

A new candidate supernova remnant G 70.5+1.9

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ABSTRACT

A compact complex of line emission filaments in the galactic plane has the appearance of those expected of an evolved supernova remnant, although non-thermal radio and X-ray emission have not yet been detected. This optical emission line region has now been observed with deep imagery and both low and high-dispersion spectroscopy. Diagnostic diagrams of the line intensities from the present spectra and the new kinematical observations both point to a supernova origin. However, several features of the nebular complex still require an explanation within this interpretation.

Key words. ISM: general – ISM: supernova remnants – ISM: individual objects: G 70.5+1.9

1. Introduction

During observations of a new candidate supernova remnant (SNR) in the Cygnus region, close to the known remnant CTB 80, a bright emission line source was detected (Mavromatakis & Strom 2002). This source, called G 70.5+1.9, is relatively compact (a few arcminutes wide), its boundaries are rather sharp and it exhibits a tail of diffuse emission to the east, only seen in the [O III] image. The southern part of this structure is visible on the POSS plates. However, no other information is available on this peculiar optical line source. Published radio surveys do not show any strong evidence for non-thermal emission at the position of the optical source, while the ROSAT All-Sky survey data do not reveal any excess X-ray emission at this position. Both could be expected if the filamentary nebula is an evolved SNR. In order to study in more detail the detected structures and to understand their origin, we performed deep optical, flux calibrated imaging observations in the low and medium ionization lines of $H\alpha + [N II]$, [S II], [O II] and [O III]. We have also obtained deep spatially-resolved low-dispersion long-slit spectra at four different slit positions of this filamentary structure in an effort to obtain more information about the actual physical conditions. Spatially resolved long-slit spectra along one slit position have also been obtained to reveal any motions typical of an evolved SNR. Information about the observations and results are given in Sect. 2, while in Sect. 3 we discuss the nature of these new filamentary structures.

2. Observations and results

2.1. Imaging

The observations presented in this paper were performed with the 1.3 m Ritchey-Cretien telescope at Skinakas Observatory,

Crete, Greece. The filamentary nebula G 70.5+1.9 was observed on July 09 and 10, 2001 with interference filters isolating the optical emission lines of $H\alpha + [N II]$, [S II], [O II], and [O III]. The 1024×1024 SITE CCD used during the observations had a pixel size of $24 \mu\text{m}$ resulting in a 8.5×8.5 field of view. Two different pointing directions were used in order to cover the area of emission seen in the wide field images of Mavromatakis & Strom (2002). The first pointing concentrated on the bright filaments to the west, while the second focused on the diffuse [O III] emission to the east. All fields were projected to a common origin on the sky and were subsequently combined to create the final mosaics. The astrometric solutions were calculated with IRAF routines and utilized the HST Guide star catalogue (Lasker et al. 1999). The log of the observations together with the filter characteristics and the exposure times are given in Table 1. All coordinates given in this work refer to epoch J2000.

We employed standard IRAF and MIDAS routines for the reduction of the data. Individual frames were bias subtracted and flat-field corrected using well-exposed twilight flat-fields. The spectrophotometric standard stars HR 5501, HR 7950, HR 7596, HR 9087, and HR 8634 were observed for absolute flux information (Hamuy et al. 1992, 1994).

2.1.1. The $H\alpha + [N II]$ and [S II] line images

The higher angular resolution images of G70.5+1.9 allow us to study in more detail the network of filaments. A basic characteristic of these filaments is their brightness in $H\alpha + [N II]$ (Fig. 1) and [S II] (Fig. 2). The bulk of the emission seems to be bounded by two very sharp filaments. These are separated by a typical distance of $\sim 2'$ along the declination axis. The southern filament is found at an almost constant declination ($\delta \sim 33^\circ 53' 30''$), while the northern is inclined by $\sim 15^\circ$ with

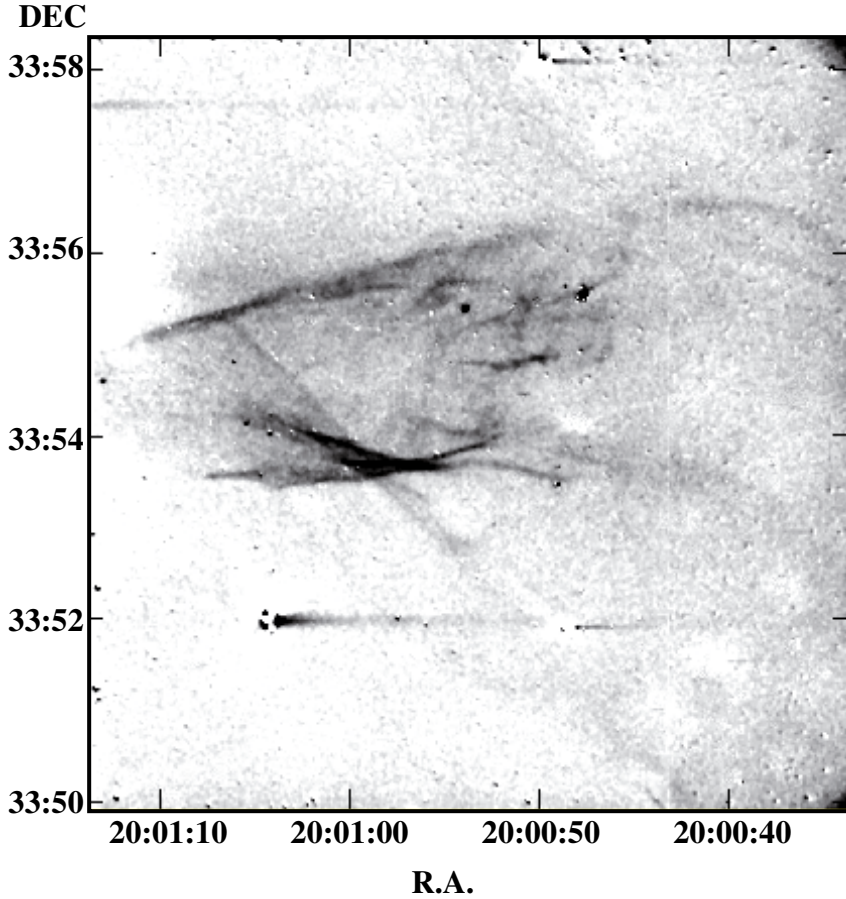


Fig. 1. The 8'5 square field covering the filamentary nebula G70.5+1.9 in the light of H α + [N II] after subtraction of the continuum image. The shadings run linearly from 0 to $20 \times 10^{-16} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ arcsec}^{-2}$. The line segments seen near over-exposed stars in this figure and the following figures are due to the blooming effect.

Table 1. Imaging and spectral log.

IMAGING				
Filter	λ_c (Å)	$\Delta\lambda$ (Å)	Total exp. time (s)	Fields ^a
H α + [N II]	6566	75	6600 (3) ^b	W
[S II]	6712	20	6600 (3)	W
[O III]	5014	28	4800 (2)	W
[O II]	3727	25	4800 (2)	W
Cont blue	5470	230	360 (3)	W
Cont red	6096	134	510 (5)	W
H α + [N II]	6566	75	2400 (1)	E
[S II]	6712	20	2400 (1)	E
[O III]	5014	28	4800 (2)	E
[O II]	3727	25	2400 (1)	E
Cont blue	5470	230	360 (3)	E
Cont red	6096	134	360 (3)	E
SPECTROSCOPY				
Area	Slit centers		Total exp. time (s)	Telescope
	α	δ		
Area 1	20 ^h 01 ^m 06.2 ^s	33°56'20"	2700 (1)	1.3-m
Area 2	20 ^h 01 ^m 05.0 ^s	33°56'15"	7200 (2)	1.3-m
Area 3	20 ^h 00 ^m 58.2 ^s	33°54'34"	3600 (2)	1.3-m
Area 4	20 ^h 00 ^m 52.0 ^s	33°55'41"	3600 (2)	1.3-m
Area 5	20 ^h 01 ^m 05.4 ^s	33°54'20"	3600 (2)	2.1-m

^a Field observed: W(=West), E(=East).

^b Numbers in parentheses represent the number of individual frames.

respect to the east-west direction. Diffuse emission and several shorter filamentary structures are detected between the north and south boundaries (Figs. 1 and 2). The detected emission line structures are not bounded in the east-west direction and weak

diffuse as well as filamentary emission is detected further to west in an arc-like shape (Fig. 3, $\alpha \approx 20^{\text{h}}00^{\text{m}}40^{\text{s}}$, $\delta \approx 33^{\circ}53'$). There is also a long filamentary structure originating roughly from $\alpha \approx 20^{\text{h}}01^{\text{m}}07^{\text{s}}$, $\delta \approx 33^{\circ}55'20''$ and extending down to $\alpha \approx 20^{\text{h}}00^{\text{m}}42^{\text{s}}$ and $\delta \approx 33^{\circ}50'02''$, after a long gap of $\sim 3'$. Finally, a patchy structure is present around $\alpha \approx 20^{\text{h}}01^{\text{m}}05^{\text{s}}$ and $\delta \approx 33^{\circ}52'15''$ to the south of the main emission area. The east field was also observed in these filters as reported in Table 1 but it is not shown here since any emission is below our detection limit.

2.1.2. The [O II] and [O III] line mosaics

The detected structures are also quite strong in the low ionization line of [O II]3727 Å (Fig. 3). The morphology in this emission line is generally similar to that of the low ionization lines discussed above. The south, bright filament ($\alpha \approx 20^{\text{h}}01^{\text{m}}$, $\delta \approx 33^{\circ}53'30''$) is characterized by a surface brightness up to $\sim 28 \times 10^{-16} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ arcsec}^{-2}$, while the northern filament displays a surface brightness in the range of $12\text{--}19 \times 10^{-16} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ arcsec}^{-2}$. The weak but arc-like shaped emission to the west is detected at a level of $1\text{--}2 \times 10^{-16} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ arcsec}^{-2}$. The typical full width at half maximum (fwhm) of the filament to the south is found in the range of $\sim 3''\text{--}6''$, while the northern filament ($\alpha \approx 20^{\text{h}}01^{\text{m}}10^{\text{s}}$, $\delta \approx 33^{\circ}55'20''$) is broader with a fwhm of $\sim 8''$. The weaker filaments seen to the interior of the bright emission are characterized by a fwhm of $\sim 3'5$.

The morphology of the structure as recorded through the [O III] interference filter (Fig. 4) differs from the morphology

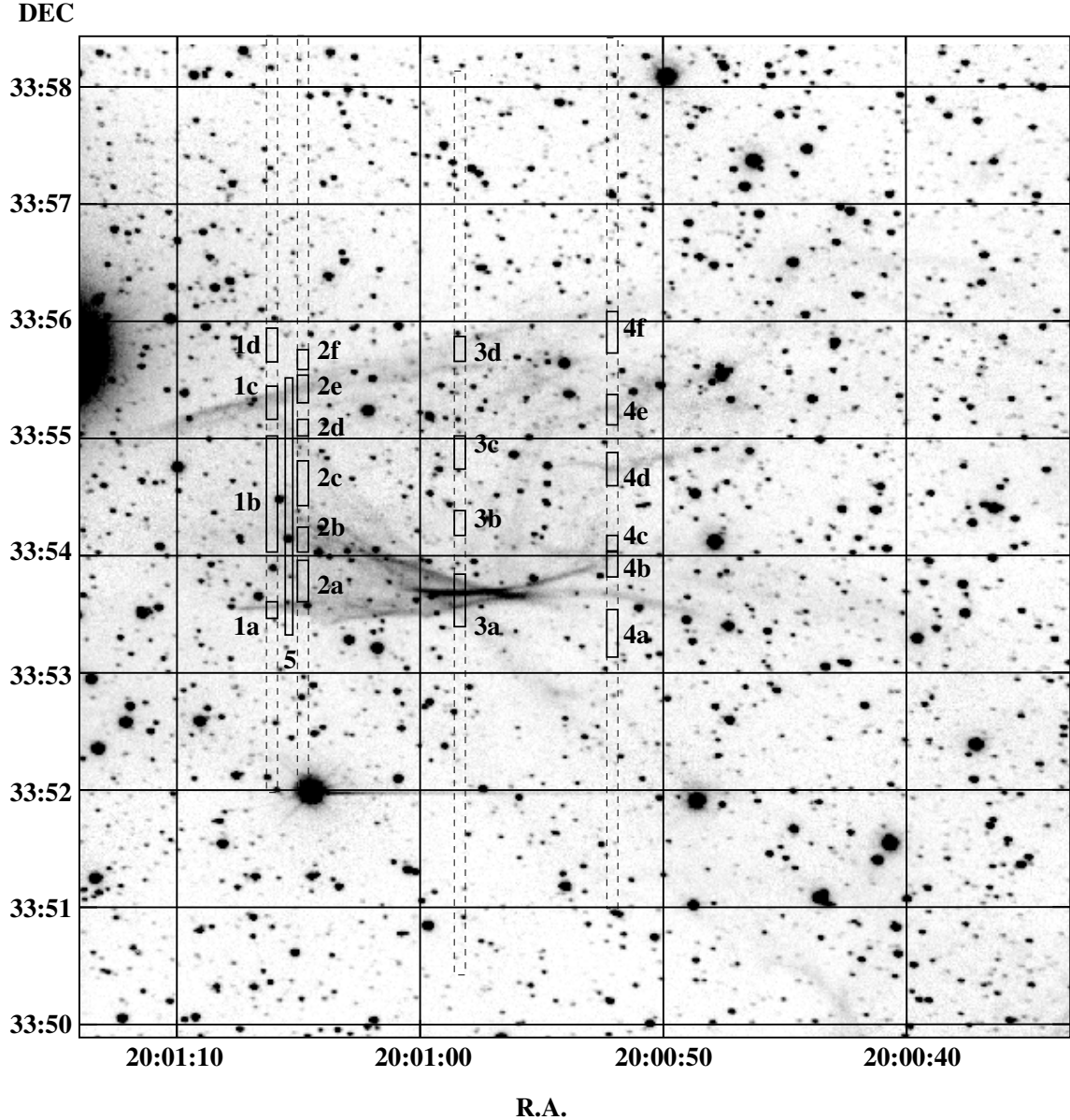


Fig. 2. The image of G70.5+1.9 in the low ionization line of [S II] without continuum subtraction so that a reference star field is clear. The long dashed rectangles show the projection of the low-dispersion slits 1–4 on the sky, while the small, individual rectangles along these represent the areas where line strengths were extracted. Slit position 5 is that part of the high-dispersion slit for which the position-velocity arrays of line profiles are shown in Fig. 5. The shadings run linearly from 0 to $8 \times 10^{-16} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ arcsec}^{-2}$.

of the low ionization images, including the [O II] image. The variety of structures detected in the $H\alpha + [\text{N II}]$, [S II] and [O II] images between the south and north filaments are not clearly present in the medium ionization line of [O III]. In fact, some structures may not be present at all or their morphology has now become so diffuse that does not allow the direct identification of the corresponding structures seen in the low ionization images. The north boundary filament appears quite sharp in this emission line, while it appears less well defined in $H\alpha + [\text{N II}]$. The inverse holds in the south-west, i.e. the filament is sharp in the $H\alpha + [\text{N II}]$ image but diffuse emission is seen in the [O III] image. A major difference between the low and medium ionization line images is the greater extent to the east of the north and south filaments (Fig. 4). The north [O III] filament extends for $\sim 43''$ further to the east from the tip of the corresponding $H\alpha + [\text{N II}]$ filament, while the south [O III] filament extends for $\sim 34''$ further to the east compared to the $H\alpha + [\text{N II}]$ emission (around $20^{\text{h}}01^{\text{m}}15^{\text{s}}$).

The images of these regions provide direct evidence for the presence of incomplete shock structures, i.e. areas where part of the recombination zone is missing, allowing high [O III]/ $H\beta$ ratios. Finally, the [O III] image displays a “tail” of mainly diffuse emission further to the east for $\sim 8'$. The typical width of this diffuse emission is less than $\sim 2'$, while its typical brightness is $\sim 1.5\text{--}2.0 \times 10^{-16} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ arcsec}^{-2}$. This diffuse component is not detected in the low ionization images and it is not clear if it is related to the filamentary emission in the west.

2.2. Spectroscopy

2.2.1. Low dispersion

Long-slit low-dispersion spectra were obtained on June 22 and 23, 2001 using the 1.3 m Ritchey-Cretien telescope at Skinakas

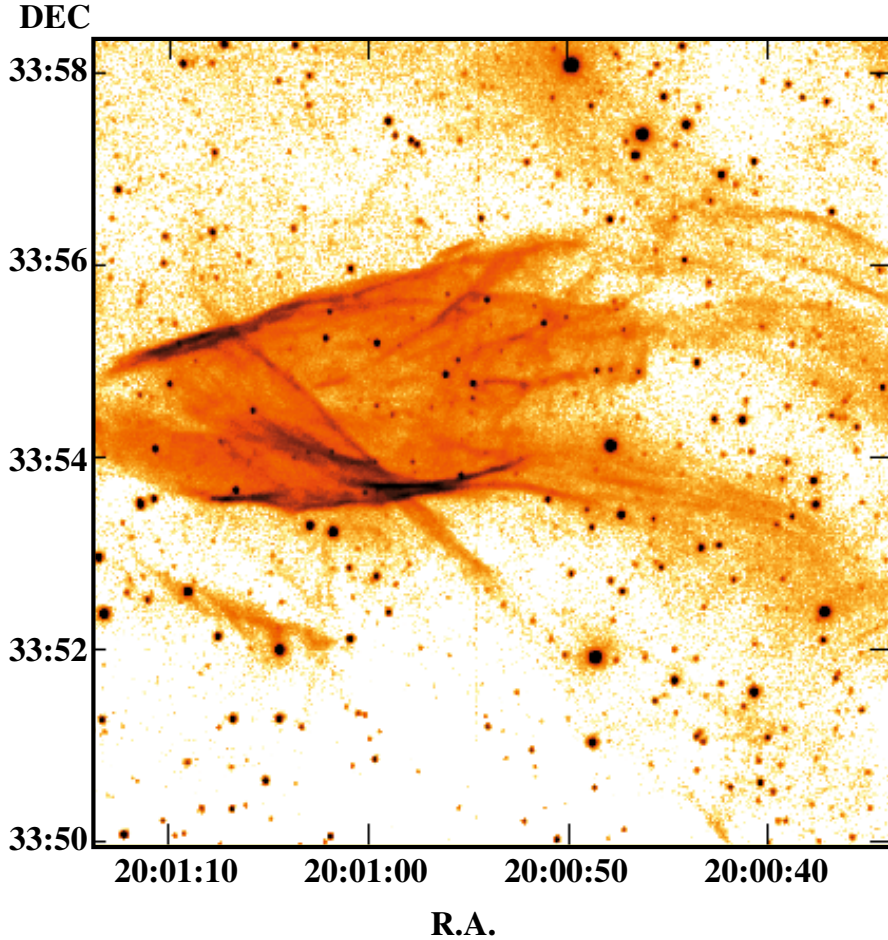


Fig. 3. The 8'5 field around G70.5+1.9 in the low ionization line of [O II]. The intensity representation is logarithmic and ranges from 0 to $30 \times 10^{-16} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ arcsec}^{-2}$. The diffuse emission seen in the west and south is at least ~ 10 times weaker than the flux of the bright filaments.

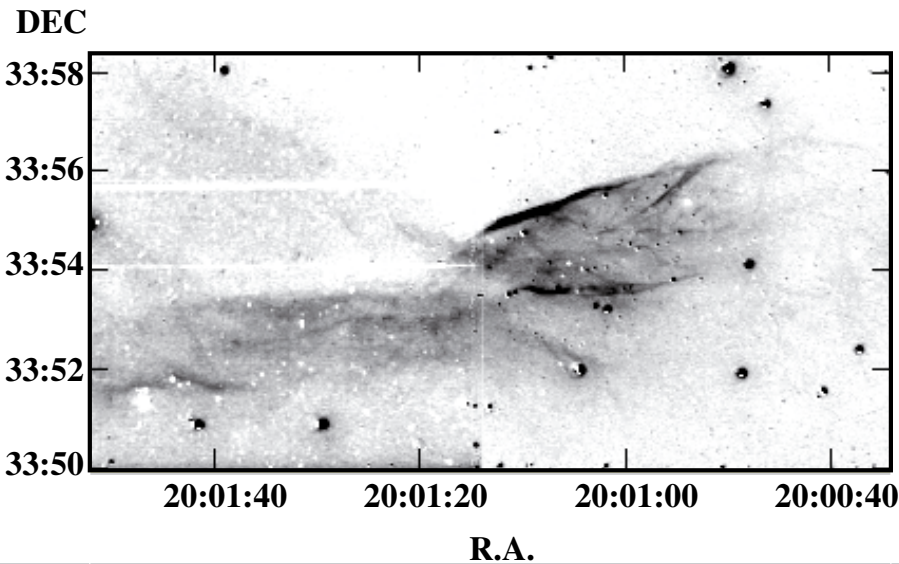


Fig. 4. The [O III] image of the whole network of filaments after subtraction of the continuum image. The north and south filamentary boundaries are also prominent in the medium ionization line of [O III]5007 Å. However, this line reveals diffuse emission further to the east which is not detected in the lower ionization images. The shading run linearly from 0 to $5 \times 10^{-16} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ arcsec}^{-2}$.

Observatory. The data were taken with a $1300 \text{ line mm}^{-1}$ grating and a 800×2000 SITe CCD covering the range of 4750–6815 Å. The slit had a width of $7''.7$ and a length of $7''.9$ and, in all cases, was oriented in the south-north direction. The coordinates of the slit centers along with the number of spectra and the total exposure times are given in Table 1. A spectrum taken on August 3, 2000 with the same hardware configuration and presented by Mavromatakis & Strom (2002) is also used here

with different apertures extracted. The spectrophotometric standard stars HR 5501, HR 7596, HR 7950, HR 9087, and HR 718 were observed for absolute flux information (Hamuy et al. 1992, 1994).

Long-slit spectra were obtained at four different positions. The relatively small size of the source ($3'$) compared with the slit length of $7''.9$ allowed very good background subtraction. In addition, we have the possibility to study the nature of the diffuse

Table 2. Relative line fluxes.

Line (\AA)	Area 1a			Area 1b			Area 1c			Area 1d			Area 2a		
	F^a	I^b	S/N^c	F	I	S/N	F	I	S/N	F	I	S/N	F	I	S/N
H β 4861	15	36	(6)	26	33	(43)	24	34	(26)	20	34	(11)	24	34	(28)
[O III] 4959	99	226	(39)	35	44	(61)	32	45	(40)	–	–	–	31	43	(39)
[O III] 5007	317	702	(114)	107	134	(177)	103	142	(115)	–	–	–	92	126	(111)
[O I] 6300	11	11	(7)	5	5	(17)	–	–	–	–	–	–	7	7	(10)
[O I] 6364	–	–	–	–	–	–	–	–	–	–	–	–	2	2	(3)
[N II] 6548	40	40	(27)	25	25	(68)	25	25	(50)	15	15	(21)	27	27	(54)
H α 6563	100	100	(47)	100	100	(158)	100	100	(133)	100	100	(105)	100	100	(147)
[N II] 6583	141	140	(72)	79	79	(148)	84	84	(35)	47	47	(53)	89	89	(144)
[S II] 6716	91	86	(55)	50	49	(119)	51	50	(87)	29	28	(41)	68	66	(125)
[S II] 6731	68	64	(40)	36	35	(76)	37	36	(67)	21	20	(23)	49	48	(94)
Absolute H α flux ^d	2			2.1			3.9			1.3			2.5		
[S II]/H α	1.50 ± 0.04			0.84 ± 0.01			0.86 ± 0.01			0.48 ± 0.01			1.14 ± 0.01		
$F(6716)/F(6731)$	1.34 ± 0.04			1.39 ± 0.03			1.38 ± 0.03			1.38 ± 0.07			1.39 ± 0.02		
[N II]/H α	1.80 ± 0.05			1.04 ± 0.02			1.09 ± 0.01			0.62 ± 0.01			1.16 ± 0.01		
[O III]/H β	25.8 ± 4.3			5.40 ± 0.13			5.50 ± 0.21			–			5.00 ± 0.18		
$c(\text{H}\beta)$	1.10 ± 0.20			0.31 ± 0.03			0.44 ± 0.05			0.67 ± 0.12			0.44 ± 0.05		
E_{B-V}	0.73 ± 0.13			0.21 ± 0.02			0.29 ± 0.03			0.44 ± 0.08			0.29 ± 0.03		
Line (\AA)	Area 2b			Area 2c			Area 2d			Area 2e			Area 2f		
	F^a	I^b	S/N^c	F	I	S/N	F	I	S/N	F	I	S/N	F	I	S/N
H β 4861	24	34	(35)	23	34	(25)	25	34	(21)	24	34	(38)	29	33	(17)
[O III] 4959	18	25	(30)	25	36	(30)	32	43	(29)	39	54	(69)	5	6	(2)
[O III] 5007	56	77	(88)	72	101	(83)	97	128	(87)	119	164	(190)	12	14	(9)
[O I] 6300	9	9	(17)	7	7	(9)	10	10	(10)	6	6	(12)	10	10	(5)
[O I] 6364	2	2	(5)	3	3	(5)	3	3	(5)	–	–	–	–	–	–
[N II] 6548	24	24	(60)	21	21	(41)	23	23	(36)	26	26	(72)	17	17	(19)
H α 6563	100	100	(174)	100	100	(140)	100	100	(109)	100	100	(197)	100	100	(77)
[N II] 6583	78	78	(159)	69	68	(111)	71	71	(89)	84	84	(189)	48	48	(44)
[S II] 6716	63	62	(159)	51	50	(92)	44	43	(64)	50	49	(130)	30	30	(44)
[S II] 6731	45	44	(102)	36	35	(69)	32	31	(49)	36	35	(99)	22	22	(25)
Absolute H α flux ^d	4.2			2.1			2.2			3.7			1.5		
[S II]/H α	1.06 ± 0.01			0.85 ± 0.01			0.74 ± 0.01			0.84 ± 0.01			0.52 ± 0.01		
$F(6716)/F(6731)$	1.40 ± 0.02			1.42 ± 0.03			1.38 ± 0.04			1.39 ± 0.02			1.36 ± 0.07		
[N II]/H α	1.02 ± 0.01			0.89 ± 0.01			0.94 ± 0.01			1.10 ± 0.01			0.65 ± 0.01		
[O III]/H β	3.00 ± 0.09			4.00 ± 0.18			5.00 ± 0.25			6.40 ± 0.17			0.60 ± 0.07		
$c(\text{H}\beta)$	0.44 ± 0.04			0.47 ± 0.05			0.38 ± 0.07			0.44 ± 0.04			0.16 ± 0.08		
E_{B-V}	0.29 ± 0.03			0.31 ± 0.03			0.25 ± 0.05			0.29 ± 0.03			0.11 ± 0.06		

The emission line ratios [S II]/H α , $F(6716)/F(6731)$, [N II]/H α and [O III]/H β are calculated using the values corrected for interstellar extinction. The errors of the emission line ratios, $c(\text{H}\beta)$ and E_{B-V} , are calculated through standard error propagation.

emission detected between the bright filaments. The absolute H α flux ranges from 0.5 to 7×10^{-16} erg s $^{-1}$ cm $^{-2}$ arcsec $^{-2}$ (Table 2). The high [S II]/H α ratios clearly demonstrate that the detected emission originates from shock heated gas ([S II]/H α \sim 0.5–1.6; Table 2). The fluxes of the individual sulfur lines may provide data on the electron density of the emitting gas (Osterbrock & Ferland 2006). However, all sulfur line ratios approach the high end of the allowable range of values suggesting low electron densities. Using the nebular package within the IRAF software (Shaw & Dufour 1995) we find that all electron densities are estimated to lie below ~ 120 cm $^{-3}$. Therefore we cannot directly estimate the preshock cloud densities but can only place upper limits, provided that there is no magnetic field to halt the compression.

The log(H α /[N II]) versus log(H α /[S II]) intensities, from Table 2, corrected for interstellar extinction, are compared with those of other well-defined phenomena in Fig. 6 (following Sabbadin et al. 1977; Cantó 1981).

2.2.2. High dispersion

Observations of G70.5+1.9 were made with the Manchester Echelle Spectrometer (MES-SPM – see Meaburn et al. 1984, 2003) combined with the 2.1-m San Pedro Martir telescope on 23 August, 2005. A SiTe CCD was the detector with 1024×1024 , 24 μm pixels although 2×2 binning was employed throughout the observations.

Spatially resolved, long-slit H α , [N II] & [O III] line profiles were obtained with the MES-SPM along the line marked 5 in Fig. 2. This is only a partial length of the full NS slit as relevant emission only occurred over small sections of the full slit length. The increments along the slit length each corresponds to $0''.63$.

In this spectroscopic mode, MES-SPM has no cross-dispersion, consequently, for the present observations, a filter of 90 \AA bandwidth was used to isolate the 87th echelle order containing the H α and [N II] nebular emission lines and one of 60 \AA bandwidth for [O III]. The slit width was always 150 μm

Table 2. continued.

Line (\AA)	Area 3a			Area 3b			Area 3c			Area 3d			Area 4a		
	F^a	I^b	S/N^c	F	I	S/N	F	I	S/N	F	I	S/N	F	I	S/N
H β 4861	25	34	(83)	28	34	(12)	26	33	(19)	25	34	(23)	27	34	(9)
[O III] 4959	9	12	(33)	26	31	(12)	32	40	(24)	14	19	(15)	11	14	(3)
[O III] 5007	25	33	(89)	84	100	(37)	109	136	(80)	46	61	(48)	57	70	(20)
[O I] 6300	7	7	(40)	–	–	–	–	–	–	–	–	–	50	51	(10)
[O I] 6364	3	3	(18)	–	–	–	–	–	–	–	–	–	17	17	(5)
[N II] 6548	22	22	(127)	27	27	(20)	18	18	(23)	24	24	(40)	–	–	–
H α 6563	100	100	(383)	100	100	(59)	100	100	(98)	100	100	(123)	100	100	(44)
[N II] 6583	72	72	(317)	82	82	(53)	65	65	(72)	76	76	(107)	43	43	(17)
[S II] 6716	65	64	(301)	55	54	(39)	45	44	(55)	44	43	(67)	95	94	(42)
[S II] 6731	46	45	(218)	36	35	(27)	31	30	(38)	32	31	(52)	64	63	(31)
Absolute H α flux ^d	6.9			1.6			1.8			3.7			0.5		
[S II]/H α	1.09 \pm 0.01			0.89 \pm 0.02			0.74 \pm 0.01			0.74 \pm 0.01			1.57 \pm 0.05		
$F(6716)/F(6731)$	1.41 \pm 0.01			1.53 \pm 0.07			1.45 \pm 0.05			1.37 \pm 0.03			1.48 \pm 0.06		
[N II]/H α	0.94 \pm 0.01			1.09 \pm 0.02			0.83 \pm 0.01			1.00 \pm 0.01			0.43 \pm 0.01		
[O III]/H β	1.32 \pm 0.02			3.80 \pm 0.32			5.30 \pm 0.28			2.40 \pm 0.11			2.50 \pm 0.31		
$c(\text{H}\beta)$	0.38 \pm 0.02			0.24 \pm 0.11			0.31 \pm 0.07			0.38 \pm 0.06			0.28 \pm 0.15		
E_{B-V}	0.25 \pm 0.01			0.16 \pm 0.07			0.21 \pm 0.05			0.25 \pm 0.04			0.19 \pm 0.10		
Line (\AA)	Area 4b			Area 4c			Area 4d			Area 4e			Area 4f		
	F^a	I^b	S/N^c	F	I	S/N	F	I	S/N	F	I	S/N	F	I	S/N
H β 4861	26	33	(30)	23	34	(22)	22	34	(37)	28	34	(30)	35	36	(29)
[O III] 4959	7	9	(10)	6	9	(7)	11	16	(20)	11	13	(15)	14	15	(14)
[O III] 5007	35	44	(46)	37	52	(39)	50	73	(92)	41	49	(50)	52	54	(48)
[O I] 6300	–	–	–	6	6	(7)	11	12	(24)	10	10	(12)	13	13	(9)
[O I] 6364	–	–	–	–	–	–	3	3	(8)	4	4	(6)	3	3	(2)
[N II] 6548	28	28	(59)	20	20	(34)	29	29	(86)	24	24	(46)	18	18	(27)
H α 6563	100	100	(156)	100	100	(118)	100	100	(215)	100	100	(144)	100	100	(109)
[N II] 6583	88	88	(148)	66	66	(90)	101	101	(233)	80	80	(127)	57	57	(73)
[S II] 6716	66	65	(127)	51	49	(78)	68	66	(182)	55	54	(101)	37	37	(58)
[S II] 6731	48	47	(100)	34	33	(56)	49	47	(131)	41	40	(81)	26	26	(42)
Absolute H α flux ^d	2.6			3.4			2.7			2.6			1.9		
[S II]/H α	1.12 \pm 0.01			0.82 \pm 0.01			1.13 \pm 0.01			0.94 \pm 0.01			0.63 \pm 0.01		
$F(6716)/F(6731)$	1.38 \pm 0.02			1.48 \pm 0.03			1.40 \pm 0.01			1.35 \pm 0.02			1.42 \pm 0.04		
[N II]/H α	1.16 \pm 0.01			0.86 \pm 0.01			1.30 \pm 0.01			1.04 \pm 0.01			0.75 \pm 0.01		
[O III]/H β	1.60 \pm 0.06			1.80 \pm 0.09			2.60 \pm 0.08			1.82 \pm 0.07			1.92 \pm 0.08		
$c(\text{H}\beta)$	0.31 \pm 0.04			0.47 \pm 0.06			0.53 \pm 0.04			0.24 \pm 0.04			0.16 \pm 0.05		
E_{B-V}	0.21 \pm 0.03			0.31 \pm 0.04			0.35 \pm 0.03			0.16 \pm 0.03			0.11 \pm 0.03		

^a Observed surface brightness normalized to $F(\text{H}\alpha) = 100$ and uncorrected for interstellar extinction.

^b Intrinsic surface brightness normalized to $F(\text{H}\alpha) = 100$ and corrected for interstellar extinction.

^c Numbers in parentheses represent the signal-to-noise ratio of the quoted fluxes.

^d In units of $10^{-16} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ arcsec}^{-2}$.

The letters next to the areas number indicate different apertures extracted along the slit.

Listed fluxes are a signal-to-noise weighted average of two fluxes for areas 2, 3 and 4.

The emission line ratios [N II]/H α and [O III]/H β are calculated using the values corrected for interstellar extinction.

which is $\approx 1.9''$ on the sky and 10 km s^{-1} spectral halfwidth. Each integration time was 1800 s.

The longslit spectra were cleaned of cosmic rays and calibrated in wavelength to $\pm 1 \text{ km s}^{-1}$ accuracy in the usual way using STARLINK FIGARO software. The greyscale representation of the position-velocity (pv) arrays of H α , [N II] & [O III] line profiles for the partial slit length shown in Fig. 2 for Slit 5 are shown in Fig. 5. As no standard star was observed, the absolute surface brightnesses are unreliable and will not be presented here.

3. Discussion

Detailed optical observations have been performed in an attempt to understand the nature of G70.5+1.9. The lower ionization images in H α + [N II], [S II], [O II] and the higher ionization image in [O III] reveal several filamentary structures. In addition,

velocity resolved profiles have been obtained along a specific slit position. The current data point to a shock heated origin of the optical emission.

The SNR origin of the proposed candidate remnant is strongly suggested by the positions of the line ratios in Fig. 6 compared with those of Herbig-Haro objects (HH-objects), HII regions and planetary nebulae (PNe). They follow closely the shape of those observed for those of shock ionized evolved SNRs. The [O III]/H β ratio is a very useful diagnostic tool for complete or incomplete shock structures (Raymond et al. 1988). All spectra suggest complete shock structures except that from Area 1a with an [O III]/H β ratio of ~ 26 . Typically, this ratio is below ~ 6 (Cox & Raymond 1985; Hartigan et al. 1987). However, this limit is easily exceeded in case of shocks with incomplete recombination zones like in Area 1a. Since the long-slit spectra do not cover the full extent of the source, we can use our flux calibrated images to map areas with incomplete shock structures. Since our H α + [N II] filter transmits equally the [N II]6548,

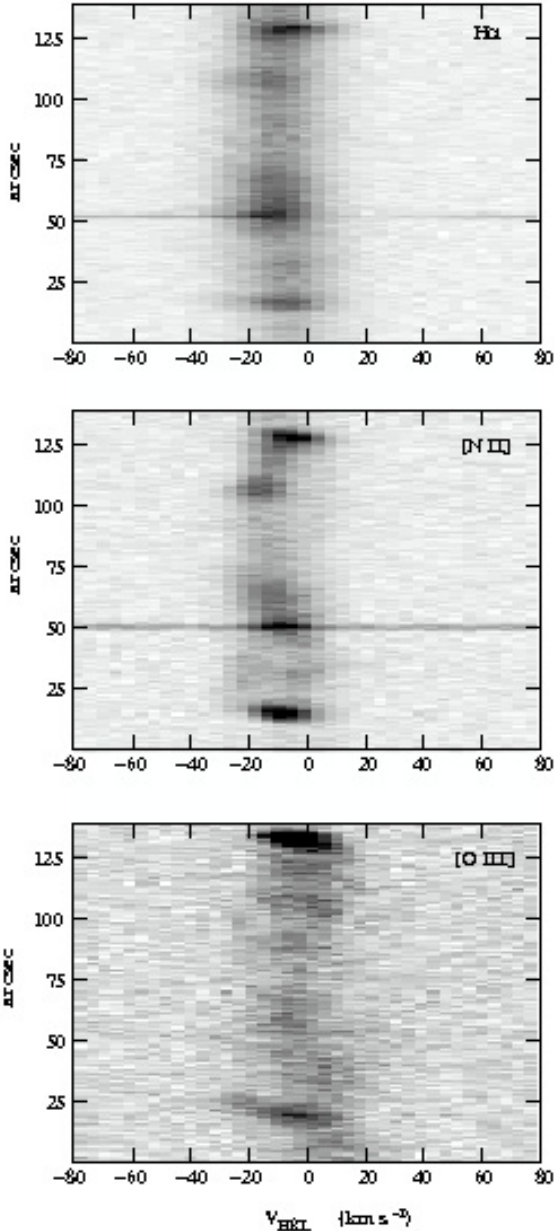


Fig. 5. Greyscale representations of the position-velocity arrays of the $H\alpha$, $[N II]$ & $[O III]$ line profiles from that part of the length of slit 5 as shown in Fig. 2.

6584 Å, and $H\alpha$ lines we can estimate the $H\alpha$ flux as $\sim 1/2$ of the flux measured in the $H\alpha + [N II]$ filter. Assuming that the $H\alpha/H\beta$ ratio is ~ 4 all over this source, as the long-slit spectra suggest, we estimate the $[O III]/H\beta$ ratio as $\sim 8 [O III]/H\alpha + [N II]$. Interestingly, we find that the area further to the east, i.e. between $\alpha \sim 20^h01^m08^s$ and $20^h01^m14^s$ is dominated by such structures.

The motions of the filaments as seen in Fig. 5 also suggest an evolved SNR origin. The line profiles are single over the bright filaments, where expansive motions are expected to be tangential along the line of sight, but become split, albeit by only a few tens of km s^{-1} towards the fainter regions.

It remains possible that this isolated and irregular group of filaments is part of a wider structure but is being seen through holes in intervening clouds, leading to patchy optical interstellar extinction. The correlation of these filaments with the structures reported by Mavromatakis & Strom (2002) is also not clear.

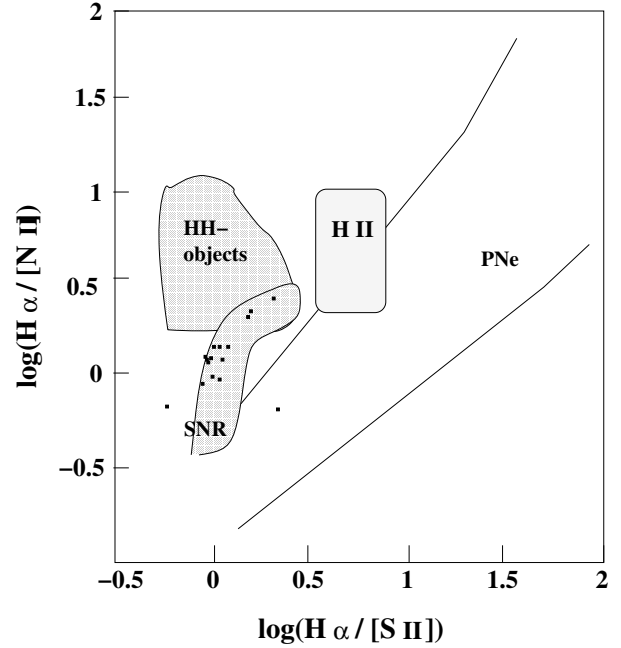


Fig. 6. Diagnostic diagram (Sabbadin et al. 1977; Cantó 1981), where the positions of line ratios listed in Table 2, from Areas 1a to 4f, are shown as black squares.

The current data alone are not sufficient to claim a correlation. The $H\alpha/H\beta$ ratios in Table 2 can be used to estimate the variations in logarithmic extinction coefficient c over this source, assuming an intrinsic ratio of 3 and the interstellar extinction curve of Kaler (1976) as implemented in the nebular package (Shaw & Dufour 1995) within the IRAF software. The signal to noise weighted average of the observed $H\alpha/H\beta$ ratios in Table 2 is $4.0 (\pm 0.1)$ derived from all apertures and spectra available. The logarithmic extinction c is then 0.40, which is equivalent to an A_V of 0.8 and an $E(B - V)$ of 0.27. Here we have assumed $E(B - V) = 0.664 c$ (Kaler 1976; Aller 1984). $E(B - V)$ values are listed in Table 2 along with other parameters, where it can be seen that statistically significant $H\alpha/H\beta$ ratios vary from 2.9 (area 4f) to 4.5 (area 4d).

The average hydrogen column density derived from the statistical relation of Predehl & Schmidt (1995) is $1.4 \times 10^{21} \text{ cm}^{-3}$, while the total galactic hydrogen column density in the direction of the candidate remnant is around $1 \times 10^{22} \text{ cm}^{-3}$ (Kalberla et al. 2005; and Dickey & Lockman 1990). It is clear that the value based on the optical data and a statistical relation is lower by a factor of ~ 6 than the estimated total NH. This implies that the detected structures are closer to us than the total distance to the outer part of the galaxy in that specific direction. The use of the code given by Hakkila et al. (1997) to calculate the visual extinction in the direction of G70.5+1.9 supports the above suggestion, given our measurements of optical extinction (with a typical color excess of 0.27). It is likely that the distance to the proposed candidate remnant is less than 1 kpc.

A crucial issue concerns the existence of radio emission. Radio emission in the area of the optical structures is detected in the low resolution ($7'$) 4850 MHz images of the Green Bank survey (Gregory & Condon 1991). Given the low resolution, it is very hard to state any spatial correlation. We have also examined the higher resolution CGPS data at 1420 MHz (Taylor et al. 2003) but no prominent emission was detected. The 3σ upper limit obtained is $\sim 4 \text{ mJy/beam}$. We note here that there are also

other SNRs that do not display radio emission, at least at the detection level of the corresponding observations (e.g. Dickel et al. 2001; Filipović et al. 2008; Stupar et al. 2008). Radio observations in different wavelengths should be performed to provide conclusive evidence of the nature of the source as an evolved SNR. Finally, we have searched ROSAT data for X-rays but none have been detected in the area of the optical emission.

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