

# Infrared integral field spectroscopy of SN 1987A<sup>★,★★</sup>

K. Kjær<sup>1</sup>, B. Leibundgut<sup>1</sup>, C. Fransson<sup>2</sup>, P. Gröningsson<sup>2</sup>, J. Spyromilio<sup>1</sup>, and M. Kissler-Patig<sup>1</sup>

<sup>1</sup> ESO, Karl-Schwarzschild-Strasse 2, 85748 Garching, Germany  
e-mail: kekjaer@eso.org

<sup>2</sup> Dept. of Astronomy, Stockholm University, AlbaNova, 106 91 Stockholm, Sweden

Received 28 March 2007 / Accepted 31 May 2007

## ABSTRACT

**Context.** SN 1987A in the Large Magellanic Cloud is close enough for a study of the very late time evolution of a supernova and its transition to a supernova remnant. Nearly two decades after explosion we are witnessing the supernova shock wave engaging the inner circumstellar ring, which had been fluorescing since being ionised by the soft X-ray flash from shock breakout.

**Aims.** We follow the interaction of the supernova shock with the ring material. The spatially resolved information provides us with insight into the individual shock regions around the ring.

**Methods.** Near-infrared integral field spectroscopy observations with SINFONI/VLT of the SN-ring interaction is presented. SINFONI's adaptive optics supported integral field spectrograph spatially resolves the ring and the data thus we obtain a better spatial understanding of the spectrum in different regions of the object.

**Results.** With a dynamical map of the interacting ring we determine parameters for its geometry. Since most of the IR emission lines originate behind the shock front we obtain an indication of the radial velocity of the shocked material after deconvolving the geometry. The ring geometry is consistent with a circle and we also derive a new, independent measurement of the systemic ring, and presumably also supernova, velocity. We find from the spatial distributions of the flux in the different emission lines the degree of cooling in the shocked material and follows the increases observed in the radio and X-rays. Emission from the ejecta is detected only in the strongest [Fe II] lines.

**Key words.** stars: supernovae: individual: SN 1987A – stars: circumstellar matter

## 1. Introduction

With the 8 m telescopes in the Southern Hemisphere SN 1987A will remain observable throughout its complete evolution from supernova to supernova remnant. While the emission from the supernova itself has been steadily fading (e.g. Leibundgut & Suntzeff 2003) the inner circumstellar ring has remained bright. The collision of the supernova shock with the circumstellar ring has been anticipated (Luo et al. 1994; Borkowski et al. 1997) and was first indicated by radio observations (Ball et al. 2001; Manchester et al. 2002). At about the same time the first hot spots appeared just inside the circumstellar ring in 1996 showing that the supernova shock had reached what are considered to be inward protrusions of the ring (Sonneborn et al. 1998; Michael et al. 1998, 2000, 2002; Pun et al. 2002). The X-ray emission started to increase as well (Park et al. 2004, 2006; Haberl et al. 2006). The detection of coronal lines in the optical spectra from 2002 indicates that the ring material is shocked to a temperature of  $\gtrsim 2 \times 10^6$  K (Gröningsson et al. 2006), consistent with the X-ray spectra (Zhekov et al. 2005, 2006).

The physics of the interaction is highly complex because of the nature of the circumstellar medium of SN 1987A (Chevalier & Dwarkadas 1995; Lundqvist 1999).

Here we present spatially resolved spectroscopy of the inner circumstellar ring from science verification observations with SINFONI on the VLT. Integral field spectroscopy yields

both imaging and spectroscopy, which allowed us to trace the shock interaction in spatially separated emission sites around the ring. The approaching and receding sides can be separated very clearly in several emission lines and we are able to study the ring dynamics.

The optical and IR provide, together with the UV and X-rays, the most important diagnostics of the temperature, density, and ionisation of the shocked gas. Therefore, the data presented here should be seen as a complement to the high resolution optical and UV spectroscopy presented in Gröningsson et al. (2006, 2007) and Pun et al. (2002), as well as the X-rays by Zhekov et al. (2005, 2006).

## 2. Observations and calibration

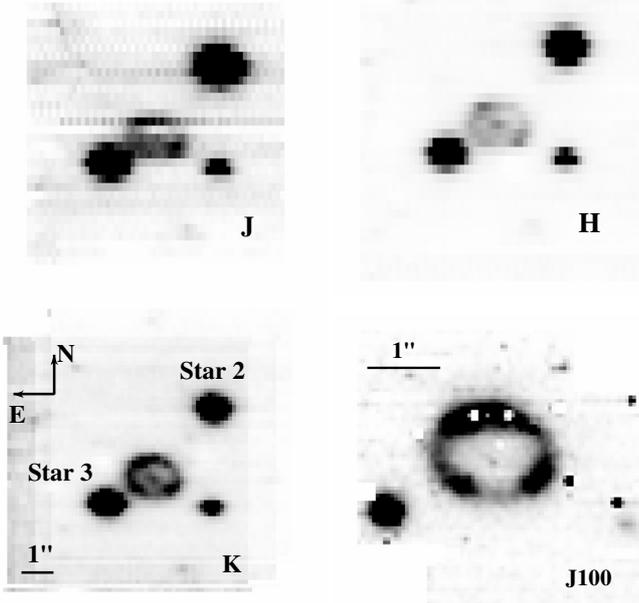
We use SINFONI Science Verification observations of SN 1987A carried out on 28 and 29 November 2004 (day 6488, preliminary results were published in Gillessen et al. 2005). Integral-field spectroscopy was obtained in *J* (1.11–1.35  $\mu\text{m}$ ), *H* (1.49–1.79  $\mu\text{m}$ ) and *K* (1.99–2.38  $\mu\text{m}$ ) bands. The observations were AO supported using Star 3 as the reference source. We have not deconvolved the images, since deconvolving extended sources is unreliable (Davies 2007). The pre-slit optics for SINFONI allows to chose different widths of the image slicing, which then leads to different fields of view on the sky. SN 1987A was observed with two different spatial resolution settings. The primary data in *J*, *H*, and *K* were obtained with a spatial resolution of 250 mas/spaxel and thus with a field of view of  $8'' \times 8''$ . In *J* we also obtained an observation with a spatial resolution of 100 mas/spaxel and a field of view of  $3'' \times 3''$  (denoted

\* Based on observations collected at the European Southern Observatory, Chile.

\*\* Tables 3 and 4 are only available in electronic form at <http://www.aanda.org>

**Table 1.** Summary of observations. The rows designate the observation date (UT), the time since the explosion of SN 1987A on 23 February 1987A, the spaxel size of the observation, the wavelength range of the spectrograph setting, the spectral resolution, the total exposure time and the image quality as measured from the combined data cube on Star 2 and Star 3. We further list the standard star used for the flux calibration, and the uncertainty in the flux as determined from a comparison of Star 2 from the corresponding black body curve, the uncertainty in the wavelength calibration as measured from night sky lines and the corresponding uncertainty in velocity.

	<i>J</i>	<i>H</i>	<i>K</i>	J100
Date	2004/11/28	2004/11/29	2004/11/29	2004/11/28
Day	6488	6489	6489	6488
Spatial Res.	250mas	250mas	250mas	100mas
Wavelength	1.11–1.35 $\mu\text{m}$	1.487–1.792 $\mu\text{m}$	1.985–2.375 $\mu\text{m}$	1.11–1.35 $\mu\text{m}$
$\lambda/\Delta\lambda$	2000	2000	4500	2000
Exptime	600s	1200s	1800s	1200s
Image quality	0.62''	0.46''	0.48''	0.32''
Std. Star	HD 36531	HD 29303	HD 29303	HD 36531
$\epsilon_{\text{Flux}}$	1 %	13 %	4 %	4 %
$\delta(\lambda)$	$8 \times 10^{-5} \mu\text{m}$	$7 \times 10^{-5} \mu\text{m}$	$7.5 \times 10^{-5} \mu\text{m}$	$8 \times 10^{-5} \mu\text{m}$
$\delta(v)$	19 km s <sup>-1</sup>	12 km s <sup>-1</sup>	11 km s <sup>-1</sup>	19 km s <sup>-1</sup>



**Fig. 1.** The  $8'' \times 8''$  FOV in the *J*, *H* and *K*-band images. Lower right is the higher resolution *J*-band  $3'' \times 3''$  FOV.

as J100 in the following). All single integrations were 600 s and Table 1 displays the total integrations per setting. When several integrations were obtained the individual exposures were offset by sub-spaxel spacings to improve the spatial resolution. The exposures were separated by sky exposures of equal integration time in an ABBA sequence.

The sizes of the spaxels for the observations are  $250 \text{ mas} \times 125 \text{ mas}$  (and  $100 \text{ mas} \times 50 \text{ mas}$  for J100). In the *K*-band images the final image quality is superior to the *J* and *H* band due to better seeing and sampling. This can be seen in Fig. 1, which shows the images (collapsed cubes) for the different bands. In the final combined data cube the spaxels are square and have the sizes  $125 \text{ mas} \times 125 \text{ mas}$  ( $50 \text{ mas} \times 50 \text{ mas}$  for the J100). The ring is fully resolved in all bands. Using the stars in the FOV of the collapsed images we measure an image quality (values listed in Table 1). The image quality measurement combines seeing, the AO correction and any smearing/dithering in combining the final cube. The image quality is in particular

important in estimating reasonable sizes of spectral extractions (apertures).

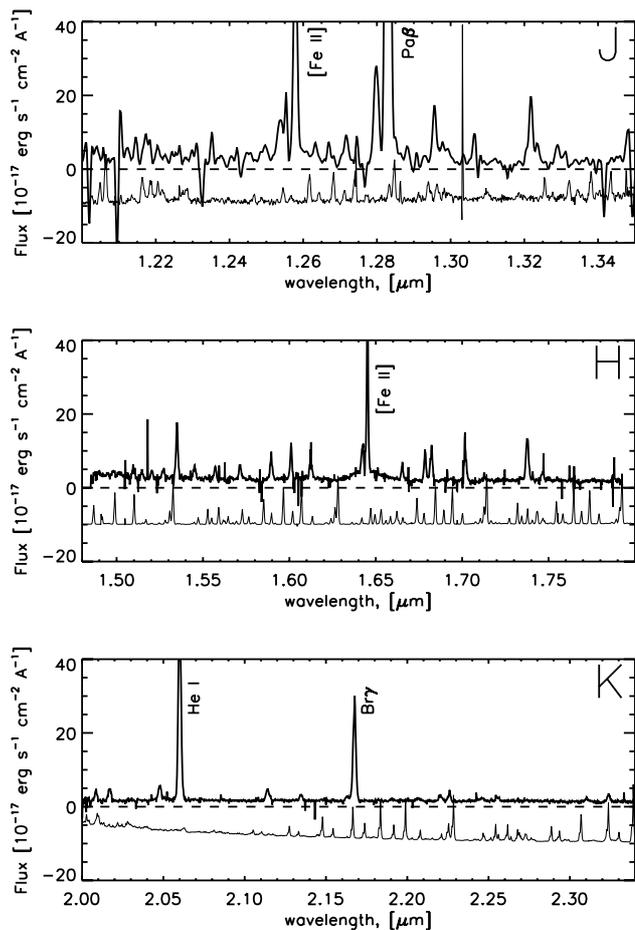
The data have been reduced and recombined into data cubes using the SINFONI pipeline (Schreiber et al. 2004; Modigliani et al. 2007). We crosschecked the wavelength solution with the atmospheric OH lines (Rousselot et al. 2000) and the resulting wavelength calibration is accurate to  $8 \times 10^{-5} \mu\text{m}$ ,  $7 \times 10^{-5} \mu\text{m}$  and  $7.5 \times 10^{-5} \mu\text{m}$ , in *J*, *H* and *K*, respectively. This corresponds to systematic velocity uncertainties of  $19 \text{ km s}^{-1}$ ,  $12 \text{ km s}^{-1}$ , and  $11 \text{ km s}^{-1}$ , as listed in Table 1.

We dereddened the spectra using the galactic extinction law from Cardelli et al. (1989), assuming  $R_V = 3.1$ , and  $E_{B-V} = 0.16$  (Fitzpatrick & Walborn 1990) for the color excess towards SN 1987A, based on  $E_{B-V} = 0.10$  from the LMC and  $E_{B-V} = 0.06$  from the Milky Way (Staveley-Smith et al. 2003). The differences between the LMC and galactic extinction law are negligible at low color excess in the near-IR.

The standard stars were observed directly before or after the supernova and have been used to remove the telluric features. We used the standard stars to construct a sensitivity curve and derive the flux calibration, based on their measured magnitudes from the 2MASS catalogue (Cohen et al. 2003). In this way the whole data cube had the same flux calibration, and we could derive calibrated spectra for Stars 2 and 3 as well. To test the absolute flux calibration we compared the spectrum of Star 2 (a B0V star NW of the SN, Walborn et al. 1993) with the black-body curve of a star with  $R = 8.3 R_\odot$  and  $T_{\text{eff}} = 22\,500 \text{ K}$  at the distance of 50 kpc. The derived flux differences are of the order of 10% and are listed as  $\epsilon_{\text{Flux}}$  in Table 1. For the J100 observations we calibrated the flux using Star 3, since Star 2 is not in this smaller FOV. As Star 3 is variable we used the observed values of Star 3 in the J250 to cross correlate the value. The total flux from Star 3 observed in J100 was 4% different from the total flux from Star 3 obtained from J250 using the same aperture in arcseconds. We combined the two *J* observations into one spectrum weighting the spectra with the exposure time. The flux error estimate for this combined spectrum is of the order of 2%. We do not think that any emission contamination from the ring will influence the result beyond the 10% we find as the accuracy of the flux.

### 3. Results

Because the data format is 3-dimensional (spatial, spatial, and wavelength) it is possible to integrate spectra over different spatial regions. Generally, it is only correct to compare line fluxes



**Fig. 2.** The integrated spectrum of the ring and the ejecta. The sky spectrum is shown offset below the SN ring spectrum. Note that the flux of the sky spectrum has been scaled down to make the spectrum fit in the same figure as the object. The line identifications are given in Table 2–4.

with respect to other line fluxes, where the number of co-added spaxels, i.e. the aperture, is the same. In order to compare with previous data we summed up the circumstellar ring and the ejecta (which corresponds to 281 spaxels co-added for the 250 mas data cubes). We extracted the spectrum in a similar way for the J100 observations and scaled this spectrum to the *J* spectrum before we combined the spaxels into a single spectrum.

We chose to scale the J100, because the aperture for the *J* band is easily compared with the ones of the *H* and *K* bands, which were observed with the same spatial scale. This spectrum should be comparable with earlier spectra of SN 1987A after a slit correction. We will refer to the spectrum extracted this way as the “integrated spectrum” and it is displayed in Fig. 2. We also show a sky spectrum offset below the “integrated spectrum”. This sky spectrum should be seen as a help to estimate where the more noisy regions are located in wavelength. The intensity of the various skylines cannot be directly compared with the integrated spectrum due to differences in aperture.

A very weak continuum flux can be discerned in the spectrum at  $(2-3) \times 10^{-17} \text{ erg}^{-1} \text{ cm}^{-2} \text{ \AA}^{-1}$ . The consistency of this continuum in all three bands makes us believe that it is real. The origin of this continuum is not clear, but it could come from large scale emission in the region. If it does indeed emerge from the ring system itself then it could be a weak H/He continuum. Most likely the continuum is a combination of several emission

origins, and it is at the moment not possible to distinguish between them.

The detected lines from the spectrum are listed in Tables 2, 3 and 4. The weakest lines we have been able to securely identify are those with a flux above  $3 \times 10^{-16} \text{ erg s}^{-1} \text{ cm}^{-2}$ . In the tables  $\lambda_{\text{obs}}$  denotes the observed wavelength,  $\lambda_{\text{air}}$  is rest wavelength in air for the line transition listed under “Identification”. The heliocentric velocity of a given line is listed in column  $v_{\text{obs}}$  and calculated as the velocity shift between the observed wavelength and the rest wavelength for the identification and then corrected for heliocentric velocity.  $\sigma_{\text{obs}}$  is the observed *FWHM* of the line and  $\sigma_{\text{deconv}}$  refers to the *FWHM* for the line corrected for the instrument resolution ( $\sigma_{\text{deconv}} = \sqrt{\sigma_{\text{obs}}^2 - \sigma_{\text{instr}}^2}$ ). The integrated flux for the line is measured by fitting a Gaussian to the line shapes and listed in the column labelled “Flux” in the tables. We chose to ignore unresolved narrow lines, which we could identify as residue sky subtraction lines using the OH line catalogue of Rousselot et al. (2000). Since we have a 3D data format we could investigate the spatial extent of each line in the integrated spectrum and check whether the line originated at a particular point in the object or whether it has a coherent appearance across the ring/ejecta. This is a very effective way to check for artefacts, such as hot/bad pixels. All observed values are integrations over the whole ring.

We have used the NIST Atomic Spectra Database version 3.1.0. (Martin et al. 2003) to identify the lines, checking the following elements and their ions: H, He, C, N, O, Na, Mg, Si, S, Ca, Fe, Co, and Ni. The ions included in the NIST database are up to ionisation XXIV for the case of iron. The other elements are complete to a similar level. The few lines which we have not yet identified did not correspond to any of the mentioned elements using the NIST database version 3.1.0, which is not necessarily complete for these elements and all their ions. The accuracies of the wavelengths in NIST are implied by the number of significant figures.

For the [Fe II] and H I lines we have verified the identification by checking line ratios using the line transition probabilities from Nussbaumer & Storey (1988) and the NIST database. The broad component in the *H*-band (Table 3) is the [Fe II]1.644  $\mu\text{m}$  line originating in the ejecta. The flux from the narrow component sitting on top of the broad component alone agrees well with the fluxes observed from the 1.257  $\mu\text{m}$  and 1.321  $\mu\text{m}$  [Fe II] lines, which come from the same upper level. However, if the broad component is the [Fe II]1.644  $\mu\text{m}$  line then we should expect to see broad components from the lines [Fe II]1.257  $\mu\text{m}$  and 1.321  $\mu\text{m}$  as well. Given the S/N in our *J* observations they are not discernible. The two lines are located in a noisy region of the *J*-band especially when compared with the *H*-band region around the 1.644  $\mu\text{m}$  line. This point can hopefully be resolved with higher signal to noise data.

When the identification transition in Tables 2–4 is preceded with a dagger ( $\dagger$ ), this transition has been observed previously and was identified by Meikle et al. (1993) or Fassia et al. (2002). Generally, the ring lines are now brighter than for those observations, as expected if the shocked component dominates the narrow component from the unshocked ring. With the higher spectral resolution of SINFONI we can now resolve some previously blended lines. However, there remain blends for lines with a true separation of less than 150  $\text{km s}^{-1}$ . We mark newly identified lines with an asterisk (\*) in the tables.

From HST spectroscopic observations (Pun et al. 2002) and our ground-based high-resolution spectroscopy (Gröningson et al. 2006, 2007), we know that the ring-ejecta system at the

**Table 2.** Emission lines of the integrated spectrum (ejecta and circumstellar ring combined corresponding to the central 281 spaxels) in the *J* band. The spectral resolution is  $150 \text{ km s}^{-1}$  in the *J* band. Errors for the values listed in the table are given either in the text or in Table 1.

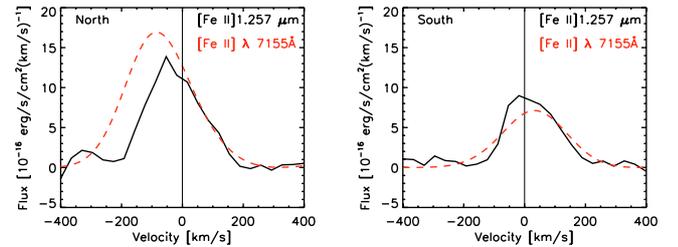
$\lambda_{\text{obs}}$ ( $\mu\text{m}$ )	$v_{\text{obs}}$ ( $\text{km s}^{-1}$ )	$\sigma_{\text{obs}}$ ( $\text{km s}^{-1}$ )	$\sigma_{\text{deconv}}$ ( $\text{km s}^{-1}$ )	Flux ( $10^{-16} \text{ erg s}^{-1} \text{ cm}^{-2}$ )	Identification	$\lambda_{\text{air}}$ ( $\mu\text{m}$ )
1.253885	408	431	404	14	*[Fe II] $a^6D_{3/2}-a^4D_{1/2}$	1.2521
1.25385	261	431	404	14	*He I $1s3s-1s4p(3S-3P)$	1.2528
Blend? The line transition probability for [Fe II] does not support this large flux. Furthermore large $\sigma_{\text{deconv}}$ .						
1.25787	277	200	132	69	†[Fe II] $a^6D_{9/2}-a^4D_{7/2}$	1.2567
1.27169	315	236	182	7	†[Fe II] $a^6D_{1/2}-a^4D_{1/2}$	1.2703
1.27983	245	305	266	43	†[Fe II] $a^6D_{3/2}-a^4D_{3/2}$	1.2788
1.27983	310	305	266	43	* He I $1s3d-1s5f(3D-3F)$	1.2785
Blend? The line transition probability for [Fe II] does not support this large flux. Furthermore large $\sigma_{\text{deconv}}$ .						
1.28300	277	281	238	367	†Pa $\beta$	1.2818
1.29562	311	190	117	18	†[Fe II] $a^6D_{5/2}-a^4D_{5/2}$	1.2943
1.30651	302	253	204	12	*N I	1.3052
1.32178	277	166	71	23	†[Fe II] $a^6D_{7/2}-a^4D_{7/2}$	1.3206
1.32901	226	339	304	10	†[Fe II] $a^6D_{3/2}-a^4D_{5/2}$	1.3278

The lines marked with † have been observed previously for this object in Meikle et al. (1993).

The lines marked with \* are new for this object or have a new identification.

moment contains three different velocity components. There is the narrow component (NC), which displays a nearly Gaussian velocity distribution, with  $\sigma \sim 10 \text{ km s}^{-1}$ , arising from the unshocked circumstellar ring. This is caused by fluorescence from the material recombining after being ionised by the supernova's UV flash at shock breakout (Fransson & Lundqvist 1989; Lundqvist & Fransson 1991). The intermediate component (IC), with  $\sigma \sim 200\text{--}300 \text{ km s}^{-1}$ , originates from the shocked ring material behind the passing shock front. The brightening has been observed in HST imaging (Michael et al. 1998, 2000; Pun et al. 2002) and arises from this material, where the ejecta collides with the inner protrusions of the ring. In our observations the NC and IC cannot be distinguished and appear as a single marginally resolved component. The SN ejecta component with an expansion velocity of  $\sim 15\,000 \text{ km s}^{-1}$  comes from the interaction with the reverse shock (Smith et al. 2005; Heng et al. 2006) not evident in our data. We also observe a broad component (BC) which will be discussed further in Sect. 3.1.

Because we do not have the spectral resolution to separate the narrow from the intermediate-velocity component (cf. Table 1) we expect some of the intermediate lines we observed to be a convolution of the narrow component with the intermediate component. In addition we have integrated spatially over the whole ring and thus the lines are also broadened by the velocity dispersion of the ring. Figure 3 shows a comparison of [Fe II] emission lines observed with UVES and SINFONI, where the UVES data has been scaled up by a factor of 100. The SINFONI observations have been convolved to the seeing of the UVES observations ( $0.5''$ ). We extracted two spectra, one North and one South on the circumstellar ring emulating the  $0.8''$  slit (PA =  $30^\circ$ ), which was the setup of the SN 1987A observations with UVES. The UVES spectra from the North and South part were convolved with the spectral resolution of SINFONI ( $\lambda/\Delta\lambda \sim 2000$  in the *J*-band) and shown in the figure as the dashed (red) curve. Taking into account that the wavelength accuracy in the *J*-band is  $\pm 19 \text{ km s}^{-1}$  then we have a good agreement for the peak velocity between the two observations. A detailed comparison of the individual elements with the high resolution optical spectrum (Gröningson et al. 2007) will further help in the separation of the different lines.



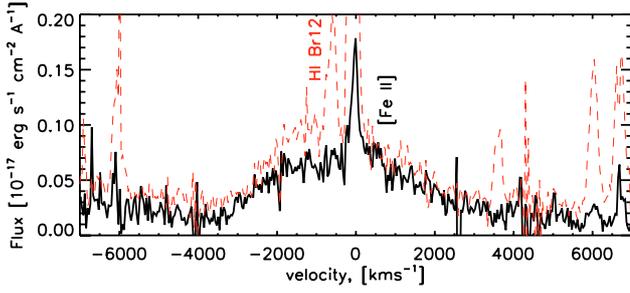
**Fig. 3.** North and South extractions from SINFONI J100 of the [Fe II]  $1.257 \mu\text{m}$ , where the SINFONI data have been spatially smoothed to the image quality of the UVES observations, and extracted along a position angle (PA =  $30^\circ$ ) of the UVES observations. The red/dashed curve are North/South extractions from UVES of the [Fe II]  $7155 \text{ \AA}$  line reduced to the spectral resolution of SINFONI and scaled up a factor 100. The vertical lines mark the systemic velocity of  $286.5 \text{ km s}^{-1}$ .

### 3.1. The broad component

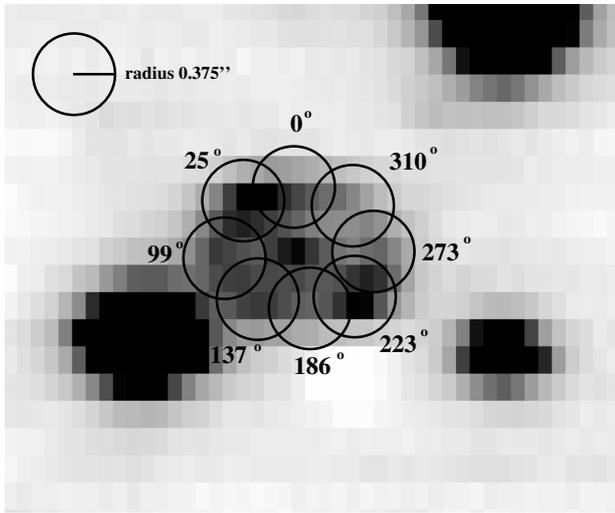
#### 3.1.1. Central emission

The only line, for which we clearly can separate the broad ejecta component from the intermediate-velocity component in the integrated spectrum is the [Fe II]  $\lambda 1.644 \mu\text{m}$  line. Figure 4 shows this line from a single spaxel centered on the ejecta in the center of the circumstellar ring (full curve), for comparison we have plotted an extraction from Spot 1 (dashed line). The emission from Spot 1 has the intermediate Br12 line which is not present in the ejecta emission, furthermore there seems to be more blue emission in Spot 1 than in the ejecta.

For the integrated spectrum when we fit a Gaussian to both the broad part and the intermediate part of this line we find a small offset in the central wavelength, so that the broad line seems less red-shifted than the intermediate line, with respect to the rest frame of the system ( $v_{\text{SN}} = 286.5 \text{ km s}^{-1}$ , Meaburn et al. 1995). The two components are individually listed in Table 3. For the broad component we derived a velocity shift of  $150 \text{ km s}^{-1}$  and a  $\sigma_{\text{deconv}} \approx 3400 \text{ km s}^{-1}$ . The bluer central wavelength for the ejecta component could indicate that the red part of the line is obscured by dust, as observed in the optical by Lucy et al. (1991).



**Fig. 4.** The [Fe II]  $\lambda 1.644 \mu\text{m}$  line from a single spaxel centered on the ejecta at the center of the circumstellar ring (aperture = 1 spaxel). The intermediate/narrow line on top is contaminating emission from faint wings of the PSF originating in the ring. The red/dashed curve is the extraction at  $25^\circ$  (aperture area  $A = \pi * (0.375'')^2$ ) scaled down with a factor of 10.

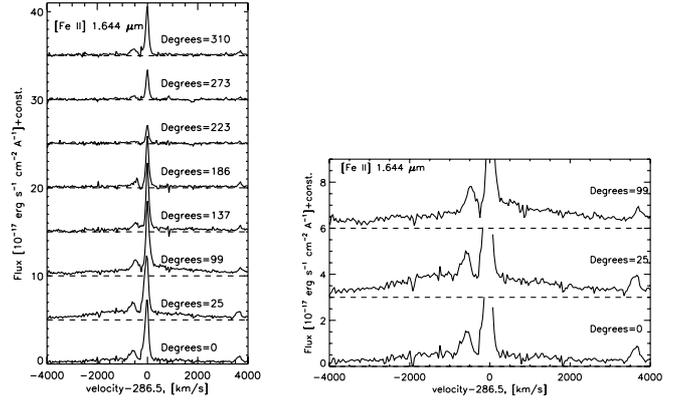


**Fig. 5.** The size of the extractions and the corresponding azimuthal positions in the  $H$  band.

This is also an indication on the faintness of the ejecta 17.8 years after the explosion. The ejecta spectrum is now dominated by energy input from  $^{44}\text{Ti}$  (Fransson & Kozma 2002; Bouchet et al. 2004, 2006).

### 3.1.2. Ring emission

We assigned an azimuthal angle for each position on the ring, defining North as zero degrees, and increasing the angle through East. Figure 5 shows the size of the extractions and their corresponding positions and azimuthal angles. Figure 6 shows the extracted spectrum of [Fe II]  $1.644 \mu\text{m}$  for different positions on the circumstellar ring and the zoom of three of the extractions. It is clear from this figure that the velocity shift of the [Fe II]  $1.644 \mu\text{m}$  line's narrow component is not changing much for different azimuthal angles (different positions on the ring). Around Spot 1 and on the eastern side in general a broad component of the line emerges ( $\sim 8000 \text{ km s}^{-1}$  FWZI). The shape of the broad component (BC) changes even on the eastern side, with a slightly more luminous blue wing at  $25^\circ$ , and a slightly more luminous red side at  $99^\circ$  (see Fig. 6). Given the accuracy of this data set we cannot definitively exclude contaminating emission from Spot 1 (from the faint wings of the point spread function (PSF) profile). Following the



**Fig. 6.** Extraction of spectra from the reduced data cube in the  $H$  band. The spectra are plotted with an offset. The faint emission on the blue side of [Fe II]  $1.644 \mu\text{m}$  is Br12.

thorough discussion of the shocks in Pun et al. (2002) we think that the broad emission in the ring emerges from the reverse shock, which would be bluer in the North part of the ring due to the inclination of the ring with respect to the line of sight, which is consistent with our observations. Again, a deeper, high signal to noise observation would solve this question.

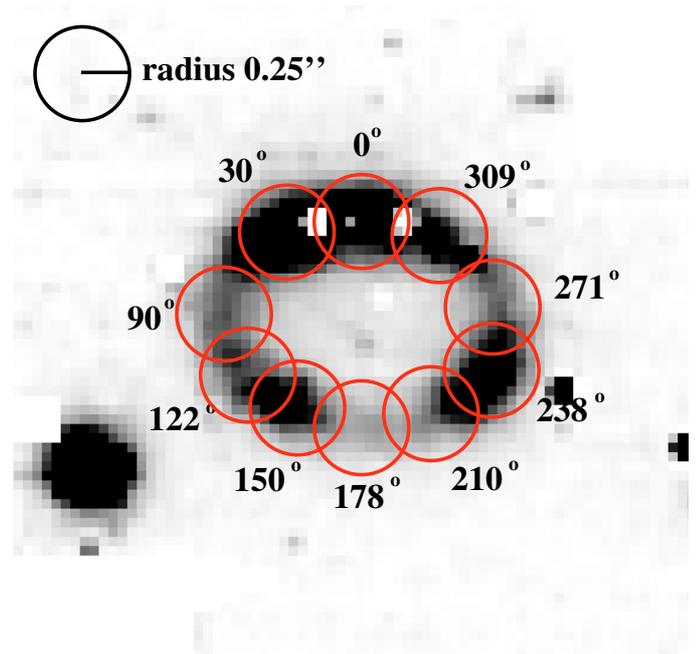
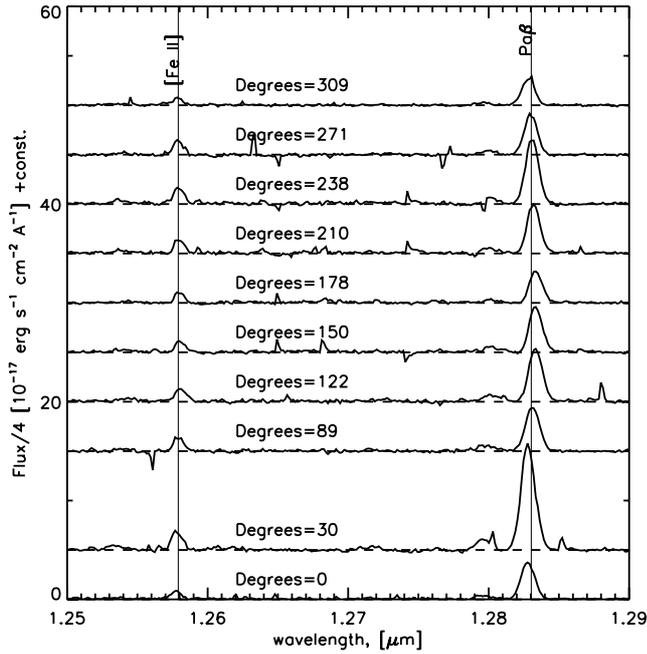
The broad line is only pronounced on the Eastern side of the ring. This East-West asymmetry has also been observed in X-rays and radio most recently in Park et al. (2006) and Manchester et al. (2002), respectively. The X-ray and radio emission is thought to emerge from the same region as the IR behind the shock front as the material cools down. We see also an East-West asymmetry with the intermediate lines, in that the Eastern side is brighter than the Western (see Fig. 10).

Whether the East-West asymmetry is due to asymmetric ejecta outflow or differences in the density of the circumstellar matter (protrusions) is still an open question, which integral field spectroscopy observations of the reverse shock would shed light on. Since the hot spots primarily first appeared on the Eastern side we could also be witnessing spot evolution. This is a question which will be addressed in an upcoming paper which presents a higher spatial resolution data set for the epoch 2005.

### 3.2. The circumstellar ring

The emission lines from the circumstellar ring are a complex combination of different emission sites. With the shock moving into the ring material an outward acceleration of the shocked material is expected, and hence a difference in the projected radial velocity shift around the ring should be observed. However, Michael et al. (2000) and Pun et al. (2002) showed that there are oblique shocks along the protrusions and the velocity field is a combination of material accelerated not only radially forward. The SINFONI data cube can be used to explore the velocity structure of the ring. We assigned an azimuthal angle for each position on the ring, defining North as zero degrees, and increasing the angle through East. The resolution in angle is then set by the spatial resolution. For each angle we coadded a spectrum with the angular position as the center and the radius of the order of the image quality. In this sampling two neighbouring extracted spectra therefore have many spaxels in common and are correlated (cf. Fig. 7).

The radius of the extraction circle was 3 spaxels for the 250 mas resolution (radius of  $0.375''$ ) and 5 spaxels for the



**Fig. 7.** *Left:* extraction of spectra from the reduced data cube close to the [Fe II] 1.257  $\mu\text{m}$  and Pa $\beta$  lines in the J100 band. The spectra are plotted with an offset. *Right:* the size of the extractions and the corresponding azimuthal positions in the J100 band. The vertical lines through the spectra mark the positions of the systemic velocity of 286.5  $\text{km s}^{-1}$  for the different lines.

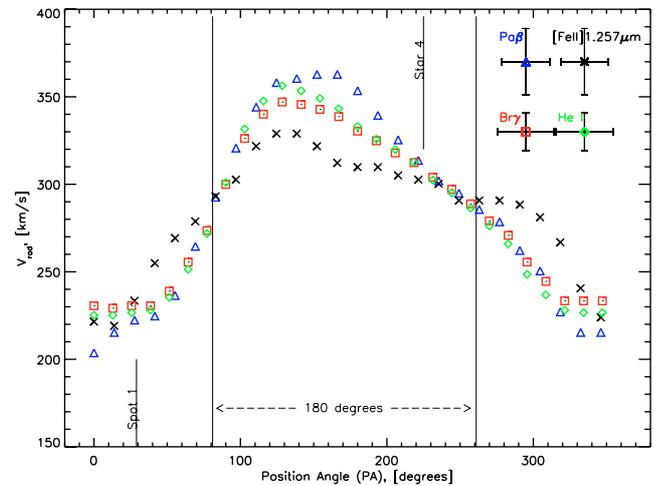
100 mas resolution (radius of 0.25''). We have here chosen the radius to be of the order of the image quality listed in Table 1, for the different observations.

The left panel of Fig. 7 shows the extracted spectrum (J100-band between 1.25  $\mu\text{m}$ –1.29  $\mu\text{m}$ ) for different positions on the circumstellar ring and the right panel shows the sizes of the extractions and their corresponding positions and azimuthal angles. It is clear that the peak of [Fe II] 1.257  $\mu\text{m}$  and the peak of Pa $\beta$  change for the different extractions.

Figure 8 shows the radial velocity shift for the Pa $\beta$ , Br $\gamma$ , He I, and [Fe II] 1.257  $\mu\text{m}$  lines, plotted vs. the position angle, PA. We have here chosen the extraction radii for the J100 extractions to be similar to that of K250 (radius = 7 spaxels  $\sim$ 0.35''). The error bars in velocity are calculated from the uncertainty in the determination of the wavelength and the error bars in angle display the size of the extraction radius for the spectra. The error bars in angle are plotted conservatively to indicate the extent of contamination between adjacent spaxels.

Only extractions with a distance of more than twice the integration radius are truly independent of each other. For *J*, *H*, and *K* the radius was 3 spaxels, which corresponds to a separation of 6 spaxels (0.75''), and 14 spaxels for J100 (0.7'').

The emission lines in Fig. 8 are all dominated by the intermediate component, coming from the shocked ring material. This is clear from the fact that the velocity of the lines differs by a large factor from the expansion velocity of the narrow lines,  $\sim$ 10  $\text{km s}^{-1}$ . It is evident from this figure that the line centers of the northern part of the ring are blue shifted, while the southern part is red shifted. This is as expected for an expanding circular ring with the Northern part tilted towards us. From HST imaging Panagia et al. (1991) find an inclination angle  $i = 42:8 \pm 2:6$  with respect to the plane of the sky (see also Sugerman et al. 2005). The theoretical line of sight velocity,  $v_r$ , for a circular ring,



**Fig. 8.** Radial velocity shift for different angular positions around the ring for the strongest emission lines. The angle increases from North to East. The location of Spot 1 and Star 4 are indicated. The J100 data has been sampled with 7 px radius in order to have a similar extraction radius as the 250 obs. The errorbars are indicated in the top right corner.

uniformly expanding with velocity  $v_{\text{exp}}$ , and tilted with respect to the plane of the sky with the angle  $i$ , is

$$v_r(\psi) = v_{\text{exp}} \sin i \cos(\psi + \phi) + v_{\text{SN}}. \quad (1)$$

Here  $\psi$  is the azimuthal angle of the ring segment, defined with  $\psi = 0$  for North. The phase shift,  $\phi$ , is introduced to account for the offset of the major axis of the ring from the East-West direction and can be interpreted as a rotation of the ring out of

the plane of the sky. The azimuthal angle,  $\psi$ , in the ring is related to the position angle, PA:

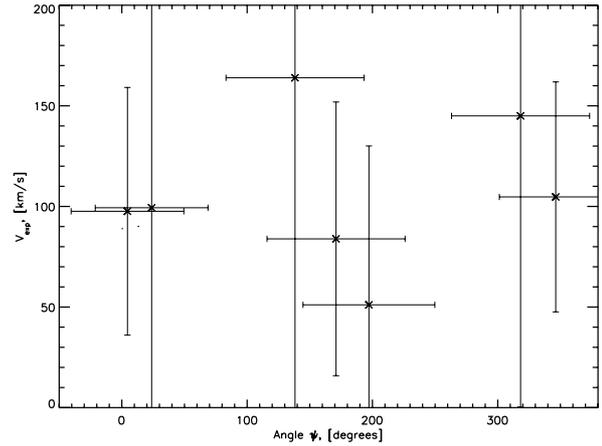
$$\tan \psi = \tan(\text{PA}) \cdot \sin i. \quad (2)$$

The observed velocity shift at a given angle is a complex convolution of the individual shock velocities at this point and the geometry. A simple model of the line profile of an individual spot was given in Pun et al. (2002). The key point of this model is that the shock velocity is likely to vary along the surface of the hot spot depending on the angle between the surface and the impacting blast wave. The highest velocities are therefore expected for a head on collision, while the tangential impacts result in lower shock velocities. Furthermore, the gas velocity, in the reference system of the shock, decreases behind the shock as the gas cools. For the observer at rest, the hot gas velocity behind the shock is  $3V_s/4$ , while that of the radiatively cooled gas is  $V_s$ . The peak velocity at a given angle therefore depends on the shock velocity which dominates the contribution to a given line. It is therefore important to realise that the peak velocity we measure is only a weighted average of the emissivity of a given line. This is then modified by the angle of the gas velocity to the line of sight.

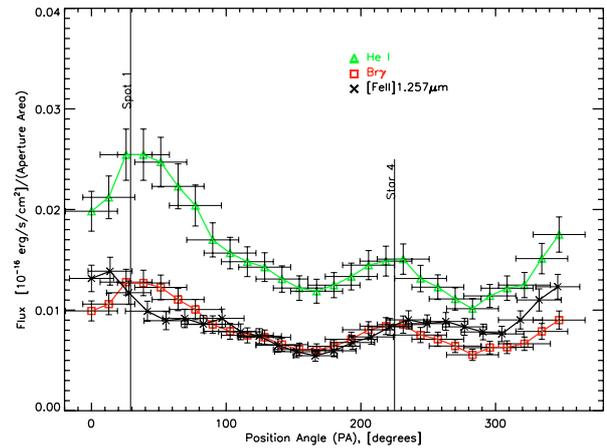
This can clearly be seen in the VLT/UVES observations, which provide line profiles with much higher spectral, but lower spatial resolution (Gröningsson et al. 2006, 2007). The line profile of e.g., the Hydrogen Balmer and He I lines from the northern part of the ring, including Spot 1, has a peak at  $\sim 192 \pm 4 \text{ km s}^{-1}$  in November 2005. This roughly agrees with that measured at the position of Spot 1 in our observations,  $\sim 223 \pm 7 \text{ km s}^{-1}$  (combining Pa $\beta$ , Bry and He I  $2.058 \mu\text{m}$ ). The H $\alpha$  line in the UVES observations, however, extends out to  $\sim 450 \text{ km s}^{-1}$ , seen only as a faint wing. While the UVES observations provide us with a good representation of the emission as function of velocity, the SINFONI observations measures the variation of the average velocity for each line along the ring. This illustrates well the complementarity of the SINFONI and UVES observations.

It is clear from Fig. 8 that the measured velocities do not follow the simple cosine as predicted by the Eq. (1).

A simplistic way to derive the bulk velocity of the material, i.e. the center of rest for the ring, would be to determine the “nodal” points ( $\psi = 90^\circ$  and  $\psi = 270^\circ$ ), where the expansion velocity is orthogonal to the line of sight, and  $\text{PA} = \psi$ . Inspection of Fig. 8 immediately shows that the two velocities are not identical at these angles for the various lines. We find that for Pa $\beta$ , Bry and He I  $v_r(90^\circ) = (300 \pm 7) \text{ km s}^{-1}$  and  $v_r(270^\circ) = (275 \pm 3) \text{ km s}^{-1}$ . For a perfect circular ring tilted out of the plane of the sky the radial velocity should vanish at points exactly  $180^\circ$  apart and these points indicate the bulk velocity of the material. We have marked the two points where this happens in Fig. 8. The mean bulk velocity is measured at a slightly different angle  $\psi + \phi = 81^\circ \pm 9^\circ$  and  $\psi + \phi = 261^\circ \pm 8^\circ$  with a bulk velocity of  $(280 \pm 7) \text{ km s}^{-1}$ . This means that the offset angle is  $\phi = -9^\circ \pm 7^\circ$ . This position of the major axis agrees well with that found by Sugerman et al. (2005) ( $\text{PA} = 81.1 \pm 0.8$ ) using imaging of light echoes and modelling the shape of the all the matter around the supernova. Our determination is not weighed by intensity and rather than having single points we make a dynamical map of the whole ring. The influence of the different hot spots or the contamination by Star 4, which is directly superposed onto the inner stellar ring, appear negligible. We have marked the location of Spot 1 and Star 4 in the diagram and there are no clear deviations from the overall behaviour at these two points. The derived velocity of the ring of  $(280 \pm 7) \text{ km s}^{-1}$  agrees with other results of  $286.5 \text{ km s}^{-1}$



**Fig. 9.** Deprojected velocity vs. angle for Pa $\beta$ , Bry and He I. 0 degrees is North, the angle increases from North to East. The curves represent the propagated uncertainties for the deprojection, and we have cut the data where the uncertainty was larger than the velocity measurement. In the deprojection we assumed a systemic velocity of  $286.5 \text{ km s}^{-1}$ .



**Fig. 10.** Flux variations for different angular positions around the ring. The aperture area is  $A = \pi * (0.375'')^2$  for the H and He lines, and  $A = \pi * (0.35'')^2$  for the [Fe II] line.

(Gröningsson et al. 2007; Meaburn et al. 1995; Crofts et al. 1995; Cumming & Meikle 1993).

Another result concerns velocity of the shocked ring material after passage of the shock. As we discussed above, the peak velocity is a function of the geometry of each shocked spot, as well as the direction of the ring plane to the line of sight. We can, however, use the geometric information above to find the deprojected, average shock velocity around the ring. Assuming an inclination angle of  $i = -42.8^\circ$ , and a systemic velocity of  $286.5 \text{ km s}^{-1}$  we have calculated the expansion velocity and binned the datapoints in Fig. 9. We derive from this a mean velocity around  $90 \text{ km s}^{-1}$ .

Figure 10 shows the flux of the [Fe II]  $\lambda 1.257 \mu\text{m}$ , the He I, and the Bry lines plotted vs. position angle. We see here that the H and He lines, as expected, follow each other, and that the ratio He/H is fairly uniform around the ring. While the narrow (intermediate) [Fe II] line flux follows the He I flux level between  $100^\circ$  and  $200^\circ$  its flux increase around Spot 1 peaks at a different angle. This difference can be understood from the fact that while the H and He lines arise in an ionised region at  $\sim 10\,000 \text{ K}$ ,

the [Fe II] lines arise at  $\sim 5000$  K, behind the ionized zone. We expect the flux of H and He to increase around the ring as the hot spots evolve. We will determine the evolution of the fluxes, as well as the velocity, around the ring in our next paper with epoch 2005 SINFONI data.

#### 4. Summary

The supernova shock interaction with the inner circumstellar ring around SN 1987A is in full swing. Our integral-field observations have revealed the extent of individual emission lines around the ring. From the high-resolution optical spectroscopy of the coronal lines from UVES (Grönningsson et al. 2006) and the soft X-ray component (Zhekov et al. 2005, 2006), we know that the shock is heating the ring material to very high temperatures. The material behind the shock is cooling rapidly, and the temporal evolution of the lines will tell us more about the conditions in these post shock areas. Our SINFONI Science Verification data only show the strongest lines and we are now investigating deeper integrations from a year later. The comparison of individual lines will be reserved for a later paper.

Our data are still consistent with a circular ring, and we have derived an average recession velocity, which agrees with other measurements. Yet, we find that there are significant differences of the velocities around the ring, with the eastern part showing larger velocities than the western segment. This corresponds well with the increase in surface brightness seen in both the radio observations (Manchester et al. 2002), Chandra observations (Park et al. 2006) and the appearance of Spot 1 in the East (Michael et al. 2002; Pun et al. 2002; Sugerman et al. 2002). By now the ring has been lit up in spots rather evenly spaced around the ring, which we also observe in the emission lines of H and He.

From the velocity shift of the lines we find an average velocity of  $\sim 100$  km s $^{-1}$  for the bulk of the H and He emitting shocked gas. We emphasise that this is just the average velocity, and considerably higher velocity shocks are present, as is directly evidenced by the UVES observations (Grönningsson et al. 2006, 2007).

The emission from the supernova itself has faded continuously. Today it is mostly the radioactive decay of  $^{44}\text{Ti}$  and reverse shock interaction, which powers the emission from the ejecta. With the integral-field spectrograph we can actually spatially isolate the ejecta emission from the ring emission. The only clear line we detect from the ejecta is [Fe II]  $1.644 \mu\text{m}$  with a width of about  $3400$  km s $^{-1}$ . This is consistent with the radioactively and reverse shock heated ejecta seen in H $\alpha$ , Ca II and Mg I (Smith et al. 2005; Grönningsson et al. 2007). Deeper observations are likely to detect more lines, as well as the broad wings from the reverse shock in the H and He lines. An advantage of the IR lines is that the [Fe II] lines are less blended than in the optical spectrum.

Monitoring the transformation of SN 1987A into a supernova remnant is a unique opportunity. It is the only such object in which we can directly observe this event while simultaneously resolving the interaction region. The structure of the circumstellar ring has been a puzzle since its discovery when it was ionised by the SN soft X-ray flash. For a complete picture we need to monitor this transition at all wavelengths. The SINFONI observations we are obtaining are unique in that they allow us to spatially investigate the emission sites around the ring and infer the local conditions of the interaction.

*Acknowledgements.* We would like to thank the Garching and Paranal Astronomers who provided support during the SINFONI Science Verification

runs. The observations could not have been carried out without the help of Sabine Mengel, Thomas Szeifert and Christophe Dumas. We also would like to thank Andrea Modigliani for valuable help in the data reduction. Special thanks go to Maria Messineo, Mariya Lyubenova, and Nina Nowak for many fruitful discussions regarding the science output from 3D spectroscopy. This work was supported by the Swedish Research Council and the Swedish National Space Board (CF, PG).

#### References

- Ball, L., Crawford, D. F., Hunstead, R. W., Klamer, I., & McIntyre, V. J. 2001, *ApJ*, 549, 599
- Borkowski, K. J., Blondin, J. M., & McCray, R. 1997, *ApJ*, 477, 281
- Bouchet, P., De Buizer, J. M., Suntzeff, N. B., et al. 2004, *ApJ*, 611, 394
- Bouchet, P., Dwek, E., Danziger, J., et al. 2006, *ApJ*, 650, 212
- Cardelli, J. A., Clayton, G. C., & Mathis, J. S. 1989, *ApJ*, 345, 245
- Chevalier, R. A., & Dwarkadas, V. V. 1995, *ApJ*, 452, L45
- Cohen, M., Wheaton, W. A., & Megeath, S. T. 2003, *AJ*, 126, 1090
- Crotts, A. P. S., Kunkel, W. E., & Heathcote, S. R. 1995, *ApJ*, 438, 724
- Cumming, R. J., & Meikle, W. P. S. 1993, *MNRAS*, 262, 689
- Davies, R. 2007, *ArXiv Astrophysics e-prints*
- Fassia, A., Meikle, W. P. S., & Spyromilio, J. 2002, *MNRAS*, 332, 296
- Fitzpatrick, E. L., & Walborn, N. R. 1990, *AJ*, 99, 1483
- Fransson, C., & Kozma, C. 2002, *New Astron. Rev.*, 46, 487
- Fransson, C., & Lundqvist, P. 1989, *ApJ*, 341, L59
- Gillessen, S., Davies, R., Kissler-Patig, M., et al. 2005, *The Messenger*, 120, 26
- Grönningsson, P., Fransson, C., Lundqvist, P., et al. 2006, *A&A*, 456, 581
- Grönningsson, P., Fransson, C., Lundqvist, P., et al. 2007, in preparation
- Haberl, F., Geppert, U., Aschenbach, B., & Hasinger, G. 2006, *A&A*, 460, 811
- Heng, K., McCray, R., Zhekov, S. A., et al. 2006, *ApJ*, 644, 959
- Leibundgut, B., & Suntzeff, N. B. 2003, in *LNP 598, Supernovae and Gamma-Ray Bursters*, ed. K. Weiler, 77
- Lucy, L. B., Danziger, I. J., Gouffes, C., & Bouchet, P. 1991, in *Supernovae. The Tenth Santa Cruz Workshop in Astronomy and Astrophysics*, held July 9–21, 1989, Lick Observatory, ed. S. E. Woosley (New York: Springer-Verlag), 82
- Lundqvist, P. 1999, *ApJ*, 511, 389
- Lundqvist, P., & Fransson, C. 1991, *ApJ*, 380, 575
- Luo, D., McCray, R., & Slavin, J. 1994, *ApJ*, 430, 264
- Manchester, R. N., Gaensler, B. M., Wheaton, V. C., et al. 2002, *PASA*, 19, 207
- Martin, W. C., Reader, J., & Wiese, W. L. 2003, *Atomic Data for X-Ray Astronomy*, 25th meeting of the IAU, Joint Discussion 17, 22 July 2003, Sydney, Australia, 17
- Meaburn, J., Bryce, M., & Holloway, A. J. 1995, *A&A*, 299, L1
- Meikle, W. P. S., Spyromilio, J., Allen, D. A., Varani, G.-F., & Cumming, R. J. 1993, *MNRAS*, 261, 535
- Michael, E., McCray, R., Pun, C. S. J., et al. 1998, *ApJ*, 509, L117
- Michael, E., McCray, R., Pun, C. S. J., et al. 2000, *ApJ*, 542, L53
- Michael, E., Zhekov, S., McCray, R., et al. 2002, *ApJ*, 574, 166
- Modigliani, A., Hummel, W., Abuter, R., et al. 2007, *ArXiv Astrophysics e-prints*
- Nussbaumer, H., & Storey, P. J. 1988, *A&A*, 193, 327
- Panagia, N., Gilmozzi, R., Macchetto, F., Adorf, H.-M., & Kirshner, R. P. 1991, *ApJ*, 380, L23
- Park, S., Zhekov, S. A., Burrows, D. N., Garmire, G. P., & McCray, R. 2004, *ApJ*, 610, 275
- Park, S., Zhekov, S. A., Burrows, D. N., et al. 2006, *ApJ*, 646, 1001
- Pun, C. S. J., Michael, E., Zhekov, S. A., et al. 2002, *ApJ*, 572, 906
- Rousselot, P., Lidman, C., Cuby, J.-G., Moreels, G., & Monnet, G. 2000, *A&A*, 354, 1134
- Schreiber, J., Thatte, N., Eisenhauer, F., et al. 2004, in *Astronomical Data Analysis Software and Systems (ADASS) XIII*, ed. F. Ochsenbein, M. G. Allen, & D. Egret, ASP Conf. Ser., 314, 380
- Smith, N., Zhekov, S. A., Heng, K., et al. 2005, *ApJ*, 635, L41
- Sonneborn, G., Pun, C. S. J., Kimble, R. A., et al. 1998, *ApJ*, 492, L139
- Staveley-Smith, L., Kim, S., Calabretta, M. R., Haynes, R. F., & Kesteven, M. J. 2003, *MNRAS*, 339, 87
- Sugerman, B. E. K., Lawrence, S. S., Crotts, A. P. S., Bouchet, P., & Heathcote, S. R. 2002, *ApJ*, 572, 209
- Sugerman, B. E. K., Crotts, A. P. S., Kunkel, W. E., Heathcote, S. R., & Lawrence, S. S. 2005, *ApJS*, 159, 60
- Walborn, N. R., Phillips, M. M., Walker, A. R., & Elias, J. H. 1993, *PASP*, 105, 1240
- Zhekov, S. A., McCray, R., Borkowski, K. J., Burrows, D. N., & Park, S. 2005, *ApJ*, 628, L127
- Zhekov, S. A., McCray, R., Borkowski, K. J., Burrows, D. N., & Park, S. 2006, *ApJ*, 645, 293

# Online Material

**Table 3.** Emission lines in the integrated spectrum (ejecta and circumstellar ring combined corresponding to 281 spaxels) in *H*. The spectral resolution is  $150 \text{ km s}^{-1}$  in the *H* band. Errors for the values listed in the table are given either in the text or in Table 1.

$\lambda_{\text{obs}}$ ( $\mu\text{m}$ )	$v_{\text{obs}}$ ( $\text{km s}^{-1}$ )	$\sigma_{\text{obs}}$ ( $\text{km s}^{-1}$ )	$\sigma_{\text{deconv}}$ ( $\text{km s}^{-1}$ )	Flux ( $10^{-16} \text{ erg s}^{-1} \text{ cm}^{-2}$ )	Identification	$\lambda_{\text{air}}$ ( $\mu\text{m}$ )
1.49825	301	320	283	2	*H I Br25	1.4967
1.50986	295	278	234	5	†He I 3s <sup>1</sup> S–4p <sup>1</sup> P <sup>0</sup>	1.5084
1.51461	253	257	209	3	* H I Br21	1.5133
1.52037	232	276	232	3	* H I Br20	1.5192
1.52705	197	334	298	4	* H I Br19	1.5261
1.53494	285	254	205	19	†[Fe II] a <sup>4</sup> F <sub>9/2</sub> –a <sup>4</sup> D <sub>5/2</sub>	1.5335
1.54512	237	272	227	5	* H I Br17	1.5439
1.55707	273	270	224	5	* H I Br16	1.5556
1.57142	257	305	266	7	†H I Br15	1.5701
1.58947	273	283	240	10	†Si I 4s <sup>1</sup> P <sup>0</sup> –4p <sup>1</sup> P	1.5880
	267				†H I Br14	1.5881
The transition probability for Br14 alone fits with the flux. Si I identification tentative.						
1.60098	281	225	168	11	†[Fe II] a <sup>4</sup> F <sub>7/2</sub> –a <sup>4</sup> D <sub>3/2</sub>	1.5995
1.61232	257	279	235	12	†H I Br13	1.6109
1.64236	298	219	160	12	†H I Br12	1.6407
1.64438	150	3393	3390	47	†[Fe II] a <sup>4</sup> F <sub>9/2</sub> –a <sup>4</sup> D <sub>7/2</sub>	1.6435
Broad component. See text for comments.						
1.64518	296	171	82	51	†[Fe II] a <sup>4</sup> F <sub>9/2</sub> –a <sup>4</sup> D <sub>7/2</sub>	1.6435
This component sits on top of the broad component (previous line).						
1.66534	282	234	180	6	†[Fe II] a <sup>4</sup> F <sub>5/2</sub> –a <sup>4</sup> D <sub>1/2</sub>	1.6638
1.67851	290	233	178	12	†[Fe II] a <sup>4</sup> F <sub>7/2</sub> –a <sup>4</sup> D <sub>5/2</sub>	1.6769
1.68227	286	268	222	15	†H I Br11	1.6807
1.70171	259	247	196	15	†He I 3p <sup>3</sup> P <sup>0</sup> –4d <sup>3</sup> D	1.7002
1.71289	307	263	216	3	†[Fe II] a <sup>4</sup> F <sub>5/2</sub> –a <sup>4</sup> D <sub>3/2</sub>	1.7111
1.73769	254	311	272	21	†H I Br10	1.7362
1.74661	286	327	291	6	†[Fe II] a <sup>4</sup> F <sub>3/2</sub> –a <sup>4</sup> D <sub>1/2</sub>	1.7449

The lines marked with † have been observed previously for this object.

The lines marked with \* are new for this object or have a new identification.

**Table 4.** Emission lines in the complete spectrum (ejecta and circumstellar ring combined corresponding to 281 spaxels) in *K*. The spectral resolution is  $67 \text{ km s}^{-1}$  in the *K* band. Errors for the values listed in the table are given either in the text or in Table 1.

$\lambda_{\text{obs}}$ ( $\mu\text{m}$ )	$V_{\text{obs}}$ ( $\text{km s}^{-1}$ )	$\sigma_{\text{obs}}$ ( $\text{km s}^{-1}$ )	$\sigma_{\text{deconv}}$ ( $\text{km s}^{-1}$ )	Flux ( $10^{-16} \text{ erg s}^{-1} \text{ cm}^{-2}$ )	Identification	$\lambda_{\text{air}}$ ( $\mu\text{m}$ )
2.00855	275	239	229	5	†[Fe II] a <sup>4</sup> P <sub>1/2</sub> –a <sup>2</sup> P <sub>1/2</sub>	2.0067
2.01696	272	327	317	8	* [Fe II] a <sup>2</sup> G <sub>9/2</sub> –a <sup>2</sup> H <sub>9/2</sub>	2.0151
2.04779	259	293	285	8	†[Fe II] a <sup>4</sup> P <sub>5/2</sub> –a <sup>2</sup> P <sub>3/2</sub>	2.0460
2.06003	275	291	283	108	†He I 2s <sup>1</sup> S–2p <sup>1</sup> P <sup>0</sup>	2.0581
2.11412	279	369	363	8	*He I 1s3p–1s4s	2.1121
2.13462	258	267	258	3	†[Fe II] a <sup>4</sup> P <sub>3/2</sub> –a <sup>2</sup> P <sub>3/2</sub>	2.1328
2.16331		416	411	4	*	
Which could be part of a broad component of Bry (not seen on the red side of Bry).						
2.16754	277	291	283	58	†H I Bry	2.1655
2.19052		575	571	3	*	
2.20659	255	381	375	2	*Na I 2p <sup>6</sup> 5p–2p <sup>6</sup> 9s	2.2047
2.20659	–58	381	375	2	or † Na I 4s <sup>2</sup> S–4p <sup>2</sup> P <sup>0</sup>	2.2070
2.22026		365	359	4	*	
2.22593	290	243	234	5	* [Fe II] a <sup>2</sup> G <sub>9/2</sub> –a <sup>2</sup> H <sub>11/2</sub>	2.2238
2.24580	286	534	530	4	†[Fe II] a <sup>4</sup> P <sub>1/2</sub> –a <sup>2</sup> P <sub>3/2</sub>	2.2436
Blend? The line transition probability for [Fe II] does not support this large flux. Furthermore large $\sigma_{\text{deconv}}$ .						
2.25525		452	447	3	*	
2.31030	304	403	397	3	*Ni II 3p <sup>6</sup> 3d <sup>8</sup> ( <sup>3</sup> F)4s–3p <sup>6</sup> 3d <sup>8</sup> ( <sup>3</sup> F)4s	2.3079
2.32381	291	284	276	5	*O I 2s <sup>2</sup> 2p <sup>3</sup> ( <sup>2</sup> D)3s–2s <sup>2</sup> 2p <sup>3</sup> ( <sup>4</sup> S)5f	2.3215

The lines marked with † have been observed previously for this object.

The lines marked with \* are new for this object or have a new identification.

The lines 2.24–2.32  $\mu\text{m}$  were observed in Meikle et al. (1993) as a blended unidentified feature.