

Radio observations and spectral index study of the SNR G126.2+1.6

W. W. Tian^{1,2} and D. A. Leahy²

¹ National Astronomical Observatories, CAS, Beijing 100012, PR China
e-mail: tww@bao.ac.cn

² Department of Physics & Astronomy, University of Calgary, Calgary, Alberta T2N 1N4, Canada

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ABSTRACT

We present new images of the low radio surface brightness Supernova Remnant (SNR) G126.2+1.6, based on the 408 MHz and 1420 MHz continuum emission and the HI-line emission data of the Canadian Galactic Plane Survey (CGPS). We find the SNR's flux densities at 408 MHz (9.7 ± 3.9 Jy) and 1420 MHz (6.7 ± 2.1 Jy) which have been corrected for flux densities from compact sources within the SNR. We list properties of the 19 brightest compact sources within G126.2+1.6 at both 408 MHz and 1420 MHz. The integrated flux density based spectral index ($S_\nu \propto \nu^{-\alpha}$) is 0.30 ± 0.41 . The respective T-T plot spectral index is 0.30 ± 0.08 . We also find spatial variations of spectral index within the SNR: 0.2–0.6. HI observations show structures probably associated with the SNR, i.e., features associated with the SNR's southeastern filaments in the radial velocity range of -33 to -42 km s⁻¹, and with its northwestern filaments in -47 to -52 km s⁻¹. This association suggests a distance of 5.6 kpc for SNR G126.2+1.6. The estimated Sedov age for G126.2+1.6 is less than 2.1×10^5 yr.

Key words. ISM: individual objects: G126.2+1.6 – ISM: supernova remnants

1. Introduction

As the third of a continuing study of supernova remnants spectral index variation (Tian & Leahy 2005; Leahy & Tian 2005), the radio spectrum of the SNR G126.2+1.6 is studied in detail in the paper. This low radio surface brightness source was first detected and classified as a SNR by Reich et al. (1979). Some of its basic physical features are unclear, such as its distance and radio spectrum. Flux densities of G126.2+1.6 and respective spectral indices have been given previously (Reich et al. 2003; Joncas et al. 1989; Fürst et al. 1984) based on observations with a lower sensitivity than the current observations. Previous studies show that better flux density measurements are needed for G126.2+1.6. In the paper, we present the SNR's continuum images at higher sensitivity than previously at 408 MHz and 1420 MHz in order to determine its flux densities and spectral index. We investigate HI-line emission at various radial velocities for detecting interactions of the remnant with the surrounding gas and estimating its distance and age.

2. Observations and analysis

The continuum and HI emission data sets come from the CGPS, which is described in detail by Taylor et al. (2003). The data sets are mainly based on observations from the Synthesis Telescope (ST) of the Dominion Radio Astrophysical

Observatory (DRAO). The spatial resolution of the continuum images is better than $1' \times 1'$ cosec(δ) (HPBW) at 1420 MHz and $3.4' \times 3.4'$ cosec(δ) at 408 MHz. The synthesized beam for the HI line images is as the same as for the continuum and the radial velocity resolution is 1.32 km s⁻¹. DRAO ST observations are not sensitive to structures larger than an angular size scale of about 3.3° at 408 MHz and 56' at 1420 MHz. Thus the CGPS includes data from the 408 MHz all-sky survey of Haslam et al. (1982), sensitive to structure greater than 51', and the Effelsberg 1.4 GHz Galactic plane survey of Reich et al. (1990, 1997), sensitive to structure with resolution 9.4' for large scale emission (the single-dish data are freely available by <http://www.mpi.fr-bonn.mpg.de/survey.html>). The low-order spacing HI data is from the single-antenna survey of the CGPS area (Higgs & Tapping 2000) with resolution of 36'. See Taylor et al. (2003) for detail of the method of combining the synthesis telescopes and single dish observations.

We analyze the continuum and HI images of G126.2+1.6 and determine its flux densities and distance using the DRAO export software package. For G126.2+1.6, integrated flux density's errors are found by comparing results for several different choices of background region. For compact sources, the flux density's errors are taken as the formal Gaussian fit errors. The influence of compact sources within the SNR is reduced by employing similar methods to Tian & Leahy (2005).

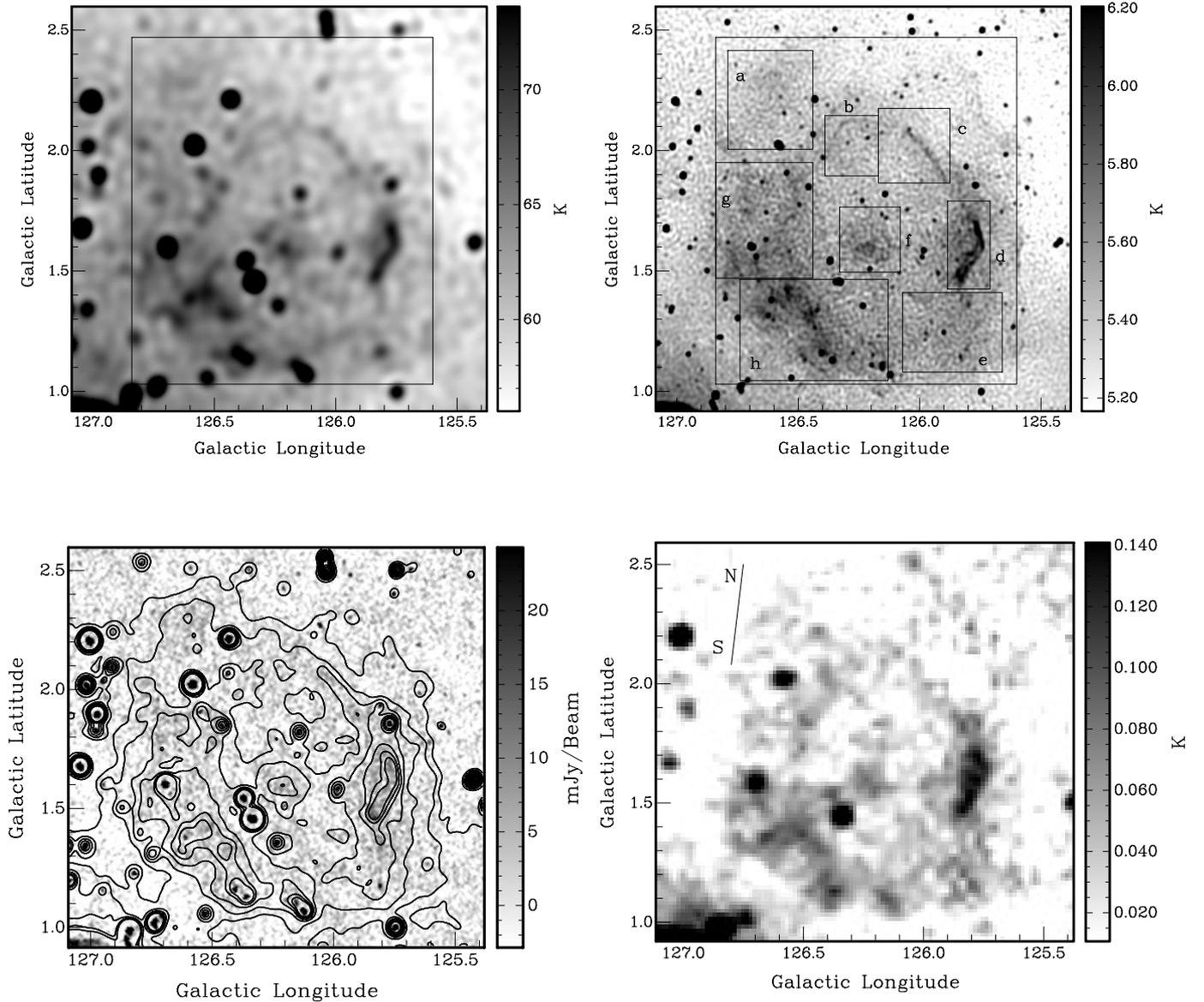


Fig. 1. The first row of images shows the CGPS maps at 408 MHz (*left*) and 1420 MHz (*right*). The left of the second row shows the WENSS image at 327 MHz (grey scale) with contours from the 1420 MHz map convolved to the same resolution as the 408 MHz resolution (contours at 5.3, 5.4, 5.5, 5.6, 5.7, 5.8, 5.9 K). The right of the second row is the 2695 MHz Effelsberg map. The box used for whole SNR T-T plots is shown in the *upper left*. The 8 boxes, labeled with letters and used for SNR sub-areas T-T plots, are shown in the *upper right*. The direction of North (N) and South (S) is marked on the *lower right* image.

3. Results

3.1. Continuum emission from G126.2+1.6

The CGPS continuum images at 408 MHz and 1420 MHz are shown in the upper left and right panels of Fig. 1. The lower left shows the WENSS (the Westerbork North Sky Survey, beam size $54'' \times 54'' \text{ cosec}(\delta)$, Rengelink et al. 1997) map at 327 MHz (grey scale) with contours from the 1420 MHz map convolved to the same resolution as the 408 MHz map. The 2695 MHz Effelsberg map is reproduced in the lower right for reference (Fürst et al. 1990). The Effelsberg map has an resolution of $4.3'$ and a sensitivity of $50 \text{ mK } T_B$.

The 408 MHz image for G126.2+1.6 reveals clearly more structure in comparison with Joncas's (1989) 408 MHz image, including a ring-like filamentary structure. Sharp and diffuse filamentary emissions are detected around the limb of G126.2+1.6, especially the bright western and northwestern sharp filaments and south to southeastern diffuse filamentary emissions. An area located at the center of the SNR shows significant diffuse emission. Compact sources distributed across the face of the SNR area are prominent, and illustrate the importance of having high enough spatial resolution to distinguish between SNR and compact source emission. The 1420 MHz map is the first published at this frequency and the best so far at any frequencies for G126.2+1.6. It shows much better detail of the fine structure in G126.2+1.6 than any previous

Table 1. List of the brightest compact sources and their integrated Flux Densities (FD) inside G126.2+1.6.

Source num.	RA(2000)	Dec(2000)	GLONG	GLAT	FD at 408 MHz	FD at 1420 MHz	Sp. index
	[h m s]	[° ' "]	deg	deg	mJy	mJy	α
1	1 22 39.55	64 7 24.2	126.331	1.458	412 ± 15	153 ± 5	0.80 (0.76 to 0.84)
2	1 25 35.93	64 39 16.0	126.579	2.025	336 ± 11	132 ± 4	0.75 (0.72 to 0.79)
3	1 24 26.25	64 51 48.0	126.429	2.216	223 ± 7	63 ± 2	1.02 (0.98 to 1.05)
4	1 26 5.20	64 13 14.1	126.689	1.602	218 ± 7	91 ± 3	0.70 (0.66 to 0.74)
5	1 23 5.06	64 12 19.3	126.367	1.545	137 ± 9	54 ± 3	0.75 (0.69 to 0.82)
6	1 22 48.99	63 48 39.6	126.386	1.150	134 ± 9	82 ± 11	0.39 (0.28 to 0.52)
7	1 20 19.86	63 45 45.7	126.119	1.070	121 ± 5	38 ± 2	0.93 (0.88 to 0.98)
8	1 23 58.46	63 41 53.8	126.527	1.054	63 ± 4	22 ± 1	0.86 (0.80 to 0.94)
9	1 21 38.46	64 2 6.0	126.231	1.357	63 ± 4	16 ± 1	1.08 (1.00 to 1.16)
10	1 20 40.53	63 47 47.8	126.153	1.108	58 ± 5	26 ± 2	0.65 (0.57 to 0.74)
11	1 17 53.45	64 34 47.7	125.767	1.854	54 ± 5	32 ± 3	0.43 (0.33 to 0.54)
12	1 19 37.81	64 17 3.8	125.985	1.580	49 ± 5	19 ± 2	0.75 (0.64 to 0.86)
13	1 21 19.04	64 30 29.9	126.141	1.823	46 ± 4	20 ± 1	0.66 (0.59 to 0.74)
14	1 22 31.42	63 47 37.6	126.356	1.129	37 ± 8	16 ± 3	0.65 (0.46 to 0.92)
15	1 26 49.86	63 49 40.5	126.823	1.224	37 ± 3	15 ± 3	0.74 (0.59 to 0.92)
16	1 24 31.14	64 2 54.6	126.542	1.409	33 ± 5	12 ± 3	0.79 (0.58 to 1.08)
17	1 24 15.65	64 29 48.3	126.457	1.850	28 ± 6	21 ± 2	0.22 (0.04 to 0.44)
18	1 21 36.35	64 16 33.9	126.199	1.596	31 ± 4	10 ± 1	0.91 (0.72 to 1.09)
19	1 25 2.05	64 1 1.1	126.602	1.385	24 ± 4	13 ± 3	0.48 (0.29 to 0.73)

images. A double filament structure appears in the northeastern part of the SNR, and multi-filamentary structures appear in the southeastern part. The inside edge of the filament distribution has a diamond-type shape more than a ring-type shape. The outlines of G126.2+1.6 at 408 MHz and 1420 MHz are very similar. The lower left of Fig. 1 shows that main features of G126.2+1.6 in the 327 MHz WENSS map are consistent with the CGPS 408 MHz and 1420 MHz maps.

3.2. T-T Plot spectral indices

Bright compact sources affect the measured integrated flux densities for G126.2+1.6 and its measured spectral index. Thus we correct for the effects of compact sources. Table 1 lists properties of the 19 brightest compact sources which are detected within G126.2+1.6 at both 408 MHz and 1420 MHz.

First we discuss spectral indices between 408 MHz and 1420 MHz based on the T-T plot method (Turtle et al. 1962). The principle of the T-T plot method is that spectral indices ($T_\nu = T_0 \nu^{-\beta}$) are calculated from a fit of a linear relation to the $T_1 - T_2$ values of all pixels within a given map region. T_1 is the brightness temperature of a map pixel at one frequency and T_2 is for the second frequency. The higher resolution image has been smoothed to the lower resolution for the T-T plot comparison. The brightness temperature spectral index β is derived from the slope of the line. The error in spectral index is derived from the uncertainty in slope of the line. The flux density spectral index α ($S_\nu \propto \nu^{-\alpha}$) is related to β by $\beta = \alpha + 2$. Spectral

index refers to flux density spectral index α in this paper unless specifically noted otherwise.

For the T-T plot analysis, first a single region for the whole SNR is used, as shown in Fig. 1. This region yields the T-T plots shown in Fig. 2. Three cases are considered: using all pixels including compact sources; using all pixels after subtracting Gaussian fits to the compact sources listed in Table 1 from the images; and excluding compact sources. The compact sources are bright compared to the SNR emission. Since the compact sources have a steeper spectrum than the SNR, they are seen in the T-T plot (Fig. 2 left panel) as the steeper lines of points extending to higher T_B . Subtracting compact sources from the image before making the T-T plot removes the lines of points associated with the compact sources (Fig. 2 middle panel). However artifacts remain in the T-T plot due to imperfect source subtraction, especially at 408 MHz. The next step is to completely remove regions of pixels including the compact sources from the analysis. Each region is taken to be a few beamwidths across, so that any contribution from the compact source is below 1% of the diffuse SNR emission. Thus any artifacts associated with the compact source are also removed. Generally, the last step produces the best results (rightmost plot of Fig. 2). Therefore, in further discussion, the second case will be not considered.

Next, smaller areas (labeled a to h in the upper right panel of Fig. 1) are selected to search for spatial variations in spectral index. Table 2 lists the results for two cases of analysis: including compact sources and removing compact sources. Visual inspection of the T-T plots confirms that the second method

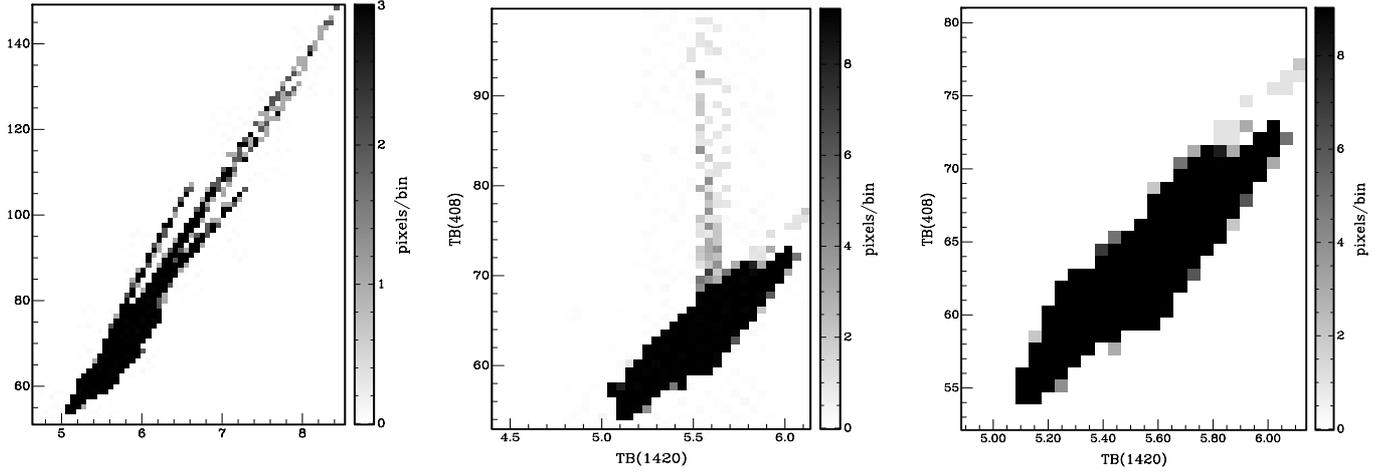


Fig. 2. Whole SNR 408-1420 MHz T-T plots. *From left to right:* plot for map including compact sources ($\alpha = 0.42 \pm 0.09$); plot for map with Gaussian fits to compact sources subtracted ($\alpha = 0.31 \pm 0.11$); plot for compact sources removed from analysis ($\alpha = 0.30 \pm 0.08$). Coordinates units of the left plots are same as the right plot.

Table 2. 408-1420 MHz T-T plot spectral indices with and without Compact Sources (CS).

Sp. index	α	α	α
Area	Including CS	CS removed	Manual fit
a	0.39 ± 0.33	0.09 ± 0.70	0.20
b	0.36 ± 0.23	0.36 ± 0.23	0.46
c	0.46 ± 0.09	0.46 ± 0.09	0.45
d	0.59 ± 0.02	0.59 ± 0.02	0.59
e	0.45 ± 0.55	0.30 ± 0.57	0.41
f	0.64 ± 0.10	0.59 ± 0.19	0.57
g	0.34 ± 0.14	0.28 ± 0.19	0.29
h	0.50 ± 0.14	0.29 ± 0.24	0.19
All G126.2+1.6	0.42 ± 0.09	0.30 ± 0.08	

Table 3. Integrated flux densities and spectral indices of G126.2+1.6 and compact sources within the SNR.

Freq. MHz	G126.2+1.6 Jy	CS of G126.2+1.6 Jy	G126.2+1.6 and CS Jy
408	9.7 ± 3.9	2.11 ± 0.12	11.8 ± 4.0
1420	6.7 ± 2.1	0.84 ± 0.06	7.5 ± 2.2
α	0.30 ± 0.41	0.74 ± 0.07	0.36 ± 0.35

produces the most reliable results. The compact sources' influence on the spectral index calculation is obvious in the T-T plots, and also seen in Table 2 for areas a, e, f, g and h. From now on we discuss spectral indices derived with compact sources removed, unless specified otherwise. From Table 2, we see spatial variation of spectral index within G126.2+1.6: 0.2–0.6. Most difference are of low significance; the most significant difference is between d and g (at 1.6σ).

Table 4. Integrated Flux Densities (FD) of G126.2+1.6 and of Compact Sources within G126.2+1.6 (CSFD).

Freq. MHz	HPBW arcmin	FD Jy	CSFD Jy	References for FD
83	59×31	35.0 ± 7.0	7.2 ± 0.4	Kovalenko et al. (1994)
408	3.4×3.8	$9.7 \pm 3.9^*$	2.1 ± 0.1	this paper
408	3.5×3.9	$12.0 \pm 2.5^*$		Joncas et al. (1989)
865	14.5×14.5	7.1 ± 1.6	1.2 ± 0.1	Reich et al. (2003)
1420	1×1.1	$6.7 \pm 2.1^*$	0.84 ± 0.06	this paper
2695	4.4×4.4	4.4 ± 0.4		Fürst et al. (1984)
2695	4.3×4.3	3.7 ± 0.4	0.52 ± 0.06	see text
4850	2.6×2.6	$1.7 \pm 0.2^{**}$	0.34 ± 0.07	Fürst et al. (1984)

* CS have been subtracted, ** lower limit.

3.3. Integrated flux densities and spectral indices

We have derived integrated flux densities of G126.2+1.6 from the 408 MHz and 1420 MHz maps. Values given have diffuse background subtracted. The resulting 408 MHz to 1420 MHz spectral index, using flux densities without compact sources, is 0.30 ± 0.41 . Table 3 lists the flux densities and spectral indices of G126.2+1.6 and the compact sources within G126.2+1.6. Compact sources contribute about 18% at 408 MHz and 13% at 1420 MHz to the SNR's flux densities, and have a significant effect on the spectral index. It is noted that the whole SNR spectral index derived from integrated flux densities is consistent with the whole SNR spectral index (0.30 ± 0.08) derived by the T-T plot method.

Published integrated flux densities and errors for the SNR at other frequencies are given in Table 4. The errors are the published errors and don't include any uncertainty due to possible differences in flux density scales. We have calculated total compact source flux densities for these other frequencies, using the 408–1420 MHz spectral index upper and lower limits and flux densities from Table 1. The compact sources' flux density

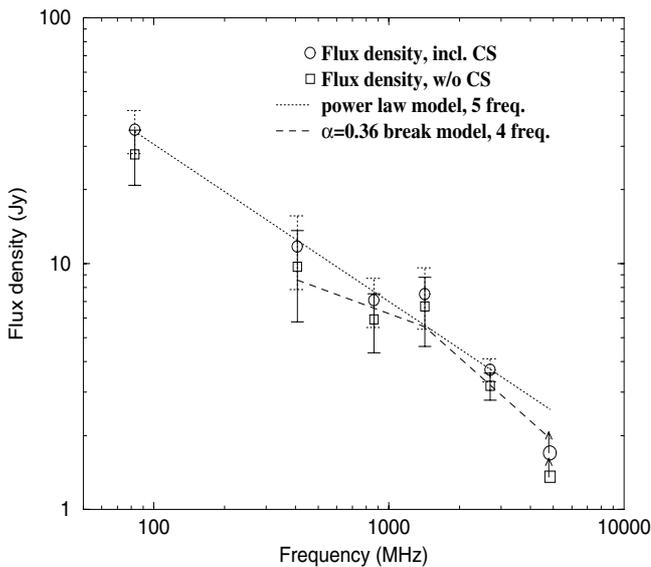


Fig. 3. Radio spectrum of G126.2+1.6. The value at 4850 MHz is a lower limit. The power law model has a best-fit spectral index of 0.62, $\chi^2 = 0.99$, for a fit to 5 frequencies 83–2695 MHz. The synchrotron aging model has a best-fit spectral index of 0.36, break frequency of 1420 MHz and $\chi^2 = 0.6$, for a fit to 4 frequencies 408–2695 MHz.

contribution is the highest at 83 MHz at 21%. It decreases to 18% at 408 MHz, 17% at 865 MHz, 13% at 1420 MHz, 14% at 2695 MHz and 20% at 4850 MHz. We have recalculated the flux density values of G126.2+1.6 at 83 MHz, 865 MHz, 2695 MHz and 4850 MHz by subtracting the compact source flux density. For 408 and 1420 MHz the values in Table 4 already have compact source flux density removed. We note that Fürst et al. (1984) gave a flux density based on the 2965 MHz Effelsberg image with resolution 4.4×4.4 arcmin, but we obtain a new value from the Effelsberg 2695 MHz image with a little higher resolution 4.3×4.3 arcmin: 3.7 ± 0.4 Jy and use this instead. We first fit the resulting flux density values with a power-law to obtain spectral index. Because the flux density at 4850 MHz is a lower limit value, we do not include it in our fits. Figure 3 shows the corrected flux densities and the best-fit power-law. For the lowest 5 frequencies, there is a best fit spectral index of $\alpha = 0.62$ with $\chi^2 = 0.99$. If we consider a synchrotron aging model for the G126.2+1.6 radio spectrum, the spectrum is steeper by 0.5 above the break frequency. The best fit to the lowest frequencies except 83 MHz has $\alpha = 0.36$, $\chi^2 = 0.6$ and a break frequency of 1420 MHz. A detail analysis about the spectrum fits has been given in discussion section.

3.4. HI emission

We have searched the CGPS radial velocity range for features in the HI which might relate to the morphology of G126.2+1.6. There is emission which is coincident with the boundary of G126.2+1.6 in the velocity range -32 to -52 km s $^{-1}$, and only in this velocity range. From -32 to -42 km s $^{-1}$, the HI emission is adjacent to the south and southeast edge of G126.2+1.6,

and from -47 to -52 km s $^{-1}$, the HI emission is adjacent to the west edge. The spatial relation with the edge of G126.2+1.6 indicates that the HI is probably associated with the supernova remnant. These HI observations show a good association of HI features with G126.2+1.6, similar to other accepted HI associations with supernova remnants and with similar HI velocity ranges (~ 20 km s $^{-1}$), e.g. HI associated with the Cygnus Loop (Leahy 2003) and with DA530 (Landecker et al. 1999).

Figure 4 shows maps of HI emission in two channels (the left and middle panels are for radial velocities -32.8 and -51.8 km s $^{-1}$ respectively) and also integrated over channels from -32.8 to -51.8 km s $^{-1}$ (the right panel). Each map has superimposed on it contours of continuum emission at 1420 MHz chosen to show G126.2+1.6.

4. Discussion

4.1. The distance and age to G126.2+1.6

There are no previous reliable distance estimates for G126.2+1.6. Because of the low surface brightness of G126.2+1.6 and due to presence of compact sources, its flux density was measured with large errors. The surface brightness-diameter relations only gives a rough distance estimates (3.6 kpc to 7.1 kpc, Joncas et al. 1989). Joncas et al. (1989) discussed the distance in detail, suggested a range of 2 to 5 kpc, and noted that a reliable distance may be obtained based on a future measure of systemic radial velocity of the remnant, either from the optical lines or from detection of an interaction with the surrounding HI. The HI yields a radial velocity of G126.2+1.6 of -42 km s $^{-1}$. We take a galactic rotation curve with $R_0 = 8.5$ kpc, $V_0 = 220$ km s $^{-1}$, and $V = 250$ km s $^{-1}$ for $R > 1.4 R_0$, and obtain a distance of 5.6 kpc for G126.2+1.6. The range of distances obtained for radial velocities between -39 km s $^{-1}$ and -45 km s $^{-1}$ is 5.3 to 5.9 kpc.

At 5.6 kpc distance, we get a mean radius of about $R = 64$ pc for G126.2+1.6. From the column density of the HI associated with G126.2+1.6, we estimate an approximate density $n_0 = 2 \times 10^{20}$ cm $^{-2}/(2R) = 0.5$ cm $^{-3}$. Applying a Sedov model (e.g. Cox 1972), for a typical explosion energy of $E = 0.5 \times 10^{51}$ erg ($\epsilon_0 = E/(0.75 \times 10^{51}$ erg) = 2/3), yields an age of 4.7×10^5 yr. This is too old since the SNR would have completely cooled by that time for $n_0 = 0.5$ cm $^{-3}$. The shock radius when the SNR has cooled is $R_s^{(c)} = 24.3(\epsilon_0/n_0)^{5/17}n_0^{-2/17}$ pc, and we require $R \leq R_s^{(c)}$. Thus we find an upper limit on n_0 is 0.1 cm $^{-3}$. This is consistent with the HI column density if the SNR exploded in a low density cavity and we are seeing the surrounding high density shell, as one sees for the Cygnus Loop (Leahy 2003). The revised Sedov age for G126.2+1.6 with $n_0 < 0.1$ cm $^{-3}$ is less than 2.1×10^5 yr. We also consider a post-Sedov model. In this case (Cox 1972), shocked gas cools and joins a dense shell which expands slowly, ~ 30 km s $^{-1}$, driven by the hot interior of the supernova remnant. However, for $n_0 = 0.5$ cm $^{-3}$ and the observed size of G126.2+1.6, the age of the supernova remnant is $\sim 10^6$ yr. This is too large, and requires that the pre-SNR density must be much lower. So again the conclusion is that the supernova exploded in a low density cavity in the interstellar medium, similar to the Cygnus Loop.

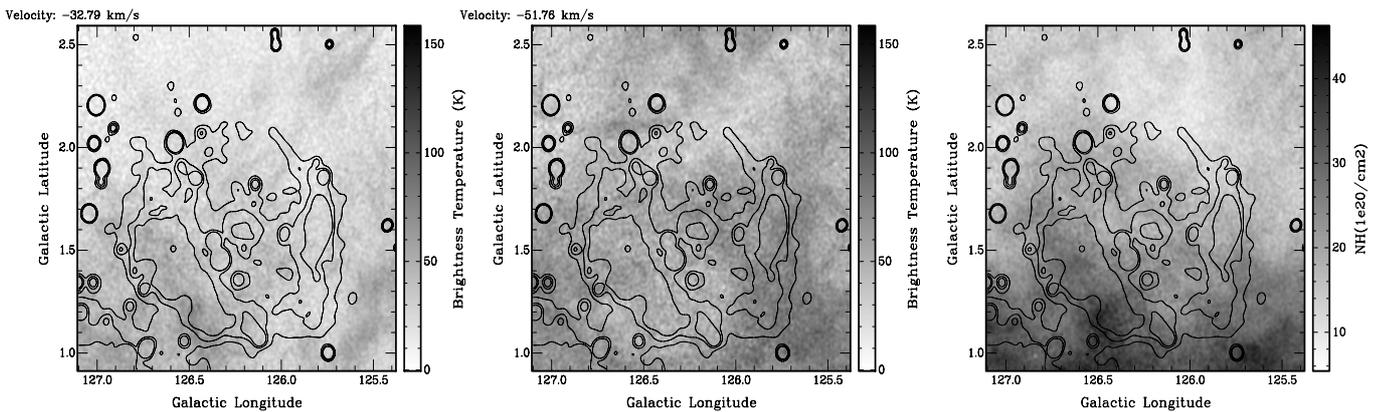


Fig. 4. Grey scale representation of HI emission in the field centered on G126.2+1.6. The radial velocity of the left and middle plots is indicated by the number at its top left corner. The grey scale bar of the left and middle plots has units of brightness temperature, of the right has units of 10^{20} atom cm^{-2} . Continuum emission at 1420 MHz is indicated by contours (5.5, 5.6, 5.7, 5.8, 5.9 K).

4.2. Radio spectrum

Generally, in the spectrum plots (Fig. 3), the 1420 MHz flux density appears high and the 83 MHz flux density appears high. Since our 1420 MHz map is the best so far for G126.2+1.6 the 1420 MHz one is the most reliable flux density. So we suspect that due to the faintness of the outer regions of G126.2+1.6, there is difficulty in determining the background level when the spatial resolution is poor. This results in increasing the uncertainty in flux densities.

The radio spectrum shown in Fig. 3 has a best-fit power-law with $\alpha = 0.62$. This is much higher than the well-determined 408–1420 MHz T-T plot spectral index and the flux density-based values (both about 0.30). Previous spectrum fits ($\alpha = 0.48$, Fürst et al. 1984; $\alpha = 0.58$, Joncas et al. 1989; $\alpha = 0.59$, Kovalenko et al. 1994; $\alpha = 0.77$, Trushkin 2002) are not reliable due to inclusion of compact sources and poorer spatial resolution. For example we obtain $\alpha = 0.36$ for 408–1420 MHz spectral index if we include compact sources (see Table 3).

Here are some alternate explanations of the observed multi-frequency radio spectrum. One alternative is that the 83 MHz observation has too high a flux density. If we omit the flux density at 83 MHz, we obtain a power-law fit with $\alpha = 0.61$ ($\chi^2 = 0.95$). This still disagrees with the 408–1420 MHz spectral index.

Another alternative is that the radio spectrum is described by a synchrotron aging model with resulting break in the radio spectrum. Fürst et al. (1984) gave a lower limit flux density at 4850 MHz and argued the spectrum of G126.2+1.6 is steepening above 1 GHz. If we include all frequencies data (except the 4850 MHz lower limit) we obtain $\alpha = 0.59$, $\chi^2 = 0.95$ and break frequency 2240 MHz. If we use the synchrotron aging model and omit the flux density at 83 MHz, we obtain $\alpha = 0.36$, $\chi^2 = 0.56$ and break frequency 1420 MHz. This is consistent with the 408–1420 MHz index and is a better fit than the pure power-law model.

The last alternative we discuss, is that the curved spectrum is partly due to the spectral index variations within

G126.2+1.6. However, when we explicitly include the flux density from regions c and d, with their known 1420 MHz flux density and known spectral index, we do not obtain significantly different results than above for either the power-law or synchrotron aging models.

In summary, we can reconcile the 0.30 ± 0.08 T-T plot spectral index with flux density-fit spectral index in two ways. (a) Including the 83 MHz data, the radio spectrum is fit by a break spectrum with break frequency 1.6 MHz, and low frequency index of $\alpha = 0.4$. This model is 1.3σ worse than the synchrotron aging model with free α ; (b) omitting the 83 MHz data, the radio spectrum is best fit by a model with break at frequency 1.4–1.5 GHz with low frequency index $\alpha = 0.36$.

5. Conclusion

We present high resolution and high sensitivity images of G126.2+1.6 at 1420 MHz and 408 MHz and show new HI-line emission data of G126.2+1.6 in this paper. The 1420 MHz continuum image is the best image of G126.2+1.6 at any frequency yet. Flux densities at 408 MHz and 1420 MHz are obtained. The T-T plot spectral index is 0.30 ± 0.08 and the flux density-based spectral index is 0.30 ± 0.41 between 408 MHz and 1420 MHz. By our analysis, a multi-frequency spectral index of about 0.36 is obtained using a synchrotron aging model. There is evidence at 1.6σ for spatial variations in spectral index in G126.2+1.6 between about 0.3 and 0.6. The association between HI brightness features and the SNR's structure suggests a distance of 5.6 kpc for G126.2+1.6. The estimated Sedov age for G126.2+1.6 is less than 2.1×10^5 yr.

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References

- Cox, D. 1972, *ApJ*, 178, 159
- Fürst, E., Reich, W., & Steube, R. 1984, *A&A*, 133, 11
- Fürst, E., Reich, W., Reich, P., & Reif, K. 1990, *A&AS*, 85, 691
- Haslam, C. G. T., Salter, C. J., Stoffel, H., & Wilson, W. W. 1982, *A&AS*, 47, 1
- Higgs, L. A., & Tapping, K. F. 2000, *AJ*, 120, 2471
- Joncas, G., Roger, R. S., & Dewdney, P. E. 1989, *A&A*, 219, 303
- Kovalenko, A. V., Pynzar, A. V., & Udal'tsov, V. A. 1994, *AR*, 38, 95
- Landecker, T. L., Routledge, D., Reynolds, S. P., et al. 1999, *ApJ*, 527, 866
- Leahy, D. A. 2003, *AJ*, 586, 224
- Leahy, D. A., & Tian, W. W. 2005, *A&A*, 440, 929
- Reich, W., Kallas, E., & Steube, R. 1979, *A&A*, 78, 13
- Reich, W., Reich, P., & Fürst, E. 1990, *A&AS*, 83, 539
- Reich, W., Reich, P., & Fürst, E. 1997, *A&AS*, 126, 413
- Reich, W., Zhang, X., & Fürst, E. 2003, *A&A*, 408, 961
- Rengelink, R. B., Tang, Y., de Bruyn, A. G., et al. 1997, *A&AS*, 124, 259
- Taylor, A. R., Gibson, S. J., Peracaula, M., et al. 2003, *AJ*, 125, 3145
- Tian, W. W., & Leahy, D. A. 2005, *A&A*, 436, 187
- Trushkin, S. A. 2002, *BSAO*, 46, 62, <http://cats.sao.ru>
- Turtle, A. J., Pugh, J. F., Kenderdine, S., & Pauliny-Toth, I. I. K. 1962, *MNRAS*, 124, 297