

X-Shooter study of accretion in Chamaeleon I[★]

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ABSTRACT

We present the analysis of 34 new VLT/X-Shooter spectra of young stellar objects in the Chamaeleon I star-forming region, together with four more spectra of stars in Taurus and two in Chamaeleon II. The broad wavelength coverage and accurate flux calibration of our spectra allow us to estimate stellar and accretion parameters for our targets by fitting the photospheric and accretion continuum emission from the Balmer continuum down to ~ 700 nm. The dependence of accretion on stellar properties for this sample is consistent with previous results from the literature. The accretion rates for transitional disks are consistent with those of full disks in the same region. The spread of mass accretion rates at any given stellar mass is found to be smaller than in many studies, but is larger than that derived in the Lupus clouds using similar data and techniques. Differences in the stellar mass range and in the environmental conditions between our sample and that of Lupus may account for the discrepancy in scatter between Chamaeleon I and Lupus. Complete samples in Chamaeleon I and Lupus are needed to determine whether the difference in scatter of accretion rates and the lack of evolutionary trends are not influenced by sample selection.

Key words. stars: pre-main sequence – stars: variables: T Tauri, Herbig Ae/Be – accretion, accretion disks – protoplanetary disks – open clusters and associations: individual: Chamaeleon I

1. Introduction

During the first Myr of evolution toward the main sequence, the disk around young stellar objects (YSOs) evolves and is dispersed mainly by disk accretion processes and photoevaporation (Alexander et al. 2014). The importance and the properties of disk accretion are constrained by studying objects located in different star-forming regions and at different evolutionary phases.

Observations of accretion in YSOs mostly focus on determining the amount of material accreted onto the central star per unit time. This mass accretion rate (\dot{M}_{acc} , e.g., Hartmann et al. 1998) may be compared to the stellar and disk properties to test the predictions of models of disk evolution. Relations between the accretion luminosity (L_{acc}) and stellar luminosity (L_{\star} , e.g., Natta et al. 2006; Clarke & Pringle 2006; Rigliaco et al. 2011a; Manara et al. 2012) and between \dot{M}_{acc} and stellar mass (M_{\star} , e.g., Muzerolle et al. 2003; Mohanty et al. 2005; Natta et al. 2006; Herczeg & Hillenbrand 2008; Antonucci et al. 2011, 2014; Manara et al. 2012; Alcalá et al. 2014; Ercolano et al. 2014; Frasca et al. 2015) are frequently derived to examine the dependence of accretion on the properties of the central star. Accretion has sometimes been found to correlate with age (e.g., Hartmann et al. 1998; Sicilia-Aguilar et al. 2010; Antonucci et al. 2014), although ages of individual young stars in a given

star-forming region are very uncertain as a result of both observational and model-dependent uncertainties (e.g., Soderblom et al. 2014), including a dependence of age estimate on the spectral type (e.g., Herczeg & Hillenbrand 2015).

Measurements of \dot{M}_{acc} are obtained spectroscopically using direct or indirect methods. Flux-calibrated spectra that span a wide wavelength range and include the Balmer continuum region yield direct estimates of the L_{acc} while also allowing for the incorporation of veiling into spectral type and extinction measurements. This approach is thought to lead to more reliable estimates for the stellar mass and radius, which are necessary ingredients when estimating \dot{M}_{acc} and ages (Manara et al. 2013b; Herczeg & Hillenbrand 2014, hereafter HH14). While the entire UV spectrum is obtained only from space telescopes and is then used to derive \dot{M}_{acc} for relatively small samples of YSOs (e.g., Gullbring et al. 2000; Ingleby et al. 2013), recent instruments mounted on very large telescopes give the community access to large quantity of spectra in the region of the visible UV from Earth, namely from ~ 300 nm to ~ 400 nm, which is sufficient for capturing the Balmer jump. Modeling the excess continuum emission at these wavelengths yields direct measurements of L_{acc} . This method has been used in the past on a substantial number of objects (e.g., Gullbring et al. 1998; Herczeg & Hillenbrand 2008; Ingleby et al. 2014), but often using non-simultaneous and low-resolution spectra at optical wavelengths to determine the photospheric properties of the target. On the other hand, the indirect method of measuring the luminosity of various emission lines (L_{line}) and converting this

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into L_{acc} through calibrated $L_{\text{acc}}-L_{\text{line}}$ relations (e.g., Muzerolle et al. 1998; Fang et al. 2009) can be efficiently used on spectra covering a relatively narrow wavelength range, such as those offered by multi-object spectrographs. Recent studies have also demonstrated that L_{acc} derived from several emission lines measured simultaneously lead to results compatible, although with a larger scatter, with more direct methods (Rigliaco et al. 2012; Alcalá et al. 2014). Therefore, measuring \dot{M}_{acc} from the luminosity of emission lines is a method that is being extensively used in the literature (e.g., Mohanty et al. 2005; Natta et al. 2006; Fang et al. 2009; Costigan et al. 2012; Antonucci et al. 2011, 2014; Biazzo et al. 2014).

The second-generation VLT instrument X-Shooter allows us to obtain simultaneous spectra of the whole wavelength range from ~ 300 nm to ~ 2500 nm at medium resolution and to thus derive L_{acc} directly from the UV-excess and stellar properties from the photospheric features of the spectra. Manara et al. (2013b) have developed a new method to model all the components of the observed X-Shooter spectra of YSOs to directly derive L_{acc} simultaneously with the stellar parameters and extinction. First studies of accretion with this instrument focused on individual targets (e.g., Rigliaco et al. 2011b; Stelzer et al. 2013) or on a few very low-mass stars and BDs in σ -Orionis (Rigliaco et al. 2012). The largest sample studied to date with this instrument is a set of 36 YSOs in the Lupus I and III clouds (Alcalá et al. 2014). This study found a significantly smaller spread of values of \dot{M}_{acc} at any given M_{\star} with respect to previous works. This result opened new questions in the field, such as whether Lupus is a special region where the initial conditions of star formation are similar for different objects, or if the intrinsic YSO variability in this region is inherently small. At the same time, the narrow \dot{M}_{acc} distribution derived using X-Shooter spectra started to show the advantage of high-quality and broad-band spectra in obtaining more precise accretion rates. A large sample of 22 transitional disks (Manara et al. 2014) has also been studied with this instrument and the method by Manara et al. (2013b), but larger samples of known full disks in various star-forming regions studied with the same method are needed to determine the general evolution of accretion in YSOs.

In this paper we present a study of a large sample of 34 Class II YSOs in the Chamaeleon I star-forming region ($d = 160$ pc, Luhman 2008) observed with X-Shooter, and of a few additional targets in Taurus and Chamaeleon II. These results are used to verify whether the small spread of \dot{M}_{acc} found in Lupus is confirmed in another star-forming region and to extend the comparison between accretion and stellar properties to solar-mass YSOs, which were not targeted by Alcalá et al. (2014). The paper is organized as follows. First, the observations and data reduction procedures are explained in Sect. 2, then the method used to derive stellar and accretion parameters is presented and compared with literature estimates in Sect. 3. The results are then presented in Sect. 4 and are discussed in Sect. 5, including a comparison with other studies with similar methods in other star-forming regions. Finally, Sect. 6 presents the main conclusions of this work.

2. Observations and data reduction

2.1. Main sample

The sample of pre-main sequence (PMS) stars presented in this study was selected from Luhman (2004, 2007), Spezzi et al. (2008), and Alcalá et al. (2008) and has been observed with X-Shooter (Vernet et al. 2011) on the Very Large Telescope

(Paranal, Chile) between January 17 and 19, 2010. X-Shooter provides simultaneous medium-resolution ($R \sim 5000-18\,000$, depending on the spectral arm and the slit width) spectroscopic observations between 300–2500 nm. It consists of three independent cross-dispersed echelle spectrographs, referred to as the UVB, VIS, and NIR arms, with the following spectral coverage: UVB (300–559.5 nm), VIS (559.5–1024 nm), and NIR (1024–2480 nm). The observations were executed by nodding the telescope along a direction perpendicular to the slit. Before the nodding observation, each target was observed in stare mode using a large slit width of $1''.5-5''$. Large-slit spectra have been used to properly flux calibrate the spectra obtained with narrow slits, which are used for the analysis because they lead to a higher spectral resolution. The observing log is reported in Table D.1 and includes the time of observation, the exposure times, the number of nodding positions, and the width of the slits used for the high-resolution observations.

The spectra were reduced with the ESO X-Shooter pipeline (Modigliani et al. 2010) version v1.3.2 following a standard reduction scheme: bias subtraction (stare mode only), flat-fielding, wavelength calibration, background removal (stare mode only), and spectrum extraction. The latter is performed manually on the rectified spectrum, and after the extraction of each component of a binary system when both were included in the slit and spatially separated. For the spectra taken in the nodding mode, the background removal is obtained by the difference of spectra pairs. The UVB and VIS arms are corrected for atmospheric extinction using the atmospheric transmission curve of Paranal by Patat et al. (2011).

The spectrophotometric standard star GD-71 was observed at the beginning of each night. The spectral response function of the instrument, which gives the conversions from ADU to $\text{erg cm}^{-2} \text{s}^{-1} \text{\AA}^{-1}$, was obtained by dividing the observed spectrum by the tabulated spectrum provided with the X-Shooter pipeline. This has been derived from observations of GD-71 for each night and each arm independently. The response function for the UVB and VIS arms does not show variations during the three nights of observation. Small variations are observed for the NIR arm, and the response function is taken as the average of the three nights. The overall flux calibration accuracy was tested on telluric standard star observations and is stable with a $\sim 2\%$ accuracy. Finally, the narrow-slit spectra are scaled upward to match the large-slit spectra. The scaling factor is computed as the ratio of the spectral continuum of the two spectra. More information on the reduction and flux calibration of the spectra obtained in this program will be discussed in Rugel et al. (in prep.).

In total, 40 bona-fide Class II members of the Chamaeleon I star-forming region were observed. Here we discuss the properties of 34 of these objects. The other six objects are transitional disks sources that were analyzed by Manara et al. (2014) with the same method as was implemented in this paper, and they are included in our analysis. The membership of 22 YSOs in the Chamaeleon I sample has been confirmed by recent proper motion studies (Lopez Martí et al. 2013). Following Luhman (2008, and references therein), the adopted distances to objects in the Chamaeleon I region is 160 pc. In addition to these targets, two objects located in the Chamaeleon II cloud ($d = 178$ pc, Luhman 2008), both confirmed members according to Lopez Martí et al. (2013), and four targets located in the Taurus molecular cloud ($d = 131-140$ pc, HH14) were observed, and their accretion properties are reported here. These targets were observed with the same observing strategy as the Chamaeleon I objects and thus are interesting objects with which to compare the accretion properties of YSOs in different star-forming regions.

2.2. Additional templates used for the analysis

The analysis presented here makes use of several photospheric templates observed with X-Shooter to estimate the stellar parameters of the targets. These templates are non-accreting Class III YSOs collected and characterized mainly by [Manara et al. \(2013a\)](#). This set of templates is almost complete for objects with SpT in the M-subclass, but includes only few objects with SpT in the K-subclass (one K5, two K7 YSOs). Thus, it is useful to include some additional templates here. Three of them have been discussed in [Manara et al. \(2014\)](#) and have SpT G4, G5, and K2. Finally, another target, HBC407, is included here. This object is a Class III YSO ([Luhman et al. 2010](#)) with SpT K0 (HH14), which was observed during the program Pr.Id.094.C-0913 (PI Manara). Here we discuss the observations, data reduction, and classification for this target.

HBC407 was observed in service mode during the night of December 8, 2014 with the following strategy. First, a complete nodding cycle ABBA was performed using the narrower slits 0.5''–0.4''–0.4'' in the three arms, respectively, with an exposure time for each exposure of 430 s in the UVB arm, 375 s in the VIS arm, and three subintegrations of 125s each in the NIR arm. Then, the object was observed immediately after this with the larger slits (5.0'') with an AB nodding cycle with single exposures of 45 s in the UVB and VIS arms and three times 15 s in the NIR arm. The latter was used to obtain a proper flux calibration of the spectra without slit losses. This spectrum was reduced using the X-Shooter pipeline version 2.5.2, and flux calibration of the spectra was made by the pipeline using the standard star FEIGE-110. We then manually rescaled the spectra obtained with the narrow slits to those obtained with the large slits and found that they agree very well with the available photometry. Telluric correction was carried out using the telluric standard Hip036392 and the procedure described by [Alcalá et al. \(2014\)](#).

This object has been classified as a K0 YSO with extinction $A_V = 0.8$ mag by [HH14](#). Using the spectral indices discussed in that work, in particular using the R5150 index, we obtain the same SpT for this target. We then compared the observed spectrum with a BT-Settl AGSS2009 synthetic spectrum ([Allard et al. 2011](#)) at $T_{\text{eff}} = 5100$ K and $\log g = -4.0$ and found a best agreement by applying a reddening correction of $A_V = 0.8$ mag using the extinction law by [Cardelli et al. \(1989\)](#) and $R_V = 3.1$. This T_{eff} is consistent with a SpT between G8 and K1 according to different SpT– T_{eff} relations ([HH14](#) and [Kenyon & Hartmann \(1995\)](#), respectively). Here a SpT K0 is assumed for this target. Finally, we calculated the luminosity of the target from the observed spectrum and the synthetic spectrum following the same procedure as [Manara et al. \(2013a\)](#), and, assuming $d = 140$ pc ([HH14](#)), we derive $\log(L_{\star}/L_{\odot}) = -0.45$ for this target, the same value as [HH14](#). However, the position of this target is below the main sequence. This could suggest that it is not a member of the Taurus region. Nevertheless, this uncertainty does not influence the estimates of stellar parameters for other targets based on this template because the information on the distance is taken into account in the procedure.

The spectrum of this target shows no signs of accretion in the typical lines tracing accretion in YSOs (e.g., $H\alpha$), but a prominent LiI $\lambda 670.8$ nm absorption line, and narrow chromospheric emission lines as a reversal in the core of the photospheric absorption lines in the Ca IR triplet.

3. Analysis method

Here we briefly explain our method to self-consistently derive stellar and accretion parameters from the X-Shooter spectra, and we report our results. All the objects discussed here have also been previously characterized by other authors (e.g., [Luhman 2007](#), see Table A.1) using red optical and infrared spectra. The comparison with these results is also presented in this section.

3.1. Determining the stellar and accretion parameters

To determine the stellar and accretion parameters, we used the method described in [Manara et al. \(2013b\)](#), which determines the SpT, A_V , L_{\star} , and L_{acc} by finding the best fit among a grid of models. This grid includes different photospheric templates (see Sect. 2.2), several slab models to reproduce the excess continuum emission due to accretion, and extinction parametrized with the reddening law by [Cardelli et al. \(1989\)](#) with $R_V = 3.1$ and values of A_V from 0 mag to 10 mag with steps of 0.1 mag. The slab model has been used by [Valenti et al. \(1993\)](#) and [Herczeg & Hillenbrand \(2008\)](#), for example, and it is not a physical model of the accretion region, unlike the shock models by [Calvet & Gullbring \(1998\)](#), but provides an accurate bolometric correction to determine the total excess flux that is due to accretion. The best fit is found by comparing the flux of the observed spectrum and that of each model in various continuum regions and in different features of the spectra, and by minimizing a χ^2_{like} distribution to determine the best match. The points included in the fitting procedure range from $\lambda \sim 330$ nm to $\lambda \sim 715$ nm. A proper fit of the Balmer continuum is needed to determine L_{acc} and the veiling of the spectra. The Balmer continuum is well determined even with spectra with very low signal-to-noise ratio (S/N) on the continuum at $\lambda \sim 330$ nm. The spectra of most of the targets discussed here have enough signal to perform the analysis. Only spectra with no signal in the Balmer continuum regions, such as those of Cha H α 9 and Cha H α 1 in this work, are not suitable for the method described here. More discussion of these and other individual targets is available in Appendix B. The stellar parameters (T_{eff} , L_{\star}) are derived from the properties of the best-fit template and L_{acc} from the slab model parameters as described by [Manara et al. \(2013b\)](#), and are reported in Table 1. The best fit of the UV excess for all the targets are shown in Fig. C.1. We have compared our direct measurements of L_{acc} from the excess Balmer continuum emission to alternative estimates from line luminosities in the same spectrum. The agreement is good when using the $L_{\text{acc}}-L_{\text{line}}$ relations by [Alcalá et al. \(2014\)](#), with only small differences for the objects with SpT K2. For further analyses of the emission lines present in the spectra, we refer to Fedele et al. (in prep.).

The HR diagram (HRD) of the Chamaeleon I sample is shown in Fig. 1 overplotted on the isochrones from the evolutionary models by [Baraffe et al. \(1998\)](#). This evolutionary model was chosen for consistency with previous analyses of X-Shooter spectra, so that the comparison is less model dependent. Most of the objects are located on this diagram around the 3 Myr isochrone, with a few of objects located closer to 10 Myr, and only three objects are located slightly below the 30 Myr isochrone, but still above the 30 Myr isochrone from other models (e.g., [Siess et al. 2000](#)). The stellar mass (M_{\star}) was obtained by comparing the position in the HRD to the [Baraffe et al. \(1998\)](#) evolutionary tracks. Finally, the stellar radius (R_{\star}) was derived from T_{eff} and L_{\star} , and \dot{M}_{acc} from the usual relation $\dot{M}_{\text{acc}} = 1.25 \cdot L_{\text{acc}} R_{\star} / (GM_{\star})$ (e.g., [Gullbring et al. 1998](#)).

Table 1. Derived stellar and accretion parameters for the Chamaeleon I sample.

Object	SpT	T_{eff} [K]	A_V [mag]	L_{\star} [L_{\odot}]	$\log L_{\text{acc}}$ [L_{\odot}]	M_{\star} [M_{\odot}]	$\log \dot{M}_{\text{acc}}$ [M_{\odot}/yr]
T3	K7	4060	2.6	0.18	-1.25	0.79	-8.62
T3-B	M3	3415	1.3	0.19	-1.66	0.37	-8.53
T4 [†]	K7	4060	0.5	0.43	<-2.24	1.03	<-9.53
TW Cha	K7	4060	0.8	0.38	-1.66	1.00	-8.96
CR Cha*	K0	5110	1.3	3.26	-1.42	1.78	-8.71
T12	M4.5	3200	0.8	0.15	-2.12	0.23	-8.77
CT Cha A	K5	4350	2.4	1.50	0.37	1.40	-6.85
ISO-ChaI 52 [†]	M4	3270	1.2	0.09	<-3.79	0.25	<-10.63
Hn5	M5	3125	0.0	0.05	-2.56	0.16	-9.27
T23	M4.5	3200	1.7	0.32	-1.65	0.33	-8.29
T24	M0	3850	1.5	0.40	-1.48	0.91	-8.68
Cha H α 1	M7.5	2795	0.0	0.00	-5.11	0.05	-11.69
Cha H α 9	M5.5	3060	4.8	0.03	-4.19	0.10	-10.85
Sz22	K5	4350	3.2	0.51	-1.03	1.09	-8.37
VW Cha	K7	4060	1.9	1.64	-0.78	1.24	-7.86
ESO H α 562	M1	3705	3.4	0.12	-2.01	0.66	-9.32
T33-A [†]	K0	5110	2.5	1.26	<-1.62	1.15	<-9.93
T33-B	K0	5110	2.7	0.69	-1.32	0.95	-8.67
ISO-ChaI 143	M5.5	3060	1.3	0.03	-3.38	0.11	-10.02
Cha H α 6	M6.5	2935	0.1	0.07	-3.86	0.10	-10.25
T38	M0.5	3780	1.9	0.13	-2.02	0.71	-9.35
T40	M0.5	3780	1.2	0.55	-0.48	0.87	-7.61
T44	K0	5110	4.1	2.68	0.62	1.40	-6.60
T45a [†]	K7	4060	1.1	0.34	<-2.59	0.97	<-9.91
ISO-ChaI 237	K5	4350	4.1	0.61	-2.47	1.15	-9.79
T49	M3.5	3340	1.0	0.29	-0.81	0.36	-7.56
CHX18N	K2	4900	0.8	1.03	-0.74	1.17	-8.06
T51	K2	4900	0.1	0.64	-0.79	0.97	-8.13
T51-B	M2	3560	0.5	0.09	-1.92	0.51	-9.12
T52	K0	5110	1.0	2.55	-0.19	1.40	-7.42
T54-A [†]	K0	5110	1.2	2.51	<-2.29	1.40	<-9.53
Hn17 [†]	M4.5	3200	0.4	0.11	<-3.05	0.22	<-9.76
Hn18 [†]	M4	3270	0.8	0.11	<-3.05	0.26	<-9.85
Hn21W	M4.5	3200	2.2	0.12	-2.40	0.22	-9.08

Notes. ([†]) Non-accreting objects according to the H α line width at 10% of the peak and H α line equivalent width. All stellar parameters have been derived using the Baraffe et al. (1998) evolutionary models apart from objects with a (*) symbol, for which the Siess et al. (2000) models were used.

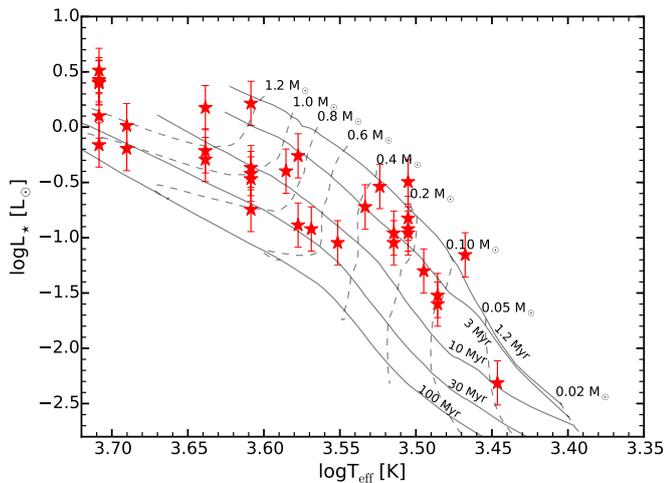


Fig. 1. H-R diagram of the Chamaeleon I objects discussed here overplotted on evolutionary models from Baraffe et al. (1998).

The distribution of the derived parameters SpT, A_V , and M_{\star} for the Chamaeleon I targets is shown in Fig. 2. As expected,

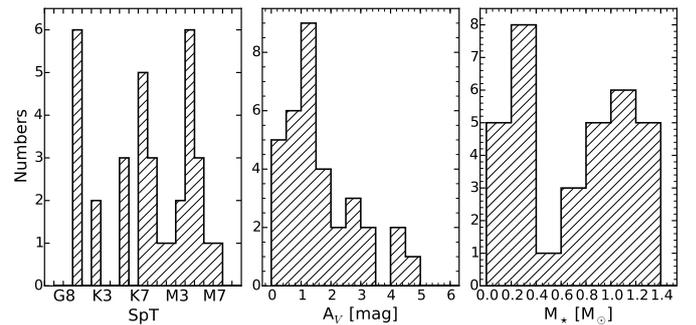


Fig. 2. Histogram of stellar properties (SpT, A_V , and M_{\star}) for the Chamaeleon I sample.

most objects have low $A_V \lesssim 3$ mag. The Chamaeleon I sample discussed here, together with the six transitional disks in Chamaeleon I discussed by Manara et al. (2014), includes more than 40% of the known Class II and transitional disks YSOs in the region, with higher completeness levels at SpT earlier than M3, corresponding to $M_{\star} \sim 0.5 M_{\odot}$, than at later SpT. This sample is thus large enough to discuss general properties of YSOs with $M_{\star} \gtrsim 0.5 M_{\odot}$ in the Chamaeleon I region alone.

Table 2. Derived stellar and accretion parameters for the other targets in this study.

Object	SpT	T_{eff} [K]	A_V [mag]	L_\star [L_\odot]	$\log L_{\text{acc}}$ [L_\odot]	M_\star [M_\odot]	$\log \dot{M}_{\text{acc}}$ [M_\odot/yr]
Taurus							
FN Tau	M3	3415	1.1	0.89	-1.77	0.33	-8.25
V409 tau	M0	3850	1.9	0.26	-1.72	0.85	-8.99
IQ Tau	M0	3850	1.7	0.75	-1.09	0.99	-8.19
DG Tau	K5	4350	2.2	1.55	-0.44	1.40	-7.65
Chamaeleon II							
Hn24	M0	3850	1.9	0.85	-1.99	1.00	-9.07
Sz50	M4	3270	1.8	0.57	-2.51	0.27	-8.96

3.2. Non-accreting targets in Chamaeleon I

All targets analyzed here have excess dust emission in *Spitzer* photometry and spectroscopy (e.g., Luhman et al. 2008; Manoj et al. 2011), indicating the presence of a disk in the system. Some of these, however, show no evidence of ongoing accretion from the lack of measurable excess in the Balmer continuum. In this section we discuss the properties of another proxy of accretion present in our spectra, the $H\alpha$ line profile. The equivalent width (EW) and width at 10% of the peak ($W_{10\%}$) of the $H\alpha$ line can be used to distinguish objects with ongoing accretion, although with discrepant results with respect to the Balmer continuum excess in some cases. According to White & Basri (2003), an object is defined as an accretor when $W_{10\%} > 270 \text{ km s}^{-1}$, and EW is higher than a threshold that depends on the SpT of the target. According to the White & Basri (2003) criterion, seven targets in the sample are non-accreting: T4, ISO-ChaI 52, T33A, T45a, T54-A, Hn17, and Hn18. The fitting of the observed spectrum was also performed for these non-accreting YSOs. In most cases, the accretion excess in the Balmer continuum with respect to the photosphere is very small, within the noise of the spectra, or absolutely negligible, and this is compatible with the non-accreting status of these targets as suggested by the $H\alpha$ line profile. The value of L_{acc} derived in these cases is to be considered as an upper limit on the accretion rate. However, it seems that the excess in the Balmer continuum is non-negligible for ISO-ChaI 52, in particular, although we measure $W_{10\%} \sim 190 \text{ km s}^{-1}$ and $EW = -10.6 \text{ \AA}$. This object has been reported to accrete at a higher rate by Antonucci et al. (2011), but their result is based on the flux of only few near-infrared emission lines, and they did not self-consistently derive the stellar parameters of the targets. It is also possible that the object is now accreting at a lower rate than in the past. For another non-accreting target, T4, Frasca et al. (2015) reported a higher value of $W_{10\%} \sim 390 \text{ km s}^{-1}$ than was measured from our spectra, $W_{10\%} \sim 230 \text{ km s}^{-1}$. This could be a variable object with a negligible accretion rate at the time of our observations.

3.3. Targets in Chamaeleon II and Taurus

While the vast majority of the targets discussed in this work are located in the Chamaeleon I region, four additional targets are located in Taurus and two are located in Chamaeleon II. These were analyzed with the same method as described in Sect. 3.1, and their properties are reported in Table 2. All these targets have a detectable accretion excess in the Balmer continuum and are

also accretors according to the EW and $W_{10\%}$ of the $H\alpha$ line criteria. Two objects, namely FN-Tau and Sz50, have $W_{10\%}$ smaller than 270 km s^{-1} , but their EWs are larger than the threshold for their SpT. Only Hn24 presents a negligible excess in the Balmer continuum. This object has a very wide $W_{10\%}$ of more than 300 km s^{-1} . We report the accretion rate detected for this target, but the small excess and the small EW of the $H\alpha$ line indicate that it might be non-accreting.

3.4. Comparison with other methods

The stellar parameters (SpT, A_V , L_\star) for all the targets discussed here have previously been derived by several studies (e.g., Luhman 2007; Spezzi et al. 2008; Daemgen et al. 2013; Frasca et al. 2015, cf. Tables A.1–A.2). Accretion rate estimates are only available for some targets (e.g., Antonucci et al. 2011, 2014; Frasca et al. 2015). Here we compare these results with our estimates.

Differences between our results and those in the literature are typically less than a spectral subclass in SpT and $\sim 0.3 \text{ mag}$ in A_V , both within the uncertainties. The largest differences in SpT occur for K-type stars, probably because our moderate-resolution blue spectra include better temperature diagnostics (e.g., Covey et al. 2007) than the low-resolution red spectra used in most previous studies of the region. The scarcity of templates in this work in the early-K spectral range may also contribute to this difference. Differences in A_V estimates may be related to different methods, as Luhman (2007) used infra-red photometry and a different R_V to derive extinction, while here the whole optical X-Shooter spectrum was used and A_V is derived together with SpT and L_{acc} .

The results derived with the method used here were also checked against other methods to derive SpT and L_\star from the observed spectra. In particular, the SpT determined here are also compared with those obtained using the spectral indices by HH14, by Riddick et al. (2007), and the TiO index by Jeffries et al. (2007). In general, the agreement is very good with only a few exceptions in cases where different indices would lead to different results. Similarly, L_\star derived here agrees well, with differences in all the cases but one smaller than a factor ~ 1.6 , with those obtained using the bolometric correction by HH14 after removing veiling.

Finally, Fig. 3 shows the comparison between the values of \dot{M}_{acc} for Chamaeleon I targets derived here and those from the POISSON survey (Antonucci et al. 2011, 2014) and the Gaia-ESO survey (GES, Frasca et al. 2015). The POISSON survey used infrared spectra and the luminosity of the Bry and Pa β emission lines to derive L_{acc} , and obtained \dot{M}_{acc} assuming the stellar parameters for the targets from Luhman (2007). These values of \dot{M}_{acc} are different from those derived here in most cases. The reason for the difference is the use of a single indirect accretion indicator, and also the relation between the luminosity of the Bry line and L_{acc} they used, which differs from the more recent one by Alcalá et al. (2014). The estimates of \dot{M}_{acc} from the GES were all derived using the EW and 10% width of the $H\alpha$ emission line. Although our accretion rates generally agree with those from GES, some large discrepancies for individual objects are present. The $H\alpha$ line, and especially its 10% width, is known not to be the best tracer of an accretion rate (e.g., Costigan et al. 2012), although it is widely used because it is the brightest emission line in YSOs. This comparison presented here confirms that single values of \dot{M}_{acc} from this indicator alone should be considered with caution.

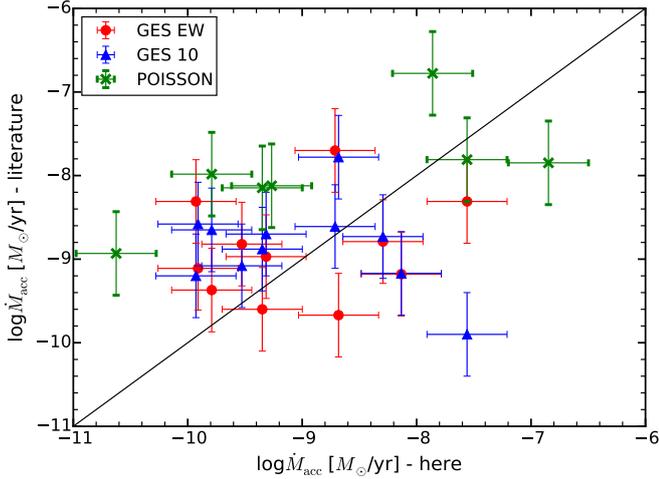


Fig. 3. Comparison of \dot{M}_{acc} derived here for targets in Chamaeleon I with those from the literature, namely the POISSON survey (Antoniucci et al. 2011, 2014) and the Gaia-ESO survey (GES Frasca et al. 2015). The latter are derived either from the equivalent width of the H α line (EW) or from the 10% width of the same line (10).

4. Results

The dependence of the accretion rates on the stellar parameters is presented here. The relation between L_{acc} and L_{\star} is first discussed for the Chamaeleon I sample and in comparison with other large X-Shooter sample of YSOs, then the $\dot{M}_{\text{acc}}-M_{\star}$ relation is presented. The results presented here are to be considered as only representative for the range of M_{\star} that is covered with high completeness by this survey, namely at $M_{\star} \gtrsim 0.5 M_{\odot}$.

4.1. Accretion luminosity and stellar luminosity dependence

The L_{acc} vs. L_{\star} relation for the Chamaeleon I sample is shown in Fig. 4. The targets mostly occupy the region between $L_{\text{acc}} = L_{\star}$ and $L_{\text{acc}} = 0.01 L_{\star}$, although not uniformly at all stellar luminosity. Indeed, objects with $L_{\star} \gtrsim 0.1 L_{\odot}$ have L_{acc}/L_{\star} ratios ranging from ~ 0.01 up to ~ 1 , or even more in two cases. These two strongly accreting targets both have $L_{\star} \sim L_{\odot}$. In contrast, objects with low stellar luminosity ($L_{\star} \lesssim 0.1 L_{\odot}$) have $L_{\text{acc}} \lesssim 0.1 L_{\star}$, and the objects with lowest luminosity have even $L_{\text{acc}} < 0.01 L_{\star}$. However, the sample discussed here comprises few objects in this range. Non-accreting objects, as expected, are located in the lower part of the plot, and they always have $L_{\text{acc}} \lesssim 0.01 L_{\star}$.

The positions of the targets on the $L_{\text{acc}}-L_{\star}$ plane was fitted with the statistics package ASURV (Feigelson & Nelson 1985) run in the IRAF¹ environment, considering the value of L_{acc} for the non-accreting objects as an upper limit. With the whole sample, we then find a best-fit relation of

$$\log L_{\text{acc}} = (1.58 \pm 0.24) \cdot \log L_{\star} - (1.16 \pm 0.21) \quad (1)$$

with a standard deviation around the best fit of 0.89. If only accreting objects are analyzed, the best-fit line is

$$\log L_{\text{acc}} = (1.72 \pm 0.18) \cdot \log L_{\star} - (0.75 \pm 0.16) \quad (2)$$

with a standard deviation of 0.62. The latter relation is shown in Fig. 4.

¹ IRAF is distributed by National Optical Astronomy Observatories, which is operated by the Association of Universities for Research in Astronomy, Inc., under cooperative agreement with the National Science Foundation.

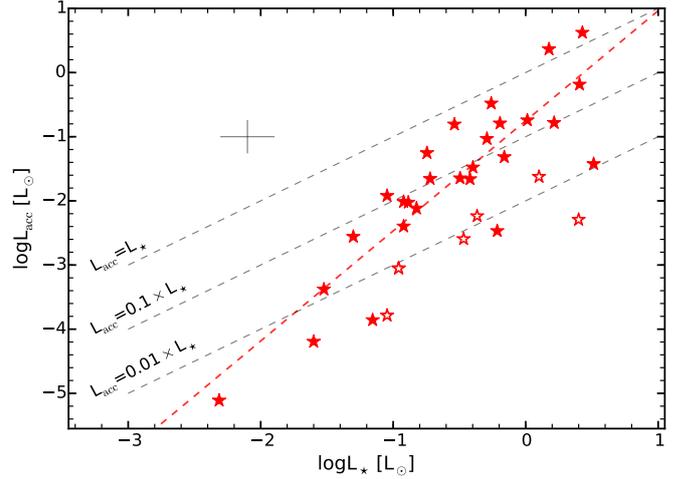


Fig. 4. Accretion luminosity as a function of stellar luminosity for the whole Chamaeleon I sample discussed here (red stars). Empty symbols denote non-accreting objects. Dashed lines are for different L_{acc}/L_{\star} ratios in decreasing steps from 1, to 0.1, to 0.01, as labeled. The best-fit relation from Eq. (2) is overlotted with a dashed red line. Typical uncertainties are shown in the upper left corner of the plot.

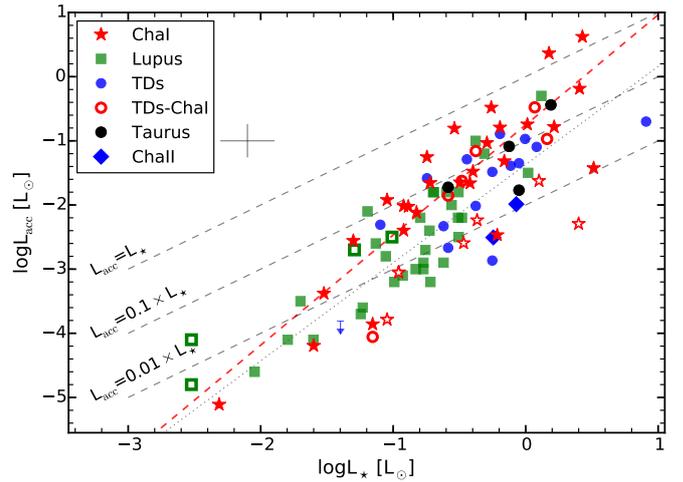


Fig. 5. Accretion luminosity as a function of stellar luminosity for the whole Chamaeleon I sample of Class II discussed here (red filled stars for accreting objects and empty stars for non-accreting objects) and for the Lupus sample of Alcalá et al. (2014), reported with green squares, and with empty green squares for subluminal targets. Symbols for objects in other regions and at other evolutionary stages, such as transitional disks, are reported in the legend. In particular, transitional disks located in the Chamaeleon I region are shown with red empty circles. Dashed lines are for different L_{acc}/L_{\star} ratios in decreasing steps from 1, to 0.1, to 0.01, as labeled. The blue upper limit is a transitional disk.

In Fig. 5 the values of L_{acc} vs. L_{\star} for the Chamaeleon I targets are compared with those in Taurus and Chamaeleon II analyzed here, with the samples of Class II YSOs in the Lupus clouds analyzed by Alcalá et al. (2014), and with transitional disks in various regions, including Chamaeleon I, by Manara et al. (2014). These samples of X-Shooter spectra comprise the largest set of flux-calibrated broadband spectra of young stars used to simultaneously measure accretion, photospheric properties, and extinction. The distribution of stellar parameters for the objects in this sample is different, and comparisons can only be

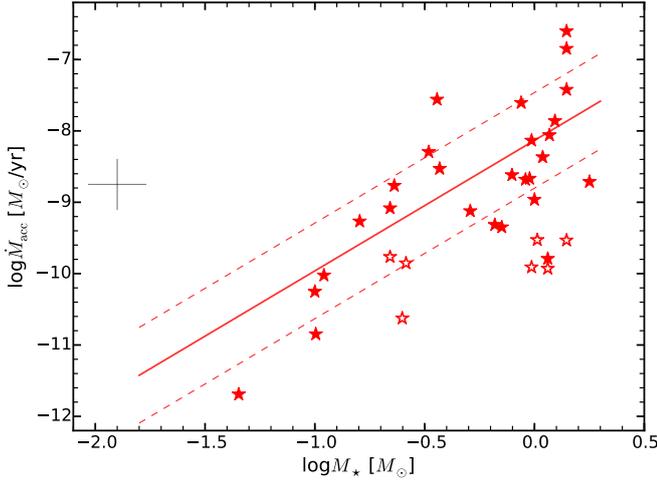


Fig. 6. Mass accretion rate as a function of the stellar mass for the whole Chamaeleon I sample discussed here (red stars). Empty symbols are used for non-accreting objects. The best-fit relation from Eq. (5) is overplotted with a red line, and the standard deviation from the best fit is shown with dashed lines. Typical uncertainties are shown in the upper left corner.

performed by extrapolating the results to other stellar parameters, for instance, L_\star ranges.

The Taurus targets analyzed here follow the $L_{\text{acc}}-L_\star$ distribution of the Lupus targets, while the Chamaeleon II targets are located in the lower part of the distribution. With respect to the best fit of the sample of Alcalá et al. (2014) (shown in the figure as a black dotted line, from Natta et al. 2014), the Chamaeleon I targets follow a steeper slope of L_{acc} vs. L_\star . However, the two values are compatible. Indeed, the slope of the best fit for this relation using the Chamaeleon I data alone and excluding non-accreting objects is 1.7 ± 0.2 , while the slope derived from the Lupus sample is 1.5 ± 0.2 . In the Chamaeleon I sample there are more strongly accreting YSOs with $L_{\text{acc}} \sim L_\star$ than in the Lupus sample. This difference in the fraction of strongly accreting YSOs in the Chamaeleon I sample compared to that of Alcalá et al. (2014) may be the result of different selection criteria (see discussion in Sect. 5.2). In any case, the sample analyzed here confirms that the relation between L_{acc} and L_\star has a slope significantly steeper than 1 and shows that there are several very strongly accreting YSOs in Chamaeleon I. In contrast, no strongly accreting YSOs are present in the Lupus sample analyzed by Alcalá et al. (2014).

The distribution of the accreting transitional disks located in the Chamaeleon I region shown in Fig. 5 also seems to follow the best fit of the $L_{\text{acc}}-L_\star$ relation of the Class II YSOs in this region fairly well, with the only exception of CHXR22E, which is a non-accreting transitional disks that has been discussed by Manara et al. (2014). At the same time, the transitional disks from other regions are also mostly located in the same part of the $L_{\text{acc}}-L_\star$ plane as the majority of Class II YSOs, mostly following the locus of Lupus targets.

4.2. Mass accretion rate and stellar mass dependence

The dependence of \dot{M}_{acc} on M_\star in the Chamaeleon I sample is shown in Fig. 6. All the non-accreting objects are located in the lower part of the distribution. Only ISO-ChaI 237 has a similar \dot{M}_{acc} as non-accreting objects, and it could also be borderline accretor. A tight correlation between \dot{M}_{acc} and M_\star is present,

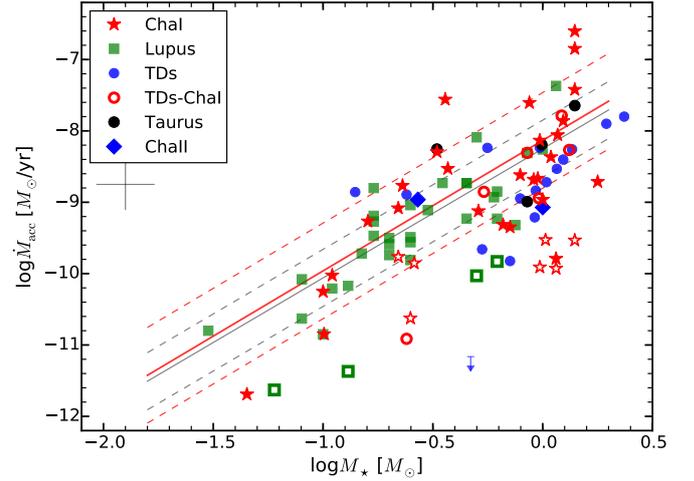


Fig. 7. Mass accretion rate as a function of mass for the whole Chamaeleon I sample of Class II discussed here (red filled stars for accreting objects and empty stars for non-accreting objects) and for the Lupus sample of Alcalá et al. (2014), reported with green squares, and with empty green squares for subluminal targets. Symbols for objects in other regions and at other evolutionary stages, such as transitional disks, are reported in the legend. In particular, transitional disks located in the Chamaeleon I region are shown with red empty circles. The blue upper limit is a transitional disk. The gray continuous line represents the linear fit of this relation for the Lupus sample (Alcalá et al. 2014), and the dashed lines represent the 1σ deviation from the fit, while the red lines are the same for the Chamaeleon I sample analyzed here.

which is a well-known result in the literature. The best fit for this relation including non-accreting objects, again considered as upper limit on \dot{M}_{acc} , is

$$\log \dot{M}_{\text{acc}} = (1.70 \pm 0.44) \cdot \log M_\star - (8.65 \pm 0.22), \quad (3)$$

with a standard deviation of 1.04. When excluding the non-accreting YSOs, the best fit is

$$\log \dot{M}_{\text{acc}} = (1.94 \pm 0.34) \cdot \log M_\star - (8.21 \pm 0.17), \quad (4)$$

with a smaller standard deviation of 0.75.

However, several targets are outliers in this plot. Excluding Cha-Ha1, which is the only brown dwarf in the sample, and ISO-ChaI 237, whose position in the plot is in the same region as the non-accreting YSOs, the following best fit of the $\dot{M}_{\text{acc}}-M_\star$ relation for the accretors in Chamaeleon I is obtained:

$$\log \dot{M}_{\text{acc}} = (1.83 \pm 0.35) \cdot \log M_\star - (8.13 \pm 0.16) \quad (5)$$

with a smaller standard deviation around the best fit of 0.67. Clearly, these two outliers influence the results substantially. This best fit and its standard deviation are shown in Fig. 6.

Figure 7 shows the $\dot{M}_{\text{acc}}-M_\star$ relation for the same samples as in Fig. 5, that is, the Chamaeleon I targets and those in Taurus and Chamaeleon II analyzed here, the Lupus sample by Alcalá et al. (2014), and the transitional disks by Manara et al. (2014). The few measurements in Taurus and Chamaeleon II are consistent with the relations found in Chamaeleon I and Lupus. The slope of this relation is very similar between the Chamaeleon I sample (1.83 ± 0.35), when excluding the outliers and non-accreting YSOs, and the Lupus sample (1.81 ± 0.20). The most relevant difference is then the larger spread around the best fit that is found in the Chamaeleon I sample (0.67 dex) with respect to the spread of the Lupus targets (only 0.4 dex). In particular,

there are some objects with \dot{M}_{acc} higher than that of objects with the same M_{\star} in Lupus at $M_{\star} \gtrsim 0.2 M_{\odot}$, but also several objects with lower \dot{M}_{acc} than the Lupus sample at $M_{\star} \sim 1 M_{\odot}$. It is also clear from Fig. 7 that there are several accreting Class II YSOs in Chamaeleon I with \dot{M}_{acc} similar to those of transitional disks, in particular similar to the transitional disks in Chamaeleon I, but also to those in other regions.

5. Discussion

5.1. Accretion rates in full and transitional disks

The results presented in Sect. 4 show that L_{acc} and \dot{M}_{acc} scale with the stellar properties, respectively L_{\star} and M_{\star} , with a high degree of correlation. In particular, the samples presented here and those from the literature include objects at different evolutionary stages of the disks, mainly Class II YSOs, that is, objects surrounded by a full disk, and transitional disks. Although the comparison sample of transitional disks by Manara et al. (2014) included here is small and biased toward known accretors, the comparison of \dot{M}_{acc} in transitional disks and full disks presented here shows no substantial differences. The values of \dot{M}_{acc} are similar for full and transitional disks even at $M_{\star} \sim M_{\odot}$. According to Najita et al. (2015), this is a result of the selected comparison sample. Indeed, they suggested that age differences between the targets might result in smaller differences in \dot{M}_{acc} , as \dot{M}_{acc} is expected to decrease with age because of the viscous evolution of the disks (e.g., Hartmann et al. 1998). However, Fig. 7 shows that transitional disks located in Chamaeleon I have \dot{M}_{acc} similar to full disks in the same star-forming region. A difference between the full- and transitional-disk Chamaeleon I samples is that there are no transitional disks that are as strongly accreting as the strongest accretor from the sample of full disks. These results confirm that the sample of accreting transitional disks by Manara et al. (2014) has similar accretion properties as full disks. However, no transitional disks in this sample have been identified as strong accretors. A possible difference between our analysis and the analysis by Najita et al. (2015) may be the prevalence of strong accretors in their sample. Again, this result is based on a small and biased sample of transitional disks and might not be valid for the whole class of objects.

5.2. Empirical differences between Chamaeleon I and Lupus

The results shown in Sects. 4.1 and 4.2 suggest that there are two main differences in the accretion rate properties of Chamaeleon I and Lupus accreting YSOs: the Chamaeleon I sample comprises more strongly accreting YSOs with $L_{\text{acc}} \sim L_{\star}$, and the spread of values of \dot{M}_{acc} at any given M_{\star} in the Chamaeleon I sample is larger than the one found by Alcalá et al. (2014) for objects located in the Lupus clouds and is more compatible with the 0.8 dex found for Taurus members by Herczeg & Hillenbrand (2008). In any case, this spread of \dot{M}_{acc} is still much smaller than what was found in previous studies (e.g., Muzerolle et al. 2003; Mohanty et al. 2005; Natta et al. 2006; Manara et al. 2012; Ercolano et al. 2014).

Before discussing the possible reasons for the difference in the accretion properties of objects in the two largest X-Shooter samples, it is important to remark that the Chamaeleon I and the Lupus X-Shooter samples both include only ~40–50% of the targets surrounded by an optically thick disk (i.e., Class II or transitional disks) in each of these regions. Therefore, selection effects are present in both samples, and some differences might be ascribed to this difference in the intrinsic stellar

properties of the targets. In particular, the distribution of M_{\star} in the two samples is different: more than 70% of the Lupus objects have $M_{\star} \lesssim 0.5 M_{\odot}$, while 60% of the Chamaeleon I targets have $M_{\star} > 0.5 M_{\odot}$. It is possible that the $\dot{M}_{\text{acc}}-M_{\star}$ relation is tighter at lower M_{\star} , as recent results in low-mass stars and brown dwarfs (BDs) in ρ -Ophiuchus suggest (Manara et al. 2015), and this would explain why there is a smaller spread of \dot{M}_{acc} in the Lupus sample. This means that it might be more plausible to have YSOs accreting with $L_{\text{acc}} \sim L_{\star}$ at higher masses, which would explain why more strongly accreting YSOs are present in the Chamaeleon I sample. Both hypotheses can be tested only with more complete samples in various star-forming regions. The selection criteria of the two samples might likewise bias the results. The Lupus targets were selected mostly to be low-mass stars with low or negligible extinction, while the Chamaeleon I targets sampled mostly objects with higher masses and bright, thus it might have been easier to include strongly accreting targets in the latter.

In the following, we discuss some possible reasons for this difference in the context of accretion and disk evolution theories.

5.3. Constraints on accretion and disk evolution theories

The slope of the $\dot{M}_{\text{acc}}-M_{\star}$ relation is a proxy for determining the main driver of the dispersal of protoplanetary disks (Clarke & Pringle 2006). The facts that the slope is significantly higher than 1 and that there is no evidence of an upper locus of objects with slope unity in the $\dot{M}_{\text{acc}}-M_{\star}$ plane disagrees with the expectations formulated by Hartmann et al. (2006). However, this locus of strongly accreting objects might not be observed in the samples analyzed here because the samples are incomplete. In the context of disk dispersal by photoevaporation, the slopes derived for both the Chamaeleon I and Lupus samples are more compatible with X-ray photoevaporation models (Ercolano et al. 2014) than EUV photoevaporation models (Clarke & Pringle 2006). If the larger spread of \dot{M}_{acc} values in Chamaeleon I with respect to Lupus is driven by the different M_{\star} range of the two samples, this spread is indeed larger for higher mass stars than for BDs; this could be explained with the models by Alexander & Armitage (2006). They explain the $\dot{M}_{\text{acc}}-M_{\star}$ relation as a consequence of initial conditions at formation of the disk. To reproduce the observed slope ~ 2 , they suggest that $t_{\nu} \propto M_{\star}^{-1}$, thus that higher mass stars evolve faster. However, observing this effect in a single star-forming region implies the presence of significant age spreads within that region and that the individual stellar ages may be reliably measured. However, ages of individual young stars are unreliable at present (Soderblom et al. 2014).

In the context of viscous evolution theory (e.g., Hartmann et al. 1998), a difference in age could explain the differences in the number of strong accretors and in the spread of \dot{M}_{acc} values between Chamaeleon I and Lupus. The difference in age between the Lupus and Chamaeleon I regions is probably smaller than 1–2 Myr, if there is any difference at all (e.g., Luhman et al. 2008; Alcalá et al. 2014). Viscous evolution models (e.g., Hartmann et al. 1998) predict a typical decrease of \dot{M}_{acc} by ~ 0.3 dex between 2 and 3 Myr for a typical low-mass YSO. This value is higher than the differences in the intercepts of the $\dot{M}_{\text{acc}}-M_{\star}$ relations in the two samples and might explain some observed differences. However, the uncertainty on the age estimates and the possibility of non-coevolution of the targets in each region lead us to consider this effect as a minor, if not negligible, explanation of the observed differences.

As variability is often referred to as a possible explanation for the observed spread of \dot{M}_{acc} values, it is worth asking if

it can play a role in explaining the differences between Lupus and Chamaeleon I discussed here. Costigan et al. (2012) have shown that the typical variability of \dot{M}_{acc} in Chamaeleon I targets is generally $\lesssim 0.4$ dex. Independent studies in other star-forming regions also confirm this finding (e.g., Venuti et al. 2014). Therefore, this effect can be relevant to explain some, even most, of the observed spread of values. If the main factor that explains the observed spread is variability, then this would suggest that the typical rate of accretion variability is lower for objects located in the Lupus cloud than for those in the Chamaeleon I region.

We also consider whether the observed differences are due to the different environment in which the targets discussed here are located. Padoan et al. (2005) suggested that the observed \dot{M}_{acc} are a consequence of Bondy-Hoyle accretion of material surrounding the YSO onto the central object. Such a description would imply a strong dependence of the values of \dot{M}_{acc} on the amount of pristine material in the region. However, the mean value of \dot{M}_{acc} is very similar in the Chamaeleon I and Lupus regions, the largest difference is the scatter of the values. Therefore, this hypothesis could be tested by checking the amount of gas present in the immediate vicinity of individual YSOs and not at the level of the whole cloud. Dullemond et al. (2006) also suggested that the spread of values in the $\dot{M}_{\text{acc}}-M_{\star}$ relation is related to different initial conditions at the formation of disks, in particular different initial core rotation rates. If this is the case, then it might be possible that there has been a larger spread of core rotation rates in Chamaeleon I than in Lupus, or, on the other hand, that higher mass stars have a larger spread of this value at formation.

Finally, it is also possible that the YSO population of Lupus is different from the populations in other regions. For example, Galli et al. (2015) have found that the lifetimes of disks may be shorter in this region than the lifetimes for objects located in the Taurus star-forming region. Although this result relies on individual age estimates, it suggests that objects located in different environments might follow different evolutionary paths.

6. Conclusions

We have presented the analysis of 34 X-Shooter spectra of YSOs in the Chamaeleon I star-forming region, together with four targets in Taurus and two in Chamaeleon II. We have derived their stellar and accretion parameters self-consistently from the same simultaneous spectra ranging from the UV excess, a direct accretion tracer, to the region at ~ 700 nm, which is full of molecular features that are proxies of SpT of the targets. We distinguished between clearly accreting YSOs and those that do not fulfill the criteria to be considered accretors.

The dependence of accretion on the stellar parameters was tested with the logarithmic $L_{\text{acc}}-L_{\star}$ and the $\dot{M}_{\text{acc}}-M_{\star}$ relations, which both show a very good correlation when considering only the accreting targets. The slopes of these relations are 1.72 ± 0.18 and 1.83 ± 0.35 , respectively. These values are compatible with those found in other star-forming regions, such as Lupus. More evolved objects, that is, transitional disks, also located in the Chamaeleon I region follow the locus of full disks in these two relations very nicely as well.

The largest discrepancy between our results and previous results with similar methods in Lupus is the larger spread of the $\dot{M}_{\text{acc}}-M_{\star}$ relation around the best fit, and, in particular, the larger amount of strongly accreting YSOs we found. These differences might suggest that there are intrinsic differences in the evolution of protoplanetary disks in these two regions, but they might

also be ascribed to the fact that these two samples cover only $\sim 50\%$ of the accreting YSOs in these regions and span a different range of stellar masses. The Chamaeleon I sample analyzed here is mostly composed of YSO with masses closer to solar mass, while the Lupus sample comprises many lower mass stars. We might thus see a different behavior in a different stellar mass range.

Our results suggest that the accretion properties of accreting YSOs located in the Chamaeleon I and in the Lupus regions might be different. This difference will be tested when complete X-Shooter samples in both star-forming regions, and possibly in more regions, will be available. At the same time, we await the results from current ALMA surveys of Chamaeleon I and Lupus that will better determine the disk properties of these targets, as this could help us in understanding the differences between the measured accretion rates for objects with similar masses. Finally, with the data release of *Gaia* we will have the possibility to refine the precision of our measurements with the newly determined distances to individual objects in these nearby star-forming regions.

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References

- Alcalá, J. M., Spezzi, L., Chapman, N., et al. 2008, *ApJ*, 676, 427
 Alcalá, J. M., Natta, A., Manara, C. F., et al. 2014, *A&A*, 561, A2
 Alexander, R. D., & Armitage, P. J. 2006, *ApJ*, 639, L83
 Alexander, R., Pascucci, I., Andrews, S., Armitage, P., & Cieza, L. 2014, *Protostars and Planets VI*, 475
 Allard, F., Homeier, D., & Freytag, B. 2011, *ASP Conf. Ser.*, 448, 91
 Antonucci, S., García López, R., Nisini, B., et al. 2011, *A&A*, 534, A32
 Antonucci, S., García López, R., Nisini, B., et al. 2014, *A&A*, 572, A62
 Baraffe, I., Chabrier, G., Allard, F., & Hauschildt, P. H. 1998, *A&A*, 337, 403
 Biazzo, K., Alcalá, J. M., Frasca, A., et al. 2014, *A&A*, 572, A84
 Calvet, N., & Gullbring, E. 1998, *ApJ*, 509, 802
 Cardelli, J. A., Clayton, G. C., & Mathis, J. S. 1989, *ApJ*, 345, 245
 Clarke, C. J., & Pringle, J. E. 2006, *MNRAS*, 370, L10
 Costigan, G., Scholz, A., Stelzer, B., et al. 2012, *MNRAS*, 427, 1344
 Covey, K. R., Ivezić, Ž., Schlegel, D., et al. 2007, *AJ*, 134, 2398
 Daemgen, S., Petr-Gotzens, M. G., Correia, S., et al. 2013, *A&A*, 554, A43
 Dullemond, C. P., Natta, A., & Testi, L. 2006, *ApJ*, 645, L69
 Ercolano, B., Mayr, D., Owen, J. E., Rosotti, G., & Manara, C. F. 2014, *MNRAS*, 439, 256
 Fang, M., van Boekel, R., Wang, W., et al. 2009, *A&A*, 504, 461
 Feigelson, E. D., & Nelson, P. I. 1985, *ApJ*, 293, 192
 Frasca, A., Biazzo, K., Lanzafame, A. C., et al. 2015, *A&A*, 575, A4
 Furlan, E., Luhman, K. L., Espaillat, C., et al. 2011, *ApJS*, 195, 3
 Galli, P. A. B., Bertout, C., Teixeira, R., & Ducourant, C. 2015, *A&A*, 580, A26
 Ghez, A. M., McCarthy, D. W., Patience, J. L., & Beck, T. L. 1997, *ApJ*, 481, 378
 Gullbring, E., Hartmann, L., Briceño, C., & Calvet, N. 1998, *ApJ*, 492, 323
 Gullbring, E., Calvet, N., Muzerolle, J., & Hartmann, L. 2000, *ApJ*, 544, 927
 Ingleby, L., Calvet, N., Herczeg, G., et al. 2013, *ApJ*, 767, 112
 Ingleby, L., Calvet, N., Hernández, J., et al. 2014, *ApJ*, 790, 47
 Kenyon, S. J., & Hartmann, L. 1995, *ApJS*, 101, 117
 Kraus, A. L., & Hillenbrand, L. A. 2007, *ApJ*, 662, 413
 Hartmann, L., Calvet, N., Gullbring, E., & D'Alessio, P. 1998, *ApJ*, 495, 385
 Hartmann, L., D'Alessio, P., Calvet, N., & Muzerolle, J. 2006, *ApJ*, 648, 484
 Herczeg, G. J., & Hillenbrand, L. A. 2008, *ApJ*, 681, 594
 Herczeg, G. J., & Hillenbrand, L. A. 2014, *ApJ*, 786, 97
 Herczeg, G. J., & Hillenbrand, L. A. 2015, *ApJ*, 808, 23
 Jeffries, R. D., Oliveira, J. M., Naylor, T., Mayne, N. J., & Littlefair, S. P. 2007, *MNRAS*, 376, 580

- Lopez Martí, B., Jimenez Esteban, F., Bayo, A., et al. 2013, *A&A*, **551**, A46
- Luhman, K. L. 2004, *ApJ*, **602**, 816
- Luhman, K. L. 2007, *ApJS*, **173**, 104
- Luhman, K. L. 2008, *Handbook of Star Forming Regions, Vol. II*, 169
- Luhman, K. L., Stauffer, J. R., Muench, A. A., et al. 2003, *ApJ*, **593**, 1093
- Luhman, K. L., Allen, L. E., Allen, P. R., et al. 2008, *ApJ*, **675**, 1375
- Luhman, K. L., Allen, P. R., Espaillat, C., Hartmann, L., & Calvet, N. 2010, *ApJS*, **186**, 111
- Manara, C. F., Robberto, M., Da Rio, N., et al. 2012, *ApJ*, **755**, 154
- Manara, C. F., Testi, L., Rigliaco, E., et al. 2013a, *A&A*, **551**, A107
- Manara, C. F., Beccari, G., Da Rio, N., et al. 2013b, *A&A*, **558**, A114
- Manara, C. F., Testi, L., Natta, A., et al. 2014, *A&A*, **568**, A18
- Manara, C. F., Testi, L., Natta, A., & Alcalá, J. M. 2015, *A&A*, **579**, A66
- Manoj, P., Kim, K. H., Furlan, E., et al. 2011, *ApJS*, **193**, 11
- Modigliani, A., Goldoni, P., Royer, F., et al. 2010, *Proc. SPIE*, **7737**
- Mohanty, S., Jayawardhana, R., & Basri, G. 2005, *ApJ*, **626**, 498
- Muzerolle, J., Hartmann, L., & Calvet, N. 1998, *AJ*, **116**, 2965
- Muzerolle, J., Hillenbrand, L., Calvet, N., Briceño, C., & Hartmann, L. 2003, *ApJ*, **592**, 266
- Najita, J. R., Andrews, S. M., & Muzerolle, J. 2015, *MNRAS*, **450**, 3559
- Natta, A., Testi, L., & Randich, S. 2006, *A&A*, **452**, 245
- Natta, A., Testi, L., Alcalá, J. M., et al. 2014, *A&A*, **569**, A5
- Nguyen, D. C., Brandeker, A., van Kerkwijk, M. H., & Jayawardhana, R. 2012, *ApJ*, **745**, 119
- Padoan, P., Kritsuk, A., Norman, M. L., & Nordlund, Å. 2005, *ApJ*, **622**, L61
- Patat, F., Moehler, S., O'Brien, K., et al. 2011, *A&A*, **527**, A91
- Riddick, F. C., Roche, P. F., & Lucas, P. W. 2007, *MNRAS*, **381**, 1067
- Rigliaco, E., Natta, A., Randich, S., Testi, L., & Biazzo, K. 2011a, *A&A*, **525**, A47
- Rigliaco, E., Natta, A., Randich, S., et al. 2011b, *A&A*, **526**, L6
- Rigliaco, E., Natta, A., Testi, L., et al. 2012, *A&A*, **548**, A56
- Robberto, M., Spina, L., Da Rio, N., et al. 2012, *AJ*, **144**, 83
- Schmidt, T. O. B., Vogt, N., Neuhäuser, R., Bedalov, A., & Roell, T. 2013, *A&A*, **557**, A80
- Sicilia-Aguilar, A., Henning, T., & Hartmann, L. W. 2010, *ApJ*, **710**, 597
- Siess, L., Dufour, E., & Forestini, M. 2000, *A&A*, **358**, 593
- Soderblom, D. R., Hillenbrand, L. A., Jeffries, R. D., Mamajek, E. E., & Naylor, T. 2014, *Protostars and Planets VI*, 219
- Spezzi, L., Alcalá, J. M., Covino, E., et al. 2008, *ApJ*, **680**, 1295
- Stelzer, B., Alcalá, J. M., Scholz, A., et al. 2013, *A&A*, **551**, A106
- Valenti, J. A., Basri, G., & Johns, C. M. 1993, *AJ*, **106**, 2024
- Venuti, L., Bouvier, J., Flaccomio, E., et al. 2014, *A&A*, **570**, A82
- Vernet, J., Dekker, H., D'Odorico, S., et al. 2011, *A&A*, **536**, A105
- White, R. J., & Basri, G. 2003, *ApJ*, **582**, 1109
- Winston, E., Cox, N. L. J., Prusti, T., et al. 2012, *A&A*, **545**, A145

Appendix A: Additional data

Additional information from the literature on the targets discussed here is reported in Tables A.1–A.2.

Table A.1. Data available in the literature for the targets in Chamaeleon I included in this work.

Object/other name	RA(2000) h:m:s	Dec(2000) ° ' "	SpT	A_J [mag]	Type	Notes	References
T3 / SX Cha	10:55:59.7	-77:24:39.91	M0/M1.5	0.79	II	Binary (2.21")	1, 6, 5
T3 B / T3 E	10:56:00.0	-77:24:40.6	M3.5/M3	0.00	...	Binary (2.21")	1, 5
T4 / SY Cha / Sz 3	10:56:30.4	-77:11:39.31	M0.5	0.23	II	Accreting	1, 10, 3
TW Cha / T7	10:59:01.1	-77:22:40.71	K8	0.34	II	Accreting, binary(spectroscopic?)	1, 6, 3, 9
CR Cha / T8 / Sz 6	10:59:06.9	-77:01:40.31	K2	0.00	II	...	1, 6
T12 / Sz10 / ISO–ChaI 10	11:02:55.0	-77:21:50.81	M4.5	0.32	II	Accreting	1, 2, 3
CT Cha A / T14 / Sz 11	11:04:09.1	-76:27:19.38	K5	0.45	II	Accreting, binary (2.67")	1, 2, 3, 5
ISO–ChaI 52	11:04:42.6	-77:41:57.13	M4	0.36	II	Accreting	1, 2, 3
Hn 5 / GX Cha	11:06:41.8	-76:35:48.98	M4.5	0.32	II	Accreting	1, 2, 3
T23	11:06:59.1	-77:18:53.57	M4.25	0.00	II	...	1, 2
T24 / UZ Cha / Sz 17	11:07:12.1	-76:32:23.23	M0.5	0.63	II	Accreting	1, 2, 3
Cha–H α –1 / ISO–ChaI 95	11:07:16.7	-77:35:53.29	M7.75	0.00	II	Accreting	1, 2, 3
Cha–H α –9 / ISO–ChaI 98	11:07:18.6	-77:32:51.66	M5.5	0.45	II	Accreting	1, 2, 3
Sz 22 / T29 / FK Cha / KG 28	11:07:57.9	-77:38:44.93	K6	1.24	Flat	Accreting, triple (17.6")	1, 2, 8, 11
VW Cha / T31 / ISO–ChaI 123	11:08:01.5	-77:42:28.85	K8	0.72	II	Accreting, binary(0.66")	1, 2, 3, 4, 5
ESO H α 562 / HO Cha / ISO–ChaI 126	11:08:02.9	-77:38:42.59	M1.25	1.35	II	Accreting, binary (0.28")	1, 2, 3, 5
T33 A / HP Cha / ISO–ChaI 135	11:08:15.1	-77:33:53.16	G7	0.85	Flat	Accreting, triple (2.4")	1, 2, 3, 4, 5
T33 B	K6/M0.5	0.52	...	Triple (2.4"), Binary (0.11")	1, 5
ISO–ChaI 143	11:08:22.4	-77:30:27.71	M5	1.13	II	Accreting, binary (18.16")	1, 2, 3, 4, 7
Cha–H α –6 / HQ Cha	11:08:39.5	-77:34:16.67	M5.75	0.38	II	Accreting	1, 2, 3, 4
T38 / VY Cha / Sz 29	11:08:54.6	-77:02:12.96	M0.5	0.90	II	Accreting	1, 2, 3
T40 / VZ Cha	11:09:23.8	-76:23:20.76	K6	0.56	II	...	1, 2
T44 / WW Cha / ISO–ChaI 231	11:10:00.1	-76:34:57.89	K5	1.35	II	Accreting	1, 2, 3
T45A / FN Cha A	11:10:04.7	-76:35:45.22	M1	0.54	II	Binary (28.32"), ISO–ChaI 237)	1, 2, 7
ISO–ChaI 237	11:10:11.4	-76:35:29.29	K5.5	1.92	II	Accreting, binary (28.32", T45A)	1, 2, 3, 7
T49 / XX Cha	11:11:39.6	-76:20:15.25	M2	0.34	II	Accreting, binary (24.38")	1, 2, 3, 7
CHX 18N	11:11:46.3	-76:20:09.21	K6	0.20	II	Accreting	1, 2, 3
T51 / CHX 20 / Sz 41	11:12:24.4	-76:37:06.41	K3.5/K7	0.00	II	Binary (1.97")	1, 2, 5
T51 B	M2.5	1.09	...	Binary (1.97")	5
T52 / CV Cha	11:12:27.7	-76:44:22.30	G9	0.42	II	Binary (11.55")	1, 2, 8
T54–A / CHX 22	11:12:42.7	-77:22:23.05	G8	0.60	II	Binary (0.24")	1, 2, 5
Hn17	11:12:48.6	-76:47:06.67	M4	0.09	II	Accreting	1, 2, 3
Hn18	11:13:24.4	-76:29:22.73	M3.5	0.14	II	Accreting	1, 2, 3
Hn21W	11:14:24.5	-77:33:06.22	M4	0.72	II	Accreting, binary (5.48")	1, 2, 3, 5

Notes. Spectral types, extinction, disk classification, accretion indication, and binarity are adopted from the following studies: 1. Luhman (2007); 2. Luhman et al. (2008); 3. Luhman (2004); 4. Costigan et al. (2012); 5. Daemgen et al. (2013); 6. Manoj et al. (2011); 7. Kraus & Hillenbrand (2007); 8. Ghez et al. (1997); 9. Nguyen et al. (2012); 10. Winston et al. (2012); 11. Schmidt et al. (2013). Complete name on SIMBAD for T# is Ass Cha T 2–#.

Table A.2. Data available in the literature for the targets in Taurus and Chamaeleon II included in this work.

Object/other name	RA(2000) h:m:s	Dec(2000) ° ' "	SpT	A_V [mag]	Type	Notes	References
Taurus ($d = 140$ pc)							
FN Tau	04:14:14.590	+28:27:58.06	M4.5	0.35	II	$d = 131$ pc for HH14	1, 2
V409 tau	04:18:10.785	+25:19:57.39	M0.6	1.00	...	$d = 131$ pc for HH14	1
IQ Tau	04:29:51.563	+26:06:44.90	M1.1	0.85	II	...	1, 2
DG Tau	04:27:04.698	+26:06:16.31	K7	1.60	II	...	1, 2
Chamaeleon II ($d = 178$ pc)							
Hn24	13:04:55.75	-77:39:49.5	M0	2.76	II	TD?	3, 4
Sz50	13:00:55.32	-77:10:22.2	M3	3.78	II	...	3, 4

Notes. Spectral types, extinction, disk classification, accretion indication, and binarity are from the following studies: 1. Herczeg & Hillenbrand (2014); 2. Furlan et al. (2011); 3. Spezzi et al. (2008); 4. Alcalá et al. (2008).

Appendix B: Comments on individual targets

As discussed in Sect. 3.1, stellar and accretion properties of the targets discussed here were determined following the procedure described by Manara et al. (2013b). In a few cases, however, the quality of the spectra was not good enough for this procedure to work, or the properties of the target did not lead to satisfyingly good fits. In these cases slightly different procedures were adopted. Additional points were included in the fit of some targets, in particular at ~ 400 nm and between ~ 500 nm and 600 nm. In two cases, namely when fitting Cha H α 1 and Cha H α 9, the procedure was totally different. These cases are described in the following.

B.1. Chamaeleon I targets

T3: this target has a best fit with templates with SpT K7. However, the derived L_\star is low and results in a position on the HRD slightly below the 30 Myr isochrone according to the Baraffe et al. (1998) models. Even including additional points in the fit did not change the results. However, the fit is good, the SpT is compatible with that from the literature (M0) and that from the TiO index by Jeffries et al. (2007) (K6) and is slightly later than the result with the indices by HH14 (K4.2). Similarly, the A_V estimated here agrees well with previous values in the literature. When using the bolometric correction by HH14 to determine L_\star , this is only ~ 1.2 times higher than the one derived from the best-fit template. This would make the object only slightly younger, thus still older than 20 Myr.

CR Cha: this object was best fit with a K0 template. Previous works reported it as an object with SpT K2, but the fit with a K2 template is not good. The value of SpT derived using the indices by HH14 is K1, but no templates with this SpT are available. Given that the position of this target on the HRD is outside the range of the evolutionary models by Baraffe et al. (1998), the stellar parameters were derived using evolutionary models by Siess et al. (2000).

CT Cha A: the best fit of this target was obtained by including additional points at ~ 400 nm and between ~ 500 nm and 600 nm. If these additional points were not included, the best fit would have been with a K2 template, but it would have been worse and leading to a significantly older age. However, the other stellar and accretion parameters would be compatible within the uncertainties. When comparing L_\star determined from the best fit with that obtained using the bolometric correction by HH14, the latter would lead to a value higher by a factor ~ 1.8 . The reason for this discrepancy is possibly the high veiling in this object due to very intense accretion.

Cha H α 1: none of the photospheric templates available in this study has the same SpT as this target. Indeed, literature estimates and spectral indices lead to an M7-M8 SpT. This object is also known to have negligible extinction (Luhman 2007), and it seems not to accrete at a high rate from the emission lines present in the spectrum. Therefore, it is possible to adopt the SpT from the spectral indices for this target, assume $A_V = 0$ mag, determine L_\star using the bolometric correction by HH14, and L_{acc} from the luminosity of the emission lines present in the spectrum and the relations by Alcalá et al. (2014) to convert these luminosity in L_{acc} . This was the procedure adopted here to derive the stellar parameters reported in Table 1 and used in this work.

Cha H α 9: the spectrum of this object has almost no flux in the whole UVB arm, thus the procedure adopted in this work for the other targets was not applicable. This target is probably accreting at a very low rate according to its H α line. Indeed, the

$EW_{\text{H}\alpha}$ is 20 Å, while the 10% width of the line is ~ 200 km s $^{-1}$. The stellar parameters were thus derived in the following way. First, the SpT was determined using several spectral indices, which led to an estimate of SpT M5.5. Then, the VIS part of the spectrum was compared to a photospheric template with SpT M5.5 that was reddened with increasing values of A_V in steps of 0.1 mag until a best match was found. The best agreement between the two spectra is with $A_V = 4.8$ mag, which is higher than that reported in the literature. Finally, L_\star was determined using the bolometric correction by HH14 and L_{acc} from the only emission line detected with high S/N in the spectrum, which is the H α line. As the width of this line is quite small, it should be noted that the L_{acc} derived here might be an upper limit of the real L_{acc} , as this object might be non-accreting.

Sz22: this object is a very complex multiple system. Only the close binary at the center (Schmidt et al. 2013) is included in the slit. This object is a confirmed member of Chamaeleon I by Lopez Martí et al. (2013) from proper motion studies. The best fit was obtained by including all the additional points listed above and with a value of A_V much lower than the one reported in the literature (Luhman 2007; Schmidt et al. 2013). No signs of the companion were found in the overall shape of the spectrum. The derived age is older than 10 Myr, but still compatible with younger ages according to other evolutionary models.

ESO H α 562: the spectrum of this target has a low S/N (~ 1) in the UVB arm until ~ 400 nm. It is reported to be a close binary (0.28'' separation) of two objects with the same SpT (M0.5, Daemgen et al. 2013). Therefore, both objects have probably been included in the slit and are not resolved. The best fit was obtained with SpT M1 and a $L_\star = 0.11 L_\odot$. The same L_\star is obtained when using the bolometric correction by HH14. This places the object low on the HRD, at an age of ~ 24 Myr. The $A_V = 3.4$ mag derived here is smaller than the $A_V > 4$ mag reported by Daemgen et al. (2013), while the L_\star determined here is the same as that reported by Daemgen et al. (2013) for component B. It is thus possible that in the VIS and UVB arm the spectrum of component B dominates, as Daemgen et al. (2013) reported that component A has $A_V = 10.5$ mag and it is as bright as component B in the near-infrared. The latter has only $A_V = 4.3$ mag, according to Daemgen et al. (2013), thus at optical wavelengths component B is significantly brighter than component A.

T33-B: the best fit was determined with a template with SpT K0, but this led to a low L_\star implying an age ~ 40 Myr according to the evolutionary models by Baraffe et al. (1998). A similar L_\star was determined when using the bolometric correction by HH14. It should be noted that this component of the system T33 is composed of two objects separated by 0.11'' (Daemgen et al. 2013), thus not resolved by our observations. From near-infrared spectra, Daemgen et al. (2013) reported a SpT M0.5 for these targets. However, both the fit and the spectral indices lead to a much earlier SpT and to an A_V similar to that of component A.

Cha H α 6: the best fit for this target was with the photospheric template with SpT M6.5, but this was obtained from a spectrum with a low S/N in the UVB. However, the spectral indices also confirm that the SpT of this target is around M6. The value of L_\star determined from the fit places the object immediately on the 1 Myr isochrone of the evolutionary models by Baraffe et al. (1998), while the L_\star from the bolometric correction by HH14 leads to a smaller L_\star . As reported earlier on, this difference might be ascribed to the SpT- T_{eff} relation used by HH14, which differs from that by Luhman et al. (2003), which was used here, at this late stellar type.

T44: this object is a very strong accretor almost totally veiled, and it is known to launch a jet (Robberto et al. 2012). The results of the fit are not perfect, for example, the slope of the Balmer continuum is not well reproduced. However, even including additional points in the fit the accretion parameters change within the uncertainties, while the stellar parameters, in particular the SpT, might change.

T45a: this is a non-accreting target, but the profile of the $H\alpha$ line is unusual for a non-accreting object. It seems rather narrow ($\sim 150\text{--}170\text{ km s}^{-1}$) and weak, but has a strong self-absorption at almost zero velocity. The excess in the UV is almost zero, and the \dot{M}_{acc} very low, consistent with being non-accreting.

ISO-Chal 237: the spectrum of this target has a low S/N in the UVB arm, thus the best fit is slightly uncertain, but leads to a SpT compatible with the values from spectral indices and from

the literature. This best fit leads to a low \dot{M}_{acc} in the $\dot{M}_{\text{acc}}\text{--}M_{\star}$ plane, in the region where non-accreting targets are located. It is difficult to determine if this object is non-accreting because its spectrum presents a broad $H\alpha$ line but with strong self-absorption and, thus, a small EW.

T49: this object is one with stronger accretion than most in the sample analyzed here. The spectrum is highly veiled, but a best fit was found and confirmed, also including additional points.

T54-A: this is a non-accreting target, composed of two components separated by $0.24''$, thus not resolved by our observations (Daemgen et al. 2013). Component A should, however, dominate the emission in this wavelength range. The SpT derived here agrees with that from the literature and from spectral indices by HH14.

Appendix C: Best fit

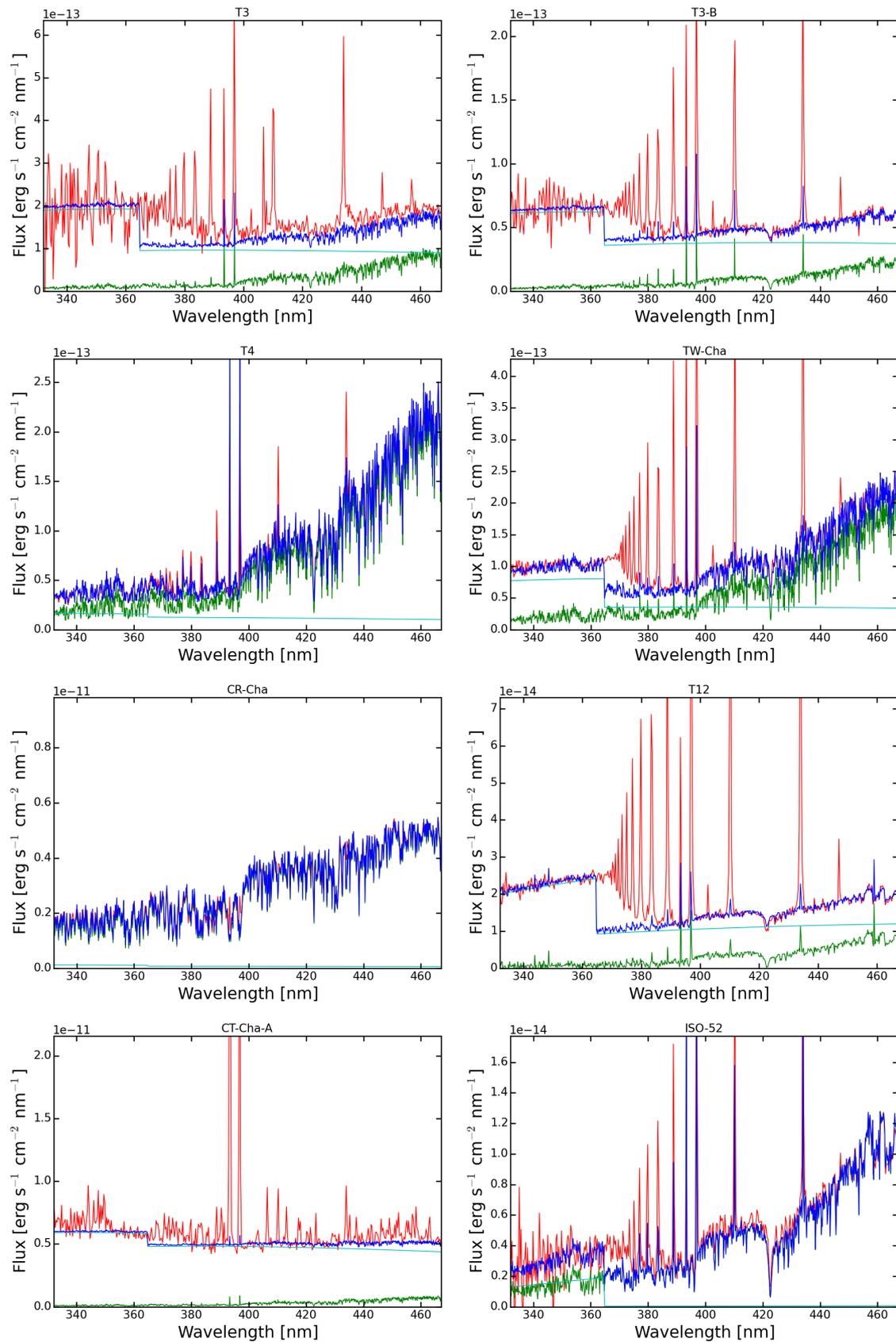


Fig. C.1. Best fit for the Chamaeleon I targets studied here. Names are reported in the title of each subplot. The red line is the reddening-corrected spectrum of the target, the blue line the best fit, which is the sum of the photospheric template (green line) and the slab model (cyan line). The input spectrum and the photospheric templates are smoothed for a better visualization.

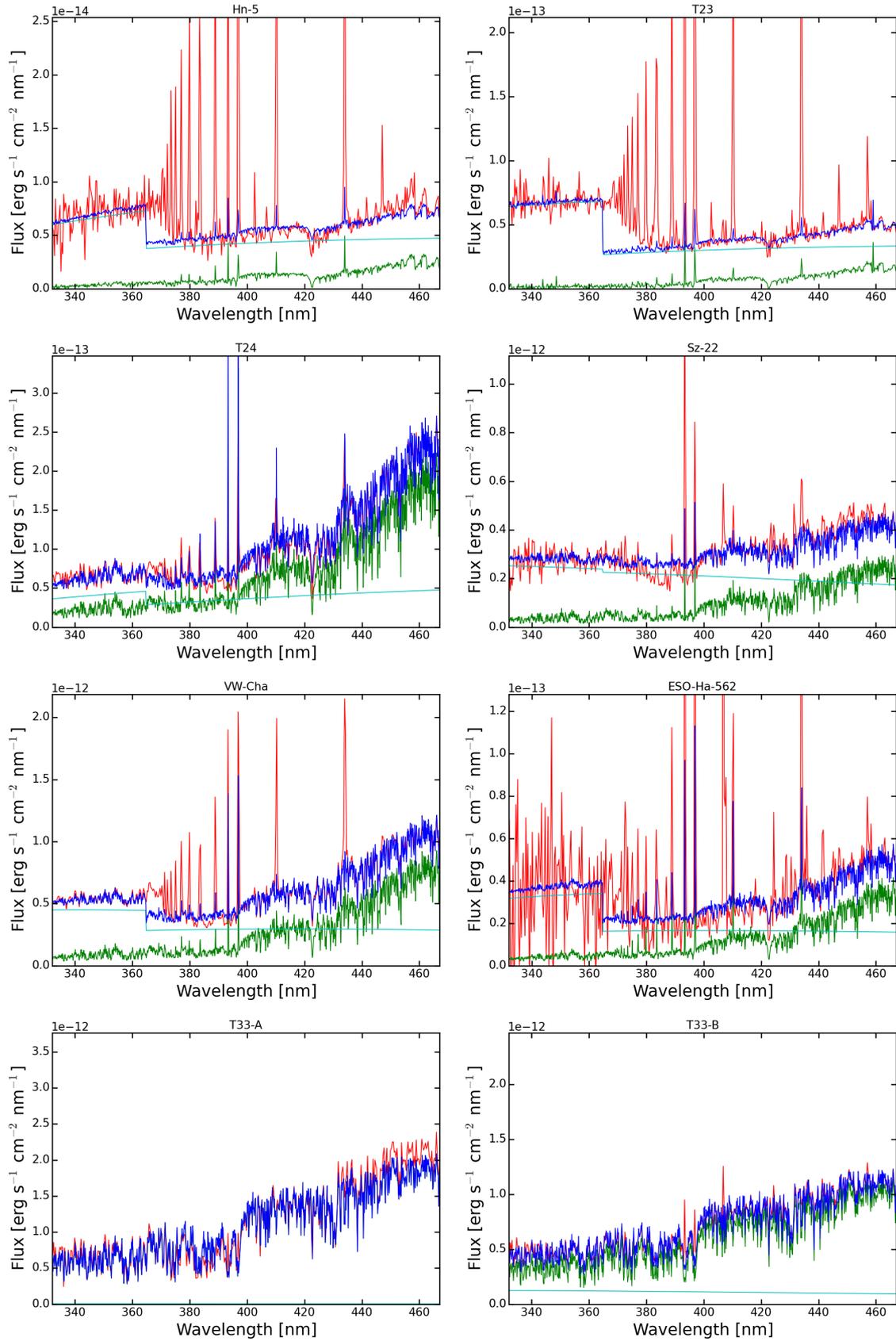


Fig. C.1. continued.

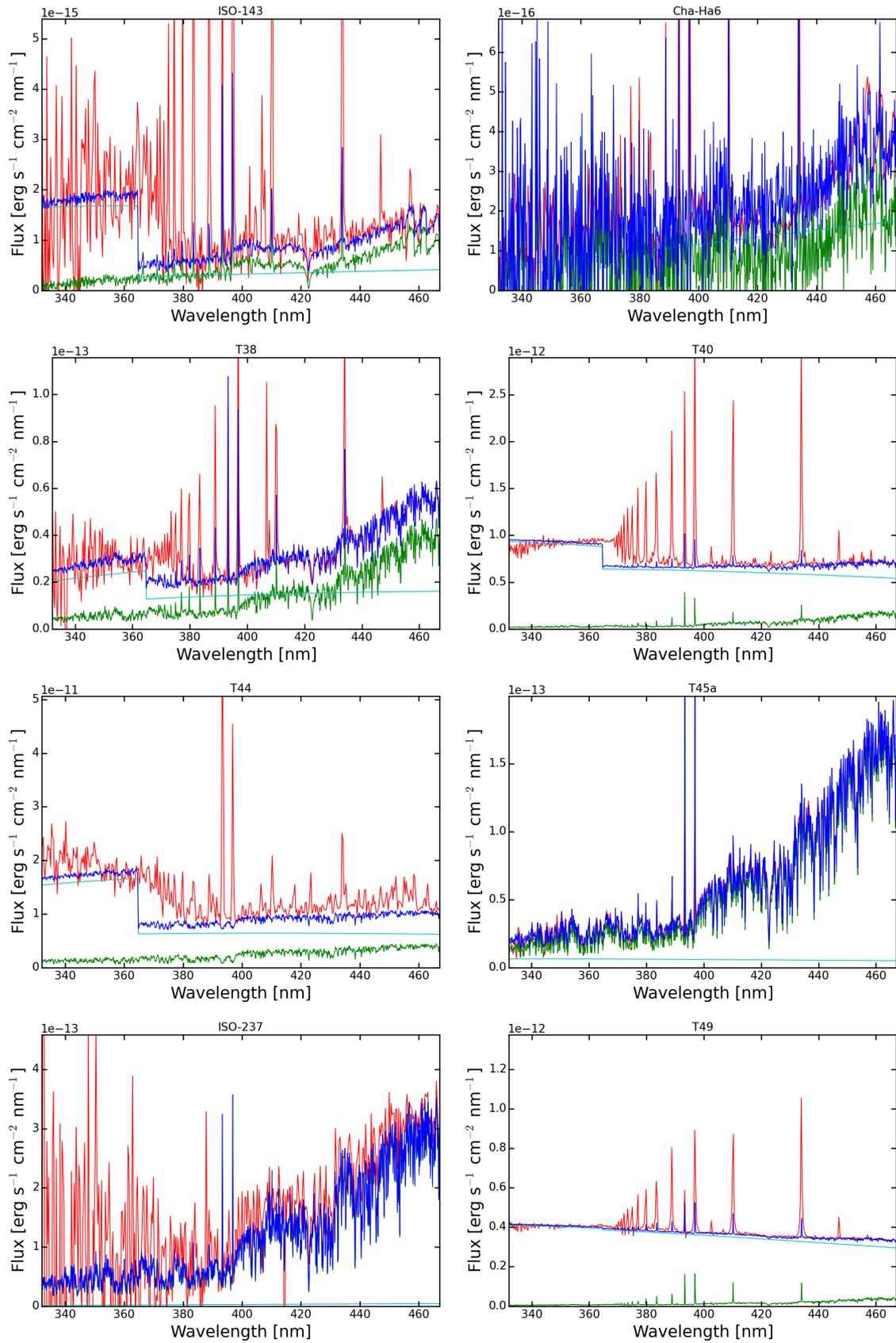


Fig. C.1. continued.

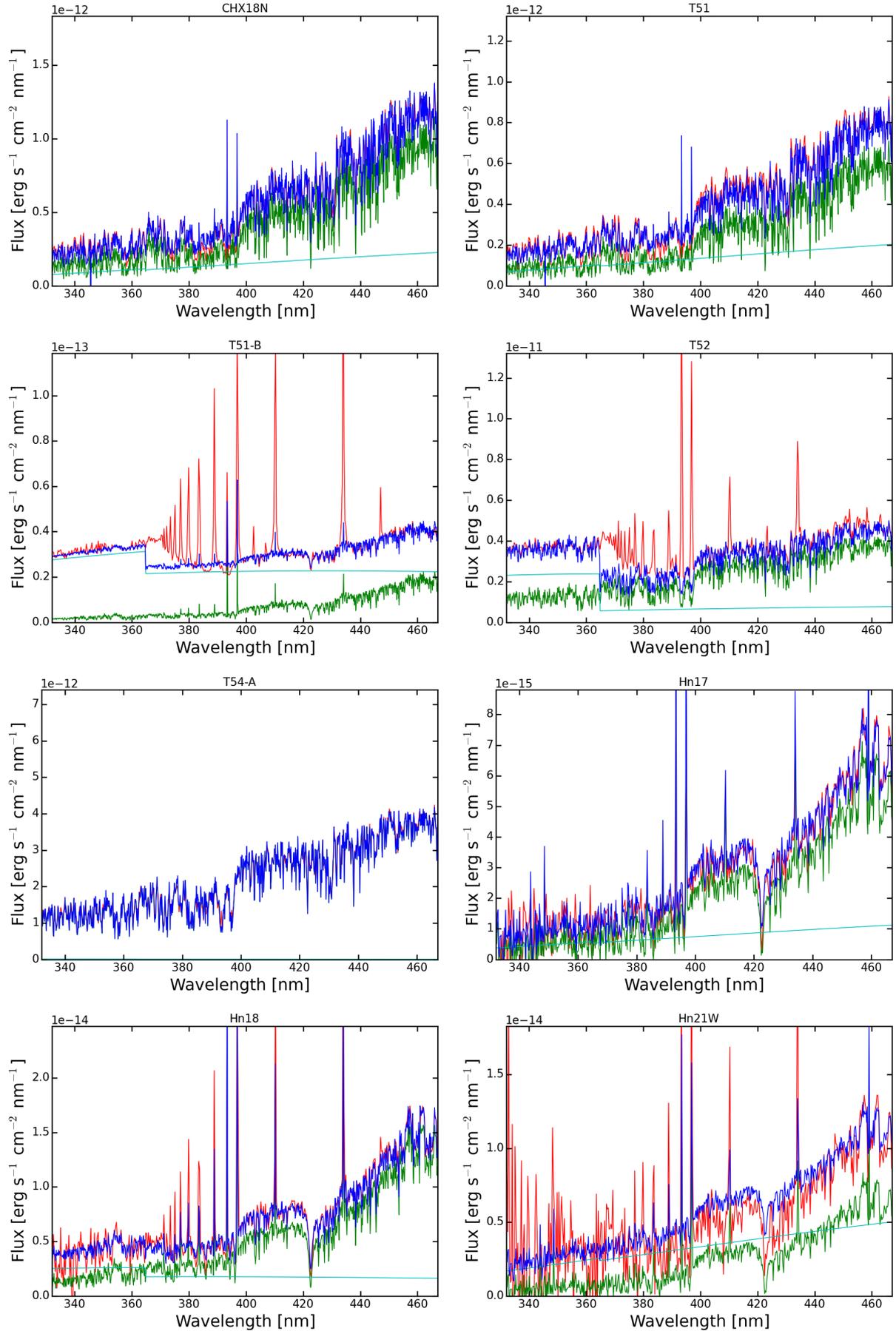


Fig. C.1. continued.

Appendix D: Observation log

Table D.1. Night log.

Star	RA(J2000)	Dec(J2000)	Date [UT]	Exposure Time [NDIT × DIT(s)]			Slit Width ["]		
				UVB	VIS	NIR	UVB	VIS	NIR
FN Tau	04:14:14.70	+28:27:55.1	2010-01-18;01:04:05.765	2 × 50	2 × 38	2 × 10	1.0	0.4	0.4
V409 tau	04:18:10.85	+25:19:55.3	2010-01-18;01:13:22.953	2 × 60	2 × 40	2 × 20	1.0	0.4	0.4
IQ Tau	04:29:51.62	+26:06:43.0	2010-01-18;01:22:07.078	2 × 90	2 × 80	2 × 10	1.0	0.4	0.4
ISO-ChaI 52	11:04:43.33	-77:41:55.2	2010-01-18;04:10:03.932	2 × 360	2 × 350	2 × 360	1.0	0.4	0.4
Cha H α 9	11:07:17.92	-77:32:52.3	2010-01-18;04:37:32.755	2 × 525	2 × 540	2 × 270	1.0	0.4	0.4
Ass Cha T 2 40	11:09:23.13	-76:23:21.0	2010-01-18;05:02:09.994	2 × 45	2 × 40	2 × 20	1.0	0.4	0.4
Ass Cha T 2 23	11:06:58.50	-77:18:53.2	2010-01-18;05:09:38.933	2 × 120	2 × 110	2 × 120	1.0	0.4	0.4
Cha H α 1	11:07:16.21	-77:35:53.3	2010-01-18;05:29:37.917	4 × 420	4 × 410	4 × 210	1.0	0.4	0.4
Hn 5	11:06:41.09	-76:35:47.1	2010-01-18;06:27:00.167	2 × 120	2 × 110	2 × 60	1.0	0.4	0.4
Cha H α 6	11:08:39.05	-77:34:15.3	2010-01-18;06:48:11.388	4 × 540	4 × 550	4 × 270	0.8	0.4	0.4
Ass Cha T 2 12	11:02:54.79	-77:21:49.5	2010-01-18;07:34:15.838	2 × 150	2 × 140	2 × 75	1.0	0.4	0.4
Hn18	11:14:25.36	-77:33:04.6	2010-01-18;08:09:13.040	2 × 530	2 × 540	2 × 270	1.0	0.4	0.4
DG Tau	04:27:04.29	+26:06:08.7	2010-01-19;01:49:03.987	4 × 240	4 × 250	4 × 15	1.3	1.2	1.2
Ass Cha T 2 4	10:55:58.62	-77:24:40.6	2010-01-19;04:43:50.328	2 × 110	2 × 120	2 × 120	1.0	0.4	0.4
ISO-ChaI 237	11:10:10.70	-76:35:29.3	2010-01-19;04:55:06.057	2 × 170	2 × 180	2 × 180	1.0	0.4	0.4
Ass Cha T 2 45a	11:10:03.76	-76:35:43.8	2010-01-19;05:09:28.923	2 × 90	2 × 100	2 × 100	1.0	0.4	0.4
Ass Cha T 2 51	11:12:25.16	-76:37:09.3	2010-01-19;05:39:42.811	2 × 300	2 × 285	2 × 60	1.0	0.4	0.4
Hn17	11:12:52.43	-76:47:06.3	2010-01-19;06:09:01.955	2 × 220	2 × 240	2 × 240	1.0	0.4	0.4
Sz22	11:07:57.19	-77:38:43.5	2010-01-19;06:25:40.422	2 × 210	2 × 210	2 × 30	1.0	0.4	0.4
ESO H α 562	11:08:02.20	-77:38:41.5	2010-01-19;06:37:21.103	2 × 360	2 × 345	2 × 180	1.0	0.4	0.4
Ass Cha T 2 54	11:12:42.35	-77:22:21.2	2010-01-19;07:26:01.448	2 × 170	2 × 180	2 × 90	1.0	0.4	0.4
Ass Cha T 2 49	11:11:39.22	-76:20:13.4	2010-01-19;07:38:58.735	2 × 200	2 × 210	2 × 210	1.0	0.4	0.4
Hn18	13:04:56.23	-77:39:47.4	2010-01-19;07:52:42.296	2 × 240	2 × 250	2 × 250	1.0	0.4	0.4
VW Cha	11:08:00.96	-77:42:27.4	2010-01-19;08:12:33.218	2 × 170	2 × 180	2 × 45	1.0	0.4	0.4
TW Cha	10:59:00.78	-77:22:39.7	2010-01-19;08:46:24.914	2 × 150	2 × 160	2 × 50	1.0	0.4	0.4
Ass Cha T 2 52	11:12:27.67	-76:44:21.0	2010-01-19;09:00:30.337	2 × 90	2 × 100	2 × 20	1.0	0.4	0.4
Hn18	11:13:23.75	-76:29:22.1	2010-01-20;04:37:16.890	2 × 310	2 × 300	2 × 150	1.0	0.4	0.4
Ass Cha T 2 4	10:56:29.59	-77:11:39.2	2010-01-20;04:54:15.929	2 × 120	2 × 110	2 × 60	1.0	0.4	0.4
CT Cha A	11:04:08.14	-76:27:19.0	2010-01-20;05:26:22.036	2 × 120	2 × 110	2 × 60	1.0	0.4	0.4
Ass Cha T 2 24	11:07:11.65	-76:32:22.2	2010-01-20;05:45:38.927	2 × 170	2 × 175	2 × 85	1.0	0.4	0.4
ISO-ChaI 143	11:08:21.60	-77:30:26.9	2010-01-20;06:02:44.956	2 × 800	2 × 780	2 × 270	1.0	0.4	0.4
Ass Cha T 2 33	11:08:14.55	-77:33:54.4	2010-01-20;06:56:21.776	2 × 90	2 × 80	2 × 10	1.0	0.4	0.4
SZ50	13:00:54.86	-77:10:21.3	2010-01-20;08:02:17.861	2 × 170	2 × 160	2 × 20	1.0	0.4	0.4
Ass Cha T 2 38	11:08:54.50	-77:02:10.1	2010-01-20;08:32:19.916	2 × 180	2 × 170	2 × 45	1.0	0.4	0.4
CR Cha	10:59:06.73	-77:01:37.0	2010-01-20;09:02:48.053	2 × 65	2 × 60	2 × 20	1.0	0.4	0.4
CHX18N	11:11:46.50	-76:20:06.4	2010-01-20;09:12:35.963	2 × 60	2 × 50	2 × 5	1.0	0.4	0.4
CHX18N	11:11:46.23	-76:20:06.9	2010-01-20;09:20:42.895	4 × 35	4 × 30	4 × 15	1.0	0.4	0.4
Ass Cha T 2 51	11:10:00.41	-76:34:55.3	2010-01-20;09:31:34.791	4 × 30	4 × 25	4 × 5	1.0	0.4	0.4