

# The HARPS search for Earth-like planets in the habitable zone

## I. Very low-mass planets around HD 20794, HD 85512, and HD 192310<sup>\*,\*\*</sup>

F. Pepe<sup>1</sup>, C. Lovis<sup>1</sup>, D. Ségransan<sup>1</sup>, W. Benz<sup>2</sup>, F. Bouchy<sup>3,4</sup>, X. Dumusque<sup>1</sup>, M. Mayor<sup>1</sup>, D. Queloz<sup>1</sup>,  
N. C. Santos<sup>5,6</sup>, and S. Udry<sup>1</sup>

<sup>1</sup> Observatoire de Genève, Université de Genève, 51 ch. des Maillettes, 1290 Versoix, Switzerland  
e-mail: Francesco.Pepe@unige.ch

<sup>2</sup> Physikalisches Institut Universität Bern, Sidlerstrasse 5, 3012 Bern, Switzerland

<sup>3</sup> Institut d'Astrophysique de Paris, UMR7095 CNRS, Université Pierre & Marie Curie, 98bis Bd Arago, 75014 Paris, France

<sup>4</sup> Observatoire de Haute-Provence/CNRS, 04870 St. Michel l'Observatoire, France

<sup>5</sup> Centro de Astrofísica da Universidade do Porto, Rua das Estrelas, 4150-762 Porto, Portugal

<sup>6</sup> Departamento de Física e Astronomia, Faculdade de Ciências, Universidade do Porto, Portugal

Received 8 April 2011 / Accepted 15 August 2011

### ABSTRACT

**Context.** In 2009 we started an intense radial-velocity monitoring of a few nearby, slowly-rotating and quiet solar-type stars within the dedicated HARPS-Upgrade GTO program.

**Aims.** The goal of this campaign is to gather very-precise radial-velocity data with high cadence and continuity to detect tiny signatures of very-low-mass stars that are potentially present in the habitable zone of their parent stars.

**Methods.** Ten stars were selected among the most stable stars of the original HARPS high-precision program that are uniformly spread in hour angle, such that three to four of them are observable at any time of the year. For each star we recorded 50 data points spread over the observing season. The data points consist of three nightly observations with a total integration time of 10 min each and are separated by two hours. This is an observational strategy adopted to minimize stellar pulsation and granulation noise.

**Results.** We present the first results of this ambitious program. The radial-velocity data and the orbital parameters of five new and one confirmed low-mass planets around the stars HD 20794, HD 85512, and HD 192310 are reported and discussed, among which is a system of three super-Earths and one that harbors a  $3.6 M_{\oplus}$ -planet at the inner edge of the habitable zone.

**Conclusions.** This result already confirms previous indications that low-mass planets seem to be very frequent around solar-type stars and that this may occur with a frequency higher than 30%.

**Key words.** planetary systems – instrumentation: spectrographs – techniques: radial velocities – stars: individual: HD 20794 – stars: individual: HD 85512 – stars: individual: HD 192310

## 1. Introduction

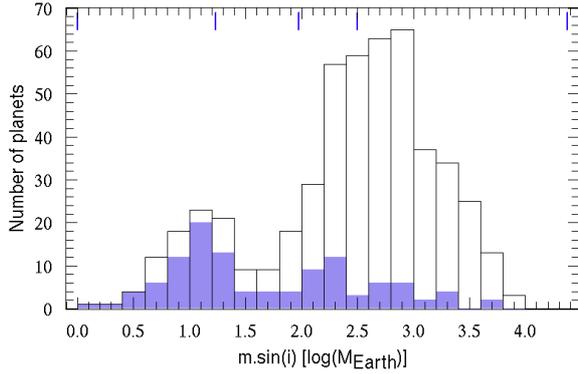
During the past years the field of extra-solar planets evolved toward the exploration of very low-mass planets down to the regime of super-Earths, i.e., to objects of only few times the Earth mass. Although finding Earth-like planets is probably the main trigger for these searches, one has to consider that their characterization contributes in a significant way to building up a general picture of how exoplanets form and evolve. Knowing the frequency and nature of these planets may enable us to distinguish between various theories and models, deliver new inputs and constraints to them, and this knowledge may contribute to refining their parameters. On the other hand, the models provide us with predictions which can be verified by observations. For instance, the presently discovered low-mass planets are predicted to be only the tip of the iceberg of a huge, still undiscovered planetary population (Mordasini et al. 2009). If this is

confirmed, hundreds of new planets will be discovered in a near future as the radial-velocity (RV) precision improves and the various programs increase their effectiveness and their time basis.

ESO's HARPS instrument (Mayor et al. 2003) has certainly played a key role by delivering more than 100 new candidates in its first eight years of operation. The most important and impressive contribution of this instrument lies more specifically in the field of super-Earths and Neptune-mass planets. Indeed, about 2/3 of the planets with mass below  $18 M_{\oplus}$  known to date have been discovered by HARPS. This new era started in 2004 with the discovery of several Neptune-mass planets such as  $\mu$  Ara c (Santos et al. 2004; see also Pepe et al. 2007, for an update of the parameters), 55 Cnc (McArthur et al. 2004), and GJ 436 (Butler et al. 2004). Many more followed, but the detection of the planetary system HD 69830 containing three Neptune-mass planets (Lovis et al. 2006), and that of HD 40307 with its three super-Earths (Mayor et al. 2009b) best illustrate the huge potential of HARPS. HARPS also revealed to us the system G1 581, with two possibly rocky planets c and d with masses of 5 and 8  $M_{\oplus}$ , respectively, both lying at the edge of the habitable zone (HZ) of their parent star (Udry et al. 2007; Selsis et al. 2007), and the G1 581 e of only  $1.9 M_{\oplus}$  (Mayor et al. 2009a). Last but not least, we should mention the system around HD 10180, with

\* Based on observations made with the HARPS instrument on ESO's 3.6 m telescope at the La Silla Observatory in the frame of the HARPS-Upgrade GTO program ID 086.C-0230.

\*\* Tables 7–9 (RV data) are only available at the CDS via anonymous ftp to cdsarc.u-strasbg.fr (130.79.128.5) or via <http://cdsarc.u-strasbg.fr/viz-bin/qcat?J/A+A/534/A58>



**Fig. 1.** Histogram of companion masses. The dark area shows the companions discovered with HARPS and includes the candidates which are in preparation for publication.

seven extra-solar planets of low mass, among which is the lightest ever detected exoplanet HD 10180 b, with  $m \cdot \sin i$  of only  $1.5 M_{\oplus}$  (Lovis et al. 2011b).

Figure 1 shows the mass distribution of companions around solar-type stars of all published planets, including those of HARPS that are in preparation. The HARPS detections (dark area) probe a mass range that suffered from strong detection biases before. It must be noticed that the planet frequency increases again below  $20 M_{\oplus}$ . This increase is highly significant, because higher-mass planets induce higher RV variations on their parent star and are therefore detected more easily. Moreover, a recent investigation of the HARPS high-precision sample has shown that about 1/3 of all sample stars exhibit RV variations, which indicates the presence of super-Earths or ice giants (Lovis et al. 2009). Indeed, planet formation models (Ida & Lin 2008; Mordasini et al. 2009; Alibert et al. 2011) show that only a small fraction (on the order of 10%) of all existing embryos will be able to grow and become giant planets. Hence, we expect that the majority of solar-type stars will be surrounded by low-mass planets. The inclusion of telluric planets can only increase the planet frequency further and consequently the probability of detection. This implies that even a small sample of target stars is likely to reveal, if followed with high enough accuracy, several Earth-like planets. Pushing the measurement precision even more farther should therefore naturally increase the probability of finding new planets, especially of very low-mass and rocky planets that are possibly in the habitable zone of their host star.

In the following section we will describe our program for the specific search for low-mass and rocky extra-solar planets that possibly lie in the habitable zone of their parent star. The observations of these targets is presented in Sect. 3. The three host stars of the program around which we have detected new planets will be described in greater detail in Sect. 4. Section 5 presents the three newly discovered planets as well as the confirmation of a recently announced planet. Finally, a short discussion and an outlook on the occurrence frequency of low-mass planets is given in Sect. 6.

## 2. Searching for Earth analogs around nearby stars

The recent HARPS discoveries made us aware that discovering Earth-like exoplanets is already within the reach of HARPS, although important questions still remain open: how frequent are low-mass planets? At what level is the detection bias? Where is

the precision limit set by stellar noise? Can we detect low-mass planets in the habitable zone of their parent star if sufficient observations on a high-precision instrument are invested? Driven by the very encouraging results of the HARPS high-precision program (Udry & Mayor 2008) and the hints for high planet occurrence announced by Lovis et al. (2009), we have decided to investigate these questions in more detail. For this purpose we have defined a specific program on HARPS for the search of rocky planets in the habitable zone HZ based on Guaranteed Time Observations (GTO).

We decided to observe ten nearby stars with the highest possible radial-velocity precision and sampling during a period of four years. Our goal is to probe the presence of extra-solar planets as similar as possible to our Earth. Therefore, we have selected targets with the following characteristics:

- the stellar magnitude must be such that we are not limited by photon noise within a 15-min exposure, i.e., ensure an RV-precision of better than  $30 \text{ cm s}^{-1}$ ;
- stars should lie within 10 pc from the Sun;
- the targets must have characteristics that guarantee the best RV precision, i.e., they should be of late-G to early-K spectral type and have low rotational velocity (a rotational velocity below what can be detected with HARPS, i.e.,  $v \sin i < 1 \text{ km s}^{-1}$ ) and low activity indicators ( $\log(R'_{\text{HK}}) < -4.85$ );
- the RV scatter measured with HARPS over the first years of operation must be lower than 2 m/s rms over several years and based on a minimum of ten recorded data points;
- the ten targets must be evenly distributed in terms of right-ascension to allow the observation of at least three to four targets at any period of the year.

Combining all these criteria resulted in a strong down-selection of the 400 possible candidates of the original HARPS high-precision program. In particular the requirement on the distance from the Sun turned out to be too limiting considering all the other requirements, so that we had to relax this requirement and allow targets to enter the list with distances up to 16 pc.

The ten selected targets are given in Table 1. The most prominent member is Alpha Cen B (HD 128621). This star is of particular interest because of its brightness and its short distance to the Sun. If planets were detected around Alpha Cen B, they would be ideal candidates for additional follow-up by spectroscopy, astrometry, photometry, etc. Furthermore, this star is part of a triple system, which makes the question even more interesting. However, what appears to be the greatest advantage generates some observational challenge: on the one hand, the bright magnitude of Alpha Cen B is a limiting factor for the telescope guiding system and may result in poorer RV precision due to incomplete light scrambling across the spectrograph's entrance slit. On the other hand, because Alpha Cen B is in a triple system, the RVs must be analyzed by considering a complete – and precise – orbital model, especially when looking at long-period planets.

Another prominent candidate is  $\tau$  Ceti (HD 10700). This star was already known to be an “RV-standard” star in the sense that it is followed-up by several groups and shows a very low RV dispersion. Indeed, the RV dispersion measured with HARPS over a time span of more than seven years and involving 157 data points is  $0.93 \text{ ms}$ ! Because of its magnitude, the spectral type, the RV stability, and its low chromospheric activity level,  $\tau$  Ceti is close to the perfect candidate for our program.

The remaining eight candidates show similar characteristics. They are all part of the original HARPS high-precision program. As such, they have all been observed since the very beginning

**Table 1.** Targets of the HARPS search for Earth analogs around nearby stars program.

Target	RA [h:m:s]	Dec [°:′:″]	Sp. type	VMag	Dist. [pc]
HD 1581	00:20:04	−64:52:29	F9V	4.23	8.59
HD 10700	01:44:04	−15:56:15	G8V	3.49	3.65
HD 20794	03:19:55	−43:04:11	G8V	4.26	6.06
HD 65907A	07:57:46	−60:18:11	G2V	5.59	16.19
HD 85512	09:51:07	−43:30:10	K5V	7.67	11.15
HD 109200	12:33:31	−68:45:20	K0V	7.13	16.17
HD 128621	14:39:35	−60:50:14	K1V	1.35	1.35
HD 154577	17:10:10	−60:43:43	K0V	7.38	13.69
HD 190248	20:08:44	−66:10:55	G5IV-V	3.55	6.11
HD 192310	20:15:17	−27:01:58	K3V	5.73	8.82

of the HARPS operations, which allowed us to make a preliminary selection also on the basis of preliminary RV-dispersion and chromospheric activity indicator values. Since the start of our specific program on April 1st, 2009, the number of measurements has increased and the mentioned values refined.

An important aspect to be pointed out is that the presented program carries some “risks”: by construction this program pre-selected only stars very stable in terms of radial velocities. In other terms, we were sure that they would not harbor any *known* planet. Even worse, their raw RV dispersion does not leave much room for “strong” RV signals. Of course, we were aware of this and that we might end up with no positive result.

There is no doubt that successful detections depend on the observational precision, in particular when searching for signals with semi-amplitudes below  $1\text{--}2\text{ m s}^{-1}$ , as is the case for the presented program. Five years of observations have proved that, on quiet and bright dwarf stars, a radial-velocity precision well below  $1\text{ m s}^{-1}$  can be achieved on both short-term (night) and long-term (years) timescales. A direct confirmation is offered by the raw RV dispersions measured for the candidates of our program and given in Table 2.

To make a rough estimation of the detection limits we may expect, let us assume a precision of  $0.7\text{ m s}^{-1}$  on each data point all included, i.e. stellar noise, instrumental errors, atmospheric effects, etc. The photon noise is “by design” well below this level and can be neglected. With 50 well-sampled data points we should then be able to detect a  $50\text{ cm s}^{-1}$  RV semi-amplitude at better than a 3-sigma level. This signal corresponds to a  $2.2 M_{\oplus}$  on a one-month orbit or a  $4 M_{\oplus}$  on a six-month orbit. When extended to three years, the same observational strategy will allow us to increase the number of data points and consequently reduce the detectable semi-amplitude to about  $30\text{ cm s}^{-1}$  or the equivalent of a  $1.3 M_{\oplus}$  planet on a one-month orbit, or a  $2.3 M_{\oplus}$  on a six-month orbit. On the other hand, it is interesting to note that such a planet would lie at the inner edge of the habitable zone of Alpha Cen B or even inside the habitable zone of HD 85512 or HD 192310, which are later-type stars.

In the previous discussion it was assumed of course that no RV signal induced by stellar activity is present at the orbital period of investigation. In the opposite case, the detection limit will not decrease proportionally to the square-root of the number of data points. Even worse, if the stellar activity signal has a coherence time similar to or longer than the observational period, it becomes indistinguishable from a possible radial-velocity signal of similar periodicity. However, other indicators such as precise photometry of the star, line bisector or stellar activity indicators can help to identify possible correlations with the radial-velocity signal and thus provide a tool to distinguish stellar noise from a

**Table 2.** Observations of the targets with HARPS.

Target	Data points	Time span [days]	RV scatter [ $\text{m s}^{-1}$ ]	$\log(R'_{\text{HK}})$
HD 1581	93	2566	1.26	$-4.93 \pm 0.003$
HD 10700	141	2190	0.92	$-4.96 \pm 0.003$
HD 20794	187	2610	1.20	$-4.98 \pm 0.003$
HD 65907A	39	2306	1.45	$-4.91 \pm 0.006$
HD 85512	185	2745	1.05	$-4.90 \pm 0.043$
HD 109200	77	2507	1.16	$-4.95 \pm 0.018$
HD 128621	171	2776	NA	$-4.96 \pm 0.014$
HD 154577	99	2289	1.05	$-4.87 \pm 0.026$
HD 190248	86	2531	1.26	$-5.09 \pm 0.009$
HD 192310	139	2348	2.62	$-4.99 \pm 0.033$

planetary signal. Another powerful tool is to verify that the RV signal is persistent and coherent over time.

### 3. Observations

#### 3.1. HARPS-Upgrade GTO

All observations presented here have been carried out using HARPS on ESO’s 3.6-m telescope at the La Silla Observatory, Chile (Mayor et al. 2003). A total of 30 nights per year over four years are allocated to this program on HARPS within our Guaranteed-Time Observations (HARPS-Upgrade GTO). Our goal was to observe each target about 50 times per season. These constraints, together with the need of observing each target several times per night to average stellar noise, set the actual number of targets we are able to follow-up. In the following we shall discuss the observational strategy. For this purpose we have to distinguish different types of measurements and adopt the following definitions: an *exposure* represents the single integration (one open-close shutter cycle). The *observation* is the average of all subsequent exposures. The *data point*, finally, denominates the average of all the exposures in a given night.

#### 3.2. Observational strategy

Thanks to the experience issued from the HARPS high-precision program we have been able to define an optimum strategy of observation to reduce the effects of stellar noise and its impact on the RV measurements. From the study by Dumusque et al. (2011b) we have learned that stellar pulsations are best averaged out by extending the measurement over a time longer than the typical stellar oscillation period. Because our targets are mainly bright G and K dwarfs, we ended up with exposure times of typically 15 min. To avoid saturation on bright objects, the exposures were split into several shorter exposures covering a total of 15 min, including overheads, leading to typical “open-shutter” time of about 10 min. From the same study we can also deduce that super-granulation noise, which is of a longer time scale, is not averaged out by the 15-min exposure. For this purpose, the daily time scale must be covered. This led us to observe each target at least twice, if possible three times during the same night, but in any case with the single observations separated as far as possible in time. A total observation time per target of 35–45 min per night had to be considered. If we assume ten hours per night in average, we end up with about 40 to 50 possible observations per target and per year.

Dumusque et al. (2011b) show that longer-period granulation noise and RV-jitter caused by spots and activity are best averaged out by evenly sampling over the observational

season. If we fix the amount of available observation time, an equidistant time grid would be the best solution. However, the visitor-observing scheme applied to ESO’s 3.6-m telescope does not allow us to obtain optimum sampling. Nevertheless, it has been possible to join our GTO time with observations of other large programs. The result is that, during nights allocated to this joint program, about two to three targets were observed on average every night, while no data are obtained during other nights, leaving some observational gap with a typical period of one to a few weeks.

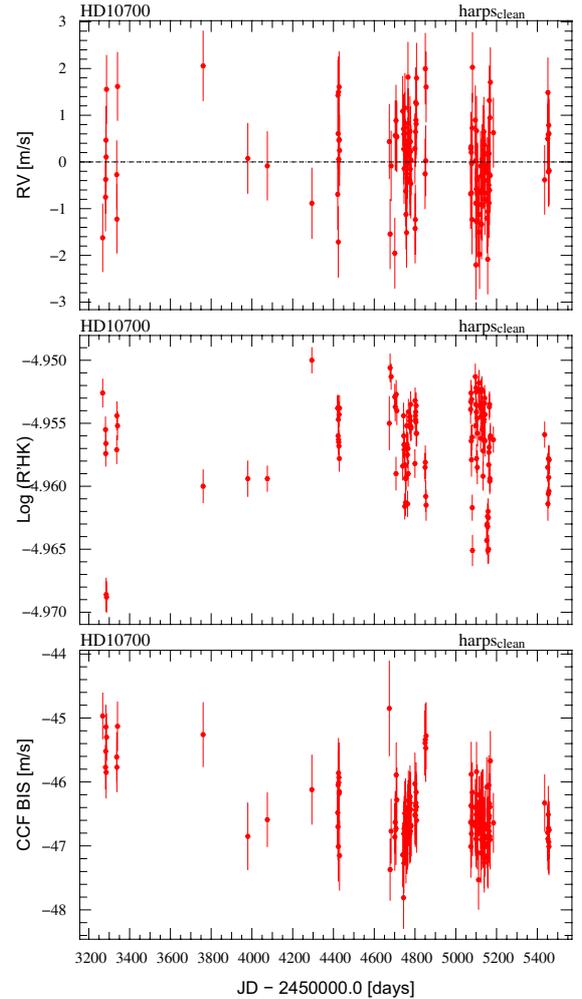
### 3.3. Obtained measurements

Table 2 summarizes the measurements made with HARPS on each of the program’s targets, including measurements issued from the original high-precision program. The second column indicates the number of data points (nightly averages of individual exposures) acquired to date. Thanks to the data-points of the original high-precision program, most of the observations span more than six years for any object. It must be noted however that the first-year observations on HARPS were not carried out using the optimum observational strategy, which resulted in data points – actually single exposures – of poorer precision. Despite this, the RV scatter of all these data points is remarkably low, which justifies our initial choice of the targets. The only target for which the RV scatter exceeds  $1.5 \text{ m s}^{-1}$  is HD 192310. Interestingly, this star has been recently announced by Howard et al. (2011a) to harbor a Neptune-mass planet. The discovery is confirmed by the HARPS data and will be described in more detail in Sect. 5. The RV scatter of Alpha Cen B (HD 128621) is not indicated in the table, because it is affected by the long-term drift caused by its stellar companion. A detailed discussion specifically of Alpha Cen B will be presented in a forthcoming paper. For all targets we give the value of the chromospheric activity indicator  $\log(R'_{\text{HK}})$ . As mentioned above, the targets where selected to have  $\log(R'_{\text{HK}}) < -4.85$ . The long-term averages and the dispersion values given in the last column confirm that all targets comply with this requirement.

### 3.4. The example of Tau Ceti

To illustrate the kind of objects we are dealing with and the level of precision that can be obtained with HARPS, we will focus a moment on HD 10700, also known as Tau Ceti. The dispersion over the RVs of the *individual exposures* (typically a few minutes) is  $1.5 \text{ m s}^{-1}$  rms. As seen in Table 2, the dispersion of the data points (*nightly averages*) is instead only  $0.92 \text{ m s}^{-1}$ , which proves that the strategy proposed by Dumusque et al. (2011b) provides good smoothing of the stellar contributions at the time scale of a night. HD 10700 turns out to be part of the most quiet stars in our sample despite the large number of data points spanning more than six years of observations. What is even more remarkable is that none of the parameters shown in Fig. 2, i.e., the radial velocity, the chromospheric activity indicator  $\log(R'_{\text{HK}})$ , and the average line bisector of the cross-correlation function *BIS*, show any trend over the six years of observations. For instance, the activity indicator shows a dispersion of only 0.003 dex rms. Finally, we underline the stability of the line bisector, which has a dispersion of about  $0.5 \text{ m s}^{-1}$ , confirming the high stability of the HARPS’ instrumental profile (IP).

In the upper panel of Fig. 3 we plotted the generalized Lomb-Scargle (GLS) periodogram of the radial velocity data of

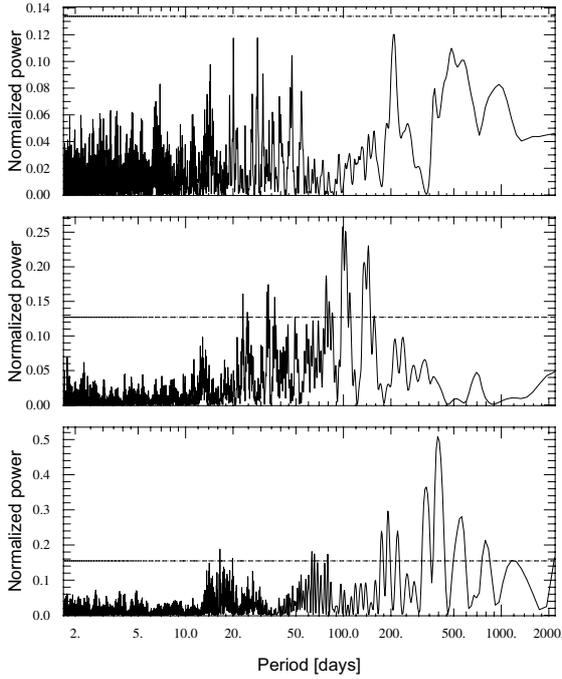


**Fig. 2.** Radial velocities, chromospheric activity indicator  $\log(R'_{\text{HK}})$ , and line bisector *CCF BIS* of HD 10700 versus time.

HD 10700. Note that none of the peaks reaches the 10% false-alarm probability (FAP) level. This demonstrates that there is no significant planetary signal in the precise HARPS data available to date and we can set a higher limit for the masses of possible planets around Tau Ceti as a function of period. For a detailed analysis and discussion we refer to Lovis et al. (in prep.).

In the central and lower panel of Fig. 3 we present the GLS periodogram of the  $\log(R'_{\text{HK}})$  and the line bisector *BIS*. The peaks are significant here, in contrast to the RV data, and indicate that there is some “signal” that clearly is above the noise. Periodograms of activity indicators like  $\log(R'_{\text{HK}})$  and line bisector *BIS* often exhibit a complicated pattern of significant peaks around the stellar rotation period and at longer periods. This can probably be explained by the combination of irregular sampling and the presence of signals with short coherence times in the data, as expected from stellar activity and spots. Indeed, solar-type stars are likely covered by a large number of small active regions that rotate with the star and have lifetimes on the order of the rotation period or below. The aspect to be pointed out here, however, is that these peaks do not have any correspondence in the RV GLS periodogram. As remarked above, the (low but significant) activity signal is not correlated to any RV variation.

One may also wonder how strongly stellar activity may influence the radial-velocities. We refer to Dumusque et al. (2011a) and Lovis et al. (2011a) for a detailed discussion. In the specific case of Tau Ceti we observe only little RV jitter although



**Fig. 3.** GLS periodograms of the HD 10700 radial velocities (*upper panel*), its  $\log(R'_{\text{HK}})$  values (*central panel*) and line bisector *BIS* (*lower panel*). The false-alarm probability (FAP) thresholds of 10% is indicated by a dashed line.

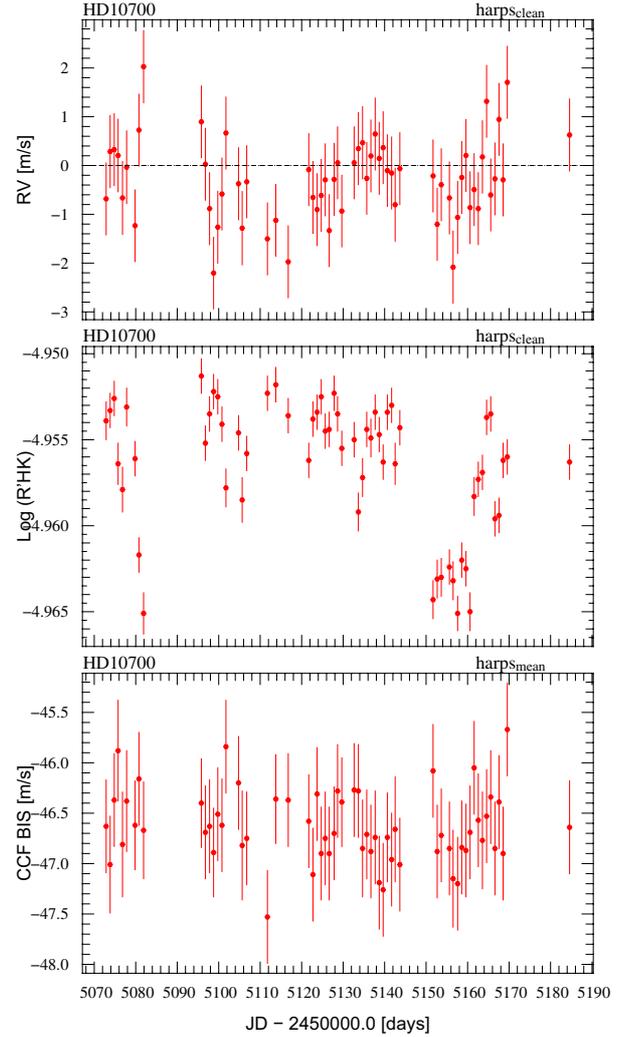
the  $\log(R'_{\text{HK}})$  seems to vary during some seasons, while it remains stable over others. As an example, Fig. 4 shows the radial velocity,  $\log(R'_{\text{HK}})$  and *BIS* over a single season. Error bars on  $\log(R'_{\text{HK}})$  values have been calculated from the expected photon statistics, but to take into account the effective noise floor (determined on the most stable stars), an uncertainty of 0.001 was quadratically added. This noise floor is likely caused by residual instrumental systematics. The variation in  $\log(R'_{\text{HK}})$  is larger than the expected error, but does not seem to follow the radial velocity, and vice-versa. This is better illustrated in Fig. 5, where the correlation between the two quantities is plotted. We deduce that in the specific case of Tau Ceti, the radial velocity is not affected by stellar activity.

HD 10700 therefore is a star that shows no detectable coherent signal in the RV data, while activity indicators exhibit some non-random, low-amplitude patterns that likely convey information on small active regions at the stellar surface. In this sense, these indicators are more sensitive to activity than RV data, which make them useful tools to better understand stellar jitter on quiet solar-type stars. Work is going on to simulate and model these auxiliary data, which may then allow us to eventually correct the RV data themselves from activity effects.

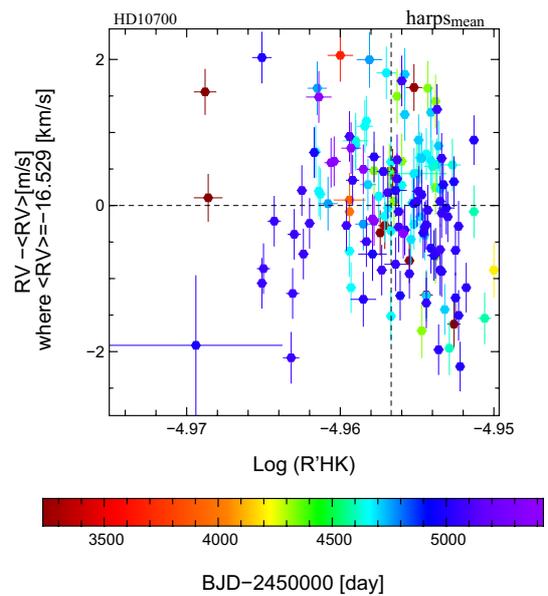
#### 4. Parameters of the planet-hosting stars

Of the ten targets presented in the previous section, three presently show planetary signatures. The stars are HD 20794, HD 85512 and HD 192310. The target HD 192310 was already announced recently by Howard et al. (2011a) to harbor a Neptune-mass planet on a 74 days orbit. Below we will confirm this detection and show that a second planet orbits this star on a much longer orbit.

The photometric and astrometric parameters of the three targets are presented in the first part of Table 3. They are all obtained from the HIPPARCOS catalog (ESA 1997). Given that



**Fig. 4.** Radial velocity, chromospheric activity indicator  $\log(R'_{\text{HK}})$  and line bisector *BIS* of HD 10700 over a single season.



**Fig. 5.** Correlation plot of radial velocity versus chromospheric activity indicator  $\log(R'_{\text{HK}})$  for HD 10700. The color scale represents the observation date.

**Table 3.** Observed and inferred stellar parameters for the observed targets.

Parameter	HD 20794	HD 85512	HD 192310
Spectral Type	G8V	K5V	K3V
$V$	4.26	7.67	5.73
$B - V$	0.711	1.156	0.878
$V$ variability	<0.005	<0.016	<0.005
$\pi$ [mas]	$165.02 \pm 0.55$	$89.67 \pm 0.82$	$113.33 \pm 0.82$
Distance [pc]	6.06	11.15	8.82
$pmRA$ [mas/yr]	$39.21 \pm 0.44$	$461.28 \pm 0.56$	$1241.35 \pm 0.78$
$pmDec$ [mas/yr]	$27.57 \pm 0.45$	$-472.87 \pm 0.57$	$-181.46 \pm 0.51$
$M_V$	5.35	7.43	6.00
$L$ [ $L_\odot$ ]	$0.656 \pm 0.003$	$0.126 \pm 0.008$	$0.385 \pm 0.007$
[Fe/H] [dex]	$-0.40 \pm 0.01$	$-0.33 \pm 0.03$	$-0.04 \pm 0.03$
$M$ [ $M_\odot$ ]	0.70	0.69	0.80
$T_{\text{eff}}$ [K]	$5401 \pm 17$	$4715 \pm 102$	$5166 \pm 49$
$\log g$ [cgs]	$4.22 \pm 0.11$	$4.39 \pm 0.28$	$4.51 \pm 0.09$
$v \sin i$ [km s $^{-1}$ ]	<3	<3	<3
$\log(R'_{\text{HK}})$ [dex]	$-4.976 \pm 0.003$	$-4.903 \pm 0.043$	$-4.994 \pm 0.027$
$P_{\text{rot}}(R'v_{\text{HK}})$ [days]	$33.19 \pm 3.61$	$47.13 \pm 6.98$	$47.67 \pm 4.88$
age( $R'_{\text{HK}}$ ) [Gyr]	$5.76 \pm 0.66$	$5.61 \pm 0.61$	$7.81 \pm 0.90$

all targets are nearby stars, the astrometric parameters are also very precise.

The spectroscopic parameters presented in the second part of Table 3 have been collected from Sousa et al. (2008), who did an in-depth spectroscopic analysis of all stars of the HARPS high-precision program. We report them here without any further analysis. We only comment that for the K5V dwarf HD 85512 the stellar parameters derived in Sousa et al. may be affected by systematic errors (see e.g. Neves et al. 2009), given the late spectral type of the star. Indeed, Sousa et al. (2008) derived an effective temperature of  $4715 \pm 102$  K for this star, which disagrees with the value of 4419 K obtained using the IRFM calibration of Casagrande et al. (2006).

To check if this difference could significantly change the derived stellar metallicity and estimated stellar mass, we proceeded with a series of tests. First, fixing the effective temperature to 4419 K and the gravity and microturbulence values derived by Sousa et al. ( $4.39$  dex and  $0.23$  km s $^{-1}$ , typical for late K-dwarfs), we derived a new value for the stellar metallicity based on the list of Fe I lines (and their equivalent widths) measured by Sousa et al. The resulting [Fe/H] is  $-0.34$ , very close to the original value of Sousa et al. Finally, we estimated the stellar mass using both effective temperature values mentioned above, and making use of the mass- $T_{\text{eff}}$ - $\log g$ -[Fe/H] calibration of Torres et al. (2010). In this procedure we fixed the stellar metallicity to  $-0.32$  and the gravity to  $4.39$ , while attributing reasonable (though conservative) uncertainties of  $0.1$  and  $0.2$  dex to [Fe/H] and  $\log g$ , respectively. The error in effective temperature was considered to be  $100$  K in each case. After a series of 10 000 Monte Carlo trials, using values of  $T_{\text{eff}}$ ,  $\log g$ , and [Fe/H] taken from a Gaussian distribution around the central value and a dispersion of 1-sigma (error), we obtained mass values of  $0.73$  and  $0.65 M_\odot$  for input temperatures of  $4715$  and  $4419$  K, respectively (the uncertainty is  $0.08 M_\odot$ , taking into account the scatter of the individual mass values and the dispersion of the Torres et al. calibration). Both values are compatible within the error bars. The former value is close to the  $0.74 M_\odot$  derived by Sousa et al. through isochrone fitting. We thus conclude that both [Fe/H] and mass values are not strongly sensitive to the choice of the effective temperature scale. We decided to adopt the average value of the two for the

remainder of the paper:  $-0.33$  for [Fe/H] and  $0.69 M_\odot$  for the mass of HD 85512.

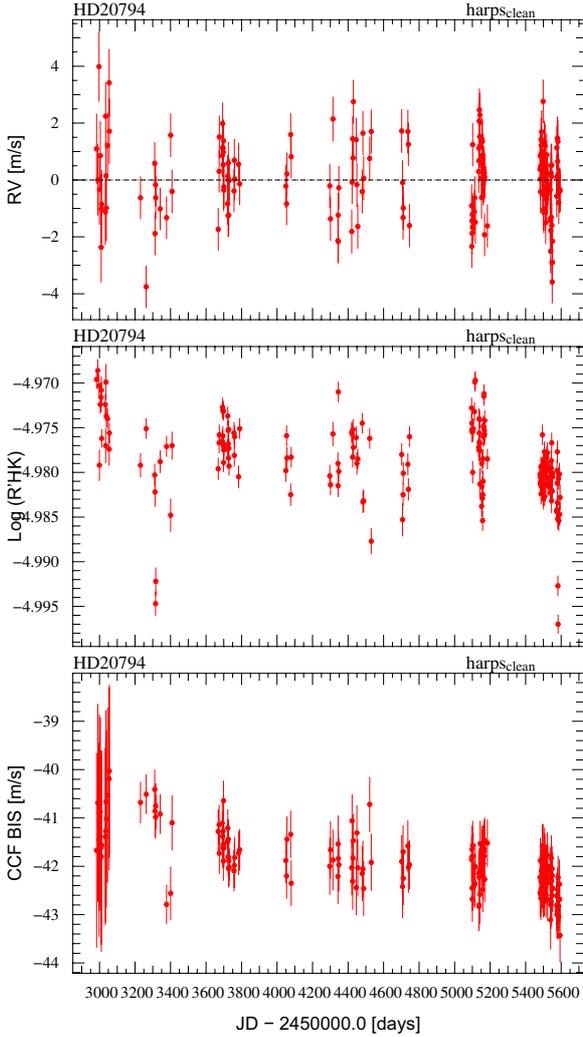
The third part of Table 3 shows additional stellar parameters: the projected stellar rotation velocity is issued from the instrumental calibration presented by Santos et al. (2002). The chromospheric activity indicator,  $\log(R'_{\text{HK}})$  has been computed in accordance with Lovis et al. (2011b). Finally, the stellar rotation period  $P_{\text{rot}}(R'v_{\text{HK}})$  and the estimated stellar age are both computed from  $\log(R'_{\text{HK}})$  following the prescription given in Mamajek & Hillenbrand (2008).

## 5. Results and orbital solutions

### 5.1. The planetary system of HD 20794

The raw radial-velocity data of the star HD 20794 are presented in Table 7 and shown in Fig. 6. As can be seen from these raw data, the star belongs to the most stable stars of the original HARPS high-precision sample. The dispersion of the 173 radial-velocity data points is of only  $1.2$  m s $^{-1}$  rms over the seven-year time span. The typical internal noise on the radial velocity data points, which takes into account photon and calibration noise, is of  $0.2$  to  $0.3$  m s $^{-1}$ . For the analysis, and to take into account any possible external noise such as instrumental, atmospheric and stellar contaminants, as well as stellar jitter, we added quadratically a systematic noise of  $0.7$  m s $^{-1}$  to the RV data. This value considers the typical dispersion obtained for the most stable stars of the program and prevents that some high SNR data are erroneously over-weighted with respect to others in the fitting process.

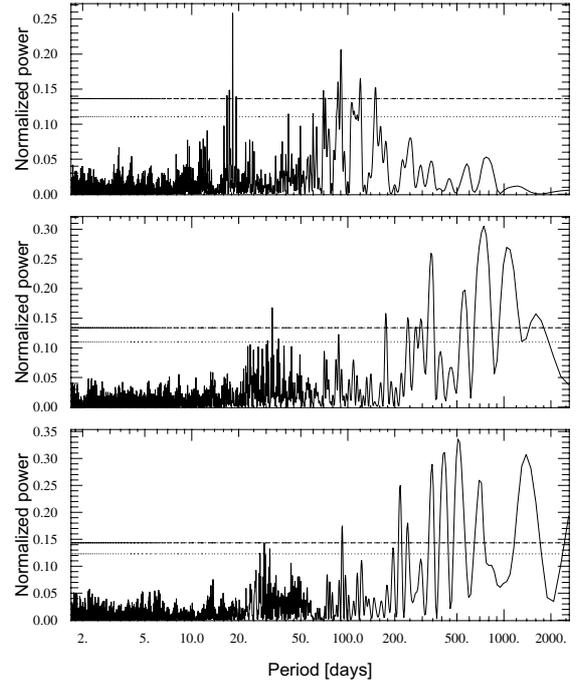
Although low, the dispersion largely exceeds the estimated total error. The GLS periodogram shows indeed a signal well above the noise level at about 18 days period (see Fig. 7). The dashed line indicates a false-alarm probability level of 1% but the  $P = 18$  days signal exceeds this level by far. In the same figure we plot for comparison the GLS periodogram of the activity indicator  $\log R'_{\text{HK}}$  and of the line bisector  $BIS$ . None of them reveal any significant power at the period indicated by the radial velocity. The  $\log(R'_{\text{HK}})$  indicator and the bisector, instead, show an excess of power at periods around 30 days, which points toward the rotational period of this star. Because the periods are significantly different from the radial-velocity



**Fig. 6.** Radial velocity,  $\log(R'_{\text{HK}})$  and *BIS* of HD 20794 as a function of time. The RV dispersion is  $1.23 \text{ m s}^{-1}$  rms.

period, we conclude that the two signals are de-correlated. Nevertheless, it must be noted that  $\log(R'_{\text{HK}})$  has a significant peak at  $P = 32$  days, which we suppose reflecting the stellar rotational period. Remarkably, this value perfectly corresponds to the value directly computed from  $\log(R'_{\text{HK}})$  using the Mamajek & Hillenbrand (2008) prescriptions.

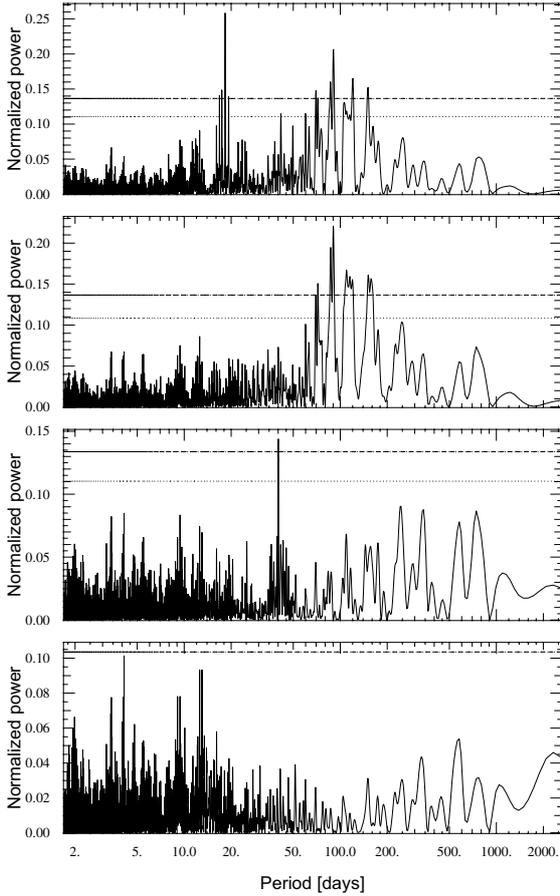
Based on the clear  $P = 18$  days signal issued from the periodogram, and reassured by the non-coincidence of the  $\log(R'_{\text{HK}})$  and *BIS* peaks with the radial-velocity peak, we fitted a single Keplerian to the radial velocities while letting all parameters free. The period of the obtained solution converges very rapidly and unambiguously towards a period of about 18 days, a semi amplitude around  $1 \text{ m s}^{-1}$ , and an eccentricity of 0.25. However, because the eccentricity error is of the same order as its value, and because the residuals almost do not depend on the eccentricity, we have decided to fix it to zero. For the  $N = 187$  data points we obtain a residual dispersion around the proposed solution of  $1.03 \text{ m s}^{-1}$ . The reduced  $\chi^2$  is thus of 2.06, indicating that some excess dispersion exists and must be caused by an additional signal or an external noise source, e.g., additional stellar noise, as suggested by the excess power in the  $\log(R'_{\text{HK}})$  and line-bisector spectra. The GLS periodogram of the residuals to the one-Keplerian fit (center of Fig. 8) indeed indicates excess



**Fig. 7.** GLS periodogram of the HD 20794 radial-velocity data with FAP levels at 10% and 1% level, respectively (*top*), the activity indicator  $\log(R'_{\text{HK}})$  (*center*) and the line bisector *BIS* (*bottom*). While a very clear peak appears at  $P = 18$  days for the radial velocities, no such signal is seen in the two other observables.

power around 90 days with a false-alarm probability well below 1%. This signal is not correlated with  $\log(R'_{\text{HK}})$  or *BIS* and is confirmed when computing the Pearson’s correlation coefficients, which is in both cases close to zero.

Encouraged by the significance of the two signals at 18 and 90 days and the absence of any correlation with the activity indicator and the line bisector, we decided to try a two-Keplerian model to fit our data. For the first run we left all parameters free. The obtained solution is quite convincing but the eccentricities are not significantly different from zero. Therefore, we fixed them to zero in the second run. The two-Keplerian model with zero eccentricities delivers two planets with a period of 18 and 90 days, both having a semi-amplitude of about  $0.8 \text{ m s}^{-1}$ . The residual dispersion of the data points around the fit is  $0.90 \text{ m s}^{-1}$  and  $\chi^2$  is equal to 1.61. It is interesting to note that the residuals to this solution still contain significant power at  $P = 40$  days. This is illustrated by the presence of a significant peak in the third plot of Fig. 8, which dominates with a FAP of 0.7%. It must be said, however, that the peak is strongly attenuated if the eccentricities of the two Keplerians “absorb” part of the 40 day-period signal. On the other hand, one could argue that a real non-zero eccentricity of the longer-period planet, if not taken into account in the fitted model, would produce harmonics. We would expect a signal around half the longer period, i.e., 45 days. The period of the peak in the residuals is significantly different, however. We conclude therefore that some additional contribution at  $P = 40$  days may still be present in the data, which has not been taken into account by the two-Keplerian model. Furthermore, the 40-day peak in the GLS periodogram does not appear to have any corresponding peak in either the activity indicator or in the line bisector. This is convincingly confirmed by the correlation plots in Fig. 9, which show a Pearson correlation coefficient of

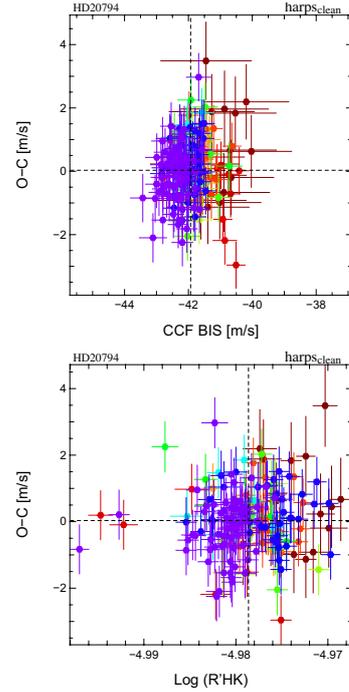


**Fig. 8.** Successive GLS periodograms of the HD 20794 time series, where the main signal is subtracted at each step by adding one more planet in the multi-Keplerian model from top to bottom, starting from the GLS of the raw data and following with the residuals to the one-Keplerian and the two-Keplerian fit. Eccentricities have been set to zero. False-alarm probability thresholds of 10% and 1% are indicated. The lowest GLS periodogram of the residuals to the three-Keplerian model shows only the 10% FAP level.

only 0.06 between the RV and activity indicator and of 0.15 between the RV and line bisector.

We studied the distribution of the RV residuals to both models (eccentric two-planet and circular three-planet) as a possible way to distinguish between, as done e.g. by [Andrae et al. \(2010\)](#). Unfortunately, both models yield distributions of normalized residuals close to Gaussian, and the Kolmogorov-Smirnov test gives similarly high probabilities that the residuals are drawn from a normal distribution (97% and 80% for the three-planet and two-planet models, respectively). Note that this result depends on the assumed level of jitter, and that reducing the error bars increases the difference between the models in favor of the three-planet model. However, it is hardly possible to choose between the models on this basis. Gathering more data may be the only way to confirm the existence of the third planet.

All these considerations lead us to adopt a three-Keplerian model. The resulting orbital parameters are presented in Table 4. The corresponding phase-folded radial-velocity curves are shown in Fig. 10. As mentioned above, the eccentricities were fixed to zero, but only after searching for generic solutions and verifying that the eccentricity values were not significant. The semi-amplitudes of the induced radial-velocity variation are *all*



**Fig. 9.** Correlation plot of the residual of the two-Keplerian model to the radial velocity of HD 20794 versus the line bisector and the activity indicator  $\log(R'_{\text{HK}})$ .

lower than  $1 \text{ m s}^{-1}$ . The orbital periods of the three components are at 18, 40, and 90 days.

The shortest and longest-period signal are very significant, while the  $P = 40$  days period is confidently detected in the periodogram but is less apparent in the phase-folded RV plot. Given the low amplitude, which would be the lowest ever detected radial-velocity variation induced by a planet, and because the orbital period is close to the supposed rotational period of the parent star, we remain cautious about this third candidate. In the frame of our program we will continue observing this star to confirm the planetary nature of the c component if possible.

The semi-major axis of the planetary orbits are 0.12 AU, 0.20 AU and 0.35 AU. For the minimum mass of the planets we obtain  $m_b \sin i = 2.68 [M_{\oplus}]$ ,  $m_c \sin i = 2.38 [M_{\oplus}]$  and  $m_d \sin i = 4.72 [M_{\oplus}]$ . We end up with three super-Earths in the same system. The error on the mass owing to only the fit uncertainties is on the order of 10% and 15%, respectively, very similar to the error set by the uncertainty on the stellar mass. It is worth noting that this system is of exactly the nature we were looking for, i.e., low-mass (possibly rocky) planets on not too close orbits. However, even for the furthest planet d, the equilibrium temperature defined by a Bond albedo of 0.3 is on the order of 388 K (see [Selsis et al. 2007](#), for details on the definition and estimation of equilibrium temperature).

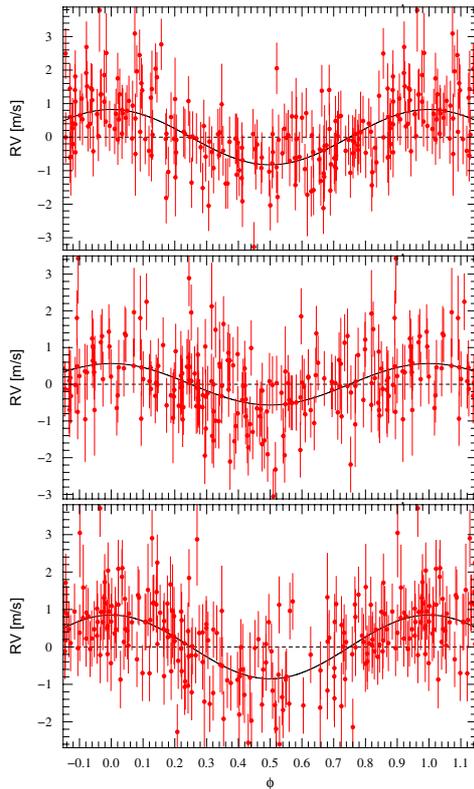
## 5.2. The planetary system around HD 85512

The raw radial-velocity data of the star HD 85512 are presented in Table 8 and shown in Fig. 11. By initial choice this star is among the most stable stars of the original HARPS high-precision sample. The dispersion of the 185 radial-velocity data points is, with  $1.05 \text{ m s}^{-1}$  rms over the 7.5-years time span, below that of HD 20794. The typical internal error on the radial velocity data points is of  $0.3 \text{ m s}^{-1}$ , to which  $0.7 \text{ m s}^{-1}$  have been

**Table 4.** Orbital and physical parameters of the planets orbiting HD 20794 as obtained from a three-Keplerian fit to the data.

Parameter	[Unit]	HD 20794 b	HD 20794 c	HD 20794 d
Epoch	[BJD]	2 454 783.40362208		
$i$	[deg]	90 (fixed)		
$V$	[km s <sup>-1</sup> ]	87.9525 ( $\pm 0.0001$ )		
$P$	[days]	18.315 ( $\pm 0.008$ )	40.114 ( $\pm 0.053$ )	90.309 ( $\pm 0.184$ )
$\lambda$	[deg]	169.0 ( $\pm 6.7$ )	149.4 ( $\pm 10.0$ )	16.2 ( $\pm 6.8$ )
$e$		0.0 (fixed)	0.0 (fixed)	0.0 (fixed)
$\omega$	[deg]	0.0 (fixed)	0.0 (fixed)	0.0 (fixed)
$K$	[m s <sup>-1</sup> ]	0.83 ( $\pm 0.09$ )	0.56 ( $\pm 0.10$ )	0.85 ( $\pm 0.10$ )
$m \sin i$	[ $M_{\oplus}$ ]	2.7 ( $\pm 0.3$ )	2.4 ( $\pm 0.4$ )	4.8 ( $\pm 0.6$ )
$a$	[AU]	0.1207 ( $\pm 0.0020$ )	0.2036 ( $\pm 0.0034$ )	0.3499 ( $\pm 0.0059$ )
$T_{\text{eq}}$	[K]	660	508	388
$N_{\text{meas}}$		187		
Span	[days]	2610		
rms	[m s <sup>-1</sup> ]	0.82		
$\chi_r^2$		1.39		

**Notes.** Error bars are derived from the covariance matrix.  $\lambda$  is the mean longitude ( $\lambda = M + \omega$ ) at the given epoch.

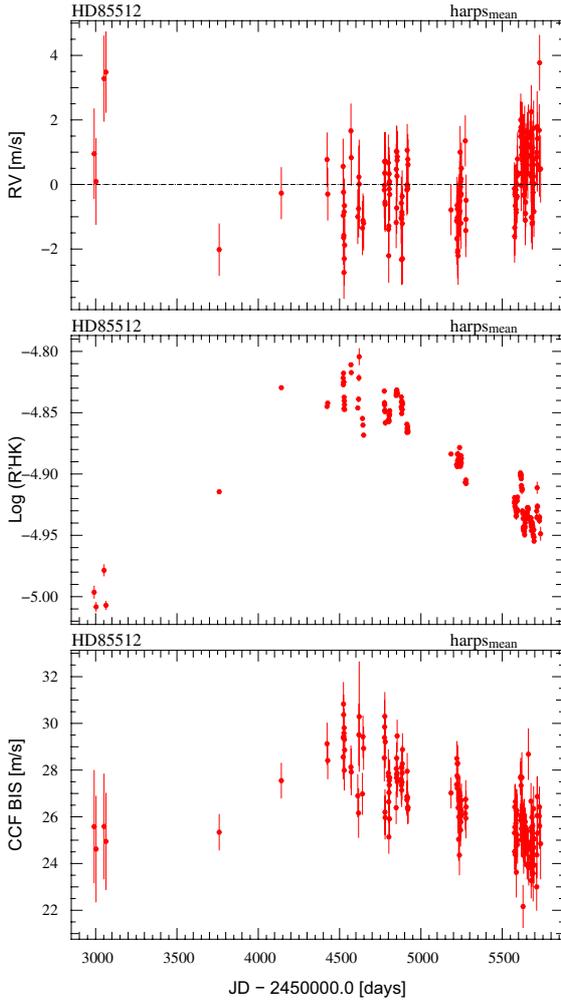


**Fig. 10.** Phase-folded RV data and fitted Keplerian solution for the three planetary components b, c, and d of HD 20794 with  $P = 18$  days,  $P = 40$  days, and  $P = 90$  days, respectively. The dispersion of the residuals is  $0.82 \text{ m s}^{-1}$  rms.

added quadratically to take into account possible instrumental and stellar noise.

The GLS periodogram shows several peaks well above noise at periods around 60–70 days (see Fig. 12). The dashed and the dotted lines indicate a false-alarm probability of 1% and 10%, respectively. The mentioned peaks clearly exceed this level. The three highest peaks are found at 58, 69, and 85 days period, which are the yearly aliases of each other. The peak at 58 days is the strongest of the series, however. Furthermore, there is significant excess power around 300 days period at above. In the same figure we plot for comparison the GLS periodogram of the activity indicator  $\log R'_{\text{HK}}$  and of the line bisector  $BIS$ . The former does not show any excess power at periods between 50 and 100 days. Instead, the excess power observed at periods above 300 days in the radial velocities seems to follow the same pattern. The reason for this pattern can be directly seen in the time domain (see Fig. 11). In both  $\log R'_{\text{HK}}$  and line bisector  $BIS$  a long-term trend is observed, which most probably reflects the magnetic cycle of the star. A similar trend, although with opposite sign and smaller amplitude, is observed in the radial-velocity data. From this we deduce that the power observed in the radial velocity at long periods is most probably caused by the stellar activity. We refer to Lovis et al. (2011a) for a general discussion, in which the specific case of HD 85512 is treated as well.

Given the absence of corresponding peaks in the activity indicator and the line bisector, we assume that the radial velocity signals at periods around 60 days are induced by a planetary companion. We therefore fitted a single Keplerian to the radial-velocity data, but also added a second-degree polynomial to take into account the long-term variation induced by the stellar activity. The solution converges very rapidly and unambiguously towards a 58.4 days orbit. The activity signal is fitted well by a linear component of  $0.5 \text{ m s}^{-1}$  per year and a quadratic component of  $0.14 \text{ m s}^{-1}$  per year<sup>2</sup>. The GLS periodogram of the residuals is very clean, as shown in Fig. 14, because none of the peaks reaches the 10% false-alarm probability indicated by the dashed line. It must be noted that all original side peaks have

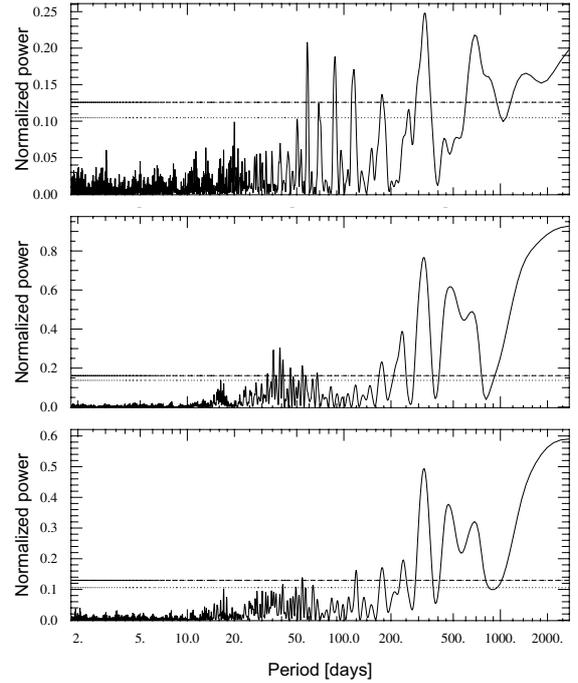


**Fig. 11.** Radial velocity,  $\log(R'_{\text{HK}})$  and BIS of HD 85512 as a function of time. The dispersion of the 185 data points is only  $1.05 \text{ m s}^{-1}$  rms.

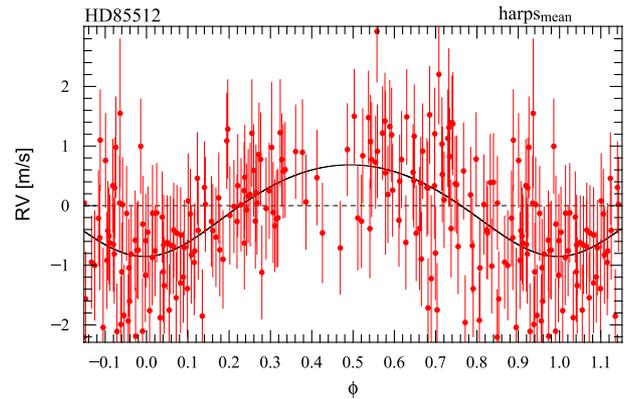
disappeared and that our assumption that they are aliases must be correct.

The eventually obtained orbital parameters of HD 85512 b are given in Table 5 and the corresponding phase-folded radial-velocity curve is shown in Fig. 13. The adopted period is  $P = 58.4$  days. The eccentricity is clearly not significant, but given the large number of data points and the good convergence of the fit, we preferred to leave it as free parameter. The obtained value is of 0.11, but we recall that setting it to zero would not change the quality of the solution at all. For the semi-amplitude we obtain the value of  $0.77 \text{ m s}^{-1}$ . The residual dispersion of the  $N = 185$  data points around the fitted solution is only  $0.77 \text{ m s}^{-1}$ . The reduced  $\chi^2$  is thus close to 1, indicating that the  $0.7 \text{ m s}^{-1}$  systematic error we added quadratically to the error bars may be a good estimate.

Using the parameters determined above we can set the semi-major axis of the planetary orbit to 0.26 AU. For the minimum mass of the planet we obtain  $m \sin i = 3.6 [M_{\oplus}]$ . The error on the mass due to only the fit uncertainties is on the order of 15%, despite the very low semi-amplitude. As in the case of the planets around HD 20794, this planet is possibly a super-Earth, i.e., a planet consisting to high degree of rocky material. Opposite to HD 20794, however, its parent star is much less luminous. The computed equilibrium temperature is only 298 K if we again



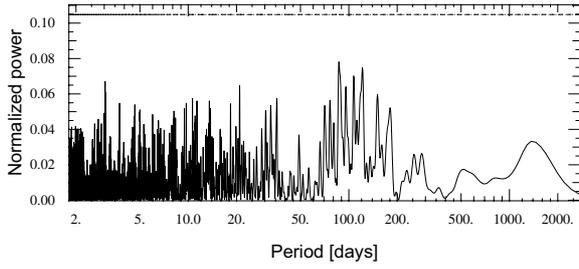
**Fig. 12.** GLS periodogram of the HD 85512 radial-velocity data with FAP levels at 10% and 1% level, respectively (top), the de-trended activity indicator  $\log(R'_{\text{HK}})$  (center) and the line bisector (bottom).



**Fig. 13.** Phase-folded RV data of HD 85512 and fitted Keplerian solution. The dispersion of the residuals is  $0.75 \text{ m s}^{-1}$  rms.

assume an albedo of 0.3. This sets this possibly rocky planet at the inner edge of the habitable zone of its parent star.

To illustrate the cleanliness of the planetary radial-velocity signal and its poor correlation with the activity indicator, we de-trended both series from the respective long-period signals. The results are best represented by the GLS periodogram in Fig. 15. The top panel shows the radial-velocity signal, which appears even more clearly after detrending. The peak at 58 days is by far the strongest and the aliases show up at the expected periods with decreasing amplitude the farther away from the main peak they are. No such “coherent” signal is observed in the  $\log(R'_{\text{HK}})$  at the mentioned periods, although significant power is still present at various (shorter) periods. This excess power is also seen in the time series of  $\log(R'_{\text{HK}})$ , but the signal’s amplitude, frequency and its phase is varying as a function of the observation epoch. We conclude that this signal is the expression of stellar rotation coupled with appearing and disappearing spots and plages, but that it cannot explain the presence of a strong and coherent



**Fig. 14.** GLS periodogram of the residuals to the  $P = 58$  days best Keplerian fit to the HD 85512 RV data.

**Table 5.** Orbital and physical parameters of the planets orbiting HD 85512 as obtained from a Keplerian fit to the data.

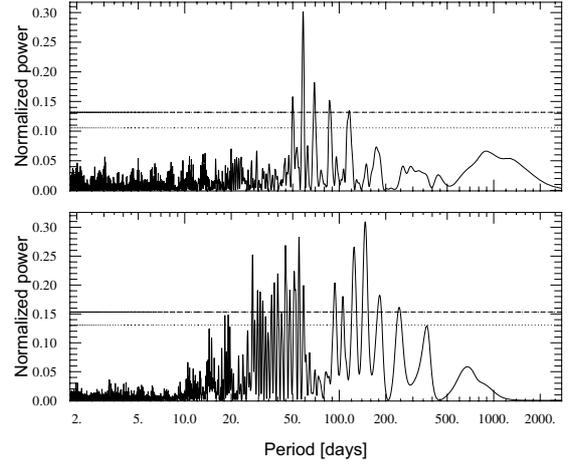
Parameter	[Unit]	HD 85512 b
Epoch	[BJD]	2 455 239.70888680
$i$	[deg]	90 (fixed)
$V$	[km s <sup>-1</sup> ]	-9.4913 ( $\pm 0.0038$ )
linear	[m s <sup>-1</sup> /yr]	+0.5085 ( $\pm 0.0003$ )
quadratic	[m s <sup>-1</sup> /yr <sup>2</sup> ]	+0.1423 ( $\pm 0.0002$ )
$P$	[days]	58.43 ( $\pm 0.13$ )
$\lambda$	[deg]	114.5 ( $\pm 6.3$ )
$e$		0.11 ( $\pm 0.10$ )
$\omega$	[deg]	178 ( $\pm 0.58$ )
$K$	[m s <sup>-1</sup> ]	0.769 ( $\pm 0.090$ )
$m \sin i$	[ $M_{\oplus}$ ]	3.6 ( $\pm 0.5$ )
$a$	[AU]	0.26 ( $\pm 0.005$ )
$T_{\text{eq}}$	[K]	298
$N_{\text{meas}}$		185
Span	[days]	2745
rms	[m s <sup>-1</sup> ]	0.77
$\chi_r^2$		1.00

**Notes.** Error bars are derived from the covariance matrix.  $\lambda$  is the mean longitude ( $\lambda = M + \omega$ ) at the given epoch.

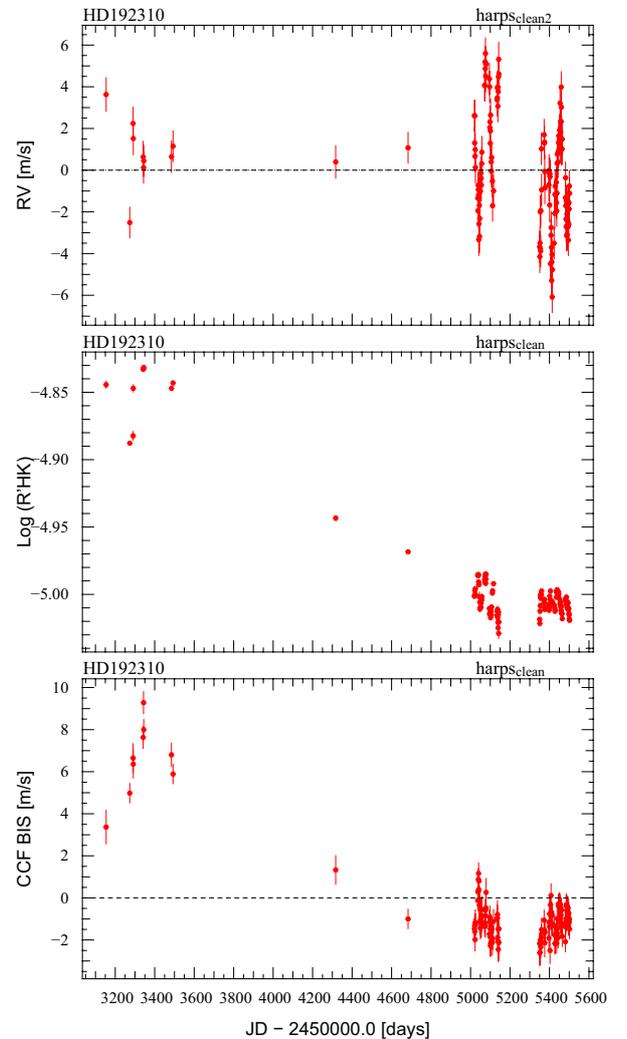
radial-velocity. This latter is therefore most likely caused by a planetary companion.

### 5.3. The planetary system around HD 192310

Recently, Howard et al. (2011a) have announced the discovery of a Neptune-mass planet around HD 192310. Because it is part of the original HARPS GTO program and of our program, we have been observing this star for several seasons. The raw radial-velocity data of HD 192310 are presented in Table 9 and shown in Fig. 16. What strikes one immediately is that, contrary to the previous objects presented in this paper, the dispersion of the 139 radial-velocity data points is  $2.6 \text{ m s}^{-1}$  rms over the 6.5-year time span. The typical photon noise on the radial velocity data points is  $0.25 \text{ m s}^{-1}$ , to which, for the analysis, we added quadratically a noise of  $0.7 \text{ m s}^{-1}$  to account for possible instrumental and stellar noise.

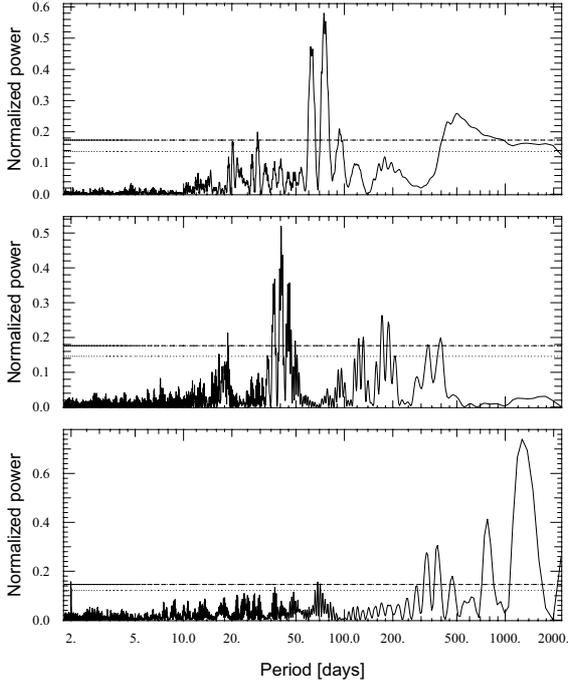


**Fig. 15.** *Top:* GLS periodogram of the de-trended radial velocities of HD 85512. *Bottom:* GLS periodogram of the de-trended activity indicator  $\log(R'_{\text{HK}})$ .



**Fig. 16.** Radial velocity,  $\log(R'_{\text{HK}})$  and  $BIS$  of HD 192310 as a function of time. The dispersion of the raw data is  $2.6 \text{ m s}^{-1}$  rms.

The dispersion clearly exceeds the internal error and the stellar jitter expected for this quiet K3 dwarf. Even by eye one can recognize a rapid radial-velocity variation. The GLS

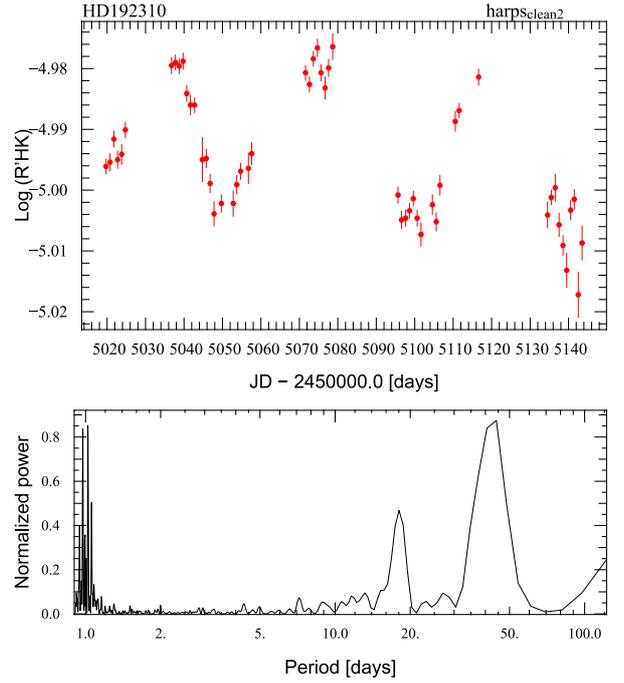


**Fig. 17.** GLS periodogram of the HD 192310 radial-velocity data with FAP levels at 10% and 1% level, respectively (*top*), the de-trended activity indicator  $\log(R'_{\text{HK}})$  (*center*) and the line bisector (*bottom*). While a very clear peak appears at  $P = 74$  days for the radial velocities, no such signal is present in the other two observables.

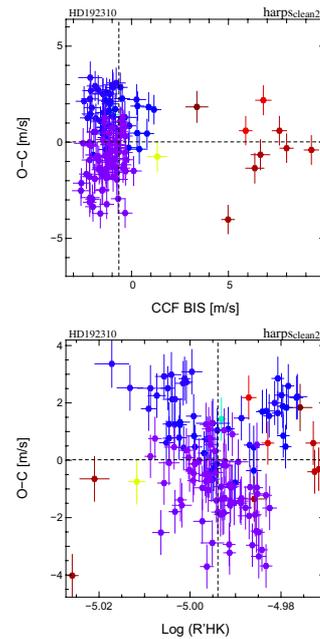
periodogram (Fig. 17) confirms the presence of a strong signal at about 74 days, which is exactly the period of the planet announced by Howard et al. (2011a). The dashed line indicates a false-alarm probability level of 1% but the  $P = 74$  days signal exceeds this level by far. In the same figure we plot for comparison the GLS periodogram of the activity indicator  $\log R'_{\text{HK}}$  and of the line bisector *BIS*. None of them reveals any significant power at the period indicated by the radial velocity.

The highest peak in the GLS periodogram for  $\log(R'_{\text{HK}})$  would be observed at very long periods ( $>1000$  days) and is produced by the stellar magnetic cycle. We therefore de-trended the data to make short-period variations appear more clearly. Some excess power then becomes visible around 40 days. A closer look at the last observational season indeed reveals a well-established periodicity in the  $\log(R'_{\text{HK}})$  time series (Fig. 18, top). The GLS periodogram of only these two seasons indeed reveals a series of significant peaks centered at 40 to 45 days (Fig. 18, bottom). It should be noted that this peak is most likely linked to the rotational period of the star and therefore delivers a direct and precise measurement of the stellar rotational period. Indeed, the period obtained in this way matches the value that is computed by directly using the value of  $\log(R'_{\text{HK}})$  and that is indicated in Table 3. On the contrary, a complete lack of power is observed at the  $P = 74$  day period, so that we can exclude any relation with the radial-velocity peak.

We therefore fitted a Keplerian signal to the raw data and obtained a solution almost identical to that presented by Howard et al. (2011a). However, we considered this to be an intermediate solution, as suggested by the GLS periodogram of the residuals to the one-Keplerian solution (Fig. 20). Indeed, a significant power peak is left over at a period around 500 days. Although there is no direct correspondence with the peaks in the bisector and  $\log(R'_{\text{HK}})$  power spectra, one may suspect that



**Fig. 18.** *Top:*  $\log(R'_{\text{HK}})$  time series of the two last observational seasons of HD 192310. *Bottom:* corresponding GLS periodogram.



**Fig. 19.** Correlation plot of the residuals of the one-Keplerian fit to the radial velocities of HD 192310 versus the line bisector and the activity indicator  $\log(R'_{\text{HK}})$ .

the residual radial-velocity signal is produced by stellar activity. Therefore, we tested the correlation between the residuals to the one-Keplerian solution, on the one hand, and the line bisector and the  $\log(R'_{\text{HK}})$ , on the other hand. This time the raw (not de-trended)  $\log(R'_{\text{HK}})$  data were used to identify possible long-period correlations. Figure 19 shows the complete absence of correlation between these quantities, the Pearson correlation coefficient being close to zero in both cases. We conclude that the residual radial velocity variations are probably caused by a second planetary companion at longer period. A closer look at

**Table 6.** Orbital and physical parameters of the planets orbiting HD 192310 as obtained from a two-Keplerian fit to the data.

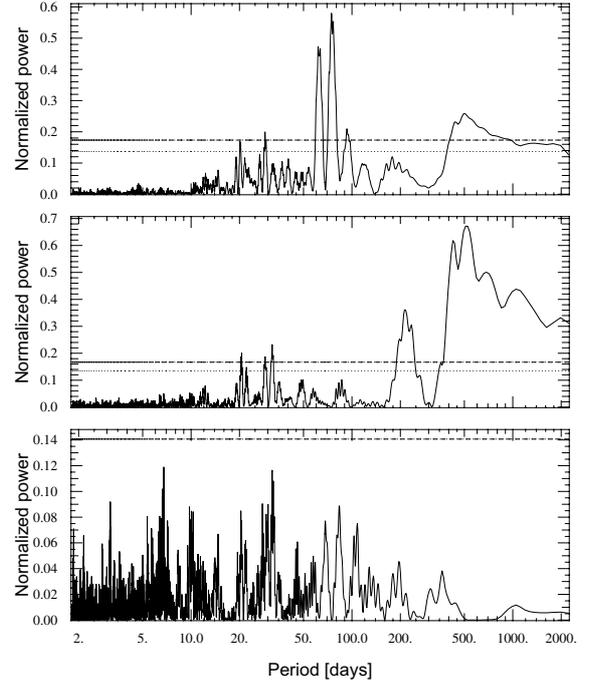
Parameter	[Unit]	HD 192310 b	HD 192310 c
Epoch	[BJD]	2 455 151.02574596	
$i$	[deg]	90 (fixed)	
$V$	[km s <sup>-1</sup> ]	−54.2232 (±0.0003)	
$P$	[days]	74.72 (±0.10)	525.8 (±9.2)
$\lambda$	[deg]	340.8 (±2.3)	1.9 (±12.3)
$e$		0.13 (±0.04)	0.32 (±0.11)
$\omega$	[deg]	173 (±20)	110 21
$K$	[m s <sup>-1</sup> ]	3.00 (±0.12)	2.27 (±0.28)
$m \sin i$	[ $M_{\oplus}$ ]	16.9 (±0.9)	24 (±5)
$a$	[AU]	0.32 (±0.005)	1.18 (±0.025)
$T_{\text{eq}}$	[K]	355	185
$N_{\text{meas}}$		139	
Span	[days]	2348	
rms	[m s <sup>-1</sup> ]	0.92	
$\chi_r^2$		1.66	

**Notes.** Error bars are derived from the covariance matrix.  $\lambda$  is the mean longitude ( $\lambda = M + \omega$ ) at the given epoch.

the raw radial-velocity data of Fig. 16 reveals a significant decrease in radial-velocity between the two last seasons, which, on the other hand, both show a trend upward. This should have suggested a second periodicity of 400 to 600 days.

Based on the above considerations, we fitted two Keplerians to the radial velocities of HD 192310 letting all the parameters free. The solution converges very rapidly and unambiguously towards a period of about 74.7 days and another of 526 days. The respective semi-amplitudes are 3.0 m s<sup>-1</sup> and 2.3 m s<sup>-1</sup>, and the eccentricities 0.13 and 0.32. The eccentricity of the second planet is significant at a 3- $\sigma$  level, and we therefore decided to let both the eccentricities as free parameter. The eventually obtained orbital parameters are given in Table 6 and the corresponding phase-folded radial-velocity curve is shown in Fig. 21. The GLS of the residuals to the two-Keplerian does not show any significant peak (lower panel of Fig. 20).

For the  $N = 139$  data points we obtain a residual dispersion around the two-Keplerians solution of 0.92 m s<sup>-1</sup>. The reduced  $\chi^2$  is 1.66. The uncertainty on the orbital period and the semi amplitude for the  $P = 74$  days planet result to be very low. For the longer-period planet, the error is somewhat larger on the semi-amplitude, which is mainly linked to the uncertainty on the eccentricity, which is in turn given by the lack of data points in the (short) descending slope of the radial-velocity curve. Period and eccentricity of the b component are, within the error bars, identical to the values given Howard et al. (2011a). The semi-amplitude – and thus the companion’s mass – is different at a level of 2–3 sigmas though. This difference may be explained, on the one hand, because the second planet had not been detected by Howard et al. (2011a). On the other hand, it may simply be caused by the higher uncertainty (0.4 m s<sup>-1</sup> on  $K$  given by Howard et al. 2011a), which is in turn due to the higher errors on the HIRES data: indeed, when fitting two planets to both data sets simultaneously by adding to both a systematic



**Fig. 20.** From top to bottom the GLS periodograms of a) the RVs of HD 192310, b) the residuals to the one-Keplerian solution and c) the residuals to the two-Keplerian solution. FAP levels of 10% and 1% level are indicated except for the last plot, where only the 10% level is shown.

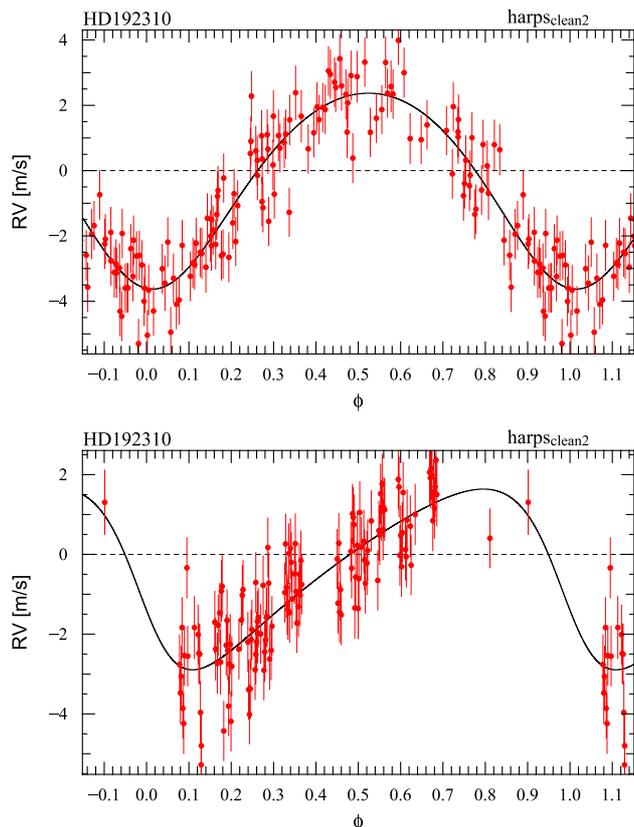
error of 1.5 m s<sup>-1</sup>, the residual dispersion of the HIRES data is 1.8 m s<sup>-1</sup>, while for the HARPS data the dispersion is only 1 m s<sup>-1</sup>. We conclude that we can undoubtedly confirm the planet HD 192310b announced by Howard et al. (2011a) but suggest a correction of the semi-amplitude and thus its minimum mass.

The semi-major axis of the planetary orbits are 0.32 AU and 1.18 AU. For the minimum mass of the planets we obtain  $m_b \sin i = 16.9 [M_{\oplus}]$  and  $m_c \sin i = 23.4 [M_{\oplus}]$ . The error on the mass due to only the fit uncertainties is on the order of 6% and 20%, respectively. While for the 74-day period planet the mass is quite well-determined thanks to the significant RV amplitude, for the long-period planet the uncertainty is fairly large, mainly because of a lack of coverage over the steeply descending orbital phase portion. Given their mass and period, both planets are probably of Neptune’s composition. The equilibrium temperature defined by a Bond albedo of 0.3 is on the order of 355 K and 185 K, which places the two planets at the inner, respectively outer edge of the habitable zone of HD 192310.

## 6. Discussion

### 6.1. Frequency of low-mass planets

By down-selecting our small sample of stars in terms of measured “raw” RV dispersion, we have explicitly excluded the possibility of finding planets typically found by ongoing RV surveys. We have reduced the probability of finding any kind of planet in our program. On the other hand, the absence of more massive planets on short and intermediate orbits leaves free space for low-mass planets, possibly in the habitable zone, which is precisely what we are looking for. We then optimized the observations to increase the probability for success by choosing the best possible instrument, selecting bright and inactive stars, and by optimizing the observation sampling and time span. Finally it should be noted that when down-selecting from the original



**Fig. 21.** Phase-folded RV data of HD 192310 and fitted Keplerian solution for the two planetary companions with  $P = 75$  days and  $P = 526$  days, respectively. The dispersion of the residuals is of  $0.91 \text{ m s}^{-1}$  rms.

HARPS high-precision sample, we did not select on spectral type or other criteria that may favor or penalize the presence of planets.

Under these conditions we have found that at least three out of the ten stars observed within our program harbor low-mass planets. Although statistics is poor over only ten targets, it is interesting to note that this 30% value was already announced by Lovis et al. (2009), who based their analysis on the larger ( $\sim 200$  stars) HARPS high-precision program. Theoretical works by Mordasini et al. (2009) actually forecasted that the frequency of small Neptunes and super-Earths on short and intermediated orbits would be considerably higher than that of Saturns and Jupiters. The recent amazing discoveries made by the KEPLER satellite using the transit technique further strengthen this. Borucki et al. (2011) report that the probability of finding low-mass planets is considerably higher than for Jupiter or Saturn-mass planets. Furthermore, when summing up the frequency of finding a planet of any mass, these authors end up with a probability of about 30%, again in perfect agreement with the results of Lovis et al. (2009). On the other hand, there is partial disagreement between the HARPS results on the one side and the results by Howard et al. (2010) and Howard et al. (2011b) on the other side, who both claim that the total planet occurrence is on the order of 15%. It has to be mentioned, however, that the studies were restricted to stars with a given spectral type and to planets with an upper period limit, and that the detection biases for low planetary masses and radii leave quite some room for making the different values compatible.

The fact that these low-mass planets were hardly ever found by the RV technique before was probably simply due to a

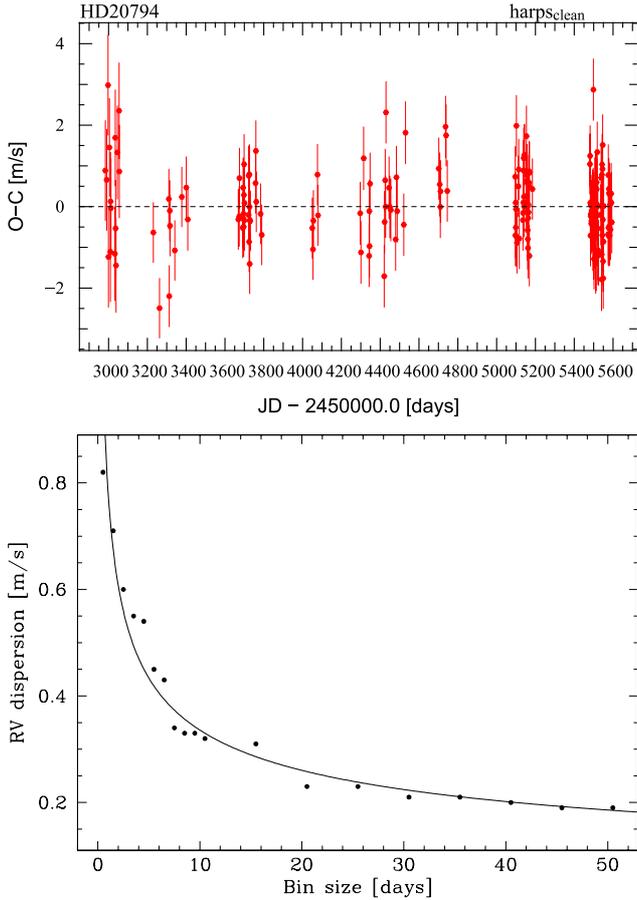
“threshold effect” arising from the reduced precision and sometimes poor time sampling of existing instruments, which had a hard time in detecting signals of few  $\text{m s}^{-1}$  or even  $\text{sub-m s}^{-1}$ . HARPS has produced impressive results by exploring this previously unexplored domain. The number and quality of new low-mass planets is amazing, and many more are still to come. Indeed, the detection of these low signals is not only a matter of precision, but also of observation strategy and time span. Furthermore, complementary analysis is particularly important if one wants to exclude that the RV signal is produced by stellar noise. Systematic long-term follow-up is required. We refer here to Mayor et al. (in prep.) for an overview and a summary of the outcomes of the HARPS program as a whole.

## 6.2. Nature of found planets

Radial velocities only provide minimum mass of the planets and contain no information on their radii. The mean density – and consequently an indication on the composition of the planets – can only be determined by combining RVs and transit measurements. Some specific discoveries, as for instance GJ 1214 b (Charbonneau et al. 2009), COROT-7 b (Léger et al. 2009; Queloz et al. 2009), KEPLER-10 b (Batalha et al. 2011), have shown that these low-mass planets have mean densities similar to our Earth and that they are probably made of rocky material. On the other hand, the atmosphere of planets on such short orbit (days) may have been evaporated by the star, contrary to planets on wider orbits. KEPLER discoveries of small planets on longer-period orbit (Lissauer et al. 2011) have indeed shown that their average density is generally lower than in the previously mentioned cases. One possible explanation is that they may have at least kept a thick gaseous envelope. By construction of our program, we have found planets of low-mass and on not too short-period orbits. Using the previous references we conclude that the presented planets are probably rocky or at least composed to a large extent of rocky material, but may have a thick atmosphere. An exception in our sample is HD 192310, which represents instead a “classical” system of two Neptune-mass planets on long-period orbits.

In the case of HD 20794 all found planets have masses in the domain of super-Earths and are probably rocky. Unfortunately, the period of the outermost planet is slightly too short to place it in the habitable zone of its star. Nevertheless, this particular example illustrates best the potential of our program and the fact that already today it is possible to find rocky planets in the habitable zone of solar-type stars.

The system of HD 85512 is much simpler in terms of dynamics, given the single planet, and does therefore not deserve any particular discussion. For this K5 dwarf the habitable zone is much closer in than for HD 20794. The planet’s equilibrium temperature supposing an albedo of 0.3 is on the order of 298 K. This value is slightly higher than that of the Earth and places the planet just at the inner edge of the habitable zone (Selsis et al. 2007). On the other hand, we might expect a thick atmosphere for this planet. If the albedo were higher, the equilibrium temperature would decrease. However, depending on the atmosphere’s composition, the planet may also run into a runaway greenhouse effect (Selsis et al. 2007). Opposite, if its atmosphere is water-poor (in the case of a “dry planet”), the habitable zone happens to be much closer in (Abe et al. 2011). Because the detail depends on many unknown parameters, it would be pure speculation to conclude on the habitability of this planets. Nevertheless, we note that HD 85512 is the lowest-mass planet discovered by RVs so far possibly lying in the habitable zone of its star.



**Fig. 22.** *Top:* residuals of the radial velocities of HD 20794 to the three-Keplerian fit as a function of time. *Bottom:* Allen analysis of the dispersion of these same residuals over the past two years.

### 6.3. Limits and possibilities of the ongoing program

The HARPS spectrograph allows us to find planetary companions to solar-type stars that induce sub meter-per-second radial velocity signatures. This proves two facts: a) instrument precision is on short *and* long term below  $1 \text{ m s}^{-1}$ ; b) for well-selected stars the stellar “jitter”, i.e., the sum of pulsation, granulation at all time scales, and activity, does not exceed  $1 \text{ m s}^{-1}$ .

To illustrate the achieved overall precision, we have analyzed the residuals to the orbital solution of HD 20794 in more detail. Figure 22 (upper panel) shows the residuals of the data to the three-Keplerian solution. Although some structure is identified in the early phases (data from the original HARPS GTO program), no secular trend at the sub- $\text{m s}^{-1}$  level is found over seven years of observations. It has to be mentioned that the observation strategy has been improved and the time sampling increased on the mentioned object in the frame of the HARPS-Upgrade GTO, which was specifically designed to obtain the best possible RV performance. The results are best illustrated by (Fig. 22, lower panel), where the dispersion of the residuals is shown as a function of binning period for the last two years of observation (optimized strategy). For no binning (one-day bin) we obtain the mentioned scatter of  $0.82 \text{ m s}^{-1}$  rms. When the binning is increased, the dispersion decreases with the bin size, demonstrating that most of the noise is indeed averaged away. The noise decreases with the power of  $-0.37$ , however. Therefore it is different from “white noise”, which is expected to decrease as the square root of the bin size. This shows that the

noise is not completely uniformly distributed over the frequency domain. Nevertheless, we end up with a scatter of only  $0.2 \text{ m s}^{-1}$  when averaging over the timescale of a month.

It is interesting to note that the stellar noise has a “color”, which reveals itself best in the GLS periodograms of  $\log(R'_{\text{HK}})$  and *BIS*. If detected in the RV data (by comparing the frequency with  $\log(R'_{\text{HK}})$  and *BIS* data, it could be “filtered”, or, in the worst case, used to identify radial-velocity signals of same frequency that are possibly induced by the star. In the former case one could benefit from the fact that the noise is colored to improve the RV data and reduce the error bars on the planets fitted parameters. This is easily done when the frequency of the planet signal is significantly different from the frequency detected in  $\log(R'_{\text{HK}})$  and *BIS*, supposing that possible stellar jitter would appear at similar frequencies (or harmonics of them) as for the activity indicator and the line bisector. Filtering of the unwanted frequencies could be performed in the Fourier domain, for instance. Alternatively, we propose to fit the planet orbital signal convolved with the window function directly in the Fourier domain.

As mentioned above and as we did in the present paper, the activity indicator and the line bisector may be used for “monitoring” and detecting possible correlations among the various observables. If we assume that RV jitter caused by stellar noise should appear at the same frequencies (or harmonics) as for the activity indicator and the line bisector, we can use this information to detect “suspect” signals. It must be pointed out, however, that both indicators remain a diagnostic rather than a proof for the presence or absence of induced RV signal, and this is best illustrated by two extreme examples: a) if no signal is detected in the *BIS*, one cannot exclude that the observed RV signal would *not* be induced by the star, as demonstrated by Santos et al. (2003); b) the cases of HD 85512 and HD 192310 show very nicely that despite the presence of a long-term signal in the  $\log(R'_{\text{HK}})$ , no correlated signal is observed in the RV. This confirms the conclusions drawn by Santos et al. (2010), who did not find any evidence for radial-velocity variations induced by variations of the stellar magnetic cycle with amplitudes significantly above  $1 \text{ m s}^{-1}$  for early K-type stars.

$\log(R'_{\text{HK}})$  and *BIS* are thus excellent diagnostic parameters, provided that they are determined to sufficient precision level. Thanks to the spectroscopic stability of HARPS, both can be determined to extreme accuracy. For the line bisector we demonstrated that it can be even more stable than the RV itself because it is only affected by asymmetry changes. The  $\log(R'_{\text{HK}})$ , on the other hand, can be measured to the milli-dex precision, a level at which it is hard to distinguish stellar effects from instrument effects. An interesting side product of this is the direct determination of the stellar rotational velocity, which, for all the three cases presented here, turned out to agree excellently with the value directly derived from  $\log(R'_{\text{HK}})$  using the procedures given by Mamajek & Hillenbrand (2008). The changes made recently on the HARPS data-reduction pipeline have greatly contributed to the improved spectroscopic performances of HARPS (Lovis et al. 2011a).

## 7. Conclusions

Of the original HARPS high-precision program we have chosen ten bright and quiet stars to be followed-up in the frame of the HARPS-Upgrade GTO program. These stars were furthermore selected because they were showing very low radial-velocity dispersion, from which we could exclude the presence of massive

and/or short-period planets. The goal of this program is to search for low-mass, possibly rocky planets in the habitable zone.

Despite this severe “down-selection”, we have found that at least three out of the ten stars harbor low-mass planetary companions. Around HD 192310 we could confirm the Neptune-mass planets recently announced by Howard et al. (2011a), but identified a second companion of similar mass on a much wider orbit and lying at the outer edge of the habitable zone. In addition, a planet of only  $m \sin i = 3.6 [M_{\oplus}]$  was found to orbit the star HD 85512. Owing to the late spectral type of this star, the  $P = 58$  day-orbit planet of probably rocky nature lies at the inner edge of its habitable zone. Finally, we reported the system around HD 20794 consisting of three super-Earths of only few Earth-masses each. While the two signals at an 18- and 90-day period leave little doubt on their planetary nature, we will need more data to definitively confirm the  $m \sin i = 2.4 [M_{\oplus}]$  on the  $P = 40$  days orbit. It must be noted, however, that the semi-amplitude induced by this planet on its parent star is only  $0.56 \text{ m s}^{-1}$ , the lowest planetary signature ever detected. All these discoveries were made possible thanks to the ultra-precise radial velocities and the impressive measurements of the line bisector and the  $\log(R'_{\text{HK}})$  delivered by the HARPS instrument.

Present and future discoveries of such low-mass planets will be of great importance per se, pushing the variety of known extrasolar planets one step further. They will furthermore offer, thanks to the proximity and magnitude of their parent star, an excellent opportunity for precise and in-depth photometric, polarimetric, spectroscopic and astrometric follow-up of the planet and its host star. A bunch of exciting and unprecedented observational data will follow, providing new input on the formation models and the understanding of the internal structure and the atmosphere of these planets. Finally, these candidates will be among the best targets to be considered by future space missions, which will need appropriate targets to look at and will have to take the best advantage of observational time.

*Acknowledgements.* We are grateful to all technical and scientific collaborators of the HARPS Consortium, ESO Head Quarter, and ESO La Silla, who have all contributed with their extra-ordinary passion and valuable work to the success of the HARPS project. This research made use of the SIMBAD database, operated at the CDS, Strasbourg, France. N.C.S. acknowledges the support by the European Research Council/European Community under the FP7 through Starting Grant agreement number 239953. N.C.S. also acknowledges the support from Fundação para a Ciência e a Tecnologia (FCT) through program Ciência2007 funded by FCT/MCTES (Portugal) and POPH/FSE (EC), and in the form of grants reference

PTDC/CTE-AST/098528/2008 and PTDC/CTE-AST/098604/2008. F.B. wish to thank the French National Research Agency (ANR-08-JCJC-0102-01. Finally, we acknowledge the support by Swiss National Science Foundation for its continuous support in our research.

## References

- Abe, Y., Abe-Ouchi, A., Sleep, N. H., & Zahnle, K. J. 2011, *Astrobiology*, 11, 443
- Alibert, Y., Mordasini, C., & Benz, W. 2011, *A&A*, 526, A63
- Andrae, R., Schulze-Hartung, T., & Melchior, P. 2010 [arXiv:1012.3754]
- Batalha, N. M., Borucki, W. J., Bryson, S. T., et al. 2011, *ApJ*, 729, 27
- Borucki, W. J., Koch, D. G., Basri, G., et al. 2011, *ApJ*, 728, 117
- Butler, R. P., Vogt, S. S., Marcy, G. W., et al. 2004, *ApJ*, 617, 580
- Casagrande, L., Portinari, L., & Flynn, C. 2006, *MNRAS*, 373, 13
- Charbonneau, D., Berta, Z. K., Irwin, J., et al. 2009, *Nature*, 462, 891
- Dumusque, X., Santos, N. C., Udry, S., Lovis, C., & Bonfils, X. 2011a, *A&A*, 527, A82
- Dumusque, X., Udry, S., Lovis, C., Santos, N. C., & Monteiro, M. J. P. F. G. 2011b, *A&A*, 525, A140
- ESA 1997, The HIPPARCOS and TYCHO catalogue, ESA-SP 1200
- Howard, A. W., Marcy, G. W., Johnson, J. A., et al. 2010, *Science*, 330, 653
- Howard, A. W., Johnson, J. A., Marcy, G. W., et al. 2011a, *ApJ*, 730, 10
- Howard, A. W., Marcy, G. W., Bryson, S. T., et al. 2011b, *ApJ*, submitted [arXiv:1103.2541]
- Ida, S., & Lin, D. N. C. 2008, *ApJ*, 685, 584
- Léger, A., Rouan, D., Schneider, J., et al. 2009, *A&A*, 506, 287
- Lissauer, J. J., Fabrycky, D. C., Ford, E. B., et al. 2011, *Nature*, 470, 53
- Lovis, C., Mayor, M., Pepe, F., et al. 2006, *Nature*, 441, 305
- Lovis, C., Mayor, M., Bouchy, F., et al. 2009, *IAU Symp.*, 253, 502
- Lovis, C., Dumusque, X., Santos, N. C., et al. 2011a, *A&A*, submitted [arXiv:1107.5325]
- Lovis, C., Ségransan, D., Mayor, M., et al. 2011b, *A&A*, 528, A112
- Mamajek, E. E., & Hillenbrand, L. A. 2008, *ApJ*, 687, 1264
- Mayor, M., Pepe, F., Queloz, D., et al. 2003, *The Messenger*, 114, 20
- Mayor, M., Bonfils, X., Forveille, T., et al. 2009a, *A&A*, 507, 487
- Mayor, M., Udry, S., Lovis, C., et al. 2009b, *A&A*, 493, 639
- McArthur, B. E., Endl, M., Cochran, W. D., et al. 2004, *ApJ*, 614, L81
- Mordasini, C., Alibert, Y., & Benz, W. 2009, *A&A*, 501, 1139
- Neves, V., Santos, N. C., Sousa, S. G., Correia, A. C. M., & Israelian, G. 2009, *A&A*, 497, 563
- Pepe, F., Correia, A. C. M., Mayor, M., et al. 2007, *A&A*, 462, 769
- Queloz, D., Bouchy, F., Moutou, C., et al. 2009, *A&A*, 506, 303
- Santos, N. C., Mayor, M., Naef, D., et al. 2002, *A&A*, 392, 215
- Santos, N. C., Udry, S., Mayor, M., et al. 2003, *A&A*, 406, 373
- Santos, N. C., Bouchy, F., Mayor, M., et al. 2004, *A&A*, 426, L19
- Santos, N. C., Gomes da Silva, J., Lovis, C., & Melo, C. 2010, *A&A*, 511, A54
- Selsis, F., Kasting, J. F., Levrard, B., et al. 2007, *A&A*, 476, 1373
- Sousa, S. G., Santos, N. C., Mayor, M., et al. 2008, *A&A*, 487, 373
- Torres, G., Andersen, J., & Giménez, A. 2010, *A&AR*, 18, 67
- Udry, S., Bonfils, X., Delfosse, X., et al. 2007, *A&A*, 469, L43
- Udry, S., & Mayor, M. 2008, in *ASP Conf. Ser.*, 398, ed. D. Fischer, F. A. Rasio, S. E. Thorsett, & A. Wolszczan, 13