

A systematic examination of particle motion in a collapsing magnetic trap model for solar flares

K. J. Grady, T. Neukirch, and P. Giuliani

School of Mathematics and Statistics, University of St. Andrews, St Andrews KY16 9SS, UK
e-mail: keithg@mcs.st-andrews.ac.uk

Received 30 January 2012 / Accepted 1 August 2012

ABSTRACT

Context. It has been suggested that collapsing magnetic traps may contribute to accelerating particles to high energies during solar flares.

Aims. We present a detailed investigation of the energization processes of particles in collapsing magnetic traps, using a specific model. We also compare for the first time the energization processes in a symmetric and an asymmetric trap model.

Methods. Particle orbits are calculated using guiding centre theory. We systematically investigate the dependence of the energization process on initial position, initial energy and initial pitch angle.

Results. We find that in our symmetric trap model particles can gain up to about 50 times their initial energy, but that for most initial conditions the energy gain is more moderate. Particles with an initial position in the weak field region of the collapsing trap and with pitch angles around 90° achieve the highest energy gain, with betatron acceleration of the perpendicular energy the dominant energization mechanism. For particles with smaller initial pitch angle, but still outside the loss cone, we find the possibility of a significant increase in parallel energy. This increase in parallel energy can be attributed to the curvature term in the parallel equation of motion and the associated energy gain happens in the centre of the trap, where the field line curvature has its maximum. We find qualitatively similar results for the asymmetric trap model, but with smaller energy gains and a larger number of particles escaping from the trap.

Key words. Sun: corona – Sun: flares – Sun: activity – Sun: surface magnetism – Sun: X-rays, gamma rays

1. Introduction

One of the most important open questions in solar physics is to identify the mechanisms by which a large number of charged particles are accelerated to high energies during solar flares. While there is general consensus that the energy released in solar flares is previously stored in the magnetic field, the physical processes by which this energy is converted into bulk flow energy, thermal energy, non-thermal energy and radiation energy are still a matter of debate (e.g. Miller et al. 1997; Aschwanden 2002; Neukirch 2005; Neukirch et al. 2007; Krucker et al. 2008; Aschwanden 2009). Based on observations, in particular of non-thermal high-energy (hard X-ray and γ -ray) radiation, it has been estimated that a large fraction of the released magnetic energy (up to the order of 50%) is converted into non-thermal energy in the form of high energy particles (e.g. Emslie et al. 2004, 2005).

Possible particle acceleration mechanisms that have been suggested include direct acceleration in the parallel electric field associated with the reconnection process, stochastic acceleration by turbulence and/or wave-particle resonance, shock acceleration or acceleration in the inductive electric field of the reconfiguring magnetic field (see e.g. Miller et al. 1997; Aschwanden 2002; Neukirch 2005; Neukirch et al. 2007; Krucker et al. 2008; Zharkova et al. 2011, for a detailed discussion and further references). As none of the proposed mechanisms can fully explain the high-energy particle fluxes within the framework of the standard solar flare thick target model, alternative acceleration

scenarios have been proposed (e.g. Fletcher & Hudson 2008; Birn et al. 2009; Brown et al. 2009).

It has been suggested (e.g. Somov 1992; Somov & Kosugi 1997) that the rapid relaxation of magnetic field lines that have been newly reconnected, but outside the actual reconnection region, could contribute to the acceleration of particles. Charged particles could be trapped within the dynamically relaxing magnetic fields, which form a collapsing magnetic trap (CMT). Based on the conservation of adiabatic constants of motion, the kinetic energy of the particles trapped in a CMT could increase due to the betatron effect, as the magnetic field strength in the CMT increases, and due to first-order Fermi acceleration, as the distance between the mirror points of particle orbits decreases due to the shortening of the field lines. Evidence of post-flare field line relaxation (field line shrinkage) has been found in observations by Yokoh (e.g. Forbes & Acton 1996) and Hinode (e.g. Reeves et al. 2008b). Field line relaxation in solar flares can be compared to the dipolarisation phase of magnetospheric substorms, which is believed to play a major role in particle acceleration during substorms (e.g. Birn et al. 1997, 1998, 2004). A more general comparison of flare and substorm/magnetotail phenomena based on observations has recently been presented by Reeves et al. (2008a). Another, albeit slightly different, particle acceleration mechanism, which also relies on the relaxation of magnetic field lines is the shrinkage of magnetic islands (plasmoids) that has been discussed, for example by Drake et al. (2006) and Karlický & Bárta (2007).

Using very basic CMT models, some of the fundamental properties of the particle acceleration process in CMTs have been investigated by Somov and co-workers (e.g. [Bogachev & Somov 2001, 2005, 2009](#); [Kovalev & Somov 2002, 2003a,b](#); [Somov & Bogachev 2003](#)). This includes, for example, the relative efficiencies of betatron and Fermi acceleration, the effect of collisions, the role of velocity anisotropies and the evolution of the energy distribution function in a CMT. [Li & Fleishman \(2009\)](#) used the results of [Bogachev & Somov \(2005\)](#) to determine the gyrosynchrotron radio emission that is to be expected from a collapsing magnetic trap. While these calculations are extremely useful for first estimates, they are based on a number of simplifying assumptions about both the CMTs and the particle orbits. Specific models are necessary for gaining a more detailed understanding of the processes in CMTs, with the disadvantage that some of the results will become model-dependent.

A number of previous papers has followed this line of investigation. Using a simple model for time evolution of the CMT magnetic field strength and a simplified equation of motion for the particles, [Karlický & Kosugi \(2004\)](#) investigated CMT properties such as particle acceleration, plasma heating and the resulting X-ray emission. [Karlický & Bárta \(2006\)](#) used an MHD (magnetohydrodynamic) simulation of a reconnecting current sheet to generate CMT-like electromagnetic fields to investigate particle acceleration using test particle calculations, in particular with a view to explain hard X-ray loop-top sources. Some indication that CMTs might be relevant for X-ray loop top sources has been provided by [Veronig et al. \(2006\)](#). [Aschwanden \(2004\)](#) used a very simple time-dependent trap model to try and explain the pulsed time profile of energetic particle injection often observed during flares. Most recently, [Minoshima et al. \(2010\)](#) presented the results of a calculation using numerical solutions of the drift-kinetic equation for a CMT model based on a time-dependent 2D magnetic field suggested by [Lin et al. \(1995\)](#) to interpret the motion of flare loops and ribbons in the framework of the standard flare model.

[Giuliani et al. \(2005\)](#) presented a general theoretical framework for more detailed analytical CMT models. This theoretical framework is based on general analytic solutions of the kinematic MHD equations, i.e. with given bulk flow profile. The theory was developed by [Giuliani et al. \(2005\)](#) for 2D and 2.5D magnetic fields, but excluding flow in the invariant direction and has recently been extended to fully 3D magnetic fields and flows by [Grady & Neukirch \(2009\)](#).

[Giuliani et al. \(2005\)](#) focus on the development of the theory and present just a few examples of model CMTs, together with a calculation of just a single example of a particle orbit based on non-relativistic guiding centre theory (see e.g. [Northrop 1963](#)). In this paper we present a systematic investigation of test particle orbits for different initial conditions, firstly by using the same symmetric CMT model as in [Giuliani et al. \(2005\)](#) and secondly by using a modified asymmetric CMT model. As in [Giuliani et al. \(2005\)](#) we will use the first order guiding centre theory. We are in particular interested in the dependence of particle energy gain on initial position in the trap, initial energy and initial pitch angle. Another interesting question is whether the energy gain mechanisms predicted using adiabatic invariants can indeed be identified using the guiding centre orbits.

The paper is structured in the following way. In Sect. 2 we summarise the basic theory and the models presented in [Giuliani et al. \(2005\)](#). An overview of the dependence of particle orbits and energy on initial conditions for the symmetric CMT model is given in Sect. 3. We take a more detailed look at two particular

particle orbits for the symmetric CMT model in Sect. 4 to study the energy gain mechanisms in more detail. In Sect. 5 we present results for an asymmetric CMT model. We finish the paper with summary and conclusions in Sect. 6.

2. Basic equations and CMT model

[Giuliani et al. \(2005\)](#) develop their theory using the ideal kinematic MHD equations,

$$\mathbf{E} + \mathbf{v} \times \mathbf{B} = \mathbf{0}, \quad (1)$$

$$\frac{\partial \mathbf{B}}{\partial t} = -\nabla \times \mathbf{E}, \quad (2)$$

$$\nabla \cdot \mathbf{B} = 0. \quad (3)$$

Under the assumption that all z -derivatives vanish, one can write the magnetic field in the form

$$\mathbf{B}(x, y, t) = \nabla A(x, y, t) \times \mathbf{e}_z + B_z(x, y, t)\mathbf{e}_z, \quad (4)$$

where $A(x, y, t)$ is the magnetic flux function. This form of the magnetic field automatically satisfies the solenoidal condition, Eq. (3).

Following [Giuliani et al. \(2005\)](#) we will assume in this paper that both B_z and v_z vanish (for an extension of the theory to 3D, see [Grady & Neukirch 2009](#)). Here the x -coordinate runs parallel to the solar surface, while the y -coordinate specifies the height above the solar surface.

The ideal Ohm's law (1) then takes the form

$$\frac{\partial A}{\partial t} + \mathbf{v} \cdot \nabla A = 0. \quad (5)$$

A CMT model is then defined by specifying the flux function A at a specific time and a velocity field $\mathbf{v}(x, y, t)$. In the present paper we will use the flux function defining the magnetic field as $t \rightarrow \infty$. Instead of defining the velocity field directly, we specify a transformation from Eulerian to Lagrangian coordinates, which is equivalent to specifying the trajectories of plasma elements perpendicular to the magnetic field.

In this paper we use the same CMT model as used by [Giuliani et al. \(2005\)](#). The final magnetic field is determined by the flux function

$$A_0 = c_1 \arctan\left(\frac{y_0 + d_1}{x_0 + w}\right) - c_1 \arctan\left(\frac{y_0 + d_2}{x_0 - w}\right), \quad (6)$$

where c_1 is used to control the strength of the magnetic field and where x_0, y_0, d_1, d_2 and w are considered to be normalised by length L , which is the characteristic size of the trap.

The corresponding magnetic field is a loop between two 2D magnetic sources (line currents) separated by a distance $2w$ and placed below the photosphere at $y_0 = -d_1$ at $x_0 = -w$ and $y_0 = -d_2$ at $x_0 = w$. The magnetic field generated by the flux function (6) is potential if regarded as a function of x_0 and y_0 . This potential field is the final field to which the CMT relaxes as $t \rightarrow \infty$. [Giuliani et al. \(2005\)](#) used the values $d_1 = d_2 = 1$ and $w = 0.5$ in their paper, which creates a magnetic loop which is symmetric with respect to $x = 0$ at all times.

The coordinate transformation used by [Giuliani et al. \(2005\)](#) is given by

$$x_0 = x, \quad (7)$$

$$y_0 = (at)^b \ln \left[1 + \frac{y}{(at)^b} \right] \left\{ \frac{1 + \tanh[(y - L_v/L)a_1]}{2} \right\} + \left\{ \frac{1 + \tanh[(y - L_v/L)a_1]}{2} \right\} y. \quad (8)$$

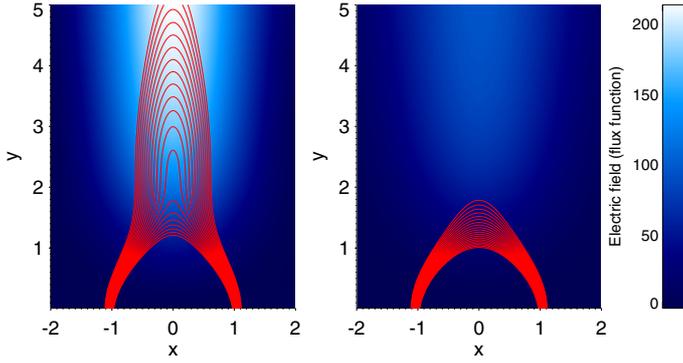


Fig. 1. The magnetic field lines (red) and the electric field (blue colour scale) of the CMT model by [Giuliani et al. \(2005\)](#) at the beginning (*left panel*) and at the end (*right panel*) of the collapse, corresponding to 95 s in the normalisation used. Lengths are normalised to $L = 10^7$ m.

This transformation stretches the magnetic field in the y -direction above a height given by L_v/L , where the transition between unstretched and stretched field is controlled by the parameter a_1 . In this paper we will use the same parameter values as [Giuliani et al. \(2005\)](#), namely $a = 0.4$, $b = 1.0$, $L_v/L = 1$ and $a_1 = 0.9$. For simplicity, the transformation depends on time only through the function $y_0(y, t)$. This time-dependence lets the field collapse to the final field described above as for $t \rightarrow \infty$, y_0 tends to y . An important feature of the transformation is that the foot points of magnetic field lines do not move during the collapse as for $y = 0$ we have $y_0 = 0$ for all t . In [Fig. 1](#) we show a plot of the magnetic field lines, and the electric field, at the start and at the end of the collapse.

For reasons of comparability, we also use the same normalisation as used by [Giuliani et al. \(2005\)](#), i.e. the typical length scale of the trap is $L = 10^7$ m, the magnetic field is normalised by 0.01 T (100 G) and the time scale for the collapse of the trap is 100 s. We remark that these are rather conservative values, giving, for example, a typical field strength of 2×10^{-3} T (20 G) and below at the initial positions of the particles. Stronger magnetic fields and smaller time-scales of CMT collapse are possible for solar flares and the consequences of varying the CMT parameters, and also the CMT model itself will be investigated in more detail in the future.

The particle orbits are calculated using first order non-relativistic guiding centre theory (see e.g. [Northrop 1963](#))

$$\left(\frac{m}{q}\right) \frac{dv_{\parallel}}{dt} = E_{\parallel} - \frac{M}{q} \frac{\partial B}{\partial s} + \left(\frac{m}{q}\right) \mathbf{u}_{\mathbf{E}} \cdot \left(\frac{\partial \mathbf{b}}{\partial t} + v_{\parallel} \frac{\partial \mathbf{b}}{\partial s} + \mathbf{u}_{\mathbf{E}} \cdot \nabla \mathbf{b}\right),$$

$$\dot{\mathbf{R}}_{\perp} = \frac{\mathbf{b}}{B} \times \left\{ -\mathbf{E} + \frac{M}{q} \nabla B + \frac{m}{q} \left[v_{\parallel} \frac{\partial \mathbf{b}}{\partial t} + v_{\parallel}^2 \frac{\partial \mathbf{b}}{\partial s} + v_{\parallel} \mathbf{u}_{\mathbf{E}} \cdot \nabla \mathbf{b} + \frac{\partial \mathbf{u}_{\mathbf{E}}}{\partial t} + v_{\parallel} \frac{\partial \mathbf{u}_{\mathbf{E}}}{\partial s} + \mathbf{u}_{\mathbf{E}} \cdot \nabla \mathbf{u}_{\mathbf{E}} \right] \right\},$$

where $M = (1/2)mv_{\mathbf{g}}^2/B$ is the magnetic moment, which is regarded as constant in this approximation and thus fixed by the initial conditions, $v_{\mathbf{g}}$ is the gyro-velocity, $\mathbf{u}_{\mathbf{E}} = (\mathbf{E} \times \mathbf{b})/B$, $\mathbf{b} = \mathbf{B}/B$, \mathbf{R} is the vector location of the guiding centre, $v_{\parallel} = \mathbf{b} \cdot \dot{\mathbf{R}}$ and $\dot{\mathbf{R}}_{\perp} = \dot{\mathbf{R}} - v_{\parallel} \mathbf{b}$ (for evolution equations for the perpendicular and parallel energies, see [Giuliani et al. 2005](#), Eqs. (49) and (50)).

In the present paper we only calculate electron orbits. For all electron orbits presented in this paper the use of the guiding centre approximation is well justified, as, for example, the typical ratio between the gyration timescale for electrons and the

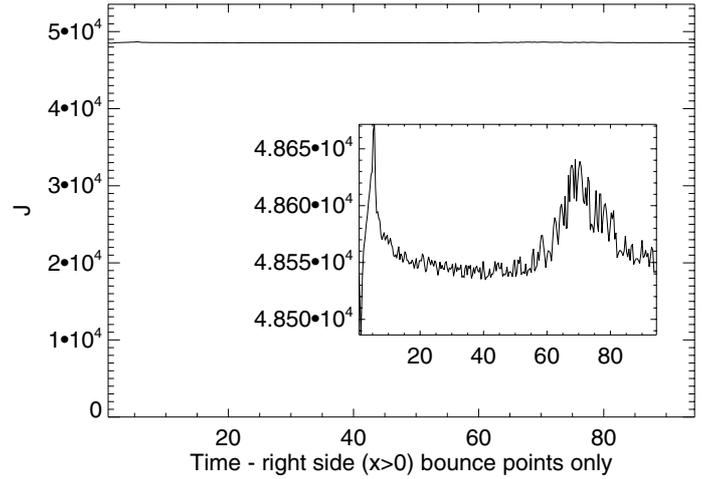


Fig. 2. Time variation of the longitudinal invariant J for the electron orbit in [Giuliani et al. \(2005\)](#). Inset shows magnified y -axis.

time-scale for the variation of the magnetic field of the CMT is of the order 10^{-3} – 10^{-4} , and the differences between typical gyroradii and the MHD length scales of the CMT are also of this order.

The longitudinal adiabatic invariant

$$J = \oint mv_{\parallel} ds, \quad (9)$$

is a good invariant for the orbits studied in this paper. As an example we show in [Fig. 2](#) the variation of J for the orbit discussed by [Giuliani et al. \(2005\)](#). We evaluated the path integral along the orbit between two consecutive mirror points on the same side of the trap. The variation of J along the orbit is less than 0.5%. This is typical for all orbits studied in this paper.

3. Electron energy gains for varying initial conditions

3.1. Discussion of initial conditions

In this section we investigate the influence of initial conditions on the energy gain of electrons in the CMT model of [Giuliani et al. \(2005\)](#). Generally, the initial position and initial velocity are varied. Regarding the initial position, only the initial x and y -values need to be varied, because the CMT model is invariant in the z -direction. Due to our use of guiding centre theory, we do not need to specify the complete initial velocity vector. In the present paper, we choose to specify the total initial kinetic energy and the initial pitch angle (α) of the particles. Together with the initial position, this fixes the magnetic moment (M) of the particles. It also implicitly fixes the initial parallel ($E_{\parallel} = \frac{1}{2}mv_{\parallel}^2$) and perpendicular ($E_{\perp} = MB$) energies of the particle.

In the following we distinguish between particle orbits that have $y > 0$ for all times (trapped particles) and particle orbits that eventually cross the lower boundary ($y = 0$; escaping particles). For escaping particles, the final energy and other quantities are recorded at the time of escape, i.e. when their orbit first reaches a value $y < 0$, whereas for particles which remain trapped the corresponding values are recorded at the final time of the calculation, i.e. when the trap is sufficiently relaxed.

To study the effects of varying the initial conditions, we use a grid of 11 by 11 equidistantly spaced initial positions for $-0.5L \leq x \leq 0.5L$ and $1L \leq y \leq 5L$ (see diamond

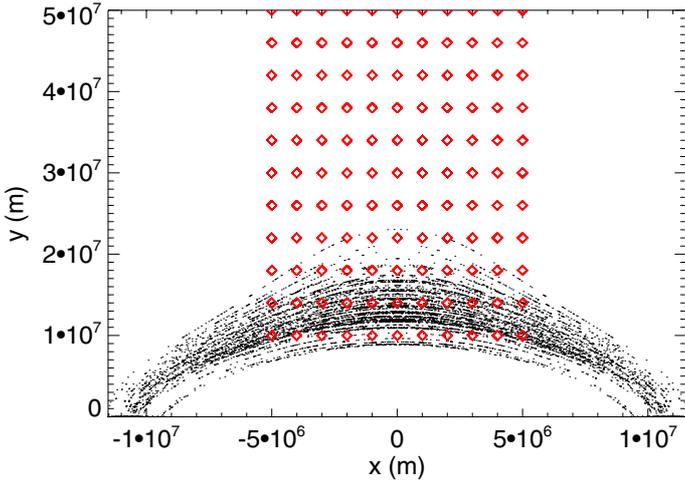


Fig. 3. Initial (diamond) and final (dot) positions for test particles. We remark that we only represent individual orbits, which are not representative of, for example, particle densities. A stronger concentration of particles in the centre of the trap at the initial time would lead to an even higher density of energetic particles trapped at the loop top at the final time.

shaped symbols in Fig. 3). For each initial position we calculate particle orbits for 11 equally spaced values between 5 keV and 6 keV for the initial energy and 10 values for the initial pitch angle between 13° and 163° . In Fig. 3 we also show the final positions of the particles remaining in the trap as dots. Particles which start at positions further away from the centre of the trap ($x = 0$) are more likely to escape quickly even for initial pitch angles relatively close to 90° , often without mirroring, whereas particles starting close to the centre are more likely to remain trapped. The reason is that outside the main region of the CMT the magnetic field strength does not vary as much as inside the CMT and mirroring is less likely to occur, i.e. the particles start with initial conditions putting them inside the local loss cone.

3.2. Dependence of energy gain on initial position

Figure 4 shows the energy gain of particles as the ratio between the final and initial energy. The values of initial energies chosen can be identified as the vertical bands on the graph. For the initial conditions investigated here, we find that the final energy can be up to 53 times the initial energy (top boundary of Fig. 4). Most particles (98.5% of the initial conditions shown) have modest energy gains of up to a factor of 10.

Furthermore, for 2% of the initial conditions shown, the particles lose energy compared to the initial state, but these are all particles which escape the trap almost immediately (within 1.5 s in the normalisation discussed above). These particles all start outwith the central region of the CMT and usually have parallel velocities which take them directly to the nearest foot point of the field line they start on. Even some of the particles staying longer within the CMT are actually never really trapped, i.e. they do not mirror before crossing the lower boundary ($y = 0$). These particles usually have an initial parallel velocity which takes them in the direction of the foot point further away from their starting position, which means they simply take longer to reach the lower boundary. It is interesting that, despite not being trapped, even some of these particles gain energy because they encounter stronger magnetic field values while travelling to the point where they leave the CMT.

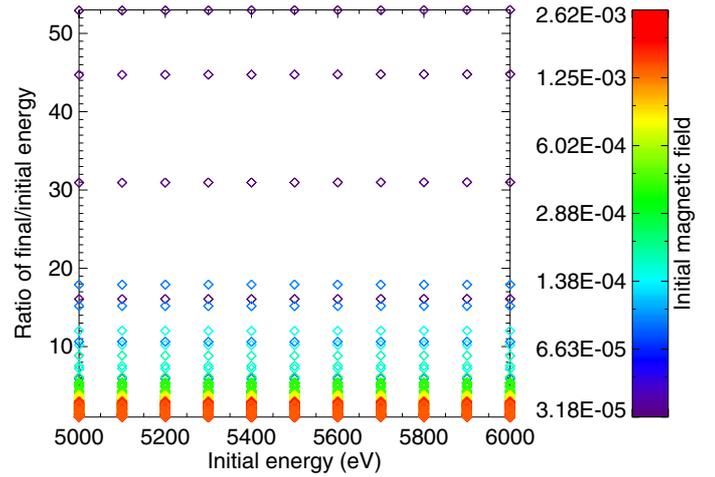


Fig. 4. Ratio of final to initial energy. Each point indicates a different test particle. Colours show, on logarithmic scale, the magnetic field strength at the initial positions of the orbits. One can see a clear trend that higher energy gains are correlated with initial positions in weak field regions. The energy gain does, however, also depend on the initial pitch angle, with orbits with initial pitch angles close to 90° gaining more energy.

A closer investigation shows that the ratio between final and initial energy is determined very strongly by the initial position and the initial pitch angle, and only to a much lesser extent by the initial energy, at least over the range of initial energies studied in the present paper (we did carry out a limited number of test particle calculations for much smaller initial energies around a few 100 eV, but found no qualitative difference in the energisation process; higher initial energies would lead to final energies close to the electron rest mass energy and would thus require a relativistic calculation). In particular, the initial position determines the initial magnetic field strength that the particle experiences. In Fig. 4 the magnetic field strength at the initial position of the particles is indicated by the colour of the symbols, with the values being shown by the colour bar. As a general trend, particles starting in regions of lower magnetic field have the higher energy gains. Although the initial pitch angle is not indicated in Fig. 4, we find that apart from starting in a region with lower magnetic field strength, the orbits with the highest energy gains also have initial pitch angles which are closest to 90° , i.e. the particles have small initial parallel energies. The nearly horizontal bands seen in Fig. 4 are actually made up of particle orbits which start at the same initial position with the same initial pitch angle, but different total initial energies. We found that the particles with the largest energy gain ratio (above 30) start in the centre of the trap (at $x = 0$, $y = 2.2 L$ for those plotted here) with a pitch angle close to 90° . These findings indicate that the betatron effect plays a major role for particles with the highest energy gains.

3.3. The effect of the initial pitch angle

To investigate more closely the effect of the initial pitch angle on the energy gain, we show in Fig. 5 the final energies (colour contours) for particle orbits starting at the same position ($x = 0.1 L$, $y = 2.0 L$), but with different initial energy (x -axis) and pitch angle (y -axis). For these orbits the initial pitch angle varies between 1.8° and 178.2° , and initial energy varies between 5 keV and 6 keV as shown. Particles that remain in the trap until the final time have initial pitch angles between 19.6°

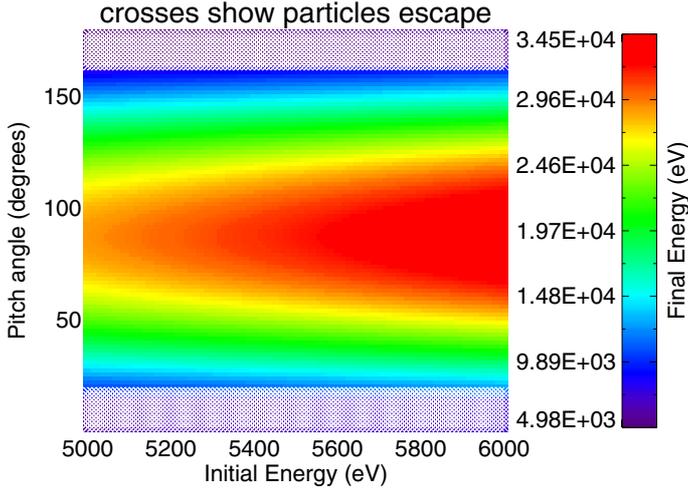


Fig. 5. Final energy (colour contours) of test particles with the same initial position ($x = 0.1 L$, $y = 2.0 L$), but different initial energies (y -axis) and pitch angles (x -axis). Crossed squares indicate particles that escape before the trap has collapsed. For this initial position the highest energy particles have pitch angles closest to 90° .

and 162.2° . Particles that eventually escape the trap had initial pitch angles $\leq 17.8^\circ$ or $\geq 164.0^\circ$ in our grid of initial pitch angle values (both the pitch angle and the loss cone angle in our CMT model are time-dependent and therefore no general simple condition for particle escape can be given). Escaping particles are indicated by crossed squares in Fig. 5. The particles ending up with the highest final energy (about 34.5 keV) have pitch angle closest to 90° and start with the highest initial energy, consistent with the conclusions of the previous section.

The effect of varying the initial pitch angle at different initial positions, for a fixed initial energy of 5.5 keV, is shown in Fig. 6. The plots show final energy distributions versus initial pitch angle for any combination of initial positions out of $x = 0.0, 0.12, 0.24, 0.36, 0.48$ and $0.6 L$ with $y = 1.0, 1.8, 2.6, 3.4, 4.2$ and $5.0 L$. Basically, every plot shown in Fig. 6 can be considered as a vertical cut through a figure similar to Fig. 5 at 5.5 keV for each of the initial positions. In Fig. 6, black dots indicate particle orbits which remain trapped, whereas red dots indicate escaping particle orbits. It is obvious that particles with pitch angles deviating substantially from 90° escape from the trap more easily. It can also be easily seen that for initial positions further away from the centre of the trap in the horizontal direction (x -direction), the range of initial pitch angles leading to escaping particle orbits becomes larger.

One can also see again that the particles with the highest final energy start with pitch angles close to 90° in the centre of the CMT ($x = 0.0 L$) and within the region of weak magnetic field (see e.g. $y = 2.6 L$). While there is still a maximum of the final energy distribution around a pitch angle of 90° in the CMT centre for other values of the initial height y , the value of the maximum energy is reduced compared to $y = 2.6 L$. Another feature of the final energy as a function of pitch angle for increasing initial height y is the development of secondary maxima at small and large pitch angles. A first indication is already visible for $y = 2.6 L$, but becomes increasingly clearer for larger initial y values. The largest energy values of the secondary maxima occur close to the point of transition from trapped to escaping particle orbits. Similar trends as for the CMT centre at $x = 0$ are also seen for the other values of x , although the final energies drop strongly in value. The maximum energy around the pitch angle

of 90° actually turns into a local minimum, with the secondary maxima for small/large pitch angle becoming the highest energies as one moves away from $x = 0$ at constant initial height y . An explanation for these features is that the largest energy gains at the CMT centre are caused by the betatron effect, because the largest increase in magnetic field with time occurs at the centre of the CMT. Particles starting close to the CMT centre with a pitch angle around 90° stay very close to the CMT centre and thus basically gain all their energy through the betatron effect. Particles with small or large pitch angle have larger oscillation amplitudes for parallel and perpendicular energy whilst inside the trap. While the trap collapses, the particles move on field lines which shorten (by becoming less curved) and the distance between successive bounces becomes shorter. These particles could therefore be mainly accelerated by the first order Fermi effect. This would explain the secondary peaks for smaller and larger pitch angles.

Particles starting away from the centre of the CMT do not experience the same large difference between initial and final magnetic field strength as the particles at the centre of the CMT, and thus the betatron effect becomes less efficient as the initial position moves away from the CMT centre. For small or large pitch angles, however, the Fermi effect could still operate, but only for particles outside the loss cone. This would explain the final energy minimum and the secondary maxima for initial positions outside the centre of the CMT.

4. Comparison of two particle orbits with different initial pitch angle

To gain a better understanding of the different acceleration processes described above and how they depend on the initial pitch angle, we investigated in detail two particle orbits with the same initial position ($x = 0 L$, $y = 4.2 L$) and energy (5.5 keV), but with different initial pitch angles. Particle orbit 1 has an initial pitch angle of 165.8° , i.e. the particle is moving initially mainly in the direction opposite to the field line. Particle orbit 2 has a pitch angle of 86.2° , so most of its initial energy is associated with the gyration motion perpendicular rather than parallel to the field. Both particle orbits are shown in Fig. 7. As is to be expected, particle orbit 1 extends far along the field line, well into the legs of the trap, whereas particle orbit 2 remains close to the centre and mirrors more frequently. Both orbits remain on the same field line at all times.

4.1. Time evolution of particle energies

The upper left panel of Fig. 8 shows the time evolution of the total kinetic energy for orbit 1. The time evolution of the total energy for this orbit shows features which are very similar to the energy evolution of the particle orbit investigated by Giuliani et al. (2005). The energy increases in steps when the guiding centre moves along the top of the field line it is on, and it decreases slightly closer to the mirror points. As shown by Giuliani et al. (2005), the steps are caused by the curvature term in the parallel equation of motion and gives rise to an initial average increase in parallel energy ($mv_{\parallel}^2/2$). This is confirmed for particle orbit 1 by the plot of the parallel energy shown in the upper right panel of Fig. 8. A clear increase is visible when looking at the envelope of maxima of the parallel energy. These maxima occur when the particle passes through the centre of the trap ($x = 0$), which is consistent with the findings of Giuliani et al. (2005). Obviously, for every trapped particle the minimum value of the

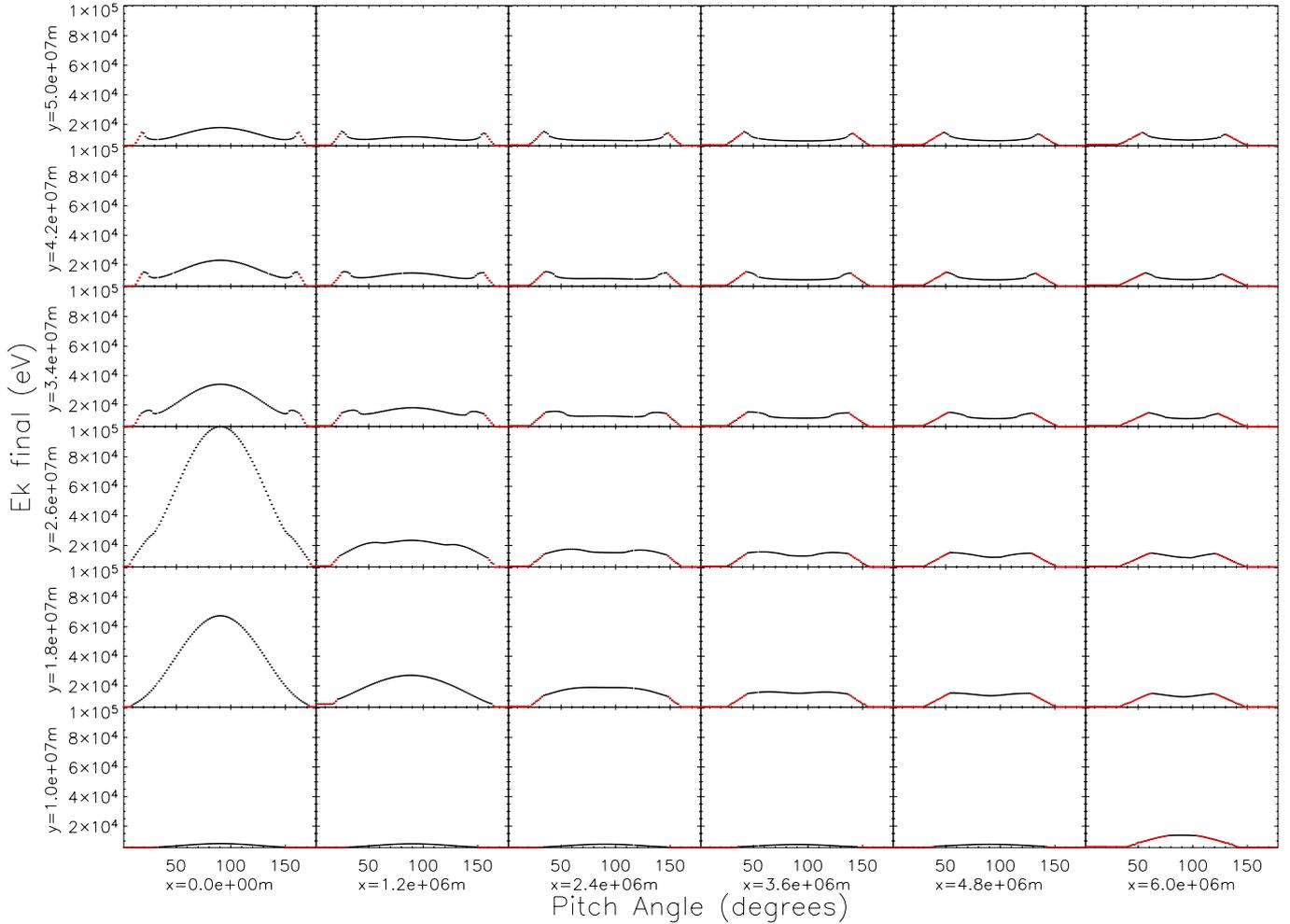


Fig. 6. Final energy vs. pitch angle for different initial positions. All particles start with the same initial kinetic energy (5.5 keV). Each plot is similar to a vertical cut through Fig. 5. Red points indicate which particles escape. Note that the initial pitch angle leading to the maximum energy gain is not always 90° , but depends on initial position.

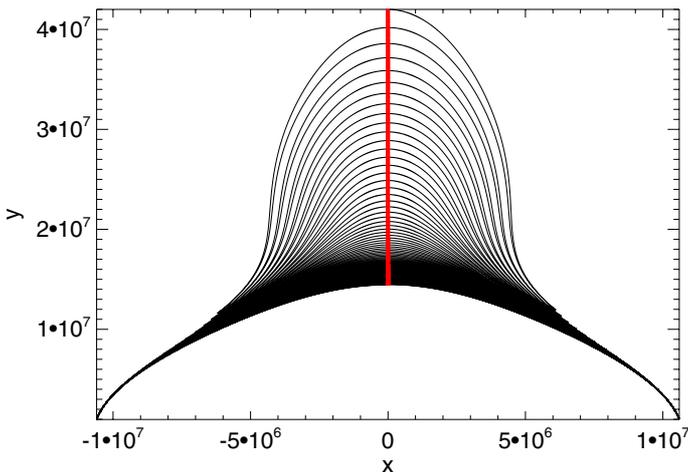


Fig. 7. The two test particle orbits. Particle orbit 1 with initial pitch angle 165.8° is shown in black, particle orbit 2 with initial pitch angle close to 86.2° is shown in red.

parallel energy is zero (at the mirror points), but on average the parallel energy increases with time.

At the same time the perpendicular energy associated with gyrational motion of the particle (MB) also increases on average,

as shown in the lower left panel of Fig. 8. The perpendicular energy has its maximum values at the mirror points and its minimum when passing through the centre of the CMT. However, even at the centre of the CMT, the perpendicular energy is increasing with time. The increase of the perpendicular energy is clearly a consequence of the collapse of the magnetic field and the corresponding increase in magnetic field strength along the particle orbit.

For comparison, we also show in the lower right panel of Fig. 8 the energy associated with the $\mathbf{E} \times \mathbf{B}$ -drift motion, $m u_E^2/2$, where

$$u_E = \frac{1}{B^2} \mathbf{E} \times \mathbf{B}. \quad (10)$$

Compared to the other parallel and perpendicular energies, the energy of the $\mathbf{E} \times \mathbf{B}$ -drift motion is insignificant (here it is smaller by a factor of about 10^{-4}). Even at the initial time, the energy due to this drift is not significant when compared to the others, contributing 0.1% at most. As the CMT collapses the energy associated with $\mathbf{E} \times \mathbf{B}$ -drift generally decreases to zero.

Obviously, due to the nature of trapped particle motion there is a constant interchange between parallel and perpendicular energy along any trapped particle orbit and the two energy forms show the corresponding oscillations between maximum and minimum values. Naturally, these oscillations are out of

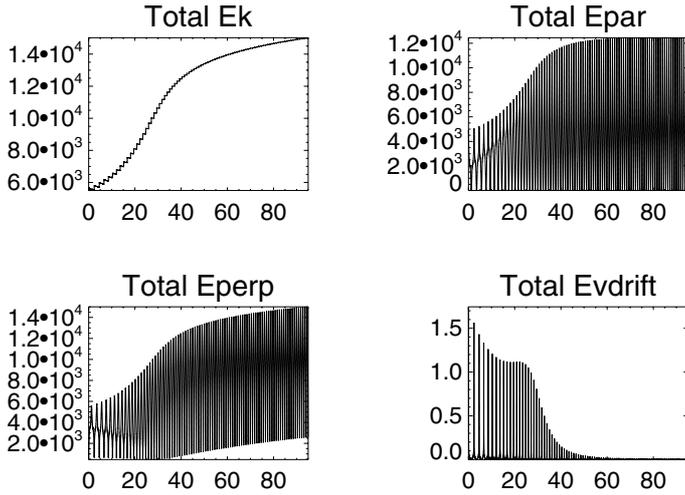


Fig. 8. Plots of time evolution of energy for particle orbit 1 (initial pitch angle 165.8°). Shown are the total kinetic energy (*upper left panel*), the parallel energy (*upper right panel*), the perpendicular energy (*lower left panel*), and the energy associated with the $\mathbf{E} \times \mathbf{B}$ -drift motion (*lower right panel*). In the normalisation discussed in the text, the numbers on the x -axis can be interpreted as seconds and the numbers on the y -axis as electron volts.

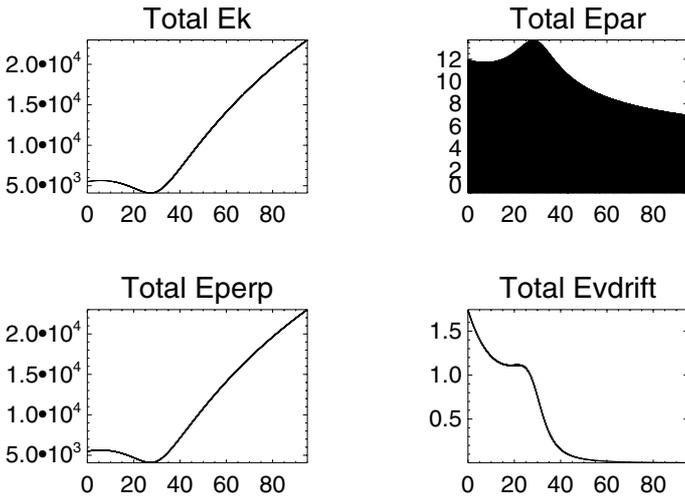


Fig. 9. Plots of time evolution of energy for particle orbit 2 (initial pitch angle 86.2°). Shown are the total kinetic energy (*upper left panel*), the parallel energy (*upper right panel*), the perpendicular energy (*lower left panel*), and the energy associated with the $\mathbf{E} \times \mathbf{B}$ -drift motion (*lower right panel*). Using the normalisation discussed in the text, the numbers on the x -axis can be interpreted as seconds and the numbers on the y -axis as electron volts.

phase and if added up lead to the total energy not having any oscillation apart from the step-like behaviour discussed above. The average perpendicular energy and the average parallel energy are comparable for this orbit.

For particle orbit 2, the energy shows a very different behaviour (see Fig. 9). The total energy (upper left panel) again shows an overall gain, but only after some energy decrease at the beginning. The step-like behaviour, seen in the total energy for orbit 1, is not visible for orbit 2. The parallel energy (shown in the upper right panel) is again periodic, but this is more difficult to see as there are far more bounces due to the particle being trapped with mirror points very close to the centre of the CMT. It should also be noted that the parallel energy for this particle orbit is three orders of magnitude smaller than the total

energy. This explains why we do not find the step-like behaviour seen for orbit 1, as it is simply too small to see on the scale of the total energy, although a closer investigation shows that it is still present, but with a much smaller amplitude than for orbit 1. We also remark that the peak seen in the parallel energy at about 28 s corresponds to the minimum in total energy around the same time.

As the parallel energy is so much smaller than the total energy, it is clear that the perpendicular energy must be the dominating contribution to the total energy, and the two are indeed almost identical (see lower left panel). As for particle orbit 1, the energy associated with the $\mathbf{E} \times \mathbf{B}$ -drift motion is negligible (see lower right panel). However, a closer investigation shows, similar to the case of orbit 1, there are still small periodic variations in the perpendicular energy, although they are not visible on the scale shown here. Because the bounce points are close to the centre of the CMT, the magnetic field does not change much over the period of a single particle oscillation, and thus $E_\perp = MB$ does not change much either. An interesting feature of the perpendicular and the total energy time evolution is that there is an energy decrease to start with and that both energies increase only after they have gone through a minimum. This feature can be explained quite easily by looking at the magnetic field structure of our CMT model. The CMT magnetic field strength has its minimum in the centre of the CMT at a height of about $y = 2L$ at the beginning. Although the magnetic field evolves in time and the minimum in magnetic field strength eventually disappears, particles initially situated above this minimum and moving mainly downwards with collapsing field lines in the centre of the CMT will pass through this minimum magnetic field strength region and their perpendicular energy will decrease accordingly. Once they have passed through that region the magnetic field will increase again and the perpendicular energy will increase as well, which is exactly what is seen in the two left panels of Fig. 9. Generally, we can conclude that for particle orbits like orbit 2, the betatron effect is the dominating mechanism of energy gain.

As orbits 1 and 2 start at the same initial position, they must both pass through the field strength minimum, although orbit 1 will only pass through it when in the centre of the CMT, i.e. when its perpendicular energy is at its minimum. A closer investigation does show that the graph of the perpendicular energy for orbit 2 has the same shape as the lower envelope of the perpendicular energy plot for orbit 1. An indication of this can be found in the lower left panel of Fig. 8. More generally, any other particle orbit starting at the same position should have a perpendicular energy graph with a lower envelope of the same shape. This shape is determined by variation of the magnetic field strength B with height at the centre of the CMT ($x = 0$). The perpendicular energy graph for any particle is given by the product of B and the magnetic moment, which is a constant in guiding centre theory, and thus the minima of the perpendicular energy correspond to the minima in B along an orbit.

4.2. Longitudinal invariant and bounce length

We already showed that the longitudinal invariant J is very well conserved for the particle orbit investigated in Giuliani et al. (2005). This is also the case for the two orbits discussed above. For particle orbit 1 the maximum value is only 0.037% larger than the minimum, and for particle 2 the maximum is only 0.062% larger than its minimum. Given that J is a good invariant for the two orbits, an interesting question is how the distance between consecutive bounce points changes during the evolution of the trap, because that could indicate the presence of the first

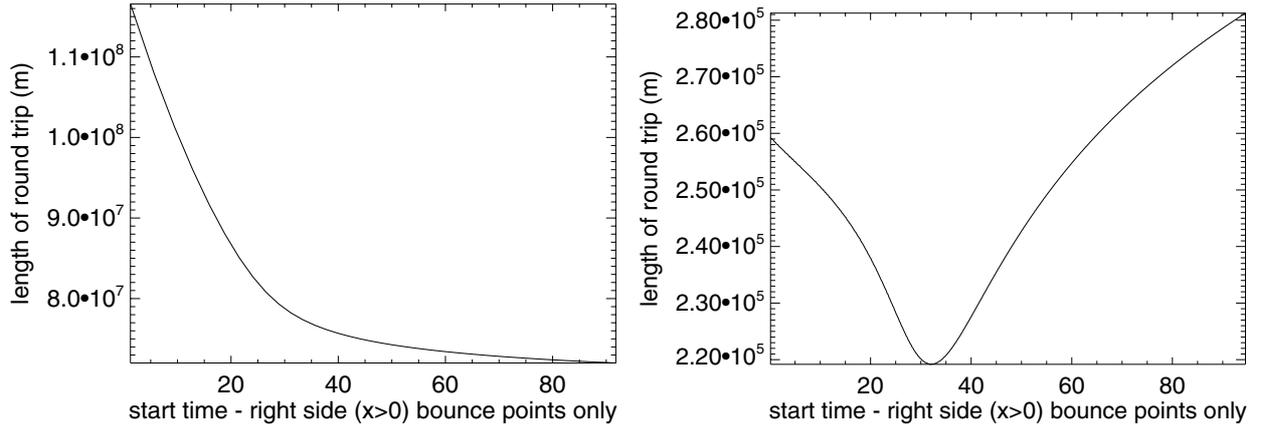


Fig. 10. Bounce length as a function of time for particle orbit 1 (*left panel*) and 2 (*right panel*).

order Fermi mechanism (see e.g. Somov & Kosugi 1997). The bounce lengths for the two orbits are shown in Fig. 10. For particle orbit 1 the length decreases all the time while the trap is collapsing. This would be consistent with interpreting at least part of the energy gain as related to the first order Fermi mechanism. One should, however, bear in mind that, as discussed in detail by Giuliani et al. (2005), the parallel energy increases mainly at the loop top due to the curvature term in the parallel equation of motion, as this gives rise to a source term in the parallel energy equation (e.g. Northrop 1963). For particle orbit 2 the bounce length decreases to a minimum and then increases again. This is consistent with the increase of average parallel energy at the beginning and decrease of the average parallel energy in the later stages of the collapse, as shown in the upper right panel of Fig. 9.

5. An asymmetric trap model

5.1. Influence of initial conditions

It is very unlikely that a solar flare would develop in perfect symmetry as we have assumed in the previous models. To investigate how the particle energisation processes change for an asymmetric CMT model, we make a simple modification to the symmetric model used so far by placing the magnetic sources at different heights below the photosphere ($y = 0$). In particular, we now choose $d_1 = 1$ and $d_2 = 1.5$ in Eq. (6). Using a larger value for d_2 means that the negative magnetic source on the right is farther below $y = 0$ so the magnetic field above it at $y = 0$ will be weaker than above the left source. Therefore, we expect the particles to penetrate deeper into the side with the weaker magnetic field.

Because of the asymmetry of the CMT, it is no longer possible to clearly define the centre of the magnetic trap as a single unique value of x as for the symmetric case with $x = 0$. We will therefore use the term “central trap region”, which we regard as the region defined roughly by the mid-points between the two mirror points of trapped particles. Due to the asymmetry, we have to use more initial positions in the x direction to study the differences between particles starting on the left side and the right side of the trap. We keep the other initial condition ranges the same as the symmetric trap. We use a grid of 22 by 11 equidistantly spaced initial positions for $-0.5 L \leq x \leq 1.6 L$ and $1 L \leq y \leq 5 L$ (see diamond shaped symbols in Fig. 11). As in the symmetric trap, for each initial position we calculate particle orbits for 11 equally spaced values between 5 keV and 6 keV for the initial energy and 10 values for the initial pitch angle

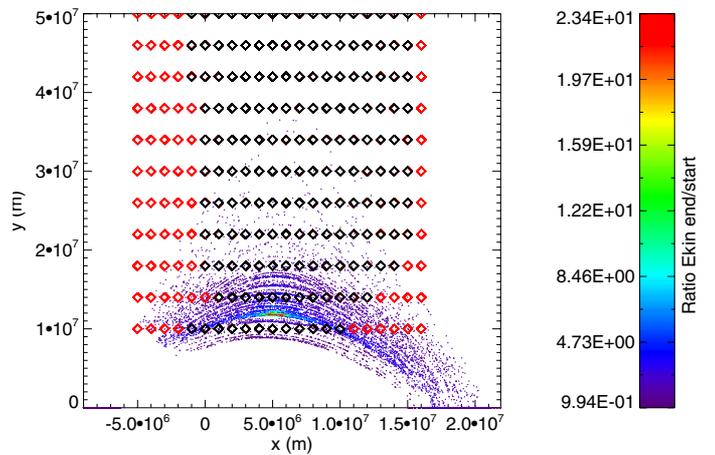


Fig. 11. Initial (diamond) and final (dot) positions of particles in the asymmetric trap. Black diamonds show that some particles starting there are trapped throughout. Colour of final position indicates energy gain.

between 16° and 163° . In Fig. 11 we also show the final positions of the particles remaining in the trap as dots, with the colour bar showing the energy gain. The highest energy particles are still trapped in the region close to the loop top.

Again, many of the particles starting far away from the central trap region escape quickly. In Fig. 11 all the particles starting at a position marked by red diamond escape from the trap before the final time. It may seem from Fig. 11 that the final positions of the particles seem to be generally higher than the final positions for the symmetric trap model shown in Fig. 3. However, this is only due to the fact that we have extended the region of starting positions in x and hence some particles starting with larger initial $|x|$ values are located on field lines that extend to larger heights in the central trap region.

One can see that as expected in this asymmetric trap model the trapped particles mirror closer to the right footpoint where the magnetic field is weaker. The escaping particles are indicated by the dots on the x -axis in Fig. 11. We see that on the right hand side the locations where the particles have crossed the x -axis is much wider than on the left-hand side, which is easily explained by the fact that the magnetic field is weaker at the right hand side foot points.

Another way to look at how the initial conditions can affect particle acceleration is to start particles with different pitch angles but all having the same initial energy from each position.

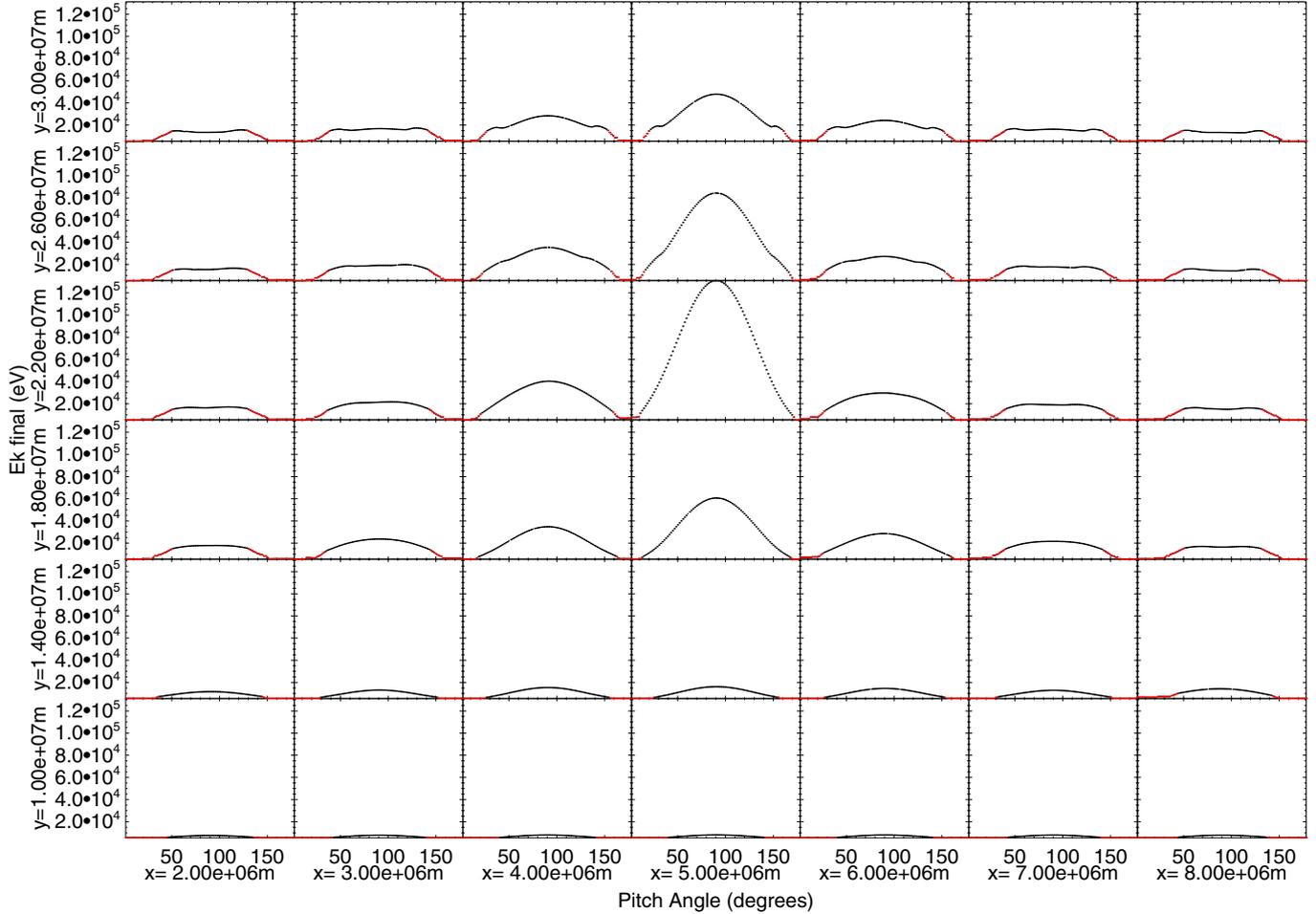


Fig. 12. Final energy vs. pitch angle for different starting positions in the asymmetric trap. All particles start with energy 5.5 keV. The region shown is centred on the maximum energy gain. The spacing is different to Fig. 6. Red shows particles that escape before the trap has collapsed.

This is shown in Fig. 12. Each box shows particles starting at different positions in the trap. The energy is plotted at the final trap time or at the time of escape. Red points indicate escape from the trap. As with the symmetric trap, the maximum energy ratio is for particles starting near the trap centre with pitch angles close to 90°. Also as in the symmetric trap, there is a secondary effect with “wings” on these graphs when the pitch angle is more field aligned and the starting position is further from the central trap region. From these points, the particles final energy can be up to 131 keV, 24 times the initial energy.

5.2. Energy gain in the asymmetric trap

Asymmetries are also noticeable in the energy graph of individual particles in the trap. Figure 13 shows the complete orbit of a particle in the asymmetric trap in the left panel and the time evolution of its energy for the first 35 s in the right panel. We show only the first 35 s of the time evolution of the total energy, because after that time the step-like fashion of the energy increase and decrease phases is strongly reduced and it is more difficult to identify the different parts of the orbit. The part of the particle orbit after 35 s is indicated in black, for completeness. This specific particle has initial energy 5 keV and pitch angle 81.8°, starting position $x = 0.1$, $y = 3.8$, but the results for this particle are representative for most particles in the asymmetric case.

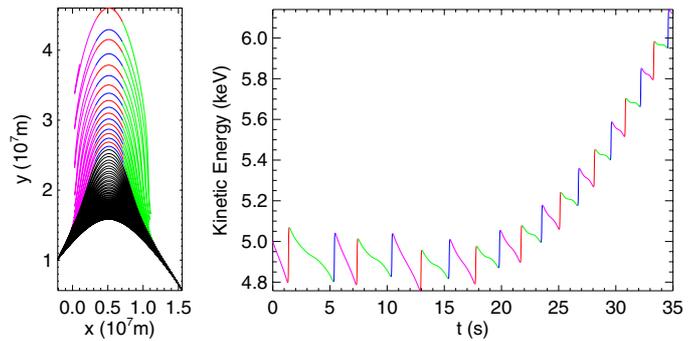


Fig. 13. Typical orbit and energy of a particle in the asymmetric trap. For the first 35 s, colours show corresponding parts where particles gain (red, blue) and lose (green, magenta) energy.

To illustrate the details of the particle energy increase and decrease along its orbit better, we present in Fig. 13 the parts of the orbit in which the particle’s energy increases and the parts where it decreases in a different colours. Specifically, the parts where the particle’s energy increases are shown in red and blue, and the parts where its energy decreases are shown in green and magenta. The reason for using two different colours for the parts of the orbit where the particle energy increases or decreases, respectively, is that this allows us to distinguish between the right- and left-hand side of the magnetic field configuration for the

parts of the orbit where the energy decreases and between the different directions in which the particle passes through the central trap region for the parts of the orbit in which the particle energy increases.

The boundaries of the different coloured sections are defined by the local minima and maxima of the total energy. Looking at the sections where the total energy decreases, we see that the particle energy decreases by a larger amount in the magenta sections (left hand side of the trap, i.e. stronger magnetic field) than the green sections (right hand side of the trap, i.e. weaker magnetic field), although the particle remains in the parts of its orbit coloured in green for a longer time than in the magenta sections. Here we ignore the initial magenta region in which the particle is located at the start.

The left panel of Fig. 13 showing the actual particle orbit provides some insight into these findings. We can see that the particle energy decreases along the parts of the orbit when the particle approaches the mirror points. As expected from the difference in magnetic field strength in the asymmetric CMT model, the particle mirrors closer to the right hand side foot-point of the magnetic field line it is located on. This corresponds to the green part of the orbit (left panel) and of the energy graph (right panel) in Fig. 13. We see also that because the particle penetrates deeper into the magnetic field in the right hand side, it spends longer in that part of the CMT than on the left hand side (magenta section), where the mirror point is located higher up due to the higher magnetic field strength. Also the particle energy decrease by a smaller amount than along the magenta section on the left.

The particle gains energy only when it is passing through the central trap region, shown in blue and red in the left panel of Fig. 13. This is consistent with the results for the symmetric CMT model, already discussed in some detail by [Giuliani et al. \(2005\)](#), who identified the terms related to the curvature of the magnetic field lines in the equation for the parallel velocity as being responsible for the energy increase of the particles. At first sight, the fact that the particle energy increases in the central region of the trap is difficult to reconcile with a simple interpretation of the energy gain in terms of Fermi acceleration model, which is usually associated with energy gain when the particle bounces off a moving obstacle (see e.g. [Longair 1994](#)), i.e. in a CMT close to the mirror points. However, the Fermi mechanism discussed for CMTs is connected to conservation of the second adiabatic invariant J , which is an integral over one complete bounce period of a particle orbit. Therefore any energy gain process related to this adiabatic invariant can only be understood as operating in an average sense over a complete bounce period of the particle.

Generally we conclude that while there are some quantitative differences between the energy gain and loss process in the symmetric and asymmetric CMT models, we did not find any significant qualitative differences.

6. Summary and discussion

We have presented a detailed study of the particle energization processes in CMTs, using in particular the symmetric CMT model of [Giuliani et al. \(2005\)](#) and a modified asymmetric CMT model based on the same theoretical framework. We found that in the particular symmetric CMT model we studied, particle energies can increase by factors of up to approximately 50, but that most particles experience a more modest energy increase. While the energy increase does not depend strongly on initial energy, it does depend on the initial position of particles in the

CMT and on the initial pitch angle. Particles with the highest increase in energy start in the region of the CMT which initially has the smallest magnetic field strength and usually have pitch angles close to 90° . The energy increase for these particles is caused mainly by the betatron effect as the trap collapses and the magnetic field strength along the orbit increases. Due to their pitch angle these particles remain trapped close to the centre of the trap, which means that at the end of the CMT collapse the highest energy population of particles is confined in a region at the top of the most collapsed magnetic loop. This is consistent with previous results using other CMT models (e.g. [Karlický & Kosugi 2004](#)).

We also found that for particles with initial pitch angles differing substantially from 90° , but outside the loss cone at any time during the collapse of the CMT, a substantial increase in parallel energy is possible. On a superficial level this could be interpreted as first order Fermi acceleration as usually the distance between mirror point is decreasing during the CMT evolution. A more careful investigation, however, corroborates the finding of [Giuliani et al. \(2005\)](#) that the particle energy increase is due to the curvature of field lines in the centre of the CMT. In particular, we found that the increase in parallel energy is caused by the curvature term in the parallel equation of motion, which in our model takes on its maximum value in the centre of the CMT and not close to the mirror points. The Fermi mechanism discussed for CMTs is associated with the conservation of the second adiabatic invariant J , which is an integral over one complete bounce period of a particle orbit and thus should be understood as operating in an average sense over a complete bounce period of the particle.

We found similar results for the asymmetric CMT model. While there are noticeable quantitative differences in the general appearance of the particle orbits that can be easily explained by the magnetic field asymmetry, as well as some differences in the maximum total energy increase that is possible, the qualitative features of particle energisation in the asymmetric CMT model were found to be the same as for the symmetric CMT model.

In view of recent findings that high energy radiation from loop tops or above loop tops ([Masuda et al. 1994](#)) is more common during solar flares than previously thought (see e.g. [Krucker et al. 2008](#), for an excellent review), it is tempting to associate the fact that the highest energy particles are trapped at the top of the loop with hard X-ray loop top sources. This has been suggested in the past by other authors (e.g. [Somov & Kosugi 1997](#); [Karlický & Kosugi 2004](#); [Karlický & Bárta 2006](#); [Minoshima et al. 2010](#)) and attempts have been made at calculating the characteristics of the hard X-ray emission expected from CMT models (e.g. [Karlický & Bárta 2006](#)). In order to assess this properly, the present CMT models should be amended to include collisions with a background plasma along the lines of previous models for hard X-ray loop top emission using static loop models (e.g. [Fletcher 1995](#); [Fletcher & Martens 1998](#)). Apart from being a possible explanation for sources of coronal high energy radiation, the trapping of high energy particles in the corona can also contribute to the explanation of observations of microwave emission from flaring loops (e.g. [Melnikov et al. 2002](#)).

Obviously, like many other models, the CMT model used in the present paper is highly simplified. Apart from the introduction of collisions with a thermal background plasma, a number of other improvements should be made in the future. We have shown in this paper that the curvature term in the parallel equation of motion plays an important role in the particle acceleration process in CMTs. For any 2D CMT model, the curvature drift (as well as, for example, the gradient- B -drift) is actually

directed into the invariant direction (i.e. along the z -direction in our coordinate system). It turns out that the particle orbits do not move too far in the z -direction compared to their motion in the x - y -plane, but it nevertheless raises the question whether the results would change for a 3D CMT model. Grady & Neukirch (2009) have recently presented a generalised theory allowing for 3D CMT models including both 3D magnetic fields and 3D flows, which could be used in the future.

As in the present paper the flow field associated with CMT models is usually assumed to be laminar. It is, however, highly unlikely that a violent event such as a solar flare will give rise to such regular behaviour. A possible improvement for future CMT models might be to add turbulent motion (and the corresponding electromagnetic fields) onto the overall laminar motion associated with the collapse. A possible way of dealing with this is to add a stochastic scattering term to the equations of motion, similar in principle, but different in detail, to a Coulomb collision term. There are several interesting questions that arise in connection with such an approach, for example: how would the particle energization in a turbulent CMT change compared to a laminar CMT? How would the energy density associated with the turbulent flow and electromagnetic fields evolve in a CMT? An interesting aspect of such models would be that they could provide a link between stochastic particle acceleration models (e.g. Miller et al. 1997) and the standard flare scenario in a similar way as proposed by e.g. Hamilton & Petrosian (1992), Park & Petrosian (1995), Petrosian & Liu (2004) and more recently by Liu et al. (2008).

Acknowledgements. We would like to thank the referee for helpful and constructive comments. The financial support by the UK's Science and Technology Facilities Council is gratefully acknowledged.

References

- Aschwanden, M. J. 2002, *Space Sci. Rev.*, 101, 1
 Aschwanden, M. J. 2004, *ApJ*, 608, 554
 Aschwanden, M. J. 2009, *Asian J. Phys.*, 17, 423
 Birn, J., Thomsen, M. F., Borovsky, J. E., et al. 1997, *J. Geophys. Res.*, 102, 2325
 Birn, J., Thomsen, M. F., Borovsky, J. E., et al. 1998, *J. Geophys. Res.*, 103, 9235
 Birn, J., Thomsen, M. F., & Hesse, M. 2004, *Phys. Plasmas*, 11, 1825
 Birn, J., Fletcher, L., Hesse, M., & Neukirch, T. 2009, *ApJ*, 695, 1151
 Bogachev, S. A., & Somov, B. V. 2001, *Astron. Rep.*, 45, 157
 Bogachev, S. A., & Somov, B. V. 2005, *Astron. Lett.*, 31, 537
 Bogachev, S. A., & Somov, B. V. 2009, *Astron. Lett.*, 35, 57
 Brown, J. C., Turkmani, R., Kontar, E. P., MacKinnon, A. L., & Vlahos, L. 2009, *A&A*, 508, 993
 Drake, J. F., Swisdak, M., Che, H., & Shay, M. A. 2006, *Nature*, 443, 553
 Emslie, A. G., Kucharek, H., Dennis, B. R., et al. 2004, *J. Geophys. Res.*, 109, 10104
 Emslie, A. G., Dennis, B. R., Holman, G. D., & Hudson, H. S. 2005, *J. Geophys. Res.*, 110, 11103
 Fletcher, L. 1995, *A&A*, 303, L9
 Fletcher, L., & Hudson, H. S. 2008, *ApJ*, 675, 1645
 Fletcher, L., & Martens, P. C. H. 1998, *ApJ*, 505, 418
 Forbes, T. G., & Acton, L. W. 1996, *ApJ*, 459, 330
 Giuliani, P., Neukirch, T., & Wood, P. 2005, *ApJ*, 635, 636
 Grady, K. J., & Neukirch, T. 2009, *A&A*, 508, 1461
 Hamilton, R. J., & Petrosian, V. 1992, *ApJ*, 398, 350
 Karlický, M., & Bárta, M. 2006, *ApJ*, 647, 1472
 Karlický, M., & Bárta, M. 2007, *A&A*, 464, 735
 Karlický, M., & Kosugi, T. 2004, *A&A*, 419, 1159
 Kovalev, V. A., & Somov, B. V. 1998, *Astron. Lett.*, 28, 488
 Kovalev, V. A., & Somov, B. V. 2003a, *Astron. Lett.*, 29, 111
 Kovalev, V. A., & Somov, B. V. 2003b, *Astron. Lett.*, 29, 409
 Krucker, S., Battaglia, M., Cargill, P. J., et al. 2008, *A&ARv*, 8
 Li, Y., & Fleishman, G. D. 2009, *ApJ*, 701, L52
 Lin, J., Forbes, T. G., Priest, E. R., & Bungey, T. N. 1995, *Sol. Phys.*, 159, 275
 Liu, W., Petrosian, V., Dennis, B. R., & Jiang, Y. W. 2008, *ApJ*, 676, 704
 Longair, M. S. 1994, *High energy astrophysics*, Vol. 2. Stars, the Galaxy and the interstellar medium (Cambridge, UK: Cambridge University Press)
 Masuda, S., Kosugi, T., Hara, H., Tsuneta, S., & Ogawara, Y. 1994, *Nature*, 371, 495
 Melnikov, V. F., Shibasaki, K., & Reznikova, V. E. 2002, *ApJ*, 580, L185
 Miller, J. A., Cargill, P. J., Emslie, A. G., et al. 1997, *J. Geophys. Res.*, 102, 14, 631
 Minoshima, T., Masuda, S., & Miyoshi, Y. 2010, *ApJ*, 714, 332
 Neukirch, T. 2005, in *The Dynamic Sun: Challenges for Theory and Observations*, ESA Spec. Publ., 600
 Neukirch, T., Giuliani, P., & Wood, P. D. 2007, in *Reconnection of Magnetic Fields*, eds. J. Birn, & E. Priest (Cambridge University Press), 281
 Northrop, T. 1963, *The Adiabatic Motion of Charged Particles* (New York: Interscience Publishers, Inc.)
 Park, B. T., & Petrosian, V. 1995, *ApJ*, 446, 699
 Petrosian, V., & Liu, S. 2004, *ApJ*, 610, 550
 Reeves, K. K., Guild, T. B., Hughes, W. J., et al. 2008a, *J. Geophys. Res.*, 113, 0
 Reeves, K. K., Seaton, D. B., & Forbes, T. G. 2008b, *ApJ*, 675, 868
 Somov, B. 1992, *Physical Processes in Solar Flares* (Kluwer Academic Publishers), *Astrophys. Space Sci. Lib.*, 172
 Somov, B. V., & Bogachev, S. A. 2003, *Astron. Lett.*, 29, 621
 Somov, B. V., & Kosugi, T. 1997, *ApJ*, 485, 859
 Veronig, A. M., Karlický, M., Vršnak, B., et al. 2006, *A&A*, 446, 675
 Zharkova, V. V., Arzner, K., Benz, A. O., et al. 2011, *Space Sci. Rev.*, 159, 357