The XMM-Newton view of IRAS 09104+4109: evidence for a changing-look Type 2 quasar?

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ABSTRACT

Aims. We report on a 14 ks XMM-Newton observation of the hyperluminous infrared galaxy IRAS 09104+4109, which harbors a type 2 quasar in its nucleus. Our analysis was aimed at studying the properties of the absorbing matter and the Fe K complex at 6–7 keV in this source.

Methods. We analyzed the spectroscopic data from the PN and the MOS cameras in the 0.4–10 keV band. We also used an archival BeppoSAX 1–50 keV observation of IRAS 09104+4109 to investigate possible variations of the quasar emission.

Results. The X-ray emission in the EPIC band is dominated by the intra-cluster medium thermal emission. We found that the quasar contributes ~35% of the total flux in the 2–10 keV band. Both a transmission- (through a Compton-thin absorber with a Compton optical depth of $\tau_c \sim 0.3$, i.e. $N_H \sim 5 \times 10^{23}$ cm$^{-2}$) and a reflection-dominated ($\tau_c > 1$) model provide an excellent fit to the quasar continuum emission. However, the value measured for the EW of Fe Kα emission line is only marginally consistent with the presence of a Compton-thick absorber in a reflection-dominated scenario, which had been suggested by a previous, marginal (i.e. 2.5σ) detection with the hard X-ray (15–50 keV) imaging BeppoSAX/PDS instrument. Moreover, the value of luminosity in the 2–10 keV band measured by the transmission-dominated model is fully consistent with that expected on the basis of the bolometric luminosity of IRAS 09104+4109. From the analysis of the XMM-Newton data we therefore suggest the possibility that the absorber along the line of sight to the nucleus of IRAS 09104+4109 is Compton-thin. Alternatively, the absorber column density could have changed from Compton-thick to thin in the five years elapsed between the observations. If this is the case, then IRAS 09104+4109 is the first “changing-look” quasar ever detected.

Key words. galaxies: individual: IRAS 09104+4109 – galaxies: active – galaxies: nuclei – X-ray: galaxies

1. Introduction

IRAS 09104+4109 is one the most powerful objects in the $z \sim 0.5$ Universe. It is a hyperluminous infrared cD galaxy of IR luminosity $>10^{12} L_{\odot}$ at $z = 0.442$, residing in the core of a rich cluster of galaxies (Kleinmann et al. 1988). Its optical spectrum shows only narrow emission lines, but broad Balmer and Mg II emission lines were observed in the polarized light (Hines & Wills 1993; Tran et al. 2000). These pieces of evidence have led to the conclusion that a dust-enshrouded type 2 quasar lies in the nucleus of IRAS 09104+4109.

The X-ray observations of IRAS 09104+4109 collected so far lend support to this suggestion. Franceschini et al. (2000; F00) reported on the analysis of a BeppoSAX observation of this source. They found that the X-ray spectrum below 10 keV is dominated by the intra-cluster medium (ICM) thermal ($kT \sim 5.5$ keV) emission. The detection of a weak signal in the 15–60 keV band with the PDS instrument was interpreted by F00 as the primary emission of the buried quasar, which emerges from an absorbing screen of $N_H \geq 5 \times 10^{24}$ cm$^{-2}$. They also reported on the marginal detection of a neutral Fe Kα emission line with an equivalent width (EW) of $\sim 1–2$ keV, consistent with a reflection-dominated scenario for the quasar emission in the 2–10 keV band. This evidence makes this source the best example of a Compton-thick type 2 quasar found to date. Exploiting the high spatial resolution of Chandra, Iwasawa et al. (2001, I01 hereafter) were able to analyze the spectral data of the central source embedded in the extended ICM emission. Although limited by the low statistics (~200 counts in the range 0.6–7 keV), these data allowed to confirm the presence of a heavily obscured quasar and, by the comparison with the BeppoSAX PDS flux, I01 concluded that the Chandra spectrum is also reflection-dominated.

2. XMM–Newton observation and data reduction

IRAS 09104+4109 was observed with XMM-Newton (Jansen et al. 2001, and references therein) on April 27, 2003 for ~14 ks. The EPIC PN and MOS observations were carried out in the full frame mode using the Medium filter. XMM-Newton data were processed with SAS v6.5. We used the EPCHAIN and EMCHAIN tasks for processing the raw PN and MOS data files, respectively, to generate the relative linearized event files. X-ray events corresponding to patterns 0–12(0–4) for the MOS(PN) cameras were selected. Hot and bad pixels were removed. The event lists were furthermore filtered to ignore
Table 1. Best-fit spectral parameters of the EPIC spectrum. See Sect. 3 for details. Column: (1) spectral model; (2) temperature of the ICM “cool core” component (keV); (3) ICM metallicity; (4) 2–10 keV flux of the quasar component (10^{-15} erg cm^{-2} s^{-1}); (5) 2–10 keV luminosity of the quasar component (10^{43} erg s^{-1}); (6) column density of the absorber (10^{21} cm^{-2}); (7) energy of the Fe Kα line (keV); (8) intensity of the Fe Kα line (10^{-6} ph/cm²/s); (9) EW of the Fe Kα line (eV); (10) reduced χ^2 and number of degrees of freedom.

<table>
<thead>
<tr>
<th>Model\footnote{a}</th>
<th>$kT$</th>
<th>$Z/Z_{\odot}$</th>
<th>$F_{2-10}$</th>
<th>$L_{2-10}$</th>
<th>$N_{H}$</th>
<th>$E_{K\alpha}$</th>
<th>$I_{K\alpha}$</th>
<th>$EW_{K\alpha}$</th>
<th>$\chi^2$/d.o.f.</th>
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</thead>
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<tr>
<td>(A)</td>
<td>3.95±0.09</td>
<td>0.47±0.10</td>
<td>4.68</td>
<td>7.95</td>
<td>4.8±1.5</td>
<td>6.38±0.06</td>
<td>8.7±0.4</td>
<td>390±12</td>
<td>0.94(219)</td>
</tr>
<tr>
<td>(B)</td>
<td>3.6±0.8</td>
<td>0.48±0.11</td>
<td>5.57</td>
<td>&gt;2.05</td>
<td>–</td>
<td>6.38±0.05</td>
<td>4.0±2.1</td>
<td>402±19</td>
<td>0.96(220)</td>
</tr>
</tbody>
</table>

\footnote{a} (A) transmission model; (B) reflection model; \footnote{b} absorption-corrected line against absorption-corrected continuum; \footnote{c} with respect to the pure reflection component.

Fig. 1. a) Left: XMM-Newton PN (top) and MOS (bottom) spectra of IRAS 09104+4109 when the transmission model is applied. The lower panel shows the deviations of the observed data from the model in unit of standard deviations. b) Right: best-fit model for the transmission scenario. The different spectral components are also plotted (e.g. Sect. 3 for further details).

periods of high background flaring according to the method presented in Piconcelli et al. (2004) based on the cumulative distribution function of background lightcurve count-rates. Final net exposures of 10.8, 13.4, and 13.4 ks were obtained for PN, MOS1 and MOS2, respectively. The source photons were extracted for the PN(MOS) camera from a circular region with a radius of 37(40) arcsec, while the background counts were estimated from a larger (i.e. ~75 arcsec radius) source-free region on the same chip. Appropriate response and ancillary files for all the EPIC cameras were created using respectively RMFGEN and ARFGEN tasks in the SAS. Combined MOS1+MOS2 spectrum and response matrix were created.

Light-curves in the 0.2–2 keV and 2–15 keV band were extracted. Both light-curves are consistent with a constant flux level over the whole XMM-Newton exposure.

3. Spectral modelling

Both PN and MOS spectra were grouped to have a minimum of 20 counts per bin to allow the use of χ² minimization technique and fitted simultaneously. Given the current calibration uncertainties, we discarded events below 0.4 keV and above 10 keV. All fits were performed using the XSPEC package (v11.3) and included the Galactic column density value of $N_{H} = 1.81 \times 10^{20}$ cm$^{-2}$ (Murphy et al. 1996). Best-fit parameter values are given in the source frame, unless otherwise specified. The quoted errors on the model parameters correspond to a 90% confidence level for one interesting parameter ($\Delta \chi^2 = 2.71$; Avni 1976). A cosmology with ($\Omega_m, \Omega_{\Lambda} = (0.3, 0.7)$ and a $H_0 = 70$ km s$^{-1}$ Mpc$^{-1}$ is assumed throughout.

The high-resolution (~1 arcsec) Chandra observation of IRAS 09104+4109 presented by I01 has definitively established that the emission from the ICM dominates the X-ray spectrum below 10 keV. Furthermore IRAS 09104+4109 is a massive cooling-flow cluster (Fabian & Crawford 1995) showing a large radial temperature gradient, from $kT \approx 3.3$ keV in the bright cool core up to $kT \approx 7.8$ keV at a distance of 200 kpc, i.e. 36 arcsec (I01). For this reason, we initially fitted the combined PN+MOS spectrum with a model consisting of two thermal plasma emission components (i.e. MEKAL in XSPEC) plus an absorbed power law. We fixed the temperature of one thermal component to 7.8 keV in order not to underestimate the total ICM emission in the hard X-ray band. Since the photon index of the power law was loosely constrained ($\Gamma \approx 1.4 \pm 1$) we fixed it to 1.8, which is the average value typically observed for radio-quiet quasars (e.g. Piconcelli et al. 2005).

This model (transmission model hereafter) gave an excellent description of the EPIC data with a $\chi^2$/d.o.f. = 0.94(219) (see Fig. 1). The best-fit values of the spectral parameters are listed in Table 1. The value of $\sim 4.8 \times 10^{23}$ cm$^{-2}$ found for the column density translates in a Compton optical depth $\tau_c \sim 0.3$ of the absorbing screen. The value of temperature ($kT = 3.9_{-0.8}^{+0.5}$ keV) in the cluster core and metallicity ($Z/Z_{\odot} = 0.47_{-0.10}^{+0.25}$) of the ICM are consistent with I01.

Works based on BeppoSAX and Chandra observations favored an interpretation of the spectrum of IRAS 09104+4109 below 10 keV in terms of reflection-dominated emission. We therefore replaced the absorbed power law in the transmission model with a Compton reflection component from neutral matter (i.e. PEXRAV model in XSPEC). For this spectral component, which is due to the reprocessing of the emission from the obscured primary X-ray source, we assumed a $\Gamma = 1.8$ for the photon index of the incident power law, along with an inclination angle of $i = 50$ deg (Tran et al. 2000) and solar metallicity for the reflector. This model (reflection model hereafter) yielded an equally good fit to the XMM-Newton spectrum with a final
\[ \chi^2_{\text{d.o.f.}} = 0.96(220) \] (see Table 1 and Fig. 2). According to both spectral models, the contribution of the quasar component to the total flux in the 2–10 keV band is \( \approx 30–35\% \).

As shown in Figs. 1a and 2a, there is a prominent line-like emission feature at 4–5 keV (observer-frame) which is broader than the instrumental resolution at this energy and most likely due to a blend of lines associated with Fe K emission. This complex is partly accounted for by the strong FeXXV Kα emission line at 6.7 keV and the FeXXVI Kα emission line at 6.97 keV (which should be likely blended with a weak Fe Kβ line at 7.06 keV due to the reprocessing of the quasar continuum) from the two-temperature ICM. However, significant (at >99\% confidence level) positive residuals are still present. We modelled this excess with an unresolved Gaussian emission line at 6.38\pm0.06 keV. This energy is consistent with a range of ionization states from FeI to FeXVI (Kallman et al. 2004), as typically observed in quasars (Jimenez-Bailon et al. 2005). In the “reflection-dominated” scenario we measure an equivalent width of the Fe Kα line at \( \approx 6.4 \) keV of \( EW_{\text{Kα}} = 402^{+312}_{-193} \) eV (calculated with respect to the Compton reflection component). For the transmission model (i.e. assuming that the line and the continuum are both absorbed) we derive \( EW_{\text{Kα}} = 390^{+380}_{-212} \) eV (see Table 1).

3.1. A comparison with broadband 1–50 keV BeppoSAX data

To investigate on the possible year-timescale variability of the overall continuum spectral shape, we plot in Fig. 3a the unfolded 2003 PN and 1998 BeppoSAX MECS+PDS spectra. These data have been unfolded through the instrument response with respect to the best-fit model found by F00 (i.e. ICM thermal emission component + absorbed powerlaw + reflection + narrow Gaussian line at 6.4 keV; model sax98 hereafter). As expected, the 1–10 keV XMM-Newton and BeppoSAX spectra have similar shape being dominated by the ICM emission. The different normalizations of the two spectra can be ascribed to the different source extraction regions. In particular, a fraction of ICM emission from the outskirt region of the cluster spreads outside the PN/MOS extraction radius. Interestingly, the PDS data above 10 keV appear to lie slightly above the extrapolation of the MECS/PN continuum level. We calculated a 20–30 keV flux level of \( F_{20–30} = 6.1 \times 10^{-13} \) and \( -1.3 \times 10^{-12} \) erg cm\(^{-2}\) s\(^{-1}\) for the transmission\(^1\) and reflection model, respectively. The 20–30 keV flux of IRAS 09104+4109 measured by BeppoSAX is \( F_{20–30} = 2.55^{+1.90}_{-1.56} \times 10^{-12} \) erg cm\(^{-2}\) s\(^{-1}\), whereby the extrapolated XMM-Newton flux in this band is fainter by a factor of \( \approx 1.6–7.3 \), if the transmission model is assumed. On the other hand, the 20–30 keV flux estimated by reflection model is consistent with the BeppoSAX value.

Figure 3b shows the iso-\( \chi^2 \) contour plot of the intensity of the Fe Kα emission line at 6.4 keV (\( I_{\text{Kα}} \)) versus the intensity of the reflection continuum for the 1998 BeppoSAX and 2003 XMM-Newton observations. The XMM-Newton values were derived by the reflection model, while for the BeppoSAX data we employed the sax98 model (temperature and abundance of the ICM).

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\(^1\) In the case of this model, we also added to the model a Compton reflection component with \( R = 1 \) (where \( R \) is the solid angle in units of 2\( \pi \) subtended by the reflecting material) as commonly observed in the 10–50 keV AGN spectra (e.g. Risaliti 2002; Reeves et al. 2006).
were permitted to vary within the 90% confidence interval measured for these parameters with XMM-Newton, e.g. Table 1). These measurements taken at different epochs are consistent within 3σ errors, as expected if Compton hump and cold iron emission arise from distant material. Furthermore, the XMM-Newton measurements are also compatible with the values of $I_{Kα} = 7.2^{+5.1}_{-1.9} \times 10^{-9}$ photons cm$^{-2}$ s$^{-1}$ and $EW_{Kα} = 1.1^{+0.9}_{-0.5}$ keV reported by I01 on the basis of a 1999 Chandra observation of IRAS 09104+4109. The mean values of $EW_{Kα}$ were significantly larger, but the errors were also very large in the low S/N BeppoSAX and Chandra spectra.

4. Discussion

The XMM-Newton observation presented here has confirmed the presence of a heavily absorbed quasar in the nucleus of the hyperluminous infrared galaxy IRAS 09104+4109. The EPIC energy range is dominated by the thermal emission of the ICM and we estimated that the quasar contributes approximately $30$–$35\%$ of the total $2$–$10$ keV flux. We were able to accurately study the Fe K complex at $6$–$7$ keV in this source, modelling it with three distinct narrow emission lines from FeI–XVI K$α$ at $6$–$7$ keV. The highly ionized Fe K lines originate in the ICM, while the line from cold iron is likely due to reprocessing of the quasar continuum off the circumnuclear environment. The better quality of EPIC data has allowed to overcome the problems of limited spectral resolution and statistics of BeppoSAX data (e.g. F00), which prevented a correct estimate of the contribution from the different ions to the Fe emission complex.

The quasar emission can be equally well fitted by either a transmission or a reflection model. The former implies a Compton-thin absorber with $N_H \approx 5 \times 10^{23}$ cm$^{-2}$, while the latter suggests a scenario where the primary X-ray continuum is blocked by a Compton-thick ($\tau_{cc} \approx 1$, i.e. $N_H \approx 10^{25}$ cm$^{-2}$) obscuring screen and the emission observed by XMM-Newton is due to indirect radiation scattered into the line of sight (it is generally assumed that the absorber and the reflector are the same material, i.e. the torus, e.g. Matt et al. 1996; Molendi et al. 2003). Using the present data these models are statistically indistinguishable. However, an $EW_{Kα} \approx 400 \pm 200$ eV assumed reflecting model is significantly lower than expected for a truly “reflection-dominated” spectrum with $[Fe/H] = 0$. In fact, the most prominent feature in the $2$–$10$ keV spectrum of a highly obscured AGN, as in the case of IRAS 09104+4109 (2 $\times$ $10^24 \lesssim N_H \lesssim 10^{25}$ cm$^{-2}$, e.g. I01), is a strong Fe K$α$ emission line with an $EW_{Kα} \gtrsim 1$ keV (Levenson et al. 2002; Guainazzi et al. 2005). This solid observational evidence matches well with the $EW$ value predicted by the theoretical calculations (e.g. Leahy & Creighton 1993; Ghisellini et al. 1994; Matt et al. 1996). However, given the uncertainty in the reflection continuum flux due to the presence of the diffuse cluster emission, we also conservatively estimated the $EW_{Kα}$ by using the lowest(highest) value in the 90% confidence interval measured for the normalization of the reflection component(Fe K$α$ emission line) and vice versa. We obtained a range of $EW_{Kα}$ values spanning from $\sim 165$ to $870$ eV, which is marginally consistent with the $EW_{Kα} \approx 1$ keV expected for a reflection-dominated scenario. As the continuum flux estimated for the obscured quasar with the reflection model ($F_{2-10} = 4.68^{+1.06}_{-1.08} \times 10^{-13}$ erg cm$^{-2}$ s$^{-1}$) is similar to the value measured with the spatially resolved Chandra spectrum ($F_{2-10} \approx 3.9 \times 10^{-13}$ erg cm$^{-2}$ s$^{-1}$, e.g. I01), the marginal discrepancy in the $EW_{Kα}$ values between the two observations arises from the different best-fit line flux intensity.

On the other hand, the $EW_{Kα} \approx 390^{+380}_{-212}$ eV (note the large error bars) found in the case of the transmission model is in agreement with the expected $EW$ value of a Fe K$α$ line transmitted through an absorbing screen with $N_{HI} \sim 5 \times 10^{23}$ cm$^{-2}$ is ($\sim 200$–$400$ eV; e.g. Awaki et al. 1991; Leahy & Creighton 1993; Ghisellini et al. 1994). A scenario with a Compton-thin absorber along the line of sight to the nucleus of IRAS 09104+4109 is therefore physically plausible. Nonetheless, the flux difference by a factor of $\sim 1.6$–$7.3$ between the 2003 XMM-Newton and 1998 BeppoSAX observation (a time interval of $3.5$ yr at the source frame after the correction for the time dilution due to the $z = 0.442$) in the $20$–$30$ keV band is puzzling. In fact, given the observed power spectral densities of low black hole mass (i.e. $M_{BH} \approx 10^{7}$ $M_\odot$) Seyfert galaxies (Markowitz et al. 2003) and scaling linearly with the black hole mass (assum- ing $M_{BH} = 2.4 \times 10^{7}$ $M_\odot$ for the nuclear black hole in IRAS 09104+4109), flux variations of a factor of $\gtrsim 2$ in $3.5$ yr should be considered unlikely (e.g. Fiore et al. 1998). However, the major problem in the interpretation of this mismatch concerns the accuracy of BeppoSAX/PDS measurements. According to our analysis, the signal-to-noise ratio of these data is poor (only $\sim 2.5\sigma$ between $15$ and $50$ keV). This implies that any conclusion based only on the PDS data about the Compton-thick or Compton-thin nature of the absorber in the nucleus of IRAS 09104+4109 could be misleading. The addition of a primary continuum power-law component modified by a Compton-thick absorber (i.e. similarly to the best-fit model proposed by F00 for the BeppoSAX broadband spectrum) might easily account for the above flux discrepancy, but, as stressed above, a scenario with a Compton-thick absorber for the EPIC spectrum is at odds with the inferred quite small value (i.e. $<1$ keV) of the Fe K$α$ $EW$.

Given the large $1.5''$ FWHM PDS field of view (FOV), PDS data might be affected by the contamination of very hard X-ray sources possibly located outside the XMM-Newton and Chandra $15$ arcmin radius FOVs. If a power-law source with $\Gamma = 1.8$ is assumed, the $20$–$30$ keV flux measured for the PDS source ($F_{20-30} \approx 2.55 \times 10^{-12}$ erg cm$^{-2}$ s$^{-1}$) is translated to $F_{2-10} \approx 7.5 \times 10^{-13}$ erg cm$^{-2}$ s$^{-1}$. According to the RXTE all-sky slew survey log N $\sim$ log $S$ function in the $2$–$10$ keV band (Revnivtsev et al. 2004), at this flux level $0.025$ sources are expected in a $1.5$ deg$^2$ area, which translates into a probability of $2.4\%$. Moreover, using the NED and SIMBAD catalogs we found a probable contaminating source in the H$_2$O maser galaxy NGC 2782 ($z = 0.008$), which is located at $\sim 50$ arcmin away from IRAS 09104+4109. The Chandra spectrum of this source (with a nuclear X-ray flux of a few times $10^{-13}$ erg cm$^{-2}$ s$^{-1}$) strongly suggests the presence of a heavily ($N_H > 10^{24}$ cm$^{-2}$) absorbed AGN (Zhang et al. 2006), whereby it is likely that NGC 2782 provides a sizable contribution to the $20$–$30$ keV flux measured with the PDS. Assuming a pure reflection model, we estimated a $F_{20-30} \sim 6 \times 10^{-13}$ erg cm$^{-2}$ s$^{-1}$, which must be considered a lower limit of the $20$–$30$ keV emission from NGC 2782, because of the likely presence of the nuclear continuum emerging after transmission through the absorber.

Alternatively, IRAS 09104+4109 could represent the first example of a “changing-look” quasar (Guainazzi et al. 2002; Matt et al. 2003, and reference therein) ever detected given that the XMM-Newton data are better explained by a transmission-dominated model with $N_H \approx 5 \times 10^{23}$ cm$^{-2}$ while, if the

2 We estimated a mass of $2.4 \times 10^7$ $M_\odot$ for the black hole in the nucleus of IRAS 09104+4109 using the measurement of the MgII line width from Hines & Wills (1993), i.e. $FWHM = 10000$ km s$^{-1}$, and the formula in Willott et al. (2003).
PDS emission is entirely due to IRAS 09104+4109, the BeppoSAX data are consistent with a reflection-dominated spectral state. In fact, similar spectral transitions from a Compton-thick to a Compton-thin state (or vice-versa) have been observed, but only in Seyfert-like AGNs so far. This scenario implies a dramatic decrease (a factor of $\geq 10$) in the line-of-sight absorbing column density during a timescale of 5 years and, in turn, suggests the presence of a largely inhomogeneous obscuring circumnuclear gas (Elvis et al. 2004; Elvis, Wilkes, McDowell, et al. 1994, ApJS, 95, 1). We estimated the line-of-sight crossing-time of an obscuring cloud in Keplerian motion around the central black hole of $2.4 \times 10^{9} M_{\odot}$ to explain the possible transition from a reflection-dominated to a transmission-dominated spectrum. We assumed a scenario similar to that described for NGC 1365 in Risaliti et al. (2007) (or NGC 3227, e.g. Lamer et al. 2003) where they found that the Compton-thick obscuring material responsible of the spectral transition is located in the broad line region (BLR) and the size of the X-ray emitting region is $\lesssim 100$ $R_{\text{G}}$. We estimated a distance of the BLR $R_{\text{BLR}} \sim 0.14$ pc using the empirical relation $R_{\text{BLR}} = M_{\text{BH}} / L_{\text{X}}$ in Kaspi et al. (2000). The redshift-corrected crossing-time (Guainazzi et al. 2002) of a Keplerian cloud covering a region of size 100 $R_{\text{G}}$ around a black hole of $2.4 \times 10^{9} M_{\odot}$ is $\sim 1.7$ yr. Such a value is therefore consistent with the 3.5 yr (source frame) elapsed between the BeppoSAX and XMM-Newton observation.

Finally, we calculated a ratio $R_{\text{X,bol}} \equiv L_{2-10 \text{ keV}} / L_{\text{bol}} = 0.016$ using the 2–10 keV luminosity of $L_{2-10 \text{ keV}} = 7.95 \times 10^{44}$ erg s$^{-1}$ measured for the transmission-scenario (see Table 1), and a bolometric luminosity of $L_{\text{bol}}(\gamma = 1) = 4.7 \times 10^{46}$ erg s$^{-1}$, which is largely dominated by the obscured quasar (e.g. Hines et al. 1999; Spoon et al. 2007). We also derived the value of $R_{\text{X,bol}}$ expected for IRAS 09104+4109 using the value of $v_{\lambda}(2500 \text{ Å}) / L_{\text{bol}}$ typical for quasars reported in Elvis et al. (1994), correcting the $L_{\text{bol}}$ value by a factor of 1/3 as suggested by Fabian & Iwasawa (1999) in order not to count twice the UV emission, and the spectral index $\alpha_{\text{OX}}$ between 2500 Å and 2 keV, defined as $\alpha_{\text{OX}} = -0.584[\ln(2 \text{ keV}) / \ln(2500 \text{ Å})]$ (e.g. Tananbaum et al. 1979). In particular, we used the relation $\alpha_{\text{OX}} = 0.137 \log(l_{\lambda}(2500 \text{ Å})) - 2.638$ reported by Steffen et al. (2006). We converted from the monochromatic value of $v_{\lambda}(2 \text{ keV})$ to the $L_{2-10 \text{ keV}}$ value by multiplying by a factor of 1.61, applying a photon index of $\Gamma = 2$. We obtained that $R_{\text{X,bol}} \equiv L_{2-10 \text{ keV}} / L_{\text{bol}} = 0.043 \times L_{\text{bol,45}}$ (with $L_{\text{bol,45}} = L_{\text{bol}}(10^{45}$ erg s$^{-1}$)). This implies that the expected value of $R_{\text{X,bol}}$ for IRAS 09104+4109 is 0.011, which is close to the value of $R_{\text{X,bol}} = 0.016$ measured using the $L_{2-10 \text{ keV}}$ derived for the transmission scenario. Since the observed luminosity of $L_{2-10 \text{ keV}} = 2.05 \times 10^{44}$ erg s$^{-1}$ in the reflection model should be just few percent of the intrinsic one (Comastri 2004; I01), this result lends further support to the hypothesis of a Compton-thin absorber along the line of sight to the nucleus of IRAS 09104+4109.

5. Summary

The XMM-Newton observation of IRAS 09104+4109 suggests the possibility that the absorber along the line of sight to the nucleus of IRAS 09104+4109 is Compton-thin. If this is the case, it implies a scenario completely different from that reported for this type 2 quasar so far. It is worth stressing, however, that previous X-ray studies of IRAS 09104+4109 inferred a reflection-dominated nature of its 2–10 keV spectrum mainly on the basis of poor signal-to-noise $\sim 15–50$ keV data taken with the non-imaging BeppoSAX/PDS detector. Future imaging spectroscopy of IRAS 09104+4109 performed in the 10–50 keV range, say with Simbol-X or IXEUS, is therefore needed to make definitive progress in measuring the exact continuum absorption from the quasar and constraining the column density of the nuclear absorber. A deep Chandra observation of IRAS 09104+4109 would also be useful to accurately quantify the strength of the Fe Kz emission line at 6.4 keV, which is a proxy for the Compton thickness of the absorber.

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BeppoSAX

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