

# Probing the 2D temperature structure of protoplanetary disks with *Herschel* observations of high- $J$ CO lines

D. Fedele<sup>1,2</sup>, E. F. van Dishoeck<sup>2,3</sup>, M. Kama<sup>3</sup>, S. Bruderer<sup>2</sup>, and M. R. Hogerheijde<sup>3</sup>

<sup>1</sup> INAF–Osservatorio Astrofisico di Arcetri, L.go E. Fermi 5, 50125 Firenze, Italy  
e-mail: fedele@mpe.mpg.de

<sup>2</sup> Max-Planck Institut für Extraterrestrische Physik, Giessenbachstrasse 1, 85748 Garching, Germany

<sup>3</sup> Leiden Observatory, Leiden University, PO Box 9513, 2300 RA Leiden, The Netherlands

Received 13 July 2015 / Accepted 5 April 2016

## ABSTRACT

The gas temperature structure of protoplanetary disks is a key ingredient for interpreting various disk observations and for quantifying the subsequent evolution of these systems. The comparison of low- and mid- $J$  CO rotational lines is a powerful tool for assessing the temperature gradient in the warm molecular layer of disks. Spectrally resolved high- $J$  ( $J_u > 14$ ) CO lines probe intermediate distances and heights from the star that are not sampled by (sub-)millimeter CO spectroscopy. This paper presents new *Herschel*/HIFI and archival PACS observations of <sup>12</sup>CO, <sup>13</sup>CO, and [C II] emission in four Herbig AeBe disks (HD 100546, HD 97048, IRS 48, HD 163296) and three T Tauri disks (AS 205, S CrA, TW Hya). In the case of the T Tauri systems AS 205 and S CrA, the CO emission has a single-peaked profile, likely due to a slow wind. For all the other systems, the *Herschel* CO spectra are consistent with pure disk emission and the spectrally resolved lines (HIFI) and the CO rotational ladder (PACS) are analyzed simultaneously assuming power-law temperature and column density profiles, using the velocity profile to locate the emission in the disk. The temperature profile varies substantially from disk to disk. In particular,  $T_{\text{gas}}$  in the disk surface layers can differ by up to an order of magnitude among the four Herbig AeBe systems; HD 100546 is the hottest and HD 163296 the coldest disk in the sample. Clear evidence of a warm disk layer where  $T_{\text{gas}} > T_{\text{dust}}$  is found in all the Herbig Ae disks. The observed CO fluxes and line profiles are compared to predictions of physical-chemical models. The primary parameters affecting the disk temperature structure are the flaring angle, the gas-to-dust mass ratio, the scale height, and the dust settling.

**Key words.** stars: pre-main sequence – stars: variables: T Tauri, Herbig Ae/Be – protoplanetary disks

## 1. Introduction

A key physical parameter of protoplanetary disks is the gas temperature,  $T_{\text{gas}}$ . Inside a disk  $T_{\text{gas}}$  controls the dynamics of the gas by setting the sound speed and, through this value, also the disk photoevaporation. At the same time,  $T_{\text{gas}}$  governs the chemical composition by regulating the reaction rates between different species. Disks are characterized by a strong temperature gradient in the radial and the vertical directions. For this reason, multiple transitions that trace different vertical layers and different orbital radii have to be observed to derive  $T_{\text{gas}}$  in disks. An ideal disk “thermometer” is the CO rotational ladder. Low- $J$  ( $J_u < 6$ ) CO rotational lines have routinely been observed from the ground in protoplanetary disks since the late 1990s (e.g., Koerner & Sargent 1995; Mannings & Sargent 1997; Dutrey et al. 1998; van Zadelhoff et al. 2001). These lines probe mostly the cold gas in the outer disk ( $r > 100$  au). Recent observations of disks with *Herschel*/PACS carried out by the DIGIT (Green et al. 2013) and GASPS (Dent et al. 2013) key programs report the detection of pure rotational high- $J$  ( $J_u > 14$ ) CO emission lines (e.g., Sturm et al. 2010; van Kempen et al. 2010; Meeus et al. 2012, 2013). These lines trace warm gas ( $E_u \geq 300$  K) located in the intermediate layers between the disk surface and the midplane at intermediate distances from the star (10–50 au) as predicted by thermochemical models of UV irradiated disks (e.g., Jonkheid et al. 2007; Gorti & Hollenbach 2008; Woitke et al. 2009; Kamp et al. 2010; Bruderer et al. 2012).

The detections of the CO high- $J$  lines allow us to estimate the gas temperature in this region of the disk for the first time. However, the PACS spectra (resolving power  $R = \lambda/\Delta\lambda \sim 10^3$ ) presented by Meeus et al. (2013) are spectrally and spatially unresolved, so that their emitting region (and the radial distribution of the gas) can only be inferred indirectly from the modeling of line fluxes (e.g., Bruderer et al. 2012). The only way to overcome the lack of spatial resolution at high THz frequencies and to determine the warm gas distribution within disks is with high-resolution spectroscopy with HIFI ( $R = 10^6$ – $10^7$ ), where Kepler’s law can be used to associate a velocity bin with a radial location in the disk.

The *Herschel* PACS and HIFI CO spectra of the Herbig Ae system HD 100546 have been presented in Fedele et al. (2013b, hereafter Paper I) in which the radial gas temperature gradient is estimated for the first time. This paper presents new *Herschel*/HIFI observations of CO  $J = 16$ – $15$  toward HD 97048, AS 205, Oph-IRS 2-48, and S CrA, and of CO and <sup>13</sup>CO  $J = 10$ – $9$  toward TW Hya, HD 100546 and HD 163296. *Herschel*/HIFI observations of [C II] (158  $\mu\text{m}$ ) in HD 97048 and HD 100546 are reported in the Appendix.

## 2. Observations and data reduction

The sample selection is based on the PACS detection of high- $J$  CO and [C II] emission (Meeus et al. 2013; Fedele et al. 2013a). The observations log is shown in Table 1. Most of the data are from program ID OT2\_DFedele\_1 (PI: D. Fedele). The HIFI

**Table 1.** *Herschel*/HIFI observations log and line properties.

Target	RA (J2000)	Dec (J2000)	Obsid 13422-	$v_{\text{LSR}}$ [km s <sup>-1</sup> ]	<i>FWHM</i> [km s <sup>-1</sup> ]	rms [K]	$d\nu$ [km s <sup>-1</sup> ]	Int. intensity [K km s <sup>-1</sup> ]	Int. flux [10 <sup>-17</sup> W m <sup>-2</sup> ]
CO $J = 16-15$ , $E_u = 751.8$ K, 1841.345 GHz, $\eta_{\text{mb}} = 0.57$ (H), 0.60 (V), <i>HPBW</i> = 11.1''									
HD 97048	11:08:03.32	-77:39:17.5	50 973	4.9	7.4	0.04	0.08	1.42 ± 0.05	2.97 ± 0.10
HD 100546	11:33:25.44	-70:11:41.2	47 519	5.3	8.2	0.06	0.08	2.89 ± 0.08	6.06 ± 0.17
AS 205	16:11:31.40	-18:38:24.5	51 072	4.8	3.7	0.04	0.08	0.76 ± 0.06	1.59 ± 0.12
Oph-IRS 48	16:27:37.19	-24:30:35.0	51 070	4.8	12.0	0.04	0.08	0.40 ± 0.04	0.84 ± 0.08
S CrA	19:01:08.60	-36:57:20.0	53 691	6.4	2.6	0.04	0.08	1.35 ± 0.07	2.83 ± 0.15
CO $J = 10-9$ , $E_u = 304.2$ K, 1151.985 GHz, $\eta_{\text{mb}} = 0.59$ (H), 0.59 (V), <i>HPBW</i> = 19.5''									
TW Hya	11:01:51.91	-34:42:17.0	10733	2.9	1.3	0.06	0.13	0.22 ± 0.04	0.34 ± 0.06
HD 100546	11:33:25.44	-70:11:41.2	35779	5.5	6.4	0.04	0.13	3.29 ± 0.05	5.21 ± 0.08
HD 163296	17:56:21.29	-21:57:21.9	51440	5.9	6.5	0.07	0.13	0.95 ± 0.23	1.51 ± 0.37
<sup>13</sup> CO $J = 10-9$ , $E_u = 290.8$ K, 1101.349 GHz, $\eta_{\text{mb}} = 0.59$ (H), 0.59 (V), <i>HPBW</i> = 19.5''									
TW Hya	11:01:51.91	-34:42:17.0	01585	2.7	1.2	0.006	0.14	0.02 ± 0.003	0.028 ± 0.004
HD 100546	11:33:25.44	-70:11:41.2	56430	5.5	7.8	0.007	0.14	0.64 ± 0.03	0.88 ± 0.06
HD 163296	17:56:21.29	-21:57:21.9	53595	5.9	9.5	0.006	0.14	0.14 ± 0.01	0.19 ± 0.01
[C II] $^2\text{P}_{3/2}^0 - ^2\text{P}_{1/2}^0$ , $E_u = 91.3$ K, 1900.537 GHz, $\eta_{\text{mb}} = 0.57$ (H), 0.60 (V), <i>HPBW</i> = 11.1''									
HD 97048	11:08:03.32	-77:39:17.5	52194	4.90	2.75	0.10	0.08	4.09 ± 0.09	9.44 ± 0.21
HD 100546	11:33:25.44	-70:11:41.2	47518	5.75	3.70	0.09	0.08	8.50 ± 0.15	19.6 ± 0.34

**Notes.** The rms is measured at the original spectral resolution; channel width is given in column “ $d\nu$ ”.

<sup>12</sup>CO and <sup>13</sup>CO  $J = 10-9$  spectra are from the WISH key program (PI: E.F. van Dishoeck) for TW Hya, and from programs OT1\_mhogerhe\_1 (HD 100546 <sup>12</sup>CO  $J = 10-9$ , PI: M. Hogerheijde), OT2\_mhogerhe\_2 (HD 100546 and HD 163296 <sup>13</sup>CO  $J = 10-9$ , PI: M. Hogerheijde), and OT1\_lpodio\_1 (HD 163296 <sup>12</sup>CO  $J = 10-9$ , PI: L. Podio).

The CO observations were executed in dual beam switch fast chopping mode with the Wide-Band Spectrometer (WBS) and the High Resolution Spectrometer (HRS) simultaneously. The spectral resolution is set to 1.1 MHz for WBS and 0.25 MHz for HRS for both polarizations. The [C II] observations were carried out in “load chop” where an internal calibration source is used in combination with an off-source calibration observation. This allows us to remove the spatially extended [C II] emission. The beam size (*HPBW*) is 11'1 at the observed frequency (Roelfsema et al. 2012). In the following, we will only refer to the WBS spectra.

The spectra are extracted from the level 2 data, which were processed with standard pipeline SPG v9.1.0. Standing waves are present in the WBS spectra. These waves were removed by fitting a set of sine functions after masking the narrow spectral features (CO or [C II]). This operation was performed with the “fitHifiFringe” script provided with Hipe. The HIFI level 2 fluxes are shown on the antenna temperature scale ( $T_A^*$ ) and were converted to main beam temperature,  $T_{\text{mb}} = T_A^* \times \eta_l / \eta_{\text{mb}}$ , with  $\eta_l$  the forward efficiency and  $\eta_{\text{mb}}$  the beam efficiency (Table 1). No major differences are present between the H and V polarizations and the two spectra were averaged together after applying the efficiency corrections and removing the continuum. High- $J$  CO lines are not contaminated by the cold cloud contribution that plagues single dish low- $J$  CO lines.

The reduction of the archival PACS data analyzed here is described in Meeus et al. (2013) and Fedele et al. (2013a).

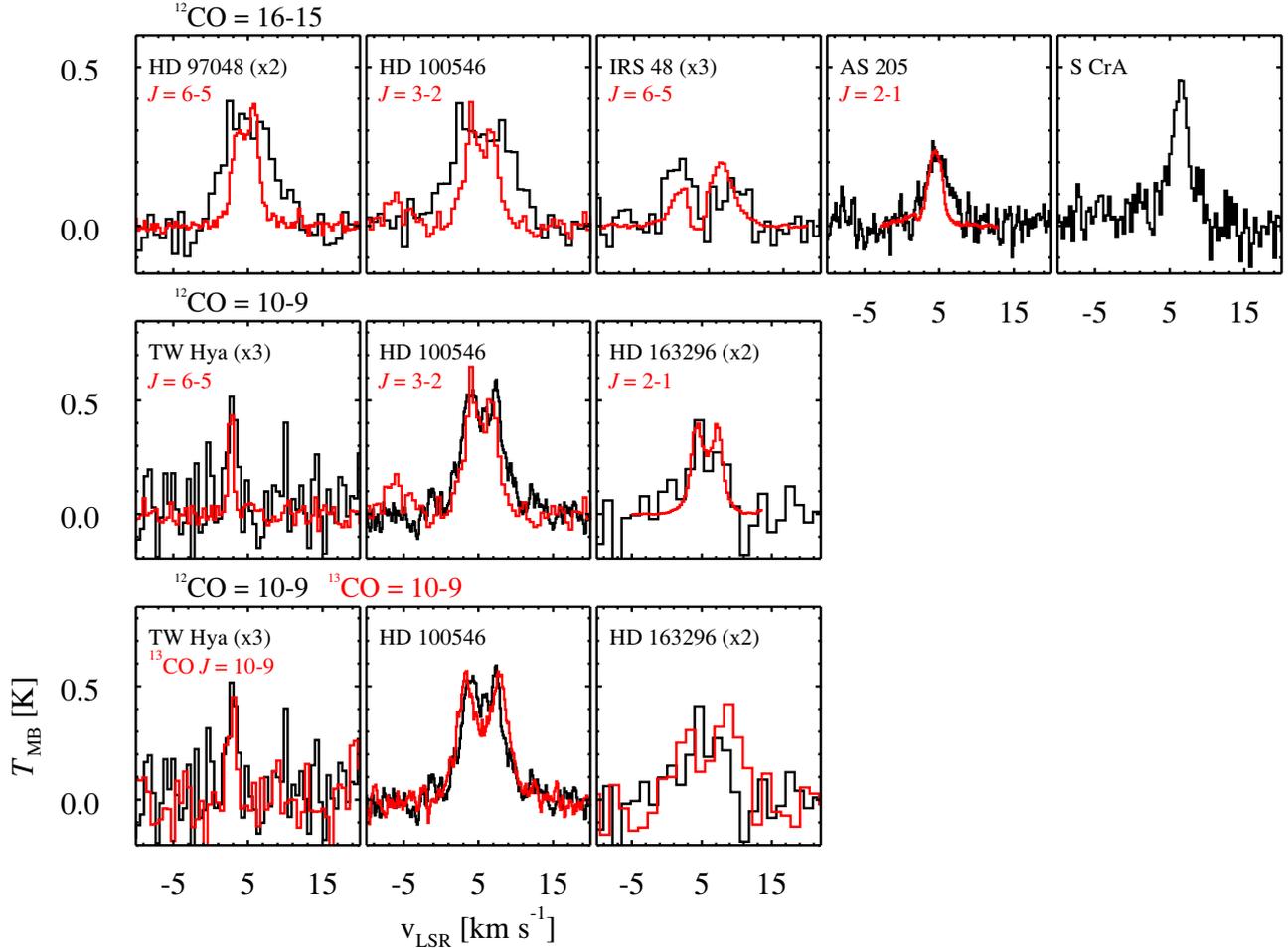
### 3. Results

The HIFI/WBS CO spectra are presented in Fig. 1 and the line parameters are given in Table 1. The [C II] spectra and analysis is presented in the Appendix. The integrated line flux (W m<sup>-2</sup>) is computed from the integrated intensity

$$\int T_{\text{mb}} d\nu = 2k \left(\frac{\nu}{c}\right)^3 \pi \left(\frac{\text{HPBW}}{2\sqrt{\ln(2)}}\right)^2 \int T_{\text{mb}} d\nu \quad [\text{K km s}^{-1}] \quad (1)$$

with  $k$  (Boltzmann constant, W s K<sup>-1</sup>),  $\nu$  the frequency (Hz),  $c$  the light speed (m s<sup>-1</sup>), and *HPBW* the beam (radians). Because of the updated values of the beam efficiency (Table 1), we note that the line intensities presented here for HD 100546 are slightly different (~3%) from the values given in Paper I.

The top panel of Fig. 1 shows the WBS spectra of the CO  $J = 16-15$  lines toward HD 97048, HD 100546, IRS 48, AS 205, and S CrA. The line is clearly detected above 5 $\sigma$  (Table 1) in all sources. The velocity profile and width are different among the five sources. The emission is broad ( $\Delta\nu > 5$  km s<sup>-1</sup>) toward the three HAeBe stars (HD 97048, HD 100546, and IRS 48): a double-peak profile is clearly visible in HD 100546 and IRS 48, while the CO velocity profile is top-flat in HD 97048. In all three cases, the WBS spectra of the CO  $J = 16-15$  transition are consistent with a Keplerian velocity field of the gas in the disk. In the case of the two T Tauri systems, AS 205 and S CrA, the CO  $J = 16-15$  emission is narrow ( $\Delta\nu < 5$  km s<sup>-1</sup>) and single-peaked. There is no evidence of Keplerian rotation. For comparison, the profiles of the



**Fig. 1.** HIFI/WBS spectra of CO  $J = 16-15$  (top),  $J = 10-9$  (middle and bottom), and  $^{13}\text{CO } J = 10-9$  (bottom). The spectra of low- $J$  lines are overlaid in red (scaled for comparison) in the top and middle rows. For clarity, some spectra were rebinned to lower resolution: for CO  $J = 16-15$   $\text{d}v = 0.8 \text{ km s}^{-1}$  in HD 97048, HD 100546, and IRS 48, and  $\text{d}v = 0.32 \text{ km s}^{-1}$  in S CrA; for CO  $J = 10-9$   $\text{d}v = 0.52 \text{ km s}^{-1}$  in TW Hya and  $\text{d}v = 1.3 \text{ km s}^{-1}$  in HD 163296; for  $^{13}\text{CO } J = 10-9$   $\text{d}v = 0.56 \text{ km s}^{-1}$  in TW Hya and  $\text{d}v = 1.4 \text{ km s}^{-1}$  in HD 163296. The remaining spectra are shown at their native resolutions.

low- $J$  CO lines are also shown when available. Spectra (either APEX or ALMA) are from Panić et al. (2010), Bruderer et al. (2014), Salyk et al. (2014) and Kama et al. (2016). In all cases, the low- and high- $J$  CO lines are centered at the same  $v_{\text{LSR}}$  with the CO  $J = 16-15$  line broader than the low- $J$  line. We note that part of the asymmetric profile of the  $J = 6-5$  line toward IRS 48 is due to extinction from the foreground cloud (Bruderer et al. 2014). In the case of AS 205, both CO lines are centered at  $v_{\text{LSR}} = 4.8 \text{ km s}^{-1}$  compared to the cloud velocity of  $v_{\text{LSR}} = 3 \text{ km s}^{-1}$ . Thus, we conclude that the high- $J$  CO emission in AS 205 arises from a slow wind/outflow, as do the low- $J$  CO (Salyk et al. 2014) and the ro-vibrational (Pontopidan et al. 2011) emission.

The narrow, single-peak, profile of the CO  $J = 16-15$  emission in S CrA also suggests a contribution from a slow wind associated with the system. The CO ro-vibrational lines toward both stellar components in this binary are broad and single-peaked, similar to AS 205 (Bast et al. 2011; Brown et al. 2013).

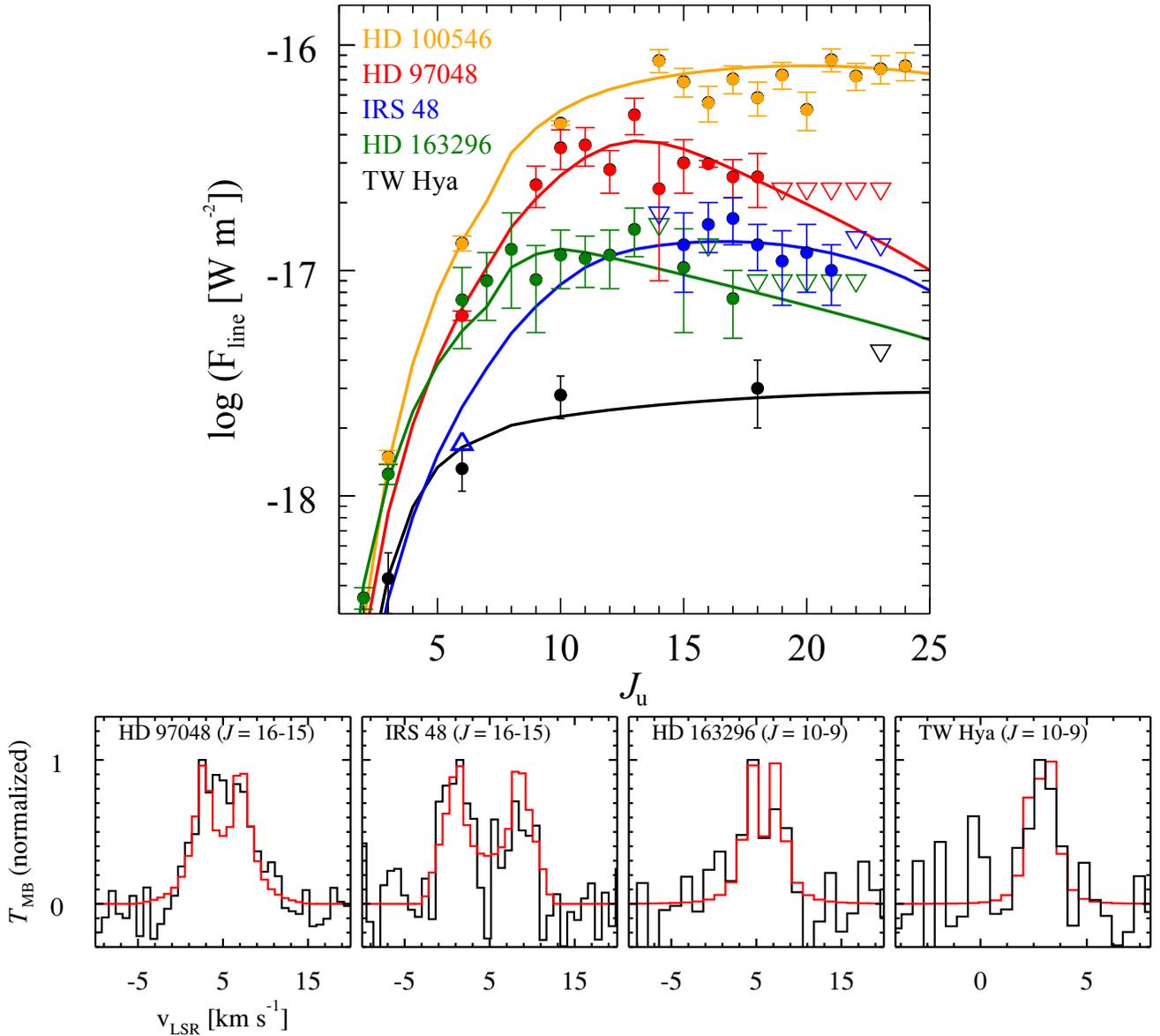
Figure 1 also shows the HIFI/WBS spectra of CO  $J = 10-9$  compared to the low- $J$  transitions (middle) and to the  $^{13}\text{CO}$  (bottom) for TW Hya, HD 100546, and HD 163296. The  $^{12}\text{CO}$  and  $^{13}\text{CO}$  lines are both clearly detected. The lines are broad and double-peaked toward HD 100546 and HD 163296 and narrow and single-peaked toward TW Hya. In all three cases, the lines

are centered on the system velocity and the narrow line profile of TW Hya is consistent with the disk being almost face-on to the plane of the sky. As noted in Paper I, the line width is narrower for lower  $J$  transitions in HD 100546. Instead, in the case of HD 163296, the width of the  $J = 2-1$  (ALMA science verification data) and  $J = 10-9$  are similar. We note that the  $^{12}\text{CO}$  and  $^{13}\text{CO } J = 10-9$  profiles appear asymmetric toward HD 163296; however, the flux difference between the two peaks is within the noise level of the spectrum.

Interestingly, the  $^{13}\text{CO}$  line toward HD 100546 and HD 163296 is slightly broader than the  $^{12}\text{CO}$  line (Fig. 1, bottom row). This implies that the line emitting region of  $^{13}\text{CO}$  extends to higher velocity regions, i.e., closer to the star. Another prominent difference is the central peak detected in the  $^{12}\text{CO } J = 10-9$  line in the HD 100546 spectrum, which is not visible in the  $^{13}\text{CO}$  spectrum. The differences in the velocity profiles are likely due to optical depth effects: the  $^{12}\text{CO}$  line becomes optically thick at lower column densities and higher up in the atmosphere than the  $^{13}\text{CO}$  line. This is discussed further in Sect. 5.

#### 4. Analysis

The CO rotational ladder is a powerful tool that can be used to assess the temperature structure of protoplanetary disks



**Fig. 2.** *Top:* CO rotational ladder. Detections are shown as filled circles and upper limits as open triangles. The solid lines represent the best-fit power-law model. *Bottom:* best-fit model profile of the CO  $J = 16-15$  (HD 97048 and IRS 48) and CO  $J = 10-9$  (HD 163296 and TW Hya). Spectral resolution as in Fig. 1.

(Bruderer et al. 2012; Fedele et al. 2013b; van der Wiel et al. 2014). The low- $J$  lines are optically thick and the lines becomes optically thin only at  $J_u \gtrsim 14$  (Bruderer et al. 2012). The advantage of the high- $J$  CO rotational lines is that they trace intermediate distances from the star (a few tens of au) and heights ( $z/r \sim 0.1-0.4$ ) above the disk midplane. The high- $J$  CO transitions are complimentary to the low- $J$  CO transitions observed at millimeter observations that trace the colder outer disk. Thus, the flux and velocity profile of the high- $J$  CO lines are crucial in order to measure the gas temperature in the disk atmosphere. The rotational ladder (Table A.1) is compiled combining low- $J$  lines ( $J_u < 6$ ) from ground-based observations, mid- $J$  ( $J_u = 7-12$ ) from SPIRE (van der Wiel et al. 2014), and high- $J$  ( $J_u > 14$ ) lines from *Herschel*/PACS (Meeus et al. 2013). In some cases, the CO line fluxes measured with SPIRE are contaminated by extended emission (cloud) as noted by van der Wiel et al. (2014). The affected lines are excluded from the fit.

The CO rotational ladders for the five disks are shown in Fig. 2. The shape of the ladder varies from object to object. We note in particular the difference between HD 100546 and HD 163296: in the first case the CO line flux increases with  $J$  and then remains almost constant, while in the case of HD 163296 the rotational ladder turns over at  $J_u \sim 10-15$ .

The analysis of HD 100546 is presented in Paper I. The sources AS 205 and S CrA are excluded from this analysis because the CO line fluxes and velocity profiles are dominated by an outflow/jet.

#### 4.1. Power-law model

Following the method described in Paper I the CO rotational ladder and the line velocity profiles ( $J = 16-15$  for HD 97048 and IRS 48,  $J = 10-9$  for TW Hya and HD 163296) are fitted simultaneously using a power-law profile for the kinetic gas

**Table 2.** Best-fit power-law model parameters.

	TW Hya	HD 97048	HD 100546 <sup>†</sup>	HD 163296	IRS 48
$M_{\star} (M_{\odot})$	0.6	2.5	2.5	2.4	2.0
$d$ (pc)	51	180	97	120	120
$i$ ( $^{\circ}$ )	7	43	42	44	50
$r_i$ (au)	0.1	11	13	0.1	20
$r_{\text{out}}$ (au)	200	400	400	600	200
$T_i$ (K)	3000–5000	500–600	750–1450	1000–1100	250–350
$q$	0.7–0.8	0.5–0.6	0.75–0.95	0.5–0.6	0.85–0.95
$N_i$ ( $\text{cm}^{-2}$ )	$1\text{--}5 \times 10^{18}$	$10^{18}\text{--}10^{19}$	$2\text{--}8 \times 10^{-17}$	$10^{19}\text{--}10^{20}$	$10^{18}\text{--}10^{19}$
$p$	0.85–0.95	0.45–0.55	0.6–1.2	0.3–0.35	0.55–0.65

**Notes.** <sup>(†)</sup> From Paper I. Best-fit parameters are found by  $\chi^2$  minimization. Ranges represent the  $1\sigma$  interval.

temperature (under the assumption that the CO excitation temperature corresponds to the kinetic temperature, which is valid given the high densities of the emitting regions) and column density:

$$T(r) = T_i \left( \frac{r}{r_i} \right)^{-q}, \quad (2)$$

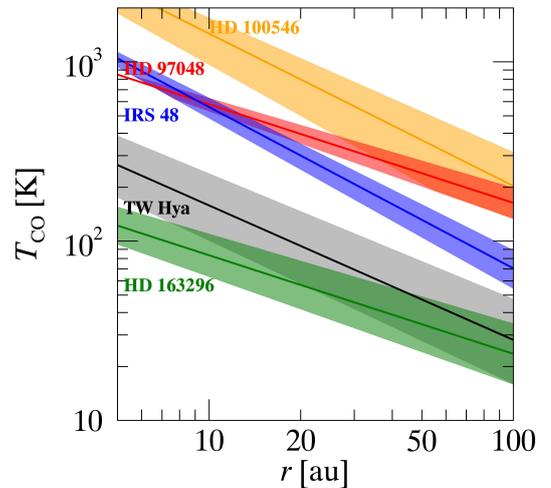
$$N(r) = N_i \left( \frac{r}{r_i} \right)^{-p}. \quad (3)$$

Here  $T_i$  and  $N_i$  are the values at the inner radius  $r_i$  of the disk (fixed, Table 2). The stellar and disk parameters are taken from the literature. In particular  $r_i = 0.1$  au for TW Hya (Pontoppidan et al. 2008), 13 au for HD 100546 (van der Plas et al. 2009; Brittain et al. 2009; Fedele et al. 2015), 11 au for HD 97048 (van der Plas et al. 2009), and 20 au for IRS 48 (Bruderer et al. 2014). The power law is truncated at the outer disk radius,  $r_{\text{out}}$  (fixed, Table 2; the choice of the outer radius does not matter as long as  $r_{\text{out}} > 100$  au). The free parameters of the model are  $T_i$ ,  $q$ ,  $N_i$ , and  $p$ . A grid of models is created for each disk varying the four parameters in the ranges  $T_i = 300\text{--}1500$  K,  $N_i = 10^{17}\text{--}10^{22}$   $\text{cm}^{-2}$ ,  $q$ , and  $p = 0.5\text{--}1.5$ . The spectra are spatially convolved with the telescope beam, represented here by a Gaussian profile. Further details about the model and the fitting procedure are given in Paper I. The best-fit parameters are found by minimizing the  $\chi^2$  between the observations and the model. The final  $\chi^2$  is given by the sum of the individual  $\chi^2$  of the CO rotational ladder and of the line profiles (one for each line), including their width and peak separation. The best-fit model parameters are listed in Table 2 and the best-fit models are overlaid on the data in Fig. 2.

The derived temperature profiles (labeled  $T_{\text{CO}}$ ) are plotted in Fig. 3. Evidence of warm disk temperature ( $T > 100$  K) is found for HD 97048, HD 100546, and IRS 48. The coldest disks are TW Hya and HD 163296, which both have  $T < 100$  K outward of 20 au.

#### 4.2. Caveats of the power-law model

The assumption of a flat disk geometry is justified by the results of physical-chemical models, which predict that most of the CO rotational lines arise from a similar vertical layer in the disk. With this assumption, however, the power-law model fails to reproduce the core of the line profile, especially in the case of HD 97048 (Fig. 2) where the line wings are accurately reproduced by the model, while the central low-velocity part of the line is not. This can be due to an optical depth effect, if the core of the line is optically thick.

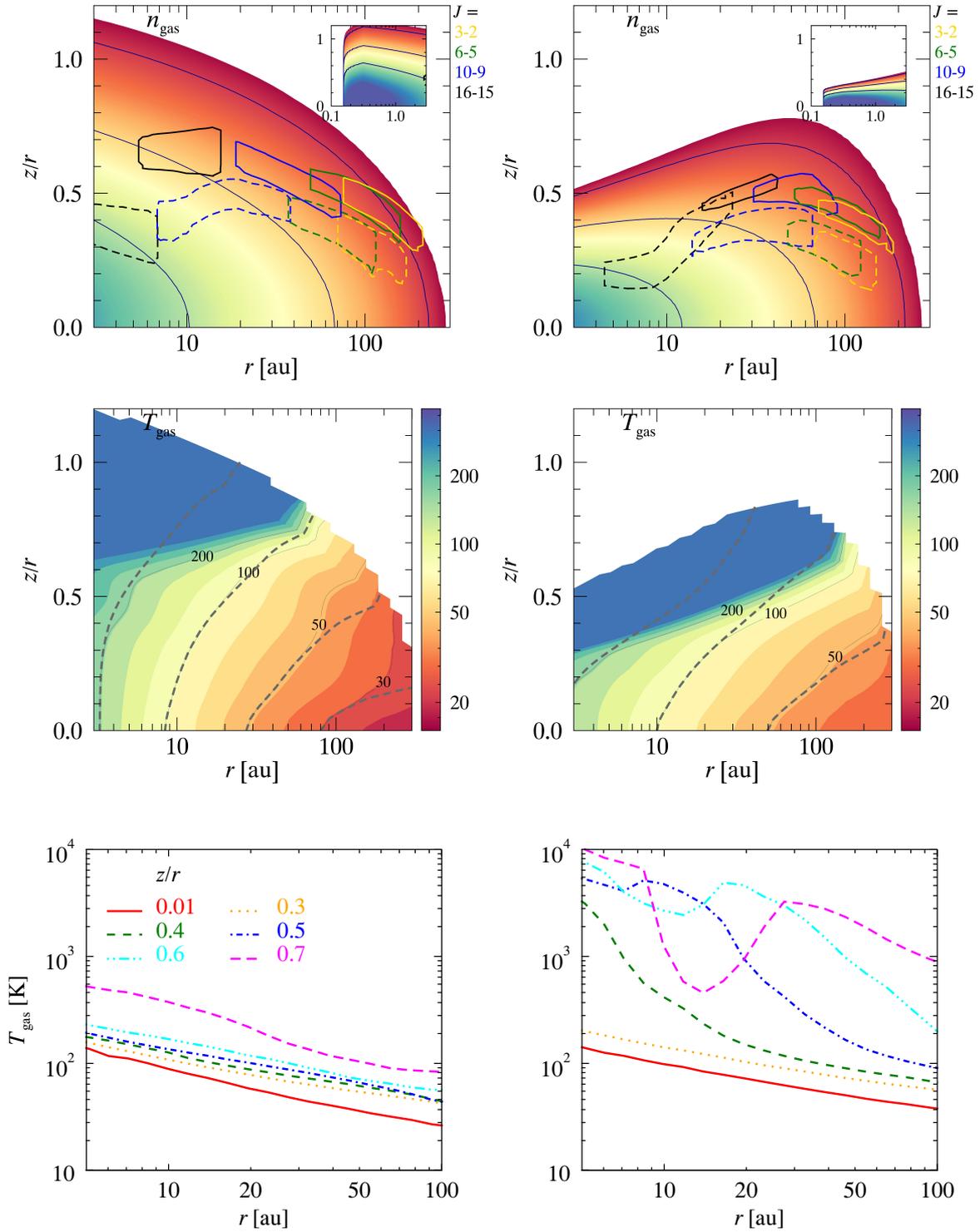


**Fig. 3.** CO temperature profile in the inner 100 au based on the best-fit power-law models (Table 2).

The second major caveat of the power-law model is that the column density profile is not constrained: most of the lines are optically thick and the CO ladder is mostly sensitive to the temperature profile of the  $\tau_{\text{line}} = 1$  layer. Thus, even if  $N_{\text{in}}$  and  $p$  are free parameters of the model, these values are to be taken with caution.

## 5. Comparison to disk models

The goal of this section is to explain the large variation in the CO rotational ladder and velocity profiles among Herbig Ae systems (Figs. 1 and 2). This study is based on the physical-chemical model, DALI (Dust and Lines; Bruderer et al. 2012; Bruderer 2013). DALI takes as input a density structure of the disk (taken to be a power-law with slope  $\gamma$  and critical radius  $R_c$  with an exponential tail) and a stellar radiation field, then solves the continuum radiative transfer and determines the dust temperature and ultraviolet radiation field at each position in the disk. The thermal balance of the gas and chemistry are subsequently solved iteratively until convergence. The output includes continuum and line emission maps, line intensities, and spectra produced via ray tracing. Dust settling is included adopting two grain size populations, small ( $a_{\text{min}} = 1 \mu\text{m}$ ) and large ( $1\text{--}1000 \mu\text{m}$ ) following D’Alessio et al. (2006), and dust cross sections from Andrews et al. (2011). The degree of settling of the large grains is controlled in DALI by the parameters  $\chi$  and  $f_{\text{large}}$ : the first



**Fig. 4.** DALI disk structure of the two representative models,  $\psi = 0.05$  (left panels) and  $\psi = 0.25$  (right panels). *Top:* gas density structure: the inset shows the inner disk structure, the dark blue curves indicate the  $n_{\text{gas}} = 10^6, 10^8,$  and  $10^{10} \text{ cm}^{-3}$  contours. The line contribution functions of a mix of low- and high- $J$  transitions of  $^{12}\text{CO}$  (solid lines) and  $^{13}\text{CO}$  (dashed) are overlaid on the  $n_{\text{gas}}$  panels. Each contour shows the layer where 50% of the line flux emerges. *Middle:* gas temperature structure, the isothermal contours are overlaid for  $T_{\text{gas}}$  and  $T_{\text{dust}}$  (dashed lines) = 20, 50, 100, 200, and 500 K. *Bottom:* gas temperature radial profile at different disk heights relevant to the disk midplane ( $z/r = 0.01$ ) and to the CO emitting layers.

defines the maximum scale height of the large grains with respect to the small ones (similar to “ $Z_{\text{big}}$ ” in D’Alessio et al. 2006), while  $f_{\text{large}}$  determines the small-to-large grain mass ratio.

Figure 4 shows the gas density (top row) structure for the two representative models (Table 3, the spectral energy distribution of the two models is shown in the Appendix). The inset

**Table 3.** DALI model parameters for Herbig Ae disks.

Parameter	Value	Unit	Description
$M_{\star}$	2	$[M_{\odot}]$	Stellar mass
$T_{\text{eff}}$	8000, <b>9000</b> , 10 500	[K]	Stellar (blackbody) temperature
$L_{\text{bol}}$	20	$[L_{\odot}]$	Bolometric luminosity
$\Delta_{\text{gas/dust}}$	1, <b>10</b> , 100		Gas-to-dust mass ratio
PAHs	<b>1</b> , 10, 50	[% (w.r.t. ISM)]	PAH abundance
$M_{\text{disk}}$	$10^{-4}$ , <b><math>10^{-3}</math></b> , $10^{-2}$	$[M_{\odot}]$	Disk mass
$\gamma$	0.8, <b>1</b> , 1.2		Surface density power-law exponent
$R_{\text{sub}}$	0.31	[au]	Sublimation radius
$R_{\text{c}}$	50, <b>75</b> , 100	[au]	Critical radius
$h_{\text{c}}$	0.1, <b>0.2</b> , 0.3	[radians]	Scale height
$R_{\text{out}}$	200, <b>400</b> , 600	[au]	Disk outer radius
$\psi$	<b>0.05</b> , 0.15, <b>0.25</b>		Flaring angle
$\chi$	<b>0.2</b> , 0.5, 1.0		Degree of settling
$f_{\text{large}}$	0.50, <b>0.85</b> , 0.999		Large-to-small grains mass ratio
$a_{\text{min}}$	<b>0.001</b> , 0.01, 0.05	$[\mu\text{m}]$	Minimum grain size
$L_{\text{X}}$	$10^{29}$	$[\text{erg s}^{-1}]$	X-rays luminosity
$d$	100	[au]	Distance
$i$	45	$[\circ]$	Disk inclination

**Notes.** Values in bold indicate the representative disk models. For model description and parameters definition see [Bruderer \(2013\)](#).

shows the inner disk structure. The line contribution functions of a set of low- $J$  and high- $J$  transitions of  $^{12}\text{CO}$  and  $^{13}\text{CO}$  are overlaid on the  $n_{\text{gas}}$  contours showing the layer where 50% of the line flux emerges. The CO emitting layer varies slightly with  $J$  and the four transitions shown here emerge from a layer between  $z/r \sim 0.4$ – $0.6$  and  $z/r \sim 0.3$ – $0.5$  for the flat and flared disk, respectively. In all cases, the  $^{13}\text{CO}$  lines emerge closer to the disk midplane and at smaller stellar distances than the  $^{12}\text{CO}$  lines do; because of the vertical and radial density gradients, the CO emission becomes optically thick higher up in the disk atmosphere and at larger distance from the star compared to the  $^{13}\text{CO}$  line. This is in excellent agreement with the broader velocity profile observed for  $^{13}\text{CO}$  (Table 1, Fig. 1, bottom row).

The gas temperature structure for the representative models is shown in Fig. 4 (middle panel). At any given position in the disk, the gas and dust temperatures increase with flaring angle. We note in particular that the dust temperature in both disks is  $T_{\text{dust}} > 20$  K everywhere in the disk. This prevents CO from freeze-out on dust grains as condensation occurs only at  $T \lesssim 20$  K in the disk interior. Without CO freeze-out, formation of complex species via surface chemistry is inhibited in such a warm disk. In both cases,  $T_{\text{gas}}$  is larger than  $T_{\text{dust}}$  in the upper layers of the disk.

The  $T_{\text{gas}}$  radial profile at different heights ( $z/r$ ) for the two representative models is shown in the bottom panel of Fig. 4. In the flat disk case, the temperature increases slightly with height with an almost identical radial dependence from the disk midplane ( $z/r = 0.01$ ) up to the disk upper layers. This is no longer true for flared disks, which show a strong dependence of  $T_{\text{gas}}(r)$  with disk height.

### 5.1. Disk models grid

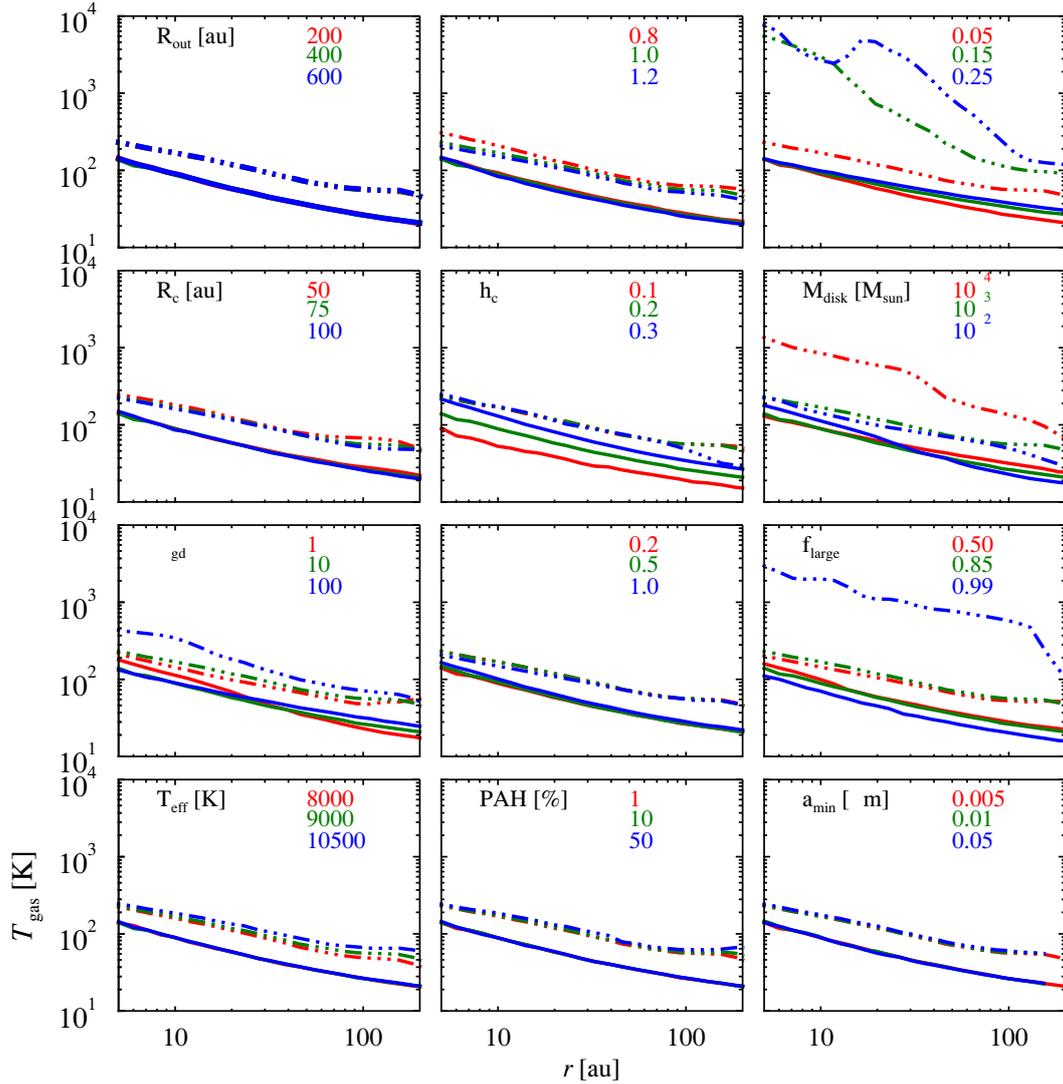
This section provides a qualitative comparison to physical-chemical models of Herbig AeBe disks. For this, a grid of DALI disk models is built to study the impact of different stellar and disk parameters on the disk thermal structure. The stellar mass and the bolometric luminosity are taken to be fixed while the stellar effective temperature ( $T_{\text{eff}}$ , assuming blackbody

emission), the disk mass ( $M_{\text{disk}}$ ), critical radius ( $R_{\text{c}}$ ), disk outer radius ( $R_{\text{out}}$ ), scale height ( $h_{\text{c}}$ ), power-law exponent ( $\gamma$ ), flaring angle ( $\psi$ ), dust settling ( $\chi$ ,  $f_{\text{large}}$ ), minimum grain size ( $a_{\text{min}}$ ), gas-to-dust mass ratio ( $\Delta_{\text{gd}}$ ), and the PAH abundances are varied (Table 3). Several of these parameters were also investigated in [Bruderer et al. \(2012\)](#), but only for a flared disk. The model grid is built around the flat ( $\psi = 0.05$ ) disk representative model varying the aforementioned parameters by the values given in Table 3. In particular, the values of  $T_{\text{eff}}$  are representative of the typical stellar temperature of Herbig Ae systems (e.g., [van den Ancker et al. 1998](#)). The far-ultraviolet (FUV; 6–13.6 eV) luminosity, relevant to the heating of the disk and the photodissociation of CO, is regulated by  $T_{\text{eff}}$  and it ranges between  $L_{\text{FUV}} = 0.45$ – $1.7 L_{\odot}$  for  $T_{\text{eff}} = 8000$ – $10\,500$  K.

To investigate the impact of dust settling on the disk thermal structure,  $\chi$  and  $f_{\text{large}}$  were varied between 0.2–1.0 and 0.5–0.999, respectively (Table 3). Lowering the value of  $\chi$  has the effect of letting the UV photons penetrate further inside the disk. As a consequence the  $\text{C}^+/\text{C}/\text{CO}$  transition layers shift deeper into the disk. For  $f_{\text{large}} = 0.5$  the dust mass is distributed equally between the small and the large dust grains, for  $f_{\text{large}} = 0.999$ , the bulk of the dust mass is in the large dust grains.

The abundance of PAHs and the gas-to-dust mass ratio are poorly constrained in disks. Model fits to observational data suggest PAH abundances that are typically 10–100 times lower than those in the interstellar medium ([Geers et al. 2006](#)). For Herbig AeBe disks, PAHs of about 100 carbon atoms survive the strong UV radiation ([Visser et al. 2007](#)). For this analysis three abundance values are considered (1%, 10%, and 50%) with respect to the abundance of the interstellar medium ([Draine & Li 2007](#)). The gas-to-dust mass ratio ( $\Delta_{\text{gd}}$ ) in disks is likely to be lower than that of the interstellar medium (e.g., [Chapillon et al. 2010](#)). In this paper  $\Delta_{\text{gd}}$  varies between 1% and 100% of the ISM value ([Bohlin et al. 1978](#)).

The effects of elemental carbon abundance and isotope selective photodissociation (relevant to  $^{13}\text{CO}$ ) are not treated here. These effects are investigated in [Bruderer et al. \(2012\)](#) and [Miotello et al. \(2014\)](#). As noted by [Miotello et al. \(2014\)](#), the  $^{13}\text{CO}$  line intensities are less affected by the isotope selective



**Fig. 5.**  $T_{\text{gas}}$  radial profile for different disk/stellar parameters. The temperature radial profile is shown for the layers  $z/r = 0.01$  (solid lines) and  $0.6$  (dot-dashed lines). In the case of  $h_c = 1$ , the CO emitting layer is  $z/r \sim 0.2\text{--}0.3$  and the dot-dashed line shows the temperature gradient at  $z/r = 0.3$ .

photodissociation than are  $\text{C}^{18}\text{O}$  and  $\text{C}^{17}\text{O}$ . It is thus reasonable to neglect this effect in this analysis. The effect of a lower carbon abundance is similar to that of a smaller gas-to-dust ratio (Bruderer et al. 2012).

## 5.2. Temperature radial gradient

Figure 5 shows the  $T_{\text{gas}}$  radial profile for all disk models at two different layers representative of the disk midplane and of the (uppermost) CO emitting layer.

The parameters affecting most the midplane temperature are the flaring angle, scale height, the dust settling, and to a lesser extent the disk mass and the gas-to-dust mass ratio. At the CO emitting layer the temperature gradient is controlled mostly by the disk flaring and less by the disk mass, the gas-to-dust mass ratio, the dust settling ( $f_{\text{large}}$ ), and the stellar temperature. We note that the  $f_{\text{large}} = 0.999$  case is a very extreme scenario in which the disk outer layers are almost completely devoid of dust as the bulk of the mass is made up of large dust grains that have settled in the disk midplane ( $\chi = 0.2$ ).

## 5.3. CO line fluxes

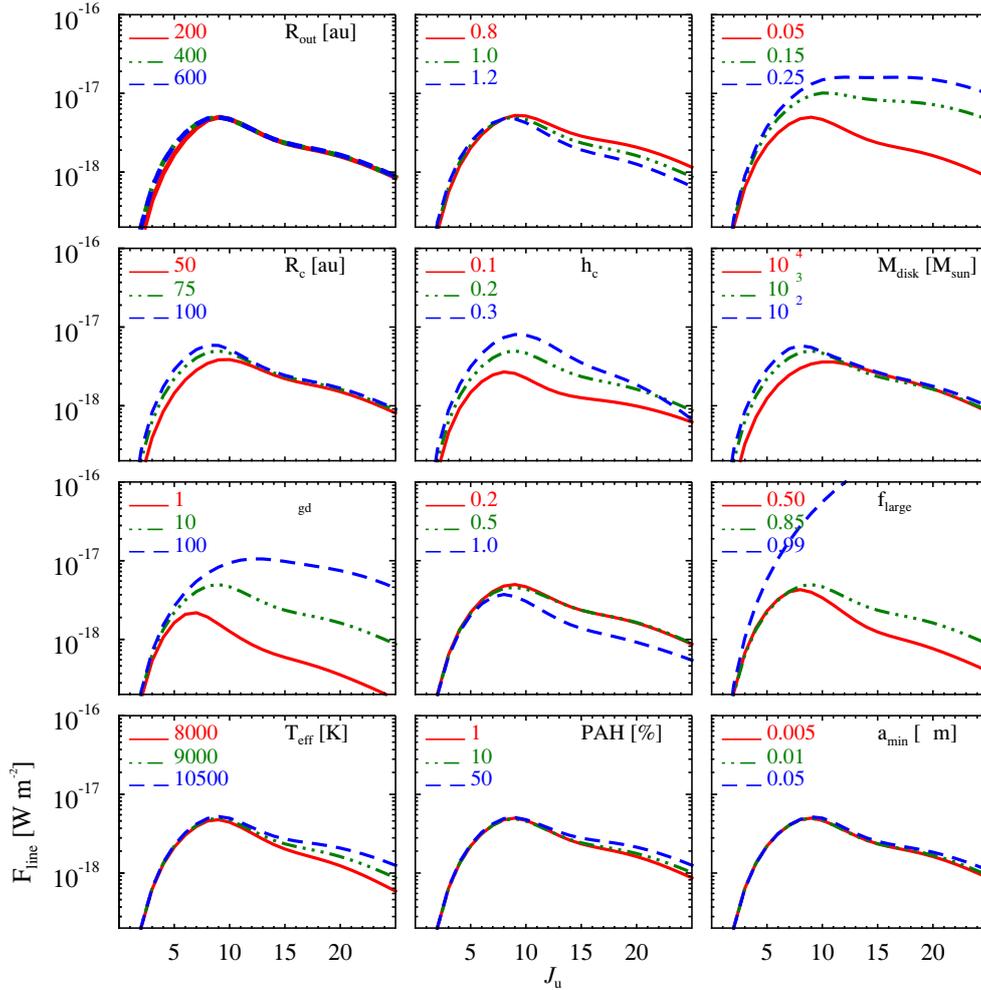
The synthetic CO ladders are shown in Fig. 6. In all cases the distance is fixed to 100 pc and the disk inclination to  $45^\circ$ .

There are two main aspects of the CO ladder that the models need to reproduce: the absolute line fluxes and especially the shape of the CO ladder, whether or not it bends over at higher  $J$ . No changes are observed for different disk sizes. The parameter that primarily affects the CO line ladder are the following:

**Flaring angle:** the CO line fluxes vary by 1–2 orders of magnitude between the flat (fainter) and the flared disks for  $J \gtrsim 10$ , and the line flux difference increases with  $J$ . This is due to the different gas temperature structure between flat and flared disks (Fig. 5);

**Critical radius:** the low- $J$  ( $J < 10$ ) lines become brighter with  $R_c$  while the flux of the high- $J$  does not vary. Since the temperature structure does not vary with  $R_c$  (Fig. 5), this is likely due to the change in vertical depth in the outer disk, which is due to the change in  $R_c$ ;

**Scale height:** all CO lines become brighter if the scale height increases. Since the temperature radial gradient does not show significant changes with  $h_c$  at the CO emitting layer,



**Fig. 6.** Synthetic CO ladder for different disk parameters (see Table 3). Fluxes measured assuming a distance of 100 pc and a disk inclination of  $45^\circ$ .

the differences in CO line fluxes are due to the shift in the  $C^+/C/CO$  transition layers deeper into the disk with decreasing scale height;

**Disk mass:** this parameter only affects the low- $J$  ( $J < 10$ ) lines, which become slightly brighter for increasing disk mass. The effect is only significant for low disk masses ( $\lesssim 10^{-4} M_\odot$ ), and is due to the change in gas column density (hence line optical depth) in the outer disk;

**Gas-to-dust mass ratio:** reducing  $\Delta_{gd}$  from 100 to 1 has the effect of lowering all line fluxes by up to one order of magnitude. We note that keeping the gas mass constant and lowering the gas-to-dust mass ratio implies a higher dust mass. Thus, the CO lines are fainter for low values of  $\Delta_{gd}$  because of the increased opacity and lower temperature (Fig. 5) induced by the higher dust mass compared to the  $\Delta_{gd} = 100$  case;

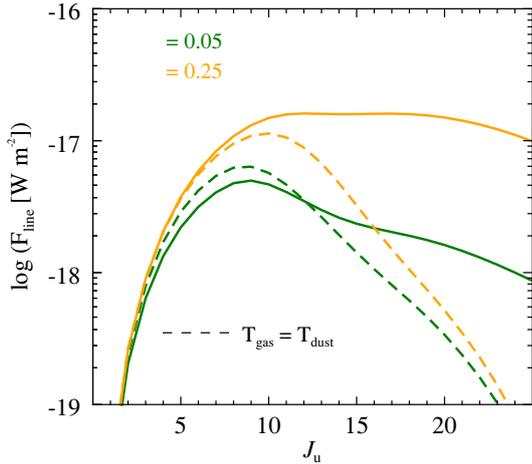
**Dust settling and large-to-small grain mass ratio:** the high- $J$  lines are brighter for a settled disk and for a higher mass ratio. This is due to the increase in gas temperature in the disk upper layers. In the extreme case of  $f_{large} = 0.999$  all the CO lines are several order of magnitude brighter.

All the other parameters analyzed here affect mostly the flux of the high- $J$  ( $J \gtrsim 10$ ) lines inducing a bending over the CO ladder for  $J \gtrsim 10$ .

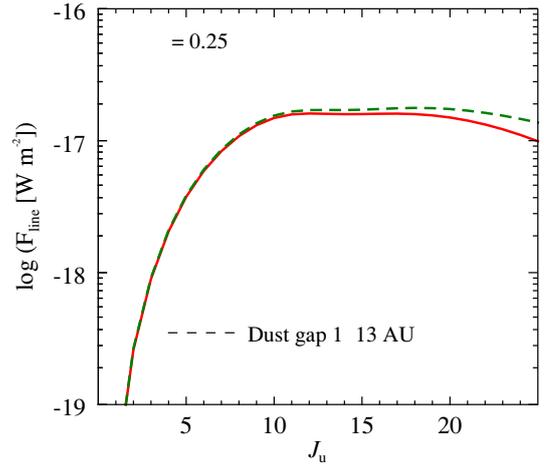
For comparison, Fig. 7 shows the CO ladder in the case of a “cold” disk atmosphere (i.e.,  $T_{gas} = T_{dust}$ , obtained by switching off the thermal balance in DALI). In this case all the CO line fluxes are substantially lower and the flux drops quickly with  $J$  for  $J_u \gtrsim 10$ . The shape of the observed CO ladder and the absolute flux level are inconsistent with the cold disk case, thus suggesting that  $T_{gas}$  exceeds  $T_{dust}$  in the upper layers of all the Herbig Ae disks studied here.

The  $^{13}CO$  ladder is shown in Fig. 8 for two representative disk models. The fluxes of the  $^{13}CO$  lines in the PACS range are all below the detection threshold of the DIGIT and GASPS surveys. The estimate of the  $^{13}CO$   $J = 10-9$  flux is in good agreement (once corrected for distance) with the value measured in this paper for HD 100546 and HD 163296 (Table 1). The fluxes of the mid- $J$   $^{13}CO$  lines reported by van der Wiel et al. (2014) toward HD 100546 (based on *Herschel*/SPIRE) are much brighter than the values estimated here. As noted by van der Wiel et al. (2014) the SPIRE measurements may be contaminated by an extended non-disk emission.

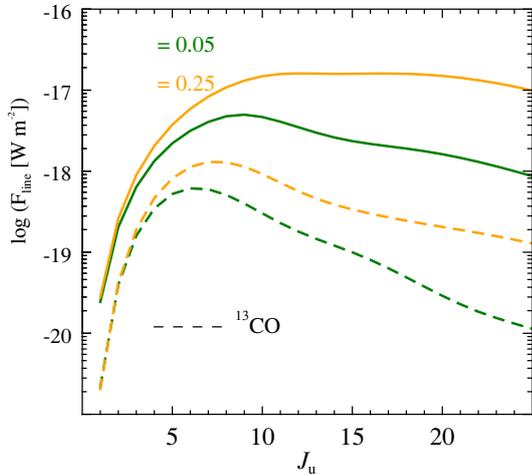
As shown by the line contribution function in Fig. 4, the  $^{13}CO$  emission probes a layer closer to the disk midplane at a smaller distance to the star than the  $^{12}CO$  emission does. The relative fluxes and line profiles of multiple high- $J$  transitions of the two isotopologues provide a strong constraint to the disk temperature structure both in the radial and in the vertical direction.



**Fig. 7.** Synthetic CO ladders in the case of  $T_{\text{gas}} = T_{\text{dust}}$  case (representative models only).



**Fig. 9.** Synthetic CO ladder for the flared disk case ( $\psi = 0.25$ ) with and without dust gap (dashed green and red curve, respectively).



**Fig. 8.** Synthetic  $^{13}\text{CO}$  ladders (representative models only).

Finally, Fig. 9 shows the effect of a dust gap in the CO ladder for the flared ( $\psi = 0.25$ ) disk case; a dust gap is included between 1–13 au with a drop in dust surface density of  $\delta_{\text{dust}} = 10^{-6}$  and a drop in the gas surface density  $\delta_{\text{gas}} = 10^{-6}$ . A structure of this kind is representative of the Herbig Ae group I systems analyzed here (e.g., HD 100546). As the figure shows, the dust gap does not affect the CO ladder for  $J < 25$  because the high- $J$  lines analyzed here are emitted at radii larger than the dust gap.

#### 5.4. CO line profiles

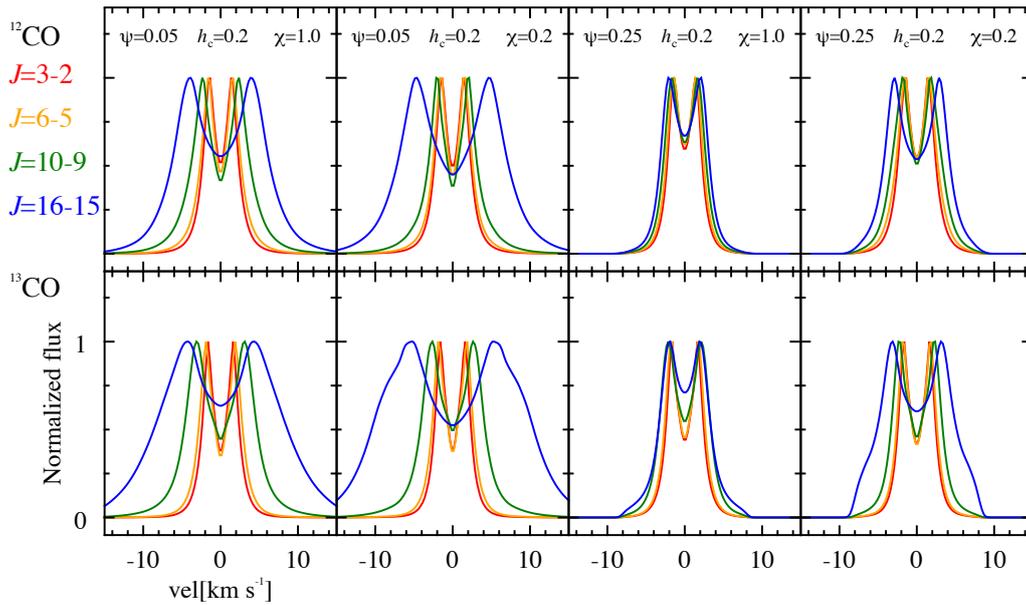
The predicted velocity profiles of different CO transitions are shown in Fig. 10 ( $^{13}\text{CO}$ ) for the representative models. A mix of low- and high- $J$  transitions are plotted for a disk inclination of  $i = 45^\circ$ . In all cases the line width increases with  $J$ . The larger differences in line width are seen for the flat disk models. We note that these differences in line width become particularly significant for  $J = 16$ – $15$ , not yet for  $J = 10$ – $9$ , demonstrating the power of these high- $J$  HIFI data. For  $\psi = 0.05$  the disk scale height and the dust settling have only minor effect on the line velocity profiles. In the case of  $\psi = 0.25$  instead, all the lines become broader if there is substantial dust settling while changing the scale height from  $h_c = 0.2$  to  $0.3$  has little impact (not shown here).

The  $^{13}\text{CO}$  lines are systematically broader than the  $^{12}\text{CO}$  lines (Fig. 10); this is particularly true for the higher  $J$  transitions because the  $^{13}\text{CO}$  transitions become optically thick at higher density and the line emitting region moves deeper into the disk and closer to the star (for a given  $J$ ). This can be seen in Fig. 4, which shows the gas density structure and the line contribution function.

## 6. Discussion and conclusion

The major finding of this paper is that the disk temperature structure varies substantially from system to system (Fig. 3). This result is based on the analysis of the high- $J$  CO line profiles and of the CO rotational ladder. We note in particular that the four Herbig Ae systems have similar stellar mass ( $M_\odot \sim 2$ – $2.5 M_\odot$ ), while the disk temperature can differ by almost an order of magnitude. The comparison of the observed (Fig. 2) and synthetic (Fig. 6) CO ladders suggest a disk temperature sequence for the four Herbig Ae systems studied here (when the different distances of the sources are taken into account). In this sequence HD 100546 is the hottest and HD 163296 the coldest disk. The primary parameters regulating the overall disk temperature are the flaring angle, the scale height, and the gas-to-dust mass ratio; varying these parameters induces a variation in the CO line flux up to 1–2 orders of magnitude. The disk temperature structure is sensitive to the disk mass only for  $M_{\text{disk}} \lesssim 10^{-4} M_\odot$ . Dust settling may also lead to changes in the disk temperature. The power-law exponent of the surface density profile, the stellar temperature and PAH abundance have secondary effects on the temperature of Herbig Ae disks affecting mostly the CO ladder shape, producing a bending in the CO rotational ladder for  $J \gtrsim 10$ . In some cases, changes in the CO ladder are not due to a different temperature structure (Sect. 5.3); for example, a different disk critical radius affects the value of the low- $J$  CO lines ( $J < 10$ ), while the overall temperature structure does not vary substantially.

An interesting outcome of the analysis performed here is that the CO ladder may help constrain some key physical parameters in disks like the gas-to-dust mass ratio, or equivalently the elemental carbon abundance (Bruderer et al. 2012). Spatially resolved millimeter images of the low- $J$  lines allow the accurate measurement of the disk flaring angle and scale height (de Gregorio-Monsalvo et al. 2013), while multifrequency dust continuum observations provide strong constraints on the grain



**Fig. 10.** CO (top) and  $^{13}\text{CO}$  (bottom) line profiles computed with DALI as a function of dust settling and flaring angle for a disk inclination of  $i = 45^\circ$ . The remaining parameters are fixed:  $h_c = 0.2$ ,  $\Delta_{\text{gd}} = 100$ ,  $T_{\text{eff}} = 10\,500\text{ K}$ , PAH = 1%,  $f_{\text{large}} = 0.85$ .

size distribution reducing the degeneracy in the CO ladder. Once these properties are measured, the major differences in the CO ladder are driven by the gas-to-dust mass ratio.

In conclusion, the CO rotational ladder and the velocity profiles of multiple- $J$  transitions are a valid diagnostic of  $T_{\text{gas}}$  in disks. Since the end of the *Herschel* operation, the investigation of the far-infrared spectrum of disks has been limited to the very bright sources observable with SOFIA. Future observations with cryogenic facilities like SPICA will allow the analysis of the CO (and isotopologues) rotational ladder to be expanded to a much larger sample of protoplanetary systems, and will routinely measure the temperature structure of both flared and flat Herbig AeBe and T Tauri disks. In particular, the line flux ratio of several, mid- and high- $J$  CO and  $^{13}\text{CO}$  will provide direct insights into the vertical (different  $\tau$ ) and radial (different  $J$ ) temperature structure of protoplanetary disks if complemented by lower- $J$  spectrally and spatially resolved data from ALMA and other instruments.

*Acknowledgements.* DF acknowledges support from the Italian Ministry of Science and Education (MIUR), project SIR (RBSI14ZRHR). The authors would like to thank the WISH team and L. Podio for providing the HIFI  $^{12}\text{CO}$  and  $^{13}\text{CO}$   $J = 10-9$  spectra of TW Hya and HD 163296. We are also grateful to the DIGIT and GASPS teams for providing the PACS spectra of CO. Astrochemistry in Leiden is supported by the Netherlands Research School for Astronomy (NOVA), by a Royal Netherlands Academy of Arts and Sciences (KNAW) professor prize, and by the European Union A-ERC grant 291141 CHEMPLAN.

## References

Andrews, S. M., Wilner, D. J., Espaillat, C., et al. 2011, *ApJ*, **732**, 42  
Bast, J. E., Brown, J. M., Herczeg, G. J., van Dishoeck, E. F., & Pontoppidan, K. M. 2011, *A&A*, **527**, A119  
Bergin, E. A., Cleeves, L. I., Gorti, U., et al. 2013, *Nature*, **493**, 644  
Bohlin, R. C., Savage, B. D., & Drake, J. F. 1978, *ApJ*, **224**, 132  
Brittain, S. D., Najita, J. R., & Carr, J. S. 2009, *ApJ*, **702**, 85  
Brown, J. M., Pontoppidan, K. M., van Dishoeck, E. F., et al. 2013, *ApJ*, **770**, 94  
Bruderer, S. 2013, *A&A*, **559**, A46  
Bruderer, S., van Dishoeck, E. F., Doty, S. D., & Herczeg, G. J. 2012, *A&A*, **541**, A91

Bruderer, S., van der Marel, N., van Dishoeck, E. F., & van Kempen, T. A. 2014, *A&A*, **562**, A26  
Chapillon, E., Parise, B., Guilloteau, S., Dutrey, A., & Wakelam, V. 2010, *A&A*, **520**, A61  
D'Alessio, P., Calvet, N., Hartmann, L., Franco-Hernández, R., & Servín, H. 2006, *ApJ*, **638**, 314  
de Gregorio-Monsalvo, I., Ménard, F., Dent, W., et al. 2013, *A&A*, **557**, A133  
Dent, W. R. F., Thi, W. F., Kamp, I., et al. 2013, *PASP*, **125**, 477  
Draine, B. T., & Li, A. 2007, *ApJ*, **657**, 810  
Dutrey, A., Guilloteau, S., Prato, L., et al. 1998, *A&A*, **338**, L63  
Fedele, D., Bruderer, S., van Dishoeck, E. F., et al. 2013a, *A&A*, **559**, A77  
Fedele, D., Bruderer, S., van Dishoeck, E. F., et al. 2013b, *ApJ*, **776**, L3  
Fedele, D., Bruderer, S., van den Ancker, M. E., & Pascucci, I. 2015, *ApJ*, **800**, 23  
Geers, V. C., Augereau, J., Pontoppidan, K. M., et al. 2006, *A&A*, **459**, 545  
Gorti, U., & Hollenbach, D. 2008, *ApJ*, **683**, 287  
Green, J. D., Evans, II, N. J., Jørgensen, J. K., et al. 2013, *ApJ*, **770**, 123  
Jonkheid, B., Dullemond, C. P., Hogerheijde, M. R., & van Dishoeck, E. F. 2007, *A&A*, **463**, 203  
Kama, M., Bruderer, S., Carney, M., et al. 2016, *A&A*, **588**, A108  
Kamp, I., Tilling, I., Woitke, P., Thi, W.-F., & Hogerheijde, M. 2010, *A&A*, **510**, A18  
Kamp, I., Thi, W.-F., Meeus, G., et al. 2013, *A&A*, **559**, A24  
Koerner, D. W., & Sargent, A. I. 1995, *AJ*, **109**, 2138  
Mannings, V., & Sargent, A. I. 1997, *ApJ*, **490**, 792  
Meeus, G., Montesinos, B., Mendigutía, I., et al. 2012, *A&A*, **544**, A78  
Meeus, G., Salyk, C., Bruderer, S., et al. 2013, *A&A*, **559**, A84  
Miotello, A., Bruderer, S., & van Dishoeck, E. F. 2014, *A&A*, **572**, A96  
Panić, O., van Dishoeck, E. F., Hogerheijde, M. R., et al. 2010, *A&A*, **519**, A110  
Pontoppidan, K. M., Boogert, A. C. A., Fraser, H. J., et al. 2008, *ApJ*, **678**, 1005  
Pontoppidan, K. M., Blake, G. A., & Smette, A. 2011, *ApJ*, **733**, 84  
Qi, C., D'Alessio, P., Öberg, K. I., et al. 2011, *ApJ*, **740**, 84  
Roelfsema, P. R., Helmich, F. P., Teyssier, D., et al. 2012, *A&A*, **537**, A17  
Salyk, C., Pontoppidan, K., Corder, S., et al. 2014, *ApJ*, **792**, 68  
Sturm, B., Bouwman, J., Henning, T., et al. 2010, *A&A*, **518**, L129  
van den Ancker, M. E., de Winter, D., & Tjin A Dje, H. R. E. 1998, *A&A*, **330**, 145  
van der Plas, G., van den Ancker, M. E., Acke, B., et al. 2009, *A&A*, **500**, 1137  
van der Wiel, M. H. D., Naylor, D. A., Kamp, I., et al. 2014, *MNRAS*, **444**, 3911  
van Kempen, T. A., Kristensen, L. E., Herczeg, G. J., et al. 2010, *A&A*, **518**, L121  
van Zadelhoff, G., van Dishoeck, E. F., Thi, W., & Blake, G. A. 2001, *A&A*, **377**, 566  
Visser, R., Geers, V. C., Dullemond, C. P., et al. 2007, *A&A*, **466**, 229  
Woitke, P., Kamp, I., & Thi, W. 2009, *A&A*, **501**, 383

**Table A.1.** CO rotational ladder fluxes

	TW Hya	HD 97048	HD 100546	IRS 48	HD 163296
	Integrated flux [ $10^{-17}$ W m $^{-2}$ ]				
2-1	0.014 ± 0.0007				0.035 ± 0.00039
3-2	0.043 ± 0.0013		0.15±0.01		0.125 ± 0.00130
4-3					
5-4					1.04 ± 0.40
6-5	0.13 ± 0.03	0.63±0.03	1.3 ± 0.1	>0.17	0.74 ± 0.29
7-6		0.8 ± 0.4			0.90 ± 0.30
8-7		1.8 ± 0.8			1.24 ± 0.55
9-8		2.4 ± 0.5			0.91 ± 0.40
10-9	0.28 ± 0.06	3.5 ± 0.7	4.5 ± 0.1		1.17 ± 0.35
11-10		3.6 ± 0.7			1.13 ± 0.35
12-11		2.8 ± 0.6			1.17 ± 0.35
13-12		4.9 ± 0.9			1.52 ± 0.40
14-13		2.4 ± 1.4	8.5 ± 1.0	<1.8	<1.60
15-14		3.0 ± 0.8	6.9 ± 1.0	1.3 ± 0.5	1.03 ± 0.50
16-15		3.0 ± 0.5	5.6 ± 1.0	1.6 ± 0.4	<1.30
17-16		2.6 ± 0.5	7.1 ± 1.0	1.7 ± 0.4	0.75 ± 0.25
18-17	0.30 ± 0.10	2.6 ± 0.5	5.8 ± 1.0	1.3 ± 0.3	<0.90
19-18		<2.3	7.4 ± 1.0	1.1 ± 0.4	<0.90
20-19		<2.3	5.2 ± 1.0	1.2 ± 0.4	<0.90
21-20		<2.3	8.6 ± 1.0	1.0 ± 0.3	<0.90
22-21		<2.3	7.3 ± 1.0	<1.4	<0.90
23-22	<0.44	<2.3	7.8 ± 1.1	<1.3	<0.90

## Appendix A: CO rotational ladder

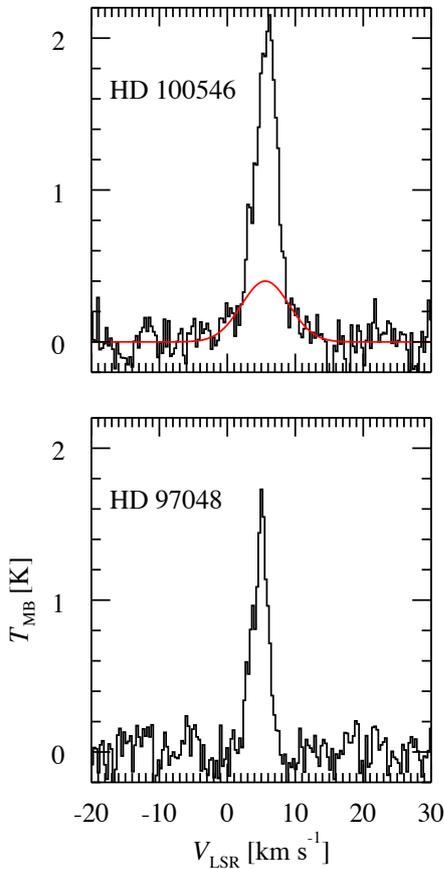
The fluxes of the CO rotational lines up to  $J = 25-24$  are listed in Table A.1. Data are from this work and from [Bergin et al. \(2013\)](#), [Bruderer et al. \(2014\)](#), [Fedele et al. \(2013b\)](#), [Kamp et al. \(2013\)](#), [Meeus et al. \(2013\)](#), [Qi et al. \(2011\)](#), [van der Wiel et al. \(2014\)](#); and references therein); and from [Kama et al. \(2016\)](#). The flux of the  $J = 6-5$  in IRS 48 is a lower limit because the line is partially obscured by the molecular cloud ([Bruderer et al. 2014](#)). Some of the high- $J$  line fluxes ( $J > 14-13$ ) based on *Herschel*/PACS observations have been re-measured with the most recent update of the instrument's calibration and some of the fluxes are slightly different than previous measurements. The  $J = 24-23$  in TW Hya is blended with H<sub>2</sub>O 4<sub>14-3</sub><sub>03</sub>.

## Appendix B: [C II]

[C II] emission is detected at high S/N toward both sources (HD 97048 and HD 100546, Fig. B.1). The line is narrow with a Gaussian FWHM of  $\sim 3\text{--}4\text{ km s}^{-1}$ . In the case of HD 97048 the peak of the emission is at  $v_{\text{LSR}} \sim 5.1\text{ km s}^{-1}$ , redshifted from the center of the CO  $J = 16\text{--}15$  line by  $\sim 0.2\text{ km s}^{-1}$ .

The analysis of the [C II] profile toward HD 100546 is presented in Paper I and the spectrum is shown here for comparison with that of HD 97048. Because of the internal load chop used in combination with an off-source calibration with HIFI, some of the large-scale line emission is removed in the HIFI spectra. Given the large beam of HIFI at this frequency ( $11''$ ), the large-scale non-disk emission contributes most of the [C II] flux and we conclude that the emission measured here is dominated by a diffuse, low-velocity gas not associated with the disk, i.e., a diffuse cloud or remnant envelope on a scale of 1000 au.

In the case of HD 100546, the base of the [C II] line is broader than the core of the line as can be seen from the high-velocity wings. This broadening may be due to an underlying disk contribution. To quantify this, the line wings are fitted with a Gaussian profile centered at the system velocity ( $V_{\text{LSR}} = 5.6\text{ km s}^{-1}$ ). The best-fit profile is shown in Fig. 1 (red curve). The fit is partly degenerate in the Gaussian height and width. Nevertheless, in order to fit the high velocity wings, the FWHM must range between  $\sim 7\text{--}8\text{ km s}^{-1}$ . The disk [C II] contribution is estimated by integrating the flux below the best-fit Gaussian profile and we find a value of  $\sim 8\text{--}9 \times 10^{-17}\text{ W m}^{-2}$ , corresponding to 4–5% of the total line flux. This value must be taken with



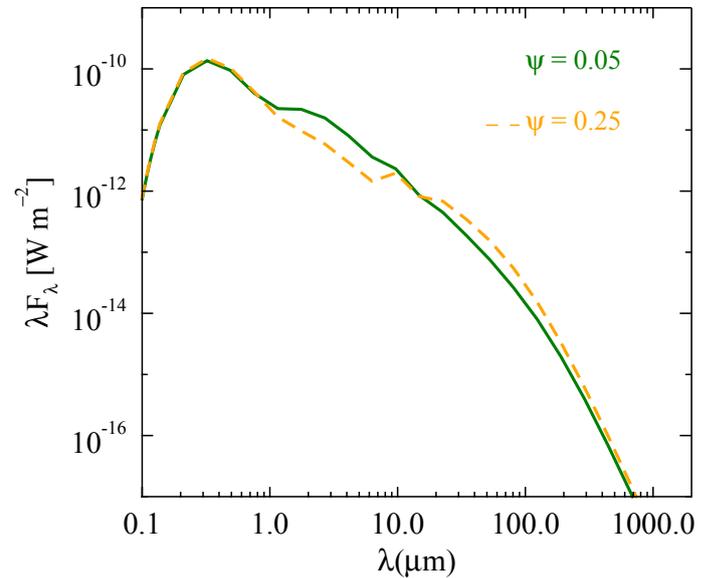
**Fig. B.1.** HIFI/WBS spectra of [C II]. The base of the [C II] line in HD 100546 is fitted with a Gaussian profile (red curve).

caution for various reasons (e.g., the intrinsic line profile is not Gaussian).

No evidence of line broadening is detectable in the HIFI [C II] spectrum of HD 97048. The upper limit to the [C II] disk contribution is measured assuming a line peak equal to  $0.3\text{ K}$  ( $3\times\text{rms}$ , Table 1) and a line width of  $7\text{ km s}^{-1}$  (similar to HD 100546). This gives a  $3\sigma$  upper limit of  $2.4\text{ K km s}^{-1}$  equivalent to  $5.5 \times 10^{-17}\text{ W m}^{-2}$ .

## Appendix C: Spectral energy distribution of representative models

The spectral energy distribution of the two representative disk models is shown in Fig. C.1.



**Fig. C.1.** Spectral energy distribution of the two representative disk models.