

Dust temperature and density profiles in the envelopes of AGB and post-AGB carbon stars from mid-infrared observations[★]

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Abstract. First mid-infrared images of a sample of AGB and post-AGB carbon stars (V Hya, IRC +10216, CIT 6 and Roberts 22) obtained at the ESO-La Silla Observatory (Chile) are reported. CIT 6 presents a cometary-like feature clearly seen in the 9.7 μm image, Roberts 22 shows an envelope slightly elongated in the north-east direction while images of V Hya and IRC +10216 are roughly spherically symmetric. Using an inversion technique, the dust emissivity was derived from the observed intensity profiles, allowing a determination of the grain temperature and density distributions inside the envelope for these stars. Dust masses and mass-loss rates were estimated for V Hya and IRC +10216. Our results are comparable to those obtained in previous studies if dust grains have dimensions in the range $\sim 0.01\text{--}0.2 \mu\text{m}$. Color maps suggest the presence of temperature inhomogeneities in the central regions of the dust envelopes. In the case of V Hya, an eccentric hot point, whose direction coincides with the jet previously seen in [S II] emission, suggests that we are observing material ejected in a previous mass-loss event. Bipolar lobes are clearly seen in the color maps of Roberts 22 and IRC +10216.

Key words. stars: AGB and post-AGB – stars: circumstellar matter – stars: mass-loss – stars: imaging – infrared: stars

1. Introduction

According to our current understanding of stellar evolution, all stars with main sequence progenitors in the mass range $1\text{--}8 M_{\odot}$ evolve via the asymptotic giant branch (AGB) phase to the planetary nebula (PN) stage. The early AGB (E-AGB) phase begins as soon as the star finishes burning He in its core and begins to burn He in a thick shell. When both H and He burning shells are active, the star enters the thermally pulsating AGB (TP-AGB) phase (Iben & Renzini 1983; Blöcker 1999).

The AGB phase is characterized by a strong mass-loss period that only stops when the outer envelope is completely lost (Habing 1996). As the star ascends the AGB, the mass-loss rate increases from solar-like values ($10^{-14} M_{\odot}/\text{yr}$) up to $10^{-4} M_{\odot}/\text{yr}$, the so-called super-wind phase. Dust grains and molecules, predominantly CO, are formed in the wind, producing a circumstellar envelope detectable in the infrared and millimetric domains. Although the radiation pressure on dust grains and on molecules is thought to play an important role in the driving mechanism of the wind, different physical aspects governing the dust formation/destruction processes are still unclear. Processed elements are dredged-up into the atmosphere as a consequence of the re-ignition of the He burning shell (E-AGB phase). A further contraction leads also to the

re-ignition of the H burning shell and the star enters a TP-AGB phase. A third dredge-up may take place during this phase and overshooting will contribute to lift carbon to the surface. These events lead to the formation of circumstellar envelopes whose chemical composition reflects that of an atmosphere enriched with nuclear processed matter. C-rich stars have envelopes composed mainly of amorphous carbon while O-rich stars have circumstellar envelopes rich in silicates. Both oxidic and carbonaceous dust particles are characterized by vibrational bands positioned in the mid-IR window (at 9.7 μm and 18 μm for silicates and at 11.3 μm for SiC) (Henning 1999). Most of the thermal radiation continuum from the dusty envelope is emitted in this wavelength range.

Mid-infrared imaging allows a direct probe of the mass distribution through thermal emission arising from dust grains. Therefore a number of imaging studies of AGB and post-AGB stars have been made at mid-infrared wavelengths and some emission regions have been resolved (Dayal et al. 1998; Meixner et al. 1999; Ueta et al. 2001a,b; Jura et al. 2002a,b; Kwok et al. 2002; Close et al. 2003; Gledhill & Yates 2003).

In the present work we report mid-infrared imaging, with a spatial resolution of about $0.7''$, of a sample of AGB and post-AGB carbon-rich stars (CIT 6, V Hya, Roberts 22 and IRC +10216) not previously imaged at these wavelengths. Our goals are to derive physical properties of the dust inside the

[★] Based on observations made at La Silla ESO Observatory (Chile).

circumstellar envelope and to estimate mass-loss rates. Most of the objects have circumstellar envelopes that are nearly symmetric, allowing the use of inversion techniques to obtain the dust emissivity and, consequently, the density and the temperature as a function of the radial distance from the star. This paper is organized as follows: in Sect. 2, observations and data reduction are described; in Sect. 3, morphological analyses of the individual objects are presented; in Sect. 4, we describe our model for the envelope and present the derived dust density and temperature profiles; in Sect. 5 we discuss our results and compare them with previous works; in Sect. 6, the main conclusions are given.

2. Observations and data reduction

Observations were performed on 1995 February 5, 7 at the ESO 3.6-m telescope in La Silla (Chile), using the mid-infrared TIMMI camera (Lagage et al. 1993) with a detector consisting of a 64×64 pixel Ga:Si array. The resulting spatial scale of the system is $0.33''/\text{pix}$. The log of the observations is given in Table 1. These observations were made in a standard mid-IR observing mode, by chopping the secondary mirror and nodding the telescope to subtract the background emission from the sky and telescope. The chopper throw was $18.3''$ toward the south and the nod beam position used was $18.3''$ north of the first position. To avoid saturation of the detector by the ambient photon background and to have good image quality, each individual nod cycle was split into many short exposures of ~ 10 ms. This procedure was repeated for as many cycles as needed to obtain the required total integration time. Nearly diffraction-limited images ($\sim 0.7''$ FWHM for point sources) resulted from these short exposures. Filter wavelengths were selected to obtain information on the dust continuum, while maximizing the detection sensitivity. The observations were carried out in four narrow-band filters, centered respectively at 8.39 (N1), 9.7 (Silicates), 9.78 (N2), and $11.65 \mu\text{m}$ (SiC) whose characteristics are given in Table 2. During the different runs the seeing was typically $0.7''$ – $0.8''$.

Data reduction was performed using IDL self-developed routines. Individual chopped frames were spatially oversampled by a factor of 4 and shifted to the nearest 0.25 pixel by using a cross-correlation algorithm to correct for turbulent motions and flexure drifts. Images were then co-added to produce a single flat-field-corrected image, comprising the average of the chop and nod differences, for each filter. Reference stars were observed, analyzed in the same way and were used to derive the instrumental point spread function (PSF) at each filter. Images were deconvolved by using the Maximum Entropy Method (Hollis et al. 1992) and fluxes were obtained by interpolation of IRAS-LRS data. Table 3 gives adopted fluxes for each observed object.

3. Image analysis of individual objects

3.1. V Hya

V Hya (IRAS 10491-2059) is an N-type AGB carbon star, with an effective temperature of 2650 K and a C6 - 7.5e spectral type

Table 1. Log of the observations.

Object	Date (UT)	Filter
CIT 6	1995 Feb. 7	N1, Silicates
V Hya	1995 Feb. 7	N2, SiC
Roberts 22	1995 Feb. 5	N1, SiC
IRC +10216	1995 Feb. 5	N1, SiC

Table 2. Description of the filters used.

Filter	Central wavelength (μm)	Filter width (μm)
N1	8.39	7.91–8.87
Silicates	9.70	9.50–10.00
N2	9.78	9.14–10.43
SiC	11.65	10.3–13.00

Table 3. Adopted mid-infrared fluxes for observed objects.

Object	Filter	log flux (Jy)
CIT 6	N1	3.58
	Silicates	3.55
V Hya	N2	2.85
	SiC	2.87
Roberts 22	N1	1.89
	SiC	2.28
IRC +10216	N1	4.42
	SiC	4.63

(Knapp et al. 1997). It is a semi-regular variable, with a period of about 529 days (Lambert et al. 1986), surrounded by a bipolar envelope as shown by CO observations (Tsuji et al. 1988; Kahane et al. 1996) and near-infrared polarization data (Johnson & Jones 1991; Trammell et al. 1994). High resolution optical spectroscopy suggests that V Hya is in rapid rotation, which could be induced by the presence of a companion (Barnbaum et al. 1995). The large infrared excess (van der Veen & Habing 1988) as well as the strong emission in molecular lines indicate that V Hya is losing mass at a rate higher than $10^{-6} M_{\odot} \text{yr}^{-1}$. Knapp et al. (1997) detected, via CO observations, a fast molecular wind ($V_w \sim 200 \text{ km s}^{-1}$) expanding from the poles of the envelope and having an increasing velocity with the distance from the star. From recent HST observations, Sahai et al. (2003) report the discovery of a newly ejected high-speed jet-like outflow in this star. These observations, combined with a previous interferometric CO ($J = 1-0$) map of V Hya, favor the picture of an expanding, tilted and dense disk-like structure, oriented north-south, present inside the inner envelope.

We obtained images of V Hya at 9.78 and $11.65 \mu\text{m}$ (N2 and SiC filters respectively). Images are roughly spherically symmetric in the central parts of the envelope but isophotes become slightly elongated in the south-west direction at intensity levels lower than 20% of the central value. The N2-image has an

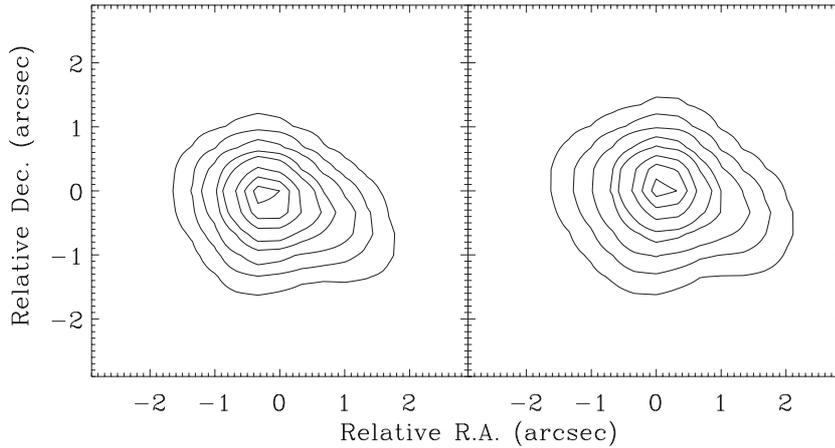


Fig. 1. Contour plots of the $9.78\ \mu\text{m}$ (left) and $11.65\ \mu\text{m}$ (right) images of V Hya. North is up and east is to the left. Contour levels are 90, 70, 50, 30, 20, 10, 5 and 2% of the peak surface brightness.

extension at a 2% intensity level of $4'' \times 4''$ whereas in the SiC filter the dimension is $4.5'' \times 4.5''$ (Fig. 1). The observed elongated feature in the south-west direction could be associated with the dust emission from material blown away perpendicular to the equatorial disk, consistent with the model proposed by Sahai et al. (2003).

Using the method proposed by Dayal et al. (1998), we derived a color map (9.78 and $11.65\ \mu\text{m}$ filters) for V Hya, which gives an indication of the mean dust temperature along the line of sight, weighted by the density distribution (Fig. 2). The meaning of such an average temperature is discussed in more detail in Sect. 4.

One striking aspect of the color map is that the maximum is off-center. If we perform a cut in the color map along the direction defined by the center of the envelope (the peak intensity coincides in both filters) and the color maximum (Fig. 3), we observe three distinct peaks: a fainter one at $0.9''$ east of the center; a slightly higher peak at the center of the envelope and a still higher peak at $0.9''$ west of the center. The direction of this cut is approximately the same as that of the jet observed by Sahai et al. (2003) through the [S II] emission as well as inferred from interferometric CO ($J = 1-0$) maps. Notice that the epoch of these observations corresponds to 2002 and 1998 respectively. The jet observed by Sahai et al. (2003) has a characteristic expansion timescale of ~ 3 years and an opening angle of $\sim 20^\circ - 30^\circ \sin i$. The estimated timescale suggests that the events observed in 1998 and 2002 are probably not the same. Thus, they suggested that V Hya undergoes periodic mass ejection events along a direction perpendicular to the equatorial disk. Our color map corresponds to data obtained in 1995, prior to the aforementioned events. If we assume an expansion velocity of $\sim 200\ \text{km s}^{-1}$ and a distance of 340 pc, we find an expansion timescale of ~ 8 yr (the peaks we observe are offset by $0.9''$ from the center and have an opening angle of roughly $40^\circ \sin i$). This timescale suggests that we are observing material ejected in a previous mass-loss episode and it is also consistent with the 12 yr period over which the high-velocity [S II] emission has been observed and monitored, from 1989 to 2000 (Lloyd Evans 2000).

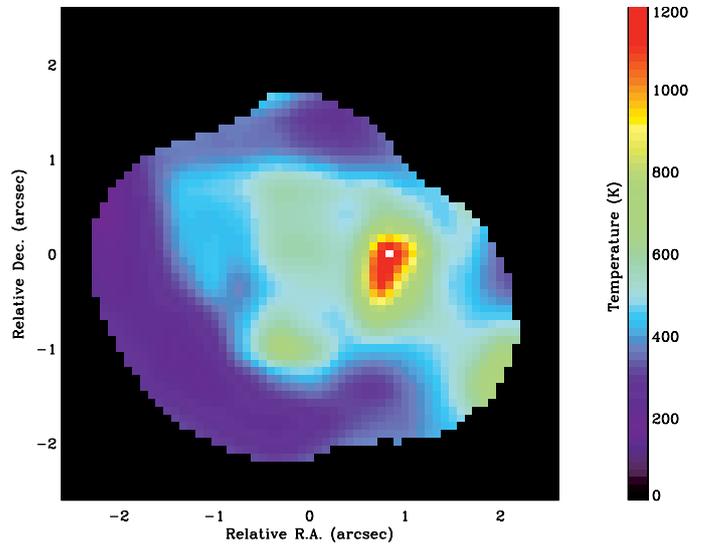


Fig. 2. Color map of V Hya. Notice that the emission peak is off-center, probably due to the presence of a jet roughly oriented in the east-west direction.

3.2. Roberts 22

Roberts 22 (IRAS 10197-5750) is an optical ($10'' \times 24''$) bipolar proto-planetary nebula (PPN), whose circumstellar dust envelope absorbs totally the light emitted by its central star (Allen et al. 1980). This object is also an OH maser source displaying strong emission at 1612 and 1665 MHz, indicating that the circumstellar envelope expands with a velocity of about $20\ \text{km s}^{-1}$. Optical spectroscopy shows that the $H\alpha$ line profile has wings extending up to $\pm 450\ \text{km s}^{-1}$, a signature of high-speed outflow. The bipolar shape of Roberts 22 is confirmed by HST observations. These observations show also a dust waist oriented along a PA of 121° , which darkens the central star and separates the polar lobes (Sahai et al. 1999). It is interesting to note that Roberts 22 display thin shell structures as seen in many PPN (e.g. CRL 2688), which could be due to discrete mass-loss episodes; In addition, a striking aspect of Roberts 22 is that C-rich and O-rich spectral features are seen as in the case of the Red Rectangle nebula (Waters et al. 1998). A possible

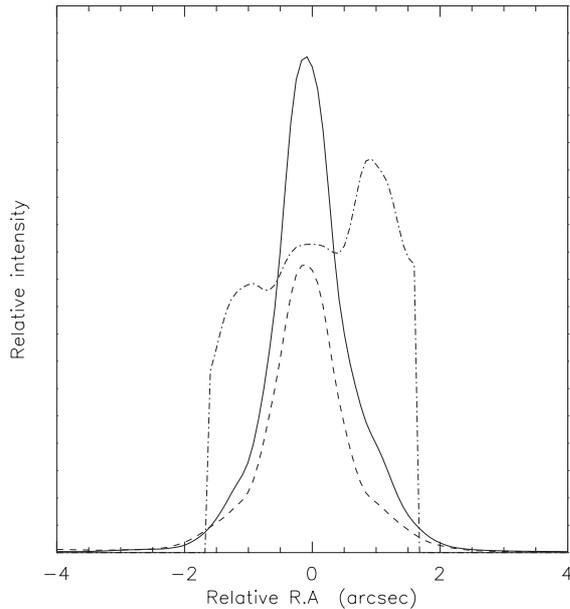


Fig. 3. Cut in the color map of V Hya along the direction defined by the center of the envelope and the color maximum (dot-dashed curve). The solid and the dashed curves are cuts in the same direction made in the N2 and SiC filter maps respectively.

explanation is that the O-rich material is a remnant of an earlier mass-loss event that occurred when the envelope of the star was O-rich, followed by the capture and formation of a disk-like structure around a companion, the C-rich material being ejected in a late evolutionary phase, when the star atmosphere became C-rich.

Our $8.39 \mu\text{m}$ and $11.65 \mu\text{m}$ images of Roberts 22 show an envelope slightly elongated in the direction north-east/south-west with PA $\sim 45^\circ$ (Fig. 5). The SiC image ($11.65 \mu\text{m}$) resembles pictures obtained by the HST, since the X-shaped structure can be inferred, but being broader than in HST images. This structure could be due to the presence of an extended dust envelope not only along the equator but also inside the lobes observed in optical diffused light. Differently from V Hya and IRC +10216, IRAS low resolution spectra of Roberts 22 (Fig. 4) clearly show two emission components, peaking respectively around $8 \mu\text{m}$ and $12 \mu\text{m}$. Our images were obtained at these wavelengths, excluding the application of our method (see Sect. 4) to derive the distributions of density and temperature.

In Fig. 6 a color map (8.39 and $11.65 \mu\text{m}$ filters) of Roberts 22 is shown. In this map the elongated north-east/south-west structure is clearly seen as some structures suggesting the existence of local temperature variations, probably produced by different physical properties of dust grains inside the envelope.

3.3. IRC+10216

IRC +10216 (IRAS 09452+1330) is the closest, and thus one of the most studied C-rich AGB stars. Its circumstellar envelope is roughly spherical but composed of many discrete and incomplete shells (Mauron & Huggins 2000). From

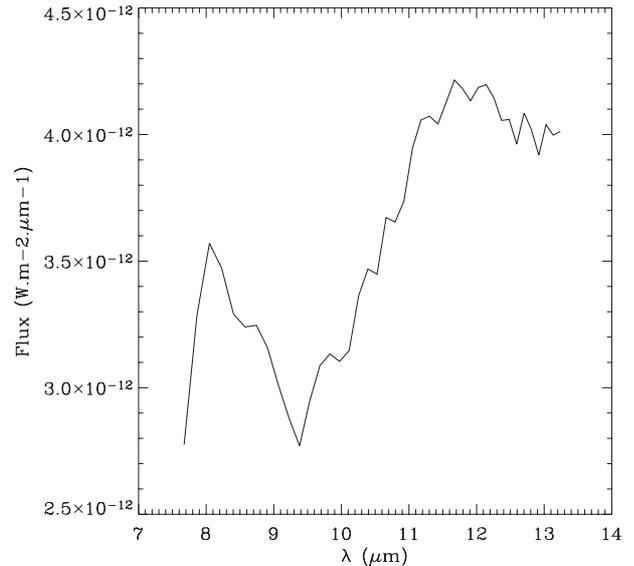


Fig. 4. IRAS low resolution spectra of Roberts 22.

near-infrared speckle-masking interferometry observations, Weigelt et al. (1998) were able to resolve five individual clumps within a distance of $0.21''$ from the central star. After six years of monitoring these structures, the onset of a fast wind phase is hardly suggested (Weigelt et al. 2002), although significant velocities might be inferred through proper motion measurements. These clumpy structures seem to be consistent with the expectations of the Interacting Stellar Winds model (ISW) of Kwok et al. (1978), which correctly describes the shape of a planetary nebula in which a fast wind starting at the end of the AGB phase interacts with low velocity fossil material. IRC +10216 would presently be in the fast wind phase whose onset is relatively recent.

A large number of observations of this star is available in the UBV filters (Mauron et al. 2003), radio (Gensheimer et al. 1995) and infrared (Weigelt et al. 1998, and references therein). However, to our knowledge, no direct mid-infrared images have been published in the literature. Here we report the first images of this star in narrow band filters at 8.39 and $11.65 \mu\text{m}$ (N1 and SiC filters). In these images, IRC +10216 appears to be roughly spherically symmetric, having a spatial extension at a 2% intensity level of $3.5'' \times 3.5''$ and $4.5'' \times 4.5''$ in the N1 and SiC filters, respectively (Fig. 7).

In Fig. 8 the color map (8.39 and $11.65 \mu\text{m}$ filters) is shown. An elongated structure with the same orientation as the near-infrared image of Weigelt et al. (1998) is clearly seen, as inhomogeneities in the temperature distribution. The cold region observed to the north-west is probably an artifact since the signal-to-noise ratio is low at these intensity levels.

3.4. CIT 6

CIT 6 (IRAS 10131+3049) is an “extreme” carbon star believed to be undergoing a transition between the AGB phase and the PN stage (Trammell et al. 1994). CIT 6 is the second brightest carbon star at $12 \mu\text{m}$ after IRC +10216. It is a long-period-variable star with a period of ~ 628 days (e.g.

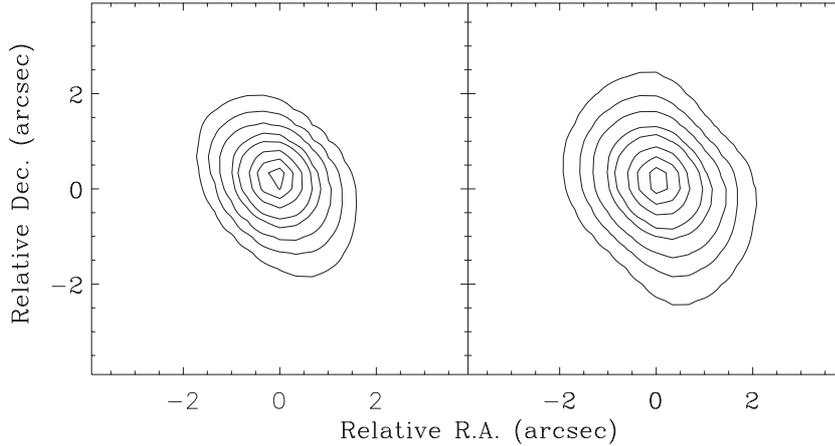


Fig. 5. Contour plots of the $8.4\ \mu\text{m}$ (left) and $11.65\ \mu\text{m}$ (right) images of Roberts 22. North is up and east is to the left. Contour levels are 90, 70, 50, 30, 20, 10, 5 and 2% of the peak surface brightness.

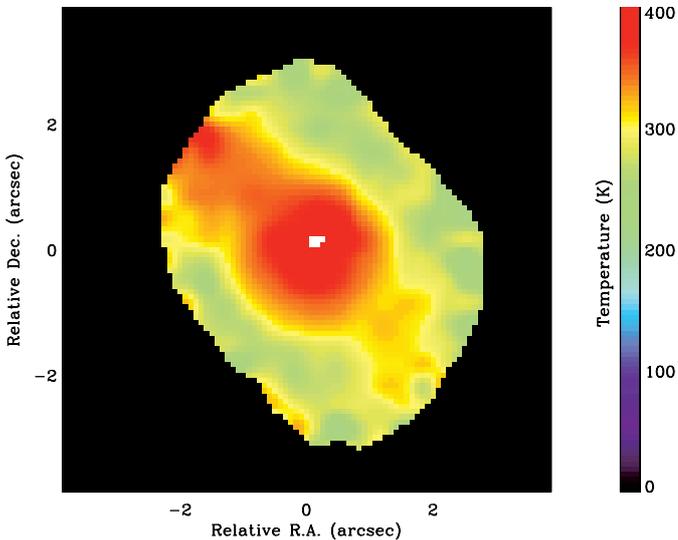


Fig. 6. Color map of Roberts 22. Warm grains are observed at the center of the nebula, as well as in a jet-like structure with PA $\sim 45^\circ$.

Alksnis 1995), and was classified as PPN by Johnson & Jones (1991). However, the estimated photospheric temperature of $\sim 2800\ \text{K}$ (Cohen 1979) implies that the star is still near the AGB. Visible and infrared polarization measurements show that the light at these wavelengths is highly polarized, with a strong variation of the polarization angle with the wavelength, indicating that the distribution of the circumstellar material around this star is asymmetric (Trammell et al. 1994, and references therein).

CIT 6 has been observed with a $\sim 50\ \text{mas}$ spatial resolution at $2.2\ \mu\text{m}$ and $3.1\ \mu\text{m}$, using aperture masking techniques at the Keck I Telescope (Monnier et al. 2000). The object was resolved into two main components: a bright elongated feature showing shape variations (north component) and a fainter bluer (south) component about $66\ \text{mas}$ away. The observed shape variations suggest an evolution of the dust envelope, which can be explained by changes in the illumination pattern. CIT 6 was also observed with WFPC2 aboard the HST, through the filters $F439W$ ($0.429\ \mu\text{m}$), $F555W$ ($0.525\ \mu\text{m}$), and

$F675W$ ($0.674\ \mu\text{m}$). Both components are seen in $F675W$ and $F555W$ images while in the $F439W$ image only the south component is seen. These observations are consistent with a binary star scenario where the north component presents a 640 day pulsation period not seen in the south component, probably a main sequence A-F star. Associated with the latter, a cometary-like feature extending along the south-west direction can be seen in the $3.1\ \mu\text{m}$ images (Monnier et al. 2000). Recently, Schmidt et al. (2002) concluded from HST imaging, ground-based spectroscopy and spectropolarimetric observations that a third closely coupled companion would be required to explain the bipolar morphology of this star.

Here we report the first mid-infrared imaging of CIT 6 at $8.39\ \mu\text{m}$ (N1) and $9.7\ \mu\text{m}$ (Silicates). The cometary-like feature is clearly seen in the $9.7\ \mu\text{m}$ image (Fig. 9) extending up to $7''$ (at 1% intensity level) from the center of the envelope at PA $\sim 230^\circ$ (the north and south components seen in the near infrared are not resolved in our images). This cometary-like feature which is less extended and less prominent in the $8.4\ \mu\text{m}$ image compares well with pictures at $3.1\ \mu\text{m}$ (Monnier et al. 2000) and is roughly orthogonal to the extended structure seen in the $F439W$ HST image. The morphology of CIT 6 seen in the optical and near-infrared domains can qualitatively be explained if the bipolar nebula is essentially seen through its scattered light. The central star is obscured by a central disk or torus of dust. The dominance of the northern pole observed in the near-infrared indicates an inclination toward us (Monnier et al. 2000). This picture is supported by linear polarization data in the infrared but it does not account for the cometary shape seen in blue WFPC2 images (Trammell et al. 1994) and in our mid-infrared images. No satisfactory flux calibration of our images was obtained.

4. The model

In spite of the presence of some asymmetries in the outer envelope of these objects, here we consider a spherical geometry as a first approximation to derive the temperature and the density profiles of the circumstellar dust around V Hya and IRC +10216. Roberts 22 and CIT 6 were excluded from

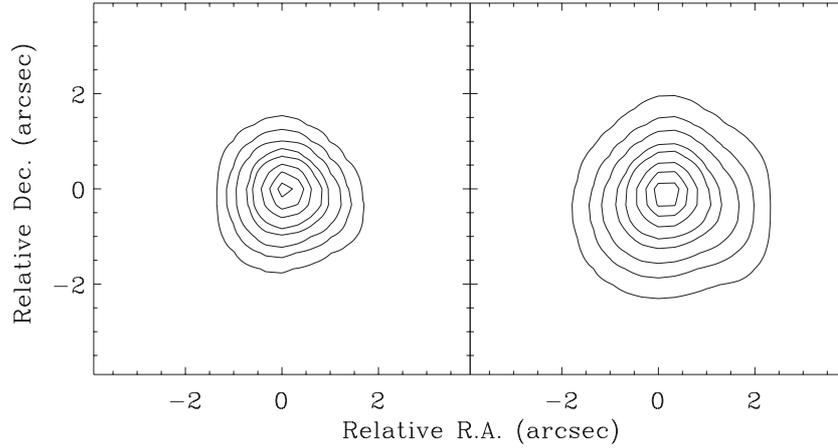


Fig. 7. Contour plots of the $8.4\ \mu\text{m}$ (left) and $11.65\ \mu\text{m}$ (right) images of IRC +10216. North is up and east is to the left. Contour levels are 90, 70, 50, 30, 20, 10, 5 and 2% of the peak surface brightness.

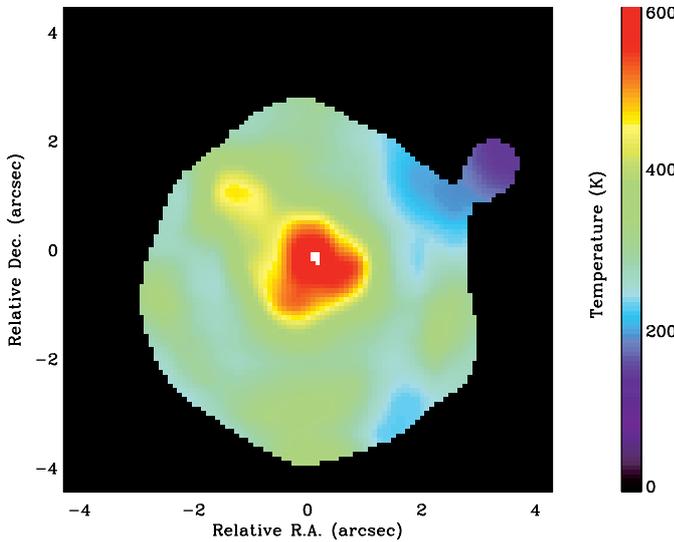


Fig. 8. Color map of IRC +10216. The envelope is roughly spherically symmetrical but displays inhomogeneities in the dust temperature distribution.

our modeling, since the former has at least two emission components and the latter shows clearly elongated isophotes. Our procedure is the following: in a first step, from the isophotal contours we derive the equivalent radius a defined by the relation $a = \sqrt{S/\pi}$, where S is the area delimited by the isophote of intensity $I(a)$.

If the outer nebular light at $\sim 10\ \mu\text{m}$ is dominated by the emission of warm dust grains and if the nebula is optically thin at these wavelengths, then the intensity at a projected distance a from the center is

$$I(a) = 2 \int_a^{R_e} j(r) \frac{r dr}{\sqrt{r^2 - a^2}} \quad (1)$$

where R_e is the outer radius of the dust envelope and $j(r)$ is the dust emission coefficient, which is a function of the distance r from the star. The equation above can easily be written in the

form of an Abel integral, whose inversion method is quite well known. Therefore, the dust emission coefficient is given by

$$j(r) = -\frac{1}{\pi} \int_r^\infty \frac{\frac{\partial I}{\partial a} da}{\sqrt{a^2 - r^2}}. \quad (2)$$

On the other hand, from its definition, the dust emissivity can explicitly be written as

$$j(r) = n_g \langle r_g^2 \rangle Q_\lambda B_\lambda(T_g), \quad (3)$$

where n_g is the dust grain density, $\langle r_g^2 \rangle$ is the dust grain mean square radius, Q_λ is the emission (absorption) efficiency and $B_\lambda(T_g)$ is the Planck function at the grain temperature T_g . If the dust emissivity is derived at two different wavelengths λ_1 and λ_2 , then their ratio is simply

$$\frac{j_{\lambda_1}}{j_{\lambda_2}} = \frac{Q_{\lambda_1} B_{\lambda_1}(T_g)}{Q_{\lambda_2} B_{\lambda_2}(T_g)}. \quad (4)$$

The grain temperature can be derived from Eqs. (2) and (4) if we assume a given chemical composition and dimension for the dust grains. Once the temperature distribution is obtained, the grain density can be evaluated from Eq. (3). The larger the difference between λ_1 and λ_2 , the higher the accuracy in the determination of the temperature distribution and the grain density.

The emission efficiency of the grains Q_λ was computed using Mie's theory, assuming homogeneous and spherical grains constituted of amorphous carbon with a mass density of $2\ \text{g cm}^{-3}$. Optical properties were taken from laboratory data (Rouleau & Martin 1991; Suh 2000; Jager et al. 1998, hereafter respectively Models 1, 2 and 3). The choice of dust parameters is somewhat arbitrary, since there are no strong observational constraints on them. Thus, grain dimensions were chosen to be comparable with those adopted in previous studies, whenever existent.

For numerical reasons, in order to invert the Abel integral, we preferred to adopt an analytical representation for the intensity $I(a)$ derived from our data. In fact, the function

$$I(a) = \frac{A}{[1 + (\frac{a}{r_0})^\gamma]} \quad (5)$$

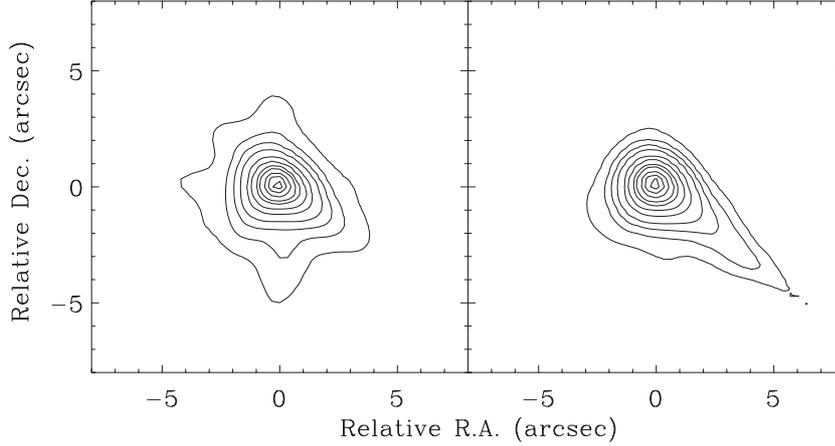


Fig. 9. Contour plots of the $8.4\ \mu\text{m}$ (left) and $9.7\ \mu\text{m}$ (right) images for CIT 6. North is up and east is to the left. Contour levels are 90, 70, 50, 30, 20, 10, 5, 3, 2 and 1% of the peak surface brightness.

Table 4. Best fit parameters appearing in Eq. (5).

Object	Filter	A ($\text{erg cm}^{-2} \text{s}^{-1} \text{arcsec}^{-2} \text{Hz}^{-1}$)	r_o (arcsec)	γ
V Hya	SiC	$4.86 \pm 0.05 \times 10^{-21}$	0.49 ± 0.05	2.66 ± 0.05
	N2	$4.79 \pm 0.05 \times 10^{-21}$	0.48 ± 0.05	2.70 ± 0.05
IRC +10216	SiC	$1.03 \pm 0.01 \times 10^{-19}$	0.73 ± 0.01	2.47 ± 0.05
	N1	$1.24 \pm 0.01 \times 10^{-19}$	0.63 ± 0.01	2.64 ± 0.07

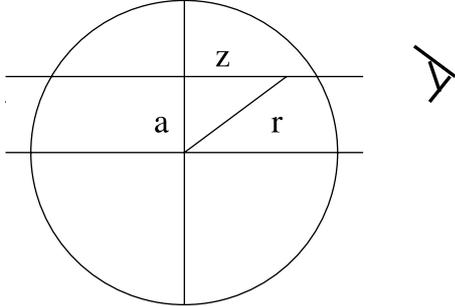


Fig. 10. Schematic view of the nebula considered as spherically symmetric. The observer is on the right. So we have $I(a) = 2 \int_0^{z_{\text{max}}} j(r) dz$, where I is the intensity received and j the emissivity of the dust grains.

quite well represents the observed intensity profiles. The resulting fit parameters for the different objects are shown in Table 4.

To verify the consistency of our assumption, namely, an optically thin envelope at the considered wavelength, the optical depth at $11.6\ \mu\text{m}$ was computed from the equation below and the derived density distribution.

$$\tau = \int \pi r_g^2 Q_\lambda n_g(r) dr. \quad (6)$$

Grains sizes ranging from $0.007\ \mu\text{m}$ to $5\ \mu\text{m}$ and dust optical properties taken from Models 1, 2 and 3 were used. The derived optical depth values (Table 5) are consistent with the hypothesis of an optically thin envelope at $\sim 10\ \mu\text{m}$, since for all models $\tau \leq 0.3$. As an additional check, other flux values in the interval $8\text{--}12\ \mu\text{m}$ were computed and compared with IRAS data. The good agreement between expected and observed values reinforces the basic assumption of our model.

Table 5. Estimated dust optical depths at $11.6\ \mu\text{m}$ for different grain sizes and optical properties. $\tau(1)$, $\tau(2)$ and $\tau(3)$ correspond to Models 1, 2 and 3 respectively.

Object	r_g	$\tau(1)$	$\tau(2)$	$\tau(3)$
V Hya	0.2	2.8×10^{-2}	2.9×10^{-2}	1.7×10^{-2}
	5.0	7.5×10^{-3}	7.7×10^{-3}	8.3×10^{-3}
IRC +10216	0.007	7.9×10^{-2}	7.9×10^{-2}	2.3×10^{-2}
	0.008	7.9×10^{-2}	2.3×10^{-2}	4.7×10^{-2}
	0.015	2.3×10^{-1}	1.4×10^{-1}	1.9×10^{-1}
	0.05	2.3×10^{-1}	1.9×10^{-1}	2.6×10^{-1}
	0.16	2.6×10^{-1}	2.2×10^{-1}	2.7×10^{-1}
	5.0	6.0×10^{-2}	6.6×10^{-2}	6.8×10^{-2}

Color maps are usually interpreted as “temperature maps” (Dayal et al. 1998). If the medium is optically thin, the intensity at a given wavelength is given by:

$$I_\lambda = \int n_g \langle r_g^2 \rangle Q_\lambda B_\lambda(T_g) dz. \quad (7)$$

Let us define:

$$\overline{Q_\lambda B_\lambda(T)} = \frac{\int n_g \langle r_g^2 \rangle Q_\lambda B_\lambda(T_g) dz}{\int n_g \langle r_g^2 \rangle dz}. \quad (8)$$

Then, from Eqs. (7) and (8) one obtains:

$$\frac{I_{\lambda 1}}{I_{\lambda 2}} = \frac{\overline{Q_{\lambda 1} B_{\lambda 1}(T)}}{\overline{Q_{\lambda 2} B_{\lambda 2}(T)}}. \quad (9)$$

Equation (9) tells us that a given color corresponds to a given temperature only if this quantity is constant along the line of

sight or, in other words, if the temperature is uniform throughout the nebula. Our model implies the presence of temperature gradients, which are clearly seen, for instance, in the color maps of Roberts 22 and IRC +10216. However, color inhomogeneities are also detected, which should be interpreted as temperature variations along the line of sight. These local variations are probably due to non-uniformities in the grain properties, which affect the equilibrium temperature and the grain emissivity (see Eq. (9)). Therefore the color maps should only be considered as a temperature “indicator”.

5. Results and discussion

We have applied the aforementioned procedure to V Hya and IRC +10216 in order to derive temperature and density profiles for the dust inside the envelope of those stars. Once the density profile is obtained, the total dust mass (M_{dust}) can be calculated from

$$M_{\text{dust}} = \int_{R_i}^{R_e} 4\pi r^2 n_g(r) m_g dr \quad (10)$$

where $m_g = 4\pi\delta r_g^3/3$ is the mass of a dust grain (δ being equal to 2.0 g cm^{-3} for the amorphous carbon), R_i and R_e are respectively the internal and external radius of the dust envelope. If we assume that the gas-to-dust mass ratio ξ and the expansion velocity V_e are constant throughout the dust envelope, then an estimate of the mass-loss rate \dot{M} can be obtained from

$$\dot{M} = \xi M_{\text{dust}} \frac{V_e}{R_e}. \quad (11)$$

Values for the gas-to-dust ratio ξ adopted by different authors range from 200 (Groenewegen 1997) up to 360 (Knapp et al. 1997). In the present calculations, the value $\xi = 300$ was adopted and for the expansion velocity, the values derived from CO observations by Loup et al. (1993) were used. Whenever the radiation pressure drives the mass-loss, dust velocities are known to be usually higher than gas velocities, but differences are expected to be small if drag forces are effective (Steffen et al. 1997). For the external radius R_e , we have adopted the values derived from our images. As we shall see later, the dust density decreases approximately as r^{-2} and, in this case, the mass-loss rate is almost independent of the value of R_e .

5.1. V Hya

Some radiative transfer models have already been proposed to represent the circumstellar envelope of this star. Kahane et al. (1996) have modeled this object using millimeter CO observations and an axisymmetric model. Distance estimates for V Hya are in the range 400 pc (derived from Hipparcos data) up to 550 pc (derived by Bergeat et al. 1998 from the period-luminosity relationship for carbon stars). However, lower values can be found in the literature. Kahane et al. (1996) adopted a distance of 340 pc and derived a total mass-loss rate of $\sim 1.5 \times 10^{-6} M_{\odot}/\text{yr}$ and a total mass (gas+dust) for the circumstellar envelope of $\sim 2.1 \times 10^{-3} M_{\odot}$. Knapp et al. (1997) modeled the dust envelope of V Hya by assuming grains constituted of amorphous carbon with dimensions of $0.2 \mu\text{m}$ and by

adopting a distance of 380 pc. They derived an envelope mass of $8.0 \times 10^{-3} M_{\odot}$ and a mass-loss rate of $2.5 \times 10^{-5} M_{\odot}/\text{yr}$. These values depend on the optical properties of the grains but are significantly higher than those obtained by Kahane et al. (1996).

In our calculations, the distance to the star was assumed to be in the range 340–550 pc. The masses of the dust envelope and the corresponding mass-loss rates obtained for different dust grain parameters and distances are given in Table 6 and in Table 7. Previous estimates are also given. The temperature and density profiles of V Hya obtained with the parameters $d = 340 \text{ pc}$; $r_g = 0.2 \mu\text{m}$ and $r_g = 5 \mu\text{m}$ are shown in Fig. 11. The temperature determination does not depend on the assumed distance but depends on the adopted grain properties. For a grain size of $0.2 \mu\text{m}$, the dust temperature ranges from 390–460 K at the inner envelope down to 255–270 K at the outer parts of the envelope. The higher temperature values correspond to optical properties of Model 3. Higher temperatures, ranging from 660 K (inner envelope) down to 350 K (outer envelope) are obtained for larger grains ($r_g = 5 \mu\text{m}$) and, in this case, no significant differences are seen for the different optical properties of Models 1, 2 and 3.

If we compare our mass-loss rates derived with grain parameters similar ($r_g = 0.2 \mu\text{m}$) to those adopted by Kahane et al. (1996), we notice that the results are comparable, in spite of differences in the methodology used by those authors with respect to our approach.

5.2. IRC +10216

We have also applied our procedure to the dust envelope of IRC +10216 for distances in the interval 130–200 pc and dust grain sizes ranging from $0.007 \mu\text{m}$ up to $5.0 \mu\text{m}$. Our results are given in Tables 6 and 7 and in Fig. 12.

For small grains ($r_g = 0.16 \mu\text{m}$), temperatures attain 600 K in the inner envelope and decrease to 220 K in the outer parts. For larger grains ($r_g = 5 \mu\text{m}$) temperatures in the inner region are higher $\sim 1200 \text{ K}$, decreasing to values around 300 K in the outer regions of the envelope.

For a distance of 200 pc, the derived dust masses are within the interval $0.7\text{--}7.9 \times 10^{-5} M_{\odot}$ and the corresponding total (gas+dust) mass-loss rates are $0.3\text{--}3.2 \times 10^{-5} M_{\odot}/\text{yr}$.

Being one of the most studied carbon stars, several models for the envelope of this object can be found in the literature. Keady et al. (1988) modeled the envelope of this star using data on the fundamental and first overtone lines of CO, assuming a distance of 200 pc and dust grains composed of amorphous carbon of size $0.05 \mu\text{m}$. They found temperatures varying from $\sim 1300 \text{ K}$ at 13.5 AU down to 300 K at $\sim 180 \text{ AU}$. They have estimated a mass-loss rate close to $2 \times 10^{-5} M_{\odot}/\text{yr}$. Groenewegen (1997) has considered the effects of a non-steady mass-loss and estimated a total dust mass of about $1.0 D^2 \text{ (kpc)} M_{\odot}$ for a variable mass-loss rate, which should be compared to $0.13 D^2 \text{ (kpc)} M_{\odot}$ derived under the usual assumption of a steady flow. An important aspect of this work is the constraint obtained for the grain sizes ($\sim 0.16 \mu\text{m}$), derived from a large body of data. Winters et al. (1994) developed a quite

Table 6. Mass of the dust envelope of V Hya and IRC +10216 for different optical models, distances (d) and grain sizes (r_g) compared to values derived from the literature. The inner and outer integration radii were 0.1'' and 4'' for V Hya, 0.1'' and 10'' for IRC +10216. Dust masses taken from the literature have been scaled to correspond to the same outer radii as those used in this present work. Expansion velocities have been taken equal to 18.7 km s⁻¹ and 14.7 km s⁻¹ for V Hya and IRC +10216 respectively.

Object	r_g (μm)	d (pc)	Present Models $M_{\text{dust}}(M_{\odot})$			Previous Models	
			Model 1	Model 2	Model 3	$M_{\text{dust}}(M_{\odot})$	Ref.
V Hya	0.2	340	2.3×10^{-6}	1.2×10^{-6}	5.4×10^{-7}	2.2×10^{-6}	(a)*
	0.2	380	2.8×10^{-6}	1.5×10^{-6}	6.7×10^{-7}	9.5×10^{-6}	(b)
	0.2	550	5.9×10^{-6}	3.2×10^{-6}	1.4×10^{-6}		
	5	340	2.4×10^{-7}	2.4×10^{-7}	3.1×10^{-7}		
IRC +10216	0.007	170	1.5×10^{-5}	7.3×10^{-6}	2.1×10^{-5}	$2.7\text{--}4.3 \times 10^{-4}$	(c)
	0.008	150	1.3×10^{-5}	1.9×10^{-5}	3.9×10^{-5}	$0.3\text{--}3.8 \times 10^{-4}$	(d)
	0.015	130	2.6×10^{-5}	8.1×10^{-6}	8.7×10^{-5}	2.2×10^{-5}	(e)
	0.05	200	6.8×10^{-5}	3.0×10^{-5}	3.2×10^{-5}		
	0.16	200	7.9×10^{-5}	3.3×10^{-5}	3.2×10^{-5}	$0.8\text{--}6.7 \times 10^{-4}$	(f)
	5.0	130	3.0×10^{-6}	3.2×10^{-6}	3.9×10^{-6}		
	5.0	200	7.2×10^{-6}	7.5×10^{-6}	9.3×10^{-6}		

References: (a) Kahane et al. (1996); (b) Knapp et al. (1997); (c) Winters et al. (1994); (d) Danchi et al. (1994); (e) Men'shchikov et al. (2001); (f) Groenewegen (1997).

* Based on radio observations. No grain model was used in these models, then estimation of the dust mass was done assuming a gas-to-dust mass ratio of 300.

Table 7. Mass-loss rates for V Hya and IRC +10216 derived from different parameters and compared with previous determinations as in Table 6.

Object	r_g (μm)	d (pc)	Present Models $\dot{M} (M_{\odot}\text{yr}^{-1})$			Previous Models	
			Model 1	Model 2	Model 3	$\dot{M} (M_{\odot}\text{yr}^{-1})$	Ref.
V Hya	0.2	340	2.0×10^{-6}	1.0×10^{-6}	4.7×10^{-7}	1.5×10^{-6}	(a)
	0.2	380	2.2×10^{-6}	1.2×10^{-6}	5.2×10^{-7}	2.5×10^{-5}	(b)
	0.2	550	3.2×10^{-6}	1.7×10^{-6}	7.6×10^{-7}		
	5	340	2.1×10^{-7}	2.1×10^{-7}	2.7×10^{-7}		
IRC +10216	0.007	170	8.0×10^{-6}	4.0×10^{-6}	1.1×10^{-5}	8.0×10^{-5}	(c)
	0.008	150	8.1×10^{-6}	1.2×10^{-5}	2.4×10^{-5}	$2.2\text{--}4.0 \times 10^{-5}$	(d)
	0.015	130	1.9×10^{-5}	5.8×10^{-6}	6.2×10^{-6}	1.6×10^{-5}	(e)
	0.05	200	3.2×10^{-5}	1.4×10^{-5}	1.5×10^{-5}	2.0×10^{-5}	(g)
	0.16	200	3.7×10^{-5}	1.5×10^{-5}	1.5×10^{-5}	4.9×10^{-5}	(f)
	5.0	130	2.2×10^{-6}	2.3×10^{-6}	2.8×10^{-6}		
	5.0	200	3.3×10^{-6}	3.5×10^{-6}	4.3×10^{-6}		

References: (a–f) as in Table 6; g) Keady et al. (1988).

sophisticated model, including a hydrodynamical gas-dust coupling, thermodynamics, radiative transfer, chemistry in the processes of dust formation and growth. Assuming a distance of 170 pc and a mass-loss rate of $8 \times 10^{-5} M_{\odot}/\text{yr}$, they found a dust mass in the envelope within the range $2.0\text{--}3.2 \times 10^{-3} M_{\odot}$. Similar dust masses and mass-loss rates have been found by Danchi et al. (1994) from mid-infrared data.

Most of the past studies suggest that mass-loss rates for this object are in the range $2\text{--}8 \times 10^{-5} M_{\odot}/\text{yr}$ if the distance is 170–200 pc. The present analysis indicates that such values can only be obtained if grain dimensions are in the

interval 0.05–0.6 μm . Larger grains lead to mass-loss rates one order of magnitude smaller than the aforementioned estimates.

6. Conclusions

In this paper, the first mid-infrared images of a sample of carbon-rich AGB stars are reported. From our data, temperature and density profiles of the dust envelope surrounding these stars were derived from a simple inversion method applied to obtain the dust emissivity.

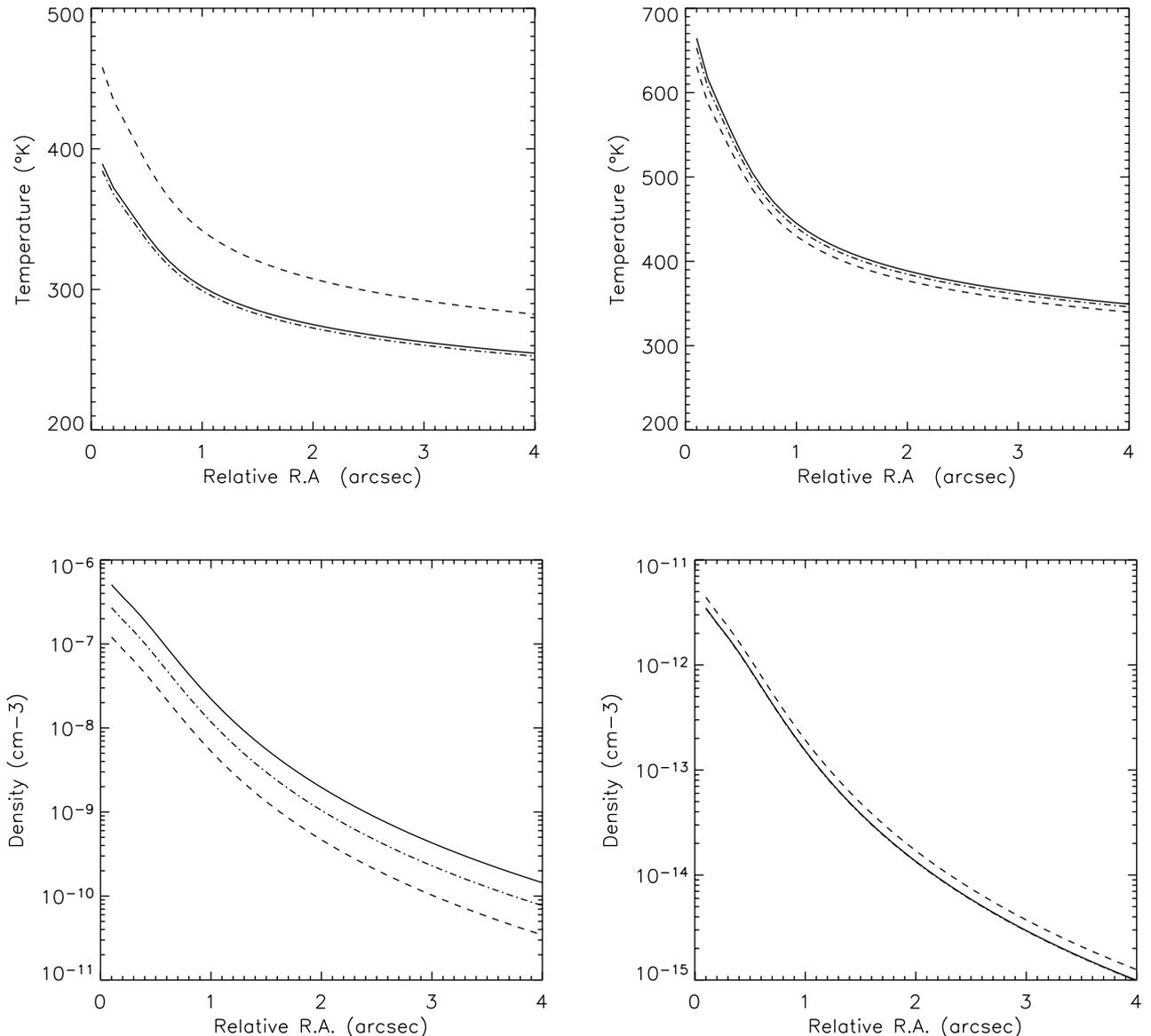


Fig. 11. Temperature and density profiles of V Hya obtained with the model parameters: $d = 340$ pc, $r_g = 0.2 \mu\text{m}$ (left) and $d = 340$ pc, $r_g = 5 \mu\text{m}$ (right) respectively. The solid, dot-dashed and dashed line curves represent results from Models 1, 2 and 3 respectively.

Color maps are compatible with the presence of temperature gradients: hot and/or warm dust grains are present in the central regions whereas cold dust grains are seen in outer parts of the envelope. Moreover, these color maps suggest the presence of temperature inhomogeneities, probably due to local variations of the optical properties of the grains. Color maps strengthen sub-structures eventually present in the envelopes as in the case of V Hya. In this star, the “hottest point” indicated by the color maximum does not coincide with the center of the envelope and the direction defined by these two points is practically orthogonal to the equatorial disk observed in the CO maps (Sahai et al. 2003). This direction coincides approximately with the jet seen in [SII] emission. The color map of Roberts 22 indicates the presence of bipolar lobes, with the north-east side being “hotter” than the south-west side (Fig. 6).

Temperatures near the central regions (angular distances $\theta \sim 1''$) are in the range 300–350 K whereas at $\theta \sim 4''$ the

values are in the range 250–280 K, if we assume grains with typical sizes of about $0.15 \mu\text{m}$. Larger grains ($r_g \sim 5 \mu\text{m}$) give higher values: 430–525 K at $\theta \sim 1''$ and 300–350 K at $\theta \sim 4''$. Comparison with previous studies, assuming similar grain properties, produces consistent results.

The optical properties of amorphous carbon grains measured in the laboratory by different experiments are not exactly the same and were characterized by Models 1, 2 and 3, according to the authors who performed the measurements. The temperature depends on the *ratio* between the grain absorption efficiency Q_λ at two (nearby) wavelengths (see Eq. (4)). Thus, the resulting values are not significantly altered according to the adopted set of measurements of the grain properties. This is not the case for the grain density, which depends directly on Q_λ (Eq. (3)). As a consequence, densities derived from Model 1 are always larger than those derived from Models 2 and 3, if

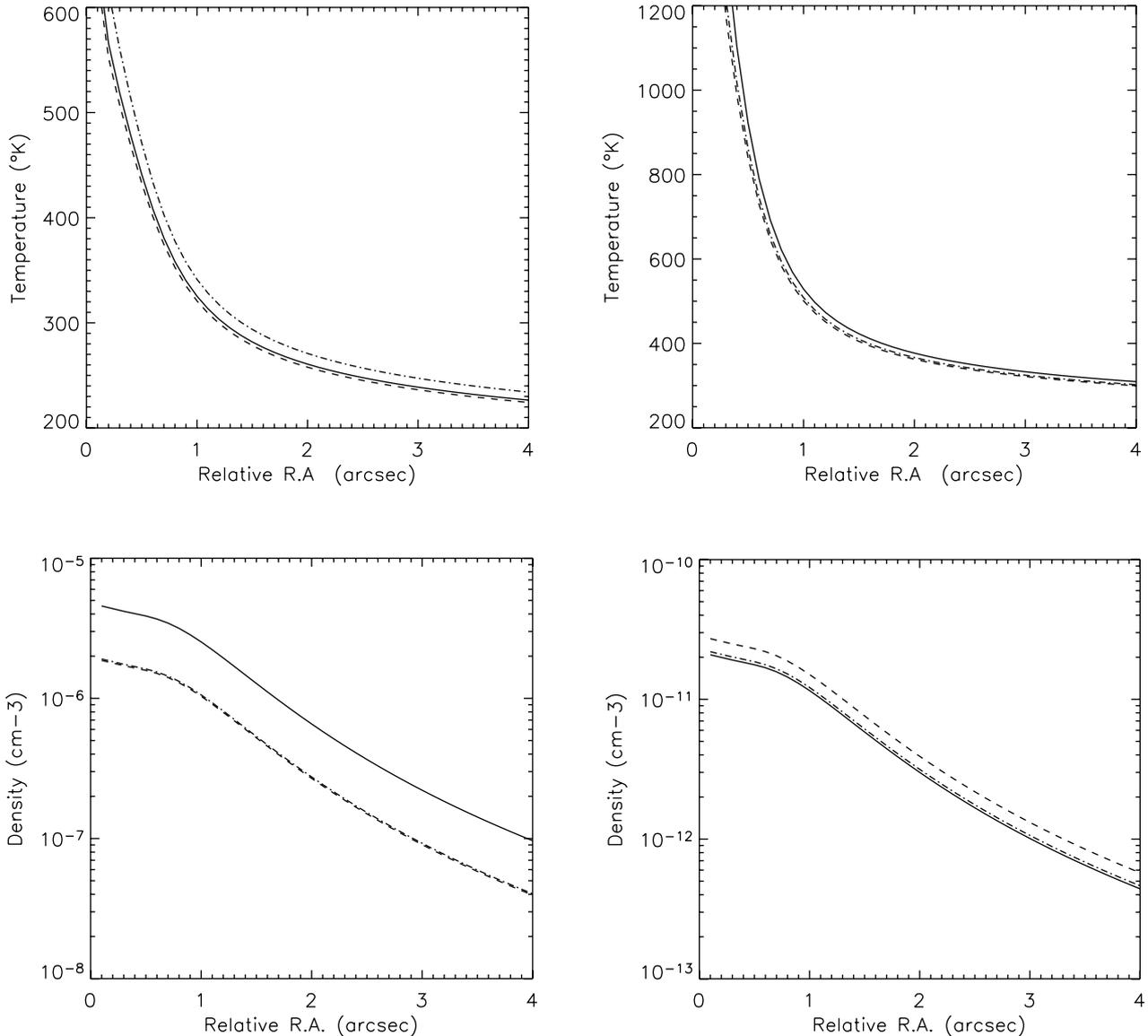


Fig. 12. Temperature and density profiles of IRC +10216 obtained with the model parameters: $d = 200$ pc, $r_g = 0.16 \mu\text{m}$ (left) and $d = 130$ pc, $r_g = 5 \mu\text{m}$ (right) respectively. The solid, dot-dashed and dashed line curves represent results from Models 1, 2 and 3 respectively.

small grains ($r_g \sim 0.1 \mu\text{m}$) are considered. For larger grains ($r_g \sim 5 \mu\text{m}$) the differences are not significant.

Our calculations are based on the hypothesis that the dust envelope is optically thin at wavelengths around $10 \mu\text{m}$. The consistency of this assumption was verified a posteriori, since optical depths are indeed less than one and the computed spectrum in the interval $8\text{--}12 \mu\text{m}$ is consistent with IRAS-LRS data, in the cases of V Hya and IRC +10216. The estimated mass-loss rate for V Hya is compatible with that derived by Kahane et al. (1996), if grain dimensions of about $0.2 \mu\text{m}$ are adopted and one uses the optical Models 1 or 2. Model 3, with the same grain size, leads to a mass-loss rate a factor of three smaller. However, we notice that for a grain size of $0.2 \mu\text{m}$, Knapp et al. (1997) derived a mass-loss rate one order of magnitude higher than our values. A similar analysis for IRC +10216 leads to the conclusion that, for distances of about 200 pc, our mass-loss rates are in agreement, in particular, with the determinations of

Keady et al. (1988) and Groenewegen (1997) if grain sizes in the range $0.05\text{--}0.16 \mu\text{m}$ are adopted, independent of the optical grain model.

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References

- Alksnis, A. 1995, *Baltic Astron.*, 4, 79
- Allen, D. A., Hyland, A. R., & Caswell, J. L. 1980, *MNRAS*, 192, 505
- Barnbaum, C., Morris, M., & Kahane, C. 1995, *ApJ*, 450, 862
- Bergeat, J., Knapik A., & Rutily, B. 1998, *A&A*, 332, 53
- Blöcker, T. 1999, *Asymptotic Giant Branch Stars*, IAU Symp., 191, 21

- Close, L. M., Biller, B., Hoffmann, W. F., et al. 2003, *ApJ*, 598, L35
- Cohen, M. 1979, *MNRAS*, 186, 837
- Danchi, W. C., Bester, M., Degiacomi, C. G., Greenhill, L. J., & Townes, C. H. 1994, *AJ*, 107, 1469
- Dayal, A., Hoffmann, W. F., Biegging, J. H., et al. 1998, *ApJ*, 492, 603
- Gensheimer, P. D., Likkell, L., & Snyder, L. E. 1995, *ApJ*, 439, 445
- Gledhill, T. M., & Yates, J. A. 2003, *MNRAS*, 343, 880
- Groenewegen, M. A. T. 1997, *A&A*, 317, 503
- Habing, H. J. 1996, *A&A Rev.*, 7, 97
- Henning, T. 1999, *Asymptotic Giant Branch Stars*, IAU Symp. 191, 221
- Hollis, J. M., Dorband, J. E., & Yusef-Zadeh, F. 1992, *ApJ*, 386, 293
- Iben, I., & Renzini, A. 1983, *ARA&A*, 21, 271
- Jager, C., Mutschke, H., & Henning, T. 1998, *A&A*, 332, 291
- Johnson, J. J., & Jones, T. J. 1991, *AJ*, 101, 1735
- Jura, M., Chen, C., & Plavchan, P. 2002a, *ApJ*, 569, 964
- Jura, M., Chen, C., & Plavchan, P. 2002b, *ApJ*, 574, 963
- Kahane, C., Audinos, P., Barnbaum, C., & Morris, M. 1996, *A&A*, 314, 871
- Keady, J. J., Hall, D. N. B., & Ridgway, S. T. 1988, *ApJ*, 326, 832
- Knapp, G. R., Jorissen, A., & Young, K. 1997, *A&A*, 326, 318
- Kwok, S., Purton, C. R., & Fitzgerald, P. M. 1978, *ApJ*, 219, L125
- Kwok, S., Volk, K., & Hrivnak, B. J. 2002, *ApJ*, 573, 720
- Lagage, P., Jouan, R., Masse, P., et al. 1993, *Proc. SPIE*, 1946, 655
- Lambert, D. L., Gustafsson, B., Eriksson, K., & Hinkle, K. H. 1986, *ApJS*, 62, 373
- Lloyd Evans, T. 2000, *IAUS*, 177, 367
- Loup, C., Forveille, T., Omont, A., & Paul, J. F. 1993, *A&AS*, 99, 291
- Mauron, N., & Huggins, P. J. 2000, *A&A*, 359, 707
- Mauron, N., de Laverny, P., & Lopez, B. 2003, *A&A*, 401, 985
- Men'shchikov, A. B., Balega, Y., Blöcker, T., Osterbart, R., & Weigelt, G. 2001, *A&A*, 368, 497
- Meixner, M., Ueta, T., Dayal, A., et al. 1999, *ApJS*, 122, 221
- Monnier, J. D., Tuthill, P. G., & Danchi, W. C. 2000, *ApJ*, 545, 957
- Rouleau, F., & Martin, P. G. 1991, *ApJ*, 377, 526
- Sahai, R., Zijlstra, A., Bujarrabal, V., & Te Lintel, H. P. 1999, *AJ*, 117, 1408
- Sahai, R., Morris, M., Knapp, G. R., Young, K., & Barnbaum, C. 2003, *Nature*, 426, 261
- Schmidt, G. D., Hines, D. C., & Swift, S. 2002, *ApJ*, 576, 429
- Steffen, M., Szczerba, R., Men'shchikov, A., & Schoenberner, D. 1997, *A&AS*, 126, 39
- Suh, K. 2000, *MNRAS*, 315, 740
- Trammell, S. R., Dinerstein, H. L., & Goodrich, R. W. 1994, *AJ*, 108, 984
- Tsuji, T., Unno, W., Kaifu, N., et al. 1988, *ApJ*, 327, 23
- Ueta, T., Meixner, M., Dayal, A., et al. 2001a, *ApJ*, 548, 1020
- Ueta, T., Meixner, M., Hinz, P. M., et al. 2001b, *ApJ*, 557, 831
- van der Veen, W. E. C. J., & Habing, H. J. 1988, *A&A*, 194, 125
- Waters, L. B. F. M., Cami, J., de Jong, T., et al. 1998, *Nature*, 391, 868
- Weigelt, G., Balega Y., Bloecker T., et al. 1998, *A&A*, 333, 51
- Weigelt, G., Balega, Y. Y., Blöcker, T., et al. 2002, *A&A*, 392, 131
- Winters, J. M., Dominik, C., & Sedlmayr, E. 1994, *A&A*, 288, 255