

Benchmarking atomic data for astrophysics: Fe VIII EUV lines

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ABSTRACT

The EUV spectrum of Fe VIII is reviewed, using new solar observations from the Hinode EUV Imaging Spectrometer (EIS), together with older solar and laboratory data. The most up-to-date scattering calculations are benchmarked against these experimental data, with the use of a large atomic structure calculation. Once adjustments are made to the excitation rates, good agreement is found between calculated and observed line intensities. All previous line identifications have been re-assessed. Several lines are identified here for the first time, most notably the strong decays from the $3s^2 3p^5 3d^2 \ ^4D_j$ levels. It is shown that they provide a new, important diagnostic of electron temperature for the upper transition region. The temperatures obtained at the base of solar coronal loops are lower ($\log T$ [K] = 5.5) than those predicted by assuming ionization equilibrium ($\log T$ [K] = 5.6), however firm measurements will only be possible once better scattering calculations are available, and the EIS radiometric calibration is properly assessed.

Key words. line: identification – atomic data – Sun: transition region – techniques: spectroscopic

1. Introduction

This paper is one of a series in which atomic data and line identifications are benchmarked against experimental data. The method and goals are described in Del Zanna et al. (2004), where Fe X was discussed. New solar observations from the Hinode EUV Imaging Spectrometer (EIS, see Culhane et al. 2007) have shown the presence of a wealth of spectral lines formed at transition region (hereafter TR) temperatures (Young et al. 2007), and offer the opportunity to benchmark atomic data for a variety of ions. The brightest TR lines in the EIS spectra are from Fe VIII, which are assessed in this paper. The assessment of Fe VII and in general of lines formed at similar temperatures is considered in a separate paper (Del Zanna 2009). The Hinode EIS observations are complemented with measurements found in the literature, and with a few new measurements from a plate (Fawcett, priv. comm.). Only EUV transitions (decays from $n = 3, 4$ levels) are considered here.

Contrary to many other ions, very little experimental work has been done on Fe VIII EUV lines. The identification of Fe VIII EUV lines started with Kruger & Weissberg (1937)[hereafter KW37], who correctly identified only two of the transitions considered here. Six of the strongest EUV transitions were identified through to the fundamental laboratory work of Fawcett and co-workers on the Zeta spectrum (see Gabriel et al. 1965, hereafter G65). The main source of accurate wavelengths and the identifications of most of the strongest EUV lines came from the excellent work of Ramonas & Ryabtsev (1980) (hereafter RR80). For this benchmark, we consider the most up-to-date electron scattering calculation for this ion by Griffin et al. (2000) (hereafter G00). For a description of previous work on electron impact excitation see G00.

2. Benchmark method and atomic structure calculations

As described in Del Zanna et al. (2004), the benchmark method starts with an assessment of the observed wavelengths λ_{obs} , from which a set of experimental level energies E_{exp} is obtained. As an aid to the identification process, atomic structure calculations are run, trying to find a set of ab-initio energies to match the experimental ones. To improve the level energies, term energy corrections (TEC) (see, e.g. Zeippen et al. 1977; Nussbaumer & Storey 1978) to the LS Hamiltonian matrix are applied, using an iterative procedure, where first-guess corrections to the LS term energies are estimated from the weighted mean of the observed level energies, whenever available. A large “benchmark” basis is normally sought, and the results in terms of level energy and oscillator strengths compared to those from the scattering calculation, to assess the validity of the scattering target. Once a good match between experimental and theoretical energies is found, line intensities are calculated by solving the level balance equations in steady state, and all the strongest spectral lines (observed or not) are examined in detail. Observed line intensities I_{ob} , whenever available, are compared to the theoretical ones by plotting the “emissivity ratio”

$$F_{ji} = \frac{I_{\text{ob}} N_e}{N_j(N_e, T_e) A_{ji}} C \quad (1)$$

for each line as a function of the electron temperature T_e . $N_j(N_e, T_e)$ is the population of the upper level j relative to the total number density of the ion, and is calculated at a fixed density N_e assuming steady state conditions. A_{ji} is the spontaneous radiative transition probability, and C is a scaling constant chosen so the emissivity ratio is near unity. If agreement between experimental and theoretical intensities is present, all lines should

Table 1. Electron configuration basis for the benchmark calculation (23C) and orbital scaling parameters.

Configurations		Scaling parameters	
even	odd		
$3s^2 3p^6 3d$	$3s^2 3p^5 3d^2$	1s	1.41654
$3s^2 3p^6 4s$	$3s^2 3p^6 4p$	2s	1.12528
$3s^2 3p^6 4d$	$3s^2 3p^6 4f$	2p	1.06749
$3s^2 3p^6 \bar{5}s$	$3s^2 3p^6 \bar{5}p$	3s	1.12766
$3s^2 3p^6 \bar{5}d$	$3s^2 3p^5 3d 4s$	3p	1.10218
$3s^2 3p^5 3d 4p$	$3s^2 3p^5 3d 4d$	3d	1.11400
$3s^2 3p^4 3d^3$	$3s^2 3p^3 3d^4$	4s	1.16978
	$3s^2 3p^4 3d^2 4p$	4p	1.13576
	$3s^2 3p^4 3d^2 4f$	4d	1.14783
	$3s^2 3p^4 3d^2 \bar{5}p$	4f	1.65984
	$3s^2 3p^6 \bar{5}f$	$\bar{5}s$	-1.17688
	$3s^2 3p^4 3d^2 \bar{5}f$	$\bar{5}p$	-1.09068
	$3s^2 3p^3 3d^3 4s$	$\bar{5}d$	-0.91633
	$3s^2 3p^3 3d^3 4d$	$\bar{5}f$	-0.64713
	$3s^2 3p^3 3d^3 \bar{5}s$		
	$3s^2 3p^3 3d^3 \bar{5}d$		

be closely spaced. If the plasma is nearly isothermal, all curves should cross at the isothermal temperature. The level populations of Fe VIII do not vary with N_e for densities higher than 10^9 cm^{-3} , so emissivity ratios are in fact only dependent on temperature for typical coronal loops densities, as is the case examined here.

It is notoriously difficult to obtain ab-initio level energies that match the observed ones for this ion. Configuration-interaction (CI) effects are very large. For the same reason, it is very difficult to obtain firm identifications for all the strongest lines. Relativistic multi-reference many-body perturbation theory calculations such as those described in [Ishikawa & Vilkas \(2008\)](#) are needed, since they have been proved to provide very good level energies.

The G00 scattering target included 33 *LS* terms from the seven configurations: $3s^2 3p^6 3d$, $3s^2 3p^5 3d^2$, $3s^2 3p^5 3d 4s$, $3s^2 3p^6 4s$, $3s^2 3p^6 4p$, $3s^2 3p^6 4d$, $3s^2 3p^6 4f$, which are the most important ones for UV and EUV lines. For the radiative rates of a few important transitions, G00 performed a larger CI structure calculation by including terms from the $3s^2 3p^4 3d^3$, $3s^2 3p^3 3d^4$, $3s^2 3p^4 3d^2 4f$ configurations, which mix strongly with the $3s^2 3p^6 3d$, $3s^2 3p^5 3d^2$, $3s^2 3p^6 4f$. As clearly shown in their Table 2, the radiative rates for strong dipole-allowed transitions reached 30% differences with the values obtained from the target used for the scattering calculation. This implies that the collisions strengths are likely to be incorrect by the same amount. The combined effect of collisions strength (which changes the population of the upper level) and of the transition probability therefore becomes significant.

[Zeng et al. \(2003\)](#) (hereafter Z03) performed large-scale multi-configuration Hartree-Fock (MCHF) calculations for Fe VIII, to show the importance of including core-valence electron correlations. In their case D (which included correlations of the type $3p^3-3d^3$), convergence in the oscillator strengths was achieved, however their level energies were not very accurate.

A ‘‘benchmark’’ configuration basis was sought, to obtain a good set of oscillator strengths and level energies. After a large number of tests, the set of 23 configurations listed in Table 1 was chosen. The AUTOSTRUCTURE code ([Badnell 1997](#)) was used. Radial scaling parameters were chosen by minimizing the equally-weighted sum of the energies of the *LS* terms providing

the lowest configurations considered here. $n = 5$ correlation orbitals were included. To further improve the energies, term energy corrections (TEC) to the *LS* Hamiltonian matrix were applied.

The energies of the benchmark calculation ($E_{\text{Bench.}}$) are in good agreement with the experimental ones ($E_{\text{Exp.}}$), in particular in terms of relative energies for strongly-mixed levels, as shown in Table 2. The levels for which it proved most difficult to obtain accurate energies were: a) the $3s^2 3p^5 3d^2 {}^2P_{3/2}$ (41) which mixes strongly with $3s^2 3p^6 4p {}^2P_{3/2}$ (43) and $3s^2 3p^5 3d^2 {}^2P_{3/2}$ (48); b) $3s^2 3p^6 4p {}^2P_{3/2}$ (43); c) $3s^2 3p^5 3d^2 {}^2P_{1/2}$ (44) which mixes strongly with the $3s^2 3p^5 3d^2 {}^2P_{1/2}$ (47) and $3s^2 3p^5 3d^2 {}^2P_{1/2}$ (23). The energies from the G00 scattering target (E_{Griffin}) and those from Z03 (case D) are also shown in Table 2. Notice that most E_{Griffin} values are accurate, but those from Z03 are not.

Table 3 lists the weighted oscillator strengths (gf) for the strongest dipole-allowed transitions, compared to the G00 and Z03 values. The G00 gf values were obtained from the published A-values and the experimental energies, following G00. The 30% differences in the gf values of some of the strongest transitions discussed by G00 are confirmed, however larger differences are present for weaker transitions. The few published gf values from Z03 are in overall agreement with those of the benchmark calculation.

The differences in the gf values imply that the G00 collision strengths at higher temperatures for various transitions are likely to be incorrect by similar amounts. It is well known that the main contribution for strong dipole-allowed transitions comes from high partial waves, where the collision strength is approximately proportional to the gf value for the transition. For the purpose of this paper, it is useful to scale the G00 collision strengths using the benchmark calculation, to see how the resulting line intensities compare with the observed ones. [Flower & Nussbaumer \(1974\)](#) were among the first to suggest scaling collision strengths according to gf (and energy) values, something which has been adopted in many cases within the CHIANTI database (e.g. for Fe X and Fe XI). The scaling has been found to improve agreement between observed and predicted intensities.

The energies and gf values were used to obtain the high-temperature limits in the scaled domain following [Burgess & Tully \(1992\)](#). An example is shown in Fig. 1 (bottom plot), where the scaled collision strengths for the 2–46 transition (dashed line, constant equal to 1.5) and its high-temperature limit are plotted. Good agreement between the G00 collision strengths and the high-temperature limits was found for all lines. The high-temperature limit for the 2–46 transition obtained from the benchmark calculation is however smaller by a factor of 0.7, and is shown in Fig. 1 (bottom plot, scaled temperature = 1). The collision strengths need to be scaled (continuous line) by this factor to converge to this limit, as shown in the Figure. The ratios between the limit values obtained from the benchmark calculation and those obtained from the G00 data were used to scale the thermally-averaged collision strengths for all the dipole-allowed transitions.

Line intensities were calculated with the scaled rates and the transition probabilities from the benchmark + TEC calculation, assuming plasma equilibrium conditions, at the temperature of peak ion abundance for Fe VIII in ionization equilibrium ($\log T[\text{K}] = 5.6$), according to the latest ionization and recombination rates published within CHIANTI¹ v.6 ([Dere et al. 1997, 2009](#)). These line intensities are listed in Table 3, in decreasing order. As usual within the benchmark process, all the

¹ <http://www.chianti.rl.ac.uk>

Table 2. Level energies for Fe VIII.

i	Conf.	Lev.	E_{exp}	$E_{\text{Bench.}+\text{TEC}}$	$E_{\text{Bench.}}$	E_{Griffin}	E_{Z03}	
1	3s ² 3p ⁶ 3d	² D _{3/2}	0.000	0.000	0.000	0.000	0.000	
2	3s ² 3p ⁶ 3d	² D _{5/2}	0.017	0.018 (−0.001)	0.018 (−0.001)	0.018 (−0.001)	0.018 (−0.001)	
3	3s ² 3p ⁵ 3d ²	⁴ D _{1/2}	3.564	3.564 (0)	3.531 (0.033)	3.509 (0.055)	3.613 (−0.049)	N
4	3s ² 3p ⁵ 3d ²	⁴ D _{3/2}	3.572	3.572 (0)	3.540 (0.032)	3.518 (0.054)	3.625 (−0.053)	N
5	3s ² 3p ⁵ 3d ²	⁴ D _{5/2}	3.585	3.585 (0)	3.553 (0.032)	3.532 (0.054)	3.642 (−0.056)	N
6	3s ² 3p ⁵ 3d ²	⁴ D _{7/2}	3.605	3.605 (0)	3.573 (0.032)	3.553 (0.052)	3.661 (−0.056)	N
7	3s ² 3p ⁵ 3d ²	⁴ G _{11/2}	–	3.726	3.726	3.721	3.846	
8	3s ² 3p ⁵ 3d ²	⁴ G _{9/2}	–	3.741	3.741	3.735	3.874	
9	3s ² 3p ⁵ 3d ²	⁴ G _{7/2}	–	3.760	3.760	3.754	3.897	
10	3s ² 3p ⁵ 3d ²	⁴ P _{5/2}	–	3.774	3.774	3.763	3.845	
11	3s ² 3p ⁵ 3d ²	⁴ G _{5/2}	–	3.780	3.781	3.773	3.915	
12	3s ² 3p ⁵ 3d ²	⁴ P _{3/2}	–	3.802	3.802	3.791	3.876	
13	3s ² 3p ⁵ 3d ²	⁴ P _{1/2}	–	3.823	3.823	3.811	3.895	
14	3s ² 3p ⁵ 3d ²	⁴ F _{3/2}	–	3.874	3.875	3.864	3.983	
15	3s ² 3p ⁵ 3d ²	⁴ F _{5/2}	–	3.885	3.885	3.876	3.982	
16	3s ² 3p ⁵ 3d ²	⁴ F _{7/2}	–	3.893	3.893	3.881	3.981	
17	3s ² 3p ⁵ 3d ²	⁴ F _{7/2}	–	3.901	3.901	3.889	3.982	
18	3s ² 3p ⁵ 3d ²	² D _{5/2}	–	3.931	3.930	3.921	4.001	
19	3s ² 3p ⁵ 3d ²	² F _{5/2}	3.930	3.928 (0.002)	3.934 (−0.004)	3.927 (0.003)	3.984 (−0.054)	
20	3s ² 3p ⁵ 3d ²	² D _{3/2}	–	3.944	3.948	3.937	4.031	
21	3s ² 3p ⁶ 4s	² S _{1/2}	–	4.223	4.223	3.951	–	
22	3s ² 3p ⁵ 3d ²	² F _{7/2}	3.960	3.962 (−0.002)	3.968 (−0.008)	3.960 (0.000)	4.024 (−0.064)	
23	3s ² 3p ⁵ 3d ²	² P _{1/2}	4.009	4.008 (0.001)	4.033 (−0.024)	4.013 (−0.003)	4.057 (−0.048)	N
24	3s ² 3p ⁵ 3d ²	² P _{3/2}	4.060	4.062 (−0.002)	4.082 (−0.022)	4.062 (−0.003)	4.084 (−0.025)	N
25	3s ² 3p ⁵ 3d ²	² H _{11/2}	–	4.069	4.069	4.069	4.173	
26	3s ² 3p ⁵ 3d ²	² F _{7/2}	4.079	4.081 (−0.002)	4.076 (0.003)	4.088 (−0.009)	4.170 (−0.090)	
27	3s ² 3p ⁵ 3d ²	² G _{7/2}	4.118	4.119 (−0.001)	4.120 (−0.002)	4.112 (0.006)	4.182 (−0.064)	
28	3s ² 3p ⁵ 3d ²	² H _{9/2}	–	4.142	–	4.138	4.224	
29	3s ² 3p ⁵ 3d ²	² G _{9/2}	–	4.167	–	4.160	4.264	
30	3s ² 3p ⁵ 3d ²	² F _{5/2}	4.186	4.186 (0)	4.181 (0.005)	4.196 (−0.010)	4.257 (−0.071)	
31	3s ² 3p ⁵ 3d ²	⁴ D _{7/2}	–	4.206	4.206	4.207	4.265	
32	3s ² 3p ⁵ 3d ²	⁴ D _{5/2}	–	4.215	4.218	4.220	4.287	
33	3s ² 3p ⁵ 3d ²	⁴ D _{3/2}	–	4.233	4.235	4.237	4.305	
34	3s ² 3p ⁵ 3d ²	⁴ D _{1/2}	–	4.253	4.253	4.254	4.316	
35	3s ² 3p ⁵ 3d ²	² D _{3/2}	–	4.364	4.403	4.409	4.432	
36	3s ² 3p ⁵ 3d ²	² D _{5/2}	4.399	4.399 (0)	4.439 (−0.039)	4.445 (−0.046)	4.462 (−0.062)	N
37	3s ² 3p ⁵ 3d ²	⁴ S _{3/2}	4.416	4.416 (0)	4.460 (−0.044)	4.469 (−0.053)	4.492 (−0.076)	
38	3s ² 3p ⁵ 3d ²	² S _{1/2}	–	4.449	4.449	4.469	4.466	
39	3s ² 3p ⁵ 3d ²	² G _{9/2}	–	4.464	4.493	4.498	4.541	
40	3s ² 3p ⁵ 3d ²	² G _{7/2}	4.468	4.469 (0)	4.499 (−0.031)	4.504 (−0.036)	4.556 (−0.087)	N
41	3s ² 3p ⁵ 3d ²	² P _{3/2}	4.634	4.640 (−0.006)	4.694 (−0.060)	4.714 (−0.080)	4.723 (−0.089)	
42	3s ² 3p ⁶ 4p	² P _{1/2}	4.650	4.661 (−0.011)	4.716 (−0.066)	4.724 (−0.075)	–	
43	3s ² 3p ⁶ 4p	² P _{3/2}	4.698	4.704 (−0.006)	4.762 (−0.064)	4.779 (−0.081)	–	
44	3s ² 3p ⁵ 3d ²	² P _{1/2}	4.746	4.745 (0.001)	4.801 (−0.055)	4.828 (−0.082)	4.831 (−0.085)	
45	3s ² 3p ⁵ 3d ²	² F _{5/2}	4.884	4.885 (−0.002)	4.901 (−0.017)	5.084 (−0.200)	5.000 (−0.116)	
46	3s ² 3p ⁵ 3d ²	² F _{7/2}	4.937	4.936 (0.001)	4.951 (−0.015)	5.135 (−0.198)	5.046 (−0.109)	
47	3s ² 3p ⁵ 3d ²	² P _{1/2}	5.394	5.394 (0)	5.408 (−0.014)	5.610 (−0.216)	5.528 (−0.133)	
48	3s ² 3p ⁵ 3d ²	² P _{3/2}	5.423	5.423 (0)	5.437 (−0.014)	5.637 (−0.214)	5.556 (−0.133)	
49	3s ² 3p ⁵ 3d ²	² D _{5/2}	5.435	5.437 (−0.001)	5.474 (−0.038)	5.723 (−0.287)	5.590 (−0.154)	
50	3s ² 3p ⁵ 3d ²	² D _{3/2}	5.441	5.439 (0.002)	5.476 (−0.035)	5.724 (−0.283)	5.592 (−0.151)	
51	3s ² 3p ⁶ 4d	² D _{3/2}	5.957	5.957 (0)	6.382 (−0.425)	6.080 (−0.123)	–	
52	3s ² 3p ⁶ 4d	² D _{5/2}	5.961	5.962 (0)	6.387 (−0.425)	6.084 (−0.122)	–	
53	3s ² 3p ⁶ 4f	² F _{5/2}	6.959	6.959 (0.001)	7.139 (−0.179)	7.080 (−0.121)	–	
54	3s ² 3p ⁶ 4f	² F _{7/2}	6.960	6.961 (0)	7.140 (−0.180)	7.081 (−0.121)	–	

Note: the columns provide: the experimental level energies E_{exp} (Ryd), those obtained from the benchmark + TEC calculation $E_{\text{Bench.}+\text{TEC}}$, those from the benchmark $E_{\text{Bench.}}$ one, alongside with those of the G00 scattering target E_{Griffin} and those of Z03 (case D). Values in parentheses indicate differences to our observed energies. N indicates a new experimental energy.

identifications of the strongest lines have been checked, using laboratory and solar spectra, as described in the following section. Line intensities, whenever available, were compared, in order to confirm identifications and assess the possible presence

of blending. The results are also shown in Table 3. Line intensities were also computed adopting the original G00 collision strengths and A-values, with the addition of $A_{2,1} = 7.83 \times 10^{-2}$,

Table 3. List of strongest Fe VIII lines in the 100–500 Å range.

$i-j$	Levels	Int	gf	G00	Z03 D	$A_{ji}(s^{-1})$	$\lambda_{exp}(\text{Å})$	$\lambda_{th}(\text{Å})$	ID
2-49	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^2 D_{5/2}$	1.0	7.06	9.1	7.2	2.8×10^{11}	168.173	159.73	G65
2-46	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^2 F_{7/2}$	0.90	3.92	5.8	4.2	9.5×10^{10}	185.213	178.09	G65
1-50	$3p^6 3d^2 D_{3/2} - 3p^5 3d^2^2 D_{3/2}$	0.58	4.56	5.9	4.6	2.7×10^{11}	167.486	159.20	G65
2-48	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^2 P_{3/2}$	0.57	3.82	4.3	3.7	2.2×10^{11}	168.544	162.16	RR80
1-45	$3p^6 3d^2 D_{3/2} - 3p^5 3d^2^2 F_{5/2}$	0.58	2.80	4.1	2.9	9.0×10^{10}	186.599	179.25	G65
1-47	$3p^6 3d^2 D_{3/2} - 3p^5 3d^2^2 P_{1/2}$	0.29	2.10	2.4	2.0	2.4×10^{11}	168.929	162.43	RR80
2-43	$3p^6 3d^2 D_{5/2} - 3p^6 4p^2 P_{3/2}$	0.29	0.67	0.60	–	3.0×10^{10}	194.661	191.39	RR80
2-54	$3p^6 3d^2 D_{5/2} - 3p^6 4f^2 F_{7/2}$	0.14	4.30	4.0	–	2.1×10^{11}	131.240	129.02	KW37
1-42	$3p^6 3d^2 D_{3/2} - 3p^6 4p^2 P_{1/2}$	0.17	0.38	0.38	–	3.3×10^{10}	195.972	192.88	? RR80
1-53	$3p^6 3d^2 D_{3/2} - 3p^6 4f^2 F_{5/2}$	9.0×10^{-2}	2.97	2.8	–	1.9×10^{11}	130.941	128.71	KW37
2-26	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^2 F_{7/2}$	0.14	0.14	0.10	–	2.2×10^9	224.305	223.88	RR80
2-6	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^4 D_{7/2}$	0.15	2.1×10^{-4}	2.2×10^{-4}	–	2.8×10^6	253.956	257.79	N
2-22	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^2 F_{7/2}$	0.12	5.6×10^{-2}	5.2×10^{-2}	–	8.8×10^8	231.097	231.15	RR80
2-9	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^4 G_{7/2}$	9.9×10^{-2}	1.0×10^{-4}	7.0×10^{-5}	–	1.4×10^6	–	243.92	
2-17	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^4 F_{7/2}$	9.5×10^{-2}	8.1×10^{-4}	4.8×10^{-4}	–	1.2×10^7	–	235.41	
2-21	$3p^6 3d^2 D_{5/2} - 3p^6 4s^2 S_{1/2}$	8.9×10^{-2}	–	0.0	–	6.5×10^5	–	231.69	
1-30	$3p^6 3d^2 D_{3/2} - 3p^5 3d^2^2 F_{5/2}$	8.1×10^{-2}	0.12	0.10	–	2.9×10^9	217.691	217.17	RR80
1-49	$3p^6 3d^2 D_{3/2} - 3p^5 3d^2^2 D_{5/2}$	6.1×10^{-2}	0.43	0.57	0.51	1.7×10^{10}	167.655	159.24	G65
2-50	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^2 D_{3/2}$	5.6×10^{-2}	0.44	0.64	0.51	2.6×10^{10}	168.003	159.69	G65
1-48	$3p^6 3d^2 D_{3/2} - 3p^5 3d^2^2 P_{3/2}$	5.5×10^{-2}	0.37	0.48	0.41	2.2×10^{10}	168.024	161.65	RR80
2-5	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^4 D_{5/2}$	8.3×10^{-2}	2.4×10^{-4}	2.2×10^{-4}	–	4.0×10^6	255.350	259.31	N
2-24	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^2 P_{3/2}$	7.3×10^{-2}	2.2×10^{-2}	1.6×10^{-2}	2.1×10^{-2}	7.2×10^8	225.396	225.29	N
2-31	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^4 D_{7/2}$	7.0×10^{-2}	1.3×10^{-4}	1.1×10^{-4}	–	2.2×10^6	–	217.51	
2-16	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^4 F_{9/2}$	7.2×10^{-2}	4.5×10^{-12}	0.0	–	45.	–	235.87	
1-19	$3p^6 3d^2 D_{3/2} - 3p^5 3d^2^2 F_{5/2}$	6.7×10^{-2}	4.1×10^{-2}	3.2×10^{-2}	–	8.5×10^8	231.884	232.05	RR80
2-18	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^2 D_{5/2}$	6.7×10^{-2}	8.8×10^{-4}	1.6×10^{-6}	–	1.8×10^7	–	233.49	
2-27	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^2 G_{7/2}$	6.3×10^{-2}	1.1×10^{-2}	3.6×10^{-2}	–	1.8×10^8	222.190	222.57	N
2-41	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^2 P_{3/2}$	5.2×10^{-2}	9.9×10^{-2}	0.19	0.13	4.3×10^9	197.362	194.03	RR80
1-21	$3p^6 3d^2 D_{3/2} - 3p^6 4s^2 S_{1/2}$	5.9×10^{-2}	–	0.0	–	4.3×10^5	–	230.66	
1-36	$3p^6 3d^2 D_{3/2} - 3p^5 3d^2^2 D_{5/2}$	5.3×10^{-2}	3.7×10^{-3}	2.2×10^{-3}	–	9.5×10^7	207.124	205.01	N
2-10	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^4 P_{5/2}$	5.9×10^{-2}	1.3×10^{-4}	1.0×10^{-4}	–	2.5×10^6	–	243.30	
1-4	$3p^6 3d^2 D_{3/2} - 3p^5 3d^2^4 D_{3/2}$	5.1×10^{-2}	1.1×10^{-4}	1.1×10^{-4}	–	2.8×10^6	255.110	259.05	N
2-20	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^2 D_{3/2}$	4.5×10^{-2}	6.0×10^{-3}	3.3×10^{-3}	–	1.9×10^8	–	232.52	
2-40	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^2 G_{7/2}$	3.9×10^{-2}	2.2×10^{-2}	1.8×10^{-2}	–	4.3×10^8	204.704	203.12	N
1-15	$3p^6 3d^2 D_{3/2} - 3p^5 3d^2^4 F_{5/2}$	4.2×10^{-2}	5.8×10^{-3}	3.9×10^{-3}	–	1.2×10^8	–	235.12	
2-32	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^4 D_{5/2}$	3.5×10^{-2}	2.8×10^{-3}	1.6×10^{-3}	–	6.5×10^7	–	216.87	
1-14	$3p^6 3d^2 D_{3/2} - 3p^5 3d^2^4 F_{3/2}$	3.8×10^{-2}	5.7×10^{-3}	4.9×10^{-3}	–	1.7×10^8	–	235.83	
1-43	$3p^6 3d^2 D_{3/2} - 3p^6 4p^2 P_{3/2}$	3.1×10^{-2}	7.1×10^{-2}	6.1×10^{-2}	–	3.2×10^9	193.968	190.69	RR80
1-23	$3p^6 3d^2 D_{3/2} - 3p^5 3d^2^2 P_{1/2}$	3.4×10^{-2}	1.5×10^{-2}	1.0×10^{-2}	–	9.4×10^8	227.290	227.10	N
2-45	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^2 F_{5/2}$	2.5×10^{-2}	0.12	0.21	0.20	3.9×10^9	187.240	179.87	RR80
1-35	$3p^6 3d^2 D_{3/2} - 3p^5 3d^2^2 D_{3/2}$	2.2×10^{-2}	2.6×10^{-3}	1.9×10^{-3}	–	9.8×10^7	–	206.68	
2-11	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^4 G_{5/2}$	2.4×10^{-2}	5.5×10^{-5}	6.2×10^{-5}	–	1.1×10^6	–	242.64	
1-3	$3p^6 3d^2 D_{3/2} - 3p^5 3d^2^4 D_{1/2}$	2.5×10^{-2}	1.6×10^{-5}	1.8×10^{-5}	–	8.0×10^5	255.684	259.66	N
1-44	$3p^6 3d^2 D_{3/2} - 3p^5 3d^2^2 P_{1/2}$	1.6×10^{-2}	4.0×10^{-2}	5.9×10^{-2}	6.0×10^{-2}	3.6×10^9	192.004	188.74	? RR80
2-12	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^4 P_{3/2}$	1.9×10^{-2}	7.6×10^{-5}	1.2×10^{-4}	–	2.2×10^6	–	241.53	
2-25	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^2 H_{11/2}$	1.610^{-2}	1.0×10^{-10}	0.0	–	1.1	–	224.95	
2-33	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^4 D_{3/2}$	1.6×10^{-2}	2.9×10^{-3}	2.3×10^{-3}	–	1.0×10^8	–	215.96	
2-15	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^4 F_{5/2}$	1.6×10^{-2}	2.3×10^{-3}	2.6×10^{-3}	–	4.6×10^7	–	236.19	
2-28	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^2 G_{9/2}$	1.5×10^{-2}	8.2×10^{-11}	0.0	–	9.7	–	221.16	
2-37	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^4 S_{3/2}$	1.3×10^{-2}	7.9×10^{-3}	6.2×10^{-3}	–	3.1×10^8	207.124	204.74	
2-30	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^2 F_{5/2}$	1.3×10^{-2}	2.0×10^{-2}	1.8×10^{-2}	–	4.7×10^8	218.564	218.09	RR80
1-10	$3p^6 3d^2 D_{3/2} - 3p^5 3d^2^4 P_{5/2}$	1.4×10^{-2}	3.1×10^{-5}	2.7×10^{-5}	–	6.0×10^5	–	242.16	
1-13	$3p^6 3d^2 D_{3/2} - 3p^5 3d^2^4 P_{1/2}$	1.4×10^{-2}	1.5×10^{-5}	2.5×10^{-5}	–	8.7×10^5	–	239.13	
2-53	$3p^6 3d^2 D_{5/2} - 3p^6 4f^2 F_{5/2}$	6.5×10^{-3}	0.22	0.20	–	1.4×10^{10}	131.257	129.03	RR80
1-12	$3p^6 3d^2 D_{3/2} - 3p^5 3d^2^4 P_{3/2}$	1.2×10^{-2}	4.8×10^{-5}	5.1×10^{-5}	–	1.4×10^6	–	240.40	
1-41	$3p^6 3d^2 D_{3/2} - 3p^5 3d^2^2 P_{3/2}$	9.6×10^{-3}	1.8×10^{-2}	2.8×10^{-2}	1.5×10^{-2}	7.8×10^8	196.650	193.30	RR80
2-35	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^2 D_{3/2}$	8.7×10^{-3}	1.0×10^{-3}	6.8×10^{-4}	–	4.0×10^7	–	207.51	
1-34	$3p^6 3d^2 D_{3/2} - 3p^5 3d^2^4 D_{1/2}$	8.9×10^{-3}	2.7×10^{-3}	2.0×10^{-3}	–	1.9×10^8	–	214.21	
2-8	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^4 G_{9/2}$	9.6×10^{-3}	1.1×10^{-12}	0.0	–	0.82	–	245.12	
2-19	$3p^6 3d^2 D_{5/2} - 3p^5 3d^2^2 F_{5/2}$	8.9×10^{-3}	5.5×10^{-3}	6.0×10^{-3}	–	1.1×10^8	232.876	233.10	RR80

Notes: the relative intensities (photons) $\text{Int} = N_j A_{ji} / N_e$ are normalised to the strongest transition and were calculated at an electron density of 10^9 cm^{-3} and $\log T[\text{K}] = 5.6$. Weighted oscillator strengths gf and A-values (s^{-1}) are from the benchmark calculation. The gf values from G00 and the Z03 (case D) calculation are also listed. λ_{exp} are our experimental wavelengths, while λ_{th} are the theoretical ones from the G00 energies. The last column (ID) provides a key to previous identifications. N indicates a new one proposed here.

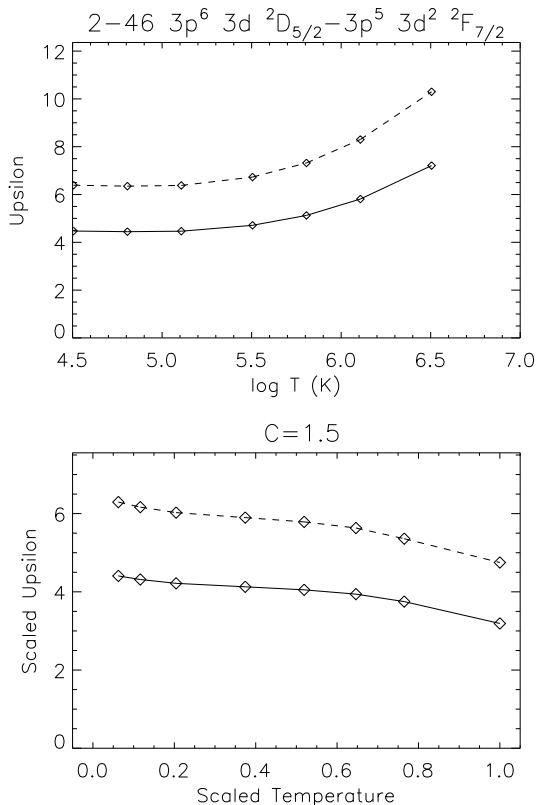


Fig. 1. *Top:* thermally-averaged collisions strengths for the 2–46 transition. The dashed line indicates the original calculated values from G00, while the continuous one the scaled values (see text). *Bottom:* the same points, plotted in the scaled domain (Burgess & Tully 1992), with the high-temperature limits (scaled temperature = 1).

a value calculated with the benchmark target and in close agreement with the NIST database² value of 7.05×10^{-2} .

3. Experimental data

The few spectral observations considered here for the benchmark are now described.

3.1. Hinode/EIS

The Hinode/EIS instrument covers two wavelength bands (SW: 166–212 Å; LW: 245–291 Å approximately). This spectral range is crowded with emission lines, and most spectral lines are significantly blended. The entire EIS database of observations was searched to find suitable observations to benchmark this ion. None was found, however an observation which was originally analysed for other purposes is presented here. It consists of a long-duration raster where the 1'' slit was moved (from west to east, between 2007 Jan. 5 21:52 UT and 2007 Jan. 6 01:07 UT) over a Sunspot while it was close to Sun centre. The exposure times were long (90 s) which allowed a good signal even in the weaker lines.

The observation is ideal in the sense that contamination from coronal lines is at a minimum, and Sunspot loops had prominent Fe VIII lines. The drawback of this observation was the small field of view, which considerably limits wavelength measurements. The data analysis to obtain line intensities is quite

straightforward, but obtaining accurate wavelengths is quite complex.

As described in Del Zanna (2008a), the main problems in the analysis of EIS data are the strong (75 km s^{-1}) orbital variation of the wavelength scale and the offsets in both N-S (18'') and E-W (2'') directions between the two channels. The offset in the E-W direction means that observations in the two channels were not simultaneous nor co-spatial. During the course of benchmarking Fe XVII, it was also found that the spectra are slanted relative to the axes of the CCD by $3.66(\pm 0.2)$ pixels end-to-end (each pixel along the slit corresponds to 1''). A small tilt is also present.

The spectra have been rotated and shifted to take into account the slant and the various misalignments. Two orbital-dependent wavelength calibrations (one for each channel) was obtained, and the spectral tilt included. These wavelengths were used, together with the *cfit* package (Haugan 1997) to fit Gaussian profiles to the brightest EIS lines. More than 200 lines were fitted and their morphology examined in detail, one by one. Figure 2 shows the resulting monochromatic images for a selection of Fe VIII lines, while Fig. 3 presents images in other lines, to show how sensitive morphology is to temperature (and to a lesser extent to density). This allows us to estimate the temperature of formation for unidentified lines, and to assess if Fe VIII lines are blended. For example, the 193.97 Å line is obviously blended with a hot line formed at temperatures close to that of Fe XVI 269.98 Å (cf. the similarities in the images in Figs. 2, 3).

Figure 5 shows an enlarged portion of the radiances of two lines recorded in the two EIS channels, showing agreement to within 1'' in the spatial alignment. This confirms that the geometrical corrections applied here are accurate. A region in a Sunspot loop leg, indicated by the crossing of the two sets of dashed lines in Fig. 5, was chosen to obtain an average spectrum to benchmark the Fe VIII lines. Portions of the spectrum where Fe VIII lines are present are shown in Fig. 4. In this spectrum, all TR lines are very strong. To account for the (small) contribution from the foreground emission in coronal lines, a nearby foreground spectrum was obtained, centred on the Sunspot, i.e. where TR and coronal lines are weak. Portions of this spectrum are also shown in Fig. 4. Notice that the strongest EIS coronal line, due to Fe XII (a self-blend at 195.12 Å identified by Del Zanna & Mason 2005), is very weak in the Sunspot spectrum. Moreover, that its ‘‘foreground’’ intensity is almost the same as that of the Sunspot loop leg. This means that the ‘‘foreground-subtracted’’ Sunspot loop leg spectrum is virtually free from coronal emission. This spectrum was used for the benchmark.

Another spectrum was needed for the wavelength calibration. The one chosen is also shown in Fig. 4. The EIS instrument does not have a reference wavelength scale so ideally one would use the spectra of a ‘‘reference’’ region to obtain the rest wavelengths. Unfortunately, as shown in Del Zanna (2007, 2008b), active regions present Doppler-shifts at all locations in all lines, in particular at TR temperatures, where Fe VIII is formed. This is clearly shown in Fig. 6. The Sunspot leg area selected presents a red-shift of about 20 km s^{-1} in all TR lines observed in both channels. To obtain good rest wavelengths, a large field of view would be needed, something that was not available. The shifts in the spectra due to thermal effects caused by the satellite orbit are difficult to correct accurately because they are usually wavelength dependent. A reference region was chosen so that it was observed at the same time as the loop leg region (i.e. the same wavelength calibration would apply) at coordinates Solar $Y = -84:-64$ (see Fig. 6).

² <http://physics.nist.gov/PhysRefData/ASD/index.html>

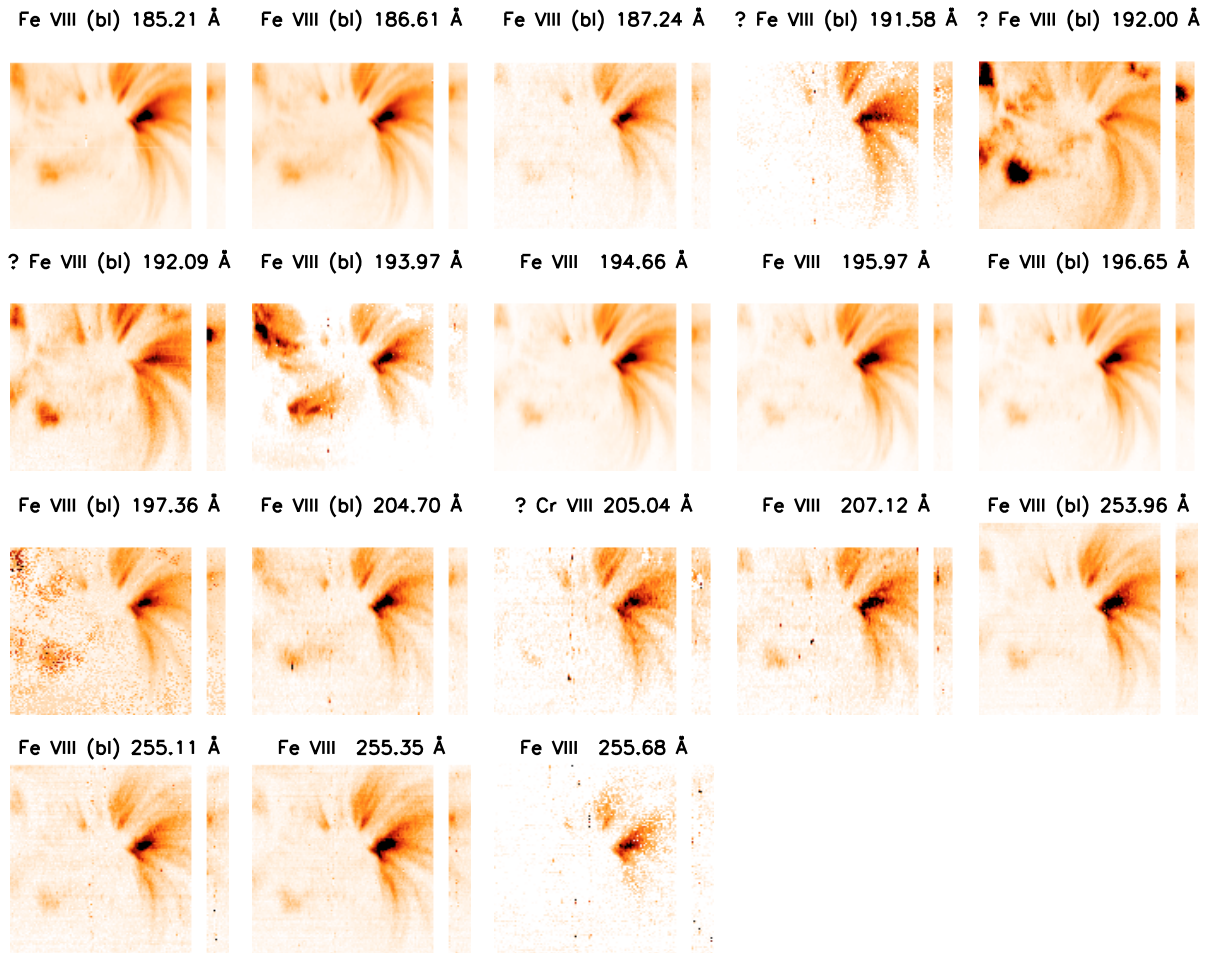


Fig. 2. Monochromatic images (negative) of the Fe VIII identified lines observed by EIS. Notice that all the Fe VIII lines have a similar morphology.

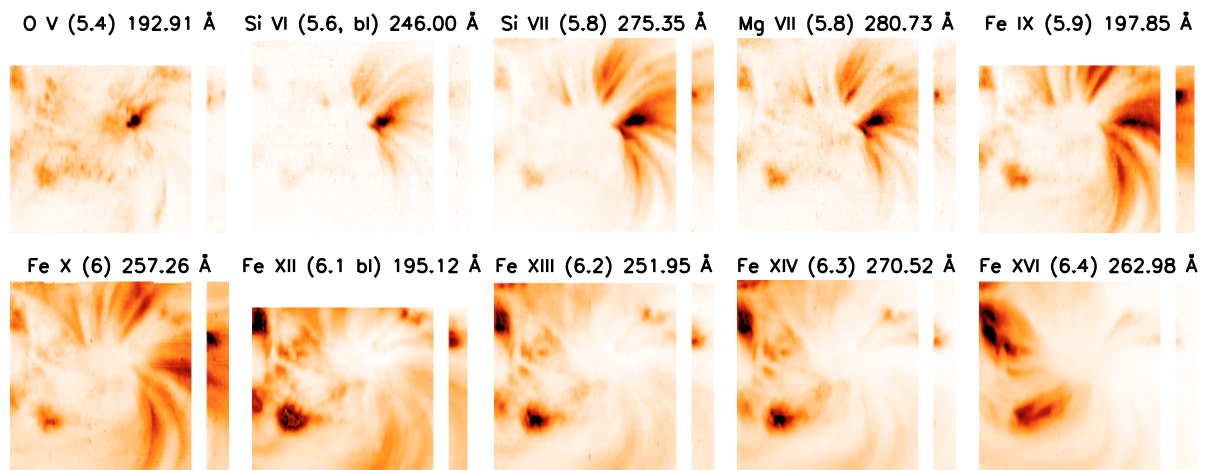


Fig. 3. Monochromatic images (negative) of a selection of EIS lines (ion and wavelength in Å are indicated together with the logarithm of the temperature (K) of formation of the lines in collisional ionization equilibrium). Notice the large morphological differences in lines formed at different temperatures.

The average reference spectra were calibrated in wavelength, using a set of about 30 lines in each channel. The spectral dispersion is such that the EIS wavelengths vary almost linearly with the CCD pixels. An EIS linear wavelength calibration was used for the SW channel, while a quadratic one was used for the LW. Results are shown in Fig. 7. Theoretical wavelengths

from version 5.2 of the CHIANTI atomic package (Landi et al. 2006) were used. These in turn rely mainly on the Behring et al. (1976) full-Sun grazing incidence rocket spectrum and various laboratory data, published in a series of papers by Fawcett. It is well known that such theoretical wavelengths need to be improved, and indeed such improved reference data are one of the

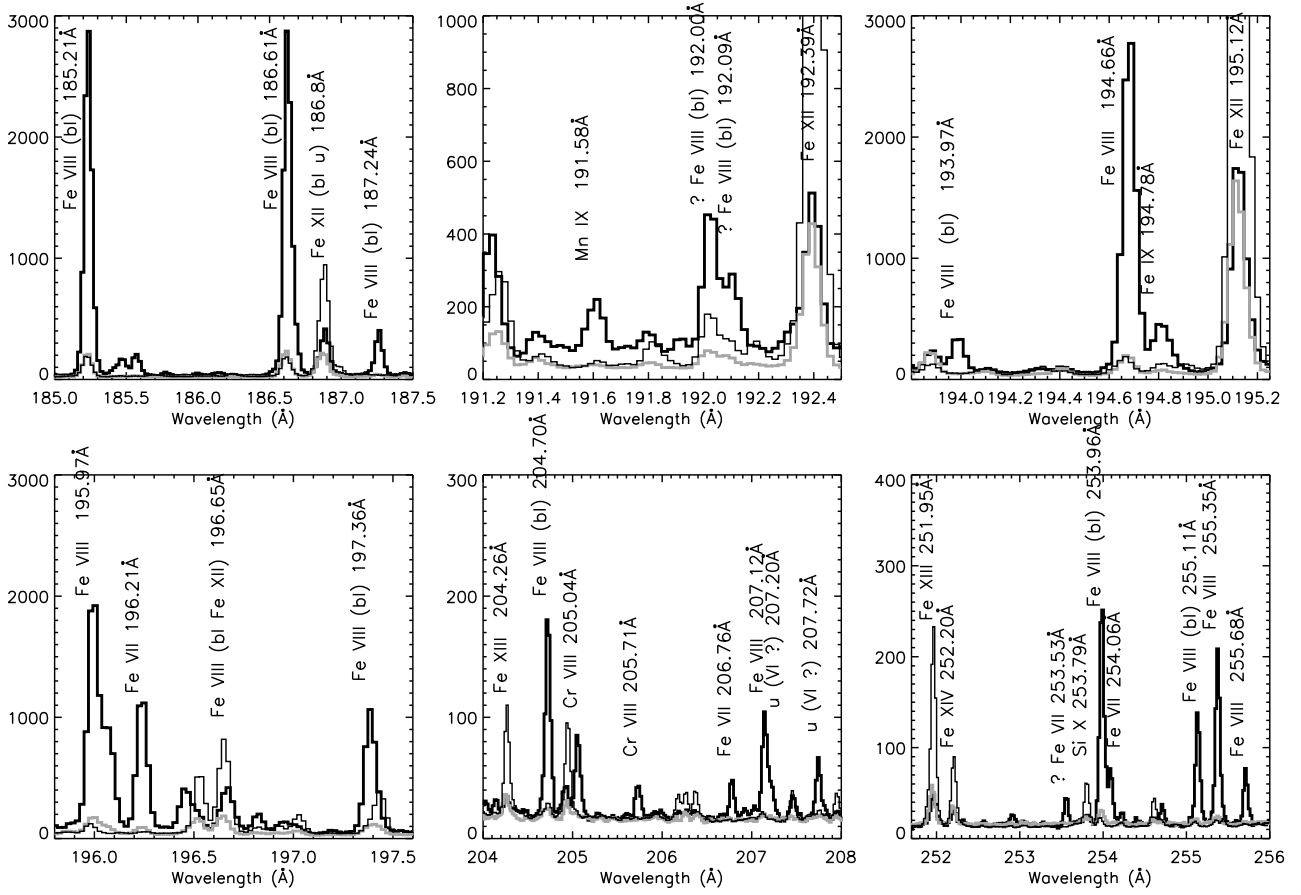


Fig. 4. Hinode EIS spectra (units are average counts per pixel) over three different areas. Thick lines refer to the spectrum over the Sunspot leg, where Fe VIII lines were clearly strong. The thick grey line shows the foreground Sunspot spectrum, while the thin black spectrum is from the reference area used for the wavelength calibration.

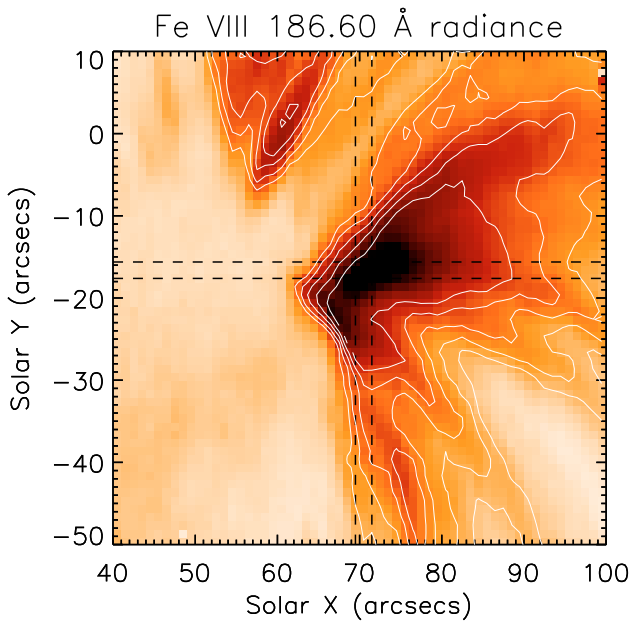


Fig. 5. Enlarged portion of the Fe VIII 186.60 Å radiance (negative image, SW channel), with contours of the Si VII 275.35 Å intensity (LW channel), showing agreement to within 1'' in the spatial alignment, after all geometrical corrections were applied to the two EIS channels. The location of the area chosen to obtain averaged spectra from a Sunspot loop leg is indicated by the crossing of the two sets of dashed lines.

intended by-products of the on-going benchmark work. Overall, as Fig. 7 shows, agreement between the EIS and the reference wavelengths is however accurate to within 6 mÅ across the entire spectral range.

The same wavelength calibration was applied to the “foreground-corrected” Sunspot loop leg spectrum, and all the measured wavelengths of TR lines formed at the Fe VIII temperatures were corrected for a red-shift of 20 km s⁻¹. The results are shown in Table 4, together with wavelength measurements from the literature and identifications (a full list is provided in Del Zanna 2009). An overall conservative uncertainty of 10 mÅ is suggested for the measured wavelengths. Note, however, that despite the lack of an appropriate reference spectrum, the measured wavelengths are well within this uncertainty for the previously-known Fe VIII, Si VI, Mg VI, Mg VI lines, while they differ significantly for Si VII. It is likely that the EIS wavelengths of Table 4 are more accurate than any previous values. Note that the wavelengths provided by Brown et al. (2008) were of limited accuracy because the effects of Doppler motion were not taken into account.

3.2. Laboratory data

One of the original plates from B.C.Fawcett was found to contain very strong transition region lines, mostly from Fe VIII, Fe IX. The plate (C12h) was obtained from an Iron Carbonyl source using a grazing incidence spectrometer. The plate was

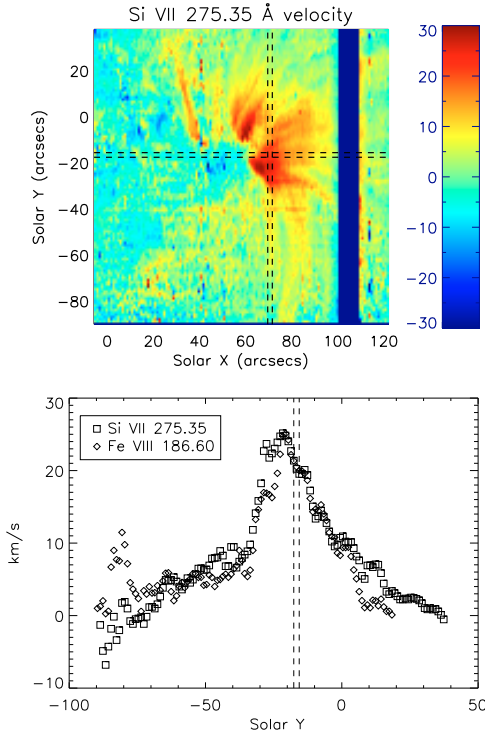


Fig. 6. *Top:* Doppler-gram of the Si VII 275.35 Å line showing strong (20 km s^{-1} or more) red-shifts in the legs of the coronal loops. The location of the area chosen to obtain averaged spectra from a Sunspot loop leg is indicated by the crossing of the two sets of dashed lines. *Bottom:* a slice along the N-S dashed lines of the Doppler-shifts in two lines in the SW and LW channels. Despite the lack of simultaneity between the channels, good agreement is found. The dashed lines indicate the location of the chosen Sunspot loop leg area (having a red-shift of about 20 km s^{-1}).

scanned, and a portion of an averaged spectrum wavelength-calibrated. Spectral windows containing Fe VIII are shown in Fig. 8. Unfortunately, it is not possible to calibrate the intensities of the spectral lines, so the assessment is mainly done considering the wavelengths. However, observed line strengths are in rough agreement with the gf values. All the Fe VIII lines observed by EIS having large gf values are also present in the plate.

3.3. The Malinovsky & Heroux (1973) spectrum

Malinovsky & Heroux (1973) published an integrated-Sun spectrum covering the 50–300 Å range with a medium resolution (0.25 Å), taken with a grazing-incidence spectrometer flown on a rocket in 1969. For a long time this has been the best available spectrum in the 150–300 Å range. One of the main reason is the fact that the spectrum was photometrically calibrated with great accuracy. Indeed previous benchmark studies have shown agreement to within a few % between predicted and observed intensities.

4. Discussion

The Malinovsky & Heroux (1973) rocket spectrum allows us to benchmark the atomic data for the strongest EUV transitions. The emissivity ratio curves are presented in Fig. 9, both using the original and scaled collision strengths and a scaling

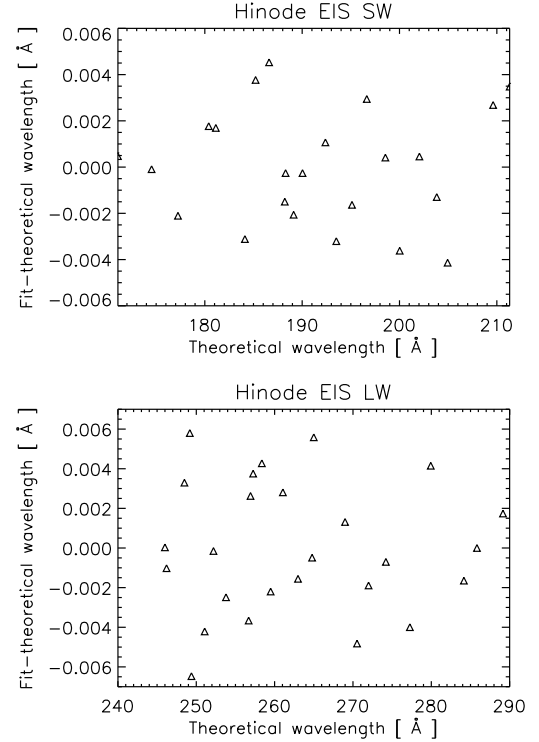


Fig. 7. Difference between fitted and theoretical wavelengths (Å) in the SW and LW channels of EIS. A linear dispersion for the SW channel was used, while a quadratic one was adopted for the LW.

constant $C = 4.5 \times 10^{19}$. The intensities were integrated over the full Sun when it was relatively quiet, so it is expected that the plasma has a broad multi-thermal distribution peaked around 1 MK; different curves are not necessarily expected to cross. However, Fig. 9 clearly shows that the relative temperature sensitivity in the lines is very small, so multi-thermal effects would also be small. Overall, the comparison between observed and expected intensities is satisfactory. There is slightly better agreement (within $\pm 20\%$) between observed and expected intensities using the scaled collision strengths, in particular when the 1–53 transition is considered. This line has a mild temperature sensitivity and provides $\log T[\text{K}] = 5.6$. However, this line is one of the three which had their intensity obtained with a deconvolution method by Malinovsky & Heroux (1973), so has a more uncertain strength.

Malinovsky & Heroux (1973) only observed the two brightest transitions at the EIS wavelengths, at 185.21, 186.60 Å, because of the lower spectral resolution and the averaging over the whole Sun of the rocket spectrum did not allow observations of the weaker lines. Notice that these strong 185.21 and 186.60 Å lines, in normal active region observations, would be partially blended with Ni XVI and Ca XIV.

The EIS spectra allow us to consider further weaker lines. Figs. 10, 11, 13 show the emissivity ratio curves for the “foreground-subtracted” Sunspot loop leg, all calculated at an electron density of 10^9 cm^{-3} and the same scaling constant $C = 5.5 \times 10^{11}$ (so they are directly comparable). The observed intensities I_{ob} are in $\text{phot cm}^{-2} \text{ s}^{-1} \text{ arcsec}^{-1}$. Figure 10 shows the curves for the six strongest transitions observed in the SW channel. The 2–43 194.66 Å and 1–42 195.97 Å are strong transitions which appear unblended. The 20% discrepancy (when using the original collision strengths) between the

Table 4. List of measured wavelengths.

$\lambda_{\text{meas}}(\text{\AA})$	$\lambda_{\text{lit}}(\text{\AA})$	ID
185.214	185.213	Fe VIII (bl oc Ni XVI 185.23)
186.607	186.601	Fe VIII (bl oc Ca XIV 186.61)
187.243	187.237	Fe VIII
191.585	191.570	Mn IX
192.000	192.004	? Fe VIII (bl oc Fe XI Fe XXIV)
192.087	–	u (VIII-IX)
193.972	193.967	Fe VIII
194.658	194.662	Fe VIII
195.969	? 195.972	Fe VIII (bl)
196.649	196.650	Fe VIII (bl oc Fe XII)
197.364	197.362	Fe VIII (bl)
204.704	–	Fe VIII (bl oc Fe XVII)
205.04	205.040	Cr VIII
205.708	205.720	Cr VIII
206.761	–	Fe VII
207.124	–	Fe VIII
207.203	–	u (VI?)
207.725	–	u (VI?)
217.71	217.691	Fe VIII
218.57	218.564	Fe VIII
222.19	–	Fe VIII
224.28	224.305	Fe VIII
225.40	–	Fe VIII
227.29	–	Fe VIII
231.03	231.097	Fe VIII
231.81	231.884	Fe VIII
246.011	246.004	Si VI (bl)
249.135	249.124	Si VI (bl oc Ni XVII)
253.520	–	? Fe VII
253.956	–	Fe VIII (bl oc S X)
254.059	–	Fe VII
255.110	–	Fe VIII (bl oc S X, Fe XXIV)
255.346	–	Fe VIII
255.684	–	Fe VIII
268.996	268.991	Mg VI
270.401	270.391	Mg VI
272.661	272.639	Si VII (bl)
275.370	275.353	Si VII
275.686	275.667	Si VII (bl)
276.145	276.153	Mg VII
280.732	280.737	Mg VII

Note: $\lambda_{\text{meas}}(\text{\AA})$ are the measured wavelengths for Fe VIII lines and a selection of lines from ions formed at similar temperatures. Results are from the Hinode EIS spectrum of the “foreground-subtracted” Sunspot loop leg and from the laboratory spectrum. Uncertainties are estimated to be about 0.01 \AA . In the second column wavelengths from the literature (λ_{lit}) are listed. A few of the main lines which are potentially blending (bl) are also noted. Some of the unidentified (u) TR lines are listed, together with the Fe ionization stage which more closely resembles their morphology. oc: blended in other plasma conditions.

2–46, 1–45 and the 2–43, 1–42 transitions confirms the benchmark calculation results: the first two transitions have too large gf values, hence too large collision strengths, which in turns increases the theoretical intensities, and lowers the emissivity ratio curves. A slightly better agreement is found when the scaled collision strengths are used. Notice that about 40% of the intensity of the 2–41 197.36 \AA line appears to be due to an Fe VII line

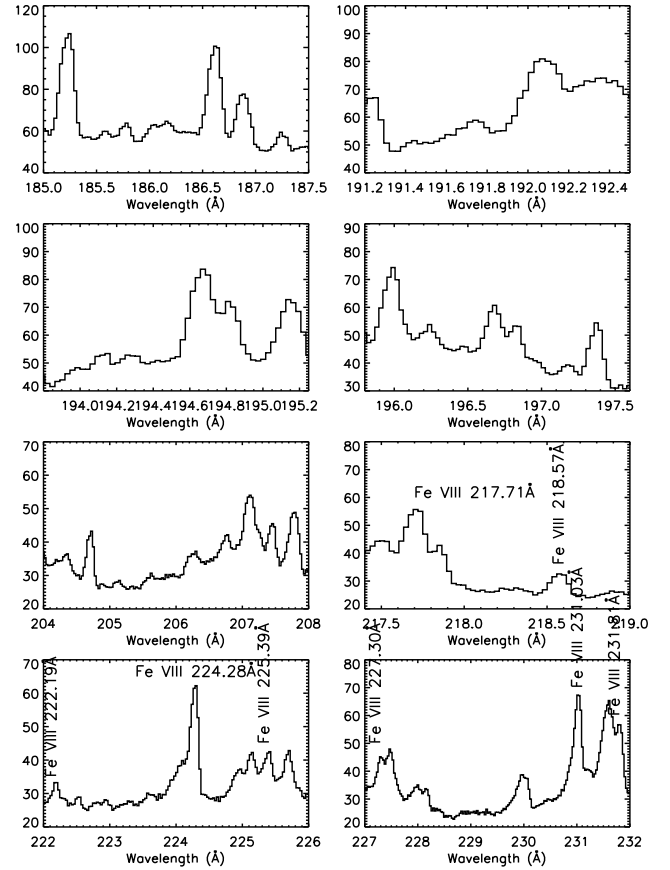


Fig. 8. Spectral windows from the laboratory plate (arbitrary units) containing Fe VIII lines. The first five wavelength ranges are the same as those reported in Fig. 4, while the last ones show regions not observed by Hinode EIS.

(Del Zanna 2009), so the observed intensity has been corrected by this amount. Very good agreement is found with the scaled collision strengths. The 1–36 $3p^63d^2D_{3/2}-3p^53d^2^2D_{5/2}$ and 2–37 $3p^63d^2D_{5/2}-3p^53d^2^4S_{3/2}$ transitions are expected to be close in wavelength and strong. They are identified here as a self-blend observed at 207.124 \AA , considering that no other nearby lines (see Fig. 4) have the appropriate intensity and morphology. About 30% of the observed intensity still appears unaccounted for.

Figure 11 shows the emissivity ratio curves for a few weaker transitions. With the exception of the 1–43 193.97 \AA transition, all lines appear to be blended. Notice that the 196.65 \AA line in normal conditions would be blended with the strong Fe XII 196.64 \AA transition, important for density diagnostics (Del Zanna & Mason 2005). A further 60% of the observed intensity must be due to an unidentified TR line. The weak 187.24 \AA also appears to be significantly blended with another TR line.

The 1–44 $3p^63d^2D_{3/2}-3p^53d^2^2P_{1/2}$ transition was only tentatively identified by RR80 with a line observed at 192.004. The EIS spectra do show a TR line at 192.000 \AA which is most probably this transition, blended by 60% with another, stronger TR line. Other possibilities are a TR line observed at 192.087 \AA , which however has a morphology closer to Fe IX, and the 191.585 \AA line, which is most probably a Mn IX transition (see Table 4, the laboratory spectrum in Fig. 8, the EIS one in Fig. 4, and the morphology in Fig. 2). Notice that the

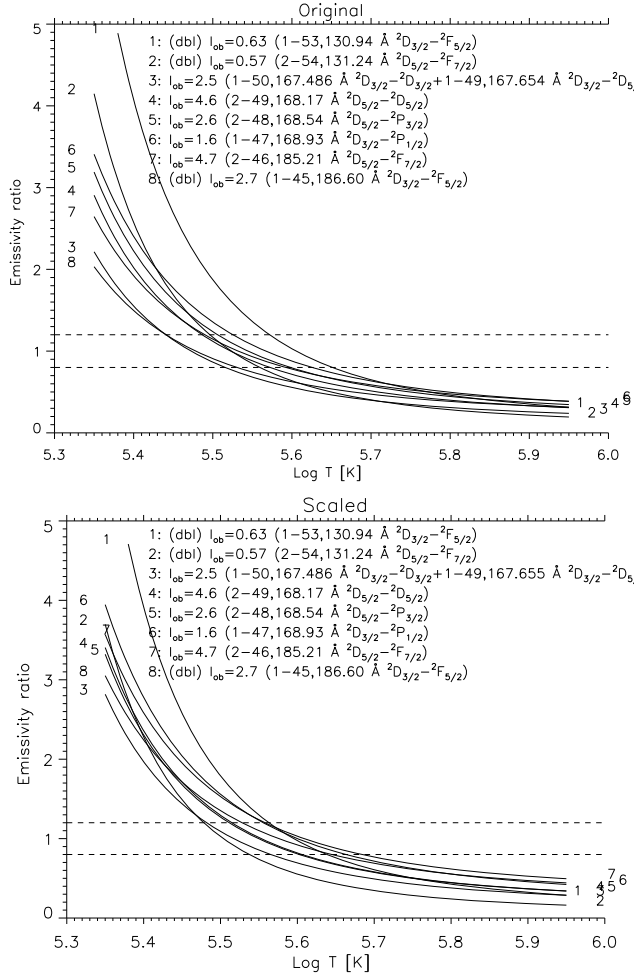


Fig. 9. The emissivity ratio curves relative to the [Malinovsky & Heroux \(1973\)](#) spectrum, using the original G00 collision strengths (above) and the scaled ones (below). I_{ob} indicates the observed flux ($10^{-3} \text{ erg cm}^{-2} \text{ s}^{-1}$). The lines labeled “(dbl)” were originally blended and their intensities obtained with a deconvolution method.

192.000 Å line is of particular importance because it blends, together with an Fe XI transition ([Del Zanna et al. 2009](#)), with the flare Fe XXIV line ([Del Zanna 2006](#)).

The 2–40 $3p^6 3d^2 D_{5/2} - 3p^5 3d^2^2 G_{7/2}$ line ought to be well observed in the EIS spectrum. The only TR lines close in wavelength and predicted intensity are those observed at 204.704 and 205.040 Å. The 205.040 and 205.708 Å lines have the correct wavelengths and intensity ratio to be identified with two strong Cr VIII transitions, firstly identified by G65 (the strongest Cr VII transition at 200.83 Å is also a prominent line in the spectrum, which confirms the presence of Cr lines). This leaves the 204.704 Å as the only possibility, although the predicted intensity is less than half of what is observed. Note that this line is important because it is blended with one of the two strongest Fe XVII lines observable by EIS ([Del Zanna & Ishikawa 2009](#)).

All the lines observed by EIS are also present in the laboratory spectrum of Fig. 8. In the 211–246 Å range, the spectrum contained 5 of the strong Fe VIII lines originally identified by RR80, plus 3 transitions identified here (see Table 4 for the measurements and Table 3 for the spectroscopic identification). These 8 transitions are labelled in Fig. 8.

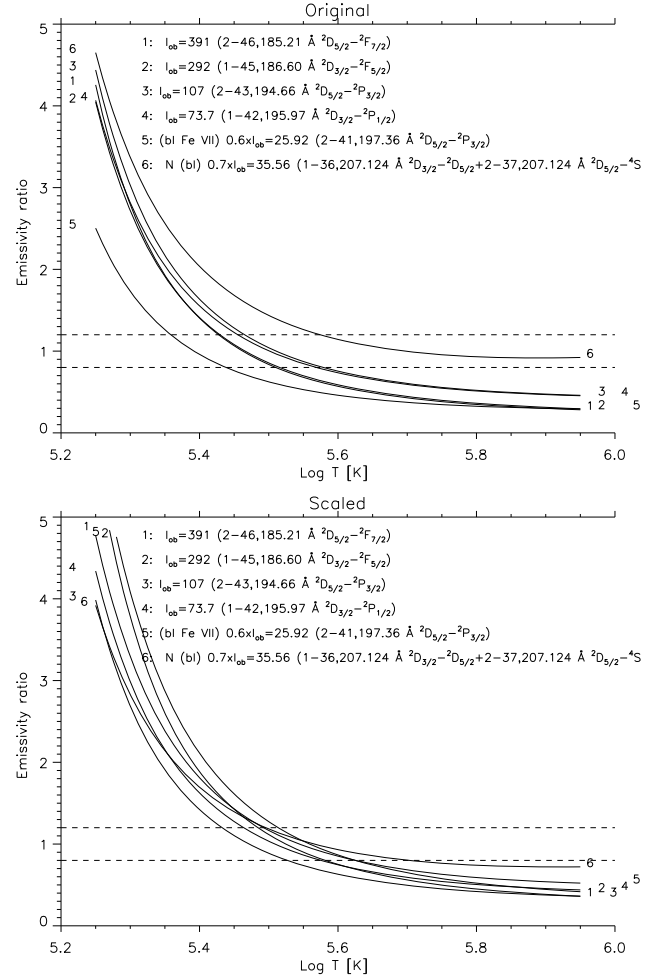


Fig. 10. The emissivity ratio curves relative to the “foreground-subtracted” Sunspot loop leg observed by EIS using the original G00 collision strengths (above) and the scaled ones (below), for the strongest transitions in the SW channel.

4.1. Identification of the 4D_j levels and the possibility of measuring T_e

The strong lines of the $3p^6 3d^2 D_j - 3p^5 3d^2 ^4D_j$ transition array were not previously identified. They are predicted to be strong and fall within the EIS wavelength range. There are only few candidate lines, based on morphology arguments, but when the theoretical splittings of the 4D_j levels is considered, there are no options. The four strong lines around 255 Å seen in Fig. 4 are clearly lines from this array. The agreement between the measured and theoretical splittings of the 4D_j levels is excellent, as shown in Table 2.

These lines are particularly important because they provide a direct way to measure the electron temperature, when combined with lines decaying from higher levels. Before discussing the emissivity ratios for these lines, however, the temperature distribution of the plasma needs to be considered.

One well-established way to estimate the temperature distribution of the emitting plasma is the so-called emission measure loci method, whereby the observed intensities, divided by their contribution functions, are plotted as a function of the electron temperature. For a description of the method, see e.g. [Del Zanna et al. \(2002\)](#). When introduced for the first time to the study of active region loops observed with the SOHO/CDS spectrometer, this method has consistently indicated that the legs of active

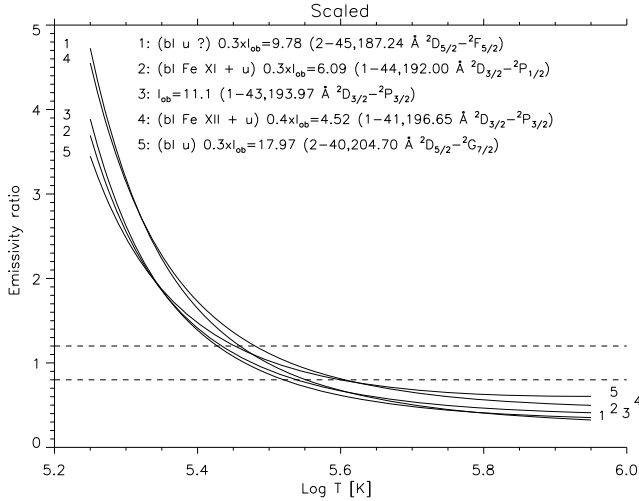


Fig. 11. The emissivity ratio curves relative to the “foreground-subtracted” Sunspot loop leg observed by EIS for a few weaker transitions, using the scaled collision strengths.

region loops seem to have a near-isothermal temperature distribution (Del Zanna 2003; Del Zanna & Mason 2003). This means that the various ions are emitting at temperatures which can be well below or well above the temperature of peak ion abundance in equilibrium.

A similar situation applies to the Sunspot loop leg examined here. Figure 12 shows the EM Loci plots for three Mg lines (Mg V 276.58 Å, Mg VI 270.39 Å, Mg VII 278.4 Å deblended from Si VII) and three Fe lines: Fe VII 176.74 Å, Fe VIII 186.61 Å, and Fe IX 171.07 Å. The atomic data used for Fe VII are described in Del Zanna (2009), while for Fe VIII the scaled collision strengths are used. For Fe IX and the Mg ions, CHIANTI v.5.2 (Landi et al. 2006) was used.

The new CHIANTI v.6 (Dere et al. 2009) ion abundances (in ionization equilibrium) were used for the EM Loci plot. The EM Loci of the Mg lines in Fig. 12 are consistent with an isothermal plasma having $\log T$ [K] = 5.65. Curves from other ions, not shown here, also cluster around $\log T$ [K] = 5.6, as are those from the Fe ions. If ionization equilibrium holds, and the ion abundances are correct, one would then expect to measure similar values from intensity ratios of lines sensitive to temperature. Moreover, the emissivity ratio curves should all cross at $\log T$ [K] = 5.65.

A few caveats however apply. First, some departures from ionization equilibrium caused by the strong down-flows (of the order of 30–50 km s⁻¹ along the structure) are to be expected. For the upper-transition-region ions such as Fe VIII, Fe IX, the down-flows would naturally lower the electron temperature at which lines are formed (Raymond & Dupree 1978). The fact that the plasma is mainly radiatively cooling (Bradshaw 2008), would also lower the measured electron temperature compared to equilibrium values. This issue will be addressed in a follow-up paper. Second, large uncertainties are associated with ionization equilibrium curves, for a variety of reasons (see e.g. some issues discussed in Del Zanna et al. 2002). Also notice that the new CHIANTI v.6 equilibrium curves are significantly different from previous ones, in particular for the Fe ions, where the peak ion abundance is now shifted toward much lower temperatures. The excellent agreement in the morphology of Si VII, Mg VII, and Fe VIII lines (see, e.g. Fig. 2, 3) suggests a possible discrepancy in the temperature calculated in collisional

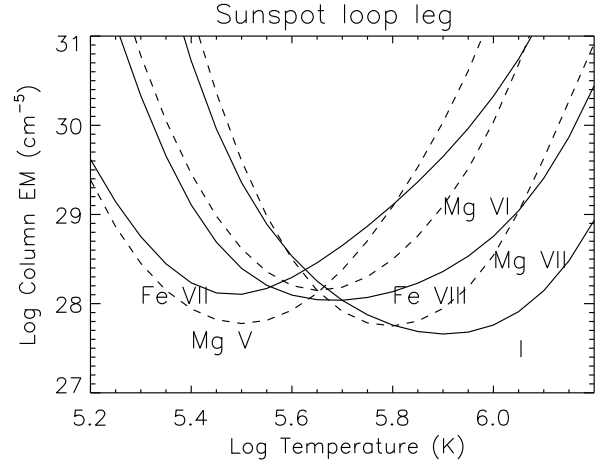


Fig. 12. Emission measure loci plot relative to the “foreground-subtracted” Sunspot loop leg observed by EIS.

equilibrium for Fe VIII ($\log T$ [K] = 5.6), compared to Si VII and Mg VII ($\log T$ [K] = 5.8). Third, even if a loop cross-section is nearly isothermal, line-of-sight effects can always affect the measurements.

The emissivity ratio curves for the lines of the $3p^6 3d^2 D_j - 3p^5 3d^2 4D_j$ transition array are shown in Fig. 13, both using the original G00 collision strengths and the scaled ones. The curves relative to the two strong decays 1–45 186.60 Å and 1–42 195.97 Å are also displayed. There are large disagreements between the observed and expected intensities when the original data are used. Also, the strongest transitions would indicate a $\log T$ [K] = 5.4, far below the temperature from the EM Loci method. The scaled data provide a much better agreement between observed and expected intensities, and higher temperatures (about $\log T$ [K] = 5.5), although the scatter in the curves is still quite large.

5. Summary and conclusions

From the various atomic structure calculations published and presented here, it is clear that the excitation rates of G00 need to be improved. A large scattering calculation for this ion is non-trivial, so for this benchmark work the G00 rates of the dipole-allowed transitions have been scaled using a large benchmark calculation. Large uncertainties in the predicted line intensities are obviously still present, however the comparison with the Hinode/EIS and the Malinovsky & Heroux (1973) intensities is very satisfactory, with agreement within $\pm 20\%$ for all the strongest transitions, when the scaled collision strengths are considered. The excellent agreement between the Hinode/EIS wavelengths and the RR80 ones is very satisfactory, and shows that with a detailed analysis, very accurate wavelengths can be obtained with the EIS spectra.

The predicted line intensities, combined with a detailed analysis of Hinode/EIS and laboratory spectra, allowed the identification of several new transitions, and confirm the previous ones from G65 and RR80. However, a few lines which should be observable in astrophysical spectra are still not identified, and in a large number of cases Fe VIII lines in the EIS spectra appear to be blended.

Of particular importance are the newly identified strong lines of the $3p^6 3d^2 D_j - 3p^5 3d^2 4D_j$ transition array because it is shown here that they can provide a direct measurement of the

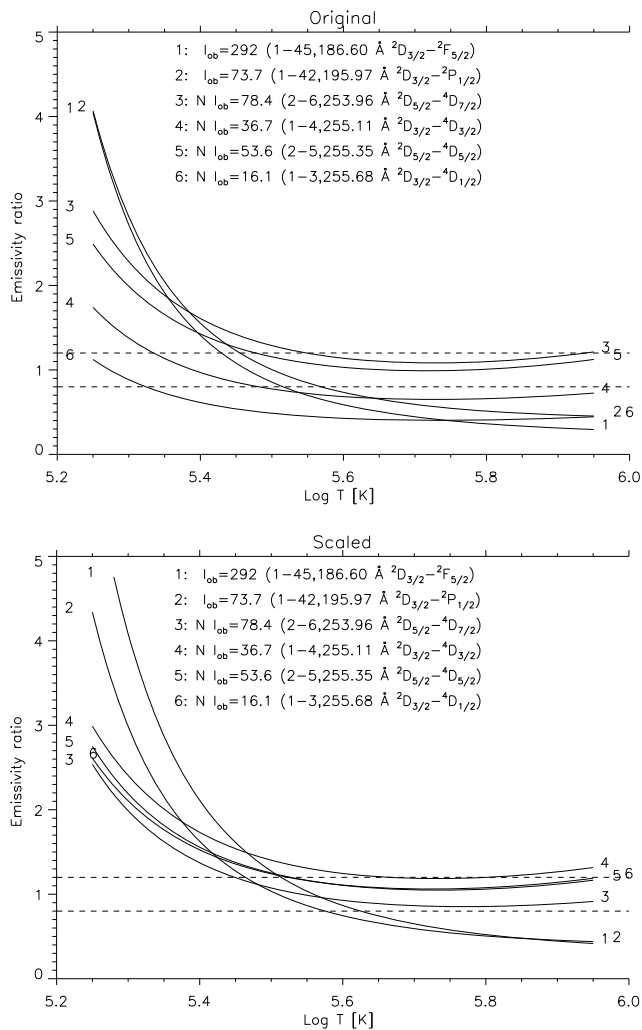


Fig. 13. The emissivity ratio curves relative to the “foreground-subtracted” Sunspot loop leg observed by EIS using the original G00 collision strengths and the scaled ones.

electron temperature. Previous *Skylab* observations produced contradictory results in terms of loop electron temperatures from line ratios in the upper transition region. For example, [Raymond & Foukal \(1982\)](#) found a very low electron temperature from the resonance vs. intercombination Ne VII ratio ($\log T$ [K] = 5.3 instead of the equilibrium 5.75) in loops observed at the limb. [Doyle et al. \(1985\)](#) used the same Ne VII ratio on a Sunspot loop but found an equilibrium temperature (consistent departures were found in lower-temperature ions).

The EIS observations indicate that loop legs have lower temperatures than predicted in ionization equilibrium for Fe VIII. This is not surprising; however, firm measurements will only be possible once better scattering calculations are available, and the EIS radiometric calibration is properly assessed. Evidence

was found in a number of cases ([Del Zanna 2008a](#); [Del Zanna & Ishikawa 2009](#)) for lines around 250 Å having significantly lower intensities than expected.

It is obviously important to attempt measurements of electron temperatures which are independent of any assumption of ionization equilibrium. Further observational constraints and non-equilibrium forward-modeling such as that of [Bradshaw et al. \(2004\)](#) will shed some light on this issue.

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