

## Evidence of early disk-locking among low-mass members of the Orion Nebula Cluster (Corrigendum)

K. Biazzo<sup>1,2,3</sup>, C. H. F. Melo<sup>2</sup>, L. Pasquini<sup>2</sup>, S. Randich<sup>1</sup>, J. Bouvier<sup>4</sup>, and X. Delfosse<sup>4</sup>

<sup>1</sup> INAF – Osservatorio Astrofisico di Arcetri, Largo E. Fermi 5, 50125 Firenze, Italy

e-mail: kbiazzo@arcetri.astro.it

<sup>2</sup> ESO – European Southern Observatory, Karl-Schwarzschild-Str. 3, 85748 Garching bei München, Germany

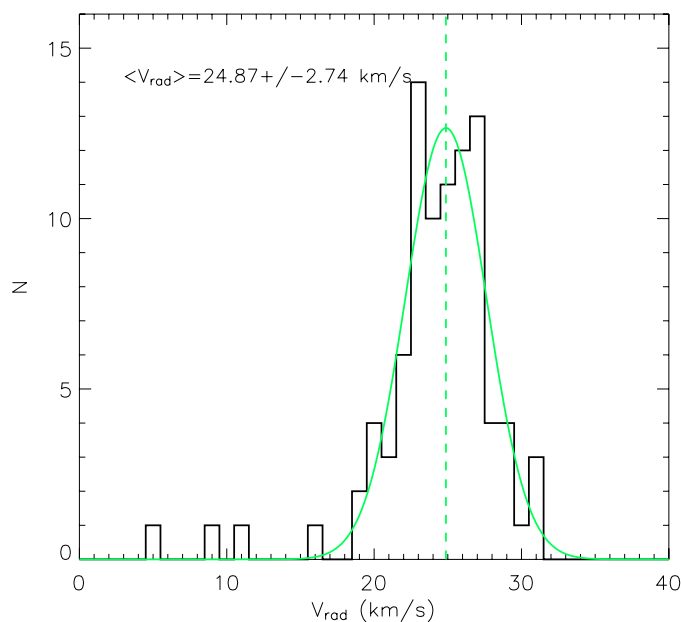
<sup>3</sup> INAF – Osservatorio Astrofisico di Catania, via S. Sofia 78, 95123 Catania, Italy

<sup>4</sup> Laboratoire d’Astrophysique, Observatoire de Grenoble, BP 53, 38041 Grenoble Cedex 9, France

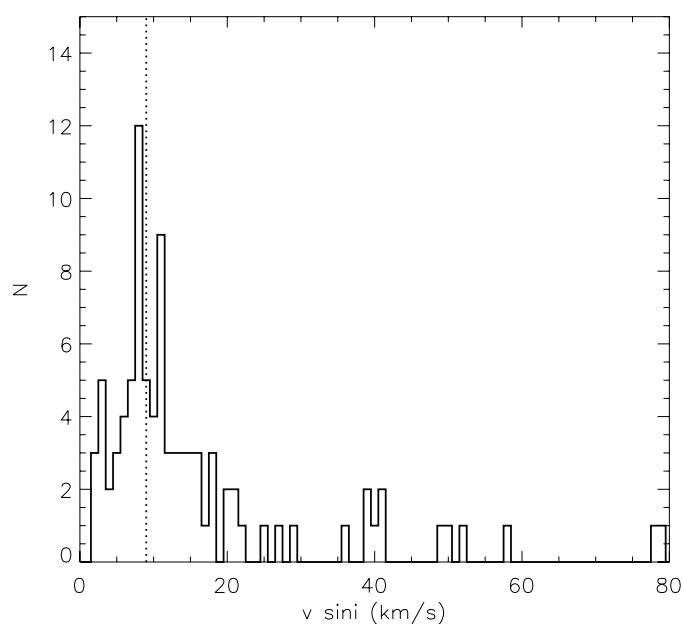
A&A, 508, 1301–1312 (2009), DOI: 10.1051/0004-6361/200913125

**Key words.** open clusters and associations: individual: Orion Nebula Cluster – stars: low-mass – stars: pre-main sequence – stars: late-type – accretion, accretion disks – errata, addenda

An error occurred during the production process. Figure 8 was published twice. The corrected Figs. 6 and 8 are published below.



**Fig. 6.** Radial velocity distribution for our sample of stars with  $M < 0.25 M_{\odot}$ . For the most probable binaries we plot the mean values of the two components (Table 2). The Gaussian fit of the distribution yields a mean  $V_{\text{rad}}$  of  $24.87 \text{ km s}^{-1}$  (with a  $\sigma_{V_{\text{rad}}}$  of  $2.74 \text{ km s}^{-1}$ ), which is fully consistent with the membership in the ONC (e.g., Stassun et al. 1999).



**Fig. 8.**  $v \sin i$  distribution for our low-mass stars. Our  $v \sin i$  limit of  $9 \text{ km s}^{-1}$  is shown with a dotted line.

### References

Stassun, K. G., Mathieu, R. D., Mazeh, T., & Vrba, F. J. 1999, AJ, 117, 2941