

High-resolution X-ray spectroscopy and imaging of the nuclear outflow of the starburst galaxy NGC 253^{*}

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ABSTRACT

Aims. Using *XMM-Newton* data, we have studied the nuclear outflow of the nearby starburst galaxy NGC 253 in X-rays with respect to its morphology and to spectral variations along the outflow.

Methods. We analysed *XMM-Newton* RGS spectra, RGS brightness profiles in cross-dispersion direction, narrow-band RGS and EPIC images, and EPIC PN brightness profiles of the nuclear region and of the outflow from NGC 253.

Results. We detect a diversity of emission lines along the outflow of NGC 253. This includes the He-like ions of Si, Mg, Ne, and O and their corresponding ions in the next higher ionisation state. Additionally transitions from Fe XVII and Fe XVIII are prominent. The derived temperatures from line ratios along the outflow range from 0.21 ± 0.01 to 0.79 ± 0.06 keV, and the ratio of Fe XVII lines indicates a predominantly collisionally ionised plasma. Additionally, we see indications of either a recombining or an underionised plasma in the Fe XVII line ratio. Derived electron densities are 0.106 ± 0.018 cm⁻³ for the nuclear region and 0.025 ± 0.003 cm⁻³ for the outflow region closest to the centre. The RGS image in the O VIII line energy clearly shows the morphology of an outflow extending out to ~ 750 pc along the southeast minor axis, while the northwest part of the outflow is not seen in O VIII due to the heavy absorption by the galactic disc. This is the first time that the hot wind fluid has been detected directly. The limb brightening in *Chandra* and *XMM-Newton* EPIC observations is only seen in the energy range that contains the Fe XVII lines (550–750 eV). In all other energy ranges between 400 and 2000 eV, no clear evidence of limb brightening could be detected.

Key words. X-rays: galaxies – galaxies: individual: NGC 253 – galaxies: spiral – galaxies: starburst – ISM: jets and outflows

1. Introduction

Starburst galaxies are known to show very complex emission in X-rays. This emission originates, on the one hand, from sources that appear to be point-like sources, like X-ray binaries, supernovae, and supernova remnants. On the other hand, emission comes from the diffuse hot component of the interstellar medium, such as diffuse emission in the disc and gaseous outflows driven out of the disc by massive stellar winds and core collapse supernovae, also called superwinds. The latter phenomenon can be quite spectacular in M 82 (e.g. Stevens et al. 2003) and NGC 253 (e.g. Strickland et al. 2000), for example, where these superwinds emerge from a starburst nucleus. In NGC 253, Fabbiano & Trinchieri (1984) first detected the southeastern part of this outflow in X-rays with *Einstein* and called it the “minor-axis component”. With *ROSAT*, Pietsch et al. (2000) also detected the part of the outflow pointing in the opposite direction. However, the spatial resolution of *ROSAT* was not yet good enough to learn more about the morphology of this outflow. Later on, observations with *XMM-Newton* (Pietsch et al. 2001), and especially with *Chandra* (Strickland et al. 2000), showed that the outflow can be explained with a limb-brightened hollow

cone structure. Temperatures of the best-fit, thin thermal plasma models are in the range 0.15–0.94 keV from *XMM-Newton* EPIC and 0.46–0.66 keV from *Chandra*. Strickland et al. (2000) conclude that the detected emission originates in the shocked region at the border of the outflow where the wind collides with interstellar medium. The wind itself, though, was thought to be too hot and too thin to be detected directly. This picture, however, disregards the possibility that the wind may be mass-loaded, entraining ambient interstellar medium (ISM), as well as infalling material. If turbulent mixing proceeds on a time scale that is larger than the flow time within a given region, such as the base of the outflow studied here, we expect some clumpiness in the outflow, imprinted on an overall *less dense* wind. As we show later, this is confirmed by our analysis, which shows that e.g. O VIII is *not* limb-brightened.

High-resolution spectra of NGC 253 and M 82, taken with the *XMM-Newton* Reflection Grating Spectrometer (RGS), were first published by Pietsch et al. (2001) and Read & Stevens (2002), respectively. Both spectra show the Ly α emission lines from Si, Mg, Ne, O, N, and their helium-like charge states. Both galaxies also show emission lines from Fe XVII and Fe XVIII, and M 82 shows lines from Fe XX, Fe XXIII and Fe XXIV. In M 82 the line ratios for neon, iron, and oxygen are quite different compared to NGC 253. In general the M 82 spectrum appears to be hotter with temperatures in the range of ~ 0.3 –1.5 keV,

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Table 1. *XMM-Newton* NGC 253 observation log.

Nr.	Obs. id.	Obs. dates	Pointing direction		P.A.	T_{exp}	$T_{\text{exp, clean}}$
(1)	(2)	(3)	RA/Dec (J2000)		(deg)	(ks)	(ks)
			(4)	(5)	(6)	(7)	(8)
1	0125960101	2000-06-03	00:47:36.74	-25:17:49.2	56.9	60.8	45.1
2	0125960201	2000-06-04	00:47:36.57	-25:17:48.7	57.0	17.5	7.0
3	0152020101	2003-06-19	00:47:36.89	-25:17:57.3	53.8	113.0	75.9

with its continuum confined more to higher energies. Its X-ray flux, as well as its X-ray luminosity, in the RGS energy band (0.35–2.5 keV) is higher than for NGC 253.

However, these spectra only give a combined spectrum of the nuclear source and the outflow. In this paper we present an analysis where we decompose the total spectrum of NGC 253 into regions containing the nucleus and different parts of the outflow, while maintaining the high spectroscopic resolution.

2. Observations and data reduction

The nuclear region of NGC 253 was observed with *XMM-Newton* (Jansen et al. 2001) during two orbits in June 2000 and June 2003 using all of the European Photon Imaging Camera (EPIC) instruments (Strüder et al. 2001; Turner et al. 2001) and the two co-aligned RGS spectrometers, RGS1 and RGS2 (den Herder et al. 2001), for a total of about 190 ks. Observation identification, observing date, pointings, orientation of the satellite (P.A.), total exposure time (t_{exp}), and exposure time after screening for high background ($T_{\text{exp, clean}}$) of these observations are shown in Table 1.

An additional archival observation (Obs. id. 0110900101) could not be used for this analysis, since the pointing of this observation was in the northwest halo of NGC 253 with the result that the outflow was not in the field of view (FOV) of the RGS.

Before we start to describe the analysis procedures, we want to place emphasis on why it is even possible to perform the following spectroscopic analysis. First, the nuclear outflow of NGC 253 is an extended X-ray object, which can be spatially resolved by *XMM-Newton* EPIC and RGS, embedded in an even larger region of X-ray emission from point-like sources and diffuse emission in the disc and halo of the galaxy. This means the RGS data for the central regions are in principle affected by the contamination from the surrounding emission, but, as shown in Fig. 1, both the nucleus and the outflow in NGC 253 are significantly brighter and well above the galaxy emission, so we can only expect a minor contamination. Moreover, as we show later, we can identify and “remove” effects due to the disc emission.

Second, since the RGS is a slitless spectrometer, the spectra of all sources in the field of view are superimposed on each other on the detector. Spatial displacement of a source along the dispersion direction corresponds to a wavelength shift in the spectrum of $2.31 \times 10^{-3} \text{ \AA arcsec}^{-1}$ with respect to a not-displaced source. Since the outflow has an extent of up to $1.4'$ in the dispersion direction, the spectral resolution is limited to $\sim 0.19 \text{ \AA}$ at 15 \AA . This is still considerably better than the energy resolution from CCD detectors.

Due to the superposition of all sources, other bright point sources in the FOV could contaminate the spectrum of the outflow. However, the effective area decreases significantly for off-axis sources and even a contribution from the brightest off-axis source at the bottom of Fig. 2 (X21 from Pietsch et al. 2000; Tanaka et al. 2005) can be neglected. The source on the

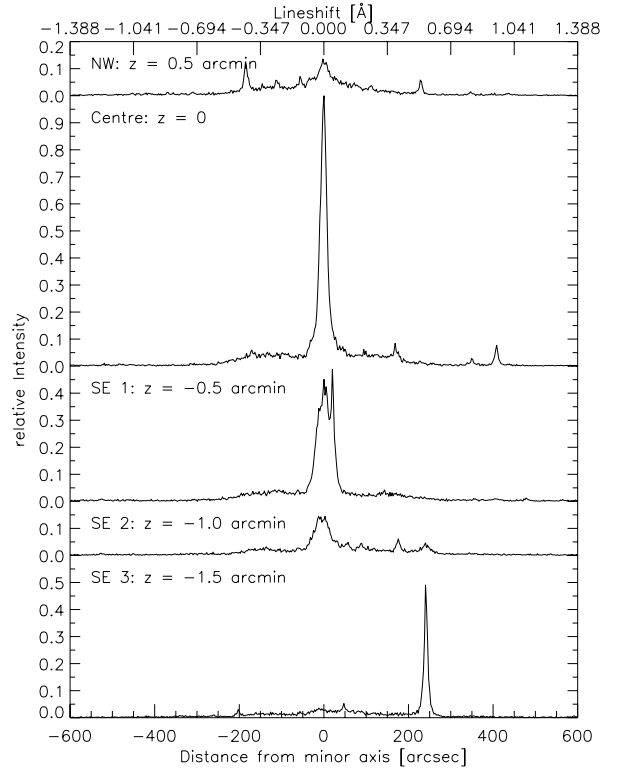


Fig. 1. EPIC PN brightness profiles along the RGS dispersion direction of each extraction region. The strong peak at $0''$ is caused by the outflow emission. In the centre region, shown in Fig. 2, this is superimposed by the nuclear source of the galaxy. Point sources in the extraction regions are seen as sharp spikes in the profile. A positive distance points parallel to the major axis towards the southwest. The distance from the galactic major axis is given by the value z in the captions of the individual extraction regions.

southwest edge of the outflow in region SE 1 (cf. Figs. 1 and 2), however, does affect the outflow spectrum. The spectrum of this source does not show line features (X33 in Pietsch et al. 2001), so its contribution to the RGS spectrum from this region is an increased continuum flux. This does not affect our conclusions.

We analysed the data using the Science Analysis System (SAS), version 6.5.0, together with the most recent calibration files available at the time of the analysis. The metatask `rgsproc 1.19.6` was used to process the RGS data. We first filtered the data sets for time periods with low contamination by low-energy protons. Then we extracted a light curve from the background region on CCD 9, the chip closest to the optical axis of the telescope, thus the most affected by background flares, to determine the threshold count rate, which we then used to filter the eventfiles. The thresholds were 0.25 ct s^{-1} for observations 0125960101 and 0125960201, and 0.20 ct s^{-1} for observation 0152020101, where the count rate was more stable during the non-flaring time intervals.

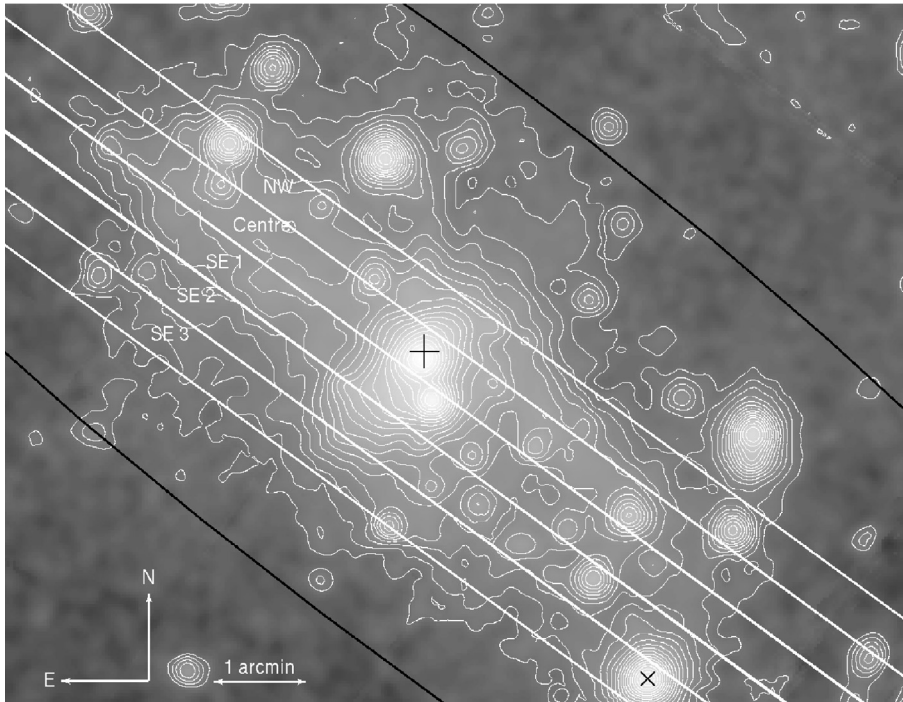


Fig. 2. Logarithmically-scaled EPIC PN+MOS image of the central region of NGC 253 in the energy band 0.5–2.0 keV. The centre of the galaxy is marked with the black cross. The D25 ellipse is overlaid in black, and brightness contours and the extraction regions for the RGS spectra are overlaid in white. The black X at the bottom of the image marks the source X21.

2.1. RGS spectra

Since the dispersion direction of the RGS was approximately aligned with the major axis of the galaxy, we were able to extract spectra for different adjacent regions along the minor axis, i.e. the cross-dispersion direction, of the galaxy (see Fig. 2). The extent of the extraction regions are $30''$ in the cross-dispersion direction. Assuming a distance to NGC 253 of 2.58 Mpc (Puche et al. 1991) this corresponds to a width of 375 pc for the extraction regions. The events in these regions were additionally filtered with a CCD pulseheight filter to select only the $m = -1$ spectral order. The reference points for the origins of the energy scales of the spectra were set on the minor axis of the galaxy, where the outflow has its peak emission (see Fig. 1).

NGC 253 is an extended source and covers most of the area of the RGS detectors. To prevent contamination in the background spectra we used the task `rgsbkgmodel 1.1.5`, which computes background spectra from RGS background templates.

To increase statistics we combined the spectra of the two RGS detectors and added up the spectra from the three observations. As the position angles of *XMM-Newton* did not differ very much in these three observations, the regions from which the spectra were extracted are only slightly tilted relative to each other. A difference in position angle between observations has the effect of degrading the spatial coincidence in the extraction regions. In our case, however, position angle differences are small enough that we can neglect this error. A correction of the spectra for effective area and the combination of the spectra from different observations and instruments was done with the task `rgsfluxer`. The task’s description states that the fluxed spectrum produced by `rgsfluxer` should not be used for any serious analysis of the data. If we use it in spite of this warning, we have to consider that the following effect will add to uncertainties: The task `rgsfluxer` neglects the redistribution of monochromatic response into the dispersion channels, so the intrinsic line broadening of the detector is not removed from the spectrum. As we do not determine line positions or line widths and since we integrate over the whole line including its wings

to derive fluxes for individual lines, this effect does not restrict our analysis. Additionally, we are only interested in relative line fluxes and not in absolute values.

To obtain acceptable statistics ($>3\sigma$) for most of the emission lines while maintaining a high spectral resolution of $\sim 0.39 \text{ \AA}$ in the spectra we combined six channels at a time.

For the observations the dispersion direction of the spectrometers was approximately aligned along the major axis of the galaxy. Therefore the spectra for different cross-dispersion areas correspond to regions with a different distance z from the galactic plane (cf. Fig. 2). The spectra obtained from these regions are shown in Fig. 3.

2.2. RGS images

The RGS is a slitless spectrometer, so the whole observed target within the FOV is imaged on the detector plane, with an offset in dispersion direction for every wavelength it emits in. This aspect can be used to extract narrow band images for various emission lines. In the cross-dispersion direction, the image in a selected emission line is directly mapped onto the detector CCD, while in the dispersion direction the observed object is compressed into a narrow region. This technique was first applied to RGS data of the supernova remnant DEM L 71 by van der Heyden et al. (2003), and with his help (private communications) we developed our own code to produce these narrow band images. The procedure is as follows. The low background eventlists are filtered for the wavelength range of the desired line and for the “banana region” in wavelength-energy-space to exclude second-order spectra and noise. By setting the wavelength range narrow enough, we made sure that no neighbouring lines would be in the extraction region. The dataset thus obtained has to be converted into spatial coordinates and to be uncompressed along the dispersion axis using the following equation as described in the SAS task `rgsangles`:

$$\Delta\beta = \frac{\sin\alpha}{\sin\beta} \Delta\phi \frac{F}{L} \quad (1)$$

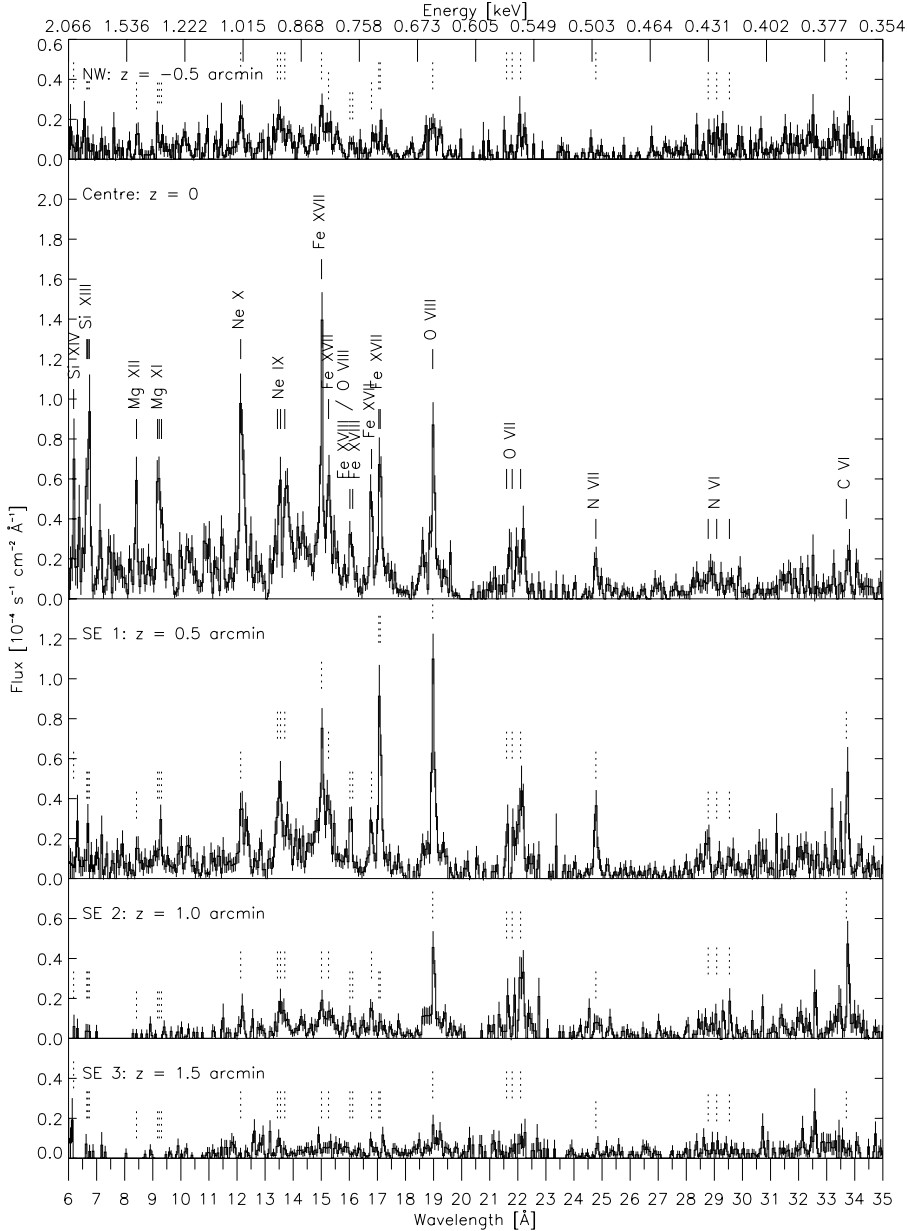


Fig. 3. Combined RGS spectra of NGC 253 extracted from different regions along the outflow. The label gives the southeast position z of the extraction region along the minor axis relative to the centre of the galaxy in arcmin.

with the change in the grating exit angle $\Delta\beta$ due to the offset in the angular component $\Delta\phi$ of an off-axis source parallel to the dispersion direction, the angle of incidence α , the grating exit angle β , the focal length F , and the distance between the Reflection Grating Assembly and the prime focus L .

The images were corrected for exposure and binned to a pixel size of $0.4''$. In the next step we included the RA-DEC coordinate system. As the reference coordinate, we chose the coordinate of the centre of NGC 253 (the position of the brightest IR source in the galaxy, $\alpha_{2000} = 0^{\text{h}}47^{\text{m}}33^{\text{s}}.3$, $\delta_{2000} = -25^{\circ}17'18''$, Forbes et al. 2000). In the cross-dispersion direction, the position of our reference coordinate on the CCD could be taken directly from the source list file that was produced in processing of RGS data. In the dispersion direction the position is given in the above procedure by the Doppler shift corrected line centre position. The galaxy's systemic velocity of 243 km s^{-1} (Koribalski et al. 2004) had to be accounted for. This shifts the reference coordinate in the dispersion direction by $\sim 0.35'' \text{ \AA}^{-1}$ times the centre wavelength in which the image is calculated, e.g. $6.6''$ for the O VIII image. The effects of the velocities of the earth with

respect to the sun and *XMM-Newton*'s orbital velocity can be neglected since they are only of the order of $0.7''$ and $4'' \times 10^{-4}$, respectively, in the O VIII image where the effect would be greatest, and therefore much smaller than the width of the point spread function. The images for the lines were created separately for each of the three observations and were then combined into one image. In a final step the images were smoothed with a Gaussian filter.

In general the method is affected by two different effects: (i) a Doppler shift due to the radial velocity component of an emitting source changes its position in the image along the dispersion direction axis. A radial velocity of 1000 km s^{-1} would correspond to $21.6''$ at a wavelength of 15.0 \AA . (ii) Assuming there is not just one but two lines in the wavelength extraction interval, we would have two images of the object in the resulting image, superimposed with an offset along the dispersion direction of $7.2' \text{ \AA}^{-1}$. For example the emission in the two Fe XVII lines at 16.780 \AA and 17.055 \AA would be superimposed with an offset of $\sim 2.0'$. The Fe XVII at 17 \AA image is the only case where we actually

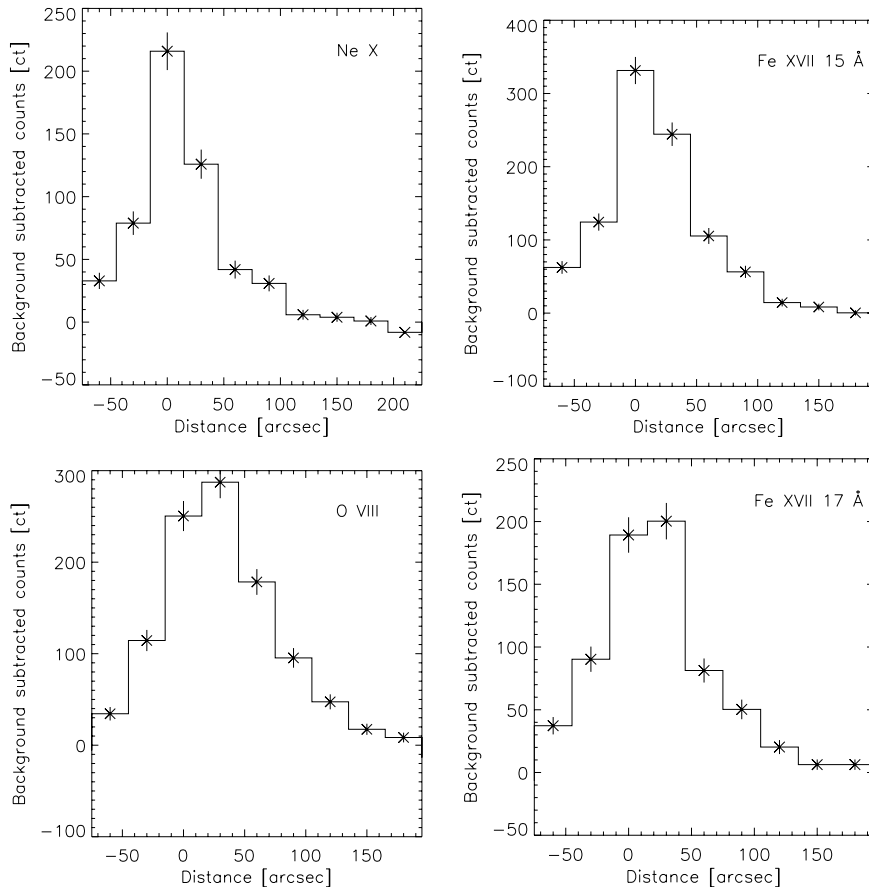


Fig. 4. Ne X (*upper left*), Fe XVII at 15 Å (*upper right*), O VIII (*lower left*), and Fe XVII at 17 Å (*lower right*) profiles of background-subtracted counts against cross-dispersion distance for the combined NGC 253 RGS1 and RGS2 data. The bin at distance zero corresponds to the extraction region “Centre”. Negative distances are towards the northwest, and positive values correspond to areas to the southeast.

have to consider that we have created an image using two lines, i.e. the Fe XVII lines at 17.055 Å and at 17.100 Å. The lines are not separated in the spectra and they appear about equally strong, which is most likely an effect of the low statistics. According to theory (Mewe et al. 1985), the line strength of the latter line should be $\sim 52\text{--}85\%$ of the former one, depending on the temperature of the plasma (1.08 keV and 0.11 keV, respectively). The separation of the lines is 0.046 Å, which corresponds to a shift in the dispersion direction of 19.9". By smoothing the image with a larger Gaussian with a *FWHM* of 20", we can account for the error we make by using both lines.

Images in the Ne X, Fe XVII and O VIII lines are shown in Fig. 5. The number of photons that were extracted from all three observations are 806, 1077, 816, and 1231 for the images in the Ne X (11.98–12.35 Å), Fe XVII at 15 Å (14.86–15.13 Å), Fe XVII at 17 Å (16.90–17.21 Å), and O VIII (18.80–19.17 Å) lines, respectively. The *FWHM* of the Gaussian filter was 12" for Ne X, Fe XVII at 15 Å, O VIII, and 20" for Fe XVII at 17 Å.

2.3. RGS cross-dispersion profiles

To get additional spatial information on the line distribution, we produced emission line profiles in the cross-dispersion direction. Therefore we extracted events from the RGS eventfiles by applying the same filters in wavelength and wavelength-energy-space as for generating the RGS images, but then we binned the counts into 30" bins to match the extraction regions that were applied to the spectra. Background counts in the respective wavelength ranges were taken from the spectra that were obtained with the RGS background model task and subtracted from the emission line profiles. Four of these profiles are shown in Fig. 4.

2.4. EPIC-PN images

To verify the results from the RGS images, we also extracted EPIC PN narrow-band images in approximately the same energy ranges. Therefore we filtered the PN eventfiles in the energy bands around Ne X (992–1052 eV), Fe XVII at 15 Å (795–844 eV), O VIII (625–690 eV), and Fe XVII at 17 Å (694–734 eV). The spectral resolution of the EPIC PN detector is ~ 70 eV, so it is possible that photons with higher or lower energies contribute to the energy band of interest. There is also contamination from higher energies due to the redistribution in the detector. Photons can lose up to 60% of their energy in the CCD before they are detected. This means that bright features in some energy range can show up to some degree in lower energy bands. The filtered eventfiles of the different observations and instruments were merged using the SAS task *merge*. We created images of the eventfiles and smoothed them with a Gaussian of 6". The resulting images are shown in Fig. 6.

2.5. EPIC-PN brightness profiles

To detect the limb brightening of the outflow, as found by Strickland et al. (2000) and Pietsch et al. (2001), we also extracted brightness profiles perpendicular to the outflowing direction from the merged PN eventfiles of all observations. Furthermore, to check for an energy dependence in the limb brightening, we subdivided them into energy bins with a width of ~ 150 eV starting from 400 eV up to 2000 eV. The emission lines of O VIII and the Fe XVII lines are included in the energy ranges 550–700 eV and 700–850 eV, respectively. Ne X is mostly in the 1000–1150 eV bin. The extraction regions match the regions we used for the RGS spectra. Furthermore we split

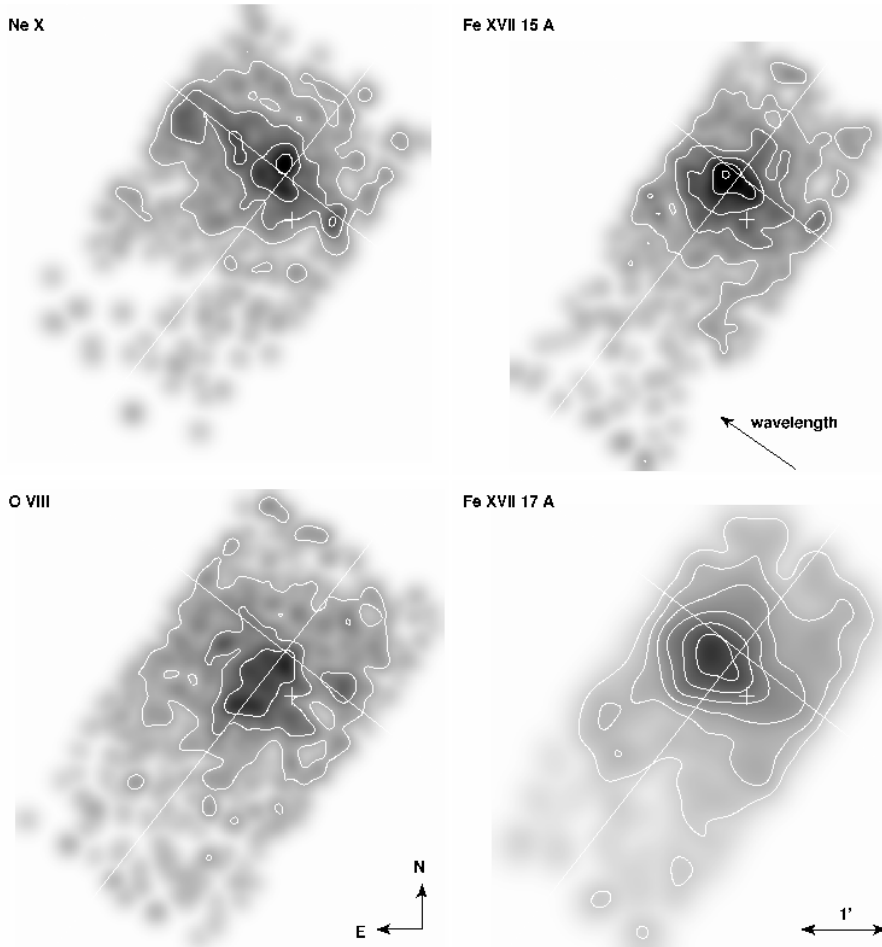


Fig. 5. RGS images of NGC 253 in the Ne X, Fe XVII and O VIII lines. The white lines mark the major (northeast to southwest) and minor axis (northwest to southeast) of NGC 253. The image in the Fe XVII line at 17 Å was smoothed with $FWHM$ of 20'', all others with 12''. White contours indicate 2σ , 3σ , 4σ , ... above the background. The white cross southwest of the nucleus marks the bright source X33 from Pietsch et al. (2000). The RGS dispersion direction is such that wavelength increases from southwest to northeast as indicated by the arrow.

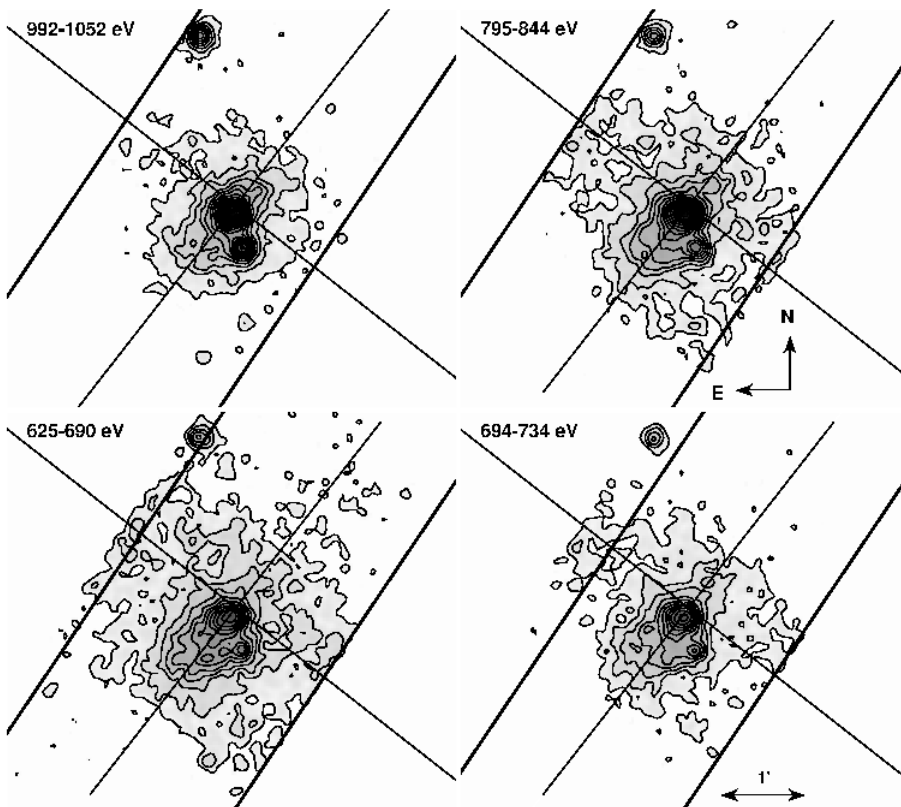


Fig. 6. EPIC PN images of NGC 253 in the energy bands around the Ne X (922–1052 eV), Fe XVII (795–844 eV and 694–734 eV), and O VIII (625–690 eV) lines. The images show the region of NGC 253 that is covered by the RGS images (Fig. 5). The thin black lines mark the major (northeast to southwest) and minor axis (northwest to southeast) of NGC 253. The area within the thick black lines gives the extraction region in the corresponding RGS line image. Black contours indicate 2σ , 3σ , 4σ , ... above the background. The images were smoothed using a Gaussian with a $FWHM$ of 6''.

Table 2. Flux values for different emission lines in the extraction regions.

Transition	$\lambda_{\text{expected}}^a$ (Å)	Flux ($10^{-6} \text{ s}^{-1} \text{ cm}^{-2}$)			
		NW	Centre	SE 1	SE 2
Si XIV	6.18		5.6 ± 4.3		
Si XIII <i>w</i>	6.65		5.8 ± 3.4		
Si XIII <i>x + y</i>	6.69		4.0 ± 2.8		
Si XIII <i>z</i>	6.74		9.2 ± 4.3		
Mg XII	7.11		6.5 ± 3.0	1.9 ± 1.6	
Mg XI <i>w</i>	7.76	1.4 ± 1.3	6.5 ± 2.7	1.3 ± 1.2	
Mg XI <i>x + y</i>	7.81		5.4 ± 2.4	1.7 ± 1.3	
Mg XI <i>z</i>	7.87		6.3 ± 2.6	2.5 ± 1.6	
Ne X	12.1	3.4 ± 2.2	16 ± 4.5	5.7 ± 2.6	2.0 ± 1.6
Ne IX <i>w</i>	13.5		4.3 ± 2.3	5.4 ± 2.5	
Ne IX <i>x + y</i>	13.6	3.0 ± 2.0	6.2 ± 2.8	6.1 ± 2.7	2.4 ± 1.8
Ne IX <i>z</i>	13.7	1.8 ± 1.6	7.6 ± 2.8	4.1 ± 2.2	
Fe XVIII15	15.0	4.5 ± 1.9	18 ± 3.8	13 ± 3.1	3.3 ± 1.6
Fe XVIII17	17.1	3.0 ± 1.7	13 ± 3.8	13 ± 3.9	
O VIII	19.0	6.7 ± 2.5	16 ± 3.8	18 ± 4.0	6.9 ± 2.5
O VII <i>w</i>	21.6		2.3 ± 2.2	3.0 ± 2.4	
O VII <i>x + y</i>	21.8		3.4 ± 2.6		
O VII <i>z</i>	22.1	5.1 ± 3.3	9.5 ± 4.4	13 ± 5.2	
N VII	24.8		2.2 ± 1.5	6.3 ± 2.8	1.5 ± 1.3
N VI <i>w</i>	28.8	1.4 ± 1.3		4.2 ± 2.4	
N VI <i>x + y</i>	29.1			2.4 ± 1.8	
N VI <i>z</i>	29.5			3.2 ± 2.0	1.8 ± 1.5
C VI	33.7	3.0 ± 2.5	4.5 ± 3.0	6.6 ± 3.3	5.7 ± 3.1

^a The references for the expected wavelengths ($\lambda_{\text{expected}}$) are Mewe et al. (1985) and Phillips et al. (1999).

the region SE 1 into the two regions “SE 1 (1)” and “SE 1 (2)” with a width of 15” in the cross-dispersion direction of the RGS. The brightness profiles, sorted by energy band, are shown in Fig. 7. Since we are interested in the emission neither from the nucleus nor from the bright source X33 southwest of the nucleus, the profiles are limited to a maximum of 250 counts. This still shows the main features in the central and SE 1 (1) region, while cutting off the peaks in some cases.

3. Results

3.1. RGS spectra

The RGS spectra show emission in many different lines (Fig. 3). Especially in the region including the major axis of NGC 253 (Centre), a large variety of lines from different elements can be identified. The spectra extend from the Si XIV line at the highest energy down to the C VI line at the low energy end ($\lambda = 6 \dots 34$ Å). All the He-like ions in this range (Si XIII, Mg XI, Ne IX, O VII, and N VI) and their corresponding ions in the next higher ionisation state (Si XIV, Mg XII, Ne X, O VIII, and N VII) can be resolved. The iron 3d–2p transitions around 15 Å, as well as the 3s–2p lines around 17 Å, are detected, i.e. their peak heights are more than twice the error in the wavelength bin. Also the lines at ~ 16.0 and ~ 16.1 Å can be clearly detected, but it is not clear whether the line at ~ 16.0 Å is from Fe XVIII or from O VIII. In the range 10 Å to 11.5 Å there is an indication of iron lines from Fe XXIII and Fe XXIV.

A spectral feature that is prominent in the spectra is a broad base at the Fe XVII lines at 15 Å and 17 Å and at the O VIII and O VII line positions with a width of up to ~ 0.5 Å. If this line broadening is caused by the velocity dispersion of the outflowing gas, it would imply deprojected velocities of up

Table 3. Temperatures of the plasma for different regions of the outflow of NGC 253 derived from line ratios of different elements.

Region	Temperature in keV			
	Si	Mg	Ne	O
NW		0.61 ± 0.08	0.51 ± 0.08	0.21 ± 0.01
Centre	0.79 ± 0.06	0.66 ± 0.04	0.43 ± 0.02	0.22 ± 0.01
SE 1		0.46 ± 0.04	0.38 ± 0.03	0.21 ± 0.01
SE 2			0.25 ± 0.02	0.31 ± 0.04

to $\sim 40\,000 \text{ km s}^{-1}$, corresponding to temperatures of $\sim 10^{11}$ K. Such high velocities disagree with outflow velocities in other galaxies. Extreme cases show values of up to $\sim 3600 \text{ km s}^{-1}$ (e.g. in NGC 3079, Veilleux et al. 1994). A more plausible explanation for the broad base is contributions from the disc emission of NGC 253: In Fig. 1 disc emission extends about 200” in the dispersion direction (major axis of the galaxy) in both directions from the centre. This corresponds to a ~ 0.5 Å shift in wavelength and can therefore explain the observed effect.

The region northwest of the centre (NW) is strongly affected by absorption from the galactic disc that lies between the outflow and the observer. Pietsch et al. (2000) derive an additional absorbing column N_{H} of $1\text{--}2 \times 10^{21} \text{ cm}^{-2}$ for this position northwest of the centre. Therefore most of the lines are only weak or not detectable at all. If we assume that the northwest and southeast outflow have a similar intrinsic spectrum, then absorption can fully account for the difference between regions NW and SE 1, at the same projected distance from the centre. The strongest lines in the NW region are from Ne X, Ne IX, Fe XVII at ~ 15 Å and ~ 17 Å, O VIII and from the forbidden line of the O VII triplet, but most of the emission originates in the disc and is smoothed out into the broad base, as mentioned above.

The regions southeast (SE 1 to SE 3) of the centre do not suffer from this absorption by the disc. One can easily follow how the lines increase or decrease in strength when going away from the galactic disc southeast along the minor axis. The lines at short wavelengths from silicon and magnesium are the first to disappear with distance from the centre. The Ne X line is seen to decrease considerably in strength southeast of the Centre region, whereas the line from the lower ionised Ne IX is not affected as much and is even stronger than the Ne X line in regions SE 1 and SE 2. All the lines from iron decrease in strength except for the Fe XVII lines at 17 Å, which grow by a factor of ~ 1.5 . Also the O VIII line increases in strength. The O VII triplet has about the same strength in region SE 1 as in the Centre region. The lines from N VII and C VI increase in strength compared to the central region. Further away from the centre, in region SE 2, the Fe XVII lines at ~ 15 Å are still detectable and the O VIII line is the strongest line in the spectrum. Also the O VII triplet is still strong. The detection of all the other lines is below 2σ , even though the lines at wavelengths longer than 12 Å can still be identified. Only very weak lines from O VIII and from the Ne IX triplet remain in region SE 3.

Unfortunately the statistics in the spectra are not good enough to allow a quantitative spectral analysis with XSPEC. However, several conclusions can be drawn.

Temperature estimates can be inferred from line ratios of different elements or of the same element in different ionisation states using model calculations. Assuming collisional ionisation equilibrium (CIE), Mewe et al. (1985) calculated line strengths for different elements and transitions depending on the temperatures of the plasma. By measuring the fluxes of two transitions in a spectrum and taking the ratio of these, the obtained value

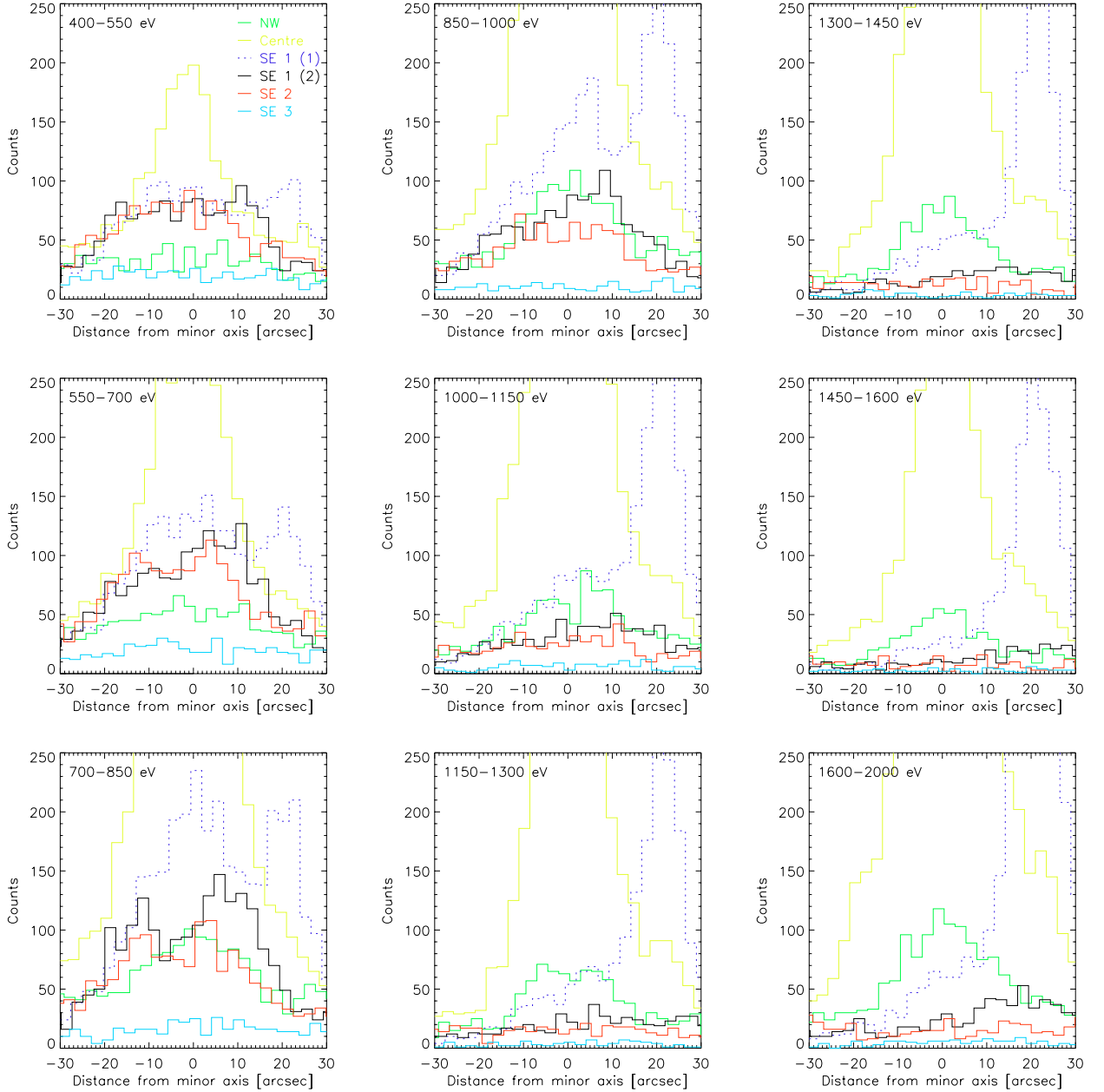


Fig. 7. EPIC PN brightness profiles of the outflow region for the RGS extraction regions sorted by energy band. Note that region SE 1 was split into the two regions “SE 1 (1)” and “SE 1 (2)” with a width of 15” in the cross-dispersion direction of the RGS. The abscissa gives the distance from the minor axis of the galaxy in arcsec. Positive values are to the southwest, negative to the northeast. The presented histograms consist of raw detector counts, i.e. the background and the detector characteristics, like quantum efficiency or effective area, were not taken into account.

can be compared with the tables in Mewe et al. (1985) and the temperature of the plasma can be derived.

We used the line strength ratio between the Ly_α state of a given element and its helium-like charge state that matched the morphology of the Ly_α state in the RGS images best. The line strengths were derived by integrating the flux of the line over the wavelength (cf. Table 2). We found that the derived temperature value and its error show only a weak dependence on how much of the wings of the line we include in the flux integration. The resulting temperatures for Si, Mg, Ne, and O and their variation along the outflow direction are shown in Table 3. Using only the peak height of the line, however, gives temperatures that are lower, except for oxygen in the regions Centre and SE 1, compared to the values shown in Table 3 by up to 50%.

Using the line flux from an emission line (cf. Table 2) and an estimate of the size of the emitting region, we can derive electron densities for the nuclear region of NGC 253 and the southeastern outflow. Mewe et al. (1985) give the line power P' normalised to the electron density for different temperatures and X-ray emission lines. The electron density can then be derived using the formula

$$n_e = \sqrt{\frac{F E_\gamma 4\pi d^2}{V P'}} \quad (2)$$

where n_e is the electron density, F the flux (in counts $\text{s}^{-1} \text{cm}^{-2}$) in an emission line with E_γ from an emitting region with volume V and distance d . The emitting region is assumed to be uniform in n_e .

For the central region we selected the Ne X line and assumed a uniform, spherical emitting region with a radius of ~ 160 pc. The outflow in region SE 1 is represented best in the O VIII line, and we assumed a uniform, cylindrical volume with a radius of ~ 200 pc and a height of 375 pc, the latter being confined by the extraction region. We selected these lines because they are strong in the spectra and because we can get a good estimate for the emitting volume from the RGS images. The resulting electron densities are $n_{e,\text{nucleus}} = 0.106 \pm 0.018 \text{ cm}^{-3}$ and $n_{e,\text{outflow}} = 0.025 \pm 0.003 \text{ cm}^{-3}$ for the nucleus and the outflow region, respectively.

Apart from the derivation of temperatures and electron densities, some selected emission lines can be used as diagnostic lines. The Fe XVII lines at 15 Å and 17 Å can be used to derive the ionising mechanism in the plasma. In the Centre region of NGC 253, as well as in the regions NW and SE 2, the line strengths indicate a predominantly collisional ionised plasma. Region SE 1, however, shows an inverted line ratio. Here the lines at 17 Å are stronger than the lines at 15 Å, which points at a photoionised plasma.

In general the helium-like line triplet of O VII can provide the electron density, the electron temperature, as well as the ionisation process (Porquet et al. 2001). However, we refrained from using the O VII triplet for the following reasons. In the combined and fluxed spectra, the significance is below 3σ for most of the spectral bins. Also the individual lines are not clearly distinguishable from each other. This is probably enhanced due to the use of the task `rgsfluxer` as described above. We also refrained from doing a simultaneous fit of the single uncombined spectra with XSPEC, as the statistics in one single spectrum are barely above 2σ for the strongest bin.

3.2. RGS cross-dispersion profiles

In the cross-dispersion profiles (Fig. 4), line emission is strongest in the Centre bin at a distance of $0''$ for Ne X and Fe XVII at 15 Å. In the O VIII and Fe XVII at 17 Å, profile this is not the case. Here the SE 1 region at a distance of $+30''$ is the brightest. However the northwest half of the central bin is already affected by absorption from the disc. When one bins the brightness profile into smaller spatial regions, the strength drops dramatically in the central bin and towards the northwest. Towards the southeast, on the other hand, it only drops slowly. Farther to the southeast with distances larger than $45''$, the Ne X and the Fe XVII profiles show a large drop in brightness, whereas the O VIII profile declines with a flatter gradient. Therefore the Ne X and Fe XVII emissions are more concentrated within the disc, whereas the O VIII emission extends farther away from the disc. The general behaviour is for the emission from higher energy lines to be more concentrated and not as extended as in M 82 (Read & Stevens 2002).

3.3. RGS images

The RGS images allow us to derive additional information about the spatial composition of the outflow. But before we describe the images in more detail, we need to point out that these images are dominated by low statistics. We therefore restrict our results and conclusions on the regions with good statistics.

The strong O VIII emission clearly traces the outflow. It reaches out to a ~ 750 pc projected distance along the southeast minor axis and has an extent perpendicular to that of ~ 400 pc. Because of the high absorption, emission from the farther side

of the disc is not detected. With the RGS spatial resolution, the outflow in O VIII does not show limb brightening. The emission is strongest close to its central axis and its intensity decreases towards the border. This suggests that the emitting ions are not concentrated on the borders of the outflow where the outflow gas interacts with the surrounding material, but that this emission comes directly from the outflowing gas. The extent to the southeast is less than the one in the EPIC images where the bright outflow emission can be followed to ~ 1.3 kpc (Fig. 6), due both to the sensitivity and to the smaller energy band ($\Delta\lambda = 1.87$ Å vs. 0.37 Å) that was used to extract the images for the RGS. The emission from Ne X is clearly concentrated in the disc and nucleus, and it does not contribute to the outflow. Both images in the iron lines show the strongest emission southeast of the nucleus and the emission is slightly extended along the outflow direction. However, they do not trace the outflow morphology as seen in the O VIII line. This is not surprising, because the excitation cross sections for the iron and O VIII lines have a different temperature dependence. Due to the work that is performed when the outflow expands against the pressure of the ambient medium, as well as to the divergence of the flow in the cone perpendicular to the disc, the temperature or more precisely the kinetic energy of the electrons should decrease (e.g. Breitschwerdt & Schmutzler 1999) with height z above the disc, explaining the relative increase in emission of O VIII to iron.

3.4. EPIC-PN images

We find that the EPIC PN narrowband images (Fig. 6) are all affected by the redistribution effect of the detector; i.e. the image is contaminated by events with higher energies from outside the energy filter boundaries. Especially very bright sources, like the central source of the galaxy, contribute strongly to this effect and all images will therefore show these sources. The bright nuclear source is clearly visible in all of the EPIC PN images. Also, the bright source $\sim 0.5'$ southwest of the nucleus is clearly visible in all EPIC PN images. For the RGS images the latter is only true for the Fe XVII at 15 Å image. This clearly shows the advantage, namely the far better energy resolution, of the RGS images compared to the EPIC PN narrowband images.

With respect to the limb brightening of the outflow emission, the image in the energy range 694–734 eV (including the Fe XVII at 17 Å line) indicates a morphology that could result from a limb-brightened outflow at a distance of $\sim 0.7'$ away from the galactic centre. Also in the energy range 625–690 eV (including the O VIII line), there is an indication of this morphology at a distance of about $1.1'$. The corresponding RGS image unfortunately has statistics too low to confirm this. The other images indicate no limb brightening.

3.5. EPIC-PN brightness profiles

In the EPIC-PN brightness profiles (Fig. 7), the peak in region SE 1 (1) that is caused by the outflow emission always coincides with the centre. A double peak with a central depression is visible in SE 1 (2) and SE 2 at 700–1000 eV (covering also the iron lines). This may point at a limb-brightened outflow in this region. An indication of this structure in region SE 2 at 550–850 eV could be a redistribution effect of the detector (see above). At other energies no indication of limb brightening can be seen. The profiles show either a flat plateau or a hump peaking around the centre of the outflow. This confirms the findings of the RGS images. Starting from 700 eV region NW comes out

and it is stronger than the two SE 1 regions together in energies above 1150 eV. The strong peak in the region SE 1 (1) at about +20'' distance is the bright point source southwest of the nucleus and not part of the outflow region.

To rule out that our findings are affected by *XMM-Newton*'s spatial resolution, we compared brightness profiles, extracted with the same regions and energy bands from *Chandra* observation 3931, to the EPIC PN profiles. There the double peak is clearly detected in region SE 1 (2) between 700 and 1000 eV, possibly also up to 1150 eV. It is not visible in regions SE 1 (1) or SE 2. The profiles in region NW show the same behaviour as in the EPIC-PN data, so taking the differences in the instruments into account *XMM-Newton* and *Chandra* give a consistent picture.

4. Discussion

4.1. Line ratios and temperatures

The extracted RGS spectra of the outflow along the minor axis of the galaxy show emission lines from many ions in different ionisation states: the Ly $_{\alpha}$ lines from Si, Mg, Ne, O, and N and also their helium-like charge states. Additionally we see emission lines from Fe XVII and Fe XVIII.

With increasing distance from the nuclear region, the relative flux in the O VIII line intensifies compared to the flux in Fe XVII; i.e. the line flux at longer wavelengths increases. The same effect can also be seen in the line flux ratio of O VII to O VIII. For the oxygen line ratio, this implies that the temperature decreases; the gas is cooling as it flows away from the nucleus. This change in temperature would also affect the excitation of Fe XVII. Since both elements, oxygen and iron, are from the same kind of sources, i.e. type II SNe, it seems unlikely that the distribution is different, unless there is a very different clumpiness in O and Fe initially. Therefore a change in the oxygen to iron line flux is more likely due to the change in temperature than due to a different radial abundance profile.

The line strengths of the Fe XVII lines at 15 Å compared to the one at 17 Å indicates a predominantly collisional ionised plasma. Region SE 1, however, shows an inverted line ratio. There, the flux in the Fe XVII lines at 17 Å is enhanced compared to the lines at 15 Å. This points to a photoionised plasma, but no strong photoionising sources were detected in the vicinity. There are two additional alternatives to create this line ratio:

A) the plasma is highly underionised compared to the ionising electrons. The time since the heating of the plasma was too short to reach an equilibrium state. In these so-called underionised plasmas, inner-shell ionisation is highly operational (Kosenko 2006) and leads to an enhanced 17.10 Å flux from Fe XVII (Doron & Behar 2002). Examples that show underionised plasmas are supernova remnants like N132D (Behar et al. 2001) and Dem L71 (van der Heyden et al. 2003).

B) the plasma is overionised. It can be produced in fast adiabatic cooling of hot ($T \sim 10^8$ K) and almost completely ionised gas expanding out of a superbubble (Breitschwerdt & Schmutzler 1999). The ionisation cross sections, as well as the recombination cross sections, are different for all ions, so when a shock propagates through the outflow region, the ionisation of different elements and different ionisation stages occurs on different time scales for each. As the expansion of the wind occurs on much shorter time scales than radiative recombination, the highly ionised atomic states remain frozen-in. These states then recombine at a later time farther away from the nucleus. This is called delayed recombination. In both regions, Centre and SE 1,

we see emission from Fe XVIII, which is produced in the hot starburst region of the galaxy. As the wind breaks out of the nuclear superbubble and expands into the lower density medium along the minor axis of the galaxy, it cools mostly adiabatically. However, the Fe XVIII in the wind recombines on a larger time scale than the wind fluid, so we have Fe XVIII in abundance, although the electron temperature would preferentially admit Fe XVII. When the Fe XVIII ions then recombine to Fe XVII, the states of Fe XVII that decay via the 17 Å lines are preferably populated compared to the ones that decay via the 15 Å lines. This enhances the flux at 17 Å and could lead to the misleading line ratio in this region.

Temperatures derived from line ratios of individual elements are in the range of 0.21 ± 0.01 to 0.79 ± 0.06 keV. This range is wider than the one found by Strickland et al. (2000) using *Chandra* ($0.46^{+0.11}_{-0.10}$ – $0.66^{+0.10}_{-0.08}$ keV). However Strickland et al. (2000) used only a single-temperature MEKAL hot plasma model for each extraction region. Given the complexity of the X-ray producing mechanisms, this approach is definitely too simple. Our derived temperature range is closer to the one given by Pietsch et al. (2001) from *XMM-Newton* EPIC spectra, who used three MEKAL models (0.15, 0.53, and 0.94 keV). The approach in this paper of deriving temperatures tries to account for the high complexity of the outflow, without the claim of completeness, by actually using one MEKAL model for each element, and by taking advantage of the superior energy resolution of the RGS. This approach already gives almost the whole measured energy range just for the Centre region, whereas Strickland et al. (2000) obtain this range of temperatures by using several different extraction regions along the outflow.

Looking at the temperatures of just one extraction region, we see that different elements give different temperatures. This could be due to sampling different regions along the line of sight, but also due to delayed recombination. The latter is a true non-equilibrium ionisation (NEI) situation and the result would be measuring different temperatures for different ions. All of the above temperatures were derived by the use of models that assume CIE for the emitting gas. This is not necessarily the case for an outflow. To avoid this assumption we would have to fit the spectra to NEI models (e.g. from Breitschwerdt & Schmutzler 1999). These models entail a higher level of complexity as they depend crucially on a detailed hydrodynamical model of the outflow. Therefore they are sensitive to the time evolution of the wind and, as such, are unique. In other words, NEI models would enable us to fully exploit the information buried in the X-ray spectra, since we could calculate a series of models with different boundary conditions and then derive suitably binned synthetic spectra for fitting the observational results. Thus a satisfactory fit would determine the physical parameters of the outflow within certain limits.

The Fe lines could be used as well to derive temperatures. However, the emission from Fe XVII at ~ 15 Å (2p–3d) and at ~ 17 Å (2p–3s) are affected by delayed recombination of Fe XVIII, which contradicts the CIE assumption. Therefore the results could be misleading and are not shown in Table 3. However from the fact that the Fe XVII at 15 Å lines are present and strong in the spectra, the temperature has to be above 0.22 keV in the regions where the Fe L shell originates; otherwise Fe XVII line formation is inefficient (Rugge & McKenzie 1985). Furthermore calculations show that, above 0.60 keV, Fe XVIII is strong while Fe XVII is weak. This is not the case in any of the regions from which we obtained spectra, therefore we expect that $kT = 0.2 \dots 0.6$ keV.

The RGS images give another indication of the cause of the large temperature spread. The different distribution of the events for each emission line strongly points to the fact that plasmas in different regions of space are responsible for the line emission. It is not surprising that these plasmas then do not have the same temperature and that the temperatures for each element are different. The time scale for attaining pressure, hence temperature equilibrium in the starburst region, is of the order of the sound crossing time. If we compare this to the average time interval between supernova explosions, we can assess the smoothness of the temperature distribution. A rate of 0.05 yr^{-1} has been quoted for NGC 253 (Colina & Perez-Olea 1992). A supernova blast wave expanding within a hot tenuous gas has a low Mach number, so that the sound crossing time, τ_{sc} , in the starburst region is roughly the time scale for the shock wave to reheat the hot gas. Adopting a value of $20''$ or $d = 250 \text{ pc}$ and a temperature of 0.6 keV or $6.6 \times 10^6 \text{ K}$ for the central region yields a value of $\tau_{\text{sc}} \sim d/c_s \approx 8.2 \times 10^5 \text{ yr}$, which is much larger than the interval of 20 years between successive explosions, even if we decrease the supernova rate by a factor of 10–100 for the starburst region. Therefore temperature inhomogeneities in the starburst region and hence variations in the ionisation stages are to be expected. As the flow moves away from the disc sources, the temperature structure should, however, become smoother with time when the flow time becomes larger than the sound crossing time τ_{sc} .

Strickland et al. (2000) argue that the spectral variation in *Chandra* spectra “along and between the northern and southern outflow cones is due to variations in the absorption column and not due to significant temperature variations along the outflow”. A higher absorption column is certainly the case for the NW region, resulting in temperature values that are too high. Although from the analysis of the *XMM-Newton* RGS spectra we cannot draw conclusions regarding the absorbing column, we find that there is a significant temperature variation along the outflow. The temperatures for Mg and Ne decrease with distance to the southeast from the centre. The temperature of oxygen seems to be constant, except for the value in region SE 2.

4.2. The morphology of the outflow

The only RGS image that clearly shows the outflow geometry is the one in the O VIII line. The shape of the southeastern outflow matches roughly the truncated cone with an opening angle of $\sim 26^\circ$ that is seen in *Chandra* and $\text{H}\alpha$ data (Strickland et al. 2000), though the angle in the O VIII image appears to be slightly smaller. Also, this emission is not limb-brightened. The image suggests that the outflow cone is filled with clumpy distributed O VIII. If this is correct, then we see the emission from the hot wind fluid itself for the first time. The clumpiness may be caused by mass-loading, i.e. the turbulent process of mixing in ambient ISM and infalling material. This detection of the wind fluid, however, contrast with the statement from Strickland & Stevens (2000) that the wind is too thin to emit efficiently enough to be detected.

In *XMM-Newton* EPIC observations by Pietsch et al. (2001) and in *Chandra* observations by Strickland et al. (2000), the outflow shows a limb-brightened morphology. In our analysis we find limb brightening in the Fe XVII at 17 \AA RGS images. The comparison between the different PN images and brightness profiles, though with lower spectral resolution, also shows an energy dependence of limb brightening and a clear detection of the emission from the outflowing gas at the energies below 700 eV .

Strickland & Stevens (2000) found from simulations that low opening angles for outflows are only possible when a thick galactic disc is present. This produces opening angles from a few degrees when the wind starts to blow out of its superbubble and then increases up to values of $\sim 60^\circ$, depending on boundary conditions. Thin-disc models typically show opening angles of $\sim 90^\circ$. However, these simulations assumed an ISM distribution in rotating hydrostatic equilibrium with the gravitational field. Simulations by de Avillez & Breitschwerdt (2004, 2005) show that the ISM in the disc has a highly complex structure. This affects the break-out dynamics of a superbubble via density and pressure gradients and also via “holes” from previous outflows that have not closed yet. Conclusions about the thickness of the disc and its collimating effect are hence difficult.

4.3. Outlook

In a forthcoming paper we will analyse the *XMM-Newton* EPIC data of the outflow and the halo emission with the NEI model from Breitschwerdt & Schmutzler (1999), which will then also consider the ionisation history of the outflowing gas from the nucleus into the halo. This will give more information on the composition, temperature and dynamical structure of the outflow.

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References

- Behar, E., Rasmussen, A. P., Griffiths, R. G., et al. 2001, *A&A*, 365, L242
 Breitschwerdt, D., & Schmutzler, T. 1999, *A&A*, 347, 650
 Colina, L., & Perez-Olea, D. 1992, *MNRAS*, 259, 709
 de Avillez, M. A., & Breitschwerdt, D. 2004, *A&A*, 425, 899
 de Avillez, M. A., & Breitschwerdt, D. 2005, *A&A*, 436, 585
 den Herder, J. W., Brinkman, A. C., Kahn, S. M., et al. 2001, *A&A*, 365, L7
 Doron, R., & Behar, E. 2002, *ApJ*, 574, 518
 Fabbiano, G., & Trinchieri, G. 1984, *ApJ*, 286, 491
 Forbes, D. A., Polehampton, E., Stevens, I. R., Brodie, J. P., & Ward, M. J. 2000, *MNRAS*, 312, 689
 Jansen, F., Lumb, D., Altieri, B., et al. 2001, *A&A*, 365, L1
 Koribalski, B. S., Staveley-Smith, L., Kilborn, V. A., et al. 2004, *AJ*, 128, 16
 Kosenko, D. I. 2006, *MNRAS*, 369, 1407
 Mewe, R., Gronenschild, E. H. B. M., & van den Oord, G. H. J. 1985, *A&AS*, 62, 197
 Phillips, K. J. H., Mewe, R., Harra-Murnion, L. K., et al. 1999, *A&AS*, 138, 381
 Pietsch, W., Vogler, A., Klein, U., & Zinnecker, H. 2000, *A&A*, 360, 24
 Pietsch, W., Roberts, T. P., Sako, M., et al. 2001, *A&A*, 365, L174
 Porquet, D., Mewe, R., Dubau, J., Raassen, A. J. J., & Kaastra, J. S. 2001, *A&A*, 376, 1113
 Puche, D., Carignan, C., & van Gorkom, J. H. 1991, *AJ*, 101, 456
 Read, A. M., & Stevens, I. R. 2002, *MNRAS*, 335, L36
 Rugge, H. R., & McKenzie, D. L. 1985, *ApJ*, 297, 338
 Stevens, I. R., Read, A. M., & Bravo-Guerrero, J. 2003, *MNRAS*, 343, L47
 Strickland, D. K., & Stevens, I. R. 2000, *MNRAS*, 314, 511
 Strickland, D. K., Heckman, T. M., Weaver, K. A., & Dahlem, M. 2000, *AJ*, 120, 2965
 Strüder, L., Briel, U., Dennerl, K., et al. 2001, *A&A*, 365, L18
 Tanaka, T., Sugihara, M., Kubota, A., Makishima, K., & Takahashi, T. 2005, *PASJ*, 57, 507
 Turner, M. J. L., Abbey, A., Arnaud, M., et al. 2001, *A&A*, 365, L27
 van der Heyden, K. J., Bleeker, J. A. M., Kaastra, J. S., & Vink, J. 2003, *A&A*, 406, 141
 Veilleux, S., Cecil, G., Bland-Hawthorn, J., et al. 1994, *ApJ*, 433, 48