

X-ray flares in the early Swift observations of the possible naked gamma-ray burst 050421

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ABSTRACT

We present the *Swift* observations of the faint burst GRB 050421. The X-ray light-curve shows at least two flares: the first flare peaking at ~ 110 s after the BAT trigger (T_0) and the second one peaking at ~ 154 s. The first flare presents a flux variation of $\delta F/F_{\text{peak}} \sim 3.7$ and a short timescale ratio $\delta t/t_{\text{peak}} \sim 0.07$. The second flare is smaller and presents a flux variation of $\delta F/F_{\text{peak}} \sim 1.7$ and a short timescale ratio $\delta t/t_{\text{peak}} \sim 0.03$. We argue that the mechanism producing these flares is probably late internal shocks. The X-ray light-curve is consistent with a rapid decline with a temporal index $\alpha \sim 3.1$, which decays from $\sim 10^{-9}$ erg cm $^{-2}$ s $^{-1}$ at $T_0 + 100$ s to $< 7 \times 10^{-13}$ erg cm $^{-2}$ s $^{-1}$ at $T_0 + 900$ s. A possible spectral softening is also observed with time, from $\beta \sim 0.1$ to ~ 1.2 . A good joint fit to the BAT and XRT spectra before $T_0 + 171$ s with $\beta_{\text{XRT-BAT}} \sim 0.2$ indicates that the early X-ray and Gamma-ray emissions are likely produced by the same mechanism. We argue that the X-ray spectral softening, if any, is due to a shift of the peak of the prompt emission spectrum down to lower energies, and that the rapid decline of the X-ray emission is probably the tail of the prompt emission. This suggests that the X-ray emission is completely dominated by high latitude radiation and the external shock, if any, is extremely faint and below the detection threshold. GRB 050421 is likely the first “naked burst” detected by *Swift*.

Key words. gamma ray: bursts

1. Introduction

The *Swift* Gamma-Ray Burst (GRB) Explorer was launched on 20th November 2004 (Gehrels et al. 2004). It is a multi-wavelength observatory covering the Gamma-ray, X-ray and UV/optical bands. After the detection of a GRB by the Burst Alert Telescope (BAT) (Barthelmy et al. 2005), the observatory slews automatically and rapidly to point the narrow field instruments: the X-ray Telescope (XRT, Burrows et al. 2005), and the UV/Optical Telescope (UVOT, Roming et al. 2005). Due to its unique capability to slew on a timescale of minutes and its high sensitivity, *Swift* is ideal to study the properties of the early afterglow, in particular the transition between the prompt emission and the afterglow.

The current Gamma-Ray Burst picture is that the GRB and the afterglow are produced by a relativistically expanding blast-wave associated with internal and external shocks (Mészáros & Rees 1993; Rees & Mészáros 1994; Kobayashi et al. 1997; Panaitescu & Mészáros 1998). In the simplest version of this model, the prompt Gamma-ray emission is produced by internal shocks. The duration of the prompt emission extends from

0.01 s to 1000 s with two distinct GRB populations: the < 2 s short GRBs peaking typically at ~ 0.2 s and the < 2 s long GRBs peaking typically at ~ 20 s (e.g. Kouveliotou et al. 1993; Dezalay et al. 1996; Paciesas et al. 1999).

Currently, it is not completely understood how the central engine can still be active over a timescale extending from 0.01 s to 1000 s for the long GRBs. Sari & Piran (1997) have shown that the time structure of GRBs must be intrinsic to their energy source. A possible explanation marking a departure from the current hypernova picture is that the extended activity of the GRBs results in the fragmentation of the collapsing stellar core in proto neutron stars and the subsequent merger of these fragments (Davies et al. 2002; King et al. 2005). Other models have also been proposed (e.g. see Mészáros 2000; Zhang & Mészáros 2004, and references in these papers).

When the blastwave shocks the surrounding medium, it is decelerated and a broad-band afterglow is produced. The afterglow emission usually fades with time as a broken power-law. For a more detailed picture of this model, see e.g. Zhang & Mészáros (2004).

Recent X-ray observations of GRBs 100–1000 s after the trigger have shown unexpected and complex light-curve behaviour, in particular a steep-to-flat transition in the early temporal decay (e.g. Tagliaferri et al. 2005; Burrows et al. 2005b; Nousek et al. 2006). An explanation of the initial steep decline of the X-ray light-curve is perhaps the curvature effect (Kumar & Panaitescu 2000; Zhang et al. 2006). This results in delayed arrival of the prompt emission from the high angular latitude regions of the blastwave relative to the observer direction.

X-ray flares have also been reported in the early X-ray light-curves. XRF 050406 was the first *Swift* GRB for which such a feature was observed by the XRT (Romano et al. 2006), peaking at ~ 210 s after the trigger. GRB 050502b also presented an X-ray flare peaking at 345 s after the trigger, but with the fluence of the flare larger in the XRT band than in the prompt Gamma-ray emission itself in the 15–350 keV band (Burrows et al. 2005b; King et al. 2005; Falcone et al. 2006). Moreover, smaller flares could be seen for this burst beyond $\sim 10^4$ s. Other *Swift* GRBs also show X-ray flares in their light-curve (e.g. Nousek et al. 2006; O’Brien et al. 2006). Most of the flares seen in the X-ray light-curves have been interpreted as evidence that the internal shock activity extends beyond the prompt Gamma-ray emission.

Piro et al. (2005) also reported the observation of X-ray flares in GRB 011121 and GRB 011211. The authors have argued that the late X-ray flares represent the onset of the afterglow, because the spectral slopes of the flares are inconsistent with those of the prompt emission, but consistent with those of the afterglow. In addition, they have shown that the tail of the X-ray flares and the light-curve of the afterglow at 1 day are connected by a single power law $\sim (t - t_0)^{-\alpha}$, t_0 corresponding to the onset of the X-ray flares.

We report in this paper the *Swift* observation of the GRB discovered by the BAT on 21st April 2005. The observatory automatically slewed to the BAT position. The XRT onboard software found and centroided on a rapidly decaying X-ray source. This source exhibited two X-ray flares in its light-curve and, at $T_0 + 900$ s, became undetectable (with an 3σ upper limit on the unabsorbed X-ray flux $F_\nu < 7 \times 10^{-13}$ erg cm $^{-2}$ s $^{-1}$ in the 0.3–10 keV band).

The paper is organised as follows: in Sect. 2, we present the characteristics of the observations and the basic steps of the data reduction. In Sect. 3, we present the spectral and temporal analysis of the multi-wavelength observations. In Sect. 4, we investigate the mechanisms producing the spectral and temporal characteristics of the burst. We also explore the possibility that the flares in the light-curve could be the evidence of late internal shocks.

By convention, we note hereafter the flux in the X-ray band is modelled as $F_\nu \propto \nu^{-\beta} t^{-\alpha}$ where β is the energy spectral index and α is the temporal index. We use the symbol Γ to refer to the Lorentz factor. The BAT spectral slope is noted as β_{BAT} and T_0 is the trigger time.

2. Observation and data reduction

2.1. BAT observations

GRB 050421 (*Swift*-BAT trigger 115135) was detected at 04:11:52 UT on 21st April 2005 at (J2000) RA = $20^{\text{h}}28^{\text{m}}58^{\text{s}}$ and Dec = $+73^{\circ}39'54''$, with an uncertainty of 4 arcmin (Barbier et al. 2005).

The BAT spectra and light-curves were extracted using the BAT analysis software (build 2.0) as described in the *Swift* BAT Ground Analysis Software Manual (Krimm et al. 2004).

2.2. XRT observations

The XRT started to observe 89 s after the trigger, following its normal sequence of readout modes: Low rate Photo-diode mode (LrPD) during the slew, Image mode (IM) when the spacecraft settled, Piled-Up Photo-Diode (PuPD), LrPD, Windowed Timing (WT), and then Photon Counting (PC) modes while pointed at the target and according to the source brightness (see Hill et al. 2004, 2005b). An uncatalogued X-ray source was identified on ground at (J2000) RA = $20^{\text{h}}29^{\text{m}}01.7^{\text{s}}$ and Dec = $+73^{\circ}39'17.4''$ with an uncertainty of 5 arcsec at a 90% confidence level (Godet et al. 2005). Recent calibration of the XRT boresight has allowed an improvement in the XRT position measurements (Moretti et al. 2006). Accounting for this, the refined position is (J2000) RA = $20^{\text{h}}29^{\text{m}}02.44^{\text{s}}$ and Dec = $+73^{\circ}39'17.8''$ with a total uncertainty radius of 3.7 arcsec at a 90% confidence level.

The XRT data were processed by the *Swift* software version 2.0¹. A cleaned event list was generated using the default pipeline, which removes the effects of hot pixels and the bright Earth.

From the cleaned event list, the source and background spectra were extracted using *xselect*. For the spectral analysis, only the WT and PC data were useful (see Sect. 3.2.2). The first 120 s of the PC data have a count rate above 1 cts s $^{-1}$ and therefore, are slightly piled-up (see Vaughan et al. 2005, for a detailed discussion about the pile-up effect). The innermost 3 pixel radius was excluded, and the source and background spectra were extracted using an annular region with an outer 30 pixel radius. The later PC data were extracted using a circular region within a 30 pixel radius. For the spectral analysis, we used the grade 0 events for the PC and WT modes, for which the response matrices are presently the best calibrated. The ancillary response files for the PC and WT modes were created using *xrtmkarf*.

The light-curves were also extracted from the cleaned event list using *xselect* over the 0.3–10 keV band. The LrPD grades 0–5, PC grades 0–12 and WT grades 0–2 events were used for the light-curves, giving slightly higher effective area at higher energies. The XRT observations are summarised in Table 1.

2.3. UVOT observations

The UVOT, which began to observe 112 s after the trigger, detected no optical fading source down to a 5σ limiting magnitude of 17.5 in V-band, 19.1 in B-band and 18.8 in U-band, using co-added images of the early time data of 9×10 s exposure (Blustin et al. 2005).

2.4. Other observations

No optical or IR transient has been detected. ROTSE-III (Rykoff et al. 2005) started to observe 67.2 s after the trigger, but no source was detected down to an unfiltered magnitude of 17.8 in 69 s. The BOOTES-2 (Burst Observer and Optical Transient Exploring System) very wide field camera started to observe 8 s after the onset of the burst, and no prompt optical flash arising from this burst was detected with a limiting (unfiltered) magnitude of 9 (Jelinek et al. 2005). The ~ 2000 s observations starting at 11:23:45 UT on the Mount Hopkins in the J, H and K bands did not reveal any new sources detected above 3σ within the XRT error circle (Bloom et al. 2005).

¹ See <http://heasarc.gsfc.nasa.gov/docs/swift/analysis/>

Table 1. Log of the XRT observations for GRB 050421 following the XRT mode sequence.

Sequence	XRT mode	Start time	End time	Start time since trigger	Exposure ^a
		(yy-mm-dd hh:mm:ss)	(yy-mm-dd hh:mm:ss)	(s)	(s)
00115135000	LrPD slew	2005-04-21 04:13:21	2005-04-21 04:13:28	89	7.2
00115135000	LrPD settling	2005-04-21 04:13:29	2005-04-21 04:13:37	97	8.2
00115135000	IM	2005-04-21 04:13:39	2005-04-21 04:13:41	107	2.5
00115135000	PuPD	2005-04-21 04:13:44	2005-04-21 04:13:45	112	1.3
00115135000	LrPD pointing	2005-04-21 04:13:45	2005-04-21 04:13:46	113	1.4
00115135000	WT	2005-04-21 04:13:47	2005-04-21 04:14:43	115	56
00115135000	PC	2005-04-21 04:14:43	2005-04-22 07:53:30	171	34 416
00115135001	PC	2005-04-23 00:58:12	2005-04-25 15:41:56	161 180	4008
00115135002	PC	2005-04-26 01:16:10	2005-04-26 19:14:57	421 546	6927

^a Only the exposure time for the PC mode is spread over several snapshots during each observation.

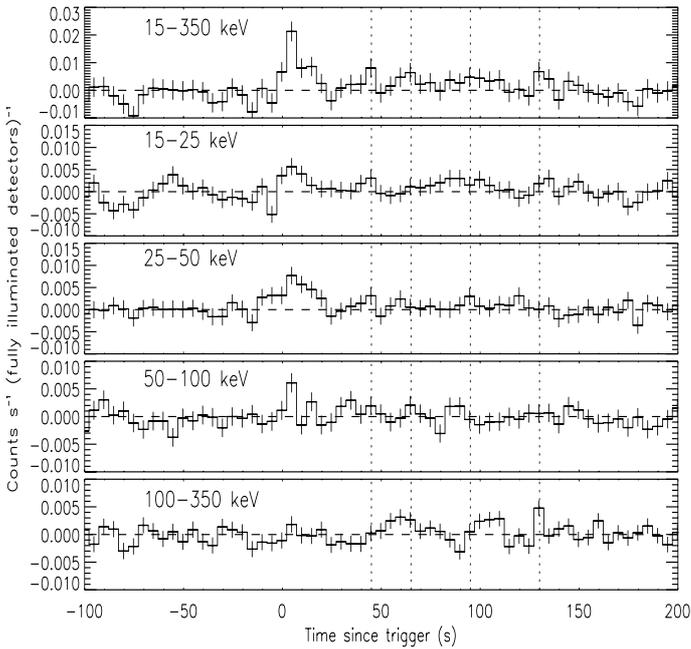


Fig. 1. BAT light-curve for 4 different energy bands from the top to the bottom: 15–350 keV, 15–25 keV, 25–50 keV, 50–100 keV and 100–350 keV. The dashed vertical lines represent possible smaller flares at ~ 45 , ~ 65 , ~ 95 –110 and ~ 130 s after the trigger. The BAT light-curve is given in units of counts s^{-1} (fully illuminated detector) $^{-1}$.

3. Data analysis results

All the errors cited below are given at a 90% confidence level for one parameter of interest (i.e. $\Delta\chi^2 = 2.706$).

3.1. Light-curve

3.1.1. Gamma-ray band

The BAT light-curve of GRB 050421 has a Fast-Rise and Exponential-Decay (FRED) structure (Sakamoto et al. 2005), with a rise time less than 1 s and a decay time of ~ 10 s (the upper panel in Fig. 1). Note that a tail or/and some further smaller flares may be present in the BAT light-curve. Figure 1 also shows the BAT light-curve in four different spectral bands.

In the 15–350 keV band, T_{50} and T_{90} are 4.8 ± 0.5 s and 10.3 ± 0.3 s respectively, while the 1-second peak photon flux is 0.5 ± 0.1 ph $cm^{-2} s^{-1}$. The fluence in the 15–150 keV band measured from T_0 to $T_0 + T_{90}$ is $1.1 \pm 0.7 \times 10^{-7}$ erg cm^{-2} .

3.1.2. X-ray band

The X-ray light-curve of GRB 050421 over the 0.3–10 keV band as shown in Fig. 2 combines data taken in all five data modes mentioned above. The light-curve from $T_0 + 97$ s to $T_0 + 200\,000$ s is fit using the BAT trigger time as the zero time.

The light-curve can be described by a simple and unbroken power law with a time decay index of $\alpha = 3.1 \pm 0.1$ with at least two flares superposed ($\chi^2 = 27.6/28$ d.o.f.). The two flares are fit by Gaussians. A possible third flare can be seen around $T_0 + 206$ s. The X-ray emission declines so rapidly that the X-ray emission became undetectable less than an hour after the trigger.

The index of the early steep decay is close to the mean of the observed range of X-ray temporal decay slopes given in Nousek et al. (2006) ($\alpha_{\text{mean}} \sim 3.4$). Note that if we only fit the XRT light-curve before $T_0 + 180$ s with the same model as discussed above, we find a shallower temporal slope of $\alpha = 2.0 \pm 0.6$.

The data for the first flare are sparse, because the flare occurred at the end of the spacecraft settling phase. The flare was then observed through a sequence of four modes (the LrPD mode during the settling of the spacecraft on the source, and then the IM, PuPD and LrPD modes when the XRT was pointed at the source). Therefore, the temporal characteristics of the flare given below should be taken with caution. The best fit indicates that the first flare peaks at $T_0 + 110 \pm 2$ s and has a fast rise with a small $\delta t_{\text{rise}}/t_{\text{peak}} \sim 0.05$. The timescale ratio for the decaying part of the flare $\delta t_{\text{decay}}/t_{\text{peak}}$ is ~ 0.03 , where δt_{rise} and δt_{decay} are the time intervals respectively from the underlying power-law decay to the peak and from the peak to the underlying power-law decay. If the underlying power-law decay is subtracted, the rise is fit as $t^{0.7 \pm 1.2}$ from $T_0 + 97$ s to $T_0 + 108$ s and the decay as $t^{-\alpha}$ with $\alpha < 10.7$ from $T_0 + 110$ s to $T_0 + 118$ s. During the flare, the source re-brightens by a factor $\delta F/F \sim 3.7$ between $T_0 + 100$ s and the peak.

The second flare is weaker, peaking at $T_0 + 154 \pm 3$ s with $\delta t_{\text{rise}}/t_{\text{peak}} \sim 0.03$ ($\delta t_{\text{decay}}/t_{\text{peak}} \sim 0.03$) and $\delta F/F \sim 1.7$ measured between $T_0 + 144$ s and the peak.

These two flares show values of $\delta t/t \ll 1$, which puts strong constraints on the physical mechanisms producing them (see Sect. 4.3).

3.2. Spectroscopy

3.2.1. Gamma-ray band

The pre-slew spectrum from $T_0 - 10$ s to $T_0 + T_{90}$ is well fit by a single power-law with the spectral index $\beta_{\text{BAT}} = 0.7 \pm 0.5$ (see Table 2), consistent with that given by Sakamoto et al. (2005).

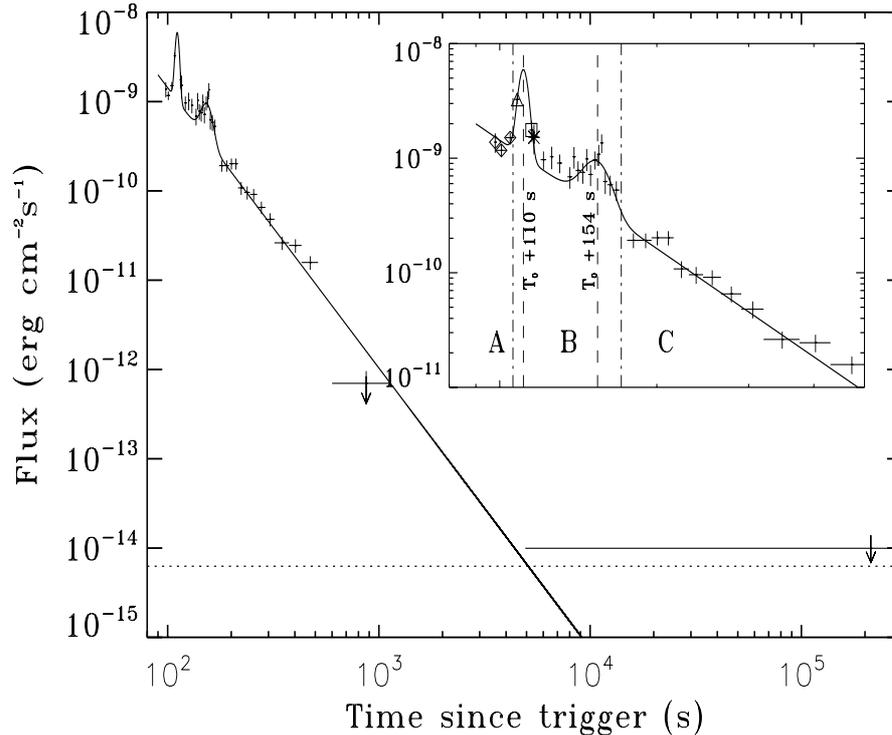


Fig. 2. Background subtracted XRT light-curve of GRB 050421 over the energy band 0.3–10 keV: **A**) LrPD slew and settling data (diamonds); **B**) IM (triangle), PuPD (square), LrPD (star) and WT pointing data; **C**) PC pointing data. The best fit to the light-curve is plotted by the solid line. The dotted line corresponds to the background level in the PC data. The upper limits are given at 3σ . The times of the two X-ray flares are also shown in the plot.

Use of a Band function (Band et al. 1993) or a cutoff power-law function did not significantly improve the fit. *Swift* GRB 050421 is the 44th faintest burst, in term of fluence, in a sample of 49 GRBs detected by the BAT² before the end of August 2005. We also extracted a spectrum from $T_0 + 20$ s to $T_0 + 100$ s to investigate a possible spectral softening of the prompt emission in the GRB tail. However, no significant spectral softening is seen in the spectrum, since the spectral slope for the spectrum from $T_0 + 20$ s to $T_0 + 100$ s is 0.8 ± 1.6 .

3.2.2. X-ray band

In the following analysis, we assume a Galactic absorption of $N_{\text{H}}^{\text{Gal}} = 1.44 \times 10^{21} \text{ cm}^{-2}$ (Dickey & Lockman 1990). The WT spectrum covers the temporal range from $T_0 + 115$ s to $T_0 + 171$ s, i.e. an exposure of 56 s. The PC data were divided into two spectra. We denote the first 120 s, from 171–291 s after T_0 , as “early” and the 480 s of data from 291–771 s after T_0 , as “late”. Beyond $T_0 + 771$ s, the count rate of the PC data becomes weak due to the rapid temporal decline and we have only upper limits (see Fig. 2). All the PC data were taken after the two flares described in the previous section. However, the data we have for the first flare were successively taken in IM, PuPD and LrPD modes. We do not have enough counts in the photo-diode modes to properly extract a spectrum, therefore no spectral information could be determined for this flare.

All the spectra were binned at 15 cts bin^{-1} within *XSPEC* v11.3.1 (Arnaud 1996). The WT mode spectrum and the “early”

and “late” PC mode spectra were fit by an absorbed power-law taking an excess absorption column (defined in the observed frame) ΔN_{H} into account. We used two photo-electric absorption (PHABS) models within *XSPEC*: one fixed to the value of the Galactic absorption column and the other kept free.

The fitted parameters for the WT spectrum are given in Table 2 as case 2. Note that the WT spectrum included the second X-ray flare. We also divided the WT data into two segments, to look for spectral softening during the flare, as reported by Burrows et al. (2005b): obtaining one spectrum before the second flare (bf) with an exposure of 27 s, and one containing the flare (f) with an exposure of 29 s. The fits suggest that the spectral slope of the flare ($\beta_{\text{f}} = 0.4 \pm 0.3$) is steeper than that of the WT data before the flare ($\beta_{\text{bf}} = 0.0 \pm 0.3$) using a fixed value of $\Delta N_{\text{H}} = 4.4 \times 10^{21} \text{ cm}^{-2}$. However, the large errors in the slopes do not require a spectral variation.

To determine whether the early X-ray emission is a continuation of the Gamma-ray emission, we fit the BAT and WT spectra (see Fig. 3) with a single absorbed power-law, but with different normalisation to take into account the temporal decay between the two spectra. We fixed the ΔN_{H} value to that found for the WT spectra. A single absorbed power-law gives a good fit (see case 3 in Table 2), consistent with the hypothesis that the X-ray and Gamma-ray emissions were produced by the same emission mechanism.

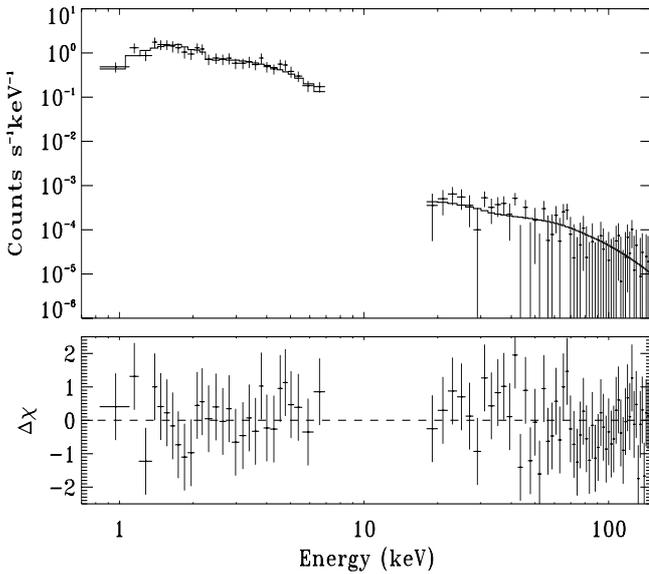
The parameters of the PC spectra are not well constrained using the χ^2 statistic, due to the low number of counts in the spectra (less than 80 counts). Therefore, we fit the “early” and “late” PC spectra using the Cash statistic (Cash 1979). Both PC spectra were binned to 1 cts bin^{-1} . The results are summarised in Table 2 as case 4a. Note that the parameter uncertainties are large due to the low statistics in the PC spectra.

² The five fainter bursts in this sample are 2 short bursts (GRB 050509b, Gehrels et al. 2005; GRB 050813, Fox et al. 2005), GRB 050815, the XRF 050406 (Romano et al. 2006) and GRB 050223 (Page et al. 2005).

Table 2. Summary of the Gamma-ray and X-ray spectral fitting parameters.

Case	Spectrum	$\Delta N_{\text{H}}(z=0)$ ($\times 10^{21}$ cm $^{-2}$)	β	$F_{2-10\text{keV}}^a$	$\chi^2(\text{d.o.f.})$
1	BAT	–	0.7 ± 0.5	–	54 (56)
2	WT	$4.4^{+4.0}_{-3.5}$	0.1 ± 0.3^d	$3.1^{+0.5}_{-0.8} \times 10^{-10}$	13 (18)
3	BAT-XRT/WT	4.4 fixed	0.2 ± 0.2	–	57 (76)
4a	PC “early”	$6.9^{+5.9}_{-3.2}$	$1.0^{+0.5}_{-0.8}$	$7.6^{+1.4}_{-2.7} \times 10^{-11}$	167^b (174) c
4a	PC “late”	$3.8^{+3.6}_{-3.2}$	$1.5^{+0.8}_{-0.7}$	$4.9^{+1.6}_{-2.2} \times 10^{-12}$	167^b (174) c
4b	PC “early”	4.4 fixed	$1.2^{+0.3d}_{-0.1}$	$5.9^{+2.0}_{-1.2} \times 10^{-11}$	148^b (174) c
4b	PC “late”	4.4 fixed	$1.2^{+0.3d}_{-0.1}$	$6.9^{+2.5}_{-2.1} \times 10^{-12}$	148^b (174) c

^a Observed fluxes in units of erg cm $^{-2}$ s $^{-1}$. ^b Cash statistic value. ^c PHA bins. ^d The uncertainties at a 99% confidence level for one parameter of interest ($\Delta\chi^2 = 6.63$) are $\beta = 0.1^{+0.6}_{-0.5}$ and $\beta = 1.2^{+0.5}_{-0.3}$ respectively for WT and PC spectra.

**Fig. 3.** Spectra from the BAT (on the right) and the WT data (on the left) fit together with an absorbed power-law ($\beta = 0.2 \pm 0.2$).

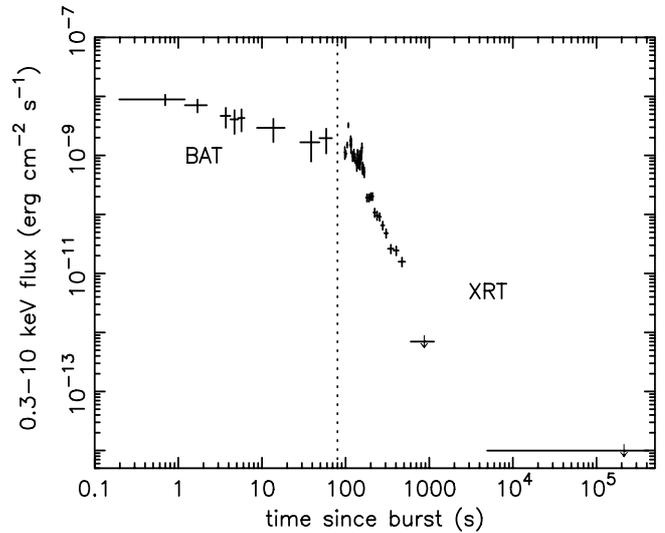
To better constrain the errors, we fit the PC spectra this time assuming the same spectral index β and ΔN_{H} values. Indeed, the β and ΔN_{H} values given in case 4a in Table 2 are consistent. We found a value of $\Delta N_{\text{H}} = 4.0^{+1.0}_{-1.2} \times 10^{21}$ cm $^{-2}$, consistent with that found for the WT spectrum. We then fixed the value of $\Delta N_{\text{H}} = 4.4 \times 10^{21}$ cm $^{-2}$ and estimated the uncertainties on the β spectral slopes (see case 4b in Table 2).

The values of the excess absorption column would be higher in the rest frame of the burst. We tried to constrain the redshift and ΔN_{H} in the rest frame of the burst using the absorption models taking into account the redshift within `xSPEC`. However, no useful constraints on these two parameters could be found.

From the best fits shown in Table 2 (cases 2 and 4b), it appears that GRB 050421 softens between the two time intervals $T_0 + 115 \rightarrow 171$ s and $T_0 + 171 \rightarrow 771$ s. If the spectral softening is real, the spectral evolution seen in our data could be due to a shift of the peak of the prompt emission down to the lower energies (Ford et al. 1995).

3.3. Combined Gamma-ray and X-ray light-curve

We used the slope of the joint BAT and XRT/WT spectra to extrapolate the BAT light-curve into the XRT 0.3–10 keV

**Fig. 4.** Combined light-curves of the BAT data extrapolated in the 0.3–10 keV energy band and the XRT data in the same band. The dotted vertical line indicates the separation between the BAT and XRT data. The unabsorbed fluxes are given in units of erg cm $^{-2}$ s $^{-1}$. The upper limits are given at 3σ .

energy range. Figure 4 shows the BAT-XRT light-curve (in units of erg cm $^{-2}$ s $^{-1}$). For the first flare, peaking at 110 seconds, we supposed what its spectral slope is the same as that for the combined XRT-BAT spectrum.

Note that the end of the BAT light-curve matches well with the beginning of the XRT light curve. The time of the transition from the gamma-ray emission to the early X-ray decay is $\sim 9\text{--}11 \times T_{90}$. From the first LrPD data, it is not completely clear to see if the decay start before or after the first flare. The T_{90} value found for this burst could then underestimate the real duration of this burst, due to skew in instrument sensitivity.

3.4. Energy of the X-ray flares

It is difficult to estimate the fluence of the first flare, because its shape is poorly known due to the lack of data (see Sect. 3.1.2). Using the XRT/BAT spectral index ($\beta \sim 0.2$) and assuming a Gaussian shape for the flare, we found a fluence of $3.1^{+5.2}_{-1.5} \times 10^{-8}$ erg cm $^{-2}$ in the XRT energy band, which corresponds to 20–75% of the GRB fluence in the 15–150 keV band.

During the second flare, the fluence in the 0.3–10 keV band was $6.1^{+4.9}_{-2.7} \times 10^{-9}$ erg cm $^{-2}$ (subtracting the underlying

power-law light-curve). This fluence corresponds to ~ 3 – 10% of the total fluence of the burst in the 15–150 keV band.

4. Discussion

4.1. Summary of the main results

The main results of the BAT and XRT temporal and spectral analysis of GRB 050421 are:

- The BAT light-curve showed one FRED peak with a duration of $T_{90} \sim 10.3$ s and a tail extending to $\sim T_0 + 70$ s. The burst is faint, with a fluence of $1.1 \pm 0.7 \times 10^{-7}$ erg cm $^{-2}$ in the 15–150 keV band. GRB 050421 is in the lowest 10% of GRBs detected by the BAT in term of fluence.
- The global XRT light-curve showed a rapid decline of the early X-ray emission ($\alpha = 3.1 \pm 0.1$). Due to the rapid decline, the X-ray emission was no longer detected beyond 1 h. The 3σ upper limit on the unabsorbed flux in the 0.3–10 keV band is 7×10^{-13} erg cm $^{-2}$ s $^{-1}$ at $\sim T_0 + 900$ s. The XRT light-curve also showed at least two flares peaking at $T_0 + 110$ s and $T_0 + 154$ s.
- There is a possible spectral softening from the early $T_0 + 115 - T_0 + 171$ s data with $\beta = 0.1 \pm 0.3$ to the later $\sim T_0 + 180 - T_0 + 771$ s data with $\beta = 1.2^{+0.3}_{-0.1}$.
- The joint fit of the BAT and early XRT data shows that they are likely to be produced by the same emission mechanism, which is consistent with an absorbed single power-law with a $\beta \sim 0.2$ slope.
- Excess absorption, $\Delta N_{\text{H}}(z = 0) \sim 4.4 \times 10^{21}$ cm $^{-2}$, over that due to our Galaxy is seen in the XRT spectra. No variation in absorption with time is seen.

4.2. The rapid decline of the X-ray emission

The rapid decline of the X-ray emission is unlikely to be produced by the emission from the forward shock. Indeed, the relations between α and β for a forward shock emission in the ISM and wind afterglow models, even taking the energy injection into account, are incompatible with the spectral and temporal characteristics of this burst (see Table 2 in Zhang et al. 2006).

It is interesting to consider the “high latitude emission” model (Kumar & Panaitescu 2000; Dermer 2004; Zhang et al. 2006). In this model, the tail of the prompt emission is produced at high angles relative to the observer direction (with angles $\theta > \Gamma^{-1}$). This radiation takes longer to reach the observer due to the curvature of the emitting shell. The curvature relation $\alpha = 2 + \beta$ can well explain the temporal decline we see after $\sim T_0 + 180$ s, since we found $\alpha = 3.1 \pm 0.1$ and $\beta_{\text{PC}} = 1.2^{+0.3}_{-0.1}$, which are consistent with the curvature relation. Moreover, using the $\alpha = 2 + \beta$ relation and the XRT/WT spectral slope ($\beta = 0.1 \pm 0.3$), we have $\alpha = 2.1 \pm 0.3$. This value of the decay slope is consistent with that derived in Sect. 3.1.2 from the fit of the WT part of the light-curve before $T_0 + 180$ s ($\alpha = 2.0 \pm 0.6$).

If the later temporal decay seen in several XRT bursts is linked to the forward shock emission (Nousek et al. 2006), then the transition between the tail of the prompt emission and the forward shock emission is not seen in our data. If a forward shock has been produced, then its emission would be very faint and below the XRT sensitivity threshold. A possible and simple explanation would be that the very local surrounding density is so low that the blastwave did not sweep up sufficient matter to produce a sufficiently bright afterglow.

Note that to date (before the end of September 2005) only 8 GRBs detected by the BAT (GRB 050416b, GRB 050418, GRB 050507, GRB 050502a, GRB 050709, GRB 050715, GRB 050911 and GRB 050925) were not detected by the XRT. In each case this was because the spacecraft did not slew or the slew was very delayed due to an observing constraint. Several authors have reported that the steep decline of the early part of the X-ray light-curve is due to high latitude emission e.g. GRB 050117a (Hill et al. 2005a) and GRB 050315 (Vaughan et al. 2006). Contrary to GRB 050421, the steep decline seen in GRBs 050117 and 050315 was followed by a shallower decline, and later by another steep decline. Thus, GRB 050421 appears unique so far in apparently not showing emission from the forward shock.

4.3. The X-ray flares as a signature of a late activity of the central engine ?

The short timescales $\delta t/t_{\text{peak}} \ll 1$ for the two flares, the large flux variations ($\delta F/F_{\text{peak}} \sim 3.7$) for the first flare, and the fact that the two flares would reasonably be interpreted by a model with the same underlying physics rule out most of the models based on the external shocks (Zhang et al. 2006). Moreover, Ioka et al. (2005) have shown that simple kinematic arguments can give limits on the timescale and amplitude variabilities in the X-ray afterglows for different cases. The timescale and amplitude variabilities for the two X-ray flares seen in GRB 050421 are located in the upper part of their Fig. 1, and therefore violate the cases b (ambient density fluctuations) and c (patchy shells) defined in Ioka et al. (2005).

When a blastwave is slowed down by the external medium, forward and reverse shocks are expected to be produced. Usually, optical, UV and radio emission are expected when the reverse shock crosses the external medium (e.g. Kulkarni et al.; Akerlof et al. 1999). Kobayashi et al. (2005) have shown that a strong flare could be produced in the X-ray band, if Synchrotron Self Compton (SSC) emission is invoked. This case would require some carefully balanced conditions. Even so it is still difficult to explain the rising and decaying slopes of the first flare, and how the reverse shocks could produce the two flares.

An alternative is then to invoke internal shocks. Assume that the central engine emits, in a single event, matter with a range of Lorentz factors, Γ . When the lower Γ matter catches up with the decelerated high Γ matter, the lower Γ matter would inject additional energy in the blastwave and a flare could be produced (Ramirez-Ruiz et al. 2002). The decaying timescale ratio in this model is $\delta t_{\text{decay}}/t_{\text{peak}} \sim 1$ and the flux level does not resume the previous level due to the additional energy injected (Zhang et al. 2006). Both of these predicted behaviours are inconsistent with our data. Moreover, the timescale and amplitude variabilities for the two X-ray flares violate the case d (refreshed shocks) defined in Ioka et al. (2005).

A more likely possibility is that the two X-ray flares are produced by late internal shocks. This explanation has already been invoked to explain the X-ray flares in several papers (e.g. Burrows et al. 2005; Falcone et al. 2006; King et al. 2005; Romano et al. 2006). The two X-ray flares of GRB 050421 are within the known range of the GRB durations (Paciesas et al. 1999).

5. Conclusion

To date, GRB 050421 was one of the faintest bursts observed by *Swift*. It was only observed in the X-ray and Gamma-ray

bands. It had a rapid temporal decline of the X-ray emission on which were superposed at least two X-ray flares. We have argued that the flares are probably produced by late internal shocks. The spectral softening seen in the X-ray data is likely due to the shift of the peak of the prompt spectrum in the XRT energy band. We have argued that the temporal decline corresponds to the tail of the prompt emission, which is produced by high latitude emission. Beyond 1 h after the BAT trigger, the source was undetectable. The 3σ upper limit on the unabsorbed flux in the 0.3–10 keV band is $<7 \times 10^{-13}$ erg cm $^{-2}$ s $^{-1}$ at $T_0 + 900$ s.

If the explanation above is correct, then GRB 050421 would be the first burst to which the XRT slewed promptly, for which no forward shock was observed. Therefore, GRB 050421 would be the first “naked burst” detected by *Swift*.

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