

The bolometric output and host-galaxy properties of obscured AGN in the XMM-COSMOS survey[★]

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ABSTRACT

We present a study of the multi-wavelength properties, from the mid-infrared to the hard X-rays, of a sample of 255 spectroscopically identified X-ray selected type-2 AGN from the XMM-COSMOS survey. Most of them are obscured and the X-ray absorbing column density is determined by either X-ray spectral analyses (for 45% of the sample), or from hardness ratios. Spectral energy distributions (SEDs) are computed for all sources in the sample. The average SEDs in the optical band are dominated by the host-galaxy light, especially at low X-ray luminosities and redshifts. There is also a trend between X-ray and mid-infrared luminosity: the AGN contribution in the infrared is higher at higher X-ray luminosities. We calculate bolometric luminosities, bolometric corrections, stellar masses and star formation rates (SFRs) for these sources using a multi-component modeling to properly disentangle the emission associated to stellar light from that due to black hole accretion. For 90% of the sample we also have the morphological classifications obtained with an upgraded version of the Zurich estimator of structural types (ZEST+). We find that on average type-2 AGN have lower bolometric corrections than type-1 AGN. Moreover, we confirm that the morphologies of AGN host-galaxies indicate that there is a preference for these type-2 AGN to be hosted in bulge-dominated galaxies with stellar masses greater than 10^{10} solar masses.

Key words. methods: statistical – quasars: general – galaxies: active – Galaxy: general – X-rays: general

1. Introduction

The formation and growth of supermassive black holes (SMBHs) and their host-galaxies are related processes. This is supported by various observational signatures: the SMBH mass correlates with the mass of the bulge of the host-galaxy (Magorrian et al. 1998; Marconi & Hunt 2003), with the velocity dispersion of the bulge (Ferrarese & Merritt 2000; Tremaine et al. 2002), and with the luminosity of the bulge (Kormendy & Richstone 1995). From theoretical models, AGN seem to be able to switch off cooling flows in clusters (Hoefl & Brüggén 2004) and star formation in galaxies (Somerville et al. 2008), with the result that the SMBH mass is related to the host-galaxy bulge mass (or vice-versa). Feedback between an accreting SMBH and the host-galaxy may play an important role in galaxy formation and evolution. Understand the role of feedback is a demanding problem for both observers and theorists. Semi-analytical models and hydrodynamical simulations have been developed

to attempt to link the formation and evolution of SMBHs to the structure formation over cosmic time. These models invoke different mechanisms to fuel the central SMBHs and to build the host-galaxy bulges, such as major/minor mergers of galaxies (e.g., Corbin 2000; Kauffmann & Haehnelt 2000; Springel et al. 2005; Hopkins et al. 2006), smooth accretion of cold gas from filamentary structures (e.g., Kereš et al. 2009; Dekel et al. 2009), or accretion of recycled gas from dying stars (e.g., Ciotti et al. 2010). Several works also consider radiative feedback which can reproduce two important phases of galaxy evolution, namely an obscured-cold-phase, when the bulk of star formation and black hole accretion occurs, and the following quiescent hot phase in which accretion remains highly sub-Eddington and unobscured (e.g., Sazonov et al. 2005; Lusso & Ciotti 2011). In some of these models, the obscured/unobscured AGN dichotomy is more related to two different phases of galaxy evolution (Hopkins et al. 2008), rather than to an orientation effect (i.e., unified model scheme).

The obscured/unobscured time dependent AGN dichotomy could be related to the bimodality in the rest-frame color distribution of host-galaxies (Rovilos & Georgantopoulos 2007;

[★] Full Table 1 is available at the CDS via anonymous ftp to cdsarc.u-strasbg.fr (130.79.128.5) or via <http://cdsarc.u-strasbg.fr/viz-bin/qcat?J/A+A/534/A110>

Nandra et al. 2007; Brusa et al. 2009; Silverman et al. 2009), namely the red-sequence (or “red-cloud”) and blue-cloud galaxies. Broad-line AGN (if the morphology of the host-galaxy is available) are likely to be associated to galaxies belonging to the blue-cloud, while obscured objects to red passive galaxies. The green valley should be populated by transition objects. The picture above is probably a too crude approximation. Moreover, one should note that red sequence galaxies may well be passively evolving galaxies without significant star formation (e.g., Rovilos & Georgantopoulos 2007; Nandra et al. 2007), rather than dusty starforming objects (e.g., Brusa et al. 2009).

Disentangling the contribution of the nuclear AGN from the host-galaxy properties in the broad band SED is fundamental to constrain the physical evolution of AGN and to place them into the context of galaxy evolution. In the standard picture the AGN energy output is powered by accretion onto SMBHs. The disk accretion emission is visible in the optical-UV as the blue-bump feature. The X-ray emission is believed to be due to a hot-electrons corona that surrounds the accretion disk, while the infrared emission is likely due to the presence of a dusty torus around the disk at few parsec from the center, which reprocesses the nuclear radiation. According to the unified model of AGN (e.g., Antonucci 1993; Urry & Padovani 1995), hot dust is located in the inner edge of the torus. However, recent studies predict and observe exceptions to the unified model. From the theoretical point of view, an alternative solution to the torus is the disk-wind scenario (e.g., Emmering et al. 1992; Elitzur & Shlosman 2006). From the observational side, AGN without any detectable hot dust emission (e.g., Jiang et al. 2010) and weak infrared emission (e.g., Hao et al. 2010) are predicted and observed. The vast majority of studies performed so far concern unobscured (type-1) AGN for which their SED is well known from low- z ($\langle z \rangle \sim 0.206$, see Elvis et al. 1994) to high- z ($\langle z \rangle \sim 1.525$, see Richards et al. 2006). An obvious complication in the study of their host-galaxy properties is that the emission of the central AGN outshines the galaxy light in UV, optical and infrared bands; therefore it is extremely difficult to derive constraints on the colors, stellar populations, and morphologies of the host. On the other hand, for obscured (type-2) AGN the host-galaxy light is the dominant component in the optical/near-infrared SED, while it is difficult to recover the AGN intrinsic nuclear emission. The lack of a proper characterization of the nuclear component of the SED of obscured type-2 AGN is a major limitation. As a consequence, the relations between stellar masses, SFR, morphologies and accretion luminosity remain poorly known. Since the relative contribution in the SED of the different components (AGN/host-galaxy) varies with wavelength, a proper decomposition can be obtained by an SED-fitting approach, complemented by a morphological analysis. This will provide a robust estimate of the nuclear emission (bolometric luminosities and bolometric corrections, absorption column density distributions, etc.) and its relation with the host-galaxy properties (mass, star formation rates, morphological classification). The AGN structure is reflected in the shape of the SED, specifically the big-blue bump and the infrared-bump are related to the accretion disk and the surrounding torus, respectively. Therefore, a densely sampled SED over a broad wavelength interval is mandatory to extract useful information from SED fitting procedures, allowing to tightly constrain physical parameters from multi-component modeling and, in particular, to properly disentangle the emission associated to stellar light from that due to accretion.

The combination of sensitive X-ray and mid-IR observatories allows us to model the obscuring gas that in type-2 AGN hides the nuclear region from the near-IR to the UV. As

supported by previous investigations, the reprocessed IR emission could be a good proxy of the intrinsic disk emission. Gandhi et al. (2009) confirm the correlation between the X-ray luminosity at [2–10] keV and the IR emission for a sample of Seyfert galaxies (see also Lutz et al. 2004). Their data are the best estimate of the nuclear (non stellar) IR flux in AGN to date. A highly significant correlation between $L_{[2-10]\text{keV}}$ and the intrinsic nuclear IR luminosity at $12.3 \mu\text{m}$ is observed in the high quality near-IR and X-ray data discussed by Gandhi et al. (2009). This reinforces the idea that “uncontaminated” mid-IR continuum is an accurate proxy for the intrinsic AGN emission.

In this work we present the largest study of the multi-wavelength properties of an X-ray selected sample of obscured AGN using the XMM-Newton wide field survey in the COSMOS field (XMM-COSMOS). Following a similar approach to that of Pozzi et al. (2007) and Vasudevan et al. (2010), we use the infrared emission to evaluate the nuclear bolometric luminosity from a multi-component fit. The paper is aimed at a detailed characterization of a large sample of obscured AGN over a wide range of frequencies. The SEDs, morphology of the host-galaxies, stellar masses, colors, bolometric luminosities and bolometric corrections for the sample of obscured AGN are presented.

This paper is organized as follows. In Sect. 2 we report the selection criteria for the sample used in this work. Section 3 presents the multi-wavelength data-set, while in Sect. 4 the method to compute average SED is described. Section 5 concerns the multi-component modeling used to disentangle the nuclear emission from the stellar light. In Sect. 6 the method used to compute intrinsic bolometric luminosities and bolometric corrections for type-2 AGN is described, while in Sect. 7 we have applied the same method to a sample of type-1 AGN. The discussion of our findings is given in Sect. 8, while in Sect. 9 we summarize the most important results.

We adopted a flat model of the universe with a Hubble constant $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_M = 0.27$, $\Omega_\Lambda = 1 - \Omega_M$ (Komatsu et al. 2009).

2. The data set

The XMM-COSMOS catalog comprises 1822 point-like X-ray sources detected by *XMM-Newton* over an area of $\sim 2 \text{ deg}^2$. The total exposure time was $\sim 1.5 \text{ Ms}$ with a fairly homogeneous depth of $\sim 50 \text{ ks}$ over a large fraction of the area (Hasinger et al. 2007; Cappelluti et al. 2009). Following Brusa et al. (2010), we excluded from our analysis 25 sources which turned out to be a blend of two *Chandra* sources. This leads to a total of 1797 X-ray selected point-like sources. We restricted the analysis to 1078 X-ray sources detected in the [2–10] keV band at a flux larger than $2.5 \times 10^{-15} \text{ erg s}^{-1} \text{ cm}^{-2}$ (see Table 2 in Cappelluti et al. 2009). The objects for which no secure optical counterpart could be assigned are often affected by severe blending problems, so that we consider in this analysis the 971 sources (hereafter 971-XMM) for which a secure optical counterpart can be associated (see discussion in Brusa et al. 2010)¹.

From the 971-XMM catalog we have selected 255 sources, which do not show broad ($FWHM < 2000 \text{ km s}^{-1}$)

¹ The multi-wavelength XMM-COSMOS catalog can be retrieved from: http://www.mpe.mpg.de/XMMCosmos/xmm53_release/, version 1st April 2010.

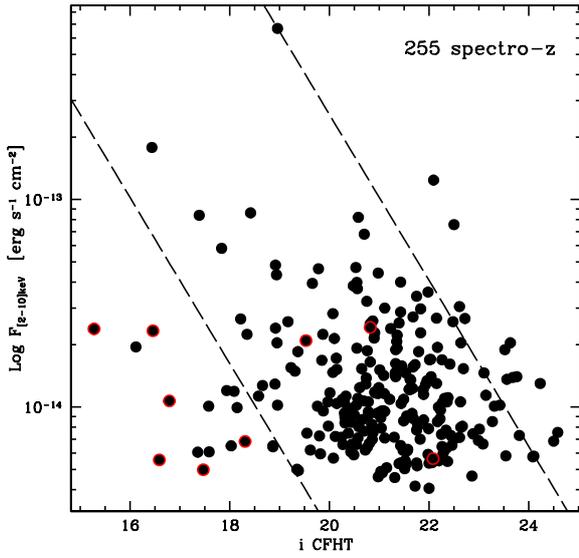


Fig. 1. Plot of the [2–10] keV flux versus the total i^* CFHT magnitude for our sample of 255 type-2 AGN. The red circles represent sources with a de-absorbed 2–10 keV luminosity lower than 10^{42} erg s $^{-1}$. The dashed lines represent a constant X-ray to optical flux ratio of $\text{Log}(X/O) = \pm 1$.

emission lines in their optical spectra² (hereafter we will refer to them as the type-2 AGN sample): 223 are classified not-broad-line AGN, while 32 are absorption-line galaxies. Not-broad-line AGN are objects with unresolved, high-ionization emission lines, exhibiting line ratios indicating AGN activity, and, when high-ionization lines are not detected, or the observed spectral range does not allow to construct line diagnostics, objects without broad line in the optical spectra. Absorption-line galaxies are sources consistent with a typical galaxy spectrum showing only absorption line.

In Fig. 1 we plot the 2–10 keV X-ray flux as a function of i^* CFHT magnitude. The dashed lines limit the region typically occupied by AGN along the X-ray to optical flux ratio $\text{Log}(X/O) = \pm 1$ ³. Nine sources have a de-absorbed 2–10 keV luminosity lower than 10^{42} erg s $^{-1}$, the conventional threshold below which the X-ray sources that can plausibly be explained by moderate-strength starbursts, hot gas in elliptical galaxies, or other sources besides accretion onto a nuclear SMBH (Hornschemeier et al. 2001). The three sources inside the dashed lines have X-ray luminosities close to 10^{42} erg s $^{-1}$, while six AGN (6/255, 2%) lie in the part of the diagram usually occupied by star-forming galaxies, and have X-ray luminosities $< 10^{42}$ erg s $^{-1}$. Their inclusion in the analysis does not affect the main results.

The type-2 AGN sample used in our analysis comprises 255 X-ray selected AGN, all of them with spectroscopic redshifts, spanning a wide range of redshifts ($0.045 < z < 3.524$) and X-ray luminosities ($41.06 \leq \text{Log} L_{[2-10]\text{keV}} \leq 45.0$). The redshift distribution of the total sample and the distribution of the

de-absorbed hard X-ray luminosities are presented in Fig. 2. The mean redshift is $\langle z \rangle = 0.76$, while the mean $\text{Log} L_{[2-10]\text{keV}}$ is 43.34 with a dispersion of 0.64.

2.1. Correction for absorption for the X-ray luminosity

For a sub-sample of 111 AGN we have an estimate of the column density N_{H} from spectral analysis (see Mainieri et al. 2007), while for 144 AGN absorption is estimated from hardness ratios (HR; see Brusa et al. 2010). For 25 sources for which a column density estimate is not available from HR we consider the Galactic column density. Therefore, we can compute the de-absorbed X-ray luminosity at 0.5–2 keV (soft band) and 2–10 keV (hard band) for all sources in our sample. In Fig. 2 (right panel) we show the distribution of column densities that ranges from 3×10^{20} cm $^{-2}$ to 1.5×10^{24} cm $^{-2}$. The mean N_{H} value is 8.5×10^{21} cm $^{-2}$ with a dispersion of 0.72 dex. The integrated intrinsic un-absorbed luminosity is computed assuming a power-law spectrum with slope, $\Gamma = 2$ and $\Gamma = 1.7$ for the 0.5–2 keV and 2–10 keV bands, respectively. The average shift induced by the correction for absorption in the type-2 sample is $\langle \Delta \text{Log} L_{[2-10]\text{keV}} \rangle = 0.04 \pm 0.01$.

3. Rest-frame monochromatic fluxes and spectral energy distributions

We used the catalog by Brusa et al. (2010) which includes multi-wavelength data from mid-infrared to hard X-rays: MIPS 160 μm , 70 μm and 24 μm GO3 data (Le Flocc’h et al. 2009), IRAC flux densities (Sanders et al. 2007), u^* and i^* CFHT bands and near-infrared K_{S} -band data (McCracken et al. 2008), J UKIRT (Capak et al. 2008), HST/ACS F814W imaging of the COSMOS field (Koekemoer et al. 2007), optical multiband photometry (SDSS, Subaru, Capak et al. 2007) and near- and far-ultraviolet bands with GALEX (Zamojski et al. 2007).

The number of X-ray sources detected at 160 μm and 70 μm is 18 and 42, respectively. For the undetected sources in these bands we consider 5σ upper limits of 65 mJy and 8.5 mJy for 160 μm and 70 μm , respectively. At 24 μm the number of detected sources is 237. For the 18 undetected sources at 24 μm , we consider 5σ upper limits of 80 μJy . All 255 sources are detected in the infrared in all IRAC bands, and only a few objects were not detected in the optical and near-IR bands: we have only 8 upper limits in the z^+ band; 1 upper limit in the B_J and i^* bands; 2 upper limits in the u^* band; 4 upper limits in the K_{S} CFHT band and 2 in the J UKIRT band. The observations in the various bands are not simultaneous, as they span a time interval of about 5 years: 2001 (SDSS), 2004 (Subaru and CFHT) and 2006 (IRAC). Variability for absorbed sources is likely to be a negligible effect, but, in order to further reduce it, we selected the bands closest in time to the IRAC observations (i.e., we excluded SDSS data, that in any case are less deep than other data available in similar bands). GALEX bands are not taken into account because, given the large aperture they can include light from close companions. All the data for the SED computation were shifted to the rest frame, so that no K-corrections were needed. Galactic reddening has been taken into account: we used the selective attenuation of the stellar continuum $k(\lambda)$ taken from Table 11 of Capak et al. (2007). Galactic extinction is estimated from Schlegel et al. (1998) for each object in the 971-XMM catalog. Count rates in the 0.5–2 keV and 2–10 keV are converted into monochromatic X-ray fluxes in the observed frame at 1 and 4 keV, respectively, using a Galactic column density

² The origin of spectroscopic redshifts for the 255 sources is as follows: 11 objects from the SDSS archive, 2 from MMT observations (Prescott et al. 2006), 70 from the IMACS observation campaign (Trump et al. 2007), 156 from the zCOSMOS bright 20k sample (see Lilly et al. 2007), 7 from the zCOSMOS faint catalog and 9 from the Keck/DEIMOS campaign.

³ $\text{Log}(X/O) = \text{Log} f_x + i^*/2.5 + 5.6$.

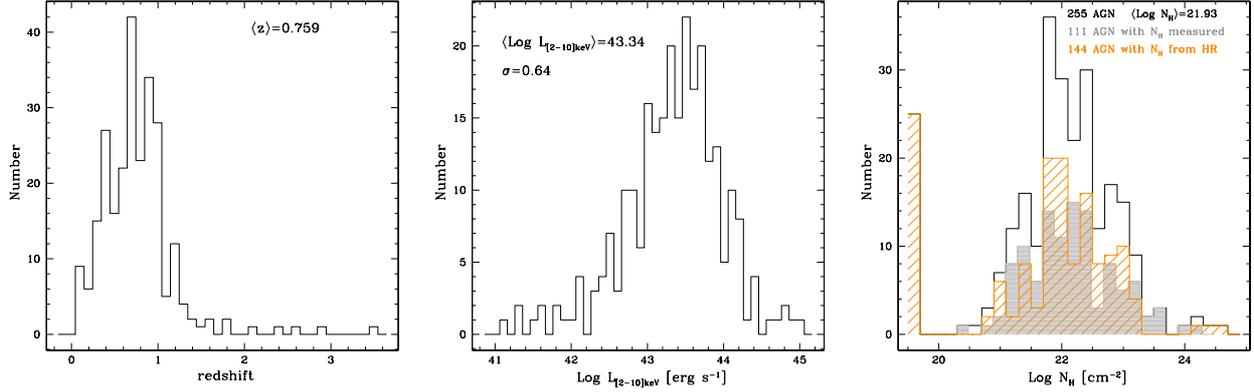


Fig. 2. *Left panel:* redshift distribution of the 255 type-2 AGN considered in this work. *Center panel:* intrinsic hard X-ray luminosity distribution of the 255 type-2 AGN considered in this work. *Right panel:* column density distribution of the 255 type-2 AGN (*black histogram*), of the 111 type-2 AGN with an N_{H} estimate from spectral analysis (*grey filled histogram*), and of the 144 type-2 AGN with an N_{H} estimate from hardness ratios (*orange hatched histogram*).

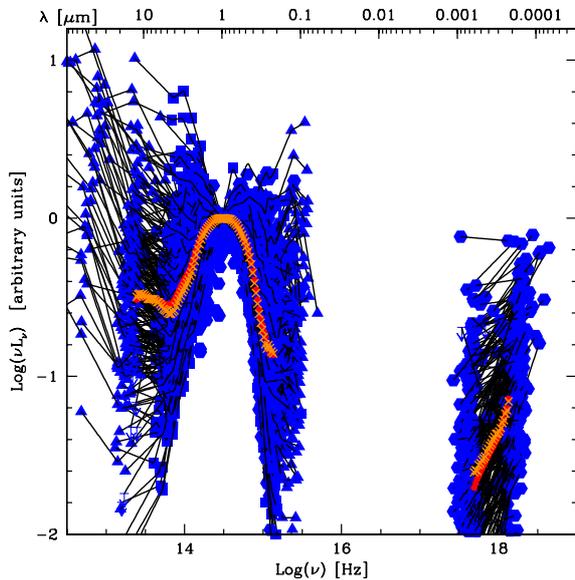


Fig. 3. Mean (*orange crosses*) and median (*red points*) SED from the total sample of 255 type-2 AGN. The blue points represent the rest-frame data, from infrared to X-ray, used to construct the average SED, while the black lines represent the interpolated SED.

$N_{\text{H}} = 2.5 \times 10^{20} \text{ cm}^{-2}$ (see Cappelluti et al. 2009), and assuming a photon index $\Gamma_x = 2$ and $\Gamma_x = 1.7$, for the soft and hard band, respectively. We do not correct these X-ray fluxes for the intrinsic column density. All sources are detected in the 2–10 keV band by definition of the sample, while in the soft band we have 70 upper limits.

4. Average SED

We have computed the individual rest-frame SEDs for all sources in the sample, following the same approach as in L10, and we have normalized all of them at $1 \mu\text{m}$. After this normalization we divided the frequency interval from $\text{Log } \nu_{\text{min}}$ to $\text{Log } \nu_{\text{max}}$ using a fixed step $\Delta \text{Log } \nu$. The minimum and maximum frequency depends on both the data and the redshift of the source considered to compute the SED, in our case we have used $\text{Log } \nu_{\text{min}} = 12 \text{ Hz}$ and $\text{Log } \nu_{\text{max}} = 20 \text{ Hz}$, with

a $\Delta \text{Log } \nu = 0.02$. We averaged data in each given interval $\text{Log } \nu_l \leq \text{Log } \nu \leq \text{Log } \nu_{l+1}$. The mean and median SEDs are obtained by taking the arithmetic mean and the median of logarithmic luminosities, $\text{Log } L$, in each bin. It is important to note that sources at different redshift contribute to the same bin. Because of the relatively wide range of redshifts, the lowest and the highest frequency bins are populated by a variable number of points. This effect may introduce relatively high fluctuations in the average luminosity in those bins. In order to minimize these effects we select a minimum number of 200 SEDs in each bin (this number depends on the total number of sources in the sample). Then, we select the mean reference frequency, $\overline{\text{Log } \nu}$, of the bin and use a binary-search algorithm to find all luminosities that correspond at $\overline{\text{Log } \nu}$ (if a source does not have a frequency that correspond to $\overline{\text{Log } \nu}$, we choose the luminosity with the closer frequency to $\overline{\text{Log } \nu}$). Finally, all adjacent luminosities in each bin are then connected to compute the final mean and median SED. In Fig. 3 the resulting mean and median SEDs are reported with orange crosses and red points, respectively. The data points are reported in order to show the dispersion with respect to $1 \mu\text{m}$. The average SED is characterized by a flat X-ray slope, $\langle \Gamma = 1.12 \rangle$, while in the optical-UV the observed emission appears to be consistent with the host-galaxy. The flattening of the X-ray slope is likely due to the fact that we do not correct for the intrinsic absorption the fluxes at 1 and 4 keV. The average SED in the mid-infrared is probably a combination of dust emission from star-forming region and AGN emission reprocessed by the dust.

Before trying to deconvolve each source using an SED-fitting code, we binned the total sample in X-ray and infrared luminosities and redshift. We used the luminosity at 4 keV and $8 \mu\text{m}$ to divide the total sample in 6 bins with the same number of sources in each bin. The wavelength of the luminosity used to bin the sample is chosen to minimize the host-galaxy contribution. In the three panels of Fig. 4 the resulting mean SEDs are shown. The redshift trend is directly connected to the luminosity trend, since at higher redshifts we are looking at the more luminous sources. The shapes of the average SEDs in the optical bands are approximately the same in all luminosity and redshift bins. As expected, there is a stronger host-galaxy contribution at lower luminosity/redshift bins, where the average SEDs have a typical galaxy shape. Moreover, there is a trend between X-ray and mid-infrared luminosity: the contribution in the infrared is higher at

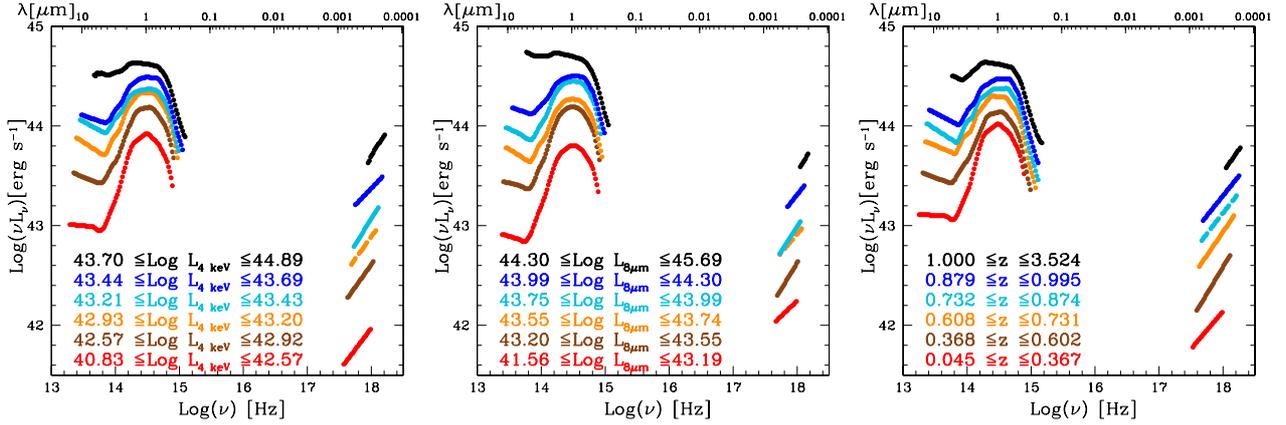


Fig. 4. Average SEDs in the rest-frame $\text{Log}(\nu L_\nu) - \text{Log} \nu$ plane. *Right panel:* mean SEDs computed binning in X-ray luminosity at 4 keV. *Center panel:* mean SEDs computed binning in infrared luminosity at $8 \mu\text{m}$. *Left panel:* mean SEDs computed binning in redshift. The color code refers to the different bins as labeled.

higher X-ray luminosities. This effect is already known for both type-1 and type-2 AGN using the intrinsic (non-stellar) emission from the AGN (e.g., Lutz et al. 2004; Gandhi et al. 2009). The observed average SED for our sample is a combination of both host-galaxy and AGN light, but the change in the average shape in the mid-infrared as a function of the X-ray luminosity suggests that most luminous sources are probably dominated by the AGN emission at those wavelengths.

5. SED-fitting

The purpose of performing SED fitting is to properly disentangle the emission associated to stellar light from that due to accretion and constrain physical parameters. Since the relative contribution of the different components varies with wavelength, a proper decomposition can be obtained through an SED-fitting approach providing a robust estimate of the nuclear emission (bolometric luminosities and bolometric corrections) and its relation with the host-galaxy properties (mass, star formation rates, morphological classification). A well sampled SED is mandatory; in particular, far-infrared observations are fundamental to sample the star-formation activity, while mid-infrared observations are necessary to sample the region of the SED where most of the bolometric luminosity of obscured AGN is expected to be re-emitted. In Fig. 5 the broad-band SEDs of four XMM-Newton type-2 AGN are plotted as examples. The lower two panels are representative of a full SED with all detections from the far-infrared to the optical. Unfortunately, there are very few detections at 160 and $70 \mu\text{m}$, so that the more representative situation is shown in the upper left panel of Fig. 5. The three components adopted in the SED-fitting code, *starburst*, *AGN torus* and *host-galaxy* templates, are shown as a long-dashed line, solid line and dotted line, respectively. All the templates used in the SED-fitting code will be described in the following Sections. The red line represents the best-fit, while the black points represent the photometric data used in the code, from low to high frequency: MIPS-Spitzer ($160 \mu\text{m}$, $70 \mu\text{m}$ and $24 \mu\text{m}$ if available), 4 IRAC bands, K CFHT, J UKIRT, optical Subaru and CFHT bands.

The observed data points from infrared to optical are fitted with a combination of various SED templates (see Sect. 5.1) using a standard χ^2 minimization procedure

$$\chi^2 = \sum_{i=1}^{n_{\text{filters}}} \left[\frac{F_{\text{obs},i} - A \times F_{\text{gal},i} - B \times F_{\text{agn},i} - C \times F_{\text{ir},i}}{\sigma_i} \right]^2 \quad (1)$$

where $F_{\text{obs},i}$ and σ_i are the monochromatic observed flux and its error in the band i ; $F_{\text{gal},i}$, $F_{\text{agn},i}$ and $F_{\text{ir},i}$ are the monochromatic template fluxes for the host-galaxy, the AGN and the starburst component, respectively; A , B and C are the normalization constants for the host-galaxy, AGN and starburst component, respectively. The starburst component is used only when the source is detected at $160 \mu\text{m}$ and $70 \mu\text{m}$. Otherwise, a two components SED-fit is used. Sixteen is the maximum number of bands adopted in the SED-fitting (only detection are considered), namely: $160 \mu\text{m}$, $70 \mu\text{m}$, $24 \mu\text{m}$, $8.0 \mu\text{m}$, $5.8 \mu\text{m}$, $4.5 \mu\text{m}$, $3.6 \mu\text{m}$, K_S , J , z^+ , i^* , r^+ , g^+ , V_J , B_J and u^* .

5.1. Template libraries

5.1.1. Optical template library

We used a set of 75 galaxy templates built from the Bruzual & Charlot (2003, BC03 hereafter) code for spectral synthesis models, using the version with the ‘‘Padova 1994’’ tracks, solar metallicity and Chabrier IMF (Chabrier 2003). For the purposes of this analysis a set of galaxy templates representative of the entire galaxy population from ellipticals to starbursts is selected. To this aim, 10 exponentially decaying star formation histories with characteristic times ranging from $\tau = 0.1$ to 30 Gyr and a model with constant star formation are included. For each SFH, a subsample of ages available in BC03 models is selected, to avoid both degeneracy among parameters and speed up the computation. In particular, early-type galaxies, characterised by a small amount of ongoing star formation, are represented by models with values of τ smaller than 1 Gyr and ages larger than 2 Gyr, whereas more actively star forming galaxies are represented by models with longer values of τ and a wider range of ages from 0.1 to 10 Gyr. An additional constraint on the age is that, for each sources, the age has to be smaller than the age of the Universe at the redshift of the source. Each template is reddened according to the Calzetti et al. (2000) reddening law. The input value is $E(B - V)$, corresponding to a dust-screen model, with $F_o(\lambda) = F_i(\lambda)10^{-0.4E(B-V)k(\lambda)}$, where F_o and F_i are the observed and the intrinsic fluxes, respectively. The extinction at a wavelength λ is related to the colour excess $E(B - V)$ and to the reddening curve $k(\lambda)$ by $A_\lambda = k(\lambda)E(B - V) = k(\lambda)A_V/R_V$, with $R_V = 4.05$ for the Calzetti’s law. The $E(B - V)$ values range between 0 and 1 with a step of 0.05.

5.1.2. AGN template library

The nuclear SED templates are taken from [Silva et al. \(2004\)](#). They were constructed from a large sample of Seyfert galaxies selected from the literature for which clear signatures of non-stellar nuclear emission were detected in the near-IR and mid-IR. After a proper subtraction of the stellar contribution, the nuclear infrared data were interpolated with a radiative transfer code for dust heated by a nuclear source with a typical AGN spectrum, and including different geometries, dust distribution, variation of the radii, density and dust grain sizes to account for possible deviations from a standard ISM extinction curve (see for more details [Granato & Danese 1994](#); [Maiolino et al. 2001](#)).

The infrared SEDs were then normalized by the intrinsic, unabsorbed X-ray flux in the 2–10 keV band, and are divided into 4 intervals of absorption: $N_{\text{H}} < 10^{22} \text{ cm}^{-2}$ for Sy1, while $10^{22} < N_{\text{H}} < 10^{23} \text{ cm}^{-2}$, $10^{23} < N_{\text{H}} < 10^{24} \text{ cm}^{-2}$ and $N_{\text{H}} > 10^{24} \text{ cm}^{-2}$ for Sy2 (see Fig. 1 in [Silva et al. 2004](#)). The main differences between the SEDs of Sy1s and Sy2s with $10^{22} < N_{\text{H}} < 10^{23} \text{ cm}^{-2}$ are the absorption in the near-IR at about $\lambda < 2 \mu\text{m}$ and the silicate absorption at $\lambda = 9.7 \mu\text{m}$, which are present in the Sy2 template. The shape of the SED in the mid-infrared with $10^{23} < N_{\text{H}} < 10^{24} \text{ cm}^{-2}$ is quite similar to that with $10^{22} < N_{\text{H}} < 10^{23} \text{ cm}^{-2}$. The Compton-thick SED ($N_{\text{H}} > 10^{24} \text{ cm}^{-2}$) shows conspicuous absorption also at $\lambda \sim 1.3 \mu\text{m}$. If a source has the N_{H} value available, this is used as a prior in the selection of the best-fit AGN template.

5.1.3. Starburst template library

We used two different starburst template libraries for the SED-fitting: [Chary & Elbaz \(2001\)](#) and [Dale & Helou \(2002\)](#). These template libraries represent a wide range of SED shapes and luminosities and are widely used in the literature. Here, we briefly describe how each of these libraries was derived and discuss the main differences between them.

The [Chary & Elbaz \(2001\)](#) template library consists of 105 templates based on the SEDs of four prototypical starburst galaxies (Arp220 (ULIRG); NGC 6090 (LIRG); M 82 (starburst); and M51 (normal star-forming galaxy)). They were derived using the [Silva et al. \(1998\)](#) models with the mid-infrared region replaced with ISOCAM observations between 3 and $18 \mu\text{m}$ (verifying that the observed values of these four galaxies were reproduced by the templates). These templates were then divided into two portions (4–20 μm and 20–1000 μm) and interpolated between the four to generate a set of libraries of varying shapes and luminosities. The [Dale et al. \(2001\)](#) templates are also included in this set to extend the range of shapes.

The [Dale & Helou \(2002\)](#) templates are updated versions of the [Dale et al. \(2001\)](#) templates. This model involves three components, large dust grains in thermal equilibrium, small grains semistochastically heated, and stochastically heated PAHs. They are based on IRAS/ISO observations of 69 normal star-forming galaxies in the wavelength range 3–100 μm . [Dale & Helou \(2002\)](#) improved upon these models at longer wavelengths using SCUBA observations of 114 galaxies from the Bright Galaxy Sample (BGS, see [Soifer et al. 1989](#)), 228 galaxies observed with ISOLWS (52–170 μm ; [Brauer 2002](#)), and 170 μm observations for 115 galaxies from the ISOPHOT Serendipity Survey ([Stickel et al. 2000](#)). All together, these 64 templates span the IR luminosity range 10^8 – $10^{12} L_{\odot}$. The total infrared template sample used in our analysis is composed of 168 templates.

6. Bolometric luminosities and bolometric corrections

The nuclear bolometric luminosities and bolometric corrections are estimated, using an approach similar to [Pozzi et al. \(2007\)](#); see also [Vasudevan et al. 2010](#); [Pozzi et al. 2010](#)), whereas the infrared luminosity is used as a proxy of the intrinsic nuclear luminosity. The appropriate nuclear template from [Silva et al. \(2004\)](#) is selected based on the absorbing column density N_{H} , when available, or from the best-fit nuclear infrared template. In order to compute the hard X-ray bolometric correction we used the standard definition

$$k_{\text{bol}} = \frac{L_{\text{bol}}}{L_{[2-10]\text{keV}}} \quad (2)$$

where the $L_{[2-10]\text{keV}}$ is the intrinsic X-ray luminosity and the bolometric luminosity is computed as the sum of the total infrared and X-ray luminosity

$$L_{\text{bol}} = L_{\text{IR}} + L_{\text{X}}. \quad (3)$$

After performing the SED-fitting, only the nuclear component of the best-fit is integrated. Hence, the total IR luminosity L_{IR} is obtained integrating the nuclear template between 1 and 1000 μm . To convert this IR luminosity into the nuclear accretion disk luminosity, we applied the correction factors to account for the torus geometry and the anisotropy (see [Pozzi et al. 2007](#)). The first correction is parameterized by the covering factor f . The covering factor is related to the geometry of the torus that obscures the accretion disk emission in the optical-UV along the line of sight, and its value is estimated from the ratio of obscured/unobscured quasars found by the X-ray background synthesis models ([Gilli et al. 2007](#)). This correction factor is ~ 1.5 . This value correspond to a typical covering factor of $f \sim 0.67$, consistent with the results based on clumpy torus models ([Nenkova et al. 2008](#)).

The anisotropy factor is defined as the ratio of the luminosity of face-on versus edge-on AGN, where the obscuration is a function of the column density N_{H} . Therefore, SEDs in [Silva et al. \(2004\)](#) have been integrated in the 1–30 μm range, after normalizing these SEDs to the same luminosity in the 30–100 μm range. The derived anisotropy values are 1.2–1.3 for $10^{22} < N_{\text{H}} < 10^{24}$ and 3–4 for $N_{\text{H}} > 10^{24}$. The same values as in [Vasudevan et al. \(2010\)](#) are adopted: 1.3 for $10^{22} < N_{\text{H}} < 10^{24}$ and 3.5 for $N_{\text{H}} > 10^{24}$.

The total X-ray luminosity L_{X} is estimated integrating in the 0.5–100 keV range the X-ray SED. We have interpolated the de-absorbed soft and hard X-ray luminosities. Since we are integrating at the rest-frame frequencies, the X-ray SED is extrapolated at higher and lower energies using the estimated X-ray slope, and introducing an exponential cut-off at 200 keV ([Gilli et al. 2007](#); see also Sect. 3.1 in L10).

7. Robustness of the method

The robustness of the method used to estimate nuclear bolometric luminosities and bolometric corrections from SED-fitting, for the sample of type-2 AGN, has been tested against the updated soft X-ray selected sample of type-1 AGN discussed in L10. The type-1 AGN sample in the L10 work was composed of 361 spectroscopically classified broad-line AGN. The recent work by [Brusa et al. \(2010\)](#) has updated the spectroscopic classification and increased the number of type-1 AGN with spectroscopic redshift, so that the final sample is composed of 395 type-1 AGN in

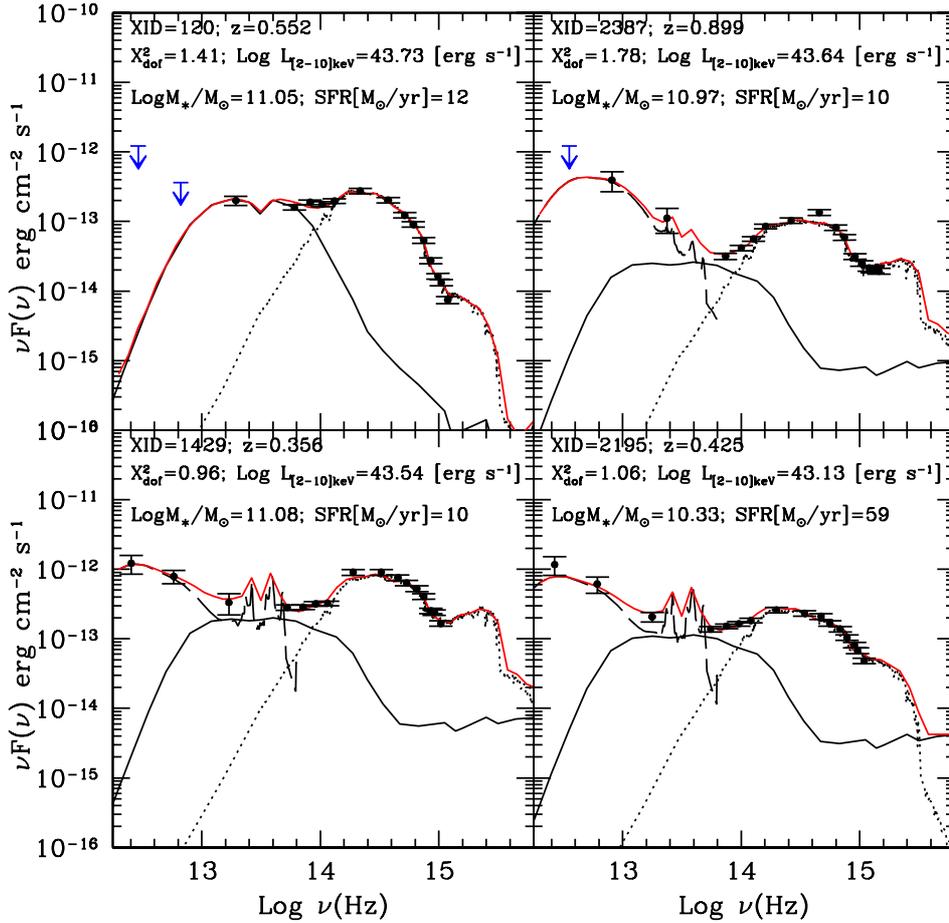


Fig. 5. Examples of SED decompositions. Black circles are the observed photometry in the rest-frame (from the far-infrared to the optical-UV). The long-dashed, solid and dotted lines correspond respectively to the starburst, AGN and host-galaxy templates found as the best fit solution. The red line represents the best-fit SED. The stellar mass and the SFR derived from the galaxy template are reported.

the redshift range $0.103 \leq z \leq 4.255$ with X-ray luminosities $42.20 \leq \text{Log } L_{[2-10]\text{keV}} \leq 45.23$. We have computed bolometric and X-ray luminosities, and bolometric corrections using the same approach as in L10 for the type-1 sample: bolometric luminosities are computed by integrating the rest-frame SEDs from $1 \mu\text{m}$ up to the UV-bump. In order to compare these estimates with the results from the SED-fitting code, we have applied to the same sample the method described in Sects. 5 and 6 to estimate bolometric parameters. To be consistent with the selection criteria of the sample discussed in this paper, we have considered only AGN with X-ray detection in the hard band, removing from the main sample 87 type-1 AGN with an upper limit at 2–10 keV. Moreover, for 2 sources the best-fit does not consider an AGN component, so we cannot compute the bolometric luminosities for them. The final test sample is composed of 306 type-1 AGN in the redshift range $0.103 \leq z \leq 3.626$ and X-ray luminosities $42.20 \leq \text{Log } L_{[2-10]\text{keV}} \leq 45.04$.

In order to select the appropriate nuclear template from [Silva et al. \(2004\)](#), we consider the SED for Sy1 AGN (no correction for anisotropy is necessary in this case) and the Sy2 SED with $10^{22} < N_{\text{H}} < 10^{23}$ for 20 AGN that have N_{H} in this range.

We present the comparison between the values of L_{bol} and k_{bol} from L10 and this work in Fig. 6. The outlier in the bottom side of the plot, XID = 357 at redshift 2.151 has $\text{Log } k_{\text{bol}} = 1.95$ from L10 and $\text{Log } k_{\text{bol}} = 1.04$ using the new approach, and

presents large error bars in the $24 \mu\text{m}$ detection, so that the total bolometric luminosity, computed using the infrared luminosity, is probably underestimated. The outlier in the right end of the distribution, XID = 5114 at redshift 0.212 ($\text{Log } k_{\text{bol}} = 2.82$ from L10 and $\text{Log } k_{\text{bol}} = 1.95$ using the new approach) has detections at 160, 70 and $24 \mu\text{m}$, $\text{Log } N_{\text{H}} = 22.68$ and $\text{Log } L_{[2-10]\text{keV}} = 42.89$. Probably this source is a star-forming galaxy, so that using the L10 approach we included stellar emission in the estimate of the nuclear bolometric luminosity, thus overestimating the nuclear bolometric luminosity and, therefore, the bolometric correction. The last notable outlier in the top/left side of the distribution, XID = 2152 at redshift 0.627 ($\text{Log } k_{\text{bol}} = 1.35$ from L10 and $\text{Log } k_{\text{bol}} = 2.27$ using the new approach) presents a significant host-galaxy contribution in the optical-UV and, therefore, the bolometric luminosity is likely to be underestimated in the L10 approach.

Although the two methods are very different, the bolometric luminosity estimates agree remarkably well, with a 1σ dispersion of 0.20 dex after performing a 3.5σ clipping method in order to avoid outliers. Bolometric luminosities from SED-fitting are on average slightly larger than those computed integrating the rest-frame SED from $1 \mu\text{m}$ to the X-ray (see the lower left side in Fig. 6). This effect is also present in the [Vasudevan et al. \(2010\)](#) work. A possible explanation is that SED-fitting underestimates the host-galaxy contribution, or that the anisotropy and

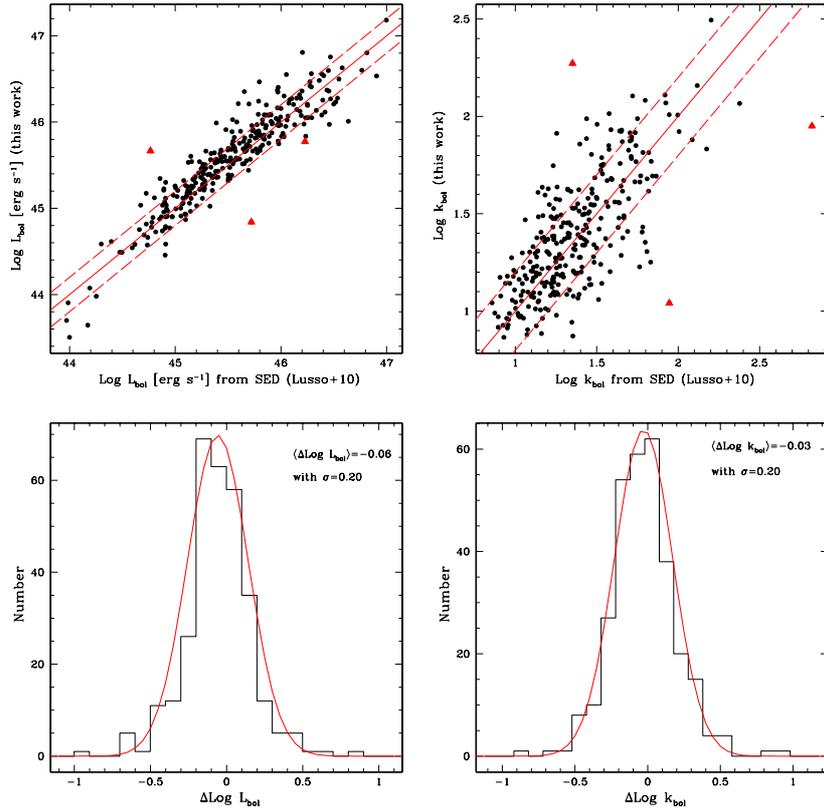


Fig. 6. *Upper panel:* comparison between the values of bolometric luminosity and bolometric correction from data presented in L10 and from this work. The three red triangles mark the outliers discussed in Sect. 7. *Lower panel:* distribution of the differences between the values of bolometric luminosity and bolometric correction from data presented in L10 and from this work.

geometry corrections are too large for some objects. The agreement between the two methods is overall quite satisfactory and in the following we will discuss our findings for the type-2 sample.

8. Results and discussion

8.1. Bolometric correction and luminosities for type-2 AGN

Bolometric luminosities and bolometric corrections have been computed for the type-2 AGN sample. Intrinsic soft and hard X-ray luminosities are estimated as described in Sect. 2.1. For 15 sources we do not have an estimate of the AGN component from the SED-fitting, and we cannot compute the bolometric luminosity for them. In Fig. 7 the distribution of the bolometric correction for the 240 type-2 AGN sample and for the 306 type-1 AGN (k_{bol} for type-1 AGN are computed using the SED-fitting code) are presented. Figure 8 shows bolometric corrections for both the type-1 and the type-2 AGN samples as a function of the hard X-ray luminosity. For both samples, bolometric parameters are estimated from the SED-fitting as discussed in Sect. 5. The green and orange curves represent the bolometric corrections and their 1σ dispersion as derived by Hopkins et al. (2007) and Marconi et al. (2004), respectively. type-2 AGN have, on average, smaller bolometric corrections than type-1 AGN at comparable hard X-ray luminosity. For example, at $43.30 \leq \text{Log } L_{[2-10]\text{keV}} \leq 44.30$ (vertical lines in Fig. 8), where both AGN types are well represented, the median bolometric correction for the type-2 AGN (134 objects) is $\langle k_{\text{bol}} \rangle \sim 13 \pm 1$, to be compared with a median bolometric correction $\langle k_{\text{bol}} \rangle \sim 23 \pm 1$ for the type-1 AGN (167 objects). The two

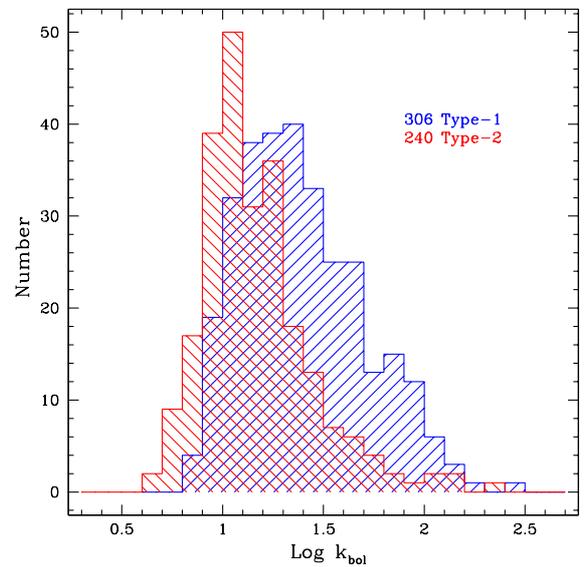


Fig. 7. Distribution of the bolometric correction for the 240 type-2 AGN sample (red hatched histogram) and for the 306 type-1 AGN (blue hatched histogram).

averages are statistically different at the $\sim 7\sigma$ level and this is consistent with the results in Vasudevan et al. (2010). The mean $L_{[2-10]\text{keV}}$ for the type-1 and type-2 AGN within this luminosity range differs by a factor 1.8, and this could in principle explain at least part of the difference in the average bolometric corrections

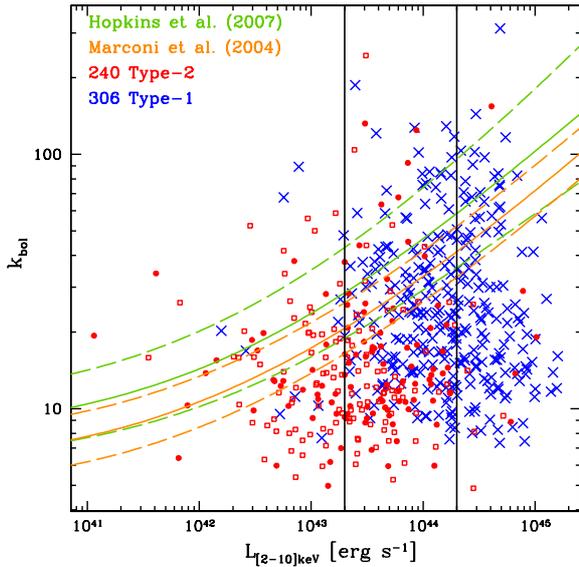


Fig. 8. Hard X-ray bolometric correction against the intrinsic 2–10 keV luminosity for 240 type-2 AGN with AGN best-fit (*red data*). The crosses represent the bolometric correction for 306 type-1 AGN, computed with the approach described in Sect. 6. The green and blue lines represent the bolometric correction and the 1σ dispersion obtained by Hopkins et al. (2007) and Marconi et al. (2004), respectively. The red points and open squares represent the 111 type-2 AGN with N_{H} from spectral analyses and the 144 type-2 AGN with N_{H} from HR, respectively.

for the two samples of AGN. However, the significance of the difference is still present if we split this luminosity range in two equal $\text{Log } L_{[2-10]\text{keV}}$ bins and perform a Kolmogorov-Smirnov test for the type-1 and type-2 AGN luminosity distributions in each bin.

Vasudevan & Fabian (2009) and Lusso et al. (2010) have shown that hard X-ray bolometric corrections are correlated with the Eddington ratios ($\lambda_{\text{Edd}} = L_{\text{bol}}/L_{\text{Edd}}$) for type-1 AGN (see also Marconi et al. 2004; Kelly et al. 2008). The $k_{\text{bol}} - \lambda_{\text{Edd}}$ relation suggests that there is a connection between the broad-band emission, mostly in the optical-UV, and the Eddington ratio, which is directly linked to the ratio between mass accretion rate and Eddington accretion rate. A high λ_{Edd} corresponds to an enhanced optical-UV emission, which means a prominent big-blue bump and therefore a higher k_{bol} . The difference between the average bolometric corrections for type-1 and type-2 AGN could be due to lower mass accretion rates in type-2 AGN, assuming the same black hole mass distribution for the two AGN populations (see Trump et al. 2011). The current theoretical framework of AGN/host-galaxy co-evolution predicts that obscured AGN are highly accreting objects and their black hole is rapidly growing. However, we note that this is true for $z = 1-3$ (see Hasinger 2008), while the majority of type-2 AGN in the present sample are relatively luminous ($42.37 \leq \text{Log } L_{\text{bol}} \leq 46.80$), and at moderately low redshift ($0 < z < 1$). Therefore, the type-2 AGN sample is likely to represent a later stage in AGN evolution history.

8.2. Infrared emission: indication of AGN activity

The re-processed infrared emission can be used as a proxy of the average disc emission, since the timescale for transfer of energy from the disk to the outer edge of the torus into infrared emission

is of the order of several years in standard AGN picture; whereas optical, UV and X-ray variability in AGN is known to occur on shorter timescales. The correlation between the 2–10 keV X-ray emission and IR emission at $12.3 \mu\text{m}$ for a sample of Seyfert nuclei has been discussed in Gandhi et al. (2009), and it could be used to estimate the intrinsic AGN power. Using X-ray data from the literature and new IR data from the Very Large Telescope’s Imager and Spectrometer for mid-Infrared (VISIR), taken specifically for addressing the issue of nuclear emission in local Seyferts, they found a tight correlation between intrinsic, uncontaminated IR luminosity and X-ray luminosity in the 2–10 keV range

$$\text{Log } \frac{L_{12.3 \mu\text{m}}}{10^{43}} = (0.19 \pm 0.05) + (1.11 \pm 0.07) \text{Log } \frac{L_{[2-10]\text{keV}}}{10^{43}}. \quad (4)$$

The relation is characterized by a small scatter with a standard deviation of 0.23 dex. The expected nuclear mid-infrared luminosity is computed from Eq. (4) using the estimate of the intrinsic unabsorbed X-ray luminosity. From the observed rest-frame SED (AGN+host-galaxy) the luminosity at $12.3 \mu\text{m}$ is computed. A comparison of the total observed luminosity at $12.3 \mu\text{m}$ and that predicted by Eq. (4) is plotted in Fig. 9 for four representative sources. In Fig. 10 the distribution of the ratio $r = \text{Log } (L_{12.3 \mu\text{m,obs}}/L_{12.3 \mu\text{m,predicted}})$ is plotted for the type-2 AGN sample. The distribution of the ratio r has a mean which is shifted from zero by ~ 0.2 . However, if we consider a gaussian distribution centered at $r = 0$ with $\sigma = 0.23$, i.e. the same dispersion observed by Gandhi et al. (2009) in their local sample, the majority of the objects are found within 2σ of the r distribution. The tail outside 2σ and extending to high r includes 73 sources (with $r \geq 0.5$) for which the predicted mid-infrared luminosity is significantly lower than observed. The hard X-ray luminosities of these 73 AGN are mainly in range $\text{Log } L_{[2-10]\text{keV}} \sim 42-44$, where the local correlation is well-sampled. There are two possible explanations for a significant ($\sim 30\%$ of the objects) tail toward high- r values: either the Gandhi relation, which was derived for a sample of local Seyfert galaxies, cannot be extended to all the sources in our sample or the SED-fitting procedure may overestimate, in a fraction of these objects, the nuclear contribution. In order to study the properties of these outliers, bolometric corrections, morphologies, stellar masses and SFR are discussed in following. We call “low- r ” AGN all sources within 2σ of the r distribution, while the “high- r ” AGN sample is populated by the sources deviating more than 2σ (see Fig. 9 for some examples).

A clear separation in bolometric corrections for these two sub-samples is found. This is shown in Fig. 11 in which bolometric corrections are plotted as a function of the 2–10 keV luminosity. At a given hard X-ray luminosity ($43 \leq \text{Log } L_{[2-10]\text{keV}} \leq 44$) the low- r sample has a median bolometric correction of $\langle k_{\text{bol}} \rangle \sim 11 \pm 1$ (110 objects), to be compared with a median bolometric correction for the high- r sample of $\langle k_{\text{bol}} \rangle \sim 26 \pm 3$ (44 objects). The two median values for k_{bol} are statistically different at the $\sim 5\sigma$ level.

Furthermore, in the high- r sample 24 type-2 AGN out of 73 have a detection at $70 \mu\text{m}$ ($\sim 33\%$, significantly higher than those for the low- r sample, $\sim 4\%$) and 9 of these 24 have also a detection at $160 \mu\text{m}$ ($\sim 12\%$ considering the total high- r sample, and only $\sim 1\%$ for the low- r sample). This denotes that the difference in the average bolometric corrections between the low- r and high- r samples is probably due to the fact that a significant fraction of the infrared emission is attributable to an incorrect modeling of the star-formation process, or the AGN contribution is somehow overestimated by the SED-fitting procedure.

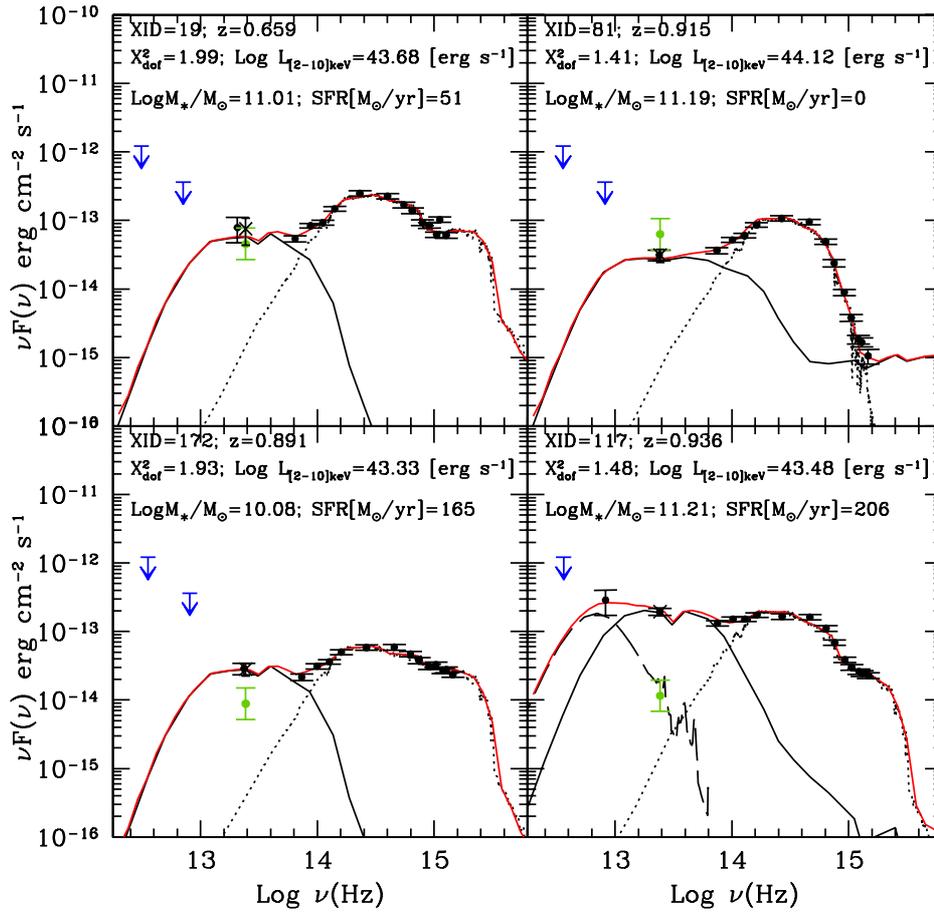


Fig. 9. Examples of SED decompositions. Black circles are the observed photometry in the rest-frame (from the far-infrared to the optical-UV). The long-dashed, solid and dotted lines correspond respectively to the starburst, AGN and host-galaxy templates found as the best fit solution. The red line represents the best-fit SED. The stellar mass and the SFR derived from the galaxy template are reported. The green point represents the nuclear mid-infrared luminosity using Eq. (4), while the cross represents the total observed luminosity at $12.3 \mu\text{m}$ computed from the rest-frame SED. XID = 19 and 81 are examples of low- r AGN, while XID = 172 and 117 represent high- r AGN.

There is no significant difference in the average nuclear absorption between the low- r and the high- r sample, while there is a possibly significant difference in SFR and stellar masses. The median stellar mass in the high- r sample is $\langle \text{Log } M_* \rangle \sim 10.93 M_\odot$ with a dispersion of 0.30, while for the low- r sample is $\langle \text{Log } M_* \rangle \sim 10.76 M_\odot$ with $\sigma = 0.30$. The two averages are statistically different at the $\sim 3\sigma$ level. The median SFR, as derived from the SED-fit, for the high- r sample is $\langle \text{SFR} \rangle \sim 17 \pm 3 M_\odot/\text{yrs}$ with a $\sigma = 0.30$, while for the low- r sample is $\langle \text{SFR} \rangle \sim 3 \pm 1 M_\odot/\text{yrs}$ with a $\sigma = 0.30$ and the two averages are statistically different at the 4.4σ level.

Overall, the SED-fitting for the 73 type-2 AGN is likely to overestimate the AGN emission in the infrared, which is probably due to the infrared emission from star-forming regions. The average bolometric correction for type-2 AGN, excluding these sources, would be even lower than what we have computed in the previous section. This reinforces the idea of lower bolometric corrections for type-2 AGN with respect to type-1 AGN. The low bolometric corrections for type-2 AGN could be also explained if a fraction of the accretion disk bolometric output is not re-emitted in the mid-infrared, but rather dissipated (e.g., by AGN-feedback). This would not be accounted in the bolometric luminosity, and could provide a plausible explanation for low k_{bol} especially if the low- r sample is considered. At this stage,

this is just a speculation, and more work is needed to verify this possibility.

8.3. Host-galaxy properties: M_* , SFR, colors and morphologies

Galaxies show a colour bi-modality both in the local Universe and at higher redshift (up to $z \sim 2$; e.g., [Strateva et al. 2001](#); [Bell et al. 2004](#)). This bi-modality (red-sequence and blue-cloud galaxies) has been interpreted as an evidence for a dichotomy in their star formation and merging histories (e.g., [Menci et al. 2005](#); but see also [Cardamone et al. 2010](#), for an alternative explanation). Color-magnitude and color-mass diagrams (e.g., rest-frame $(U - V)$ versus stellar mass) have been used as tools in galaxy evolution studies, and since many models invoke AGN feedback as an important player in such evolution, it is interesting to locate the hosts of type-2 AGN in those diagrams. Using the galaxy component obtained from the best fit of the type-2 AGN, it is possible to derive rest-frame colors for the host that, linked to the stellar mass and the morphology, can provide hints on AGN feedback. Several studies found that the hosts of obscured AGN tend to be redder than the overall galaxy population in the rest-frame $(U - V)$ color (e.g., [Nandra et al. 2007](#)). There are at least two possible and

Table 1. Properties of the type-2 AGN sample.

XID	Redshift	Log $L_{[2-10]\text{keV}}$ [erg s^{-1}]	Log L_{bol} [erg s^{-1}]	k_{bol}	Log M_* [M_{\odot}]	SFR [M_{\odot}/yrs]	M_U	M_V	M_J	Morphological class ^a
67	0.367	42.79	43.80	10.24	9.68	1.86	-18.63	-19.80	-20.72	0
65	0.979	43.83	44.85	10.62	10.19	38.60	-20.50	-21.63	-22.71	7
64	0.686	43.54	44.52	9.57	10.45	32.82	-20.17	-21.56	-23.05	10
63	0.355	42.98	44.28	19.88	10.79	4.94	-20.61	-22.13	-23.26	2
54	0.350	42.58	43.38	6.32	11.18	0.11	-20.57	-22.62	-23.78	12
45	0.121	41.90	42.91	10.31	9.39	0.00	-16.76	-18.70	-19.77	1
43	1.162	44.22	45.28	11.49	11.30	1.47	-21.55	-23.48	-24.61	2
19	0.659	43.67	44.76	12.44	11.07	35.88	-21.57	-22.88	-24.02	1
117	0.936	43.47	45.59	132.11	11.24	199.40	-21.81	-23.34	-24.80	23
116	0.874	43.49	44.50	10.17	10.56	3.77	-20.19	-21.81	-22.90	0
112	0.762	43.65	44.65	9.97	10.93	0.62	-20.00	-22.13	-23.52	3
104	0.623	44.08	45.11	10.75	10.79	0.45	-20.25	-22.18	-23.30	1
101	0.927	43.70	44.68	9.56	10.92	0.61	-21.19	-22.93	-23.78	0
100	0.270	42.61	43.54	8.59	10.21	0.00	-19.14	-20.98	-21.92	0
99	0.730	43.54	44.78	17.47	9.97	116.26	-21.61	-22.26	-23.15	11
85	1.001	43.46	44.84	23.99	10.13	38.52	-20.15	-21.34	-22.56	8
81	0.915	44.11	45.05	8.65	11.18	0.00	-20.35	-22.59	-24.07	2
70	0.688	44.00	45.60	39.74	10.65	542.60	-20.87	-22.32	-24.31	3
152	0.895	43.75	44.82	11.72	9.87	92.88	-21.06	-21.82	-22.85	7
150	0.740	43.29	44.25	9.27	10.64	0.55	-19.38	-21.28	-22.50	3

Notes. This table is presented entirely in the electronic edition; a portion is shown here for guidance. ^(a) The morphological classification of the type-2 AGN hosts is coded from 0 to 23: 0 = elliptical; 1 = S0; 2 = bulge-dominated; 3 = intermediate-bulge; 4 = disk-dominated; 5 = irregular; 6 = compact/irregular; 7 = compact; 8 = unresolved/compact; 9 = blended; 10 = bulge-dominated/close-companion; 11 = intermediate-bulge/close-companion; 12 = S0/close-companion; 23 = possible mergers.

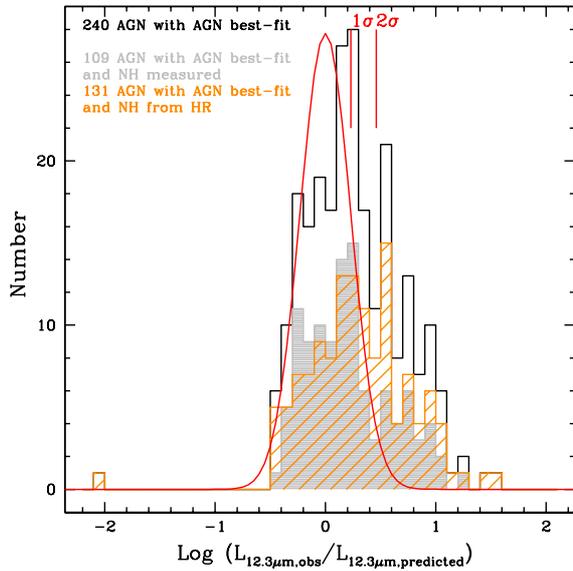


Fig. 10. Histogram of the ratio between the total observed luminosity at $12.3 \mu\text{m}$ and the mid-infrared luminosity predicted by Eq. (4). The red curve represents a gaussian with mean equal to zero and standard deviation 0.23. The 1σ and 2σ standard deviations of the correlation are also reported.

significantly different interpretations for this observational result: the observed red colors are mainly due to dust extinction, so that a significant fraction of obscured AGN would live in massive, dusty star-forming galaxies with red optical colors (e.g., Brusa et al. 2009); or red sources are linked with passive systems (e.g., Rovilos & Georgantopoulos 2007; Schawinski et al. 2009; Cardamone et al. 2010). Therefore, accurate

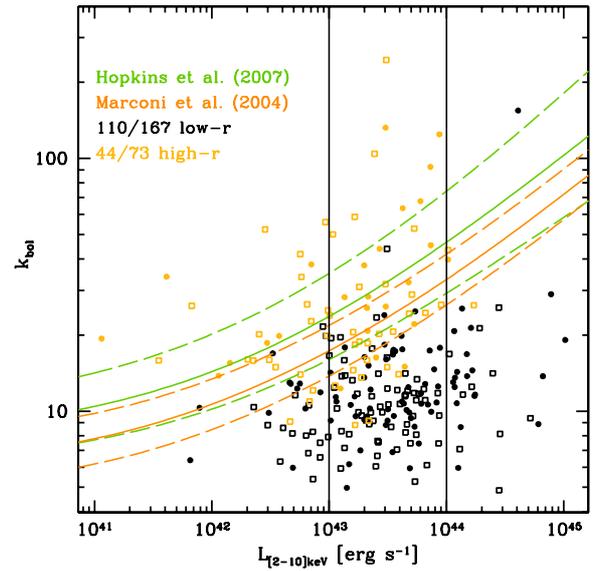


Fig. 11. Hard X-ray bolometric correction against 2–10 keV luminosity for 240 type-2 AGN with AGN best-fit. The 240 type-2 sample is divided into subsamples: low- r AGN sample (*black data*) and high- r AGN sample (*yellow data*). The green and orange lines represent the bolometric correction and 1σ dispersion obtained by Hopkins et al. (2007) and Marconi et al. (2004), respectively. In the de-absorbed hard X-ray luminosity range highlighted by the solid lines, we have 167 low- r and 73 high- r (in the infrared) sources.

stellar mass and SFR estimates, together with detailed galaxy morphologies, are of particular importance to discriminate between the two alternative possibilities.

The very high resolution and sensitivity of ACS-HST imaging in the COSMOS survey provides resolved morphologies

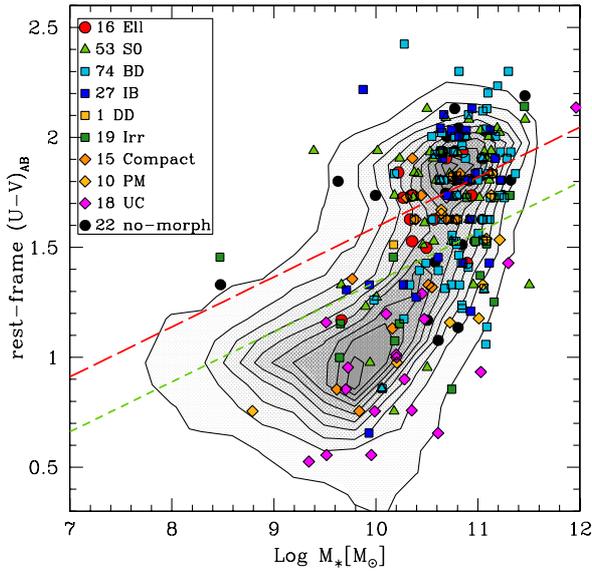


Fig. 12. The morphology distribution (using the ZEST+ code) of the 233 AGN host-galaxies on the $(U - V)$ colour-mass diagram. We also plotted the 22 sources without morphological information. The $(U - V)$ color and stellar masses are computed using the SED-fitting code. We overplot the contours of about 8700 galaxies in z COSMOS (colours and stellar masses from *Hyperz*). The morphology classification is labeled as follows: elliptical (Ell), S0, bulge-dominated galaxy (BD), intermediate-bulge galaxy (IB), disk-dominated galaxy (DD), irregular (Irr), Compact, possible mergers (PM) and unresolved compact (UC). The red dashed line represents the red sequence cut defined by [Borch et al. \(2006\)](#), while the green short dashed line defines an approximate green valley region, both lines are calculated at redshift ~ 0.76 , which is the average redshift of the main type-2 sample.

for several hundreds of thousands galaxies with $i_{\text{acs}} \leq 24$ (see [Scarlata et al. 2007](#), for details). Galaxy morphologies were obtained with an upgraded version of the Zurich Estimator of Structural Types (ZEST; [Scarlata et al. 2007](#)), known as ZEST+ ([Carollo et al., in prep.](#)). Relative to its predecessor, ZEST+ includes additional measurements of non-parametric morphological indices for characterising both structures and substructures. For consistency with the earlier versions, ZEST+ uses a Principal Component Analysis (PCA) scheme in the 6-dimensional space of concentration, asymmetry, clumpiness, M_{20} (second-order moment of the brightest 20% of galaxy pixels), Gini coefficient, and ellipticity. ZEST+ classifies galaxies in seven morphological types located in specific regions of the 6-dimensional space: elliptical, S0, bulge-dominated disk, intermediate-bulge disk, disk-dominated, irregular, compact. The different types were then visually inspected. For 19 objects ZEST+ is unable to give any information on morphology because these sources lie off the edge of the ACS tiles and 4 sources are blended. As a result of the ZEST+ procedure and visual inspection of the other 233 galaxies in our sample, we find that 16 are ellipticals (Ell), 53 are S0s, 74 are bulge-dominated (BD) disks, 27 are intermediate-bulge (IB) disks, just 1 is disk-dominated (DD), 19 are irregular galaxies (Irr), 15 are compact galaxies (i.e. the structural parameters computed for these galaxies from the HST-ACS images are highly affected by the instrumental PSF) and 18 are unresolved compact galaxies (UC, i.e. essentially point-like sources). Ten galaxies show distortions and potential signatures of ongoing or recent mergers (PM). At the typical magnitudes of the objects in our sample, the ZEST+

classification is highly reliable for galaxies with redshift ≤ 1 . At higher redshifts, morphological k-correction and, to a lesser extent, resolution effects can adversely affect measurements of quantitative morphological parameters (note that 16% of the main type-2 AGN sample have $z > 1.0$). However, broad morphological bins (e.g., early/late type galaxies) should be relatively robust. For high- z galaxies, resolution might also have an impact on the classification for mergers, and ACS images could not be deep enough to distinguish merger features (see also [Mainieri et al. 2011](#)). Moreover, inclination might also affect the morphological classification (e.g., [Nair & Abraham 2010](#)). However, a detailed study of systematics and biases in the morphological classification is beyond the purposes of the present paper. In [Table 1](#) we list the main properties of the sample.

The rest-frame $(U - V)$ color encompasses the 4000 Å break, and it is particularly sensitive to age and metallicity variations of stellar population in galaxies (e.g., [Sandage & Visvanathan 1978](#); [Bell et al. 2004](#); [Borch et al. 2006](#); [Cardamone et al. 2010](#)). In [Fig. 12](#) the distribution of the rest-frame $(U - V)$ colors, which are computed directly from the best-fit galaxy template, and stellar masses (from the SED-fitting code) are reported for the entire type-2 AGN sample. In the same figure, the background contours for a sample of ~ 8700 galaxies in z COSMOS ($i_{\text{acs}} < 22.5$, 240 type-2 are detected in the i_{acs} band, 183/240 type-2 AGN (76%) have $i_{\text{acs}} < 22.5$) are also plotted, where colours and stellar masses are computed using the *Hyperz* code ([Bolzonella et al. 2000](#)).

AGN are known to reside in massive galaxies (e.g., [Silverman et al. 2009](#); [Brusa et al. 2009](#)) and this is fully confirmed by the present analysis. The morphologies of the host-galaxies and the stellar masses indicate that there is a preference for these type-2 AGN to be hosted in bulge-dominated and S0 galaxies ($\sim 50\%$) with stellar masses greater than $10^{10} M_{\odot}$. This result is consistent with the previous studies on type-2 AGN by [Silverman et al. \(2008\)](#), see also [Kauffmann et al. 2003](#); [Bundy et al. 2008](#); [Schawinski et al. 2011](#)).

It should be noted that no correction for the internal extinction has been applied to the $(U - V)$ colors of both background galaxies in z COSMOS and type-2 AGN hosts. This correction could be important as shown in [Cowie & Barger \(2008\)](#); see also [Cardamone et al. 2010](#)). In that work star formation and galactic stellar mass assembly are analyzed using a very large and highly spectroscopically complete sample selected in the rest-frame NIR bolometric flux in the GOODS-N. They found that applying extinction corrections is critical when analyzing galaxy colors; nearly all of the galaxies in the green valley are $24 \mu\text{m}$ sources, but after correcting for extinction, the bulk of the $24 \mu\text{m}$ sources lie in the blue cloud. This correction introduces an average shift in color of ~ 0.2 mag for the most extincted/star-forming galaxies. However, to be consistent with the colors of the background galaxies no correction for intrinsic extinction is considered.

AGN host-galaxies belong to the red-sequence if their $(U - V)$ color is above the threshold ([Borch et al. 2006](#)):

$$(U - V)_{\text{AB,rest-frame}} > 0.277 \text{Log } M_* - 0.352z - 0.39. \quad (5)$$

Sources in the green-valley are approximately defined shifting this relation by 0.25 downward towards bluer colors. With these definitions, $\sim 42\%$ (108/255) and $\sim 25\%$ (63/255) of the total sample are included in the red-cloud and the green-valley, respectively. For all sources the Specific Star-Formation Rate (SSFR) is estimated, defined as the ratio of the SFR per unit of galaxy stellar mass ($SSFR = SFR/M_*$). The inverse of the $SSFR$,

$SSFR^{-1}$, is called ‘‘growth time’’

$$SSFR^{-1} = M_*/\dot{M}_*, \quad (6)$$

and corresponds to the time required for the galaxy to double its stellar mass, assuming its SFR remained constant. Actively star-forming galaxies are defined as sources with growth time smaller than the age of the Universe at their redshift ($SSFR^{-1} < t_{\text{Hubble}}$), while sources with $SSFR^{-1}$ larger than the age of the Universe can be considered passive galaxies (see also Fontana et al. 2009; Brusa et al. 2009). Figure 13 shows $SSFR^{-1}$ as a function of the stellar mass in three different redshift bins for the AGN host-galaxies in the red-sequence, in the green-valley and in the blue-cloud and for the zCOSMOS galaxies in same redshift ranges. The horizontal lines mark the age of the Universe at the two redshift boundaries of the chosen intervals. At face value, almost all the sources in the red-sequence have $SSFR^{-1}$ larger than the age of the Universe at their redshift, which is consistent with passive galaxies. However, the value of $SSFR^{-1}$ has to be considered only as an approximate indication of the star-formation activity; in fact, there is some possible evidence of some residual star-formation, in red-cloud AGN host-galaxies, as witnessed by their morphologies. In the red-sequence 8 and 28 sources are classified as ellipticals and S0s, respectively; all together they represent 34% of the host-galaxy population in the red-sequence. About 42% is represented by disk galaxies (both bulge-dominated and intermediate-bulge), which are probably still forming stars but not at high rates. In fact, 15 over 108 sources ($\sim 14\%$) have a detection at $70 \mu\text{m}$ and 5 have also a detection at $160 \mu\text{m}$ ($\sim 6\%$).

For these objects, the SFR inferred from the far-infrared detections is significantly higher than the SFR derived from the SED-fitting procedure. Indeed, an SED-fitting over the UV, optical and near-infrared bands is not always able to discriminate between the red continua of passive galaxies and those of dusty star-forming galaxies. Therefore, we decided to include another indicator in the present analysis broadly following the procedure described in Cardamone et al. (2010, i.e., the $(U - V) - (V - J)$ color diagram). Near-infrared emission can distinguish between red-passive or dust-obscured galaxies: given a similar $0.5 \mu\text{m}$ flux, a star-forming galaxy has more emission near $\sim 1 \mu\text{m}$ than a passive galaxy. A sub-sample of galaxies is selected in the same redshift range explored by Cardamone et al. (2010), we find 92 AGN host-galaxies with $0.8 \leq z \leq 1.2$. Figure 14 shows both inactive galaxies and AGN host-galaxies in the same redshift range and the thresholds considered to divide galaxies in the red-sequence and in the green-valley (we consider an average redshift of 1 to define the threshold for the red-sequence and the green-valley). Thirty-five out of 92 AGN hosts are found to lie in the red-sequence ($\sim 38\%$) and 23 in the green-valley ($\sim 25\%$); while for inactive galaxies about 32% and 21% lie in the red-sequence and green-valley, respectively. In Fig. 15 the $(U - V) - (V - J)$ color diagram for the 92 type-2 AGN hosts is presented. From a preliminary analysis of the rest frame $(U - V)$ against the rest-frame $(V - J)$ color (see Fig. 15; but see also Fig. 2 in Cardamone et al. 2010), only $\sim 9\%$ of the AGN host-galaxies in the red-sequence and $\sim 30\%$ of AGN host-galaxies in the green-valley are moved in the region populated by dusty star-forming galaxies in the color-color diagram. To be compared with 20% AGN host-galaxies in the red-sequence and 75% AGN host-galaxies in the green-valley found by Cardamone and collaborators. The fractions of dust-obscured galaxies among the red-cloud and green-valley AGN in our sample, at $0.8 \leq z \leq 1.2$, are lower than those in the Cardamone et al. (2010) sample.

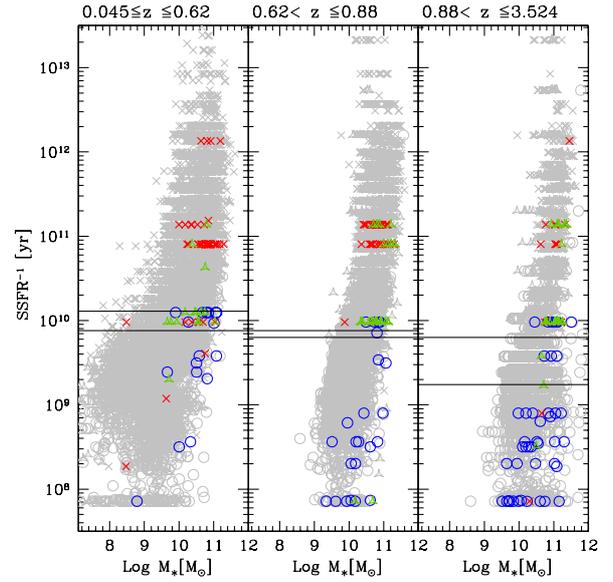


Fig. 13. Inverse of the SSFR rate as a function of the stellar mass of the AGN host-galaxies in three different redshift bins for the zCOSMOS galaxies and for the type-2 AGN sample in the red-sequence (red crosses), in the green-valley (green triangles) and in the blue-cloud (blue open circles). The horizontal lines mark the age of the Universe at the two redshift boundaries of the chosen intervals.

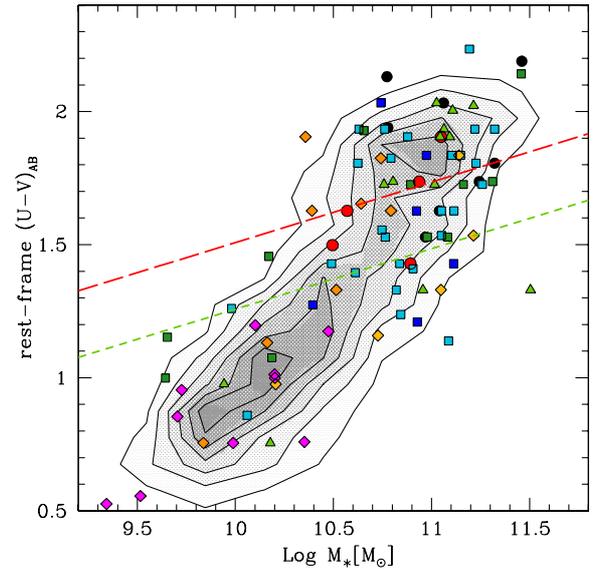


Fig. 14. Distribution of the stellar masses as a function of the rest-frame $(U - V)$ colors in the redshift range $0.8 \leq z \leq 1.2$. The red dashed line represents the red sequence cut defined by Borch et al. (2006), while the green short dashed line defines an approximate green valley region, both lines are calculated at redshift ~ 1 . The points are color coded as in Fig. 12.

However, the global fractions of AGN hosts, tentatively associated to passive galaxies, are very similar ($\sim 50\%$) in the two samples. The fractions for both AGN host-galaxies and inactive galaxies are reported in Table 2.

9. Summary and conclusions

A detailed analysis of the SEDs of 255 spectroscopically identified hard X-ray selected type-2 AGN from the XMM-COSMOS survey is presented. In obscured AGN, the optical-UV nuclear

Table 2. AGN hosts and galaxies properties.

Sample	N	Red-sequence	Green-valley		Blue-cloud	
$0.045 \leq z \leq 3.452$						
type-2 AGN	255	108 (42%)	96 (89%) P 12 (11%) D	63 (25%)	52 (82%) P 11 (18%) D	84 (33%)
Galaxies	8742	2535 (29%)	2306 (91%) P 229 (9%) D	1923 (22%)	1596 (83%) P 327 (17%) D	4284 (49%)
$0.8 \leq z \leq 1.2$						
type-2 AGN	92	35 (38%)	32 (91%) P 3 (9%) D	23 (25%)	16 (70%) P 6 (30%) D	34 (37%)
Galaxies	1836	587 (32%)	569 (97%) P 18 (3%) D	385 (21%)	269 (70%) P 116 (30%) D	864 (47%)

Notes. P = Passive, D = Dusty.

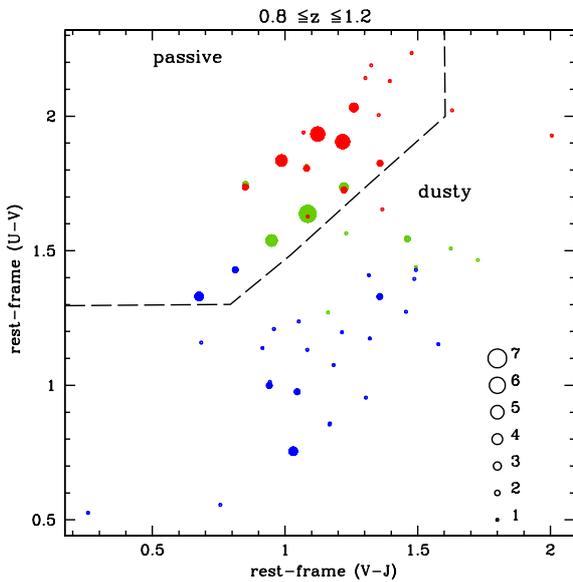


Fig. 15. Distribution of type-2 AGN hosts in the rest-frame ($U - V$) against the rest-frame ($V - J$) color. Color coded as in Fig. 13. Sources with the same best-fit galaxy template and the same extinction lie in the same position in the color-color diagram. Point size is keyed to the number of objects.

luminosity is intercepted along the line of sight by the dusty torus and reprocessed in the infrared, so that what we see in the optical-UV is mostly the light from the host-galaxy. On the one hand, this allows us to study the galaxy properties, on the other hand it makes difficult to estimate the nuclear bolometric power. An SED-fitting code has been developed with the main purpose of disentangling the various contributions (starburst, AGN, host-galaxy emission) in the observed SEDs using a standard χ^2 minimization procedure (the starburst component is only used in the case of detection at $70 \mu\text{m}$). The code is based on a large set of starburst templates from Chary & Elbaz (2001) and Dale et al. (2001), and galaxy templates from the Bruzual & Charlot (2003) code for spectral synthesis models, while AGN templates are taken from Silva et al. (2004). These templates represent a wide range of SED shapes and luminosities and are widely used in the literature. The total (nuclear) AGN bolometric luminosities are then estimated by adding the X-ray luminosities integrated over the 0.5–100 keV energy range to the infrared luminosity between 1 and $1000 \mu\text{m}$. The total X-ray luminosity is computed integrating the X-ray SED using the de-absorbed soft and hard

X-ray luminosities. The SED is extrapolated to higher energies using the observed X-ray slope, and introducing an exponential cut-off at 200 keV. The total infrared luminosity is evaluated integrating the infrared AGN best-fit and then converted into the nuclear accretion disk luminosity applying the appropriate correction factors to account for the geometry and the anisotropy of the torus emission. The reprocessed IR emission is considered to be a good proxy of the intrinsic disk emission and this is supported by previous investigations (Pozzi et al. 2007; Gandhi et al. 2009; Vasudevan et al. 2010). In the distribution of the ratio $r = \text{Log} (L_{12.3 \mu\text{m,obs}}/L_{12.3 \mu\text{m,predicted}}$; see Eq. (4)) the majority of the objects are within 2σ of the r distribution. The tail outside 2σ and extending to high r includes 73 sources (with $r \geq 0.5$) for which the predicted mid-infrared luminosity is significantly lower than observed. We call “low- r ” AGN all sources within 2σ of the r distribution, while the “high- r ” AGN sample is represented by the sources deviating more than 2σ .

Our main results are the following:

1. The average observed SED is characterized by a flat X-ray slope, ($\Gamma = 1.12$), as expected for obscured AGN (not corrected for absorption), while in the optical-UV the observed light appears to be consistent with the host-galaxy emission. The average SED in the mid-infrared is more likely a combination of dust emission from star-forming region and AGN emission reprocessed by the dust.
2. The full sample is split into four bins of different X-ray and infrared luminosities and redshift. The shapes of the average SEDs in the optical bands are approximately the same in all luminosity and redshift bins. There is a stronger host-galaxy contribution at lower luminosity/redshift bins, where the average SEDs have a typical galaxy shape. Moreover, there is a trend between X-ray and mid-infrared luminosity: the contribution of the AGN in the infrared (around 8–15 μm) is higher at higher X-ray luminosities.
3. type-2 AGN appear to have smaller bolometric corrections than type-1 AGN. At the same hard X-ray luminosity, $43.30 \leq \text{Log} L_{[2-10]\text{keV}} \leq 44.30$, where both samples are well represented, we find that the median bolometric correction for type-2 AGN (134 objects) is $\langle k_{\text{bol}} \rangle \sim 13 \pm 1$, to be compared with a median bolometric correction $\langle k_{\text{bol}} \rangle \sim 23 \pm 1$ for type-1 AGN (167 objects). The two averages are statistically different at the $\sim 7\sigma$ level.
4. A clear separation in bolometric corrections for the low- r and the high- r samples is found. The relation provided by Gandhi and collaborators is valid for the majority of objects, while for 30% of the sample SED-fitting procedure

may underestimate the non-nuclear contribution. At a given hard X-ray luminosity ($43 \leq \text{Log } L_{[2-10]\text{keV}} \leq 44$) the low- r sample has a median bolometric correction of $\langle k_{\text{bol}} \rangle \sim 11 \pm 1$ (110 objects), to be compared with a median bolometric correction for the high- r sample of $\langle k_{\text{bol}} \rangle \sim 26 \pm 3$ (44 objects). The two median values for k_{bol} are statistically different at the $\sim 5\sigma$ level.

5. Host-galaxies morphologies and the stellar masses indicate that type-2 AGN are preferentially hosted in galaxies which have a bulge, irrespective of the strength of the bulge or if the galaxy is on the red sequence or blue cloud, and with stellar masses greater than $10^{10} M_{\odot}$.
6. Almost all the sources in the red-sequence have $SSFR^{-1}$ larger than the age of the Universe at their redshift, which is consistent with passive galaxies. Following the same approach as in Cardamone and collaborators (i.e., combining the rest-frame $(U - V)$ vs. $\text{Log } M_{*}$ and the rest-frame $(U - V)$ vs. $(V - J)$ color diagrams), we find that, consistent with their results, $\sim 50\%$ of AGN hosts lie in the passive region of this diagram. In contrast from Cardamone et al. (2010), only $\sim 30\%$ of AGN host-galaxies in the green-valley in our sample are consistent with dust-obscured sources in $0.8 \leq z \leq 1.2$.

It is clear that the mid and far-infrared parts of the SED are under-sampled with respect to the optical part. The ongoing Herschel survey over various fields at different depths (100 μm and 160 μm in the COSMOS field) and the upcoming ALMA surveys will allow us to gain an optimal multiwavelength coverage also in the far-infrared.

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References

- Antonucci, R. 1993, ARA&A, 31, 473
 Bell, E. F., Wolf, C., Meisenheimer, K., et al. 2004, ApJ, 608, 752
 Bolzonella, M., Miralles, J., & Pelló, R. 2000, A&A, 363, 476
 Borch, A., Meisenheimer, K., Bell, E. F., et al. 2006, A&A, 453, 869
 Brusa, M., Civano, F., Comastri, A., et al. 2010, ApJ, 716, 348
 Brusa, M., Fiore, F., Santini, P., et al. 2009, A&A, 507, 1277
 Bruzual, G., & Charlot, S. 2003, MNRAS, 344, 1000
 Bundy, K., Georgakakis, A., Nandra, K., et al. 2008, ApJ, 681, 931
 Calzetti, D., Armus, L., Bohlin, R. C., et al. 2000, ApJ, 533, 682
 Capak, P., Aussel, H., Ajiki, M., et al. 2007, ApJS, 172, 99
 Capak, P., Aussel, H., Ajiki, M., et al. 2008, VizieR Online Data Catalog, 2284, 0
 Cappelluti, N., Brusa, M., Hasinger, G., et al. 2009, A&A, 497, 635
 Cardamone, C. N., Urry, C. M., Schawinski, K., et al. 2010, ApJ, 721, L38
 Chabrier, G. 2003, ApJ, 586, L133
 Chary, R., & Elbaz, D. 2001, ApJ, 556, 562
 Ciotti, L., Ostriker, J. P., & Proga, D. 2010, ApJ, 717, 708
 Corbin, M. R. 2000, ApJ, 536, L73
 Cowie, L. L., & Barger, A. J. 2008, ApJ, 686, 72
 Dale, D. A., & Helou, G. 2002, ApJ, 576, 159
 Dale, D. A., Helou, G., Contursi, A., Silbermann, N. A., & Kolhatkar, S. 2001, ApJ, 549, 215
 Dekel, A., Sari, R., & Ceverino, D. 2009, ApJ, 703, 785
 Elitzur, M., & Shlosman, I. 2006, ApJ, 648, L101
 Elvis, M., Wilkes, B. J., McDowell, J. C., et al. 1994, ApJS, 95, 1
 Emmering, R. T., Blandford, R. D., & Shlosman, I. 1992, ApJ, 385, 460
 Ferrarese, L., & Merritt, D. 2000, ApJ, 539, L9
 Fontana, A., Santini, P., Grazian, A., et al. 2009, A&A, 501, 15
 Gandhi, P., Horst, H., Smette, A., et al. 2009, A&A, 502, 457
 Gilli, R., Comastri, A., & Hasinger, G. 2007, A&A, 463, 79
 Granato, G. L., & Danese, L. 1994, MNRAS, 268, 235
 Hao, H., Elvis, M., Civano, F., et al. 2010, ApJ, 724, L59
 Hasinger, G. 2008, A&A, 490, 905
 Hasinger, G., Cappelluti, N., Brunner, H., et al. 2007, ApJS, 172, 29
 Hoeft, M., & Brüggem, M. 2004, ApJ, 617, 896
 Hopkins, P. F., Hernquist, L., Cox, T. J., et al. 2006, ApJS, 163, 1
 Hopkins, P. F., Richards, G. T., & Hernquist, L. 2007, ApJ, 654, 731
 Hopkins, P. F., Hernquist, L., Cox, T. J., & Kereš, D. 2008, ApJS, 175, 356
 Hornschemeier, A. E., Brandt, W. N., Garmire, G. P., et al. 2001, ApJ, 554, 742
 Jiang, L., Fan, X., Brandt, W. N., et al. 2010, Nature, 464, 380
 Kauffmann, G., & Haehnelt, M. 2000, MNRAS, 311, 576
 Kauffmann, G., Heckman, T. M., Tremonti, C., et al. 2003, MNRAS, 346, 1055
 Kelly, B. C., Bechtold, J., Trump, J. R., Vestergaard, M., & Siemiginowska, A. 2008, ApJS, 176, 355
 Kereš, D., Katz, N., Fardal, M., Davé, R., & Weinberg, D. H. 2009, MNRAS, 395, 160
 Koekemoer, A. M., Aussel, H., Calzetti, D., et al. 2007, ApJS, 172, 196
 Komatsu, E., Dunkley, J., Nolte, M. R., et al. 2009, ApJS, 180, 330
 Kormendy, J., & Richstone, D. 1995, ARA&A, 33, 581
 Le Floc'h, E., Aussel, H., Ilbert, O., et al. 2009, ApJ, 703, 222
 Lilly, S. J., Le Fèvre, O., Renzini, A., et al. 2007, ApJS, 172, 70
 Lusso, E., & Ciotti, L. 2011, A&A, 525, A115
 Lusso, E., Comastri, A., Vignali, C., et al. 2010, A&A, 512, A34
 Lutz, D., Maiolino, R., Spoon, H. W. W., & Moorwood, A. F. M. 2004, A&A, 418, 465
 Magorrian, J., Tremaine, S., Richstone, D., et al. 1998, AJ, 115, 2285
 Mainieri, V., Hasinger, G., Cappelluti, N., et al. 2007, ApJS, 172, 368
 Mainieri, V., Bongiorno, A., Merloni, A., et al. 2011, A&A
 DOI: [10.1051/0004-6361/201117259](https://doi.org/10.1051/0004-6361/201117259)
 Maiolino, R., Marconi, A., & Oliva, E. 2001, A&A, 365, 37
 Marconi, A., & Hunt, L. K. 2003, ApJ, 589, L21
 Marconi, A., Risaliti, G., Gilli, R., et al. 2004, MNRAS, 351, 169
 McCracken, H. J., Radovich, M., Iovino, A., et al. 2008, VizieR Online Data Catalog, 2286, 0
 Menci, N., Fontana, A., Giallongo, E., & Salimbeni, S. 2005, ApJ, 632, 49
 Nair, P. B., & Abraham, R. G. 2010, ApJS, 186, 427
 Nandra, K., Georgakakis, A., Willmer, C. N. A., et al. 2007, ApJ, 660, L11
 Nenkova, M., Sirocky, M. M., Nikutta, R., Ivezić, Ž., & Elitzur, M. 2008, ApJ, 685, 160
 Pozzi, F., Vignali, C., Comastri, A., et al. 2007, A&A, 468, 603
 Pozzi, F., Vignali, C., Comastri, A., et al. 2010, A&A, 517, A11
 Prescott, M. K. M., Impey, C. D., Cool, R. J., & Scoville, N. Z. 2006, ApJ, 644, 100
 Richards, G. T., Lacy, M., Storrie-Lombardi, L. J., et al. 2006, ApJS, 166
 Rovilos, E., & Georgantopoulos, I. 2007, A&A, 475, 115
 Sandage, A., & Visvanathan, N. 1978, ApJ, 225, 742
 Sanders, D. B., Salvato, M., Aussel, H., et al. 2007, ApJS, 172, 86
 Sazonov, S. Y., Ostriker, J. P., Ciotti, L., & Sunyaev, R. A. 2005, MNRAS, 358, 168
 Scarlata, C., Carollo, C. M., Lilly, S., et al. 2007, ApJS, 172, 406
 Schawinski, K., Virani, S., Simmons, B., et al. 2009, ApJ, 692, L19
 Schawinski, K., Treister, E., Urry, C. M., et al. 2011, ApJ, 727, L31
 Schlegel, D. J., Finkbeiner, D. P., & Davis, M. 1998, ApJ, 500, 525
 Silva, L., Granato, G. L., Bressan, A., & Danese, L. 1998, ApJ, 509, 103
 Silva, L., Maiolino, R., & Granato, G. L. 2004, MNRAS, 355, 973
 Silverman, J. D., Mainieri, V., Lehmer, B. D., et al. 2008, ApJ, 675, 1025
 Silverman, J. D., Lamareille, F., Maier, C., et al. 2009, ApJ, 696, 396
 Soifer, B. T., Boehmer, L., Neugebauer, G., & Sanders, D. B. 1989, AJ, 98, 766
 Somerville, R. S., Hopkins, P. F., Cox, T. J., Robertson, B. E., & Hernquist, L. 2008, MNRAS, 391, 481
 Springel, V., Di Matteo, T., & Hernquist, L. 2005, MNRAS, 361, 776
 Stickel, M., Lemke, D., Klaas, U., et al. 2000, A&A, 359, 865
 Strateva, I., Ivezić, Ž., Knapp, G. R., et al. 2001, AJ, 122, 1861
 Tremaine, S., Gebhardt, K., Bender, R., et al. 2002, ApJ, 574, 740
 Trump, J. R., Impey, C. D., Gabor, J. M., et al. 2007, in The Central Engine of Active Galactic Nuclei, ed. L. C. Ho, & J.-W. Wang, ASP Conf. Ser., 373, 726
 Trump, J. R., Impey, C. D., Kelly, B. C., et al. 2011, ApJ, 733, 60
 Urry, C. M., & Padovani, P. 1995, PASP, 107, 803
 Vasudevan, R. V., & Fabian, A. C. 2009, MNRAS, 392, 1124
 Vasudevan, R. V., Fabian, A. C., Gandhi, P., Winter, L. M., & Mushotzky, R. F. 2010, MNRAS, 402, 1081
 Zamojski, M. A., Schiminovich, D., Rich, R. M., et al. 2007, ApJS, 172, 468