

## Lithium and zirconium abundances in massive Galactic O-rich AGB stars<sup>★,★★</sup>

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### ABSTRACT

Lithium and zirconium abundances (the latter taken as representative of s-process enrichment) are determined for a large sample of massive Galactic O-rich AGB stars, for which high-resolution optical spectroscopy has been obtained ( $R \sim 40\,000$ – $50\,000$ ). This was done by computing synthetic spectra based on classical hydrostatic model atmospheres for cool stars and using extensive line lists. The results are discussed in the framework of “hot bottom burning” (HBB) and nucleosynthesis models. The complete sample is studied for various observational properties such as the position of the stars in the IRAS two-colour diagram ( $[12] - [25]$  vs.  $[25] - [60]$ ), Galactic distribution, expansion velocity (derived from the OH maser emission), and period of variability (when available). We conclude that a considerable fraction of these sources are actually massive AGB stars ( $M > 3\text{--}4 M_{\odot}$ ) experiencing HBB, as deduced from the strong Li overabundances we found. A comparison of our results with similar studies carried out in the past for the Magellanic Clouds (MCs) reveals that, in contrast to MC AGB stars, our Galactic sample does not show any indication of s-process element enrichment. The differences observed are explained as a consequence of metallicity effects. Finally, we discuss the results obtained in the framework of stellar evolution by comparing our results with the data available in the literature for Galactic post-AGB stars and PNe.

**Key words.** stars: AGB and post-AGB – stars: abundances – stars: evolution – nuclear reactions, nucleosynthesis, abundances – stars: atmospheres – stars: late-type

### 1. Introduction

The asymptotic giant branch (AGB) is formed by stars with initial masses in the range between  $0.8$  and  $8 M_{\odot}$  in a late stage of their evolution. The internal structure of an AGB star consists of an electron-degenerate C–O core surrounded by a shell of He. During most of the time H burning is the main source of energy for the AGB star but, occasionally, the inner He shell ignites in a “thermal pulse”; and eventually, the byproducts of He burning may reach the outer layers of the atmosphere in a “dredge-up” of processed material (the so-called third dredge-up). Repeated thermal pulses add carbon to the stellar surface. As a consequence, originally O-rich AGB stars can turn into

C-rich AGB stars ( $C/O > 1$ ). This scenario would explain the observed spectral sequence M–MS–S–SC–C in AGB stars (e.g. Mowlavi 1999).

Another important characteristic of AGB stars is the presence of neutron-rich elements (s-elements such as Sr, Y, Zr, Ba, La, Nd, Tc, etc.) in their atmospheres. These species are formed by slow-neutron captures in the intershell region. After a dredge-up episode, according to theoretical models, protons can be partially mixed in the  $^{12}\text{C}$ - and  $^4\text{He}$ -rich region between the hydrogen- and helium-burning shells (e.g. Straniero et al. 1995, 1997; Mowlavi 2002; Lattanzio 2003) and can react with  $^{12}\text{C}$  to form  $^{13}\text{C}$  during the interpulse phase. The  $^{13}\text{C}(\alpha, n)^{16}\text{O}$  reaction releases neutrons, which are captured by iron nuclei and other heavy elements, forming s-elements that can later be dredged up to the stellar surface in the next thermal pulse (Straniero et al. 1995, 1997; Busso et al. 2001).

Another neutron source,  $^{22}\text{Ne}$ , needs to be considered. In this case the activation takes place during the convective thermal pulses (e.g. Straniero et al. 1995, 2000; Vaglio et al. 1999; Gallino et al. 2000).  $^{22}\text{Ne}$  can be formed from  $^{14}\text{N}$ ,

\* Based on observations at the 4.2 m William Herschel Telescope operated on the island of La Palma by the Isaac Newton Group in the Spanish Observatorio del Roque de Los Muchachos of the Instituto de Astrofísica de Canarias. Also based on observations with the ESO 3.6 m telescope at La Silla Observatory (Chile).

\*\* Tables 1–9 are only available in electronic form at <http://www.aanda.org>

and  $^{14}\text{N}$  is formed by the CNO cycle in the hydrogen shell. The  $^{22}\text{Ne}(\alpha, n)^{25}\text{Mg}$  neutron source requires higher temperatures and typically occurs at higher neutron densities than the  $^{13}\text{C}(\alpha, n)^{16}\text{O}$  reaction. Thus, a different s-element pattern is expected depending on the dominant neutron source. According to the most recent models, the  $^{13}\text{C}(\alpha, n)^{16}\text{O}$  reaction is the preferred neutron source for masses around  $1\text{--}3 M_{\odot}$ , while for more massive stars (i.e.  $M \gtrsim 3\text{--}4 M_{\odot}$ ) neutrons are thought to be mainly released through the  $^{22}\text{Ne}(\alpha, n)^{25}\text{Mg}$  reaction (see, for example, Busso et al. 1999 and Lattanzio & Lugaro 2005, for a recent review). In the literature, there is strong evidence that most Galactic AGB stars enriched in s-process elements have masses around  $1\text{--}3 M_{\odot}$ , where the  $^{13}\text{C}(\alpha, n)^{16}\text{O}$  reaction is the neutron donor (e.g. Lambert et al. 1995; Abia et al. 2001). Unfortunately, a confrontation of the predictions made by these models with observations of more massive AGB stars in our Galaxy is not yet available.

In the case of the more massive O-rich AGB stars ( $M > 3\text{--}4 M_{\odot}$ ), the convective envelope can penetrate the H-burning shell thereby activating so-called “hot bottom burning” (hereafter, HBB), which takes place when the temperature at the base of the convective envelope is hot enough ( $T \geq 2 \times 10^7$  K) and  $^{12}\text{C}$  can be converted into  $^{13}\text{C}$  and  $^{14}\text{N}$  through the CN cycle (Sackmann & Boothroyd 1992; Wood et al. 1983). The HBB models (Sackmann & Boothroyd 1992; D’Antona & Mazzitelli 1996; Mazzitelli et al. 1999) also predict the production of the short-lived  $^7\text{Li}$  by the chain  $^3\text{He}(\alpha, \gamma)^7\text{Be} (e^-, \nu)^7\text{Li}$ , through the so-called “ $^7\text{Be}$  transport mechanism” (Cameron & Fowler 1971). One of the predictions of these models is that lithium should be detectable, at least for some time, on the stellar surface.

Activation of HBB in massive O-rich AGB stars is supported by studies of AGB stars in the Magellanic Clouds (hereafter, MCs) (Plez et al. 1993; Smith & Lambert 1989, 1990a; Smith et al. 1995), which show a lack of high-luminosity C stars beyond  $M_{\text{bol}} \sim -6$ . Instead, the more luminous AGB stars in the MCs are O-rich. Detection of strong Li overabundances, together with strong s-element enhancement in these luminous AGB stars in the LMC and in the SMC, is the signature that these stars are indeed HBB AGB stars that have undergone a series of thermal pulses and dredge-up episodes in their recent past. The HBB nature of these stars is also confirmed by the detection of a very small  $^{12}\text{C}/^{13}\text{C}$  ratio ( $\approx 3\text{--}4$ ), expected only when HBB is active (Mazzitelli et al. 1999). Unfortunately, current HBB theoretical models have been tested almost exclusively using the results of studying the more massive AGB stars in the MCs but have never been applied to Galactic sources, mainly because of the lack of observations available but also because of inaccurate information on distances within our Galaxy.

Indeed, only a handful of Li-rich stars have been found in our Galaxy so far (e.g. Abia et al. 1991, 1993; Boffin et al. 1993) and, unlike those detected in the MCs, they are not very luminous ( $-6 \leq M_{\text{bol}} \leq -3.5$ ). Most of them are low-mass C-rich AGB stars (Abia & Isern 1996, 1997, 2000) and intermediate-mass S- and SC-stars (Abia & Wallerstein 1998) and not O-rich M-type stars. A few are less luminous red giant branch (RGB) stars at the bump (Charbonnel 2005). Some of these AGB stars are among the most Li-rich stars in our Galaxy, the so-called “super Li-rich” AGB stars (Abia et al. 1991). Under these conditions, however, HBB is not expected to be active and the Li production is not well-understood. Note that HBB is expected to be active only in the most massive (and luminous) AGB stars (from  $\sim 4$  to  $7 M_{\odot}$ ) (Mazzitelli et al. 1999), which should not be C-rich, but O-rich. Actually, the best candidates in our Galaxy

are the so-called OH/IR stars, luminous O-rich AGB stars that are extremely bright in the infrared, showing a characteristic double-peaked OH maser emission at 1612 MHz. These stars are also known to be very long period variables (LPVs), sometimes with periods of more than 500 days and large amplitudes of up to 2 bolometric magnitudes. However, they experience very strong mass loss rates (up to several times  $10^{-5} M_{\odot} \text{ yr}^{-1}$ ) at the end of the AGB, and most of them appear heavily obscured by thick circumstellar envelopes, making optical observations very difficult (e.g. Kastner et al. 1993; Jiménez-Esteban et al. 2005a, 2006a). Thus, no information exists yet on their lithium abundances and/or possible s-process element enrichment.

In this paper we present results from a wide observational programme based on high-resolution optical spectroscopy of a carefully selected sample of Galactic AGB stars thought to be massive from their observational properties. The criteria followed to select the sources in the sample are explained in Sect. 2. The high-resolution spectroscopic observations made in the optical and the data reduction process are described in Sect. 3, while the main results are presented in Sect. 4. In this section we also show how spectral synthesis techniques were applied to derive the atmospheric parameters, as well as the lithium and zirconium abundances, of the stars in our sample. A discussion of these results in the context of HBB models, nucleosynthesis models, and stellar evolution can be found in Sect. 5. Finally, the conclusions derived from this work are given in Sect. 6.

## 2. Selection of the sample

A large sample of long period (300–1000 days), large amplitude variability (up to 8–10 mag in the V band), late-type ( $>M5$ ) O-rich AGB stars displaying OH maser emission with a wide range of expansion velocities (from just a few  $\text{km s}^{-1}$  to more than  $20 \text{ km s}^{-1}$ ) was carefully selected from the literature. These sources are all very bright in the infrared and, as such, included in the IRAS Point Source Catalogue (Beichman et al. 1988). Some of them are so strongly reddened that they do not show any optical counterpart in the Digitized Sky Survey<sup>1</sup> (DSS) plates, although in a few cases the detection of these sources in the optical at a different epoch has been reported in the literature.

Stars were included in the sample if they satisfied at least one of the above conditions (ideally as many of them as possible), which guarantees that they are actually relatively massive stars. In Table 1 we list the 102 stars selected for analysis, together with their IRAS names and other being names from the literature, the Galactic coordinates, the variability type taken from the Combined General Catalogue of Variable Stars (GCVS, Kholopov 1998), the spectral type taken from Kwok et al. (1997) and references therein, and the IRAS colour indices  $[12] - [25]$  and  $[25] - [60]$ <sup>2</sup>. Similar information for a few well-known Galactic M-supergiants and S-, SC-, and C-type AGB stars that were also observed for comparison purposes is presented in Table 2. In addition, nine stars classified as Mira-type stars in SIMBAD<sup>3</sup> and included in the initial sample were found to exhibit observational properties corresponding to C-rich stars. These are listed in Table 3.

<sup>1</sup> The Digitized Sky Survey was produced at the Space Telescope Science Institute under US Government grant NAG W-2166.

<sup>2</sup> Defined as  $[12] - [25] = -2.5 \log(F_{12}/F_{25})$  and  $[25] - [60] = -2.5 \log(F_{25}/F_{60})$ , where  $F_{12}$ ,  $F_{25}$  and  $F_{60}$  are the IRAS flux densities at 12, 25 and 60  $\mu\text{m}$ , respectively.

<sup>3</sup> Set of Identifications, Measurements, and Bibliography for Astronomical Data, operated at CDS, Strasbourg, France.

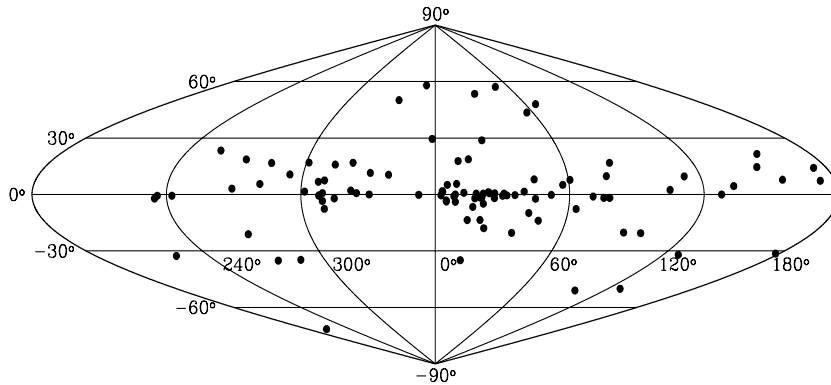


Fig. 1. Galactic distribution of the sources in the sample.

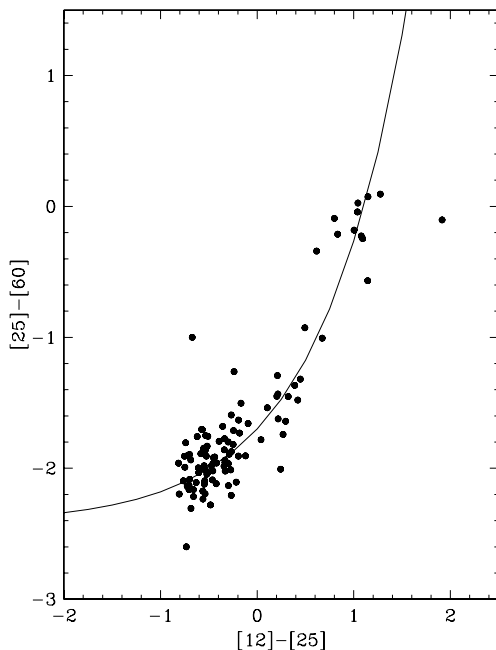


Fig. 2. IRAS two-colour diagram  $[12] - [25]$  vs.  $[25] - [60]$ , where the positions of all the sources in the sample are represented (black dots). The continuous line is the “O-rich AGB sequence” (see text) which defines the sequence of colours expected for O-rich AGB stars surrounded by envelopes with increasing thickness and/or mass loss rates (Bedijn 1987).

The Galactic distribution of the stars in our sample is shown in Fig. 1. The distribution observed clearly suggests that most of them must belong to the disc population. A few sources with bright optical counterparts located at high Galactic latitudes may represent a subset of nearby stars.

In Fig. 2 we show the position of these stars in the IRAS two-colour diagram  $[12] - [25]$  vs.  $[25] - [60]$ . They fall along the observational “O-rich AGB sequence” (García-Lario 1992), as was expected and with very few exceptions. This sequence agrees very well with the model predictions by Bedijn (1987) and is interpreted as a sequence of increasing thickness of the circumstellar envelopes and/or mass loss rates in O-rich AGB stars.

### 3. Observations and data reduction

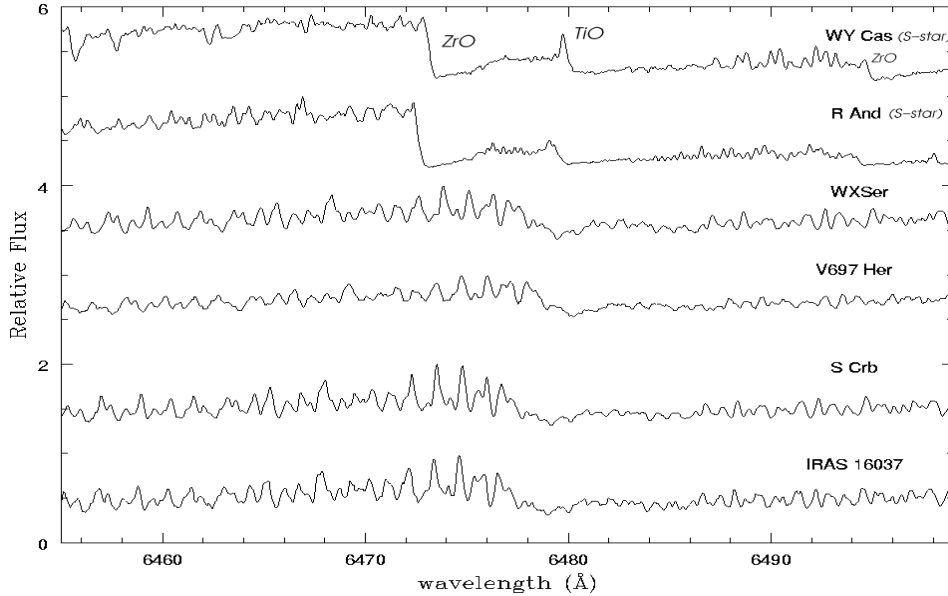
The observations were carried out during several observing periods in 1996–97. High-resolution spectra were obtained using the Utrecht Echelle Spectrograph (UES) installed at the Nasmyth

focus of the 4.2 m William Herschel Telescope at the Spanish Observatorio del Roque de los Muchachos (La Palma, Spain) over three different observing runs in August 1996 (run #1), June 1997 (run #2) and August 1997 (run #3) and the CAspec Echelle Spectrograph (CASPEC) of the ESO 3.6 m telescope at the European Southern Observatory (La Silla, Chile) in February 1997 (run #4). The full log of the spectroscopic observations is shown in Table 4, including more detailed information on the telescopes, dates of the observations, instrumentation used in each run, spectral dispersion, as well as on the spectral range covered.

We used a TEK  $1124 \times 1124$  CCD detector during the first run at the 4.2 m WHT with UES in August 1996 and a SITE1  $2148 \times 2148$  CCD during the second and third runs in June 1997 and August 1997, respectively. Since we were mainly interested in the spectral range between  $6000 \text{ \AA}$  and  $8200 \text{ \AA}$ , we used the  $31.6 \text{ line/mm}$  grating in order to provide full coverage of this spectral region in a single spectrum. With a central wavelength around  $6700 \text{ \AA}$  our spectra extended over  $\sim 4000 \text{ \AA}$  during the first run and  $\sim 6000 \text{ \AA}$  during the second and third runs. The spectra were taken with a separation between orders of around 21 pixels (or  $7.5''$ ), which is large enough to allow sky subtraction, taking into account that the targets in our sample are non-extended. The resolving power was around 50 000, equivalent to a spectral resolution of  $0.13 \text{ \AA}$  around the Li I line at  $6708 \text{ \AA}$ . The selected set-up covered the spectral regions from  $5300 \text{ \AA}$  to  $9400 \text{ \AA}$  in about 47 orders with small gaps in the redder orders for the first run while the  $4700\text{--}10300 \text{ \AA}$  region was covered in about 61 orders without any gaps for the second and third runs.

CASPEC spectra, taken at the ESO 3.6 m telescope, covered the wavelength range from  $6000 \text{ \AA}$  to  $8200 \text{ \AA}$ . The red cross-disperser ( $158 \text{ line/mm}$ ) was used, which gave a resolving power of  $\sim 40\,000$  (equivalent to a spectral resolution of  $\sim 0.17 \text{ \AA}$  around the Li I line at  $6708 \text{ \AA}$ ) and an adequate interorder separation, using the TEK  $1024 \times 1024$  CCD. With this set-up, the selected spectral region is completely covered in 27 orders with small gaps only between the redder ones.

The observational strategy was similar in all runs. Since the stars of the sample are known to be strongly variable, typically 3–4 mag in the R-band and sometimes more than 8 mag in the V-band, we determined the exposure time according to the brightness of the source at the telescope. The typical exposure times ranged from 10 to 30 min. Two 30-min spectra were taken only for a few very faint stars (and later co-added in order to increase the  $S/N$  ratio of the final spectrum). The goal was to achieve an  $S/N$  ratio of 50–150 in the region around  $6708 \text{ \AA}$



**Fig. 3.** High-resolution optical spectra of sample stars displaying the lack of the ZrO absorption bands at 6474 and 6495 Å compared with two Galactic S-stars (WY Cas and R And), where these bands are prominent. WX Ser (IRAS 15255+1944) and V697 Her (IRAS 16260+3454) are Li-detected while S Crb (IRAS 15193+3132) and IRAS 16037+4218 are Li undetected. The absorption band at  $\sim 6480$  Å corresponds to the TiO molecule.

(the lithium line) in order to resolve this narrow absorption line, usually veiled at these wavelengths by the characteristic molecular bands which dominate the optical spectrum of these extremely cool stars. Note that, because of the very red colours of the sources observed, the S/N ratios achieved for a given star can strongly vary from the blue to the red orders (e.g. 10–20 at  $\sim 6000$  Å while  $>100$  at  $\sim 8000$  Å).

At the telescope, all the 102 O-rich stars listed in Table 1 were tried, but useful spectra were obtained only for 57 (56%). The remaining 45 (44%) sources were either too red to obtain any useful information on the strength of the Li I line at 6708 Å or the optical counterpart was simply not found, in the most extreme cases. Some targets with known, relatively bright optical counterparts in the DSS plates were too faint to be detected on the date of the observations. In contrast, other targets with no optical counterpart in the DSS plates appeared as considerably bright stars. A few Galactic M-supergiants together with some C-, SC-, and S-type AGB stars were also observed for comparison (see list in Table 2). Some of the stars in this latter group are among the most Li-rich stars in our Galaxy. In addition, as we have already mentioned in Sect. 2, a few stars in the initial sample turned out to be carbon stars when observed, and not O-rich as initially suspected (see list in Table 3). The total number of objects observed was 120.

For wavelength calibration, several Th–Ar lamp exposures were taken every night during the UES runs. In the case of the CASPEC observations, these calibration lamp exposures were taken before or after any science exposure at the position of the target in order to keep control of possible dispersion changes with the telescope position. Note that with UES the system is more stable because the instrumentation is located at the Nasmyth focus of the 4.2 m WHT telescope. Finally, the corresponding bias and flatfield images were also taken at the beginning (or at the end) of the night.

The two-dimensional frames containing the echelle spectra were reduced to single-order one-dimensional spectra using the standard ECHELLE software package as implemented in

IRAF<sup>4</sup>. Basically, the data reduction process consists of bias-level and scattered-light subtraction, the search and trace of apertures using a reference bright star, the construction of a normalized flatfield image to remove pixel-to-pixel sensitivity fluctuations as well as the extraction of the 1D spectra from the 2D frames. For the wavelength calibration we selected non-saturated Th–Ar emission lines and third or fourth order polynomials for the fitting. The calibration accuracy reached was always better than 20 mÅ. Finally, we identified the terrestrial features (telluric absorption lines) comparing the target spectra with the spectrum of a hot, rapidly rotating star observed on the same night. It should be noted, however, that the majority of the spectral ranges used in the abundance analysis presented in this paper are not significantly affected by these features.

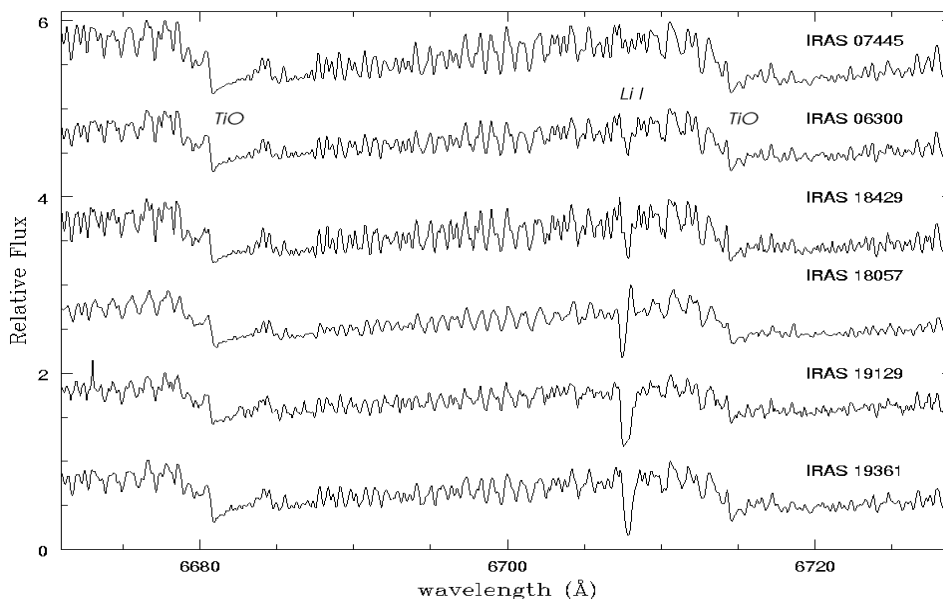
## 4. Results

### 4.1. Overview of the main spectroscopic properties

In general, all stars show extremely red spectra with the flux level falling dramatically at wavelengths shorter than 6000 Å. In addition, the spectra are severely dominated by strong molecular bands mainly due to titanium oxide (TiO), as a consequence of the very low temperature and the O-rich nature of these stars<sup>5</sup>. The bandheads of TiO at  $\sim 6651$ , 6681, 6714, 7055, and 7125 Å are clearly present in all spectra. Interestingly, the bandheads of ZrO at  $\sim 6378$ , 6412, 6474, 6495, 6505, and 6541 Å seem to be absent. These ZrO bandheads (together with those corresponding to other s-element oxides such as LaO or YO) are very strong in Galactic S-stars (see Fig. 3).

<sup>4</sup> Image Reduction and Analysis Facility (IRAF) software is distributed by the National Optical Astronomy Observatories, which is operated by the Association of Universities for Research in Astronomy, Inc., under cooperative agreement with the National Science Foundation.

<sup>5</sup> In contrast, the spectra of the few C-rich AGB stars found in our spectroscopic survey (see list in Table 3) are dominated by CN and C<sub>2</sub> absorption bands.



**Fig. 4.** High-resolution optical spectra of sample stars displaying increasing strength of the Li I line at 6708 Å. Note the P cygni-type profile of the Li I line in IRAS 18057–2616. The jumps at 6681 and 6714 Å correspond to bandheads of the TiO molecule.

The TiO veiling effect is so intense that it is very difficult to identify individual atomic lines in the spectra of these stars with the exception of the Li I line at 6708 Å, the Ca I lines at 6122 and 6573 Å, the K I line at 7699 Å, the Rb I line at 7800 Å, and a few strong Fe I lines. The K I and Rb I resonance lines sometimes show complex profiles, emission over absorption, and blue-shifted components which may have their origin in the expanding circumstellar shell. Actually, the difference between the mean radial velocities of the stellar and circumstellar line components is found to be of the order of the expansion velocity derived from the OH maser emission. In a few cases the Li I line can also show a similar behaviour (see Fig. 4). Finally, some stars display H $\alpha$  emission, which in this type of stars is generally interpreted as the consequence of the propagation of shock waves through the outer layers of the stellar atmosphere.

We detected the presence of the Li I resonance line at 6708 Å in 25% of the sources in the sample (25 stars) with a wide variety of strengths, while in 31% of them (32 stars) we did not find any signature of this line. The remaining 44% (45 stars) were either too red or the optical counterpart was simply not found at the moment of the observations. Sample spectra of stars showing increasing strength of the Li I line at 6708 Å are presented in Fig. 4.

The observed sample has been divided into two different groups on the basis of the detection or non-detection of the lithium line in Tables 5 and 6, respectively. A third group is formed by the stars that were too red (or which did not show any optical counterpart at the telescope), and these are listed in Table 7. The IRAS name, the run in which the stellar field was observed, the OH expansion velocity, and the pulsational period (if available) are given for every source in these tables. The same information is also presented for the comparison stars in Table 8 and for the sample of peculiar C-rich AGB stars in Table 9 (in this latter case the expansion velocities are derived from CO data).

The detection of strong circumstellar Li I and Rb I lines in S Per (IRAS 02192+5821), sometimes classified in the literature as an M-type Galactic supergiant, is remarkable. M-type supergiant stars are not expected to show a strong Li I line (e.g.

Luck & Lambert 1982) or any s-process element enhancement (e.g. Smith et al. 1995). According to our data, S Per looks like a possible genuine O-rich AGB star. In addition, we find a cool effective temperature of  $\sim 3000$  K which is quite different from Levesque et al. (2005) and further investigation is needed. Similarly, PZ Cas (IRAS 23416+6130), another source usually classified as an M-type supergiant (e.g. Levesque et al. 2005), is for the first time clearly identified by us as a possible S-,SC-type AGB star. PZ Cas seems to be slightly enriched in Zr and displays some strong Ba atomic lines but it is not enriched in Li, as expected if PZ Cas is a low-mass S-,SC-type AGB star. Consistent with our possible AGB identification for the latter two sources, they show variability amplitudes of more than 3 mag in the V-band and are the only two stars with periods well beyond 500 days in Table 8. A more detailed analysis of these two stars is beyond the scope of this paper but it would be needed to reach a definitive conclusion. Also interesting is the new detection of a strong lithium line in the peculiar mixed chemistry (C-rich and O-rich) star IRAS 09425–6040, which is studied in a separate paper (García-Hernández et al. 2006a). Lithium was also detected in the C-rich AGB stars IRAS 04130+3918, IRAS 20072+3116 (V1969 Cyg), and IRAS 23320+4316 (LP And), for which a detailed analysis will be presented elsewhere. Finally, the S-star IRAS 10436–3459 (Z Ant), observed for the first time with high-resolution spectroscopy, shows, as expected, intense molecular bands of s-element oxides (like ZrO and LaO) but no lithium. This star will also be analysed in detail in a separate publication.

#### 4.2. Modelling strategy

Atmospheres of cool pulsating AGB stars present a major challenge for realistic, self-consistent modelling. At present, two types of self-consistent models exist. Classical hydrostatic model atmospheres include a sophisticated treatment of microphysical processes and radiative transfer but neglect the effects of dynamics due to pulsation and winds (e.g. Plez et al. 1992). Time-dependent dynamical models for stellar winds (Winters et al. 2000; Höfner et al. 2003, and references therein) include

pulsation, dust formation but a more crude description of radiative transfer. Propagating shock waves caused by stellar pulsation modify the structure of the atmosphere on local and global scales, which give rise to strong deviations from a hydrostatic stratification. The radiative field is dominated by the effect of molecular opacities or even by the dust grains forming in the cool outer layers of these atmospheres. Chemistry and dust formation may be severely out of equilibrium.

Hydrostatic model atmospheres based on the MARCS code (Plez et al. 1992; Gustafsson et al. 2003) reproduce well the observed spectra of C-rich AGB stars (e.g. Loidl et al. 2001) and O-rich AGB stars (e.g. Alvarez et al. 2000) in the optical domain and were thus adopted for the analysis. Indeed, the optical spectra of massive O-rich AGB stars in the MCs have successfully been modelled using this code (e.g. Plez et al. 1993). A comparison of these models with high-resolution observational data in the optical for massive O-rich AGB stars in our Galaxy is unfortunately still lacking.

Our analysis combines state-of-the-art line-blanketed model atmospheres and synthetic spectroscopy with extensive line lists. For this we need to estimate first the range of values of the stellar parameters that can reasonably be adopted for the stars in our sample: effective temperature  $T_{\text{eff}}$ , surface gravity  $\log g$ , mass  $M$ , metallicity  $Z = [\text{Fe}/\text{H}]$ , microturbulent velocity  $\xi$ , and C/O ratio.

Spherically symmetric, LTE, hydrostatic model atmospheres were calculated using the MARCS code for cool stars (Gustafsson et al. 2003). The models are designed with the notation ( $T_{\text{eff}}$ ,  $\log g$ ,  $M$ ,  $Z$ ,  $\xi$ , C/O), where  $T_{\text{eff}}$  is the effective temperature of the star in K,  $g$  is the surface gravity in  $\text{cm s}^{-2}$ ,  $M$  is the stellar mass in  $M_{\odot}$ ,  $Z$  is the metallicity,  $\xi$  is the microturbulent velocity in  $\text{km s}^{-1}$ , and C/O is the ratio between the C and O abundances.

Synthetic spectra were generated with the TURBO-SPECTRUM package (Alvarez & Plez 1998) which shares much of its input data and routines with MARCS. Solar abundances of Grevesse & Sauval (1998) were always adopted, except for the iron abundance, which was taken as  $\log \varepsilon(\text{Fe}) = 7.50$ . Most opacity sources for cool M-type stars were included and taken from the literature. In addition, updated line lists were produced by us for TiO (Plez 1998), ZrO (Plez et al. 2003), and VO (Alvarez & Plez 1998).

For atomic lines, the primary source of information was the VALD-2 database (Kupka et al. 1999). The NIST<sup>6</sup> atomic spectra database was also consulted for comparison. Where possible,  $gf$  values were checked by fitting the solar spectrum. VALD-2  $gf$  values taken from Kurucz (1993, 1994) were often erroneous and did not fit the atomic lines of the solar spectrum. The identification of features was made using the solar atlases by Moore et al. (1966) and Wallace et al. (1993, 1998) and the observed solar spectrum of Neckel (1999). The TURBOSPECTRUM program was run using the abundances from Grevesse & Sauval (1998) together with the solar model atmosphere by Holweger & Müller (1974) with parameters  $T_{\text{eff}} = 5780$  K,  $\log g = 4.44$ , and a variable microturbulence as a function of the optical depth.

The whole machinery was tested on the high-resolution optical spectrum ( $R \sim 150\,000$ ) of the K2 IIIp giant Arcturus ( $\alpha$  Boo) from Hinkle et al. (2000). A MARCS model atmosphere was used with the fundamental parameters determined by Decin et al. (2003a):  $T_{\text{eff}} = 4300$  K,  $\log g = 1.50$ ,  $Z = -0.5$ ,  $M = 0.75 M_{\odot}$ ,  $\xi = 1.7$   $\text{km s}^{-1}$ , C/N/O = 7.96/7.55/8.67, and  $^{12}\text{C}/^{13}\text{C} = 7$ . These parameters are in excellent agreement with other determinations

reported in the literature. Some Kurucz  $gf$  values of metallic lines of Fe I, Ni I, Co I, Cr I, Ti I, Zr I, and Nd II, which were too weak in the solar spectrum, were adjusted so as to yield a good fit to the Arcturus spectrum in several regions (especially in the 7400–7600 Å and 8100–8150 Å regions); otherwise we used the  $gf$  values from the VALD-2 database. This exercise was very useful in confirming the lack of s-process atomic lines in the 7400–7600 and 8100–8150 Å spectral windows in the Galactic O-rich AGB stars in our sample, as we shall see later.

### 4.3. Spectral regions of interest

For the abundance analysis we concentrated our attention on the following spectral regions.

#### 4.3.1. Lithium

We used the Li I resonance line at  $\sim 6708$  Å to derive the lithium abundances, which can be used as a signature of HBB. The selected spectral region ranges from 6670 to 6730 Å and covers the TiO molecular bandheads at  $\sim 6681$  and 6714 Å, which are sensitive to variations in the effective temperature. In addition, the spectral regions around the subordinate Li I lines at  $\sim 6103$  and 8126 Å were also inspected. However, we decided to drop them from the analysis because the latter lines are much weaker. Moreover, the S/N ratio is usually very low around 6103 Å in our sources and the contamination by telluric lines at 8126 Å is very strong.

#### 4.3.2. Rubidium and potassium

We also synthesized the spectral regions around the resonance lines of Rb I at  $\sim 7800$  Å and K I at 7699 Å, with the intention of using the elemental abundances derived from these lines as neutron density and metallicity indicators, respectively. In addition, these spectral regions were also used to check that the model parameters adopted to reproduce the lithium region also provided a reasonably good fit at other wavelengths. Two intervals covering  $\sim 60$  Å in the regions 7775–7835 Å and 7670–7730 Å were selected for this purpose. Unfortunately, the frequent detection of circumstellar components to these lines prevented the accurate determination of the K I elemental abundances, and they will not be discussed here. To a lesser extent the problem also affects the Rb I line. A detailed analysis of the Rb abundances will be presented in a forthcoming paper (García-Hernández et al. 2006b).

#### 4.3.3. s-process elements

The study of the s-process elements was carried out through the analysis of several ZrO absorption bands in the region from 6455 to 6499 Å. In particular, we used the ZrO bandhead around  $\sim 6474$  Å to determine severe upper limits to the Zr elemental abundance, taken as representative of the overall s-process element enrichment. Surprisingly, this band was not detected in any star in our sample, neither did we find any signature from atomic Zr or from atomic lines corresponding to any other abundant s-process element in the 7400–7600 Å and 8100–8150 Å spectral windows, where many atomic lines of Zr I, Nd II, Ba II might potentially have been found, as we show later.

<sup>6</sup> National Institute of Standards and Technology (NIST); [http://physics.nist.gov/cgi-bin/AtData/lines\\_form](http://physics.nist.gov/cgi-bin/AtData/lines_form).

#### 4.4. Derivation of the stellar parameters

In order to analyze how the variations in stellar parameters influence the output synthetic spectra, we constructed a grid of MARCS model atmospheres and generated the associated synthetic spectra with the following specifications: i) the stellar mass was in all cases taken to be  $2 M_{\odot}$ ; ii)  $T_{\text{eff}}$  values ranging from 2500 to 3800 K in steps of 100 K; iii) C/O ratio values of 0.5, 0.7, 0.8, and 0.9; iv)  $\log g$  between  $-0.5$  and  $1.6$  dex in steps of 0.3 dex; v) microturbulent velocity  $\xi$  between 1 and  $6 \text{ km s}^{-1}$  in steps of  $0.5 \text{ km s}^{-1}$ ; vi) metallicity  $Z$  between 0.0 and  $-0.3$  dex; vii)  $\log \varepsilon(\text{Zr})^7$  from 1.6 to 3.6 dex in steps of 0.25 dex; viii) CNO abundances shifted  $\pm 1.0$  dex (in steps of 0.5 dex) from the solar values of Grevese & Sauval (1998); and ix)  $^{12}\text{C}/^{13}\text{C}$  ratios of 10 (as expected for HBB stars) and 90 (the solar value). Finally, the synthetic spectra were convolved with a Gaussian profile (with a certain *FWHM* typically between 200 and  $600 \text{ m}\text{\AA}$ ) to account for macroturbulence as well as instrumental profile effects.

This, however, results in an extremely large grid of synthetic spectra containing thousands of possible combinations of the above parameters! In order to reduce the number of spectra to be considered in our analysis, further constraints were imposed by studying the sensitivity of our spectral synthesis to changes in these parameters.

##### 4.4.1. Stellar mass

As we have already mentioned, the stellar mass was in all cases considered to be constant and equal to  $2 M_{\odot}$ . This is because the temperature and pressure structure of the model atmosphere is practically identical for a  $1 M_{\odot}$  and a  $10 M_{\odot}$  model atmosphere (see Fig. 1 in Plez 1990). Increasing the mass will just have a marginal effect on the output synthetic spectrum through a decrease, for a given gravity, of the atmospheric extension.

##### 4.4.2. Effective temperature

The synthetic spectra are, in contrast, particularly sensitive to  $T_{\text{eff}}$ , which determines the overall strength of the molecular absorption over the continuum (mainly, from TiO molecules but also from VO, and from ZrO if the Zr elemental abundance is increased above a certain limit). Thus, we decided to keep the whole range of values initially considered in our grid, i.e.  $T_{\text{eff}}$  from 2500 to 3800 K in steps of 100 K. Actually, the depth of the molecular bands increases considerably with decreasing temperature. The effective temperature has also a large impact on the strength of the Li I line at  $6708 \text{ \AA}$  in the synthetic spectrum. A decrease of  $T_{\text{eff}}$  has the effect of increasing the TiO veiling, as well as the strength of the Li I line, as a consequence of the changes in the Li I/Li II population equilibrium.

##### 4.4.3. C/O ratio

The C/O ratio basically determines the prevalence of O-rich molecules ( $\text{C/O} < 1$ ) against C-rich molecules like CN,  $\text{C}_2$ , etc. In the case of our O-rich AGB stars, the precise value of the C/O ratio adopted, always  $< 1$ , affects mainly the strength of the TiO veiling and of the Li I line in the synthetic spectrum. An increase of the C/O ratio will decrease the TiO veiling and increase the Li I line strength relative to the adjacent continuum in a very

similar way as a decrease in  $T_{\text{eff}}$ . This is a good example of how a set of parameters providing a good match to the observations for a given spectral region is not necessarily unique. A good fit may be obtained for another combination of parameters that can correspond to quite different Li abundances. Fortunately, one of these two sets of parameters usually does not provide acceptable fits for other spectral regions. In particular, the presence of detectable atomic lines in the synthetic spectrum is very sensitive to variations in the C/O ratio: an increase in the C/O ratio will lower the TiO veiling, making the detection of atomic lines easier. According to the models, the immediate effect should be the detection of much stronger atomic lines of K I, Fe I, Zr I, Nd II, etc. Since we do not see any such effect in our spectra, we conclude that C/O must always be  $\leq 0.75$  in our stars. Abundances derived from models with C/O between 0.15 and 0.75 show actually little differences (Plez et al. 1993). Given that all stars in our sample are clearly O-rich and taking into account the above considerations, we decided to use  $\text{C/O} = 0.5$  in all cases, as a constant value. This selection is in agreement with other more detailed determinations made in massive O-rich AGB stars in the MCs (e.g. Plez et al. 1993; Smith et al. 1995) and in a few low-mass M-type stars studied in our Galaxy (Smith & Lambert 1985, 1990b)<sup>8</sup>.

##### 4.4.4. Surface gravity

The surface gravity also affects the appearance of the output synthetic spectra but its effect is small compared to that of the effective temperature. The influence of a change in the value adopted for the surface gravity within the range of values ( $\log g$  between  $-0.5$  and  $0.5$  dex) under consideration is not as severe as a temperature change. The molecular absorption becomes slightly weaker at higher surface gravities. For example, an increase in the surface gravity of 0.5 dex has approximately the same impact as an increase of 100 K in effective temperature. However, for a fixed temperature we found that the lowest gravities in general fit better both the TiO band strength and the pseudo-continuum around the  $6708 \text{ \AA}$  Li I line. As the appearance of the spectra is not so dependent on the surface gravity, and considering that its value must be low in these mass-losing stars, a constant surface gravity of  $\log g = -0.5$  was adopted for all the stars in the sample. This selection seems to be also appropriate if it is compared with the values adopted by Plez et al. (1993) of  $-0.33 \leq \log g \leq 0.0$  for O-rich AGB stars in the MCs with effective temperatures between 3300 and 3650 K.

##### 4.4.5. Microturbulent velocity

The microturbulent velocity is usually derived by demanding no correlation between the abundance of iron derived from individual lines and their reduced equivalent widths. Previous studies of O-rich AGB stars found microturbulent velocities between 2 and  $4 \text{ km s}^{-1}$  (Plez et al. 1993; Smith & Lambert 1985, 1989; Smith et al. 1995; Vanture & Wallerstein 2002). Unfortunately, we cannot estimate the microturbulent velocity in the stars of our sample due to the lack of useful Fe I atomic lines (the small number of detectable Fe I lines does not cover a wide range of

<sup>7</sup> The zirconium abundance is given in the usual scale  $12 + \log N(\text{Zr})$ . On this scale the solar zirconium abundance is 2.6 dex.

<sup>8</sup> This is the only other chemical analysis previously carried out in Galactic O-rich AGB M-type stars (only six stars!) before the work presented here. However, these were low-mass stars ( $M \lesssim 1.5\text{--}2 M_{\odot}$ ) and their study was focussed on the determination of CNO abundances through near-IR spectra and the s-process enrichment, but they did not study the lithium enrichment in particular.

equivalent widths). Thus, we will assume in the following a microturbulent velocity  $\xi = 3 \text{ km s}^{-1}$ , which is a typical value generally adopted for AGB stars (e.g. Aringer et al. 2002). Higher values of  $\xi$  would strengthen all atomic lines in the synthetic spectra and we do not see this effect in our data.

#### 4.4.6. Metallicity

The metallicity of the Galactic O-rich AGB stars in our sample is assumed to be solar, as expected for stars belonging to the disc population of our Galaxy. A lower metallicity is unlikely since all the stars are expected to be at least of intermediate mass ( $M \gtrsim 2\text{--}3 M_{\odot}$ ), if not more massive. This assumption is also in good agreement with the typical metallicities derived for other AGB stars in our Galaxy (e.g. Abia & Wallerstein 1998; Vanture & Wallerstein 2002). The Ca I lines at  $\sim 6122 \text{ \AA}$  and  $6573 \text{ \AA}$  generally used for this determination in other studies are unfortunately not sensitive enough to metallicity nor to surface gravity at the very low temperature of our sample stars (see Fig. 7 of Cenarro et al. 2002). This was checked by running different test models and spectral syntheses in the  $6100\text{--}6160 \text{ \AA}$  and  $6535\text{--}6585 \text{ \AA}$  regions where these lines fall. A decrease in the metallicity implies less availability of metals like Ti, and thus, lowers the TiO veiling, making the detection of atomic lines easier. Actually, this may be the main reason why they are easily detected in O-rich AGB stars at the metallicity of the MCs (e.g.  $Z = -0.5$  in the SMC, Plez et al. 1993), while we do not detect them in our spectra.

#### 4.4.7. Zirconium

The synthetic spectra are also very sensitive to changes in the zirconium elemental abundance, especially in those wavelength regions where ZrO molecular bands are present. For a set of fixed stellar parameters, an increase in the zirconium abundance leads to stronger bands. Since the zirconium abundance was a priori unknown in our stars, we decided to use a large set of zirconium abundances with  $\log \varepsilon(\text{Zr})$  between 1.6 and 3.6 dex in steps of 0.25 dex. The effect of a variation in the zirconium elemental abundance on the ZrO bands is in general much stronger than any other effect associated with variations in the  $T_{\text{eff}}$  or C/O ratio. In contrast, even at relatively high zirconium abundance, the corresponding Zr atomic lines are more difficult to detect at low  $T_{\text{eff}}$  or low C/O ratio because the TiO and ZrO lines become stronger and more numerous under these conditions compared to the Zr atomic lines.

#### 4.4.8. CNO abundances and $^{12}\text{C}/^{13}\text{C}$ ratio

In the synthetic spectra, relative variations in the carbon, nitrogen, and oxygen abundances and of the  $^{12}\text{C}/^{13}\text{C}$  ratio are completely negligible with respect to other stellar fundamental parameters in the spectral regions under analysis. Thus, we decided to fix them and adopt their solar values in the following analysis.

### 4.5. Abundance determination

#### 4.5.1. Overall strategy

After imposing these further constraints, the grid of MARCS model spectra to consider is composed of “only” a few hundred spectra, with effective temperatures in the range  $T_{\text{eff}} = 2500\text{--}3800 \text{ K}$  in steps of 100 K,  $\log \varepsilon(\text{Zr})$  between 1.6

and 3.6 dex in steps of 0.25 dex, and a variable value of the *FWHM* in the range  $200\text{--}600 \text{ m\AA}$  in steps of  $50 \text{ m\AA}$ , keeping all the other stellar parameters fixed:  $\log g = -0.5$ ,  $M = 2 M_{\odot}$ , solar metallicity  $Z = [\text{Fe}/\text{H}] = 0.0$ ,  $\xi = 3 \text{ km s}^{-1}$ , C/O = 0.5, solar CNO abundances, and solar  $^{12}\text{C}/^{13}\text{C}$  ratios.

In order to find the synthetic spectrum which better fits the observed spectrum of a given star a modified version of the standard  $\chi^2$  test was used. When fitting observed data  $Y_{\text{obsi}}$  to model data  $Y_{\text{synthi}}$ , the quality of the fit can be quantified by the  $\chi^2$  test. The best fit corresponds to that leading to the minimum value of  $\chi^2$ . This test is used here in a modified way and is mathematically expressed as:

$$\chi^2 = \sum_{i=1}^N \frac{[Y_{\text{obsi}} - Y_{\text{synthi}}(x_1 \dots x_M)]^2}{Y_{\text{obsi}}} \quad (1)$$

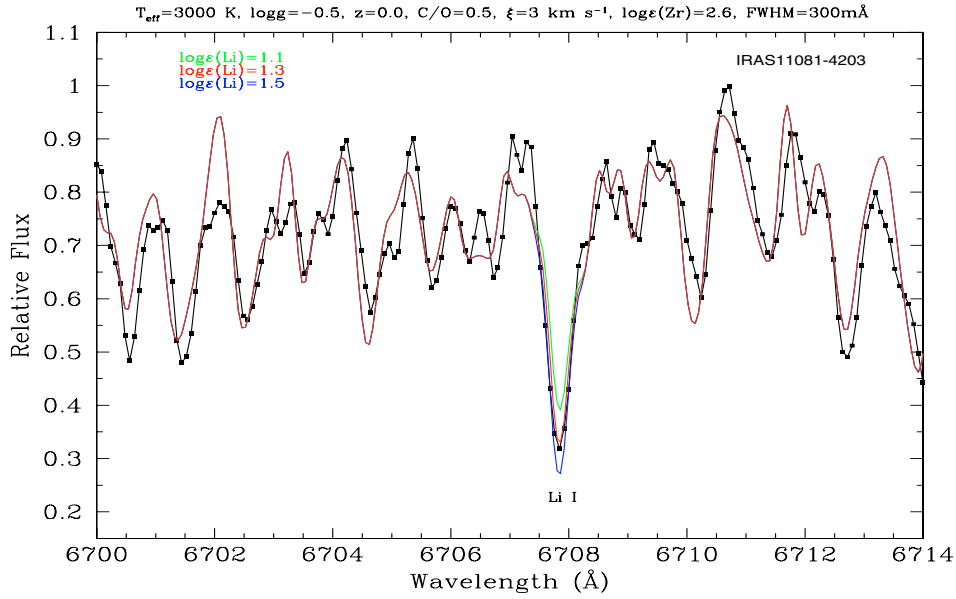
with  $N$  the number of data points and  $M$  the number of free parameters.

The observed spectra, once fully reduced, were shifted to rest-wavelength using the mean radial velocity shift derived from the Li I and Ca I lines. In addition, they were re-binned to the same resolution as the synthetic ones ( $0.06 \text{ \AA}/\text{pix}$ ) and normalized in order to make the comparison easier. The observed spectra were then compared to the synthetic ones in intervals of  $\sim 60 \text{ \AA}$  which allowed the analysis of their overall characteristics as well as of the relative strength of the TiO molecular bands.

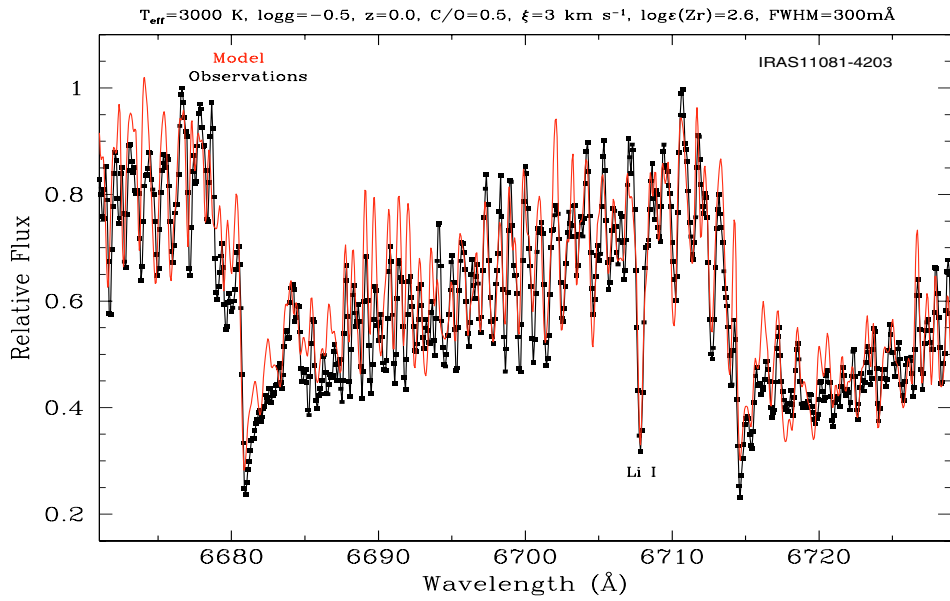
The fitting procedure makes a special emphasis on the goodness of the fit in the lithium region ( $6670\text{--}6730 \text{ \AA}$ ). In this spectral range the goal was to fit the relative strength of the TiO bandheads at  $\sim 6681$  and  $6714 \text{ \AA}$ , which are very sensitive to the effective temperature, together with the pseudo-continuum around the Li I line. We first determined by  $\chi^2$  minimization which of the spectra from our grid of models provided the best fit, mainly fixing  $T_{\text{eff}}$ . The best fits resulting from the  $\chi^2$  minimization were also judged by eye in order to test the method. The lithium content was then estimated by changing the lithium abundance. This procedure was repeated on each star of the sample for which an acceptable *S/N* ratio was achieved around the Li I  $6708 \text{ \AA}$  line. Unfortunately, we could not analyse a few AGB stars with a low signal-to-noise ratio in their spectra at  $6708 \text{ \AA}$  (see Table 7). IRAS 18025–2113, IRAS 03507+1115, and IRAS 18304–0728 showed unusual spectra, very different from the rest of the stars in our sample and a good fit from our grid of MARCS model spectra was not possible. The spectra of these stars show unusual TiO band strengths which could only be reproduced with a low  $T_{\text{eff}}$  but the rest of the spectrum (e.g. the local continuum level) seems to be hotter. In addition, the atomic lines are broader (IRAS 18025–2113 and IRAS 18025–2113) or narrower (IRAS 03507+1115) than the *FWHM* needed to describe the TiO bandheads and lines. Their complicated spectra look like a combination of two temperatures and *FWHMs* suggesting that they could probably be double-lined spectroscopic binaries or strongly affected by shock waves propagating in their atmosphere.

The best model fit in the Li I spectral region ( $6670\text{--}6730 \text{ \AA}$ ) is also usually found to fit the wavelength regions around the ZrO bandhead and the K I and Rb I resonance lines reasonably well. In general, the overall shape of the spectrum (including the TiO bandheads) is very well reproduced. As an example, the fits made in different spectral regions for the star IRAS 11081–4203 are presented in Figs. 5 to 9. The effective temperatures of the best-fitting model spectra together with the lithium abundances





**Fig. 5.** Best model fit and observed spectrum around the Li I line (6708 Å) for the star IRAS 11081-4203. The Li abundance derived from this spectrum was  $\log \varepsilon(\text{Li}) = 1.3$  dex. The synthetic spectra obtained for Li abundances shifted +0.2 dex and -0.2 dex from the adopted value are also shown. The parameters of the best model atmosphere fit are indicated in the top label.

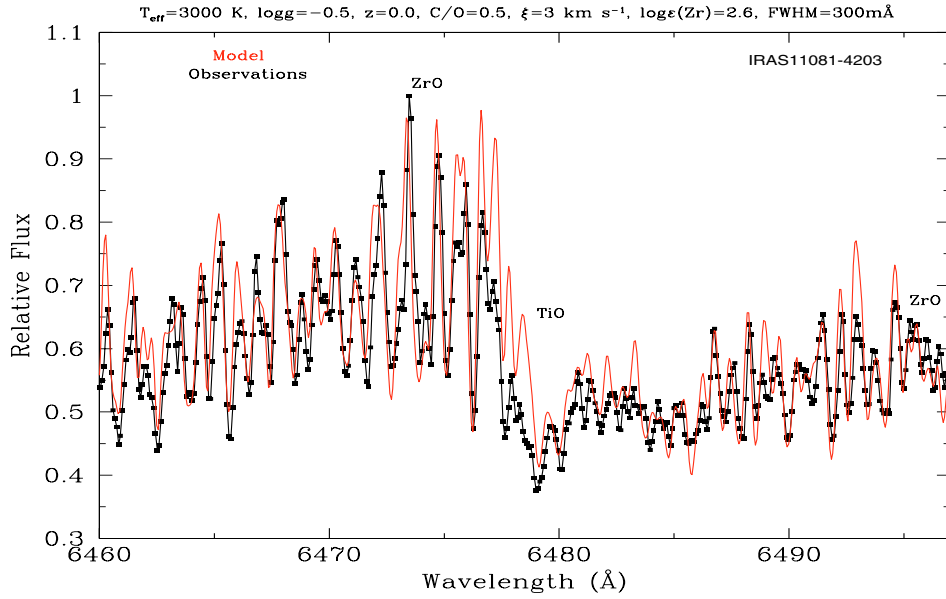


**Fig. 6.** Best model fit and observed spectrum in the region 6670–6730 Å for the star IRAS 11081-4203. The  $T_{\text{eff}}$  derived from this spectrum was 3000 K. The parameters of the best model atmosphere fit are indicated in the top label.

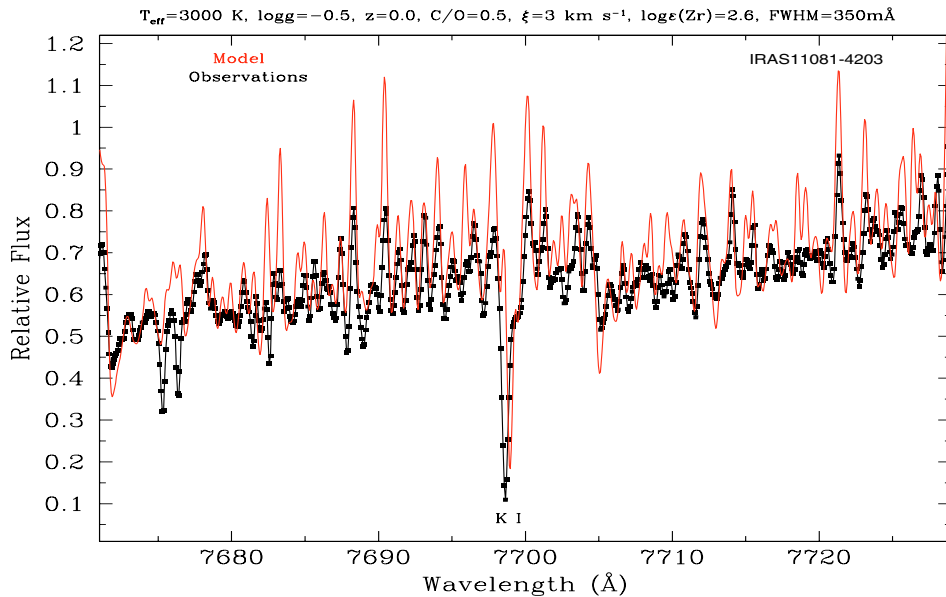
(or upper limits) derived following the above procedure are listed in Table 10, where we have separated the Li detected stars from the Li undetected ones.

For the measurement of the zirconium abundance we concentrated our attention on the synthesis of the 6455–6499 Å spectral region. In this case, the goal was to obtain the best fit around the ZrO molecular bands at 6474 Å and 6495 Å. The non-detection of these features in any of the stars under analysis imposes severe upper limits to the zirconium abundance. To be sure that there was no other way to interpret our data, we also checked the presence of atomic zirconium in the 7400–7600 and 8100–8150 Å spectral regions, where strong Zr I atomic lines should have been detected in that case. Unfortunately, the local continuum in these latter wavelength regions is not well

reproduced in the O-rich AGB stars in our sample, in contrast with the perfect agreement obtained in some of the comparison stars observed, such as the S-star IRAS 10436–3454, for which a perfect fit is obtained both for the ZrO bandheads at 6455–6499 Å (see Fig. 12) as well as for the Zr I atomic lines in the 7400–7600 and 8100–8150 Å spectral regions (not shown). Sample spectra of the region around 7565 Å are shown in Fig. 10, where the position of some atomic lines of Zr I, Nd II, and Fe I are indicated. As we can see, the atomic lines corresponding to s-process elements (e.g. Zr, Nd, La, etc.) were not detected in any of the sample stars observed. In contrast, these s-element atomic lines appear very strong in the Galactic S-star IRAS 10436–3454. Note, however, that the stars in our sample are the coolest O-rich AGB stars ever studied (with effective



**Fig. 7.** Best model fit and observed spectrum in the region 6460–6499 Å for the star IRAS 11081–4203. The Zr abundance derived from this spectrum was  $\log \epsilon(\text{Zr}) = 2.6$  dex, which corresponds to the solar value. The parameters of the best model atmosphere fit are indicated in the top label. Note the non-detection of the characteristic strong ZrO bandheads typical of Galactic S-stars.



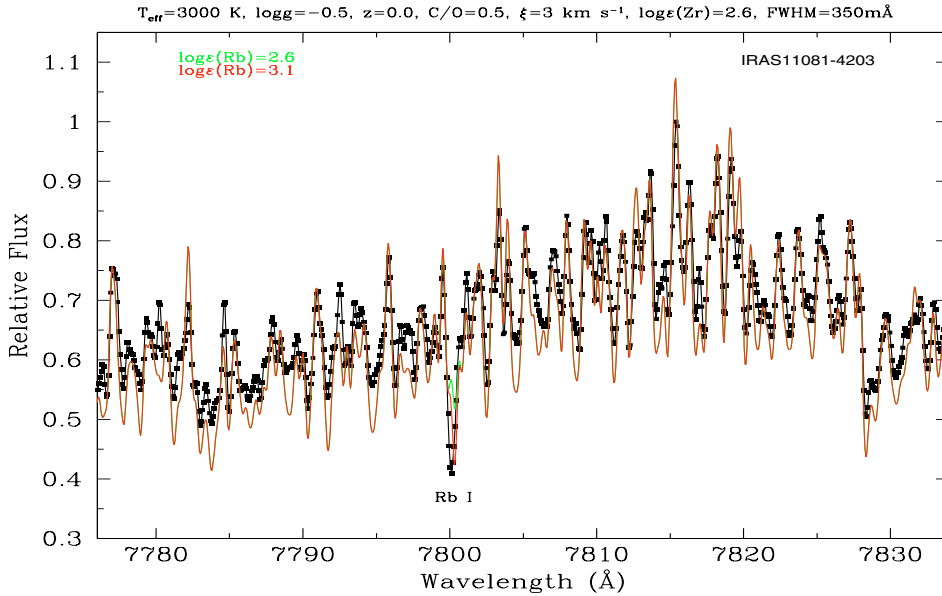
**Fig. 8.** Best model fit and observed spectrum in the region 7670–7730 Å (around the K I line at 7699 Å) for the star IRAS 11081–4203. The synthetic spectrum obtained for solar K abundance ( $\log \epsilon(\text{K}) = 5.12$  dex) is shown. The parameters of the best model atmosphere fit are indicated in the top label. Note the radial velocity shift, which is interpreted as a result of the partially circumstellar origin of the K I line. The narrow features seen at  $\sim 7675$ – $77$  Å, which are not adjusted by the model, are telluric lines.

temperatures between 2700 K and 3300 K; see Table 10). We suspect that the strong discrepancies observed between model and observations may be due to the non-inclusion of other O-rich molecules, such as H<sub>2</sub>O, in our line lists. The effect of H<sub>2</sub>O may be ignored in S-stars, where the C/O ratio is close to unity, but it might become dominant in our O-rich stars at these very low temperatures, preferentially at the longer wavelengths (e.g. Allard et al. 2000; Decin et al. 2003a,b). In addition, lanthanum oxide also has a strong absorption band around 7500 Å and it could have been a good candidate for detection in the spectral range here considered. As for ZrO, the non-detection of the strong LaO molecular bandhead, usually very strong in S-stars at

these wavelengths, is another indication of the lack of s-process elements in our sample stars.

#### 4.5.2. Lithium and s-element abundances

The derived Li abundances are displayed in Table 10, where we can see that almost all Li detected stars show enhanced Li abundances  $\log \epsilon(\text{Li}) \geq 1$ , i.e. larger than solar, but smaller than those found in the so-called “super Li-rich” AGB stars (with  $\log \epsilon(\text{Li}) > 3$ – $4$ ; e.g. Abia et al. 1991) in our Galaxy. A very similar range of Li overabundances was found in the massive O-rich AGB stars studied in the MCs (Plez et al. 1993;



**Fig. 9.** Best model fit and observed spectrum in the region 7775–7835 Å (around the Rb I line at 7800 Å) for the star IRAS 11081–4203. The synthetic spectra obtained for solar Rb abundance ( $\log \varepsilon(\text{Rb}) = 2.6$  dex) and shifted +0.5 dex ( $\log \varepsilon(\text{Rb}) = 3.1$  dex) are shown. The parameters of the best model atmosphere fit are indicated in the top label. Again, the radial velocity shift observed suggests a partially circumstellar origin of the Rb I line.

Smith & Lambert 1989, 1990a; Smith et al. 1995). The errors in the derived Li abundances mainly reflect the sensitivity of our models to changes in the adopted stellar parameters. In particular, they strongly depend on the uncertainties in the determination of the effective temperature ( $\pm 100$ – $200$  K; larger for the coolest stars) and metallicity ( $\pm 0.3$ ), while they are less sensitive to the microturbulent velocity ( $\pm 1$  km s $^{-1}$ ), surface gravity ( $\pm 0.5$ ), and  $FWHM$  ( $\pm 50$  mÅ) uncertainties. The effect of the uncertainty in the location of the pseudo-continuum around the Li I line on the measured abundances is negligible compared with any other change in the adopted stellar parameters. If we consider all these uncertainties as independent sources of error, the resulting Li abundances given in Table 10 are estimated to be affected by errors of the order of 0.4–0.6 dex. As an example, the changes in the derived Li abundance induced by slight variations of each of the atmospheric parameters used in our modelling for IRAS 15255+1944 are shown in Table 11.

Note, however, that the estimated errors do not reflect possible non-LTE effects, dynamics of the stellar atmosphere, or errors in the model atmospheres or in the molecular/atomic linelists themselves. In particular, over-ionization and over-excitation of Li is predicted to occur under non-LTE conditions both in C-rich and O-rich AGB stars (Kiselman & Plez 1995; Pavlenko 1996) although the effect is expected to be more pronounced in metal-deficient stars (i.e. MC AGBs). Use of LTE in Li-rich AGB stars is likely to result in underestimation of the Li abundance (Kiselman & Plez 1995; Abia et al. 1999).

Concerning s-process elements, the non-detection of the ZrO molecular bands at 6474 Å and 6495 Å is found to be consistent with upper limits for the Zr abundance around  $[\text{Zr}/\text{Fe}] = 0.25 \pm 0.25$  dex with respect to the solar value of  $\log \varepsilon(\text{Zr}) = 2.6$  dex. The synthetic spectra predict detectable ZrO bandheads at 6474 Å and 6495 Å for Zr abundances above 0.00–0.25 dex for  $T_{\text{eff}} \geq 3000$  K and above 0.25–0.50 dex for  $T_{\text{eff}} < 3000$  K. Actually, the fits are also reasonably good even if we do not include any ZrO in the synthesis. We found the same results both in the Li non-detected stars and in the Li detected ones.

This result is in strong contrast with the higher Zr abundances ( $[\text{Zr}/\text{Fe}] > 0.5$ ) found in Galactic MS-, S-stars, and in massive O-rich AGB stars in the MCs. The effect of varying the Zr abundance for the Li detected star IRAS 11081–4203 is displayed in Fig. 11, as an example. The best model spectrum fit for the Galactic S-star IRAS 10436–3454 is also shown in Fig. 12 for comparison. As we can see, strong ZrO bandheads are detected in IRAS 10436–3454, while these are completely absent from the spectrum of IRAS 11081–4203. According to our models, a modest Zr enhancement (with respect to the solar value) would be enough for the ZrO molecular bands to show up in our sample stars despite the strong veiling produced by the TiO molecule (see Fig. 11). An even lower Zr abundance is enough to produce strong ZrO bands in S-stars, where the C/O ratio is  $\sim 1$  because of the weaker TiO veiling, as we can see in the spectrum of IRAS 10436–3454 displayed in Fig. 12. This also favours the visibility of the ZrI atomic lines in the 7400–7600 and 8100–8150 Å spectral regions in these stars, as has already been mentioned.

## 5. Discussion

In this section, we will first try to get some information about the absolute luminosities and progenitor masses of stars in our sample. Then, we will discuss our abundance results in the framework of HBB and nucleosynthesis models and we will compare our results with the previous studies made on massive AGB stars in the MCs. Finally, we will interpret the overall results obtained in the context of stellar evolution.

### 5.1. Absolute luminosities

In contrast to the studies made in the past of AGB stars in the MCs, for which a common distance can be assumed leading to relatively well-determined absolute luminosities, the estimation of absolute luminosities for the stars in our Galactic sample is very difficult, mainly because of the large uncertainties

**Table 10.** Spectroscopic  $T_{\text{eff}}$  and Li abundances<sup>1</sup> derived for the subgroup of Li detected (left) and Li non-detected stars (right).

Li detected			Li non-detected		
IRAS name	$T_{\text{eff}}$ K	$\log \varepsilon(\text{Li})$ $12 + \log N(\text{Li})$	IRAS name	$T_{\text{eff}}$ K	$\log \varepsilon(\text{Li})$ $12 + \log N(\text{Li})$
01085+3022	3300	2.4	03507+1115 <sup>2</sup>	...	...
02095-2355	3300	1.6	05098-6422	3000	<-1.0
04404-7427**	...	...	05151+6312	3000	<+0.0
05027-2158	2800	1.1	07080-5948	3000	<+0.5
05559+3825	2900	0.6	07445-2613	2900	<-1.0
06300+6058	3000	0.7	10189-3432	2900	<-1.0
07222-2005**	...	...	10261-5055	3000	<-1.0
07304-2032	2700	0.9	13379-5426	2900	<+0.0
09429-2148	3300	2.2	13442-6109 <sup>4</sup>	3000	<-1.0
11081-4203	3000	1.3	13475-4531	2900	<-1.0
11525-5057	3300	0.9	14086-0730**	...	...
12377-6102**	...	...	14086-6907**	...	...
14337-6215	3300	2.4*	14247+0454	2700	<+0.5
15211-4254	3300	2.3	14266-4211	2900	<+0.0
15255+1944	2900	1.0	15193+3132	2800	<+0.0
15576-1212	3000	1.1	15586-3838	3000	<+0.0
16030-5156	3000	1.5	16037+4218	2900	<-1.0
16260+3454	3300	2.7	16503+0529	2800	<-1.0
18025-2113 <sup>2</sup>	...	...	17034-1024	3300	<+0.0
18057-2616 <sup>3</sup>	3000	...	17359-2138**	...	...
18413+1354	3300	1.8	18050-2213	2900	<-1.0
18429-1721	3000	1.2	18304-0728 <sup>2</sup>	...	...
19129+2803	3300	3.1*	18454-1226 <sup>4</sup>	3300	<+0.0
19361-1658 <sup>5</sup>	3000	1.9/2.0	19147+5004	3000	<+0.0
20052+0554	3300	2.6	19157-1706	3300	<+0.0
			19412+0337	3300	<+0.0
			20343-3020	3000	<-1.0
			20350+3741	3000	<+0.0
			22180+3225	3300	<+0.0
			22466+6942	3300	<+0.0
			22510+3614	3000	<-1.0
			23492+0846 <sup>4</sup>	3000	<-1.0

<sup>1</sup> The uncertainty in the derived Li abundances is estimated to be around  $\pm 0.4$ – $0.6$  dex.<sup>2</sup> Possible double lined spectroscopic binary.<sup>3</sup> The lithium line has a P cygni-type profile.<sup>4</sup> Possible non-AGB star.<sup>5</sup> Observed Li abundances in two different runs.

\* The abundance value must be treated with some caution because the line is resolved in two components (circumstellar and stellar) and the abundance estimate corresponds to the photospheric abundance needed to fit the stellar component.

\*\* The S/N ratio at 6708 Å is very low to derive any reliable  $T_{\text{eff}}$  and/or Li abundance estimate.

involved in the determination of distances within our Galaxy<sup>9</sup>. As a first approach, we can try to estimate luminosities using the period-luminosity relationship (for those sources with a well determined period). If the P–L relationships for LMC LPVs found by Hughes & Wood (1990) are applied to our sample of OH/IR stars, we obtain  $M_{\text{bol}}$  between  $\sim -5$  and  $-6.2$  for those stars with periods of 400–600 days while an  $M_{\text{bol}}$  of  $\sim -4.6$  is derived for periods around  $\sim 300$  days. However, if we take those stars with the longer periods in the sample ( $\geq 1000$  days), we derive  $M_{\text{bol}} \lesssim -7.9$ , which seems unrealistically high. This suggests that the period-luminosity relationship may not hold for the most extreme OH/IR stars (Wood et al. 1998). On the other hand, there is recent observational evidence for the existence of low metallicity (e.g. in the LMC) HBB AGB stars with luminosities brighter than the predictions of the core-mass

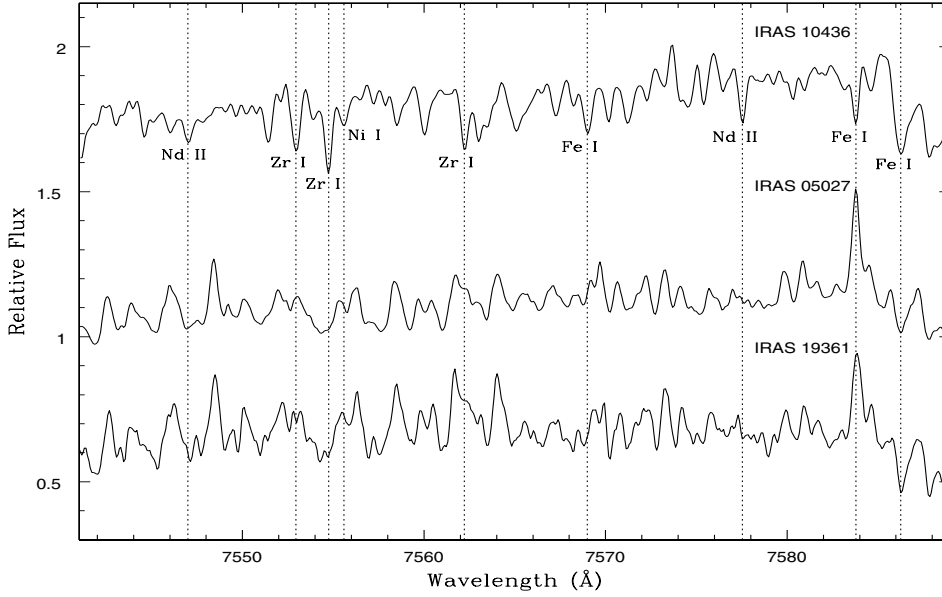
luminosity relation that have been attributed to an excess flux from HBB (Whitelock et al. 2003). A similar effect could explain the detection of AGB stars with such high  $M_{\text{bol}}$  also in our Galaxy.

## 5.2. Progenitor masses

Determining the progenitor masses of our Galactic O-rich AGB stars is not a simple task either. Several observational parameters, however, such as the OH maser expansion velocity or the variability period, have been proposed in the past as useful distance-independent mass indicators for this class of stars, and we will make use of them in the following.

The assumption is based on the fact that the group of OH/IR stars with longer periods and larger expansion velocities shows a Galactic distribution that corresponds to a more massive population (Baud et al. 1981; Baud & Habing 1983; Chen et al. 2001; Jiménez-Esteban et al. 2005b) suggesting that the periods of variability and the expansion velocities of the circumstellar

<sup>9</sup> Note that, in addition, AGB stars are also strongly variable. As an example, Engels et al. (1983) found an  $M_{\text{bol}}$  variation of 2 mag between the minimum and maximum light (from  $-5.0$  to  $-7.1$  mag respectively) in the Galactic OH/IR star OH 32.8–0.3.



**Fig. 10.** High-resolution optical spectra around the 7565 Å region of two stars in our sample where we show the lack of s-process element atomic lines (e.g. from Zr and Nd) in comparison with the Galactic S-star IRAS 10436–3459. IRAS 05027–2158 and IRAS 19361–1658 show lithium lines with very different strength. The positions of some atomic lines of Zr, Nd, and Fe are indicated. It seems that only the Fe I line at  $\sim 7586$  Å is clearly detected in all stars.

**Table 11.** Sensitivity of the derived Li abundances (in dex) to slight changes in the model atmosphere parameters for IRAS 15255+1944.

Adopted value	Change	Li abundance
$T_{\text{eff}} = 2900$ K	$\Delta T_{\text{eff}} = \pm 100$ K	$\Delta \log \epsilon(\text{Li}) = \pm 0.3$
$Z = 0.0$	$\Delta Z = \pm 0.3$	$\Delta \log \epsilon(\text{Li}) = \pm 0.2$
$\xi = 3$ km s $^{-1}$	$\Delta \xi = \pm 1$ km s $^{-1}$	$\Delta \log \epsilon(\text{Li}) = \mp 0.1$
$\log g = -0.5$	$\Delta \log g = +0.5$	$\Delta \log \epsilon(\text{Li}) = +0.1$
$FWHM = 300$ mÅ	$\Delta FWHM = \pm 50$ mÅ	$\Delta \log \epsilon(\text{Li}) = \pm 0.1$

envelopes of AGB stars must be closely correlated with the progenitor mass.

Actually, stars in our sample follow the expected trend, as is illustrated in Fig. 13. Although the limited number of stars considered prevents us from deriving any statistical conclusion based on our data alone, it seems clear that the Galactic latitude distribution of the stars in our sample becomes narrower as a function of  $v_{\text{exp}}(\text{OH})$ , as one would expect if the higher expansion velocities do correspond to a more massive population. A similar trend is observed when the variability period is analysed (not shown).

Figure 14 (left panel) shows the distribution of the sources in our sample on the IRAS two-colour diagram [12] – [25] vs. [25] – [60] for which OH expansion velocities are available. As we can see in this diagram, on average, the sources with intermediate OH expansion velocities, between 6 and 12 km s $^{-1}$ , show redder [12] – [25] and [25] – [60] colours than those with low OH expansion velocities (i.e. below 6 km s $^{-1}$ ), which implies thicker circumstellar envelopes or larger mass loss rates, but they are bluer than the subsample of stars with OH expansion velocities beyond 12 km s $^{-1}$ . Redder IRAS colours can be interpreted as the consequence of a more advanced stage on the AGB if we assume that mass loss rate increases along this evolutionary phase. Alternatively, they may also be interpreted as corresponding to a more massive population of AGB stars since in principle they should be able to develop thicker circumstellar envelopes even at a relatively early stage in their AGB evolution. A

similar behaviour is observed when the variability period is analysed (Fig. 14; right panel), although the number of stars with known periods is significantly smaller. In this case we can see that the reddest sources in the diagram belong to the subgroup showing the longer periods, in excess of 700 days.

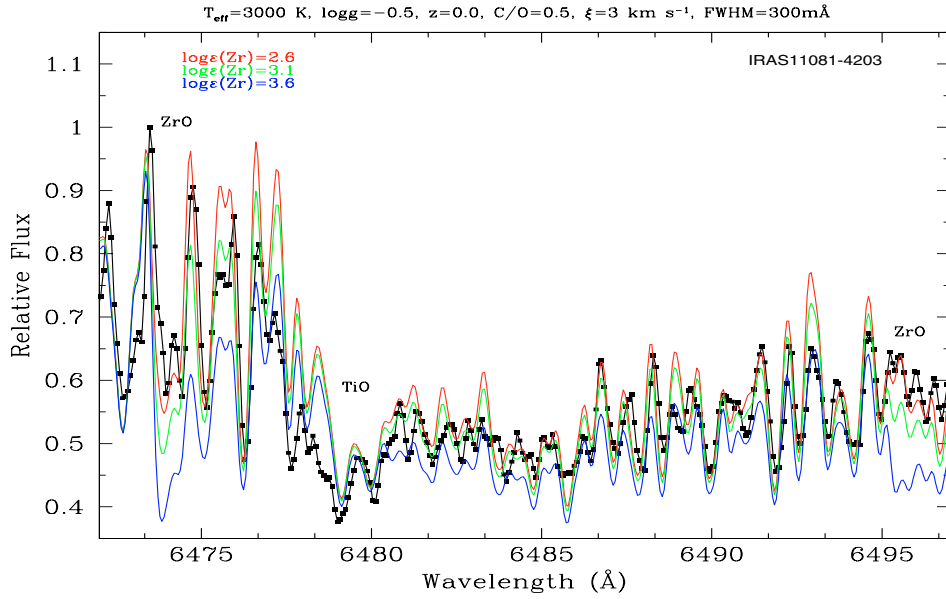
The distribution of variability periods and OH expansion velocities is found to be remarkably different for each of the three subgroups identified in our sample. Histograms showing the results obtained are shown in Figs. 15 and 16, respectively. The same results are presented in a different way in Table 12, where the distribution of sources among the different subtypes is presented as a function of their OH expansion velocity and variability properties.

In Table 12 we can see that most of the AGB stars with the lower OH expansion velocities ( $v_{\text{exp}}(\text{OH}) < 6$  km s $^{-1}$ ) are among those undetected in the Li I line (92%) while 44% of stars with OH expansion velocities between 6 and 12 km s $^{-1}$  (but still observable in the optical) are detected as Li-strong sources. However, there are also stars with non-detection of lithium (24%) and stars too red or not found (32%) within this group. Finally, a majority of the stars in the group of heavily obscured sources (too red or not found) shows OH expansion velocities higher than 12 km s $^{-1}$  (66%).

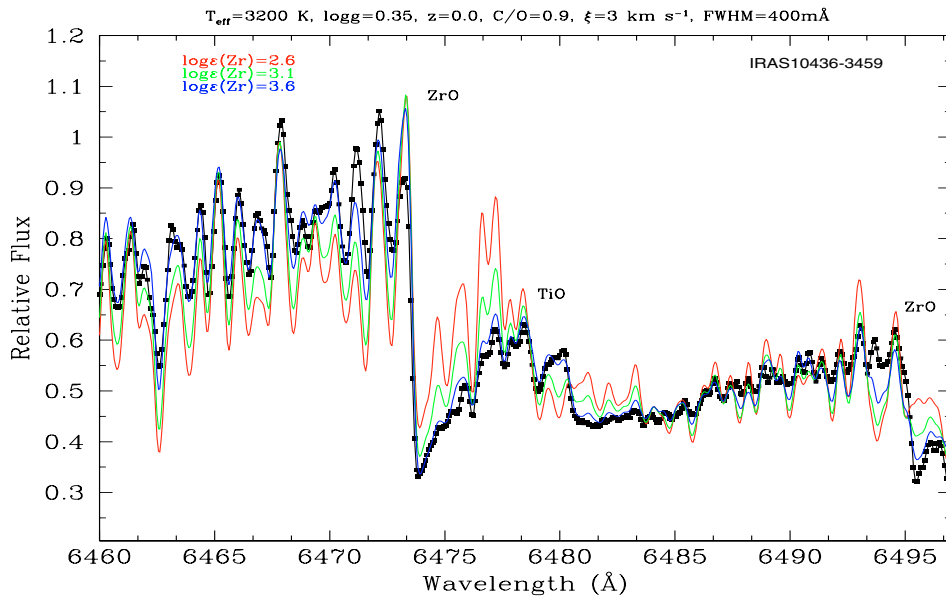
Similar statistics are obtained for the variability period. Sources with relatively short periods, i.e. below 400 days, are predominantly non-Li detections (84%), while almost half of the stars showing periods between 400 and 700 days are Li detected (46%). Again, the same observational problems arise when we try to analyse the stars with the longer periods ( $P > 700$  days), since most of them are completely obscured in the optical (88%).

### 5.3. Period and OH expansion velocity versus HBB

In Fig. 17 we can see that there is actually a clear correlation between the Li abundance and the variability period in our target stars. Remarkably, for periods below 400 days ( $\log P < 2.6$ ), only one star out of 12 is Li detected. In contrast, almost all stars



**Fig. 11.** Synthetic and observed spectra in the region 6460–6499 Å for the star IRAS 11081–4203, one of the Li detected stars in our sample. The synthetic spectra corresponding to  $[Zr/Fe] = +0.0, +0.5,$  and  $+1.0$  dex (or  $\log \epsilon(Zr) = 2.6, 3.1,$  and  $3.6$  dex, respectively) are shown. The parameters of the best model atmosphere fit are indicated in the top label.



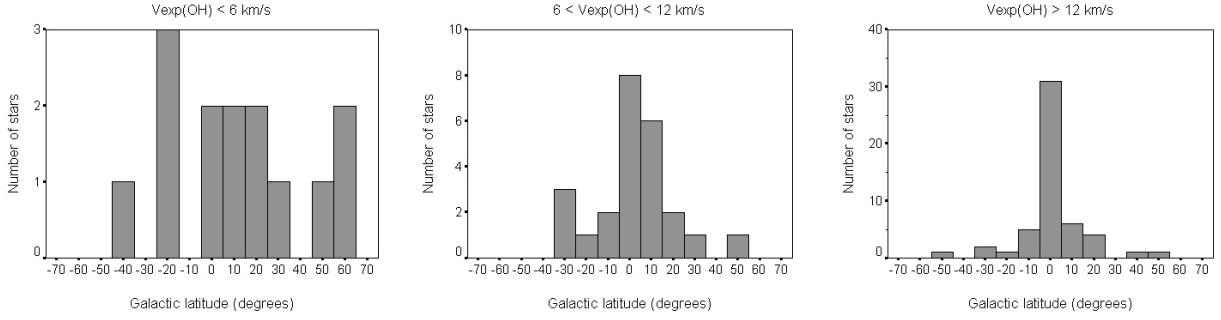
**Fig. 12.** Synthetic and observed spectra in the region 6460–6499 Å for the Galactic S-star IRAS 10436–3454. The synthetic spectra corresponding to  $[Zr/Fe] = +0.0, +0.5,$  and  $+1.0$  dex (or  $\log \epsilon(Zr) = 2.6, 3.1,$  and  $3.6$  dex, respectively) are shown. The parameters of the best model atmosphere fit are indicated in the top label. Note the strong overabundance of  $[Zr/Fe] = +1.0$  dex ( $\log \epsilon(Zr) = 3.6$  dex) needed to fit the observations.

with  $P > 500$  days ( $\log P > 2.7$ ) are Li-rich, the only exception being IRAS 18050–2213 (=VX Sgr), which is a peculiar, very cool semi-regular pulsating star, not showing a proper Mira variability in the last decade, which some authors identify as a massive red supergiant, and not as an AGB star (Kamohara et al. 2005).

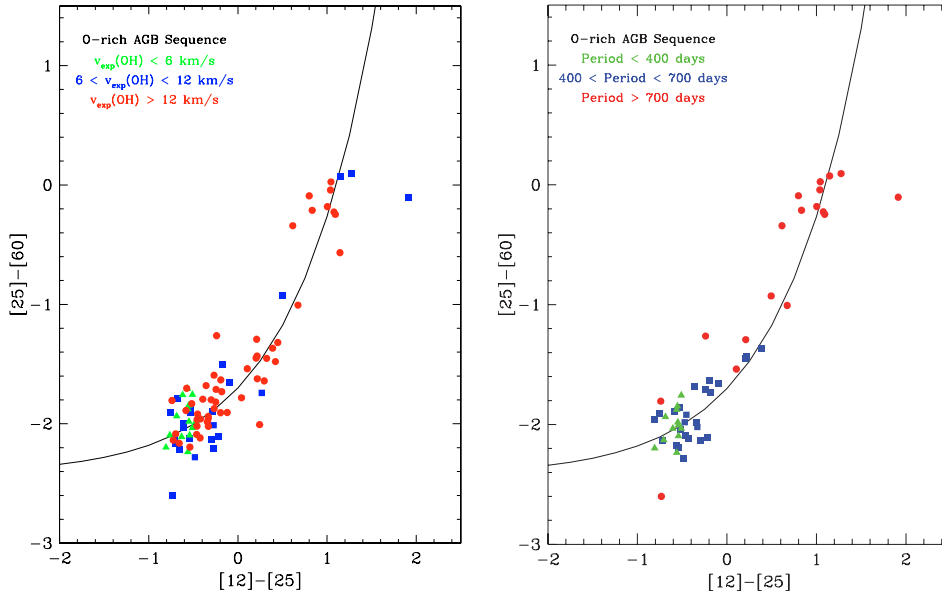
A similar correlation, although not so clear, is found when the expansion velocities derived from the OH maser measurements are considered (see Fig. 18). In this case we can see that the maximum Li abundance increases with the OH expansion velocity, suggesting that the most massive AGB stars can experience a higher lithium enrichment. However, the more interesting

result is the non-detection of Li in any star of the sample with  $v_{\text{exp}}(\text{OH}) < 6 \text{ km s}^{-1}$ . IRAS 18050–2213 is again one of the two outliers showing no Li and a very high expansion velocity, while the other one is IRAS 13442–6109, for still unknown reasons.

At present, it is generally accepted that lithium production in massive O-rich AGB stars is due to the activation of hot bottom burning at the bottom of the convective mantle, which prevents the formation of luminous C-rich AGB stars (see details in Sect. 1). If the detection of lithium is taken as a signature of HBB and the period and the OH expansion velocity are accepted as valid mass indicators, our results indicate that there are no HBB stars among the subgroup of AGB stars in our sample with



**Fig. 13.** Galactic latitude distribution of the sources in our sample as a function of their OH expansion velocities:  $v_{\text{exp}}(\text{OH}) < 6 \text{ km s}^{-1}$  (left panel);  $6 < v_{\text{exp}}(\text{OH}) < 12 \text{ km s}^{-1}$  (middle panel); and  $v_{\text{exp}}(\text{OH}) > 12 \text{ km s}^{-1}$  (right panel).



**Fig. 14.** Distribution of the sources included in our sample in the IRAS two-colour diagram  $[12] - [25]$  vs.  $[25] - [60]$  showing the location of the stars in our sample as a function of the OH expansion velocities (left panel) and variability periods (right panel). The continuum line is the “O-rich AGB sequence” (see text), which indicates the sequence of colours expected for O-rich AGB stars surrounded by envelopes with increasing thickness and/or mass loss rates (Bedijn 1987).

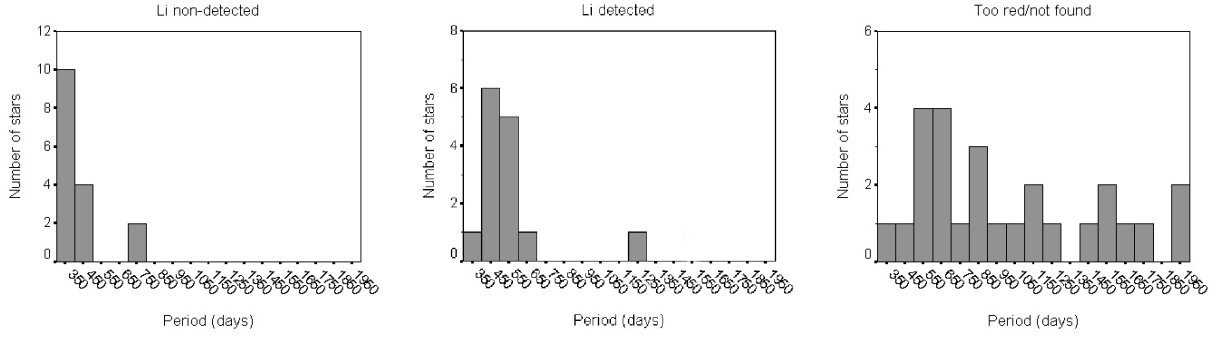
periods lower than 400 days and  $v_{\text{exp}}(\text{OH}) < 6 \text{ km s}^{-1}$ , while most of the AGB stars with periods beyond 500 days and high expansion velocities are quite probably developing HBB.

#### 5.4. Lithium production and HBB model predictions

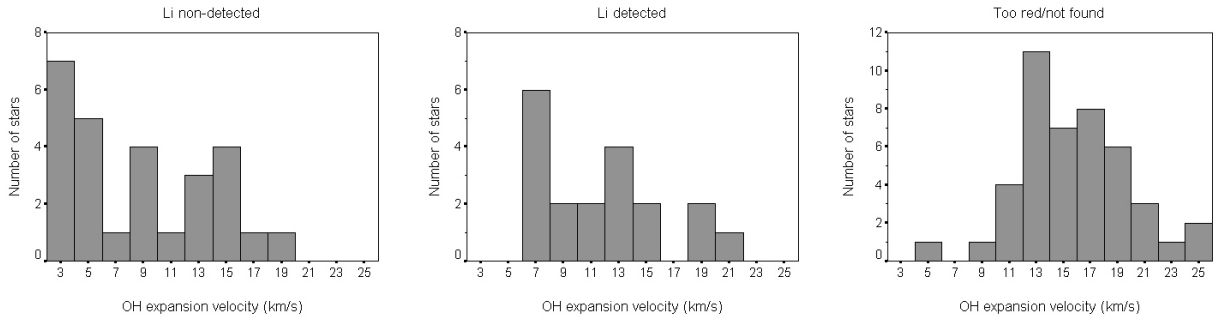
Modelling of HBB started thirty years ago with envelope models including non-instantaneous mixing coupled to the nuclear evolution (Sackmann et al. 1974). More recent models (Sackmann & Boothroyd 1992; Mazzitelli et al. 1999, hereafter SB92 and MDV99, respectively) seem to work well in reproducing the lithium production observed in massive O-rich AGB stars in the MCs. Although the precise determination of the lithium production depends on the input physics (stellar mass, metallicity, mass loss rate, overshooting, etc.), the strongest dependence is with the convection model assumed. SB92 models predict lithium production in intermediate-mass stars assuming the so-called “mixing length theory” (MLT) framework for convection. But these models require a fine tuning of the mixing length parameter  $\alpha$ . MDV99 explored lithium production in intermediate-mass stars according to a more modern treatment of turbulent convection, known as the “Full Spectrum of Turbulence” (FST) model

(D’Antona & Mazzitelli 1996, and references therein). Actually, MDV99 models are able to reproduce the strong Li enhancements observed in massive O-rich LMC AGB stars in a satisfactory way and provide a theoretical explanation to the existence of a few sources in the LMC with extreme luminosities of up to  $M_{\text{bol}} = -7.3$  and  $-7.6$ , which are known to be long period, obscured AGB stars (Ventura et al. 2000). However, these models have not yet been tested against their Galactic analogues.

For solar metallicity, MDV99 predict lithium production as a consequence of the activation of HBB for stellar masses  $\geq 4.5 M_{\odot}$  without core overshooting and for  $M \geq 4 M_{\odot}$  including core overshooting. The amount of lithium produced during the AGB does not depend on the assumed initial lithium abundance as the star structure loses all memory of the previous history of lithium at the beginning of the AGB phase. As an example, according to this model the lithium surface abundance drops to  $\log \epsilon(\text{Li}) \sim 0.3$  for a  $6 M_{\odot}$  star after the second dredge-up, at the beginning of the AGB phase. If overshooting from below the convective envelope is allowed, an important decrease in the number of thermal pulses is predicted by the model and lithium production takes place before the first thermal pulse. The variation with time of the lithium abundance, total luminosity,



**Fig. 15.** Distribution of variability periods for the Li non-detected (*left panel*), Li detected (*middle panel*), and too red/not found (*right panel*) sources.



**Fig. 16.** Distribution of OH expansion velocities for the Li undetected (*left panel*), Li detected (*middle panel*), and too red/not found (*right panel*) sources.

**Table 12.** Distribution of sources among the different subgroups identified in the sample as a function of their OH expansion velocity and variability properties.

Period (days)	Observed	Li undetected	Li detected	Too red/not found
<400	12	10 (84%)	1 (8%)	1 (8%)
400–700	26	4 (15%)	12 (46%)	10 (39%)
>700	17	1 (6%)	1 (6%)	15 (88%)
$v_{\text{exp}}(\text{OH})$ (km s <sup>-1</sup> )	Observed	Li undetected	Li detected	Too red/not found
<6	13	12 (92%)	0 (0%)	1 (8%)
6–12	25	6 (24%)	11 (44%)	8 (32%)
>12	53	10 (19%)	8 (15%)	35 (66%)

and temperature at the base of the convective envelope ( $T_{\text{bce}}$ ) for different masses is shown in Fig. 15 of MDV99. A stronger Li-overabundance, faster increase of the luminosity at the beginning of the AGB phase, and higher  $T_{\text{bce}}$  are predicted when the mass of the star is increased.

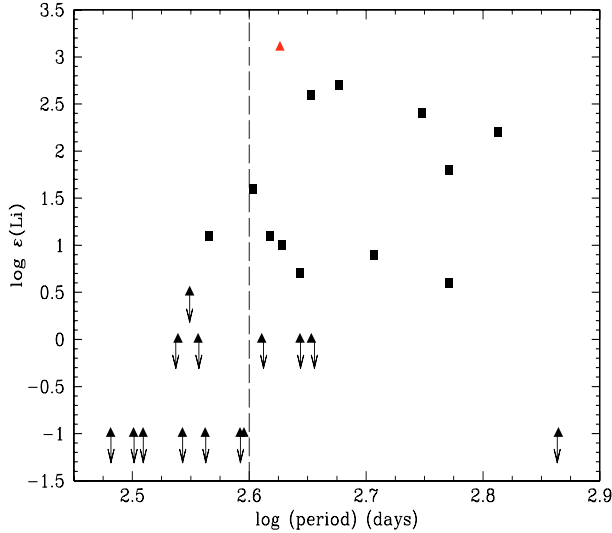
Mass loss rates and metallicities play also an important role according to MDV99. The minimum mass needed for the ignition of HBB decreases with decreasing mass loss rates. In addition, a low mass loss rate leads to a longer run of thermal pulses and to a lithium abundance which remains large for the entire thermal pulsing AGB (TP-AGB, hereafter) lifetime. In contrast, for high mass loss rates, the HBB process is stopped only after a few pulses ( $\sim 5$  for a  $6 M_{\odot}$  star). Models with strong mass loss show oscillations of the Li abundance by orders of magnitude on short timescales of  $\sim 10^4$  years (even within a single thermal pulse!). The temperature at the bottom of the convective envelope,  $T_{\text{bce}}$ , varies so much that the lithium production is temporarily stopped while the lithium already produced has already been diluted by convection. When  $T_{\text{bce}}$  increases, the lithium abundance increases again. The consequence is that, at least for 20% of the time, there is negligible lithium in the envelope (at a level of approximately  $\log \epsilon(\text{Li}) \lesssim 1.0$ ). This implies that the distribution of Li abundances derived from the observations can only be analysed in a statistical way.

On the other hand, the activation of HBB takes place at a lower mass limit of only  $3.0\text{--}3.7 M_{\odot}$  (depending on the mass loss rate considered) at the metallicity of the LMC (Ventura et al. 2000). Low metallicity models predict larger values of  $T_{\text{bce}}$  and higher luminosities, and as a consequence, higher lithium production for a given stellar mass. In addition, this accelerates Li-production and a maximum value of  $\log \epsilon(\text{Li}) \sim 4.3$  is reached in a shorter timescale with respect to the solar metallicity case. The variation with time of the surface lithium abundance, stellar luminosity, and  $T_{\text{bce}}$ , computed for a  $6 M_{\odot}$  model and different metallicities is shown in Fig. 10 of MDV99. Lower metallicity models essentially produce lithium faster due to the larger temperature at the base of the convective envelope, but show high Li abundances for a shorter time.

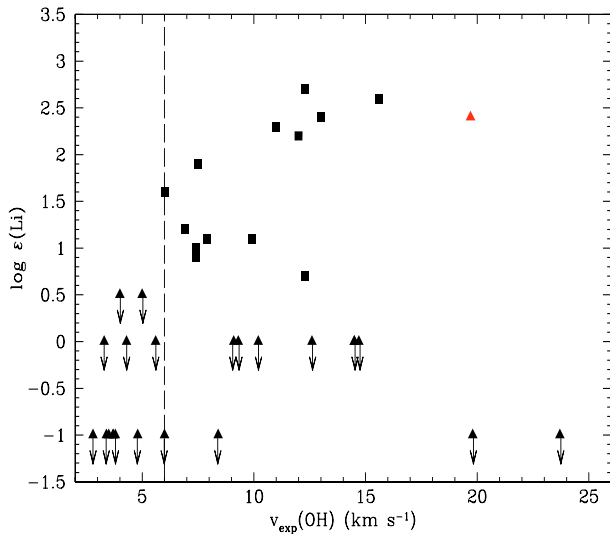
### 5.5. Theory versus observations

Globally considered, our results indicate that below a certain value of the period and the OH expansion velocity our stars do not show any Li enhancement. However, not all the stars above these values show detectable Li. The strong IR excesses associated with most of the stars in our sample suggest that they are actually at the end of their evolution as AGB stars. Surprisingly, none of the stars in our sample shows a strong s-process





**Fig. 17.** Observed Li abundance vs. variability period. Upper limits to the Li abundance are shown with black triangles and marked with arrows. Abundance values derived from Li I lines which are resolved in two components (circumstellar and stellar) are shown with red triangles and correspond to the photospheric abundance needed to fit the stellar component. The dashed vertical line corresponds to a variability period of 400 days (see text).



**Fig. 18.** Observed Li abundance vs. OH expansion velocity. Upper limits to the Li abundance are shown with black triangles and marked with arrows. The abundance values derived from Li I lines which are resolved into two components (circumstellar and stellar) are shown with red triangles and correspond to the photospheric abundance needed to fit the stellar component. The dashed vertical line corresponds to an OH expansion velocity of 6 km s<sup>-1</sup> (see text).

enhancement (e.g. from Zr, Y, Sr) as would have been expected if they were stars that have experienced a significant third dredge-up during their AGB evolution. In the case of the more massive stars in our sample one could argue that the mass loss is so strong at the high metallicity of our Galaxy (compared to the MC metallicity) that their AGB lifetime may become shortened significantly, so that they will end this phase experiencing too few thermal pulses before they completely lose their envelope. However, the lack of s-process enhancement is also observed in the group of stars in which no lithium has been detected.

In order to interpret these results, several scenarios can be proposed:

- i) all the stars in the sample are HBB stars, and thus, considerably massive ( $M \gtrsim 4 M_{\odot}$ ). The difference between the Li detected and the Li undetected ones would be just that the undetected ones are in a very early stage as AGB stars. This hypothesis would explain the absence of s-process enhancement (e.g. Zr, Y, Sr) in all the stars in the sample, but it would not be consistent with the non-detection of lithium in the group of stars with the shorter periods and lower OH expansion velocities. HBB models predict a considerable Li enhancement just after the arrival at the AGB. Statistically, there should be no difference in the behaviour of the Li abundances for any star in the sample, and this is not observed.
- ii) the sample is composed by a combination of low-mass ( $M \lesssim 1.5 M_{\odot}$ ) and high-mass ( $M \gtrsim 4 M_{\odot}$ ) stars. In this scenario, the optically bright O-rich AGB stars in our sample with the shorter periods and lower OH expansion velocities, undetected in Li, would belong to the population of low-mass stars in the Galaxy that do not experience HBB and evolve all the way along the AGB as O-rich stars. However, these stars are expected to evolve very slowly and develop a large number of thermal pulses during the AGB. At the end of their AGB evolution they are expected to become strongly enhanced in s-process elements as a consequence of the dredge-up of processed material induced by the many thermal pulses experienced. But the Li undetected stars in our sample do not show this enhancement in s-process elements. The only way to reconcile this result with their status as low-mass stars would imply assuming that they are at a very early stage as AGB stars, which again seems to be incompatible with the strong IR excess detected by IRAS. Such a strong excess is only expected in low-mass stars at the very end of the AGB. Alternatively, they could experience less thermal pulses than predicted by the models or the dredge-up process could be very inefficient. But this is not what we observe in other low- and intermediate-mass stars in the Galaxy, such as the S-stars and many C-rich stars, which show the strong s-process enhancement that is not observed in our stars. Therefore, we must conclude that the hypothesis of these Li undetected stars being identified as low-mass AGB stars seems to be implausible.
- iii) all the stars in the sample are intermediate-mass stars ( $M \gtrsim 3 M_{\odot}$ ) but only a fraction of them, the more massive ones ( $M \gtrsim 4 M_{\odot}$ ), experience HBB. This scenario would explain why the sources with lower periods and OH expansion velocities in our sample are undetected in the Li I line, while at the same time, they do not show any s-process enhancement. These stars, although slightly below the minimum mass needed to develop HBB, would also experience a very strong mass loss rate and only a few thermal pulses ( $\sim 5$ – $10$ ) before the early termination of the TP-AGB as a consequence of the strong mass loss. The initial envelope of these stars would also be quite massive and thus a slow change of the C/O ratio is expected because of the huge dilution of the dredged up material. But how can we explain the large spread of Li abundances observed in the group of sources in our sample with longer periods and larger OH expansion velocities? These are all supposed to be HBB stars. And HBB models predict a large fluctuation in the Li abundance as the star evolves along the TP-AGB (see Fig. 15 of MDV99 again) between two consecutive pulses. Actually, the timescale of the Li production phase in the TP-AGB is

predicted to be of the order or slightly smaller than the timescale between pulses ( $\sim 10^4$  years for a massive star with  $M \gtrsim 4 M_{\odot}$ ). In addition, as the AGB evolution proceeds, the surface Li abundances decrease again very rapidly as a consequence of the  ${}^3\text{He}$  becoming almost completely burned in the envelope, and thus, the Li production ceases (e.g. MDV99; Forestini & Charbonnel 1997). This would be consistent with the detection of strong Li abundances in  $\sim 50\%$  of the stars in our sample for which we suspect that HBB is active. In this scenario, the strongly obscured sources for which we were not able to obtain an optical spectrum could also be interpreted as a population of high mass stars developing HBB with even more extreme observational properties. They are sources with longer periods ( $>700$  days) and larger OH expansion velocities ( $>12 \text{ km s}^{-1}$ ) which may have developed optically thick circumstellar envelopes as a consequence of the huge mass loss experienced in the AGB. These stars must also be detectable as Li-rich at least for some time during their AGB evolution but, unfortunately, they are so strongly reddened in the optical range that their study is only possible in the infrared, where no strong Li lines are present. We propose that the HBB nature of these obscured stars can be elucidated by determining their  ${}^{12}\text{C}/{}^{13}\text{C}$  ratios, which can also be used as an HBB indicator, through near infrared spectroscopy.

### 5.6. Comparison with the O-rich AGB stars in the MCs

The situation is quite different for the AGB stars in the MCs. In contrast with their Galactic analogues, the more luminous AGB stars in the MCs (with  $-7 \leq M_{\text{bol}} \leq -6$ ) are O-rich stars showing s-process element enhancements. In addition, a higher proportion of them ( $\sim 80\%$  compared to  $\sim 50\%$  in our Galaxy) also shows lithium overabundances. The enrichment in lithium indicates that they are also HBB stars, but why are these stars also enriched in s-process elements?

The answer to this question must be related to the different metallicity of the stars in the MCs with respect to our Galaxy. Actually, theoretical models predict a higher efficiency of the dredge-up in low metallicity atmospheres (e.g. Herwig 2004) with respect to those with solar metallicity (e.g. Lugaro et al. 2003).

Moreover, there is observational evidence that lower metallicity environments are also less favourable to dust production, as is suggested by the smaller number of heavily obscured AGB stars in the MCs (Trams et al. 1999; Groenewegen et al. 2000), compared to our Galaxy. This is supported by the lower dust-to-gas ratios derived by van Loon (2000) in the few obscured AGB stars in the MCs for which this analysis has been made.

If mass loss is driven by radiation pressure on the dust grains, it might be less efficient with decreasing metallicity (Willson 2000). In that case, longer AGB lifetimes would be expected, which could increase the chance of nuclear-processed material reaching the stellar surface. This would also explain the larger proportion of luminous C-rich stars (up to  $M_{\text{bol}} \sim -6$ ) observed in the MCs (Plez et al. 1993; Smith & Lambert 1989, 1990a; Smith et al. 1995) with respect to the Galaxy. The slow evolution predicted for AGB stars in the MCs as a consequence of the less efficient mass loss leaves time for more thermal pulses to occur during the AGB lifetime and, therefore, a more effective dredge-up of s-process elements to the surface can be expected before the envelope is completely gone at the end of the AGB. This would explain why even the more massive stars in the MCs show

a strong s-process enrichment in contrast to their Galactic counterparts. In our Galaxy the only AGB stars showing a similar overabundance in s-process elements seem to be the result of the evolution of low- to intermediate-mass stars ( $M \lesssim 1.5\text{--}2.0 M_{\odot}$ ), while no, or very little, s-process enhancement is observed in Galactic AGB stars with higher main sequence masses.

In addition, the lower critical mass (which also decreases with decreasing mass loss rate, see Sect. 5.4) needed to develop HBB ( $M \gtrsim 3 M_{\odot}$  at the metallicity of the LMC, compared to the  $\sim 4 M_{\odot}$  limit in our Galaxy) favours the simultaneous detection of s-process elements and Li enrichment in a larger number of AGB stars in the MCs. In our Galaxy, the Li-rich sample of AGB stars is restricted to the fraction of stars with main sequence masses  $M \gtrsim 4 M_{\odot}$ . In contrast to their MC counterparts, these stars evolve so rapidly (because of the strong mass loss) that there is no time for a significant enhancement in s-process elements, as the results here presented seems to demonstrate.

Obviously, more refined s-process calculations taking into account the effects of HBB and mass loss, using a large grid of stellar evolutionary models with different masses and metallicities are strongly encouraged.

### 5.7. Evolutionary connections

The well known existence of at least two different chemical branches of AGB stars (C-rich and O-rich) and the dependence of this chemical segregation on the main sequence mass and on the environmental conditions in which these stars evolve (such as the metallicity) is a long debated issue in stellar evolution. Moreover, the connection with the chemical composition observed in Galactic post-AGB stars and/or well evolved PNe (which are expected to be the result of the evolution of AGB stars with a wide variety of masses) has never been addressed in detail from an overall perspective.

Stars classified as post-AGB in the literature are mainly low-mass stars, according to their wide Galactic distribution. They are thought to represent the fraction of stars which evolve so slowly that they are still visible for some time as stars with intermediate spectral types before they become PNe. Many of these stars are C-rich and show strong s-process element enhancements [ $s\text{-process}/\text{Fe}$ ]  $\sim +1.5$  (van Winckel & Reyniers 2000; Reddy et al. 1999; Klochkova et al. 1999), in agreement with what we would expect if they were the final product of the evolution of low-mass AGB stars. However, other post-AGB stars do not show the characteristic signatures of the third dredge-up. In the non-enriched sources the s-process element abundances can be as low as  $[\text{Zr}/\text{Fe}] \lesssim 0.2$  (Luck et al. 1990; van Winckel 1997). These objects might be stars with even lower masses ( $M \sim 0.8\text{--}1.0 M_{\odot}$ ) which have evolved off the AGB before the third dredge-up could occur (Marigo et al. 1999). Some of these sources are high Galactic latitude, hot B-type post-AGB stars (e.g. Conlon et al. 1993; Mooney et al. 2001) that show strong carbon deficiencies. In the more extreme cases they will probably never become PNe.

On the other hand, and as we have shown in this article, the more massive AGB stars may also show a similar lack of s-process elements, but in this case as a consequence of the rapid evolution induced by the strong mass loss. These stars are more difficult to observe in the optical as they are expected to evolve as heavily obscured sources from the TP-AGB to the PN stage but may be the progenitors of most Galactic PNe.

Actually, a similar chemical segregation is also observed in Galactic PNe. Classically, they are classified as type I, II or III

as a function of their chemical abundance properties (Peimbert 1978). They are expected to cover a wide range of progenitor masses and thus show different chemical signatures as a consequence of their previous passage along the AGB.

In particular, the higher mass fraction of stars (with  $M > 4 M_{\odot}$ ) identified in this paper as stars developing HBB and showing a strong lithium enhancement are expected to arrive at the PN stage still as heavily obscured stars. These stars would evolve into N-rich type I PNe, which are characterized by their large He abundances ( $\text{He}/\text{H} > 0.10$ ) and N/O ratios (Manchado 2003, 2004; Mampaso 2004, and references therein). They are thought to represent the fraction of the more massive PNe in the Galaxy, as is confirmed by their strong concentration at low Galactic latitudes. Their strong N overabundances would be consistent with their identification as PNe resulting from the evolution of HBB AGB stars.

## 6. Conclusions

As a result of a high-resolution spectroscopic survey carried out in the optical for a large sample (102) of massive Galactic O-rich AGB stars, we have detected the Li I resonance line at 6708 Å in 25 stars, while 32 stars did not display this Li I line. The rest of the stars observed (45 stars) were heavily obscured by their thick circumstellar envelopes and, as such, they were too red or were not found at optical wavelengths.

We have classified the sample into three subgroups on the basis of their OH expansion velocity and variability period. From their relative distribution in the IRAS two-colour diagram [12] – [25] vs. [25] – [60] and Galactic distribution, we conclude that they must represent populations of Galactic O-rich AGB stars with different progenitor masses.

By combining MARCS model atmospheres and synthetic spectroscopy with extensive line lists we were able to derive the fundamental stellar parameters ( $T_{\text{eff}}$ ,  $\log g$ ,  $\text{C}/\text{O}$ ,  $\xi$ , etc.) and obtain the Li and Zr abundances in those stars for which an optical spectrum had been obtained. Our chemical abundance analysis shows that half of stars show Li overabundances in the range between  $\log \varepsilon(\text{Li}) \sim +0.5$  and  $+3.0$  dex. This is interpreted as a signature of the activation of HBB, confirming that they are actually massive AGB stars ( $M > 4 M_{\odot}$ ). However, these stars do not show any zirconium enhancement (taken as representative of the s-process element enrichment), indicating that they behave differently from the more massive (and luminous) AGB stars in the MCs.

Assuming that the variability period and the OH expansion velocity can be taken as distance-independent mass indicators and by comparing our results with the theoretical predictions of HBB and nucleosynthesis models, we conclude that the O-rich AGB stars in our sample are all very massive ( $M > 3 M_{\odot}$ ), even those non-detected in lithium; but only the more massive ones ( $M > 4 M_{\odot}$ ; with periods longer than 400 days and  $v_{\text{exp}}(\text{OH}) > 6 \text{ km s}^{-1}$ ) experience HBB. The lack of lithium enrichment found in some of the sources thought to belong to the group of more massive AGB stars in the sample is explained as (i) the consequence of the timescale of the Li production phase being of the order of or slightly smaller than the interpulse time ( $\sim 10^4$  years) or, alternatively, (ii) by the  $^3\text{He}$  exhaustion in the envelope interrupting the Li production, in agreement with the predictions made by the models.

The group of stars displaying the most extreme observational properties (with periods sometimes longer than 700 days and  $v_{\text{exp}}(\text{OH}) > 12 \text{ km s}^{-1}$ ) are believed to represent the more massive AGB stars in our Galaxy. Unfortunately, these stars are

strongly obscured by their thick circumstellar envelopes so we could not carry out any type of analysis in the optical range. This very high-mass population should also experience HBB and show strong Li enhancement but no s-process element overabundances. We propose that the HBB status of these obscured stars can be determined by measuring their  $^{12}\text{C}/^{13}\text{C}$  ratios in the near-infrared, which can also be used as an HBB indicator.

Our results suggest that the dramatically different abundance pattern found in AGB stars belonging to the MCs and to our Galaxy can be explained in terms of the different metallicity conditions under which these stars evolved. This constitutes strong observational evidence that the chemical evolution of massive AGB stars is strongly modulated by metallicity.

Globally considered, our results can also be used to establish evolutionary connections between TP-AGB stars, post-AGB stars, and PNe. In particular, we suggest that the HBB AGB stars identified in our sample must be the precursors of N-rich type I PNe.

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# Online Material

**Table 1.** The sample of Galactic O-rich AGB stars.

IRAS name	Other names	Galactic Coordinates		Variability	Spectral Type	[12]–[25]	[25]–[60]
01037+1219	WX Psc	128.64	–50.11	Mira	M10	–0.192	–1.632
01085+3022	AW Psc	127.96	–32.04	Mira	M9	–0.331	–2.021
01304+6211	V669 Cas	127.81	–0.02	Mira	...	+0.494	–0.927
02095–2355	IRC–20029	206.68	–71.55	...	M6	–0.748	–1.992
02316+6455	V656 Cas	133.61	+4.46	Mira	M8	–0.463	–2.089
03507+1115	IK Tau	177.95	–31.41	Mira	M9e	–0.725	–2.137
04404–7427	SY Men	286.97	–34.74	Mira	...	–0.297	–2.132
05027–2158	T Lep	222.67	–32.71	Mira	M7e	–0.706	–2.129
05073+5248	NV Aur	156.44	+7.83	Mira	M10	+0.204	–1.450
05098–6422	U Dor	274.30	–35.16	Mira	M7e	–0.510	–1.756
05151+6312	BW Cam	148.28	+14.57	Mira	M9	–0.698	–2.084
05559+3825	V373 Aur	173.24	+7.27	Mira	M9–10	–0.563	–2.174
06297+4045	IRC+40156	174.11	+14.12	...	M8	–0.094	–1.658
06300+6058	AP Lyn	154.31	+21.52	Mira	M7+	–0.357	–1.680
07080–5948	NSV 3436	270.50	–21.14	...	...	–0.631	–2.109
07222–2005		234.58	–2.16	...	M9	–0.733	–2.600
07304–2032	Z Pup	235.90	–0.69	Mira	M5e	–0.752	–1.908
07445–2613	SS Pup	242.43	–0.69	Mira	M6e	–0.807	–2.197
09194–4518	MQ Vel	269.09	+3.10	Mira	...	–0.273	–2.012
09429–2148	IW Hya	255.80	+23.35	Mira	M9	–0.217	–2.106
10189–3432	V Ant	271.03	+18.61	Mira	M7e	–0.561	–2.235
10261–5055	VZ Vel	281.31	+5.57	SRa	M6e	–0.542	–2.096
11081–4203		283.71	+16.75	...	Me	–0.699	–1.895
11525–5057		294.00	+10.65	...	...	–0.299	–1.964
12377–6102		301.63	+1.52	Mira	...	–0.195	–1.907
12384–4536		301.08	+16.94	...	...	–0.120	–1.905
13203–5536		307.45	+6.71	...	...	–0.249	–1.818
13328–6244		308.00	–0.56	...	...	+0.295	–1.641
13341–6246		308.14	–0.63	L	...	–0.300	–1.799
13379–5426		310.14	+7.48	...	...	–0.290	–1.893
13442–6109	OH 309.6+0.7	309.62	+0.73	...	M10	–0.337	–1.774
13475–4531	V618 Cen	313.60	+15.88	Mira	Me	–0.764	–2.096
13517–6515		309.51	–3.46	...	...	+0.040	–1.782
14086–0730	IO Vir	334.78	+50.12	Mira	M8	–0.333	–1.938
14086–6907		310.04	–7.61	...	...	–0.333	–1.955
14247+0454	RS Vir	352.67	+57.97	Mira	M5e	–0.553	–1.847
14266–4211		321.67	+16.84	...	...	–0.547	–2.121
14337–6215		314.87	–2.09	...	...	–0.518	–1.831
15099–5509	RAFGL 4212	322.26	+2.10	...	...	–0.267	–1.873
15193+3132	V S CrB	49.47	+57.17	Mira	M6e	–0.510	–2.034
15211–4254		330.43	+11.47	...	...	–0.609	–1.997
15255+1944	WX Ser	29.51	+53.48	Mira	M8e	–0.483	–2.280
15303–5456		324.83	+0.69	...	...	+0.421	–1.479
15576–1212	FS Lib	358.44	+29.52	Mira	M8–9	–0.535	–1.860
15586–3838	NSV 7388	338.79	+10.48	...	...	–0.672	–1.785
16030–5156	V352 Nor	330.50	+0.01	Mira	M4e	–0.438	–1.914
16037+4218	V1012 Her	66.88	+48.00	...	M8	–0.549	–1.981
16260+3454	V 697 Her	56.37	+43.53	Mira	M9	–0.245	–1.712
16503+0529	RX Oph	23.69	+28.78	Mira	M5	–0.687	–1.936
17034–1024	V850 Oph	10.68	+17.76	Mira	M8	–0.543	–2.007
17103–0559		15.57	+18.70	...	...	–0.539	–2.195
17239–2812		352.61	–0.18	...	...	–0.268	–1.593

**Table 1.** continued.

IRAS name	Other names	Galactic Coordinates		Variability	Spectral Type	[12]–[25]	[25]–[60]
17359–2138		5.40	+5.09	...	M7	–0.705	–2.163
17433–1750	GLMP 637	9.56	+5.61	...	...	+1.144	–0.567
17433–2523		3.10	+1.68	...	...	–0.421	–1.960
17443–2519	RAFGL 5386	3.29	+1.52	...	...	+0.322	–1.452
17501–2656	V4201 Sgr	2.58	–0.43	SR	M8–9	–0.239	–1.262
18025–2113	IRC–20427	8.93	–0.01	...	cM4	–0.167	–1.503
18050–2213	VX Sgr	8.34	–1.00	SRc	M4eIa	–0.740	–1.805
18057–2616	SAO 186357	4.89	–3.18	...	M	–0.575	–1.703
18071–1727	OH 12.8+0.9	12.75	+0.89	...	...	+1.275	+0.094
18083–2630	NSV 24316	4.97	–3.73	...	M6	+0.243	–2.007
18172–2305		8.95	–3.89	...	...	–0.459	–2.020
18198–1249	OH 18.3+0.4	18.30	+0.43	...	...	+1.092	–0.246
18257–1000	V441 Sct	21.46	+0.49	Mira	...	+1.038	–0.042
18273–0738	V436 Sct	23.73	+1.23	Mira	...	+0.105	–1.538
18276–1431	V445 Sct	17.68	–2.03	Mira	...	+1.914	–0.103
18304–0728	IRC–10434	24.21	+0.63	...	M5r	–0.394	–1.794
18312–1209		20.18	–1.70	...	...	+0.447	–1.319
18314–1131	NSV 24500	20.76	–1.46	...	M6	–0.533	–1.750
18348–0526	V437 Sct	26.54	+0.62	Mira	...	+0.615	–0.341
18413+1354	V837 Her	44.56	+8.02	Mira	M8III	–0.424	–2.118
18429–1721	V3952 Sgr	16.84	–6.60	Mira	M9	–0.527	–1.907
18432–0149	V1360Aql	30.71	+0.42	Mira	...	+0.799	–0.091
18437–0643	V440 Sct	26.42	–1.93	Mira	...	+0.389	–1.367
18454–1226	IRC–10463	21.53	–4.92	...	M6	–0.685	–2.306
18460–0254	V1362 Aql	30.09	–0.68	Mira	...	+1.003	–0.181
18488–0107	V1363 Aql	31.99	–0.49	Mira	...	+1.043	+0.026
18549+0208	OH 35.6–0.3	35.57	–0.34	...	...	+0.833	–0.212
18560+0638	V1366 Aql	39.71	+1.50	Mira	...	+0.210	–1.292
19059–2219	V3880 Sgr	14.66	–13.61	Mira	M8	–0.343	–1.984
19129+2803		60.67	+7.74	...	...	–0.521	–2.045
19147+5004	TZ Cyg	81.22	+16.80	Lb	M6	–0.659	–2.163
19157–1706	RAFGL 2361	20.53	–13.55	...	...	–0.458	–1.965
19161+2343	RAFGL 2362	57.12	+5.12	...	...	+0.217	–1.623
19192+0922	NSV 24761	44.79	–2.31	...	...	+0.215	–1.433
19254+1631	GLMP 915	51.80	–0.22	...	...	+1.077	–0.226
19361–1658	NSV 24833	22.74	–17.93	...	M8	–0.268	–2.207
19412+0337	V1415 Aql	42.34	–9.86	Mira	M9	–0.583	–1.889
19426+4342		77.48	+9.73	...	...	–0.659	–2.216
20052+0554	V1416 Aql	47.37	–13.95	Mira	M9	–0.453	–1.919
20077–0625	V1300 Aql	36.36	–20.42	Mira	M	–0.182	–1.731
20109+3205	V557 Cyg	70.47	–1.08	Mira	M7	–0.562	–1.876
20181+2234	V371 Vul	63.44	–7.72	Mira	...	+0.268	–1.741
20272+3535	GLMP 999	75.27	–1.84	...	...	+1.147	+0.074
20343–3020	RT Mic	13.59	–34.91	Mira	...	–0.605	–2.033
20350+3741	V1828 Cyg	77.90	–1.84	Mira	M7	–0.814	–1.961
22177+5936	NSV 25875	104.91	+2.41	...	...	+0.673	–1.006
22180+3225	YZ Peg	89.60	–20.24	...	...	–0.620	–1.758
22466+6942		112.69	+9.59	...	...	–0.472	–1.981
22510+3614	NSV 14347	97.90	–20.56	...	M7	–0.567	–1.705
23492+0846	DO 8089	99.10	–51.03	...	M5	–0.716	–1.900

**Table 2.** The sample of comparison stars.

IRAS name	Other names	Galactic Coordinates		Variability	Spectral Type	[12]–[25]	[25]–[60]
00213+3817	R And	117.07	–23.98	Mira	S5/4.5e	–0.723	–2.104
02192+5821	S Per	134.62	–2.19	SRc	M4eIa	–0.407	–1.897
10436–3459	Z Ant	275.86	+20.94	SR	S5.4	–0.830	–2.172
18397+1738	V821 Her	47.78	+10.01	Mira	Ce	–0.873	–1.499
19133–1703	T Sgr	20.32	–13.01	Mira	S5/6e	–1.125	–1.372
20166+3717	WX Cyg	75.43	+0.87	Mira	C9,2eJ	–1.191	+1.445
23416+6130	PZ Cas	115.06	–0.05	SRc	M3vIa	+0.070	–1.538
23554+5612	WY Cas	115.57	–5.62	Mira	S6/6e	–0.633	–1.286
23587+6004	WZ Cas	116.76	–1.91	SRc	SC7/10e	–1.280	–1.348

**Table 3.** The sample of peculiar carbon AGB stars.

IRAS name	Other names	Galactic Coordinates		Variability	Spectral Type	[12]–[25]	[25]–[60]
00247+6922	V668 Cas	120.86	+6.87	Mira	C	–0.774	–1.762
04130+3918	RAFGL 6312	160.41	–8.09	...	C	–1.273	–1.633
09425–6040	GLMP 260	282.04	–5.88	...	C	+0.790	–1.045
17297+1747	V833 Her	40.80	+25.30	Mira	C	–0.340	–1.859
19321+2757	V1965 Cyg	62.57	+3.96	Mira	C	–0.703	–1.598
19594+4047	V1968 Cyg	76.52	+5.59	Mira	C	–0.285	–1.594
20072+3116	V1969 Cyg	69.35	–0.87	Mira	C	–0.644	–1.483
23166+1655	LL Peg	93.53	–40.35	Mira	C	+0.101	–1.238
23320+4316	LP And	108.46	–17.15	Mira	C	–0.777	–1.555

**Table 4.** Log of the spectroscopic observations.

Run	Telescope	Instrumentation	Dates	Dispersion (Å/pix)	Spectral Range (Å)
#1	4.2 m WHT	UES	16–17 August 1996	0.065	5300–9400
#2	4.2 m WHT	UES	10–12 June 1997	0.065	4700–10300
#3	4.2 m WHT	UES	18–20 August 1997	0.065	4700–10300
#4	3.6 m ESO	CASPEC	21–24 February 1997	0.085	6000–8200



**Table 5.** Galactic O-rich AGB stars where the lithium line was detected.

IRAS name	Run	$v_{\text{exp}}(\text{OH})$ ( $\text{km s}^{-1}$ )	Ref.	Period (days)	Ref.
01085+3022	3	13.0	3	560	1
02095–2355	1	...	...	...	...
04404–7427*	4	7.7	1	534	2
05027–2158	4	7.9	1	368	1
05559+3825	3	single peak	4	590	3
06300+6058	3	12.3	2	440	4
07222–2005*	4	8.2	1	1200	10
07304–2032	4	7.4	5	509	1
09429–2148	4	12.0	2	650	1
11081–4203	4	single peak	1	...	...
11525–5057	4	single peak	1	...	...
12377–6102*	4	20.4	1	...	...
14337–6215	4	19.7	9	...	...
15211–4254	4	11.0	1	...	...
15255+1944	1	7.4	2	425	1
15576–1212	4	9.9	2	415	1
16030–5156	4	single peak	1	...	...
16260+3454	1	12.3	2	475	1
18025–2113	1	10.6	1	...	...
18057–2616	1	18.7	1	...	...
18413+1354	2	14.7	4	590	3
18429–1721	1	6.9	1	...	...
19129+2803	2	single peak	1	420	10
19361–1658	2,3	7.5	1	...	...
20052+0554	2	15.6	12	450	3

\* Tentative Li detection.

References for the OH expansion velocities:

1) te Lintel Hekkert et al. (1991); 2) te Lintel Hekkert et al. (1989); 3) Chengalur et al. (1993); 4) Sloatmaker et al. (1985); 5) Sivagnanam et al. (1989); 6) Lewis et al. (1990); 7) Lewis et al. (1995); 8) Groenewegen et al. (2002) (from CO data); 9) Sevenster et al. (1997); 10) Groenewegen & de Jong (1998); 11) Nyman et al. (1992); 12) Lewis (1994).

References for the periods:

1) Combined General Catalogue of Variable Stars (GCVS), Kholopov (1998); 2) Whitelock et al. (1994); 3) Jones et al. (1990); 4) Lockwood (1985); 5) Engels et al. (1983); 6) Nakashima et al. (2000); 7) Sloatmaker et al. (1985); 8) Le Bertre (1993); 9) Herman & Habing (1985); 10) Jiménez-Esteban et al. (2006b).

**Table 6.** Galactic O-rich AGB stars where the lithium line was not detected.

IRAS name	Run	$v_{\text{exp}}(\text{OH})$ ( $\text{km s}^{-1}$ )	Ref. <sup>1</sup>	Period (days)	Ref. <sup>1</sup>
03507+1115	3	16.5	2	470	1
05098–6422	4	6.0	2	394	2
05151+6312	1	14.5	2	...	3
07080–5948	4	5.0	1	...	...
07445–2613	4	4.8	2	391	1
10189–3432	4	3.4	2	303	1
10261–5055	4	3.7	1	317	1
13379–5426	4	10.2	1	...	...
13442–6109**	4	23.7	1	...	...
13475–4531	4	3.7	1	...	...
14086–0730*	4	13.7	2	...	...
14086–6907*	4	13.4	1	...	...
14247+0454	4	4.0	2	354	1
14266–4211	4	9.1	1	...	...
15193+3132	1	3.3	2	360	1
15586–3838	4	9.3	1	...	...
16037+4218	1	3.5	1	360	10
16503+0529	2	3.8	3	323	1
17034–1024	2	single peak	1	346	1
17359–2138*	2	8.5	1	...	...
18050–2213	3	19.8	2	732	1
18304–0728	3	15.3	2	...	...
18454–1226**	3	single peak	1	...	...
19147+5004	1	12.6	1	...	...
19157–1706	2	14.7	1	...	...
19412+0337	2	14.5	6	440	3
20343–3020	3	8.4	1	349	2
20350+3741	3	...	...	450	3
22180+3225	1	4.3	3	...	...
22466+6942	1	5.6	1	408	6
22510+3614	1	2.8	3	...	...
23492+0846**	1	...	...	...	...

\* Poor S/N ratio.

\*\* Possibly not an AGB star.

<sup>1</sup> See applicable references in Table 5.

**Table 7.** Galactic O-rich AGB stars for which no analysis of the lithium line was possible.

IRAS name	Run	$v_{\text{exp}}(\text{OH})$ ( $\text{km s}^{-1}$ )	Ref. <sup>1</sup>	Period (days)	Ref. <sup>1</sup>	Notes
01037+1219	1,3	18.2	2	660	2	R
01304+6211	3	11.0	2	1994	7	...
02316+6455	3	15.5	7	534	3	R
05073+5248	3	16.9	2	635	1	R
06297+4045	3	12.0	2	520	10	R
09194-4518	4	11.3	1	...	...	R
12384-4536	4	13.2	1	...	...	R
13203-5536	4	12.7	1	...	...	...
13328-6244	4	25.1	1	...	...	...
13341-6246	4	16.7	1	...	...	R
13517-6515	4	17.9	1	...	...	...
15099-5509	4	18.8	1	...	...	R
15303-5456	4	24.0	1	...	...	...
17103-0559	2,3	15.3	1	420	10	...
17239-2812	3	20.7	1	...	...	...
17433-1750	3	15.0	1	...	...	R
17433-2523	3	19.9	1	...	...	R
17443-2519	3	16.1	1	...	...	...
17501-2656	3	23.0	1	928	8	R
18071-1727	3	11.0	2	1488	9	...
18083-2630	2,3	18.2	2	...	...	...
18172-2305	3	15.9	1	...	...	...
18198-1249	3	15.3	2	845	9	...
18257-1000	2,3	19.1	2	1975	9	...
18273-0738	2	14.3	2	700	5	R
18276-1431	2	12.0	2	890	9	R
18312-1209	3	12.3	1	...	...	...
18314-1131*	3	...	1	...	...	R
18348-0526	2,3	13.9	2	1575	3	...
18432-0149	2,3	17.5	2	1140	5	...
18437-0643	2,3	12.1	2	652	9	...
18460-0254	2,3	20.2	2	1730	5	...
18488-0107	3	20.5	1	1540	4	...
18549+0208	3	13.7	1	840	9	...
18560+0638	1,2,3	16.0	2	1033	5	...
19059-2219	2,3	13.3	2	510	1	R
19161+2343	2,3	17.9	2	...	...	...
19192+0922	2,3	16.5	2	552	9	...
19254+1631	2	19.0	2	1162	9	R
19426+4342	1,3	8.6	1	...	...	R
20077-0625	1,3	12.3	2	680	1	R
20109+3205	3	5.7	1	382	1	...
20181+2234	1,3	11.2	2	...	...	...
20272+3535	3	12.0	2	1603	9	...
22177+5936	3	15.1	2	1215	9	...

\* Single peak OH maser.

R: optical counterpart too red;

...: optical counterpart not found.

<sup>1</sup> See applicable references in Table 5.**Table 8.** Comparison stars.

IRAS name	Run	$v_{\text{exp}}(\text{OH})$ ( $\text{km s}^{-1}$ )	Ref. <sup>1</sup>	Period (days)	Ref. <sup>1</sup>
00213+3817	1	9.0	10	409	1
02192+5821	1	14.5	2	822	1
10436-3459	4	11.3	1	104	1
18397+1738	2	13.4	8	511	3
19133-1703	1	10	10	395	1
20166+3717	3	...	...	410	1
23416+6130	1	26.5	2	925	1
23554+5612	1	15.5	10	477	1
23587+6004	1	...	...	186	1

<sup>1</sup> See applicable references in Table 5.**Table 9.** Peculiar C-rich AGB stars.

IRAS name	Run	$v_{\text{exp}}(\text{CO})$ ( $\text{km s}^{-1}$ )	Ref. <sup>1</sup>	Period (days)	Ref. <sup>1</sup>	Notes
00247+6922	3	17.2	8	650	3	...
04130+3918	3	...	...	470	10	B
09425-6040	4	...	...	...	...	B
17297+1747	2	15.2	11	520	3	R
19321+2757	2	24.4	8	625	3	R
19594+4047	3	20.5	8	783	3	...
20072+3116	2	25.6	8	550	3	R
23166+1655	3	15.1	8	520	3	...
23320+4316	3	14.7	8	620	3	R

B: bright optical counterpart;

R: optical counterpart too red;

...: optical counterpart not found.

<sup>1</sup> See applicable references in Table 5.