

Eight more ultra luminous X-ray candidates unmasked

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ABSTRACT

Aims. We aim to identify the nature of several objects previously catalogued as ultra-luminous X-ray (ULX) source candidates.

Methods. Long-slit spectroscopy of the optical counterparts has been carried out.

Results. The eight objects are at higher redshift than their putative parent galaxies.

Conclusions. It is found that these eight previously catalogued ULX candidates, most of them projected in galactic halos are actually background unrelated objects. This result confirms the high degree of contamination present in current catalogues of ULXs.

Key words. galaxies: quasars: general – X-rays: general

1. Introduction

Ultra-luminous X-ray sources (ULXs) are point-like objects that are distributed around nearby (parent) galaxies; from the redshift of their putative parent galaxies and their X-ray flux, it is possible to estimate their X-ray luminosity (10^{39} – 10^{41} erg s⁻¹). The luminosities of ULXs fill a gap between those of accreting compact stars $\leq 10^{38}$ erg s⁻¹, and those of active galactic nuclei (AGN; 10^{43} – 10^{44} erg s⁻¹). ULXs were discovered by the EINSTEIN satellite and were extensively found by subsequent X-ray missions such as ROSAT, Chandra and XMM. ULXs are a relatively rare phenomenon; Ptak & Colbert (2004), analyzing ROSAT data, estimated that only ~12% of galaxies host at least one of such objects. Major efforts from the theoretical and observational sides have been dedicated to elucidating the still obscure physical nature of such sources. X-ray variability over short time scales found in some objects indicates the compact nature of these sources (e.g., Strohmayer et al. 2007). It has been proposed that ULXs could be due to super-Eddington accretion, supermassive X-ray binaries, or beamed emission (King et al. 2001; Körding et al. 2002). Alternatively, ULXs could represent the manifestation of accreting black holes of intermediate mass (IMBHs) located in the disks or halos of nearby galaxies. Although there is strong evidence that relates ULXs with star-forming activity (e.g., Gao et al. 2003), some inactive galaxies seem to host also many ULXs (Watson et al. 2005). Some ULXs are associated with recent supernovae (Roberts et al. 2003). Observational progress in this field (and subsequent identification of the most promising IMBH candidates) requires more observations in the X-ray region (variability studies, subarcsecond images, and spectra), and identification and characterization of possible counterparts in other spectral domains (e.g., Abolmasov et al. 2006; Ramsey et al. 2006).

There are several major catalogues of ULX candidates by Colbert & Ptak (2002), Swartz et al. (2004), Liu & Mirabel (2005) and Liu & Bregman (2005), which constitute a useful database for conducting statistical studies or analyzing individual objects. One problem that has been addressed by many authors (e.g., Ptak & Colbert 2004; Irwin et al. 2004; Swartz 2006; López-Corredoira & Gutiérrez 2006) is the high degree of contamination by foreground (stars in our own galaxy) and

mainly background (AGN) objects existing in these compilations. In fact, this contamination can be on average as high as 50% (López-Corredoira & Gutiérrez 2006). Similar conclusions have been reached by Ptak & Colbert (2004) and Liu et al. (2006). Analyzing the spatial distribution of ULX candidates, the contamination seems to be different according to the type of the putative host galaxy. In fact, Irwin et al. (2004) claim that elliptical galaxies do not host any ULXs with X-ray luminosities $\geq 2 \times 10^{39}$ erg s⁻¹. So, until direct confirmation of redshift, existing compilations of such objects must be seen only as ULX candidates. This makes the identification of optical counterparts essential. There are several reasons why these identifications are difficult: some hypothetical counterparts are intrinsically faint, angular resolution (in the case of ROSAT images) is poor, and ULXs usually reside in crowded regions, thus making identification of the real counterpart ambiguous.

In previous studies (Arp et al. 2004; Gutiérrez & López-Corredoira 2005; Gutiérrez 2006), we have identified and characterized the optical counterparts of about 13 ULX candidates. We showed that about 50% of the sources listed in the Colbert & Ptak catalogue (CP02 hereafter) have an optical counterpart ≤ 20 mag. So far, nearly all of them have turned out to be background/foreground contaminants. In general, these studies are biased towards the identification of objects with the brightest counterparts and towards objects situated within more isolated environments. ULX candidates embedded in HII regions offer probably more chances of being real ULXs. However, in these cases the optical counterpart is surrounded by the diffuse emission, thus rendering identification very difficult. In such cases the use of Chandra and HST data is mandatory (Ptak et al. 2006). Here, we extend our observational programme with the identification and characterization of eight new objects.

2. Sample selection and observations

2.1. Identification of optical counterparts

We select the objects from the catalogue of ULX candidates by CP02 and Swartz et al. (2004). The objects analyzed in this paper correspond to the optical counterparts of the sources IXO 29, 43, 54, 55, 56, 57, and 70 (we follow the same notation as that

Table 1. Optical and X-ray properties of ULXs. 1–2. main and alternative identification of the ULX candidates analyzed in this work (LM and LB stand for Liu & Mirabel 2005 and Liu & Bregman 2005, respectively); 3–4. X-ray RA (J2000) and Dec (2000) coordinates; 5. Log of the Luminosity in the 2–10 keV band (for object SW78 the quoted luminosity is in the 0.5–8 keV band) assuming that the X-ray sources were at the distance of the parent galaxy; 6–8. identification, morphological type and distance of the parent galaxy; 9. difference in arcsec between X-ray and nominal optical coordinates.

ID	Alternative ID	RA (J2000) (hh:mm:ss.s)	Dec. (J2000) (°:′:″)	L_x (erg s^{-1})	ID Gal.	Type	D (Mpc)	Δ (arcsec)
IXO 29	(LM) ULX1 NGC 1961	05:41:43.3	+69:20:46	40.5	NGC 1961	+5.0	52.40	0.9
IXO 43	(LM) NGC 4151, X1; (LB) NGC4151 X6	12:10:07.9	+39:23:12	39.3	NGC 4151	+2.0	20.30	1.0
IXO 54	(LM) NGC4438 ULX1; (LB) NGC4438 X2	12:27:57.5	+13:02:30	39.9	NGC 4438	+0.0	16.80	5.2
IXO 55	(LM) NGC4472 ULX1	12:29:13.1	+07:57:40	39.1	NGC 4472	-5.0	16.80	3.7
IXO 56	(LM) NGC4472 ULX2	12:29:22.3	+07:53:31	39.4	NGC 4472	-5.0	16.80	2.7
IXO 57	(LM) NGC4472 ULX3; (LB) NGC4472 X5	12:29:23.9	+07:54:00	39.8	NGC 4472	-5.0	16.80	2.7
IXO 70	(LM) NGC4649 ULX4	12:44:07.2	+11:35:25	39.3	NGC 4649	-5.0	14.85	3.5
SW78		12:25:17.17	+18:13:46.7	39.4	NGC 4382	SA(s)0+pec	14.40	0.1

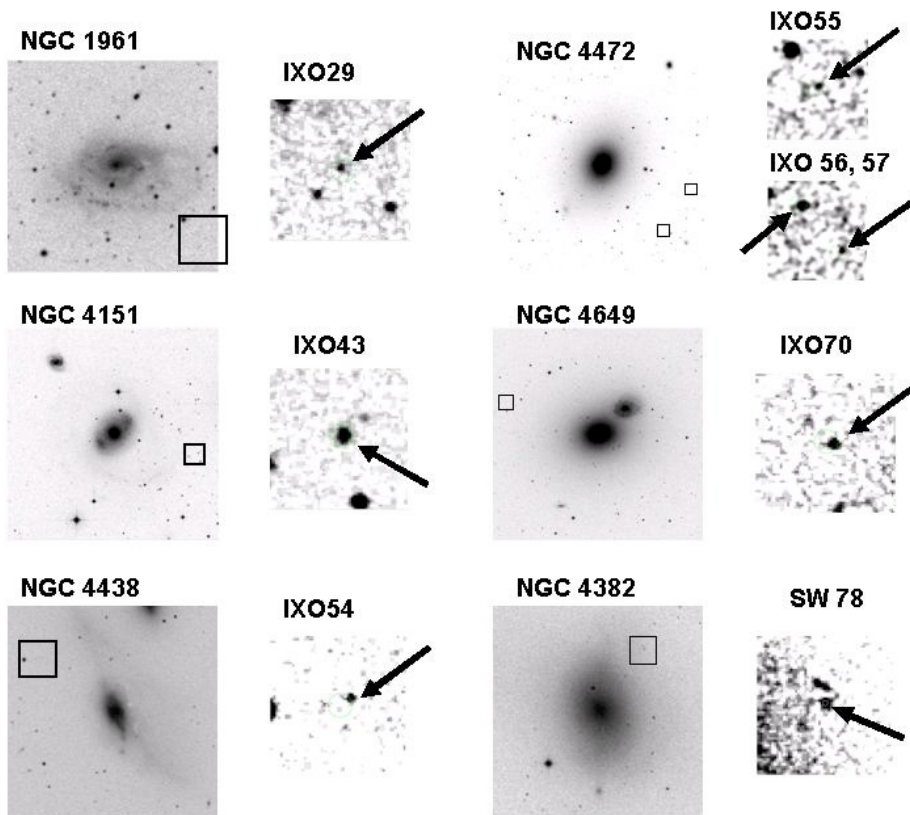


Fig. 1. DSS images centred on the putative host galaxies (*left*). The small black squares correspond to a one square arcmin centred on the ULX candidates analyzed in this paper. These regions are expanded on figures to the *right*. The arrows identify the optical counterpart of the X-ray sources. Small circles are centred on the nominal X-ray position and have radius of 5 and 1 arcsec for ROSAT and Chandra observations respectively. North is up and east to the left.

of Colbert & Ptak), and CXOU J122517.1+181346 (which we denote as SW 78 hereafter). This is the only object taken from the Swartz et al. catalogue. Table 1 summarizes the main properties of such sources and those of their candidate host galaxies. We look for possible optical counterparts to these X-ray sources in the Digital Sky Survey (DSS) plates, the USNO catalogue, and the released Sloan Digital Sky Survey (SDSS) data. For the eight cases considered here, we have identified a single optical source compatible with the X-ray positions. The last column in Table 1 lists the offset between the optical and X-ray coordinates. Figure 1 shows the DSS images with the identification of the optical counterpart.

2.2. Observations

The observations analyzed here were taken on the 1.93 m OHP¹ and on the WHT². Table 2 presents a summary of the observations. The run at OHP was in March 2005; we did long-slit

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² The William Herschel Telescope (WHT) is operated by the Isaac Newton Group and the IAC in Spain's Roque de los Muchachos Observatory.

Table 2. Observations. 1. identification of the ULXs. 2. telescope and instrument used. 3. epoch of observations. 4. exposure time.

Main ID	Telescope	Date	Exposure time (s)
IXO 29	4.2 m WHT ISIS	October 2006	3 × 1800
IXO 43	1.9 m OHP CARELEC	March 2005	2 × 1800
IXO 54	1.9 m OHP CARELEC	March 2005	3 × 1800
IXO 55	4.2 m WHT ISIS	November 2006	2 × 600 + 1 × 500
IXO 56	4.2 m WHT ISIS	November 2006	3 × 600
IXO 57	1.9 m OHP CARELEC	March 2005	1 × 1800
IXO 70	4.2 m WHT ISIS	November 2006	1 × 400
SW78	4.2 m WHT ISIS	November 2006	2 × 600

optical spectroscopy with the instrument CARELEC, whose grating provides a resolution of 133 Å/mm. The slit was 2 arcsec. This instrumental configuration provided a sampling of 1.78 Å and an effective resolution of 5 Å as measured from arc and sky lines. More details on the instrumental configuration can be found in the telescope web pages (www.obs-hp.fr/www/guide/carelec/carelec-eng.html). With WHT we used the ISIS spectrograph with the gratings R300B and R158R in the blue and red arm respectively. The observations were carried out in October and November 2006. In November the conditions were photometric with very good seeing (~0.5 arcsec) and high atmospheric transparency. The slit width was 1 arcsec. For wavelength calibration we took a few Cu–Ar and Cu–Ne lamps at the beginning and at the end of the night. The stability of the wavelength calibration during the night was checked with the main sky lines. The sampling was 0.86 Å and 1.64 Å in the blue and red arms respectively. The effective resolution was 3.8 and 8 Å respectively. The OHP and WHT spectra were analyzed following a standard procedure using IRAF³ that comprises bias subtraction, flat field correction, coaddition of exposures of the same field, extraction of the spectra and wavelength calibration. We used standard spectroscopic stars from the catalogue by Oke (1990) to correct for the response of the configuration to different wavelengths.

3. Analysis

The spectra of the optical counterparts show features which allow a clear identification and characterization. All of them, apart from IXO 43, show the presence of broad emission lines typical of AGN/QSOs. In the case of IXO 43 the spectrum shows only narrow emission lines. The position of these main features is presented in Table 3. Analysis of each spectrum is presented below.

3.1. IXO 29

This is one of the brightest sources in CP02. DSS plates show an object at RA = 05^h41^m43.4^s, Dec. = +69°20′46.7″ (J2000) with SNR in the peak ~8 in the red plates. However this object is not listed in the USNO catalogue. The spectral features identified allow us to classify the object as an AGN/QSO at redshift of $z = 0.8414$, so that it is not physically associated with the putative host galaxy NGC 1961. The real X-ray luminosity of the object is 3.3×10^{44} erg s⁻¹, which is within typical range of values for AGNs.

³ IRAF is the Image Reduction and Analysis Facility, written and supported by the IRAF programming group at the national Optical Astronomy Observatories (NOAO) in Tucson, Arizona.

Table 3. Main features in the optical spectra. 1. identification of the ULXs; 2. observed wavelength; 3. spectral line identification; 4. redshift.

ID	Spectral (Å) feature	Identification	Redshift
IXO 29	5154	MgII(λ 2799 Å)	0.8415
	6866	OII(λ 3728 Å)	0.8414
	7126	NeIII(λ 3870 Å)	0.8415
	8960?	H β	0.8426
	9222	OIII(λ 5007 Å)	0.8413
IXO 43	4617	OII (λ 3728 Å)	0.2384
	6022	H β	0.2384
	6141	OIII(λ 4959 Å)	0.2381
	6203	OIII(λ 5007 Å)	0.2385
IXO 54	4646	MgII(λ 2799 Å)	0.6601
	6184	OII(λ 3728 Å)	0.6586
	6421	NeIII(λ 3870 Å)	0.6593
IXO 55	3380	OIV(λ 1402 Å)+SiIV(λ 1397 Å)	
	3735	CIV (λ 1549 Å)	1.4109
	4603	CIII (λ 1909 Å)	1.4114
	6758	MgII (λ 2799 Å)	1.4146
IXO 56	3529	CIII(λ 1909 Å)	0.8491
	4574	OII(λ 2471 Å)	0.8511
	5178	MgII (λ 2799 Å)	0.8503
	5901	HeI(λ 3189 Å)?	0.8507
	6890	OII(λ 3728 Å)	0.8463
	7153	NeIII(λ 3870 Å)	0.8468
IXO 57	5165	MgII(λ 2799 Å)	0.8453
	5914?	HeI(λ 3189 Å)?	0.8546
	6910?	OII(λ 3728 Å)	0.8538
	7175	NeIII(λ 3870 Å)	0.8540
IXO 70	3658	CIV(λ 1549 Å)	1.3628
	3863	HeII(λ 1640 Å)	1.3548
	4493	CIII(λ 1909 Å)	1.3535
	6545?	MgII(λ 2799 Å)	1.3397
SW78	3358	CIV(λ 1549 Å)	1.1674
	4153	CIII(λ 1909 Å)	1.1760
	6126	MgII(λ 2799 Å)	1.1887

3.2. IXO 43

The X-ray source is situated close to an open arm of the Seyfert galaxy NGC 4151. There is an object at 1 arcsec from the X-ray position in the DSS plates that is listed in USNO with photographic magnitudes 18.8 and 18.1 in *b* and *r* respectively. This object appears also in SDSS classified as a red galaxy with *ugriz* magnitudes of 20.04, 19.20, 18.42, 18.06, and 17.84. This object was first considered by Arp (1997); it was named NGC4151: 14 by this author and it was found to have a redshift of $z = 0.24$. However, no spectra were presented in that work. So the spectrum presented here is the first to be published. The inferred redshift is $z = 0.2384$, a value that confirms the original estimate by Arp. The object, then, is not associated with NGC 4151. The value of $\log \text{OIII}/\text{H}\beta = 0.33$ does not allow this galaxy to be accurately classified (Veilleux & Osterbrock 1987). There are also

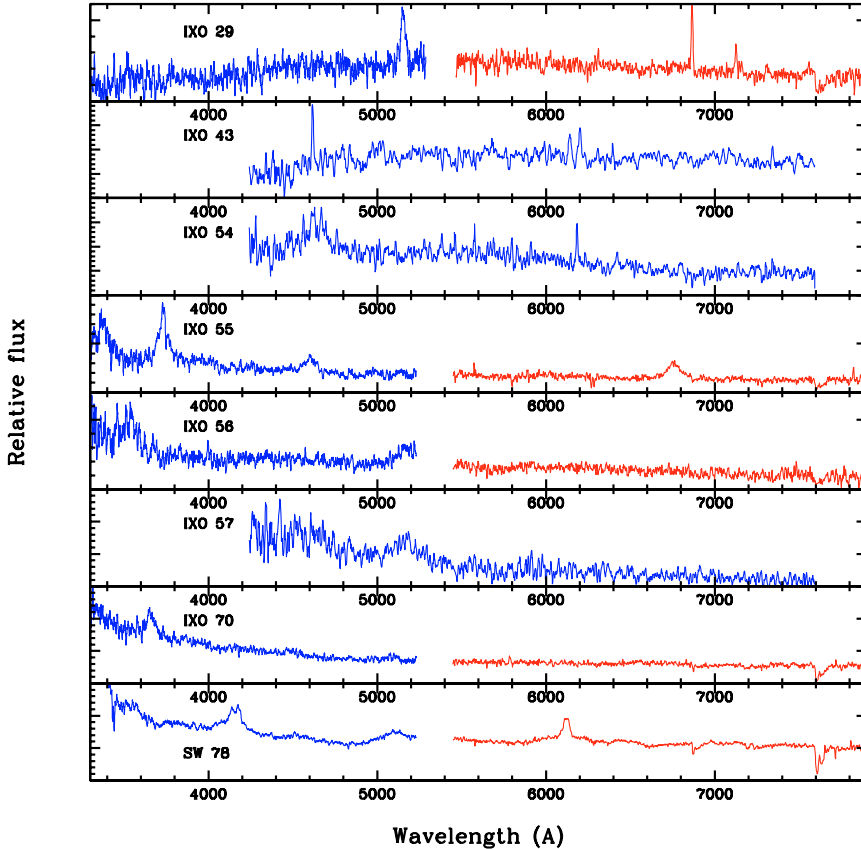


Fig. 2. Optical spectra of the counterparts of ULX sources. The y -axis is the flux in arbitrary units (counts per angstrom). Broad absorption features (obvious in most of the spectra) centered at ~ 6900 and ~ 7600 Å are telluric A and B bands respectively of molecular oxygen.

some features that seem to be associated with the stellar continuum (CaH & K, G -band, etc.) at the same redshift as the emission lines. The actual luminosity needs to be computed properly according to the real cosmological distance. Doing that, we obtained an X-ray luminosity of 6.8×10^{42} erg s $^{-1}$.

3.3. IXO 54

An optical counterpart having magnitudes of 19.2 and 17.5 in b and r respectively is listed in the USNO catalogue at ~ 5 arcsec from the nominal X-ray position. No other source listed in the USNO is compatible with the position of the X-ray source, so we tentatively identify this as the optical counterpart. The same object has been identified in the SDSS as a stellar-type object denoted by J122757.20+130232.8 having $ugriz$ 19.89, 19.17, 18.97, 18.85, and 18.84 respectively. The same object is denoted in 2MASS as 12275718+1302327 and has magnitudes of 16.762, 16.159 and 15.379 in J , H and K respectively. We identify the broad emission feature at ~ 4646 Å as MgII($\lambda 2799$ Å), and the narrow emission line at 6184 Å as OII($\lambda 3728$ Å). The apparent feature at 5577 Å corresponds to residual subtraction of a sky line. Using the OII line we determine a redshift of $z = 0.6586$. With this redshift, we identified also the emission line of NeIII($\lambda 3870$ Å). The object then turns out to be an AGN/QSO with a X-ray luminosity of 4.3×10^{44} erg s $^{-1}$ and an absolute magnitude (without K -correction) $M_r = -24.07$.

3.4. IXO 55

This ULX is in the halo of the elliptical NGC 4472. There are no sources listed in USNO in the position of this X-ray source. A

close inspection of the DSS plates reveals the presence of a faint object at $SNR \sim 3$ in the peak with position RA = $12^{\text{h}}29^{\text{m}}12.9^{\text{s}}$, Dec. = $+07^{\circ}57'41.4''$, and then at ~ 3.7 arcsec from the nominal ROSAT position. The object has also been detected by SDSS with magnitudes $ugriz$ 20.57, 20.66, 20.41, 20.36 and 20.60 respectively, and with coordinates RA = $12^{\text{h}}29^{\text{m}}12.88^{\text{s}}$ and Dec = $+07^{\circ}57'41.7''$. The spectrum is dominated by several broad emission features typical of an AGN/QSO. These are easily identified as CIV($\lambda 1549$ Å), CIII($\lambda 1909$ Å) and MgII($\lambda 2799$ Å) at 3735, 4603 and 6758 Å respectively. Another broad feature at 3380 Å is due to OIV($\lambda 1402$ Å) with possibly some contribution from SiIV ($\lambda 1397$ Å). Based on these identifications, the redshift is $z = 1.411$. This corresponds to a luminosity distance of 10 200 Mpc; the actual X-ray luminosity of the object is 4.6×10^{44} erg s $^{-1}$. We also found the typical doublet of a strong MgII absorber at 6275 and 6291 Å, giving a redshift for this absorption system of 1.2446.

3.5. IXO 56

We identify the optical counterpart as the source SDSS J122922.15+075332.5. This source is not listed in USNO. However, as in the previous case, a faint object is located in the DSS plates at ~ 2.7 arcsec from the ROSAT position. The two most obvious spectral features are two broad emission lines at 3529 and 5178 Å that correspond to CIII($\lambda 1909$ Å) and MgII($\lambda 2799$ Å) respectively. Based on those identifications, the object is an QSO/AGN with $z = 0.849$. We also found the narrow emission line OII($\lambda 2471$ Å) at 4574 Å. In the red part of the spectrum, we tentatively identify weak emission in HeI($\lambda 13189$ Å), OII($\lambda 3728$ Å) and NeIII($\lambda 3870$ Å). With a luminosity distance of 5410 Mpc, the X-ray luminosity is 2.6×10^{44} erg s $^{-1}$.

3.6. IXO 57

An optical counterpart having $r = 18.7$ mag is located at 2.7 arcsec from the nominal X-ray position in the DSS plates. No other object is compatible with the X-ray position. The optical counterpart is identified as the star J122923.73+075359.2 in the SDSS survey with magnitudes *ugriz* 19.56, 19.11, 19.04, 19.15, 19.01 respectively. During the preparation of this work⁴ the SDSS DR5 was released in which the object appeared spectroscopically measured as a QSO at $z = 0.8540$. Our spectrum is in agreement with this result. In fact, the most relevant spectral feature is a broad-band emission feature centered at ~ 5165 Å, which corresponds to MgII ($\lambda 2799$ Å) at $z = 0.845$. Based on this, we identify two possible additional narrow emission lines centered at 6910 Å and 7175 that we identify as OII($\lambda 3728$ Å) and NeIII($\lambda 3870$ Å). There is also some evidence for the presence of CII and NeIV in the blue part of the spectrum. The difference in redshift between the broad and the narrow features could be due mostly to uncertainties in the determination of the centroid of MgII; in any case, it would correspond to a difference of ~ 3000 km s⁻¹. The X-ray luminosity of the object is 6.6×10^{44} erg s⁻¹. The apparent *r*-band magnitude (19.04) measured with SDSS corresponds to $M_r = -24.66$ (without *K*-correction).

3.7. IXO 70

The object has been observed in imaging and spectroscopy by the SDSS collaboration, who labeled it as SDSS J124406.97+113524.3 and classify it as a QSO. The *ugriz* magnitudes are 19.48, 19.31, 19.13, 18.96 and 18.86 respectively. However, according to the NED database, the determination of the redshift was uncertain, with the values $z = 4.41$ and 1.34 depending on the technique and lines used. There are no sources listed in the USNO compatible with the position of the X-ray emission. In this case, this is rather surprising because the object is rather conspicuous in the DSS plates having $SNR \sim 20$ in the peak. The optical counterpart is situated 3.5 arcsec SW of the X-ray position. We took long-slit spectra of the optical counterpart with the 1.93 m OHP telescope in March 2005; however the data were very noisy, making the identification of spectral features very uncertain. We observed the object again with the WHT in November 2006 (these are the data presented here). The spectrum shows a broad emission feature at 3658 Å. Two other fainter and similar features are found at 4493 and 6545 Å. These allow us to classify the object as an AGN/QSO and to identify these lines as CIV($\lambda 1549$ Å), CIII($\lambda 1909$ Å) and MgII($\lambda 2799$ Å) respectively. From the position of the CIV line (by far the most intense) we obtain a redshift $z = 1.363$. The X-ray luminosity is then 8.6×10^{44} erg s⁻¹.

3.8. SW78

The object was catalogued by Swartz et al. (2004) and is the only object with Chandra coordinates considered in this paper. An object coincident with the X-ray coordinates is listed in the USNO catalogue as having magnitudes $b = 18.4$ and $r = 18.6$. The object is well above the noise in the DSS plates and seems to be elongated. We assume this to be the optical counterpart of the X-ray source. This is the same object

discovered by Knezek & Bregman (1998) as QSO1225+182 with a tentative redshift of $z = 1.19$. An extended object (or perhaps a couple of close objects) is located at ~ 8 arcsec from the optical counterpart. The spectrum of the optical counterpart shows three major broad emission lines that we identify as CIV($\lambda 1549$ Å), CIII($\lambda 1909$ Å) and MgII($\lambda 2799$ Å) respectively. There are two absorption lines at 3433 and 3441 Å that we identify as the signature of a strong Mg II absorber at redshift 0.2277. From the emission lines, the object turns out to be an AGN/QSO at redshift $z = 1.177$ (this is the mean value of the redshifts obtained from the three main lines). The actual X-ray luminosity is 8.2×10^{44} erg s⁻¹.

4. Discussion and conclusions

In the eight cases considered here, the object identified as the optical counterpart of the ULX candidate is the only present in DSS plates, the USNO catalogue or SDSS images compatible with the X-ray position. Liu & Bregman (2005) have determined that the positional error of Rosat HRI observations are well described by a two-dimensional distribution with $\sigma = 3.62 \pm 0.02$ arcsec. This is consistent with the differences (2.8 ± 0.6 arcsec) we find between the position of the X-ray sources and their optical counterparts (see Table 1, last column). The positional accuracy for Chandra observations is ~ 0.6 arcsec (Aldcroft et al. 2001, <http://cxc.harvard.edu/cal/ASPECT/celmon/>) consistent also with the difference between X-ray and optical positions we found for SW78. This and the fact that all of them turn out to be emission line galaxies (at least seven of them AGN/QSOs), objects which have relative low density, gives confidence in these identifications. In fact, a rough calculation demonstrates that a misidentification for any of the cases of the counterparts of Rosat sources that have been classified as QSO/AGNs (i.e. IXO 29, 54, 55, 56, 67 and 70) is unlikely. The mean density of QSO/AGNs brighter than mag 20 in random fields is ~ 17 per square degree (e.g. Croom et al. 2004; López-Corredoira & Gutiérrez 2004), and assuming a very conservative error in the X-ray nominal position of 10 arcsec, the probability to have detected at least one AGN/QSO by chance is $< 2.5 \times 10^{-3}$. This is certainly an upper limit to the real probability because the limit in magnitude and positional errors are very conservative. Due to the better resolution of Chandra observations, a misidentification in the case of SW78 is even more unlikely. The optical counterpart of IXO 43 has been identified as an emission line galaxy with $r = 18.1$ and $b = 18.8$ mag in DSS plates. Although the density of such objects is about an order of magnitude higher than the one of AGNs (e.g. Ho et al. 1997), a misidentification for this object is also very unlikely ($\sim 7 \times 10^{-4}$).

Burbidge et al. (2006) claim that the number of ULXs identified as QSOs exceeds the background expectations. However, their analysis does not take into account that the total number of galaxies surveyed (most of them without any ULX candidate) is much more larger. The density of QSOs inferred from the identification presented in this paper (and taken into account that only 12% of the galaxies explored by CP02 hosted an ULX candidate) is roughly in agreement with the density of QSOs up to $b = 19$ mag. Optical expected background densities and further discussion on possible anomalies to these densities were given in López-Corredoira & Gutiérrez (2006).

The actual X-ray luminosities of the objects identified unambiguously as AGN/QSOs are in the narrow range $2.6 \times 10^{44} - 8.6 \times 10^{44}$ erg s⁻¹, which is in the range typical for AGN/QSOs (e.g. Boyle et al. 1994). The small relative separation (~ 35 arcsec),

⁴ Preliminary estimates of the redshift for this and some other objects analyzed in this paper were presented in advance in Gutiérrez & López-Corredoira (2006) and López-Corredoira & Gutiérrez (2006).

and the similar redshifts indicate that objects IXO 56 and IXO 57 are physically related, perhaps members of a larger structure.

It is important to stress that these identifications are clearly biased, and likely they are not representative of the nature of the full sample of ULX candidates. In fact, the objects analyzed here belong to the subsample of ULX candidates that are relatively bright in the optical (thus having a higher optical to X-ray luminosity ratio). None of these belongs to the subsample of 25 objects which López-Corredoira & Gutiérrez (2006) assign higher probabilities to be true ULXs. Most of the ULXs unmasked here were found projected on the halo of early-type galaxies and thus reinforce previous claims (e.g., Swartz 2006) that ULXs are absent in this type of galaxies. Including the objects presented here, there are 40 contaminants already identified from the list of 99 objects presented in López-Corredoira & Gutiérrez (2006).

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