

# Ultra-high energy cosmic rays detected by the Pierre Auger observatory

## First direct evidence, and its implications, that a subset originate in nearby radiogalaxies

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### ABSTRACT

*Context.* The Pierre Auger Collaboration has reported 27 Ultra-High Energy Cosmic Ray Events (UHECRs) with energies above  $56 \times 10^{18}$  eV (56 EeV) and well determined arrival directions as of 2007 August 31. They find that the arrival directions of these UHECRs are not isotropic, but instead appear correlated with the positions of nearby Active Galactic Nuclei (AGNs) from the catalog of Véron-Cetty & Véron.

*Aims.* Our aim was to determine the sources of these UHECRs by comparing their arrival directions with more complete and/or comprehensive astronomical source catalogs.

*Methods.* We have cross correlated the arrival directions of the UHECRs with the positions of supernovae, radio supernovae, galaxies, active galaxies, radiogalaxies, and clusters of galaxies, all at distances within  $\sim 200$  Mpc.

*Results.* Four (eight) of the 27 UHECRs with energy greater than 56 EeV detected by the Pierre Auger Observatory have arrival directions within  $1.5^\circ$  ( $3.5^\circ$ ) of the extended ( $\geq 180$  kpc) radio structures of nearby radiogalaxies or the single nearby BL Lac with extended radio structure. Conversely the radio structures of three (six) of all ten nearest extended radiogalaxies are within  $1.5^\circ$  ( $3.5^\circ$ ) of a UHECR; three of the remaining four radiogalaxies are in directions with lower exposure times. This correlation between nearby extended radiogalaxies and a subset of UHECRs is significant at the 99.9% level. For the remaining  $\sim 20$  UHECRs, an isotropic distribution cannot be ruled out at high significance. The correlation found by the Auger Collaboration between the 27 UHECRs and AGNs in the Véron-Cetty & Véron catalog at  $D \lesssim 71$  Mpc has a much lower significance when one considers only the  $\sim 20$  UHECRs not “matched” to nearby extended radiogalaxies. No clear correlation is seen between UHECRs and supernovae, supernova remnants, nearby galaxies, or nearby groups and clusters of galaxies.

*Conclusions.* Nearby extended radiogalaxies are the most likely source of at least some UHECRs detected by the Pierre Auger Observatory. The remaining UHECRs are not inconsistent with an isotropic distribution; their correlation to nearby AGNs is much less significant than earlier estimated. This is the first direct observational proof that radio galaxies are a significant source of UHECRs. The primary difference between the UHECR detections at the Pierre Auger Observatory and previous experiments, e.g. AGASA, may thus primarily be that the Southern Hemisphere is more privileged with respect to nearby radiogalaxies with highly extended radio jets and lobes.

**Key words.** ISM: cosmic rays – galaxies: active – galaxies: jets

## 1. Introduction

Ultra-High Energy Cosmic Rays (UHECRs) are protons or light nuclei with energies greater than about  $10^{19}$  eV (10 EeV). When these highly energetic particles enter the earth’s atmosphere they produce a shower of secondary particles and excite the Nitrogen molecules in the atmosphere; both effects can be detected from the ground. Recent advances in ground shower detectors allow a precise measurement of both the initial energy and the arrival direction of the UHECR above the earth’s atmosphere, e.g. AGASA (Takeda et al. 1999) and the Pierre Auger Observatory (Abraham et al. 2004; de Mello Neto et al. 2007). While UHECRs are rare they have a unique astronomical potential: UHECRs with energies  $\geq 50$  EeV are not expected to be deflected significantly by Galactic or intergalactic magnetic fields. Their measured arrival directions can therefore be traced directly back to the originating source within the observatory measurement uncertainties.

Suspected sources of high energy cosmic rays include core collapse Supernovae (SNII), Supernova Remnants, Pulsars, Active Galaxies – especially radio jets and radio lobes, and gamma rays bursts (for reviews and lectures see Biermann 1993; Kachelriess 2008; Hillas 1999, and references therein). In these sources, the cosmic rays are posited to undergo successive accelerations via scattering off energetic charged particles and/or in shocks. More exotic explanations include an origin in dark matter annihilation.

The main energy loss for UHECR propagating over cosmological distances is expected to be pion-production, the so called Greisen Zatsepin Kuzmin (GZK) effect (Greisen 1966; Zatsepin & Kuzmin 1966). In this process, a UHECR interacts with a CMB photon and loses an estimated  $\sim 30\%$  of its energy. Both the energy loss and the mean free path between interactions depend on several details; energy loss length predictions vary between 20 and 100 Mpc (see e.g. Stanev 2004; Kachelriess 2008;

PA08). If the GZK effect is at work then UHECRs with energy above  $\sim 50$  EeV are expected to originate primarily in sources closer than this predicted energy loss length.

The Pierre Auger Observatory (hereafter PAO) in Argentina (Abraham et al. 2004; de Mello Neto et al. 2007), an array of 1600 Cerenkov detectors spread over  $3000 \text{ km}^2$  plus six optical telescopes, is designed to measure arrival directions and energies of cosmic rays via their secondary particle showers and their atmospheric fluorescence. Latest details on the instrumentation and exciting science results being produced by the Auger Observatory and Collaboration can be found at the Auger web-site <http://www.auger.org>. The array has been in partial operation since January 2004. The Auger collaboration has reported 81 events between 2004 January 1 and 2007 August 31 with reconstructed energies above 40 EeV and zenith angles smaller than  $60^\circ$ ; of these, 27 have energies above 56 EeV (Pierre Auger Collaboration et al. 2007, 2008, hereafter PA07, PA08). The origins of the latter 27 UHECRs are the focus of this paper.

The Auger Collaboration found that the arrival directions of the 27 UHECRs with energies above 56 EeV are not isotropic at a 99% significance level (PA07, PA08). Instead, they found that the arrival directions are correlated with the positions of AGNs within  $\sim 71$  Mpc (PA07, PA08, Mollerach et al. 2008). For this analysis they used all AGNs (and galaxies with H II nuclei) in the catalog of Véron-Cetty & Véron, 12th edn. (Véron-Cetty & Véron 2006). In their statistical correlation of all 81 UHECRs at energy above 40 EeV with the Véron-Cetty & Véron catalog of AGNs, PA07, PA08 found that the highest correlation between UHECRs and AGNs was obtained for a maximum AGN distance of 71 Mpc, an angular separation of  $3.2^\circ$  between “matched” AGNs and UHECRs, and an energy threshold for UHECRs of 56 EeV. This maximum distance, 71 Mpc, is in line with the expectations of the GZK effect. Gorbunov et al. (2008), in an astro-ph comment, disagreed with this finding, pointing out that weighting each AGN with its distance would predict a very different distribution of UHECRs than that found by the PAO.

Our interest in this topic was provoked when a quick examination of the 27 UHECR events showed that the direction of at least five of the 27 UHECR events were directly in the line of sight to nearby extended radiogalaxies. In this article, we attempt a more comprehensive correlation between the arrival directions of the 27 UHECRs with energies above 56 EeV and catalogs of potential sources of UHECRs, both Galactic and Extragalactic. Section 2 describes the data used in our study, Sect. 3 describes the principle results obtained, and Sect. 4 contains a brief discussion and conclusions of our study. Distances to galaxies are calculated using  $72 \text{ km s}^{-1} \text{ Mpc}^{-1}$ , except for relatively nearby galaxies for which we use distances as referenced.

## 2. Data

Positions and energies of the 27 UHECRs detected by the PAO with energies  $\geq 56$  EeV were taken from PA08. The figures in this publication are slightly different from those of the Auger collaboration: PA07 and PA08 used the Aitoff projection (an equidistant projection with 2:1 aspect ratio), while we use the Aitoff-Hammer projection (an equal area projection). The latter is more commonly used in astronomy and is typically referred to as simply the “Aitoff” projection, e.g. the IDL procedure `aitoff.pro`. Positions and energies of the UHECRs detected by AGASA are taken from Hayashida et al. (2000) which is an update of the data listed in Takeda et al. (1999).

Our catalogs of astronomical sources are drawn from various publications and web-based source lists. In all cases we attempted to assemble as complete a sample as possible. Further, following the results of PA07, PA08, and the expectations from the GZK effect, we considered only objects within a distance of  $\sim 200$  Mpc. For matches to the PAO UHECRs we consider only those sources which rise above elevation  $30^\circ$  at PAO since the PAO UHECR list includes only UHECRs detected above this elevation.

Our list of extragalactic supernova remnants, galaxies, and galaxy clusters were drawn from NED (Madore et al. 1992) using the “All Sky Search” option. The list of Gamma Ray Bursts (GRBs) was taken from the Gamma-ray Burst Real-Time Sky Map<sup>1</sup>. For Galactic Supernova Remnants we used the list maintained by Green<sup>2</sup>. Extragalactic Radio Supernova were taken from Kurt Weiler’s list at the web-site of the Naval Research Laboratory<sup>3</sup>. For extragalactic supernova, we used the Sternberg Astronomical Institute (SAI) Supernova Catalog<sup>4</sup> maintained by Tsvetkov et al. In the case of radiogalaxies and radio sources we used several surveys and catalogs, including NVSS (Condon et al. 1998), SUMMS (Bock et al. 1999), and NED. The list of extragalactic jets is based on that compiled by Liu & Zhang (2002); to this list we added data on the galaxies as given in NED, and re-measured the total flux and total extent of the radio emission from NVSS and SUMMS maps. The total extent of the radio emission therefore includes both radio jets and any radio lobes. This extent is typically referred to as the Largest Angular Size (LAS) or Largest Linear Size (LLS). In this work we added the LLS of the individual jets and lobes as scalars instead of vectors in order to discount the effects of jet and lobe bending.

To recreate the results of the Auger Collaboration (PA07, PA08), we used the 12th edn. of the Véron-Cetty & Véron (2006) catalog of AGNs. The list of AGNs (`agn.dat`) in this catalog includes nuclei with H II spectra which were previously misclassified as AGNs. These should ideally be left out of the correlation analysis. From the AGN statistics quoted in PA07, PA08, it appears that the Auger Collaboration included these H II nuclei in their analysis, so for consistency we do the same. Additionally, the list of QSOs (`qso.dat`) lists seven galaxies with redshift = 0. Four of these are truly nearby galaxies, and the other three are probably high redshift quasars. Again, for consistency with PA07 and PA08 we keep the latter three in our analysis.

We have also used the Fanaroff-Riley (FR) classification for radiogalaxies (Fanaroff & Riley 1974). This division is based on whether the jets increase (FR II) or decrease (FR I) in brightness with increasing distance from the nucleus. A clear division in radio luminosity is found between the two classes, with FR IIs being more luminous in the radio than FR Is (Owen & Ledlow 1994).

## 3. Results

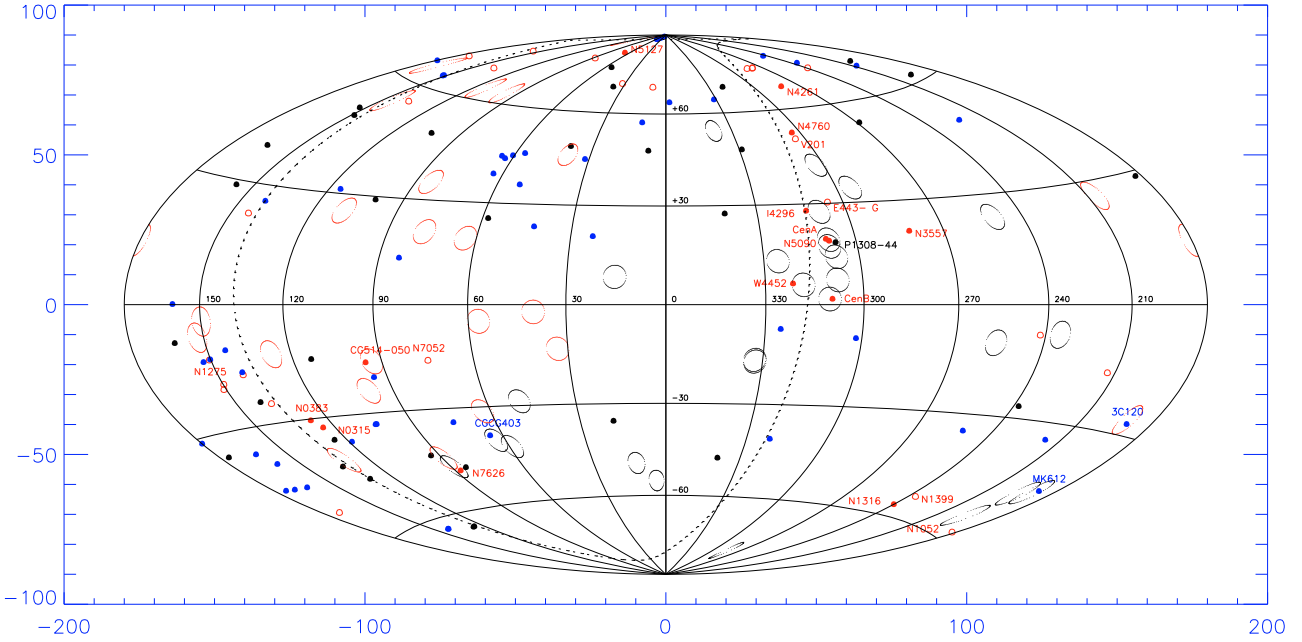
A comparison of the arrival directions of the 27 UHECRs and the positions of nearby galaxies with radio jets (all galaxies listed in Liu & Zhang 2002) is shown in Fig. 1. We have divided the galaxies with radio jets into three redshift bins: “nearby” ( $D \leq 75$  Mpc), “intermediate” ( $75 \text{ Mpc} < D \leq 150$  Mpc), and

<sup>1</sup> <http://grb.sonoma.edu>

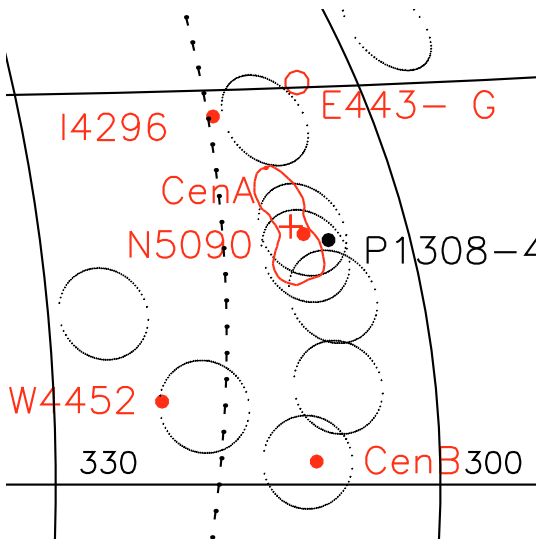
<sup>2</sup> <http://www.mrao.cam.ac.uk/surveys/snrs/snrs.data.html>

<sup>3</sup> <http://rsd-www.nrl.navy.mil/7213/weiler/kwdata/RSNtable.txt>

<sup>4</sup> <http://www.sai.msu.su/sn/sncat/>



**Fig. 1.** A comparison of the arrival directions of the 27 UHECRs with the positions of galaxies hosting radio jets. PAO UHECR events with energies above 56 EeV are shown as black circles with radius  $3.5^\circ$ . AGASA UHECR events with energies above 56 EeV are shown as red circles with radius  $3.5^\circ$ . The colour symbols mark the positions of galaxies with radio jets at  $D \leq 75$  Mpc (red),  $75 \text{ Mpc} < D \leq 150$  Mpc (blue), and  $150 \text{ Mpc} < D \leq 210$  Mpc (black). For galaxies at  $D \leq 75$  Mpc, filled symbols are used for galaxies with radio structures more extended than 180 kpc, and open symbols for galaxies with radio structures less extended than 180 kpc. Galaxy names are marked in some of the less crowded regions in the plot, in the same color as the corresponding symbol. The SuperGalactic plane is marked by the dashed line. In this and following figures, we use an Aitoff Hammer Projection. The orientation and labelling of the longitude axis follows PA07 and PA08 for easier comparison.



**Fig. 2.** A detail of Fig. 1 showing the region around the radiogalaxy Cen A. Symbols are the same as in Fig. 1 except in the case of Cen A, for which we use a cross to mark the galaxy center and a polygon to outline its radio jet structure as seen in a 403 MHz map.

“distant” ( $150 \text{ Mpc} < D \leq 210 \text{ Mpc}$ ). In the closest redshift bin,  $D \leq 75 \text{ Mpc}$ , we distinguish between galaxies with radio structures more extended than 180 kpc (“extended”) and those with radio structures less extended than 180 kpc (“compact”). We calculated the total extent of the radio structure of any given galaxy by summing the lengths of the two radio jets and/or lobes as scalars instead of vectors. The total radio extent includes radio jets and any radio lobes.

There is no intrinsic reason to use a radius  $\psi = 3.5^\circ$  for a “match” between UHECR arrival direction and the position of an astronomical source. The instrumental resolution of the PAO is about  $1^\circ$  at these UHECR energies, to which one must add a small angle – perhaps up to a few degrees (Stanev 2004) – to account for the deflection of the UHECR by magnetic fields. The best correlation between UHECRs and nearby AGNs found by PA07, PA08 is based on considering  $\psi \sim 3.2^\circ$  as a match. However, as we show later (e.g. Fig. 7), the value of  $\psi$  in the analysis of PA07, PA08 should really be  $\psi \sim 3.1^\circ$ – $4.3^\circ$ . For the above reasons, we consider two values of  $\psi$  for a match:  $\psi = 1.5^\circ$  and  $\psi = 3.5^\circ$ . The distance limit for “nearby” galaxies – 75 Mpc – is similar to the value  $d_{\text{max}} \sim 71 \text{ Mpc}$  found by PA07 and PA08. The division between “compact” and “extended” radio structures – 180 kpc – is the one “new” free parameter in our analysis. We note that this value is close to the minimum linear size of the radio structures of a typical radio galaxy (e.g. Hardcastle et al. 1998).

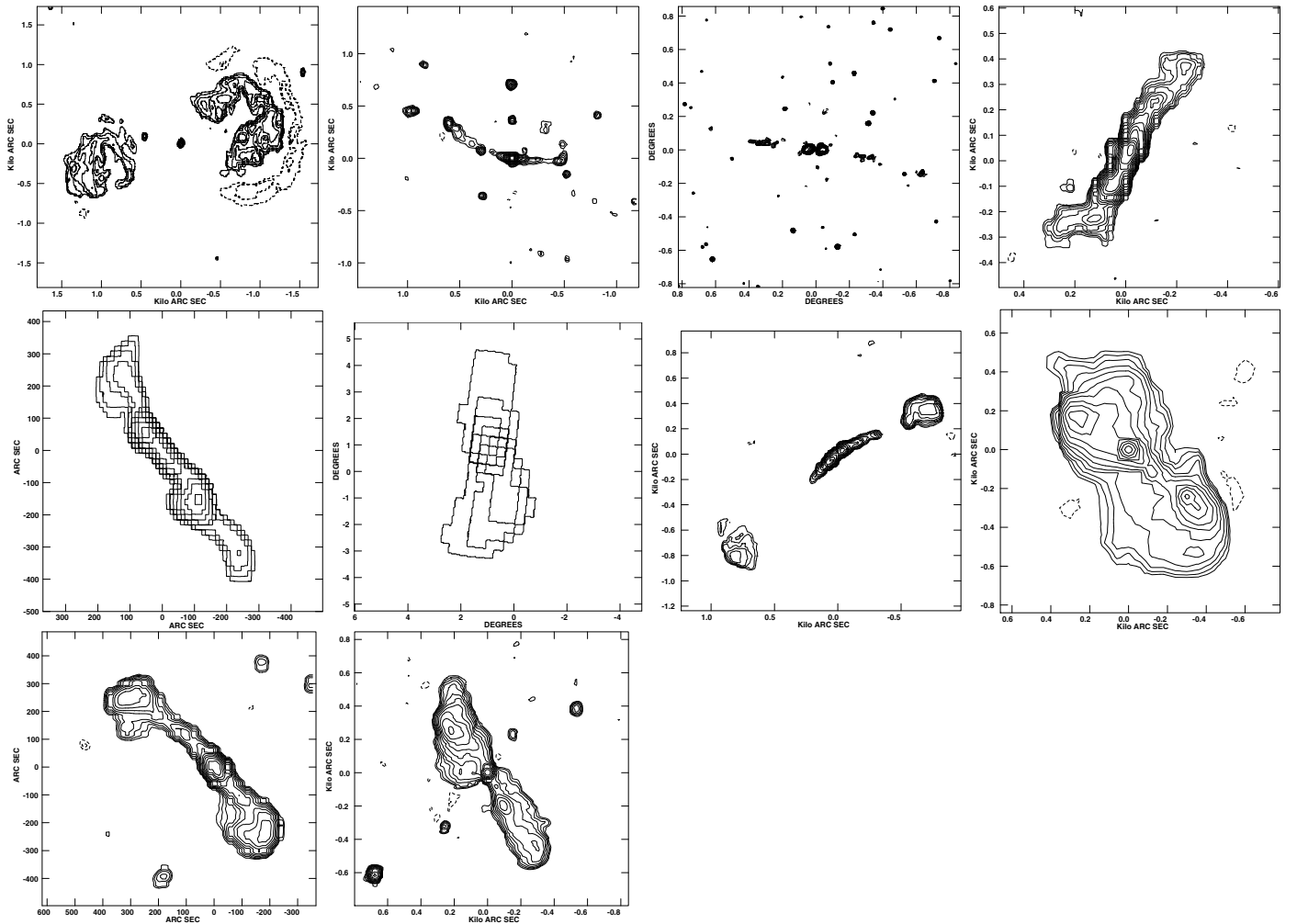
Of all 27 UHECRs with energy above 56 EeV detected by the PAO four (eight) are within  $\psi = 1.5^\circ$  ( $\psi = 3.5^\circ$ ) of the radio structures of nearby extended radiogalaxies in the field of view of the PAO. Conversely, of all ten nearby extended radiogalaxies in the field of view of the PAO three (six) have radio structures which can be matched to within  $1.5^\circ$  ( $3.5^\circ$ ) of a UHECR event. Of the remaining four “unmatched” radiogalaxies, three – NGC 1316, NGC 4261 and NGC 4760 – fall outside the area of maximum exposure at the PAO. For the intermediate and distant redshift bins there is little match between extragalactic radio jets and UHECRs; the only two matches within  $3.5^\circ$  are CGCG 403-019, a Seyfert/BL Lac at 112 Mpc, and Mrk 612, a Seyfert galaxy at 85 Mpc. Under the assumption of an isotropic distribution of UHECR arrival directions, one would expect less than 1 match to the nearby extended radiogalaxies instead of the

**Table 1.** Nearby extended radiogalaxies in the field of view of the Pierre Auger observatory.

Object name (1)	$l$ (Degrees) (2)	$b$ (Degrees) (3)	Dist. (Mpc) (4)	$S_{4.85 \text{ GHz}}$ (Jy) (5)	$S_{1.4 \text{ GHz}}$ (Jy) (6)	$S_{843 \text{ MHz}}$ (Jy) (7)	$S_{408 \text{ MHz}}$ (Jy) (8)	Size (kpc.) (9)	Morph. (10)	Ref. (11)
NGC 1316 (Fornax A)	240.162919	-56.690021	24.5	...	...	169	259	375	Rel.Dbl	2, 3
MRK 0612	187.717486	-45.039107	85.2	...	0.013	...	...	2.6	Sey.2	8
NGC 3557	281.578358	21.089037	42.5	0.32	0.63	...	0.96	310	FR I	2
NGC 4261 (3C 270)	281.804700	67.372841	31.1	4.6	19.4	...	41.5	436	FR I	2
NGC 4760	303.609402	52.375264	66.2	0.37	0.9	...	4.7	322	FR I	5, 7
NGC 5090	308.611065	18.838673	47.6	1.98	5.8	7.16	...	187	FR I	3
NGC 5128 (Cen A)	309.515874	19.417325	3.4	62.84	278	392	2746	453	Int. FRI/II	3
IC 4296	313.538349	27.973003	51.9	1.88	...	14.9	34	1564	Int FRI/II	2, 3
PKS 1343-60 (Cen B)	309.721407	1.731590	53.8	27	79	...	...	256	FR I	6
WKK 4452	321.594046	6.291192	66.7	...	3.2	3.8	...	358	FRI+FR II	1, 3
CGCG 403-019	64.701023	-38.631307	112.5	0.74	0.78	...	1.64	187	BLLac/Sey.1	4
NGC 7626	87.859111	-48.378945	47	0.25	0.86	...	1.9	212	FR I	2

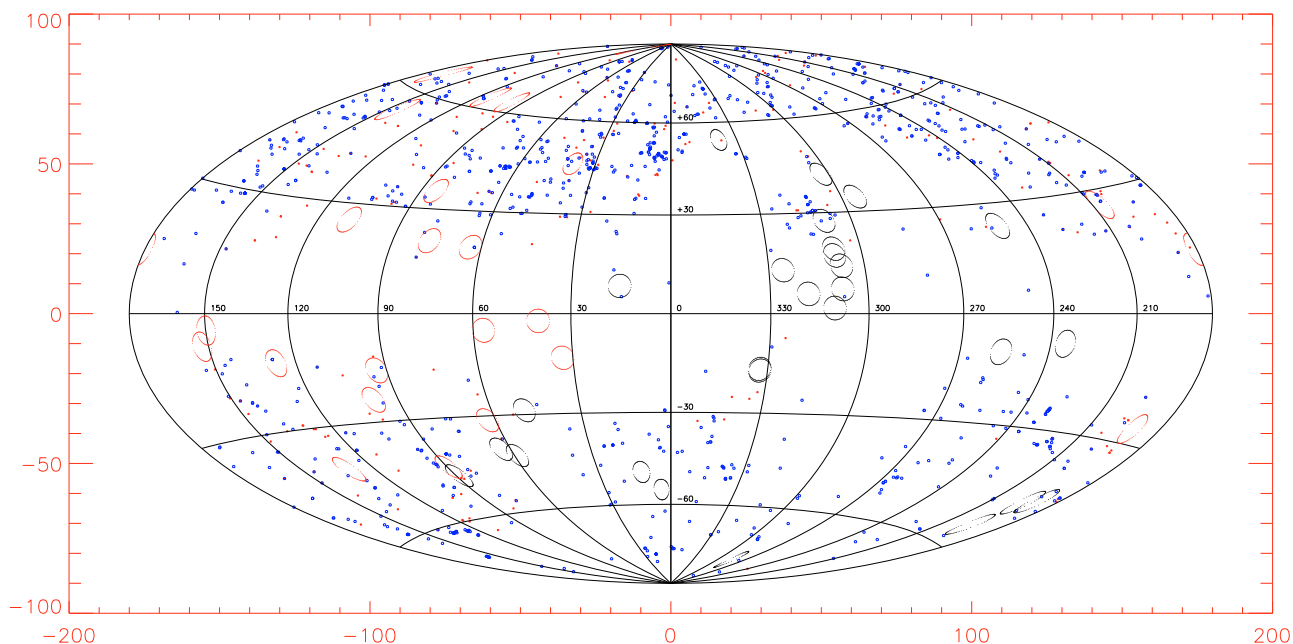
Note. This table lists some relevant properties of all ten nearby and extended radiogalaxies (see text) in the “field of view” (see text) of the Pierre Auger Observatory, plus two Active Galaxies which potentially match to UHECRs detected by the Pierre Auger Observatory.

Note. Columns are: (1) galaxy name; (2) and (3) galactic longitude and latitude; (4) distance; (5) to (8) approximate total flux in low resolution radio maps, taken from NED, measured from NVSS or SUMMS, or from the references in Col. 11; (9) largest linear size, in kpc, of the radio structures; (10) jet morphology (see references in Col. 11). Here FRI/II is intermediate between FR I and FR II, FRI+FR II implies one jet has an FR I morphology and the other a FR II morphology, Rel. Dbl. is a “relaxed double”, and Sey. denotes Seyfert. (11) References: (1) Jones (1986); (2) Birkinshaw & Davies (1985); (3) Jones & McAdam; (4) Ulvestad & Johnston (1984); (5) Saikia et al. (1986); (6) Jones et al. (2001); (7) NVSS; (8) Nagar et al. (1999).

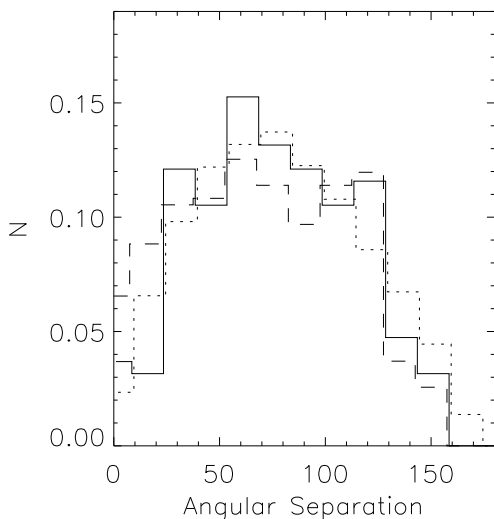


**Fig. 3.** The radio morphologies of the radiogalaxies in Table 1. Unless otherwise specified, the radio maps are from NVSS. From left to right, the contour maps are *top row*: NGC 1316 (SUMMS), NGC 3557, NGC 4261, NGC 4760; *middle row*: NGC 5090 (SUMMS), Cen A (408 MHz), IC 4296, Cen B (SUMMS); *bottom row*: WKK 4452 (SUMMS), NGC 7626. Radiogalaxies which match the positions of UHECRs (*middle and bottom rows*) are preferentially those with intermediate FRI-FR II type characteristics. Those not matched to positions of UHECRs (*top row*) are FR Is or in the case of NGC 1316, a “relaxed double”.





**Fig. 4.** A comparison of the arrival directions of the 27 UHECRs with the positions of clusters of galaxies (as listed and defined by NED). The symbols for UHECRs are the same as in Fig. 1. Red symbols are used for galaxy clusters at  $D \leq 75$  Mpc, and blue symbols for galaxy clusters at  $75 \text{ Mpc} < D \leq 200$  Mpc (see the online manuscript for a color figure).



**Fig. 5.** Histograms of the angular separation between UHECRs detected by the PAO. The dashed line is for all 27 UHECRs with energy above 56 EeV. The solid line is for the subset of 19 UHECRs which are not matched to nearby radiogalaxies with extended radio structures (see text). The dotted line is the expectation for a uniform distribution of the 27 UHECRs given the exposure time weighting in declination for the PAO. The histograms are offset by a few degrees in  $x$  for clarity.

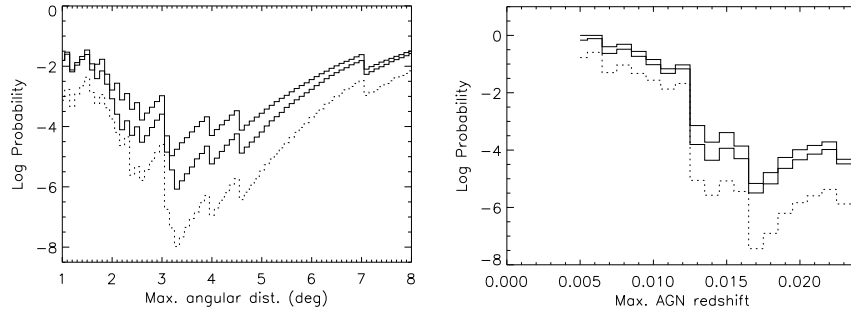
8 matches obtained. This is the main result of this publication: a correlation at 99.9% significance between a subset of UHECRs and nearby extended radiogalaxies.

The match between a UHECR and CGCG 403-019 deserves further comment. CGCG 403-019 is a BL Lac with a Seyfert 1 spectrum (Veron-Cetty & Veron 1993). In the 12th AGN Catalog of Veron-Cetty & Veron (2006) it is listed by its alias PKS 2201+04 and is classified as a Seyfert 1. It is a powerful radio source with a core-jet structure with total extent  $\sim 6'$  or  $\sim 190$  kpc (Ulvestad & Johnston 1984; Laurent-Muehleisen et al. 1993). Within the unified scheme of radiogalaxies, BL Lacs are

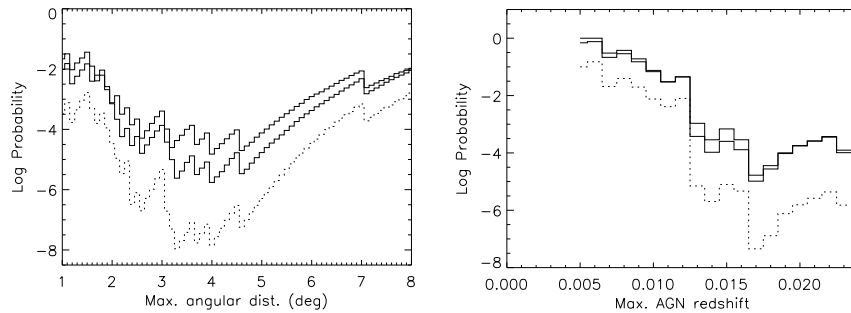
posited to be pole-on FR Is, implying a much larger deprojected radio extent. We therefore consider CGCG 403-019 as an “intermediate” distance “extended” radiogalaxy. Further, CGCG 403-019 is unique in being the nearest BL Lac with extended radio structure. The 12th AGN Catalog of Veron-Cetty & Veron (2006) includes only three confirmed BL Lacs – V Zw 331, RXS J05055+0416, and TEX 0554+534 – and only one probable BL Lac – RXS J21231-1036 – at  $D \leq 150$  Mpc. All four show compact radio emission in NVSS radio maps.

The Auger Collaboration noted that the region around Cen A has an unusually large number of UHECRs; this is also the area with the maximum number of nearby extended radiogalaxies. As seen in Fig. 2, the two UHECRs centered on Cen A can be explained as originating in either of Cen A, NGC 5090 ( $D = 48$  Mpc; total radio extent  $\sim 187$  kpc), or even PKS 1308-441 ( $D = 211$  Mpc; total radio extent  $\sim 1140$  kpc). The event just further south can be matched to the southern jet of Cen A, or to the relatively distant PKS 1308-441. Cen B ( $D = 54$  Mpc; total radio extent  $\sim 250$  kpc) is almost perfectly centered on a UHECR, and WKK 4552 ( $D = 67$  Mpc; total radio extent  $\sim 360$  kpc) is within  $\sim 3^\circ$  of another UHECR. The UHECR just to the north of Cen A is bracketed by IC 4296 ( $D = 52$  Mpc; total radio extent  $\sim 1560$  kpc) and ESO 443-G024 ( $D = 71$  Mpc; total radio extent  $\sim 60$  kpc). The very extended radio structure of the former actually intersects with the  $3.5^\circ$  circle around the UHECR event.

Are the nearby extended radiogalaxies matched to PAO UHECRs different from those not matched to PAO UHECRs? A comparison of the relevant properties of the ten nearest extended radiogalaxies is given in Table 1, and their radio contour maps in Fig. 3. Table 1 also lists the relevant radio properties of CGCG 403-019 and Mrk 0612, the two intermediate distance AGNs which match to a UHECR. There is some difference in the combination of two factors: the radiogalaxies in the direction to UHECRs include all of the subset which have radio jet morphologies closer to FR II, and are all in the area where the PAO has obtained a higher



**Fig. 6.** Probability of the null hypothesis (isotropic distribution) vs. maximum angular distance  $\psi$  (*left panel*) and maximum AGN redshift  $z_{\max}$  (*right panel*). In this figure we followed the analysis of PA07, PA08, i.e. correlating UHECRs to *all* galaxies at  $z \leq 0.024$  in the AGN Catalog 12th edn. (Véron-Cetty & Véron 2006). The dashed lines are the results of using all 27 UHECRs, i.e. the same plots as in PA07 and PA08. The middle solid line is the result of deleting all 4 UHECRs within  $1.5^\circ$  of the radio structures of nearby extended radiogalaxies and the upper solid line is the result of deleting all 8 UHECRs within  $3.5^\circ$  of the radio structures of nearby extended radiogalaxies (see text).



**Fig. 7.** The same as Fig. 6, but this time, of all galaxies listed in the AGN Catalog 12th edn., we do not consider galaxies with H II type nuclei and the three high redshift QSOs which are listed as  $z = 0$  sources (see text). The main difference from Fig. 6 is that the minimum value for the null hypothesis in  $\psi$  is now over a broader range:  $\psi \sim 3.1^\circ$  to  $4.3^\circ$ .

exposure time. In any case, given the small number statistics, i.e. typically  $1 \pm 1$  UHECR events per matched radiogalaxy, it is not surprising that UHECRs are not yet detected towards the remaining four radiogalaxies. Clearly more UHECRs are needed to form any firm conclusion.

Figure 4 compares the arrival directions of UHECRs with the distribution of nearby groups and clusters of galaxies. While there is an over-density of galaxies, clusters of galaxies and UHECRs around the region of Cen A, the overall distributions of nearby galaxies and nearby galaxy clusters do not correlate with the arrival directions of UHECRs. This argues against an origin of UHECRs in material associated with large mass haloes, e.g. Dark Matter. The more detailed analysis of Gorbunov et al. (2008) also predicts a different distribution of UHECRs when the distances to galaxies and clusters are taken into account.

We have also compared the arrival directions of the 27 UHECRs with the positions of Galactic and nearby extragalactic supernovae and supernova remnants, of nearby extragalactic radio supernovae, of nearby BL Lacs, and of gamma-ray bursts. In all cases there appears to be no more than a random overlap.

The analysis of PA07 and PA08 considered a single source population for the UHECRs. How would their results change if one were to consider multiple source populations, at least one of which is nearby extended radiogalaxies? Figure 5 compares the histograms of the angular distribution of the 27 UHECRs with energies above 56 EeV, the subset of 19 UHECRs not towards nearby extended radiogalaxies, and the expectations from a random distribution of 27 UHECRs. The deviation from isotropy discussed by PA07 and PA08 (dashed line in the figure) comes primarily from the clustering of events around Cen A. Deleting

the UHECRs which lie within  $3.5^\circ$  of nearby radiogalaxies (solid line in the figure) gives a distribution in which isotropy cannot be ruled out at high significance.

The Auger Collaboration cross correlated the positions of the 27 UHECRs with the AGNs in the 12th catalog of Véron-Cetty & Véron (2006) by varying the maximum AGN distance ( $z_{\max}$ ), the angular separation between AGN and UHECR to be considered a match ( $\psi$ ), and the threshold energy of UHECRs. They then looked for the minimum cumulative probability of chance coincidences. Note that they apparently included H II nuclei from the catalog in their analysis. In this process (see PA08 for details) they found that the best match was found for  $z_{\max} = 0.017$  ( $D \sim 71$  Mpc) and  $\psi = 3.2^\circ$ . We repeated their cumulative probability scans in the parameters  $z_{\max}$  and  $\psi$ ; we were unable to scan in energy as we do not have access to the position and energy data on the UHECRs with energies less than 56 EeV. The resulting scans are shown in Fig. 6. The dashed line shows the result if all 27 UHECRs are used, i.e. the result obtained by PA08. In this case 19 of the 27 UHECRs match to some AGN or H II nuclei in the catalog. The middle solid line of each panel shows the result of considering only the 23 UHECRs remaining after deleting the four UHECRs which best match (within  $1.5^\circ$ ) the radio structures of nearby extended radiogalaxies (2 UHECRs for Cen A, and one each for Cen B and NGC 7626). In this case 13 of the 23 UHECRs match to some AGN or H II nuclei in the catalog. The upper solid lines show the results of considering only the 19 UHECRs remaining after deleting the four UHECRs mentioned above, plus the UHECRs matched within  $3.5^\circ$  to the South jet of Cen A, to CGCG 403-019, to WKK 4452, and to IC 4296. In this case 11 of the 19 UHECRs match to some AGN or H II nuclei in the catalog. While the overall shape of the scans in  $\psi$  and  $z_{\max}$  remain roughly

the same, the significance of the results decrease significantly. In the case of the scan in  $\psi$ , the minimum at around  $3.2^\circ$  is no longer a clear and unique minimum. We repeated the above exercise after making two modifications to the AGN list: we deleted all galaxies with H II type nuclei (50 galaxies at  $z \leq 0.024$ ) and the three probable high redshift QSOs which are listed as having  $z = 0$  in the AGN Catalog 12th edn. The results are shown in Fig. 7. The main difference with the previous figure is the broader valley in the  $\psi$  scan. Instead of a clearly defined minimum at  $\psi = 3.2^\circ$  one now has a greater range  $\psi \sim 3.2^\circ - 4.3^\circ$ .

#### 4. Discussion and concluding remarks

The arrival directions a subset of the 27 UHECRs with energies above 56 EeV detected by the PAO are statistically most closely related to nearby radiogalaxies with extended radio jets and/or lobes. It is thus likely that nearby (due to the GZK effect) radiogalaxies with extended (radio extent  $\geq 180$  kpc), two sided jets and lobes are the source population of a subset of the UHECRs detected by the PAO. Additionally, there is weak evidence that hosting a jet with morphology closer to FR II makes a radiogalaxy more likely to be a UHECR source. Interestingly, all these factors are consistent with the lack of UHECR detections towards M 87.

The results of the previous section were based only on UHECRs detected by the PAO. We have also tested for correlations against the UHECRs with energy above 56 EeV detected by AGASA (red circles with radius  $3.5^\circ$  in Fig. 1). Among the nearby ( $D \leq 75$  Mpc) galaxies with extended radio jets, the only clear match is to CGCG 514-050, a radiogalaxy at  $D = 72.2$  Mpc with radio extent  $\sim 250$  kpc. The remaining three nearby galaxies with extended radio structures, NGC 315, NGC 383, and NGC 5127, are not matched to within  $3.5^\circ$  of any AGASA UHECR event, though there are several UHECRs at slightly larger distances. Additionally, one AGASA UHECR is close to the location of NGC 7626 and another is centered on the radiogalaxy 3C 120, at  $D = 135$  Mpc and total radio extent  $\sim 315$  kpc.

Among the UHECRs detected by PAO, an isotropic distribution of the  $\sim 20$  UHECRs which are not in the line of sight to nearby extended radiogalaxies cannot be ruled out at high significance. The significance of the correlation found by PA07 and PA08 between UHECR arrival directions and nearby AGNs from the catalog of Véron-Cetty & Véron is much lower when one only considers the  $\sim 20$  UHECRs not matched to nearby extended radiogalaxies and even lower when H II nuclei in the the AGN Catalog are deleted. The findings of PA07, PA08 were probably largely influenced by the correlation between UHECRs and nearby radiogalaxies which we have shown above.

The UHECR arrival directions are not strongly correlated with either supernovae, extragalactic radio-supernovae, or nearby groups and clusters of galaxies.

The main difference between the results of the Pierre Auger Collaboration and those from previous studies of UHECRs, e.g. AGASA, could primarily be that the southern location of the PAO is more privileged with respect to nearby extended radiogalaxies.

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