

LETTER TO THE EDITOR

Transient horizontal magnetic fields in solar plage regions

R. Ishikawa^{1,2}, S. Tsuneta², K. Ichimoto², H. Isobe³, Y. Katsukawa², B. W. Lites⁴, S. Nagata⁵, T. Shimizu⁶,
R. A. Shine⁷, Y. Suematsu², T. D. Tarbell⁷, and A. M. Title⁷

¹ Department of Astronomy, University of Tokyo, Hongo, Bunkyo-ku, Tokyo 113-0033, Japan
e-mail: ryoko.ishikawa@nao.ac.jp

² National Astronomical Observatory of Japan, 2-21-1 Osawa, Mitaka, Tokyo 181-8588, Japan

³ Department of Earth and Planetary Science, University of Tokyo, Hongo, Bunkyo-ku, Tokyo 113-0033, Japan

⁴ High Altitude Observatory, National Center for Atmospheric Research, PO Box 3000, Boulder CO 80307-3000, USA

⁵ Kwasan and Hida Observatories, Kyoto University, Yamashina, Kyoto 607-8471, Japan

⁶ ISAS/JAXA, Sagamihara, Kanagawa 229-8510, Japan

⁷ Lockheed Martin Solar and Astrophysics Laboratory, B/252, 3251 Hanover St., Palo Alto, CA 94304, USA

Received 8 November 2007 / Accepted 21 December 2007

ABSTRACT

Aims. We report the discovery of isolated, small-scale emerging magnetic fields in a plage region with the Solar Optical Telescope aboard Hinode.

Methods. Spectro-polarimetric observations were carried out with a cadence of 34 s for the plage region located near disc center. The vector magnetic fields are inferred by Milne-Eddington inversion.

Results. The observations reveal widespread occurrence of transient, spatially isolated horizontal magnetic fields. The lateral extent of the horizontal magnetic fields is comparable to the size of photospheric granules. These horizontal magnetic fields seem to be tossed about by upflows and downflows of the granular convection. We also report an event that appears to be driven by the magnetic buoyancy instability. We refer to buoyancy-driven emergence as type 1 and convection-driven emergence as type 2. Although both events have magnetic field strengths of about 600 G, the filling factor of type 1 is a factor of two larger than that of type 2.

Conclusions. Our finding suggests that the granular convection in the plage regions is characterized by a high rate of occurrence of granular-sized transient horizontal fields.

Key words. Sun: faculae, plagues – Sun: granulation – Sun: magnetic fields – Sun: photosphere

1. Introduction

Sunspots and active regions are believed to form from magnetic flux globally emerging from the solar interior into the solar atmosphere. The emergence of magnetic flux is due to the buoyancy of large-scale, coherent toroidal flux tubes in the convection zone (Parker 1955). Emerging horizontal fields become more or less vertical to the photosphere, and the vertical fields provide natural connectivity with observed X-ray loops in the corona. This established paradigm is widely supported by observations and theory, but a concern has been raised on theoretical grounds that flux rising buoyantly through the convection zone will fail to maintain a coherent magnetic structure, and therefore fail to emerge (Magara 2001; Moreno-Insertis et al. 1995). Here we report on observations at granular scales of widespread emergence of isolated magnetic fields that may be related to the concern.

Our observations were made with the Solar Optical Telescope (SOT, Tsuneta et al. 2007) aboard Hinode (Kosugi et al. 2007). This telescope performs high-resolution ($\sim 0''.3$) polarimetric observations with a very stable point spread function. Its Spectro-Polarimeter (SP) obtains the complete polarization state (Stokes I , Q , U , and V) of the absorption profiles of the Zeeman-sensitive photospheric lines FeI 630.15 nm and FeI 630.25 nm. We observed a plage region located near the

center of the solar disk. A narrow area of $2''.5$ (8 slit positions, East-West) \times $164''$ (North-South) was rapidly scanned with a coarse pixel size of $0''.32$ by the SP, resulting in the moderate spatial resolution of $\sim 0''.6$, and very high time cadence of 34 s over 40 min.

2. Ubiquitous transient fields in a plage region

We obtain the continuum intensity, I_c , linear polarization, $LP = \int \sqrt{Q^2 + U^2} d\lambda / \int I_c d\lambda$, and circular polarization, $CP = \int V d\lambda / \int I_c d\lambda$. The integration for LP is performed between -21.6 pm and $+21.6$ pm from the average center of the Stokes I profiles (FeI 630.25 nm) for pixels without magnetic signals, and the integration for CP is between $+4.32$ pm and $+21.6$ pm from the center of each Stokes I profile. In the weak field limit, LP is proportional to squares of the transverse field strength, and CP to the line-of-sight field strength.

We discovered the frequent appearance and disappearance of nearly horizontal magnetic structures in the plage region: we find 52 occurrences of transient horizontal magnetic fields in this short time series. These events appear to be randomly located spatially and temporally in the $2''.5 \times 164''$ region observed for a duration of 40 min. These events are not associated with

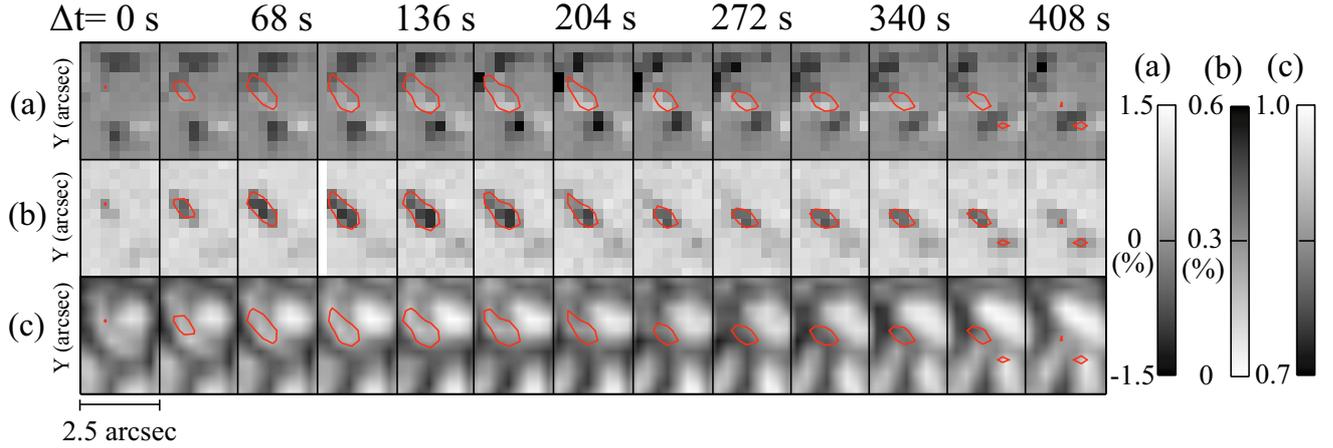


Fig. 1. The plage region located near disk center ($72''$ W, $76''$ S) was observed from 16:00 to 16:40 UT on February 10, 2007. The Stokes profiles at two slit positions with integration time of 1.6 s each and scan step of $0''.16$ each are summed, and 2 pixels along the slit with a pixel size of $0''.16$ are also summed to obtain polarization accuracy better than 0.1%: the images consist of 8 slit positions ($0''.32$ width) with a total scan time of 34 s. The pixel size along the slit is $0''.32$. The evolution of physical quantities for the plage region are shown: **a)** CP (vertical magnetic field), **b)** LP (horizontal magnetic field), **c)** I_c . The region where LP is larger than 0.3% is enclosed by red lines. The emergence of the horizontal magnetic flux starts at $\Delta t = 0$ s. Solar north is up and east to the left in all images of this report.

existing long-lived magnetic fields, and appear in a region of insignificant vertical fields. An episode of such transient horizontal magnetic fields is shown in Fig. 1. The maximum size of the transient horizontal magnetic field is $\sim 1''.4 \times 1''$, and its duration is about 6 min. A remarkable feature is that the horizontal magnetic structure first appears inside the granule, subsequently moves to the inter-granular lane, and finally disappears in the inter-granular lane. The horizontal magnetic structure continues to be smaller than or at most as large as the size of the granule where it appears. Horizontal magnetic fields with negative vertical magnetic fields at one end appear at $\Delta t = 34$ s (black patch in Fig. 1), where there are no pre-existing vertical magnetic fields at $\Delta t = 0$ s. Vertical magnetic fields with positive polarity start to be seen clearly at the opposite end at $\Delta t = 102$ s. After $\Delta t = 238$ s, it is difficult to identify the bipolar configuration around the shrinking horizontal magnetic fields due to the weak vertical magnetic signals.

The transient horizontal fields discovered in the plage regions have lifetimes ranging from a minute to about ten minutes, comparable to the lifetime of granules. Among 52 events, 43 horizontal magnetic structures appear inside the granules, and four appear in inter-granular lanes with the remaining five events ambiguous in position. Since 52 events are detected in the $2''.5 \times 164''$ observing area during the 40 min, a new horizontal field appears every 46 s in the same observing region. The turnover time of the granules is approximately 1000 s with a velocity of 2 km s^{-1} and with a depth comparable to the horizontal scale. There are approximately 182 granules in the observing area, assuming that the size of the granules is $1''.5 \times 1''.5$. 84% of these granules are not associated with stable strong vertical magnetic fields, and we use this smaller sample for estimating the frequency of events. Were every granule to have one embedded horizontal magnetic field structure, the horizontal field would have appeared at the surface every 6.6 s ($\sim 1000 \text{ s}/152$ granules) in the observing area. This shows that more than approximately 10% of the granules have embedded horizontal fields, suggesting a relatively common occurrence of horizontal magnetic fields underneath the plage region.

We examined the Stokes profiles along the strong horizontal magnetic structure at $\Delta t = 136$ s (Fig. 2). The Stokes Q

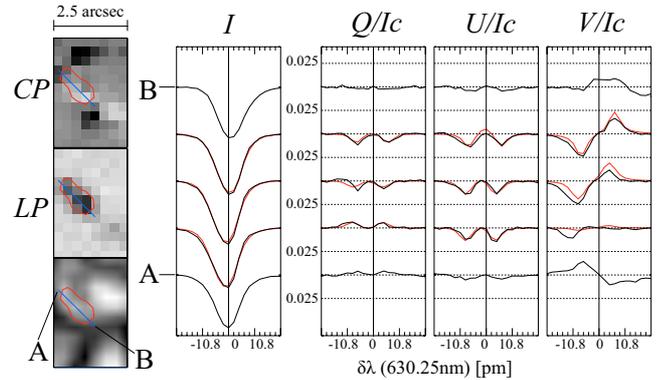


Fig. 2. The frame at $\Delta t = 136$ s in Fig. 1 is enlarged on the left. Stokes profiles (I , Q , U , and V) along the major axis (blue line) are shown to the right with black lines. The fitted profiles are shown with red lines. The profiles marked by A and B correspond to the NE and SW ends of the horizontal field structure, respectively. The vertical line in the Stokes profiles shows the average center of the Stokes I absorption profiles (Fe I 630.25 nm) without magnetic fields, and provides a fiducial for the zero point in velocity. The tick mark indicates 2.16 pm.

and U profiles representing the horizontal magnetic fields are blue-shifted by $0\text{--}1 \text{ km s}^{-1}$ with respect to the line center of each Stokes I profiles. The strong horizontal magnetic field rises along with the upward convective motion of the granule in the early phase of the event; these magnetic fields would be subject to the convective flows. The polarity of the very weak vertical field component detected by the Stokes V profile reverses along the horizontal magnetic structure, indicating either convex- or concave-upward magnetic fields, depending on the resolution of the 180-degree ambiguity in the direction of the horizontal field. Because we follow the emergence of this event from its inception, the only logical configuration is convex-upward; i.e., an emerging omega-shaped loop. This scenario is consistent with the higher upward flow near the center of the granular cell.

The transient horizontal fields reported here differ from the typical emerging flux driven by the Parker instability (Parker 1955; Shibata et al. 1989; Strous et al. 1996; Lites et al. 1998;

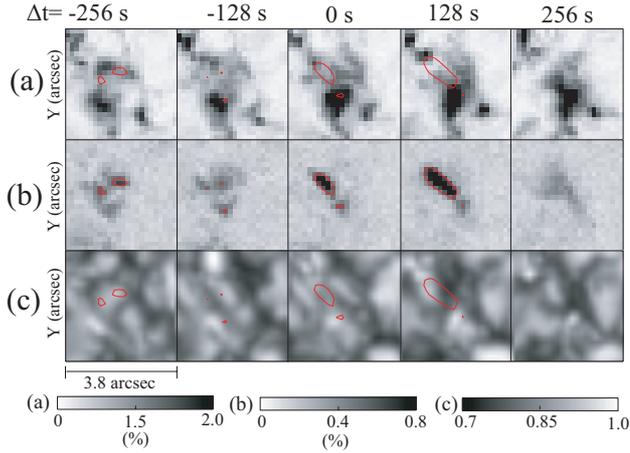


Fig. 3. NOAA AR 0931 located not far from disk center ($289''$ E, $83''$ S) was observed from 17:35 to 18:25 UT on December 28, 2006. The images consist of 24 consecutive slit positions ($0'.16$ width) with a total scan time of 128 s. The pixel size along the slit is $0'.16$. **a)** nCP (vertical magnetic field), **b)** LP (horizontal magnetic field), **c)** I_c . The region where LP is larger than 0.5% is enclosed by red lines.

Kubo et al. 2003, and references therein). This difference can be seen in an observation made with SOT which we describe next.

3. Comparison with buoyancy-driven emerging flux

We also present an example of an emerging flux event in a remnant active region located around the disk center. The properties of the event are considerably different from the previous example. We find only one such event in the observing area of $3'.8 \times 82''$ during 50 minutes' observation. The episode starts with the appearance of horizontal magnetic fields, that appear to cause a considerable deformation of the surrounding granules (Fig. 3). The maximum size of the horizontal structure is $\sim 1'.6$ along its major axis. At $\Delta t = -128$ s, prior to the emergence of magnetic flux, the granules become elongated with their alignment parallel to the orientation of the emerging fields. At $\Delta t = 0$ s, strong horizontal magnetic field structure appears in the region between the elongated granules and existing vertical fields, and vertical magnetic components are detected at both ends of the horizontal magnetic structure. Note that we employ unsigned circular polarization, $nCP = \int |V| d\lambda / \int I_c d\lambda$, instead of CP because highly deformed Stokes V profiles, some of which have three lobes, are seen in the region. The integration is carried out with the same wavelength range as LP . The horizontal flux emerges with separating vertical components at both ends. At $\Delta t = 128$ s, the horizontal magnetic structure becomes longer and stronger, and the vertical magnetic fields also strengthen. The bright point appears where the vertical field is strong. In the next frame ($\Delta t = 256$ s), the horizontal magnetic structure disappears (most likely rising above the line formation height), and we are left with vertical bipolar components. The series of observations indicate that the emerging magnetic field structure is stiff against the granular motion since it affects the surrounding granules even before its emergence into the photosphere.

The Stokes profiles along the horizontal magnetic structure at $\Delta t = 128$ s are shown in Fig. 4. The Stokes Q and U profiles show that the direction of the magnetic vector is roughly parallel to the major axis of the flux tube. The Stokes U profiles representing the horizontal magnetic structure are blue-shifted by

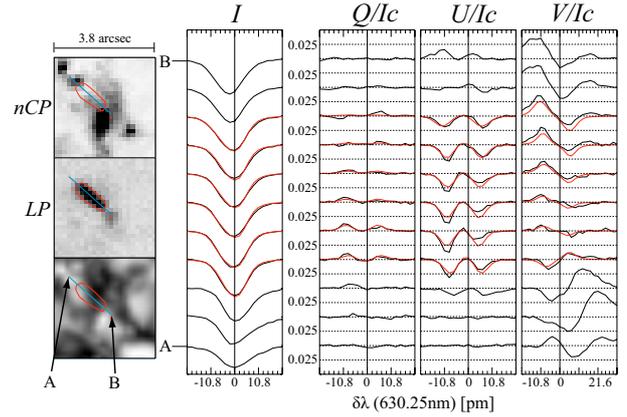


Fig. 4. The frame at $\Delta t = 128$ s in Fig. 3 enlarged on the left. The caption is the same as Fig. 2.

$1\text{--}2 \text{ km s}^{-1}$ at the top of the loop with respect to the line center of the Stokes I profiles, while the Stokes V profile is strongly red-shifted at one end of the horizontal magnetic structure with a velocity up to 5 km s^{-1} , which is below but close to the local sound speed ($6\text{--}7 \text{ km s}^{-1}$). The strong downward flow may be related to the convective collapse (Bellot Rubio et al. 2001). The expansion velocity is estimated to be about 2 km s^{-1} from the change of the length of the horizontal magnetic structure. This velocity is comparable to the upward velocity at the top of the loop. The high-velocity downflows at one end of the flux tube continuing over 4 min show that plasma descends in the flux tube in association with the rising flux tube. At the other footpoint, a few km s^{-1} upflow is detected mainly at $\Delta t = 128$ s. This may be a reverse flow due to the local enhancement of gas pressure triggered by a downflow near the footpoint. Another possibility is siphon flow driven by the pressure gradient of the two foot points (Cargill & Priest 1980). Note that even with the upflow at one end of the footpoints, the higher downflow dominates, and reduces the net mass in the flux tube. The magnetic buoyancy will be sustained as a result. This property is closer to that of emerging flux by the Parker instability than the previous one.

Given the significant difference in the two episodes presented here, we hereafter call this episode type 1, and the previous episode type 2. We do not find any events of type 1 in the data set where we find many type 2 events. The occurrence rate of type 1 is much lower than that of type 2. The three components of the magnetic field are estimated by a Milne-Eddington inversion for the Stokes spectra shown in the Figs. 2 and 4. (The black lines indicate the observed Stokes spectra, while the red lines the fitted spectra.) The average horizontal magnetic field strengths are 560 Gauss (type 1) and 580 Gauss (type 2). The field lines are roughly aligned along the major axis of the strong horizontal magnetic field region. For type 2, we have chosen the event having the strongest LP , so that we were able to perform a reliable Stokes inversion. Other type 2 events generally have smaller LP .

A remarkable difference between the types 1 and 2 is in the average filling factors, which are 0.44 and 0.17, respectively. The filling factor is an areal fraction of the magnetic atmosphere within a resolution element. We estimate the maximum diameter of the flux tube under the assumption of a single horizontal flux tube partially occupying a pixel. The width is equal to the size of the telescope's point spread function ($0'.32$) multiplied by the filling factor. The size of the type 1 horizontal flux tube thus derived is less than $\sim 100 \text{ km}$, while that for the type 2 is less than $\sim 40 \text{ km}$. They are much smaller than the scale height of the line

forming layer, indicating that actual diameter of the flux tube is not as small as this. Nevertheless, it is true that the type 2 tube is much thinner than the type 1 tube.

4. Discussion

The occurrence of isolated small-scale, transient horizontal magnetic fields in the quiet Sun was discovered using spectropolarimetric data from the Advanced Stokes Polarimeter at a spatial resolution of 1'' (Lites et al. 1996). These horizontal fields are interpreted to be convection-driven or buoyancy-driven emerging bipolar fields. Very recently, the high time variability and ubiquity of horizontal fields in the quiet Sun were clearly shown in data from the SOLIS and GONG instruments (Harvey et al. 2007). Subsequently, Hinode data has permitted us to directly see the ubiquitous presence of horizontal magnetic fields in the quiet Sun at unprecedented small scales (Lites et al. 2008; Centeno et al. 2007; Orozco Suárez et al. 2007). It is as yet unclear if the granular-size emergence phenomenon in the plage region reported herein pertains to flux elements having a stronger degree of polarization than much of that reported for the quiet Sun. It is important to understand whether the plage and quiet Sun horizontal fields have a similar origin.

The equi-partition field strength B_e is the field strength where magnetic energy is equal to kinetic energy: $\frac{B_e^2}{8\pi} = \frac{1}{2}\rho v^2$. The typical equipartition field strength B_e is about 400 Gauss for granules with a velocity of $v = 2 \times 10^5 \text{ cm s}^{-1}$, and the density ρ is $3 \times 10^{-7} \text{ g cm}^{-2}$ at the solar surface. The magnetic field strengths of both type 1 and type 2 are apparently stronger than the equi-partition field. The magnetic field strength of type 1 higher than the equi-partition field strength is consistent with the observed robustness against granular motion (Cheung et al. 2007). The type 1 event is interpreted as essentially the instability-driven emergence of flux tubes. As strong a field for type 2 cannot be readily explained. It appears that the horizontal flux tubes of the type 2 event are carried by the upflow of convection to the photosphere, and then the horizontal flow of the convection transports them to the inter-granular lane. The type 2 event is apparently subject to granular motion.

The buoyant force is proportional to $(\rho_e - \rho_i)\pi g a^2$, where ρ_e and ρ_i are the plasma density inside and outside the horizontal flux tubes, and a the radius of the flux tube. The drag force is given by $0.5C_d\rho_e u^2 a$, where C_d is an aerodynamic coefficient of order unity, and u the turbulent velocity. The magnetic tension force is $\pi a^2 B^2 / 4\pi L$, where L is the curvature radius of horizontal fields. We substitute a granular size of 10^3 km for L to obtain

the upper limit of magnetic tension. The three forces become comparable for a of 50 km and B of 600 G. The type 2 tube diameter of 40 km means that the effect of buoyancy is relatively smaller than the drag force. This implies that a thinner flux tube is more strongly influenced by turbulent motion.

Convective collapse (Parker 1978) would not work due to its horizontal nature throughout the evolution for the type 2 event. Indeed, Abbett (2007) shows that a convective surface dynamo can generate horizontally directed magnetic structures. The transient horizontal fields revealed by Hinode in our study suggest the possibility of such a local dynamo process (Cattaneo 1999; Vögler & Schüssler 2007). An alternative physical picture is presented by the extended fields caused by an exploding magnetic structure (Moreno-Insertis et al. 1995) and/or by horizontal magnetic fields that failed to emerge (Magara 2001).

Acknowledgements. The authors thank B. C. Low for his valuable comments and encouragement. Hinode is a Japanese mission developed and launched by ISAS/JAXA, with NAOJ as a domestic partner and NASA and STFC (UK) as international partners. It is operated by these agencies in co-operation with ESA and NSC (Norway). The National Center for Atmospheric Research is sponsored by the National Science Foundation.

References

- Abbett, W. P. 2007, ApJ, 665, 1469
 Bellot Rubio, L. R., Rodríguez Hidalgo, I., Collados, M., Khomenko, E., & Ruiz Cobo, B. 2001, ApJ, 560, 1010
 Cargill, P. J., & Priest, E. R. 1980, Sol. Phys., 65, 251
 Cattaneo, F. 1999, ApJ, 515, L39
 Centeno, R., Socas-Navarro, H., Lites, B., et al. 2007, ApJ, 666, L137
 Cheung, M. C. M., Schüssler, M., & Moreno-Insertis, F. 2007, A&A, 467, 703
 Harvey, J. W., Branston, D., Henney, C. J., & Keller, C. U. 2007, ApJ, 659, L177
 Kosugi, T., Matsuzaki, K., Sakao, T., et al. 2007, Sol. Phys., 243, 3
 Kubo, M., Shimizu, T., & Lites, B. W. 2003, ApJ, 595, 465
 Lites, B. W., Leka, K. D., Skumanich, A., Martínez Pillet, V., & Shimizu, T. 1996, ApJ, 460, 1019
 Lites, B. W., Skumanich, A., & Martínez Pillet, V. 1998, A&A, 333, 1053
 Lites, B. W., Kubo, M., Socas-Navarro, H., et al. 2008, ApJ, 672, 1237
 Magara, T. 2001, ApJ, 549, 608
 Moreno-Insertis, F., Caligari, P., & Schuessler, M. 1995, ApJ, 452, 894
 Orozco Suárez, D., Bellot Rubio, L. R., del Toro Iniesta, J. C., et al. 2007, ApJ, 670, L61
 Parker, E. N. 1955, ApJ, 121, 491
 Parker, E. N. 1978, ApJ, 221, 368
 Shibata, K., Tajima, T., Steinolfson, R. S., & Matsumoto, R. 1989, ApJ, 345, 584
 Strous, L. H., Scharmer, G., Tarbell, T. D., Title, A. M., & Zwaan, C. 1996, A&A, 306, 947
 Tsuneta, S., Suematsu, Y., Ichimoto, K., et al. 2007, Sol. Phys., submitted
 Vögler, A., & Schüssler, M. 2007, A&A, 465, L43