

# Isolated galaxies: residual of primordial building blocks?

G. Galletta<sup>1</sup>, G. Rodighiero<sup>1</sup>, D. Bettoni<sup>2</sup>, M. Moles<sup>3</sup>, and J. Varela<sup>2</sup>

<sup>1</sup> Dipartimento di Astronomia, Università di Padova, Vicolo dell'Osservatorio 2, 35122, Padova, Italy  
e-mail: giuseppe.galletta@unipd.it

<sup>2</sup> INAF – Osservatorio Astronomico di Padova, Vicolo dell'Osservatorio 5, 35122, Padova, Italy  
e-mail: bettoni@pd.astro.it, e-mail: jesus.varela@oapd.inaf.it

<sup>3</sup> Instituto de Astrofísica de Andalucía (C.S.I.C.) Apartado 3004, 18080 Granada, Spain  
e-mail: moles@iaa.es

Received 15 February 2006 / Accepted 25 April 2006

## ABSTRACT

**Context.** The mass assembly is believed to be the dominant process of early galaxy formation. This mechanism of galaxy building can proceed either by repeated major mergers with other systems, or by means of accretion of matter from the surrounding regions.

**Aims.** In this paper we compare the properties of local disk galaxies that appear isolated, i.e., not tidally affected by other galaxies during the last few Gyr within the volume given by  $cz \leq 5000 \text{ km s}^{-1}$ , with those galaxies at  $z$  values from 0.25 to 5.

**Methods.** Effective radii for 203 isolated galaxies and 1645 galaxies from the RC3 have been collected and the two samples have been analyzed statistically. A similar comparison has been made with half light radii studied at high  $z$  from the literature.

**Results.** We found that isolated galaxies are, in general, smaller than other present epoch galaxies from the RC3. We notice the lack of systems larger than 7 kpc among them. Their size distribution appears to be similar to that of galaxies at  $1.4 \leq z \leq 2$ . The models of the merging history also indicate that the isolated galaxies did stop their merging process at about that redshift, evolving passively since then. The galaxy density seems to have remained unchanged since that epoch.

**Conclusions.** Isolated galaxies appear to be the end products of the merging process, as proposed by the hierarchical accretion scenario at around  $z = 1.4$ . For this class of galaxies, this was the last significant merging event in their lives, and they have evolved passively since then. This is confirmed by the analytical estimate of the merging fraction with  $z$  and by the comparison with sizes of distant galaxies.

**Key words.** galaxies: formation – galaxies: evolution – galaxies: high-redshift – galaxies: spiral

## 1. Introduction

The mass assembly is believed to be the dominant process of early galaxy formation. This mechanism of galaxy building can proceed either by repeated major mergers with other similar systems, or by means of accretion of matter from the surrounding regions. Some of these processes are still active at the present epoch, as in the cases of interacting pairs (Arp 1966; Arp & Madore 1987; Vorontsov-Velyaminov 1959; Moles et al. 2004), while other galaxies show the mark of gas- or star-accretion processes that happened in the past, in the form of polar rings or of matter rotating with spin opposite to the main galaxy body (counterrotation). In all cases, the accretion process increases the final galaxy mass and powers its evolution.

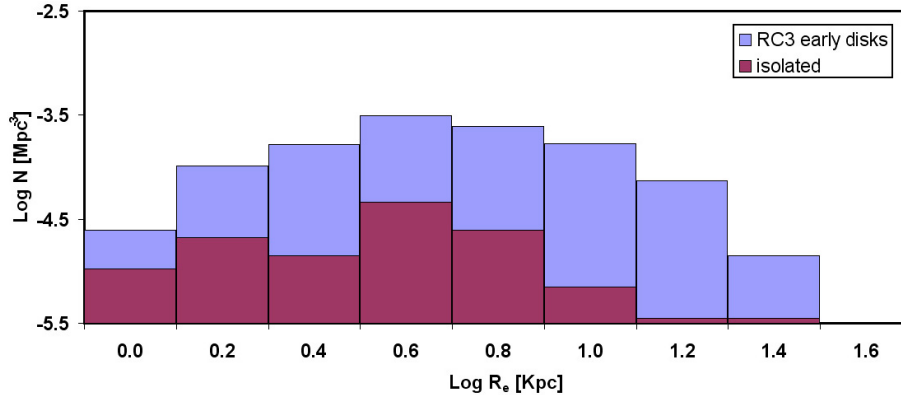
The merging or accretion processes are, in addition, independent from the original galaxy morphological type, so, in principle, galaxies of every morphological type may increase their sizes with cosmic time, from the first, smaller building blocks to the giant galaxies at the center of higher density environments.

According to the most extended view, galaxies form by accretion, very probably starting from clumpy systems similar to those observed in the Hubble Ultra Deep Field, later evolving to small exponential disks that grow from the inside out. They start with a dense halo and dense disk with spiral arms. The exponential structure of these earlier galaxies is predicted by model simulations, where different feeding mechanisms are considered (Abadi et al. 2003; Governato et al. 2004; Robertson et al. 2006).

Once started, the small spirals should grow, by adding material to the outer regions. The galaxy increases its visible size by accreting material and feeding the star formation processes from the center outwards. Often, strong interactions between galaxies may destroy these structures by merging or cannibalizing. This picture is supported by the observation of an increase in size for galaxies at different redshifts, as we shall discuss in the next sections (Ferguson et al. 2004; De Mello et al. 2006; Papovich et al. 2005). The building, from the inside out, may also be supported by a color gradient detected in galaxies at lower redshift (Papovich et al. 2005). The color evolution, according to the process of galaxy formation described above also has been investigated by semi-analytical models (Cattaneo et al. 2006). According to some authors (Elmegreen et al. 2005; Noeske et al. 2006), the small disk galaxies observed in the Ultra Deep Field may be the dense, inner regions of present-epoch small disk galaxies.

The efficiency of the process of building large galaxies depends on the initial local density, so galaxies that are initially in lower density environments have a lower chance of having large merging events with other galaxies or of accreting matter and should, in principle, be smaller and less evolved than the average of galaxies.

In a previous paper (Varela et al. 2004, hereafter Paper I), we presented a catalog of disk galaxies that appear not tidally affected by other galaxies in their evolution during the last few Gyr and, on this basis, are defined isolated. Within the volume given



**Fig. 1.** Comparison between the histogram of isolated galaxies with effective radius  $>0.5$  kpc and the galaxies with the same morphological type contained in RC3.

by  $cz \leq 5000 \text{ km s}^{-1}$ , 12% of the original sample galaxies satisfy this criterion. We found that they markedly differ in size, morphology, luminosity, and color indices from galaxies that are in higher density environments.

Within the hierarchical galaxy formation frame these properties indicate that the building process was stopped, long time ago, and they were left to evolve only passively from then on. In other words, if the building of the galaxies has proceeded by nested merging of smaller building blocks, these isolated galaxies, selected because they were not involved in significant merge events in the last part of their life, are in evolutionary stages as close as can be to their high redshift counterparts.

In the present work we consider these galaxies and their properties in light of the recently available data on high redshift systems to see whether both kind of objects similarly reflect the galaxy population at those epochs.

## 2. The size of isolated galaxies

### 2.1. Comparison with normal galaxies

In Paper I we compared the size and other properties of the isolated galaxies in the catalog we built with that of a sample of perturbed galaxies, in mild interaction with their neighbors. We found that isolated galaxies are smaller, bluer, less massive, and more frequently late type-spirals than the perturbed galaxies. These facts agree with the scenario of early disk formation described in the Introduction. Here we follow the analysis comparing the size of isolated objects in our sample with a larger sample of stellar systems with similar morphological types extracted from the RC3 (de Vaucouleurs et al. 1991) and LEDA (Paturel et al. 1997) catalogues. We discarded all the galaxies already classified as isolated from these large samples so, even if this cannot exclude the isolated galaxies that may exist among the galaxies present in these huge compilations, they are a fair representation of non-isolated systems.

Isolated galaxies in Paper I were extracted from the CfA2 catalogue (Huchra et al. 2000) with a clear physical criterion, namely, having no sizable companions within a volume large enough to ensure that they could not be affected in their evolution by other systems in the last Gyr. Other practical criteria were 1) having the redshift available, to establish the absolute magnitude and physical size; 2) being within a volume limited by  $cz < 5000 \text{ km s}^{-1}$ ; 3) having morphological types from S0 to Sc ( $-3 \leq t \leq 5$ ); and 4) being at Galactic latitude  $|b| \geq 40^\circ$  and declination  $\delta \geq -2.5$ . All the selected 203 galaxies have

measured diameters at the 25 mag arcsec<sup>2</sup> isophote,  $D_{25}$ , whereas only 38 also have the effective diameter,  $A_e$ , measured.

We have applied the same kind of constraints, excluding that of being isolated, to the RC3 catalogue, to build a big enough comparison sample. We found data for 1645 systems with 826 disk galaxies. Among them, only 483 (325 disks) have measured effective diameters. After the conversion of the  $R_e$  radii in kpc, we built histograms of size distributions, taking the RC3 and isolated galaxy samples binned every  $\Delta \log R_e = 0.2$  kpc and normalized to the volume of space defined by our adopted selection criteria. The histograms showing the difference in effective radii between RC3 and isolated galaxies are shown in Fig. 1, in the range of galaxy sizes from  $\sim 1$  kpc to 30 kpc.

It is clear that the two distributions look different, since systems with  $R_e$  larger than about 10 kpc are very rare among isolated galaxies. This translates in a difference in their respective median values that amounts to 0.489 (3.08 kpc) for the RC3 galaxies, whereas for isolated galaxies it is only 0.369 (2.34 kpc). Both the T-test and the F-Test indicate that the difference between the median values is statistically significant at the 95% level.

The clear difference in size distribution between isolated and non-isolated objects would then reflect the difference in evolutionary histories between both families, which, following the standard views, is dominated by merging and accretion. If the larger galaxies are formed as the final step of a series of collisions and merging processes between close systems, this difference indicates that this process stopped much earlier for isolated objects than for the others.

### 2.2. Merging fractions at different $z$ values

The previous result indicates that isolated galaxies were not altered by any significant merging process for a large fraction of their histories. To identify the epoch when these galaxies were excluded from the more general evolving processes, we have considered the models that describe the merging process and its history with the redshift. The merging fraction of galaxies at different epochs is generally expressed by (Patton et al. 2000)

$$f(z) = f_0 \times (1 + z)^m,$$

where  $f_0$  and  $m$  are parameters whose values change from author to author. Thus, Patton et al. (2000) assumes a constant merging fraction of the form  $f(z) = 0.011$ , while Le Fèvre et al. (2000), using more data, finds  $f_0 = 0.019$  and  $m = 3.25$ . More recently,

**Table 1.** Percentage of galaxies  $f(\leq z_m)_{\text{surv}}$  that survive the merging process at different  $z$ , according to two models. See the text for details.

$f_0$	$m$	$f(\leq z_m)_{\text{rem}}$	$f(\leq z_m)_{\text{surv}}$	$z$	Notes
Conselice et al. (2003)					
0.06	0.5	88.0%	12.0%	3.8	$M_B \leq -18$
0.07	0.4	88.4%	11.6%	3.7	$M_B \leq -19$
0.07	0.7	88.5%	11.5%	2.9	$M_B \leq -20$
0.004	3.7	–	–	>5	$M_B \leq -21$
Le Fèvre et al. (2000)					
0.019	3.25	88.1%	11.9%	1.7	–

Conselice et al. (2003) have found that both values are different if massive or small galaxies are considered at different redshift ranges. They found  $f_0$  ranging from 0.004 to 0.07, and  $m$  ranging from 0.5 to 3.7 or higher.

Note that this approach assumes a constant local density ( $m$  constant) or describes the local density changes, assuming a different  $m$  slope for merging rates with  $z$ . More accurate models may be considered for a fine-tuning of the resulting numbers.

Assuming the literature values, we estimated the fraction of galaxies that have undergone mergers since  $z = z_m$ , following the method of Patton et al. (2000) by means of the formula:

$$f(\leq z_m)_{\text{rem}} = 1 - \prod_{z=0}^{z=z_m} \frac{1 - f(z)}{1 - 0.5f(z)}.$$

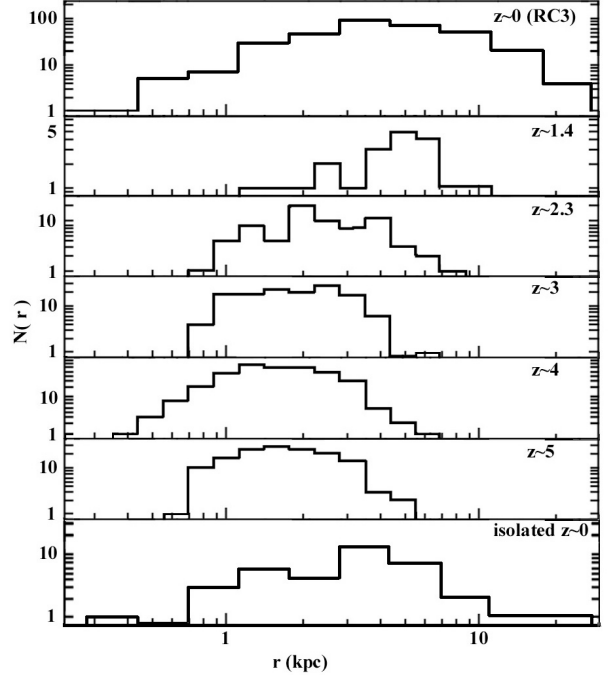
The percentage of surviving galaxies that “avoided” a merger since  $z_m$  is then

$$f(\leq z_m)_{\text{surv}} = 1 - f(\leq z_m)_{\text{rem}}.$$

We may estimate when the merging process may have stopped for them by calculating, with different values of  $f_0$  and  $m$ , at what redshift  $z$  the percentage of surviving galaxies  $f(\leq z_m)_{\text{surv}}$  is equal to the observed percentage of isolated galaxies, that is,  $\sim 11.62\%$  of the original CfA2 sample, being 203 over 1745 (Varela et al. 2004). This is shown in Table 1, assuming different values for  $f_0$  and  $m$  from the literature (Conselice et al. 2003; Le Fèvre et al. 2000). The model of Conselice et al. (2003) is divided in absolute magnitude bins, characterized by different  $m$  values. The values in the table indicate that all the galaxies have undergone several mergers since their formation because for  $z > 5$  the  $f(\leq z_m)_{\text{rem}}$  value is greater than 100%. Therefore, according to this result, our isolated galaxies may hardly be residuals of the original blocks that have built the present-epoch galaxies, and none of those blocks could have survived with its pristine properties until now. However, it appears that the present fraction of isolated systems is close to that at redshift between 1.7 and 3.8, which would imply that the merging process was stopped at that epoch for such a fraction of galaxies. This prevented them from merging again and producing larger size galaxies, as shown in Fig. 1, and from evolving like the other non-isolated galaxies.

### 2.3. Comparison with the sizes of galaxies at higher redshifts

Looking back in time, we may expect to see an epoch in which the galaxy size distribution was similar to that of present day isolated galaxies. This latter may be then compared with that of galaxies observed at higher redshifts. Ferguson et al. (2004), Papovich et al. (2005) and De Mello et al. (2006) have shown that an evolution of the physical size exists for galaxies from  $z \sim 1$  to  $z \sim 5$ . Comparing our data with those, we plotted all the values in linear scale (kpc).

**Fig. 2.** The size distribution of present epoch galaxies from the RC3 (top panel) and galaxies at different redshifts from Ferguson et al. (2004) (intermediate panels). We scaled our effective radii to the  $r$  radius in arcsecond at  $z = 4$  used by these authors, to have a direct comparison. In the bottom panel we add the sizes histogram for isolated galaxies.

The resulting histograms for local and distant samples are shown in Fig. 2, where the sizes of the galaxies in the largest, local existing sample are plotted together with the sizes of galaxies at several redshift intervals (Ferguson et al. 2004).

Indeed, galaxies at different  $z$  intervals are normally observed in different rest wavelength intervals, so the comparison is not so straightforward. Ferguson et al. (2004) selected galaxies from the HST Advanced Camera Survey (ACS) that are observed in their UV rest-frame. The local galaxies are, on the contrary, studied in B images. The Ferguson et al. (2004) selection may favor galaxies with large UV luminosities that are morphologically different from present-day systems and whose dimensions measured at the UV band may not represent the whole galaxy size, but only that of the star-forming regions (De Mello et al. 2006; Elmegreen et al. 2005). However, according to Lotz et al. (2005), the galaxy morphology does not appear to change if galaxies are observed between  $1.5 \leq z \leq 4$ . This means that clumpy objects will appear with similar morphologies in this redshift interval.

By comparing sizes of galaxies selected from UV data and scaled to a B-rest frame with that of galaxies selected in the NIR and scaled to the same rest frame, De Mello et al. (2006) find that the median size of UV selected galaxies appear larger by  $\sim 21\%$  than that of systems selected in the NIR. This may result from selecting the brightest and biggest starburst galaxies in FUV, and missing the smaller and less active objects. However, this is not always true for all the samples in the literature. Thus, the size distribution plotted in Fig. 4 of Papovich et al. (2005) around  $z \sim 2.3$  (B-rest frame) has a peak at 2 kpc like that of Ferguson et al. (2004), but narrower.

Another source of error in comparing effective radii available for RC3 and isolated galaxies with the half radii adopted by Ferguson et al. (2004) may derive from the different methods used. Effective radii are computed in large galaxies by means

of an integration of the total luminosity in nested apertures, and then using a numerical method to extrapolate the total magnitude and the radius containing half of the luminosity. On the other side, the study of galaxies at high redshift and small apparent sizes is often performed with different techniques. In the case of the galaxies studied by Ferguson et al. (2004), the program SExtractor has been used, with particular assumptions on the fit parameters (see Ferguson et al. 2004). Adopting the same values for the parameters as Ferguson et al. (2004), we have calculated the effective radii for the 713 galaxies observed in the field of the Abell 85 cluster (Varela et al., in preparation) using SExtractor, but using also GASPHOT (Pignatelli et al. 2006), a program that performs a surface brightness distribution and then gives more accurate effective radii. In general, effective radii calculated with SExtractor are systematically smaller than the values obtained using other packages, being lower than 20% for galaxies smaller than 5 arcsec in the analyzed images (0.4 arcsec/pixels scale).

The sizes plotted by Ferguson et al. (2004) derive from the HST Advanced Camera for Surveys data with a scale of 0.05 arcsec pix<sup>-1</sup>, and the apparent size of their farthest galaxies falls in the above range of pixel sizes. So, we may expect that the histograms plotted in Fig. 2 may be shifted toward higher sizes by 20% at maximum, or from 7 kpc to 8.4 kpc in intrinsic size. This shift, however, does not significantly change our results and, therefore, we can confirm that large systems are very scarce in the distribution of the size of galaxies in the range  $1.4 \leq z \leq 5$ . They are, in that sense, similar to local isolated galaxies.

When plotted on the same scale, the peak of the size distribution of distant galaxies moves toward higher values, while larger and larger galaxies are built at the expense of the smallest stellar systems, which are phagocytated inside bigger ones. The comparison with RC3 confirms this trend. The present-day galaxies of RC3 have a peak size at effective radii bigger than galaxies at higher  $z$ . Isolated galaxies have a peak size falling close to that of the distant galaxies of Ferguson et al. (2004) at redshifts between 1 and 2.

In addition to the shift of the peak toward larger radii with decreasing  $z$ , in Fig. 2 a lack of systems with effective radii larger than 7 kpc (corresponding to 1 arcsec at  $z = 4$ ) in the distant samples is also visible. This indicates that galaxies larger than 7 kpc (approximately the size of a galaxy like M33) are not yet formed in the Universe or are very rare, still at  $z = 1.4$ . On the contrary, in RC3, at  $z = 0$ , they represent 22.5% of the total, and in the isolated galaxies sample, 10.5%. We may deduce that isolated galaxies have not taking part in the merging processes that happened after  $z \geq 1$ .

Our data can be compared with the theoretical values of the mean galaxy size with redshift calculated, assuming different cosmological scaling laws. These laws describe the size evolution in an expanding Universe and are related to the assumed cosmological geometry. The geometric effect may be described by simple functions (e.g.,  $R_e \propto H^{-1}$  or  $R_e \propto H^{-2/3}$ ) and adds to the gradual increase with time of the physical size, depicted by hierarchical models, due to the accretion of matter from the environment or from the merger with smaller companions along the galaxy lifetime. In the literature, the scaling laws predict an increase of this parameter with  $z$  (Ferguson et al. 2004). This is also confirmed by various recent observations up to redshift  $z \sim 5$  (Ferguson et al. 2004; Cassata et al. 2005).

Our resulting mean size of the whole local disk population ( $R_e = 3.07$  kpc from RC3), which is greater than the mean sizes of all the galaxies samples at higher redshifts (Ferguson et al. 2004; Cassata et al. 2005), is in full agreement with the

predictions of these models at low redshift. In addition, the mean size of isolated galaxies ( $R_e = 2.37$  kpc) is consistent with the mean galaxy size expected and observed at  $z$  between 1.5 and 3.

#### 2.4. Density evolution

We further check this hypothesis by looking at the density evolution of disk galaxies. Ravindranath et al. (2004) have computed the observed number density of disk galaxies in the GOODS CDF-S field and they found no significant evolution from redshift  $z = 0.25$  to  $z = 1.25$ . They divided the sample into small disks ( $R_e < 4$  kpc) and large disks ( $R_e \geq 4$  kpc). Within the uncertainties, neither sample shows a strong evolution and remains almost constant.

The  $R_e < 4$  kpc selection criterion for small disks matches the typical size of our isolated disks sample (median  $R_e = 2.34$  kpc). As we mentioned above, along the cosmic time, smaller galaxies should have undergone only minor merging episodes, as suggested by the hierarchical growth of galaxy structures. If the number density of small disks is constant from  $z = 0.25$  to  $z = 1.25$ , this might support our suggestion about the isolated population and about the disk population in general. Ravindranath et al. (2004) conclude that luminous disk galaxies were present with roughly the same abundances at  $z = 1$  as at low redshift ( $z = 0.25$ ), and are likely to have undergone only a modest luminosity evolution (and then a modest number of merger events).

To better constrain the results in the lower redshift bin, where the GOODS statistics are quite poor, we adopt the RC3 sample to estimate the local density of disks. Assuming the same value of Ravindranath et al. (2004), to discriminate between small and large disks, we found a density that is still consistent with a non-evolutionary scenario for the density of disk galaxies. This is shown in Fig. 3, where we extend the plot of Fig. 5 of Ravindranath et al. (2004) to  $z < 0.25$ .

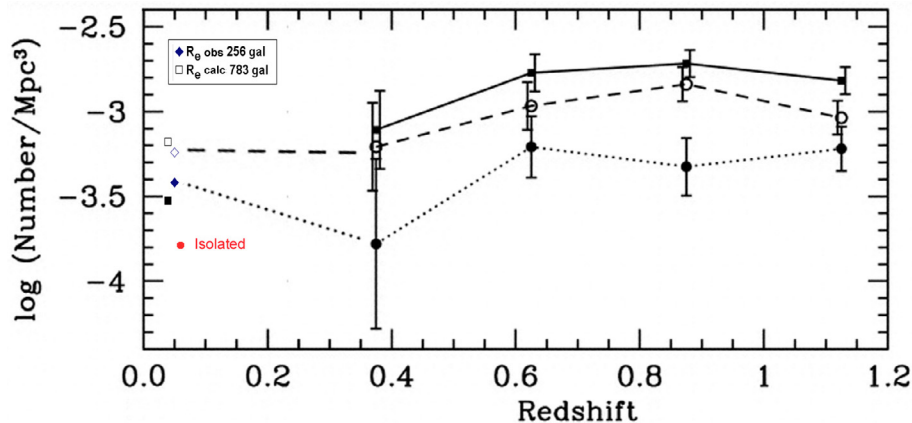
We adopt the same size value for  $R_e = 4$  kpc to separate local small disks (filled lower symbols) and local large disks (upper empty symbols).

### 3. Conclusions

From the above discussion, it appears that isolated galaxies could be important for understanding the history of the galaxy building processes at redshift larger than 1.4. The similarity of properties between high redshift galaxies and the local isolated objects we have selected, indicates that, within the frame of hierarchical galaxy formation, these local isolated galaxies would have been out of the process since that epoch.

Indeed, our isolated systems are rather small, with only a fraction of 10% having  $R_e > 7$  kpc, at variance with the fraction of 22.5% that we find in the sample of RC3 non-isolated galaxies selected within the same volume and with the same specifications. The median effective radius amounts to 2.3 kpc for the isolated galaxies, while for RC3 this value is 3.1 kpc.

The fraction of isolated galaxies in the large sample we explored amounts to 11.7% (see Paper I). This figure can be compared with the predictions from models of hierarchical galaxy formation that allow us to calculate the percentage of “surviving” galaxies, that is, the fraction of galaxies not affected by merging at a given epoch. Even if different authors use different parameter values to describe the merging evolution, we have found that for a rather large set of parameters the results are similar, around 12% at  $z \sim 1.7$ . This is very close to the percentage



**Fig. 3.** Observed number densities of disk galaxies as a function of redshift. The number densities of small disks ( $R_e < 4$  kpc, open symbols, dashed lines) and large disks ( $R_e \geq 4$  kpc, filled symbols, dotted lines) by Ravindranath et al. (2004) with  $z > 0.25$  do not show significant evolution in their relative abundance with redshift. Filled squares and a solid line represent the total. Our representative point, for local systems, has been arbitrary plotted at  $z = 0.1$ , to be visible far from the axis. Data from the RC3 effective radii (256 galaxies) are shown as diamonds; data calculated using  $R_{25}$  (783 galaxies) are plotted as squares.

of isolated local galaxies we found, which let us to conclude that isolated galaxies had experienced their last merging event by that epoch.

To further explore this possibility, we have compared the sizes, given by the effective radii, of the isolated galaxies with these of systems at  $z \sim 1.4$  and at  $z \sim 2.3$ . The result is that similar sizes are found for both families in a rather narrow  $z$ -range, close to 1.4. Therefore, isolated galaxies with sizes larger than 7 kpc such as those in our RC3 sample, continued the building process built after  $z \leq 1$ , at variance with our local isolated galaxies.

We also considered the possible density evolution of galaxies as a function of their size. We find that neither the density of galaxies smaller than 4 kpc, nor that of galaxies larger than 4 kpc have evolved since  $z = 0.4$ . Our RC3 data confirms this lack of evolution to now.

All the above results suggest that isolated galaxies were definitively assembled at  $z \sim 1.4$ , when the merging process stopped due to the low density environment, and have evolved passively since then.

The merging history could thus be retraced by looking at local galaxies in gradually richer environments.

*Acknowledgements.* This paper has been partially funded by the contract 60A02-2077/05 Fondi 2005 – University of Padua. M.M. acknowledges a grant by the Spanish Ministerio de Educacion y Ciencia, reference AYA2002.01241. We would like to thank Dr. Prada for helpful discussions. G.R. acknowledge Prof. Franceschini for useful discussions. The authors also thank the referee for useful comments.

## References

- Abadi, M. G., Navarro, J. F., Steinmetz, M., & Eke, V. R. 2003, *ApJ*, 597, 21
- Arp, H. 1966, *Atlas of Peculiar Galaxies*, California Institute of Technology, Pasadena, CA
- Arp, H. C., & Madore, B. F. 1987, *A Catalog of Southern Peculiar Galaxies and Associations* (Cambridge University Press)
- Cassata, P., Cimatti, A., Franceschini, A., et al. 2005, *MNRAS*, 357, 903
- Cattaneo, A., Dekel, A., Devriendt, J., Guiderdoni, B., & Blaizot, J. 2006 [arXiv:astro-ph/0601295]
- Conselice, C. J., Bershad, M. A., Dickinson, M., & Papovich, C. 2003, *AJ*, 126, 1183
- De Mello, D. F., Wadadekar, Y., Dahlen, T., Casertano, S., & Gardner, J. P. 2006, *AJ*, 131, 216
- de Vaucouleurs, G., de Vaucouleurs, A., Corwin, H. G., et al. 1991, *Third Reference Catalogue of Bright Galaxies (RC3)* (New York: Springer-Verlag)
- Elmegreen, B. G., Elmegreen, D. M., Vollbach, D. R., et al. 2005, *ApJ*, 634, 101
- Ferguson, H. C., Dickinson, M., Giavalisco, M., et al. 2004, *ApJ*, 600, L107
- Governato, F., Mayer, L., Wadsley, J., et al. 2004, *ApJ*, 607, 688
- Huchra, et al. 2000, *CfA2 Catalogue of galaxy redshift*, Center for Astrophysics
- Le Fèvre, O., Abraham, R., Lilly, S. J., et al. 2000, *MNRAS*, 311, 565
- Lotz, J. M., Madau, P., Giavalisco, M., Primack, J., & Ferguson, H. C. 2006, *ApJ*, 636, 592
- Moles, M., Bettoni, D., Fasano, G., Kjærgaard, P., & Varela, J. 2004, *A&A*, 418, 495
- Noeske, K. G., Koo, D. C., Phillips, A. C., et al. 2006, *ApJ*, 640, L143
- Papovich, C., Dickinson, M., Giavalisco, M., et al. 2005, *ApJ*, 631, 101
- Patton, D. R., Carlberg, R. G., Marzke, R. O., et al. 2000, *ApJ*, 536, 153
- Paturel, G., Andernach, H., Bottinelli, L., et al. 1997, *A&AS*, 124, 109
- Pignatelli, E., Fasano, G., & Cassata, P. 2006, *A&A*, 446, 373
- Ravindranath, S., Ferguson, H. C., Conselice, C., et al. 2004, *ApJ*, 604, L9
- Robertson, B., Cox, T. J., Hernquist, L., et al. 2006, *ApJ*, 641, 21
- Varela, J., Moles, M., Márquez, I., et al. 2004, *A&A*, 420, 873
- Vorontsov-Velyaminov, B. A. 1959, *Atlas and catalog of interacting galaxies*