \sim 4.6 \text{ km s}^{-1}$ and $V_{\text{LSR}} \sim 10.7 \text{ km s}^{-1}$, respectively. The integrated velocity widths of the blueshifted and redshifted components are 1.5 km s^{-1} and 3.5 km s^{-1} in the CO (2–1) emission, and those are 1 km s^{-1} and 3 km s^{-1} in the CO (3–2) emission. Crosses denote the continuum peak position, and filled blue ellipses show the sizes of the synthesized beams. Contours are from 3σ in steps of 2σ . 1σ noise levels in the CO (2–1) maps of the blueshifted and redshifted components are 0.14 and $0.2 \text{ Jy beam}^{-1} \text{ km s}^{-1}$, and those are 0.15 and $0.24 \text{ Jy beam}^{-1} \text{ km s}^{-1}$ in the CO (3–2) maps.

τ , and T_k . We considered $n_{\text{H}_2} > 10^5 \text{ cm}^{-3}$, typical of the average gas density on a 100 AU scale around young stellar objects. The gas temperature T_k is estimated to be 13–18 K for the blueshifted component and 24–40 K for the redshifted component. The CO lines in both components are thermalized. The CO (2–1) and (3–2) lines in the blueshifted component are both optically thick, yielding only a lower limit to the CO column density. In the redshifted component, the CO (2–1) line is optically thin to marginally optically thick, and the CO (3–2) line is close to optically thick. Its CO column density is estimated to be $(2 - 3) \times 10^{16} \text{ cm}^{-2}$.

4.2.2. Continuum emission

Figure 9 presents the 1.3 mm and 0.9 mm continuum images of HH 250-IRS. The peak position is measured to be $\alpha(2000) = 19:21:22.62$, $\delta(2000) = 10:53:56.9$ at both wavelengths, which is within about $0''.1$ of the 2MASS position. The apparent full-width-half-maximum (FWHM) sizes of the continuum emission are measured to be $1''.4 \times 0''.9$ at 1.3 mm and

$0''.8 \times 0''.6$ at 0.9 mm by fitting a two-dimensional Gaussian distribution with the MIRIAD task *imfit*. The observed sizes are consistent with the angular resolutions, and the deconvolution by *imfit* results in a point source. By fitting a point source to the visibility data, the flux is measured to be $17 \pm 1.4 \text{ mJy}$ at 1.3 mm and $81.5 \pm 5.6 \text{ mJy}$ at 0.9 mm. The spectral index α of the continuum emission between 1.3 mm and 0.9 mm is derived to be 3.5 ± 0.2 , and the frequency dependence of the dust opacity $\beta = \alpha - 2$ is 1.5 ± 0.2 .

The mass of the circumstellar dust traced by the continuum emission (M_{dust}) can be estimated as

$$M_{\text{dust}} = \frac{F_{\lambda} d^2}{\kappa_{\lambda} B(T_{\text{dust}})}, \quad (1)$$

where F_{λ} is the continuum flux at the wavelength λ , d is the distance to the source, κ_{λ} is the dust mass opacity at the wavelength λ , T_{dust} is the dust temperature, and $B(T_{\text{dust}})$ is the Planck function at a temperature of T_{dust} . By adopting the frequency function for dust mass opacity, $\kappa_{\lambda} = 10 \times (0.3 \text{ mm}/\lambda)^{\beta} \text{ cm}^2 \text{ g}^{-1}$

Table 4. Gaussian fitting to the CO (2–1) and (3–2) spectra.

Line	Blueshifted component			Redshifted component		
	T_b (K)	V_c (km s ⁻¹)	ΔV (km s ⁻¹)	T_b (K)	V_c (km s ⁻¹)	ΔV (km s ⁻¹)
CO (2–1)	9.9 ± 1.2	4.5 ± 0.1	1.8 ± 0.3	11.8 ± 1.1	10.4 ± 0.1	2.5 ± 0.3
CO (3–2)	6.7 ± 0.6	4.6 ± 0.1	1.8 ± 0.2	13.4 ± 0.5	10.9 ± 0.1	2.4 ± 0.1

Table 5. Gas temperature estimated from the CO (2–1) and (3–2) spectra.

n_{H_2} (cm ⁻³)	T_k (K)	N_{CO} (cm ⁻²)	CO (2–1)			CO (3–2)		
			T_{ex} (K)	τ	T_b (K)	T_{ex} (K)	τ	T_b (K)
			Blueshifted component					
10 ⁵	15 ⁺⁴ ₋₁	>1.2 × 10 ¹⁶	15 ⁺³ ₋₁	>0.9	9.0 ⁺⁰ _{-1.0}	14 ⁺⁴ ₋₁	>1.1	6.9 ⁺⁰ _{-0.7}
≥10 ⁶	14 ⁺⁴ ₋₁	>1.3 × 10 ¹⁶	14 ⁺⁴ ₋₁	>1	8.8 ⁺⁰ _{-0.8}	14 ⁺⁴ ₋₁	>1.1	7 ^{+0.2} _{-0.7}
			Redshifted component					
10 ⁵	32 ⁺⁸ ₋₆	(2.3 ^{+0.8} _{-0.3}) × 10 ¹⁶	31 ⁺⁸ ₋₆	0.6 ^{+0.5} _{-0.2}	11.7 ^{+1.6} _{-1.3}	30 ± 6	0.9 ^{+0.6} _{-0.3}	13.4 ^{+0.3} _{-0.1}
≥10 ⁶	28 ⁺⁸ ₋₄	(2.5 ^{+0.8} _{-0.4}) × 10 ¹⁶	28 ⁺⁸ ₋₄	0.7 ^{+0.5} _{-0.3}	11.8 ^{+1.4} _{-1.5}	28 ⁺⁸ ₋₄	1.1 ^{+0.2} _{-0.4}	13.3 ± 0.1

Notes. n_{H_2} is the H₂ number density, T_k is the gas kinematic temperature, N_{CO} is the CO column density, T_{ex} is the excitation temperature, τ is the optical depth, and T_b is the peak brightness temperature. In our spectral fitting with the radiative calculation, T_k and N_{CO} are the free parameters, and n_{H_2} is a fixed parameter.

(Beckwith et al. 1990), we obtain $\kappa_{1.3\text{mm}} = (1.1 \pm 0.3) \text{ cm}^2 \text{ g}^{-1}$ and $\kappa_{0.9\text{mm}} = (1.9 \pm 0.6) \text{ cm}^2 \text{ g}^{-1}$ with the measured β value¹.

Given the coincidence of the peak of the continuum emission with the peak of the blueshifted component of the CO emission, both being unresolved, it appears reasonable to assume that both trace the same structure. Therefore we adopt the derived temperature of the dust as 14 K, assuming that dust and gas are thermally well coupled. In this way, M_{dust} is estimated to be $(1.6 \pm 1.1) \times 10^{-4} M_{\odot}$ from the 1.3 mm measurement, and $(2.1 \pm 0.9) \times 10^{-4} M_{\odot}$ from the 0.9 mm measurement. Assuming a standard gas-to-dust ratio of 100 by mass (Draine et al. 2007) this points to a rather massive envelope in excess of $10^{-2} M_{\odot}$. However, the crude assumption of a single dust temperature, existing evidences for gas depletion in circumstellar disks (Ansdell et al. 2016), and the lack of an independent estimate of the gas mass of the envelope through an optically thin tracer refrains us from making any definitive claims in this regard.

5. Discussion

5.1. A circumbinary disk around HH 250-IRS

The distinct kinematical components detected in CO emission, together with their different spatial distributions and temperatures, suggest that they trace different structures in the HH 250-IRS system. The CO emission is clearly elongated in the direction defined by both components of the binary, and shows hints of a double-peaked spatial distribution of the intensity in the CO (3–2) map with both peaks separated by an angular distance

¹ It should be noted that we use the opacity per unit dust mass, rather than the more commonly used opacity per unit total mass, since we only estimate the dust mass here.

of $\sim 1''$ that exceeds the angular distance between the components of the binary. This leads us to assume that the CO emission associated with the redshifted component probably comes from an extended, torus-shaped circumbinary envelope surrounding both members of the HH 250-IRS system, with a dimension along the diameter of the torus of ~ 250 AU. The unresolved blueshifted gas component that we attributed in Sect. 4.2.2 to a circumstellar envelope around one of the components of HH 250-IRS seems most likely related to HH 250-IRS NW, the most embedded member of the system displaying the strongest excess at mid-infrared wavelengths. However, the close coincidence between the peak position of the continuum emission with the 2MASS position rather points to HH 250-IRS SE, the brightest member of the pair at the 2MASS bands, as the dominant source at millimeter wavelengths. Higher-resolution observations with accurate astrometry are needed to clarify the association.

A crude estimate of the mass of the circumbinary disk can be obtained from the CO column density derived from the spectral fit (Table 5), integrating it over the angular area of the redshifted component (3.3 arcsec^2), and assuming a H₂-to-CO ratio $N_{H_2}/N_{CO} \approx 10^4$ (France et al. 2014) and a mean molecular weight of $2.33m_H$. In this way we obtain $M \sim 1.7 \times 10^{-5} M_{\odot}$. The non-detection of continuum emission associated with the circumbinary ring sets in turn an upper limit of $M_{\text{dust}} \lesssim 10^{-5} M_{\odot}$. Given the estimated gas mass, the non-detection of the circumbinary ring in the continuum is thus expected for a normal gas-to-dust ratio of 100.

Binary systems where both circumstellar and circumbinary disks coexist are a predicted outcome of the fragmentation process leading to multiple star formation when the accreting material possesses a high angular momentum

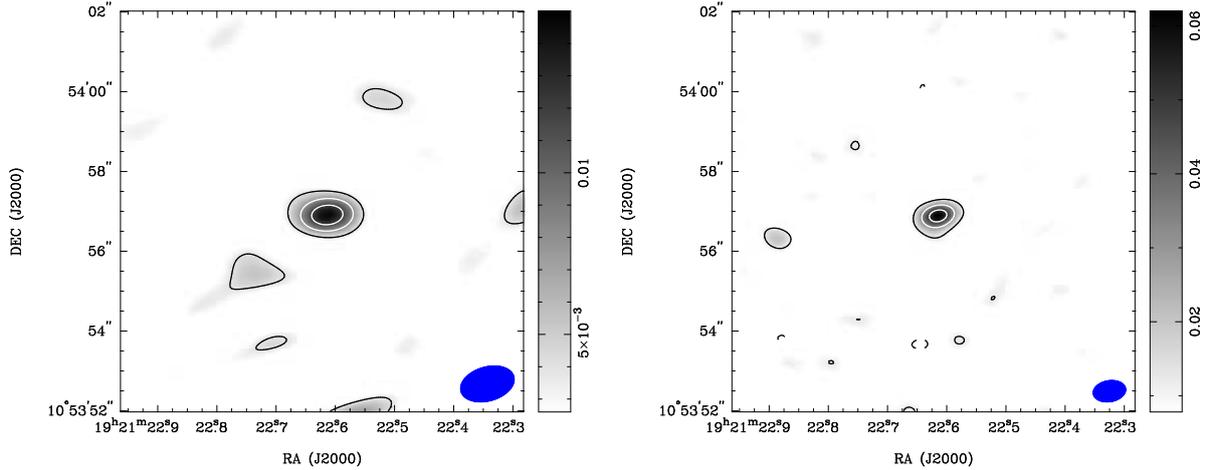


Fig. 9. 1.3 mm (left) and 0.9 mm (right) continuum images of HH 250-IRS A obtained with our SMA observations. The images are not corrected for the primary beam attenuation. Contour levels start from 3σ in steps of 3σ in the 1.3 mm image and in steps of 5σ in the 0.9 mm image, where 1σ is 1.3 and $3.9 \text{ mJy Beam}^{-1}$, respectively. Filled blue ellipses show the sizes of the synthesized beams.

(Bate & Bonnell 1997). The circumstellar disks around each component are truncated due to dynamical interaction with the companion, and their long lifetimes inferred from observed systems requires their sustained feeding through accretion streams from the circumbinary disk that acts as a reservoir (Mathieu et al. 2000). The material accreted by each member of the binary greatly depends on the mass ratio of the protostellar seeds and the specific angular momentum of both the accreting gas and the secondary. This can lead to widely different accretion paths for both components of the binary, making possible that only one of the components is deeply embedded in a circumstellar envelope. The remarkably different spectral energy distributions of HH 250-IRS NW and SE, and the detection of only one of the circumstellar envelopes in our observations, probably reflects this.

Simulations of the dynamical interaction between a binary system and a massive circumbinary disk show that the latter strongly influences the orbital characteristics of the former (Artymowicz et al. 1991), with the production of a highly eccentric system as one of the main outcomes. This leads to the expectation that HH 250-IRS NW and SE may undergo periodic episodes of proximity much closer than the present 120 AU of projected distance, causing major disturbances to their respective disks and possibly major events of mass loss. It is possible that a past periastron passage may have been responsible for the formation of HH 250. Assuming a combined mass of $\sim 1 M_{\odot}$ for the binary and a semimajor orbital axis of half the current projected separation of 120 AU, we obtain an orbital period of ~ 500 years, which within the significant uncertainties compares quite favorably with the kinematic age of 3,500 yr of HH 250 (Sect. 3) that sets an upper limit to the time since the last outburst. We speculate that major perturbations in the circumstellar disk structure might also be responsible for the cavity around HH 250-IRS hinted by its drop in flux at $3.8 \mu\text{m}$. Evidence of companions causing large dust-free cavities that masquerade as transition disks has been reported by Ruíz-Rodríguez et al. (2016), showing that this mechanism is at work in other binary systems.

5.2. A comparison with other similar systems

The observations presented here indicate that HH 250-IRS is an addition to the handful of binary or multiple systems known to

harbor both circumstellar and circumbinary disks. The best studied of such systems is GG Tau, where both circumstellar disks and the circumbinary disk have been observed in much detail; see Dutrey et al. (2016) for a recent review. That system is probably older than HH 250-IRS, with an age of 1.5 Myr derived from the fit of pre-main sequence evolutionary tracks to the properties of GG Tau Aa and Ab. Similarly to HH 250-IRS only one of the two components, GG Tau Aa, displays submillimeter continuum emission. The inner radius of the circumbinary disk is 180 AU, comparable to the size of the HH 250-IRS circumbinary disk when taking into account the distance uncertainty of the latter.

UY Aur (Close et al. 1998; Tang et al. 2014) is another binary, or possibly triple system, with a projected separation of 120 AU between its two main components (Hartigan & Kenyon 2003) and a circumbinary disk with a radius of 520 AU. High resolution submillimeter images presented by Tang et al. (2014) show that its circumbinary disk is well detected in CO but only marginally in dust continuum. Like in GG Tau, the photospheres of the UY Aur binary are well observed, suggesting an older age than that of HH 250-IRS.

A closer and perhaps even younger analog to HH 250-IRS that has attracted recent attention is the L1551 NE system (Reipurth et al. 2002; Takakuwa et al. 2012), a system containing two Class I members with a projected separation of 70 AU, one of which drives an outflow traced by [FeII] emission, and surrounded by a circumbinary disk with a radius of 300 AU (Takakuwa et al. 2014). High resolution images of L1551 NE resolving both the circumbinary and circumstellar disks obtained with ALMA, and a comparison with numerical simulations of the gas flow in the system, are presented by Takakuwa et al. (2017).

Other systems confirmed or suspected to have circumbinary disks are listed in Monin et al. (2007) and Dutrey et al. (2016). The best studied example where a well developed jet exists is HH 30, a system where the binarity of the driving source and its properties (masses and separation) are inferred from the wiggling of the jet, rather than by direct observation due to its blocking by a disk seen nearly edge-on (Estalella et al. 2012). This makes HH 250-IRS a rare example among the selected group of binary systems surrounded by circumbinary disks, where the connection of the observed characteristics of the system with its outflow activity in the recent past may be more directly established.

6. Conclusions

In this paper we report the first dedicated study of HH 250-IRS, the driving source of HH 250, a little studied Herbig-Haro object in Aquila. Our conclusions are summarized as follows:

- Obvious evolution is observed in the region around HH 250-IRS and HH 250 by comparing $H\alpha$ images obtained in 1994 and 2016. The appearance of the most prominent feature, the arc-shaped HH 250A, is dominated by a bright spot in the 2016 image that was inconspicuous in 1994. The overall proper motion of HH 250A is estimated to be $0''.046 \text{ yr}^{-1}$ (corresponding to $\sim 50 \text{ km s}^{-1}$) pointing away from HH 250-IRS and confirming its origin in that source. Different parts of the reflection nebula in the surroundings of HH 250-IRS have faded or brightened over that time span, suggesting a clumpy envelope around the illuminating source.
- The powering source of HH 250-IRS is shown to be a binary system separated by $0''.53$. Both components have Class I spectral energy distributions, but clear differences appear between both. The near-infrared spectrum of the SE component is closer to the Class I/II transition, and the spectral energy distribution suggests a strong silicate feature in emission, while it seems to be in absorption in the NW component. A drop in the L' -band flux of the SE component is tentatively attributed to the dynamical interaction with the NW component.
- Millimeter-wave observations reveal two kinematical components of the gas traced by CO emission separated by a difference of $\approx 6 \text{ km s}^{-1}$ in radial velocity. The redshifted component is resolved, elongated along the same direction as that defined by the HH 250-IRS NW-SE pair, shows hints of being doubly-peaked, and has a temperature of $\approx 30 \text{ K}$. The blueshifted component remains unresolved by our observations and has a temperature of $\approx 14 \text{ K}$. The extent and hints of structure of the redshifted component suggest that it traces a circumbinary disk around the HH 250-IRS system, while the blueshifted component may correspond to a circumstellar disk around one of the two components of the binary system.
- Cold dust traced by continuum emission is detected as an unresolved source centered on HH 250-IRS, being spatially coincidental with the blueshifted component seen in CO emission.

The intriguing characteristics of the HH 250-IRS binary system described in this work suggest its potential for the study of accretion from a circumbinary disk onto circumstellar disks in a system in which both components have in turn a strong dynamical interaction. At the same time, the existence of a large-scale outflow revealed by the existence of the HH 250 bow shock offers the possibility of further investigating the link between binarity and collimated outflows. While the observations presented here are in many respects tantalizing due to their limited resolution and sensitivity, higher quality observations can make HH 250-IRS a cornerstone for the understanding of the interplay between the central binary, its circumstellar envelopes, the circumbinary material and the jets. Given the small number of such systems known to date and the diversity of outcomes resulting from the possible combinations of initial parameters that simulations predict, HH 250-IRS is a precious addition to the sample.

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Appendix A: An adjacent binary

While examining our L' -band images of HH 250-IRS we noticed another binary $21''$ to the northeast, which is identified as 2MASS 19212367+1054108 (hereafter 2M192123). This is a fainter object ($K = 10.56$) than HH 250-IRS, and a tighter binary with a separation of only $0''.15$. Given the proximity to HH 250-IRS in a very high-extinction region, we decided to examine the possibility that 2M192123 and HH 250-IRS might form a bound ϵ -Lyrae type quadruple system in the early stages of formation. We therefore obtained the same kind of spectral data of 2M192123 as for HH 250-IRS. Figure A.1 shows a SINFONI spectrum of the combined light from the two components of 2M192123. The spectrum clearly indicates a cool photosphere, with recognizable absorption features due to NaI ($2.206 \mu\text{m}$),

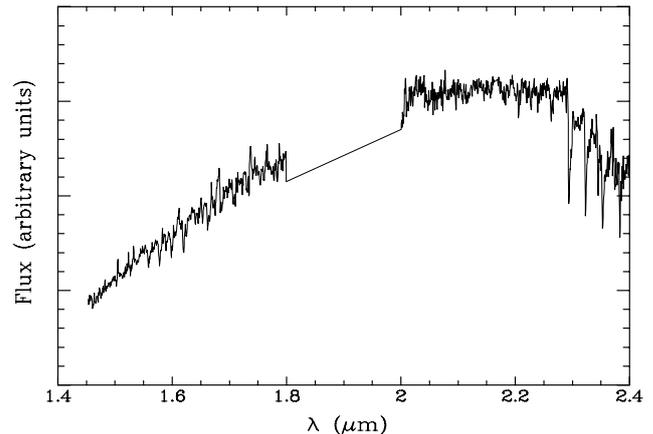


Fig. A.1. SINFONI spectrum of the close double star located $21''$ north-east of the HH 250-IRS binary, initially suspected to be physically related to the system. Despite the low photospheric temperature indicated by the appearance of the NaI ($2.206 \mu\text{m}$) and CaI ($2.262 \mu\text{m}$) lines and the deep CO bands, the spectrum is distinctly different from that of a low-mass young stellar object, being instead typical of a late-type giant or supergiant star. The apparent emission at $2.166 \mu\text{m}$ is an artifact due to the presence of Bry absorption in the spectrum of the star used for telluric features correction.

CaI ($2.262 \mu\text{m}$) and, most notably, CO. However, the spectrum rules out the possibility that 2M192123 may be a low-mass young stellar object, as the broad water steam absorption wings that should dominate the H - and K -band spectrum in that case are clearly absent. Furthermore the depth of the CO absorption, for which we measure an equivalent width of 27 \AA in the $\nu(2-0)$ band starting at $2.293 \mu\text{m}$, indicates a very low surface gravity characteristic of a supergiant (Lançon & Wood 2000), over two orders of magnitude lower than for a young stellar object. Finally, WISE observations do not show any evidence for infrared excess in the $3.5\text{--}20 \mu\text{m}$ range². This is further strengthened by the complete absence of any detection in the 225 and 351 GHz SMA map of the region. We conclude that 2M192123 is a background object, probably a binary containing at least one M supergiant member, unrelated to the HH 250-IRS binary.

² The AllWISE Data Release lists $W3$ and $W4$ magnitudes which, taken at face value, would indicate substantial infrared excess, $(W2 - W3) = 1.55$ and $(W3 - W4) = 5.62$. However, inspection of the WISE images shows that the object is undetected by WISE in both the $W3$ and $W4$ bands, and that the listed values are actually contaminated, most probably by HH 250-IRS itself.