

Very high energy gamma rays from the composite SNR G 0.9+0.1

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Abstract. Very high energy (>100 GeV) gamma-ray emission has been detected for the first time from the composite supernova remnant G 0.9+0.1 using the HESS instrument. The source is detected with a significance of $\approx 13\sigma$, and a photon flux above 200 GeV of $(5.7 \pm 0.7_{\text{stat}} \pm 1.2_{\text{sys}}) \times 10^{-12} \text{ cm}^{-2} \text{ s}^{-1}$, making it one of the weakest sources ever detected at TeV energies. The photon spectrum is compatible with a power law ($dN/dE \propto E^{-\Gamma}$) with photon index $\Gamma = 2.40 \pm 0.11_{\text{stat}} \pm 0.20_{\text{sys}}$. The gamma-ray emission appears to originate in the plerionic core of the remnant, rather than the shell, and can be plausibly explained as inverse Compton scattering of relativistic electrons.

Key words. ISM: supernova remnants – ISM: individual objects: G 0.9+0.1 – gamma-rays: observations

1. Introduction

G 0.9+0.1 is a well known composite supernova remnant, recognised as such from its radio morphology (Helfand & Becker 1987). It exhibits a bright compact core ($\sim 2'$ across) surrounded by an $8'$ diameter shell. The radio spectrum of the core is significantly harder ($\alpha \approx 0.12$) than that of the shell ($\alpha \approx -0.77$) (LaRosa et al. 2000). X-ray observations of the nebula with BeppoSAX (Mereghetti et al. 1998), Chandra (Gaensler et al. 2001) and XMM-Newton (Porquet et al. 2003) have unambiguously identified the core region as a pulsar wind nebula (PWN). A plausible candidate for the central pulsar is the hard spectrum point source CXOU J174722.8–280915; however no pulsed emission has been detected. Observations with XMM have revealed a softening of the X-ray spectrum with increasing distance from the centre of the PWN, a signature of energy loss of electrons within the nebula. The location of G 0.9+0.1 in the Galactic Centre (GC) region suggests a distance of ~ 8.5 kpc (Mezger et al. 1996). At this distance the size of the shell implies an age of a few thousand years for the SNR. At gamma-ray energies G 0.9+0.1 has not been detected previously. It is not associated with any EGRET source and the only published observation in the TeV domain is an unconstraining upper limit from the HEGRA collaboration (flux $< 4.6 \times 10^{-12} \text{ cm}^{-2} \text{ s}^{-1}$ above 4.2 TeV, Aharonian et al. 2002).

The High Energy Stereoscopic System (HESS) is a newly completed instrument designed to study astrophysical gamma radiation in the energy range 100 GeV–10 TeV. HESS consists of four imaging atmospheric Cherenkov telescopes (Hinton 2004; Bernlöhr et al. 2003; Vincent et al. 2003; Aharonian et al. 2004a); shower images are required in at least two of them to trigger the detector (Funk et al. 2004). Using the stereoscopic technique the instrument reaches an angular resolution of $\sim 0.1^\circ$ and a point source sensitivity of 1% of the flux from the Crab Nebula (5σ in 25 h).

Observations of the GC region were made with the partially complete HESS system in 2003. These data revealed emission from the Sgr A region, consistent with the position of Sgr A* (Aharonian et al. 2004b). Within the same field of view indications for a signal from the supernova remnant G 0.9+0.1 were seen at the 4σ level. A relatively deep HESS observation of this region with the complete telescope system was performed in March–September 2004. The results from this dataset relevant to G 0.9+0.1 are presented here. A detailed discussion of the implications of the 2004 dataset for the region around Sgr A* will be presented elsewhere.

2. HESS observations and results

The large field of view of the HESS cameras ($\sim 5^\circ$) provides good sensitivity for point sources at an angular distance up to $\sim 2^\circ$ from the pointing direction of the telescope system. The bulk of observations presented here were taken in *wobble mode* around the position of Sgr A*. In this mode 28-min runs are taken pointing 0.5° away from the nominal source in alternating directions. Approximately 20% of the data were taken in wobble mode around G 0.9+0.1 itself. The observations have an average offset of 0.9° from the SNR. The off-axis sensitivity of the system derived from Monte Carlo simulations has been confirmed via observations of the Crab Nebula (Aharonian et al. 2004c). The total live time of the observations after run selection is 50 h, from a total observation time of 60 h. The mean zenith angle of observation was 18° , resulting in an energy threshold of 170 GeV after standard cuts.

Standard event reconstruction was applied to these data: tail-cuts image cleaning, Hillas parameterisation and stereoscopic direction reconstruction based on the intersection of image axes (Aharonian et al. 2005). Cuts on the scaled width and length of images (optimised on gamma-ray simulations and off-source data) are made to suppress the hadronic background. An additional cut on the minimum size of images (200 photoelectrons) is made to further suppress the background and select a subset of events with a superior angular resolution (0.07° rms), but with an increased energy threshold of 350 GeV. Events with directions reconstructed within an angle θ of a trial source location are considered on-source where $\theta < 0.1^\circ$. A ring of radius 0.5° and an area 7 times that of the on-source region is used to derive a background estimate. The on-source and background region counts, together with a normalisation factor for the different acceptance of these regions, are used to derive the statistical significance of any excess following the likelihood method of Li & Ma (1983).

Figure 1 shows a significance sky map for the field of view of the HESS GC observations derived as described above. Two sources of very high energy (VHE) gamma-rays are clearly visible: the Galactic Centre (HESS J1745–290) and a second source (HESS J1747–281) coincident with the SNR G 0.9+0.1. The statistical significance of HESS J1745–290 is 35σ in this dataset, compared with 11σ in our 2003 data (Aharonian et al. 2004b). This increase is expected (for a steady source) from the greater sensitivity of the full 4-telescope array with respect to the 2-telescope configuration used in 2003, and from the increased exposure time in 2004.

The HESS 2003 dataset showed evidence for gamma-ray emission at the position of G 0.9+0.1 at the 4σ level. This evidence is confirmed by the presence of a 13σ excess at the

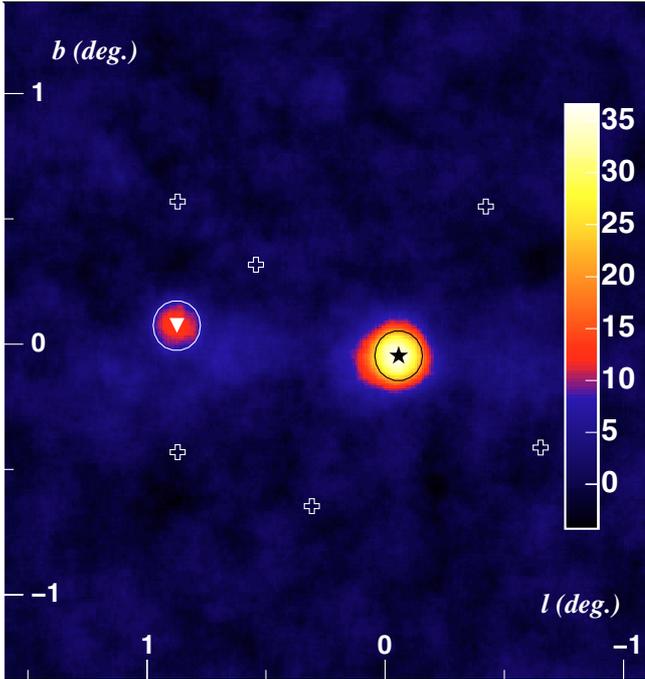


Fig. 1. Gamma-ray point source significance map for the region around the GC. The position of G 0.9+0.1 is marked with a triangle. Sgr A* is marked with a star. The circles show the integration regions contributing to the quoted significance at the positions of these two objects. The six telescope pointings are shown as crosses.

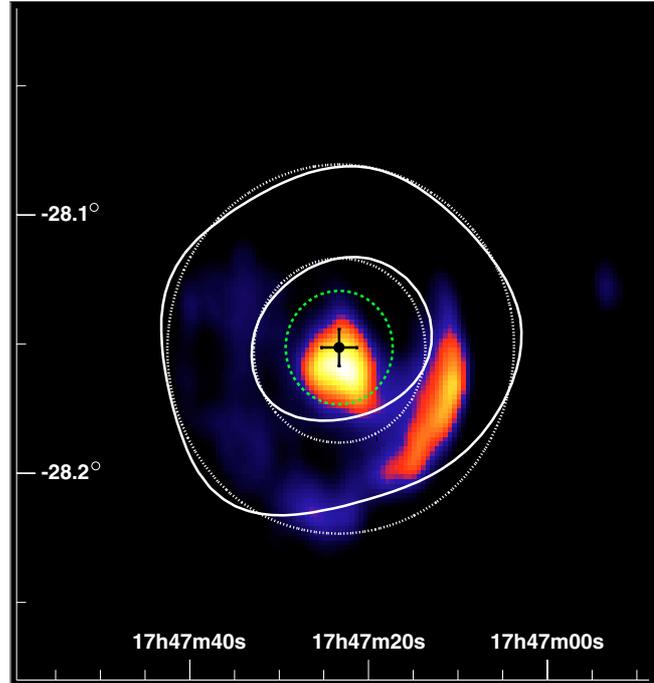


Fig. 2. 90 cm radio flux map of G 0.9+0.1 from LaRosa et al. (2000) (colour scale), overlaid with contours (at 40% and 80% peak brightness) of the smoothed and acceptance corrected count map of gamma-ray candidates (solid lines). The HESS data are smoothed with a Gaussian of rms 0.03° to reduce statistical fluctuations. The simulated point-spread function of the instrument is shown with dotted lines (also 40% and 80%). The best fit position of the VHE gamma-ray excess is shown with error bars showing combined statistical and systematic errors. The innermost circle (dashed) illustrates the 95% confidence limit on the rms size of the emission region.

same position in the independent 2004 dataset. The gamma-ray like excess is well fit by the point spread function of the instrument, yielding a best fit position of $l = 0.872^\circ \pm 0.005^\circ$, $b = 0.076^\circ \pm 0.005^\circ$, consistent within statistical errors with the position of the PWN in G 0.9+0.1 ($l = 0.871^\circ$, $b = 0.0772^\circ$) (Gaensler et al. 2001). The systematic pointing error is approximately 20'' in each direction for these data.

The simulated point-spread function of the instrument provides an adequate description of the excess. Assuming a radially symmetric Gaussian emission region ($\rho \propto \exp(-\theta^2/2\sigma_{\text{source}}^2)$) a 95% confidence limit on the extension $\sigma_{\text{source}} < 1.3'$ is obtained. For emission from a uniform thin shell a limit of 2.2' on the shell radius can be derived at the same confidence level. Figure 2 shows a radio map of G 0.9+0.1 at 90 cm (LaRosa et al. 2000) showing the SNR shell and the central compact source (colour scale). Contours of VHE gamma-ray emission are superimposed along with the point spread function of the instrument for comparison. The reconstructed gamma-ray source position is marked as a cross. Given the close positional coincidence, we identify the VHE gamma-ray source HESS J1747–281 with G 0.9+0.1.

Figure 3 shows the reconstructed gamma-ray spectrum of G 0.9+0.1 using two sets of event selection cuts. A size cut of 80 photoelectrons (pe) is applied to extend the spectrum to lower energies. This dataset can be fit by a power law in energy with photon index $2.40 \pm 0.11_{\text{stat}} \pm 0.20_{\text{sys}}$ and a flux above 200 GeV of $(5.7 \pm 0.7_{\text{stat}} \pm 1.2_{\text{sys}}) \times 10^{-12} \text{ cm}^{-2} \text{ s}^{-1}$. The fit has a $\chi^2/\text{d.o.f.}$ of 3.5/5. Using the harder image size cut of 200 pe used elsewhere in this paper yields consistent results (photon

index $2.29 \pm 0.14_{\text{stat}}$ and a flux of $(5.5 \pm 0.8_{\text{stat}}) \times 10^{-12} \text{ cm}^{-2} \text{ s}^{-1}$ above 200 GeV). This flux represents only 2% of the flux from the Crab Nebula above the 200 GeV spectral analysis threshold. Although G 0.9+0.1 is one of the weakest sources ever detected at TeV energies, the detection presented here is statistically highly significant (13 σ), demonstrating the sensitivity of the HESS instrument.

3. Discussion

X-ray emission from the shell of G 0.9+0.1 is only marginally detected above the local GC diffuse emission with a flux of $2 \times 10^{-12} \text{ erg cm}^{-2} \text{ s}^{-1}$ in the range 2–10 keV, and is consistent with both a thermal or non-thermal origin (Porquet et al. 2003). The non-thermal flux from the PWN is $5.8 \times 10^{-12} \text{ erg cm}^{-2} \text{ s}^{-1}$ in the same energy range, corresponding to a luminosity of $\sim 5 \times 10^{34} \text{ erg s}^{-1}$ assuming a distance of 8.5 kpc. The lack of strong non-thermal X-ray emission from the shell and the point-like nature of the HESS detection argue against an origin of the VHE emission in the shell. Thus the PWN appears to be a more compelling source for the VHE gamma-rays.

The total power radiated by G 0.9+0.1 in VHE gamma-rays is $\sim 2 \times 10^{34} \text{ erg s}^{-1}$ (at 8.5 kpc), compared with $\sim 4 \times 10^{34} \text{ erg s}^{-1}$ for the Crab Nebula (at 2 kpc) in the energy band 0.2–10 TeV.

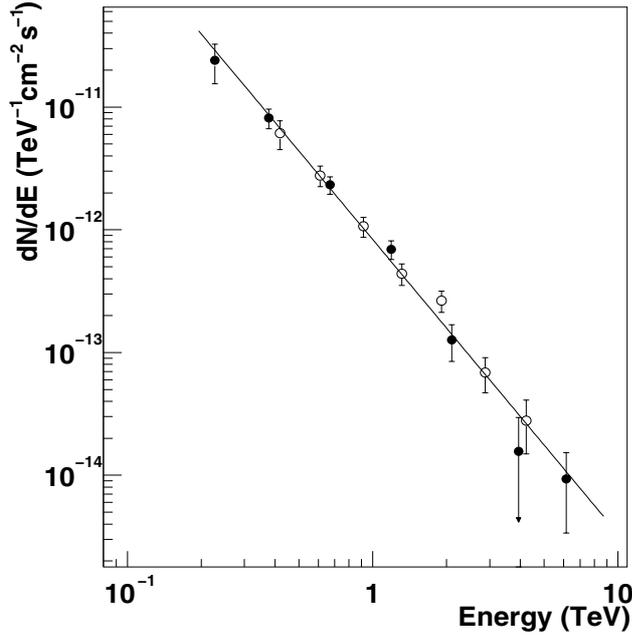


Fig. 3. Reconstructed VHE gamma-ray spectrum of G 0.9+0.1. Empty circles show the spectrum derived using the rather tight cuts used for Figs. 1 and 2. Solid symbols show the spectrum derived using an image size cut of 80 photoelectrons (optimised for maximal significance for 10% Crab strength sources) rather than 200 photoelectrons. The line shows a power-law fit to the 80 pe data set.

Until now the Crab Nebula represented the only firm identification of VHE emission from a nebula powered by a young neutron star.

A simple one-zone inverse Compton model for the VHE gamma-ray emission (Khélifi 2002) can be used to explain the observed data. A parent population of accelerated electrons with a broken power-law spectrum is assumed. The magnetic field strength B is assumed to be uniform within the PWN and is a free parameter. The maximum energy of the electron spectrum is kept fixed at 500 TeV as it can not be derived due to the absence of measurements in the hard X-ray/soft γ -ray band. The synchrotron and inverse Compton (IC) emission of these electrons are calculated according to Blumenthal & Gould (1970). The Klein-Nishina effect on the IC cross section is taken into account.

For sources in the GC region the presence of radiation densities considerably higher than typical Galactic Plane values leads to substantially enhanced IC emission (de Jager et al. 1995). The seed photons for IC scattering are the cosmic microwave background radiation (CMBR), the galactic diffuse emission around $100\mu\text{m}$ from dust and around $1\mu\text{m}$ from starlight. The energy spectra of the dust and starlight components are taken from a recent estimation of the interstellar radiation field (ISRF) used for the GALPROP code (Strong et al. 2000). To account for possible local variations of the starlight component the energy density of this component is kept as a free parameter. The energy density of the dust ISRF is kept fixed at 0.23 eV cm^{-3} as used in GALPROP. The starlight ISRF energy density, the electron spectrum parameters and B

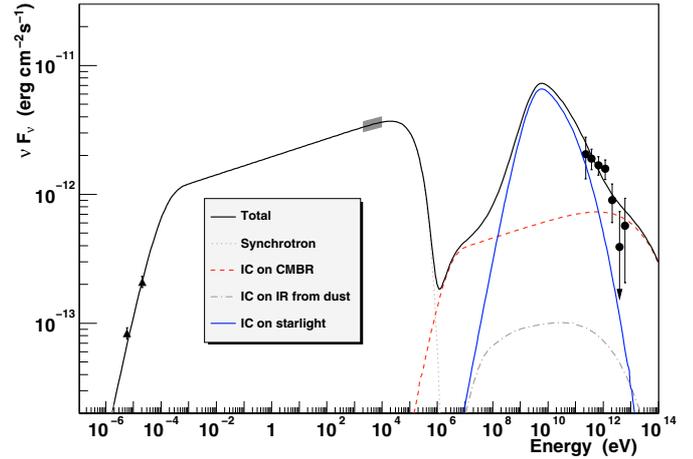


Fig. 4. The spectral energy distribution of the PWN in G 0.9+0.1 from radio to VHE gamma-ray. Radio data (triangles) are taken from (Helfand & Becker 1987), X-ray data (shaded box) from Porquet et al. 2003. The circles show the HESS data from this work. The solid curve shows a fit of a one-zone inverse Compton model to all data. Contributions of the CMBR, IR from dust and starlight photon fields to the IC emission are shown.

are adjusted such that the computed synchrotron and IC spectra match the observations.

Figure 4 shows the broad-band spectral energy distribution (SED) of G 0.9+0.1, together with the best model fit. The fitted mean magnetic field strength is $\approx 6\mu\text{G}$. The model electron spectrum has a low energy index of 0.6 and a high energy index of 2.9 with a break at 25 GeV. The best fit energy density of the starlight ISRF is about 5.7 eV cm^{-3} , that is 50% smaller than the value used in GALPROP. This derived value of B is very close to the equipartition magnetic field ($\approx 5\mu\text{G}$), calculated from the whole fitted electron spectrum and given a source of radius l' at 8.5 kpc. Given the observed variation of the X-ray spectrum within the PWN, this one zone model is clearly somewhat simplistic. However, such a model, describing the average properties of the PWN, can describe the available data with reasonable physical parameters. Future measurements in the GeV regime with GLAST (sensitivity $2 \times 10^{-12}\text{ erg s}^{-1}\text{ cm}^{-2}$ at 1 GeV) are needed to resolve the ambiguity between the magnetic energy density and that of target radiation fields in this region. Alternative explanations based on a hadronic origin of the VHE emission (for example via the decay of neutral pions produced in p - p collisions) are not excluded.

The detection of VHE gamma-ray emission from G 0.9+0.1 provides the first direct evidence for the acceleration of very energetic particles in this object. The position and point-like nature of the gamma-ray emission, combined with the broad band SED of G 0.9+0.1, strongly suggest an origin of the VHE emission in the compact central source (previously identified as a pulsar wind nebula).

This detection represents the first step towards the study of an emerging population of VHE gamma-ray emitting PWN. The detection of new PWN with accurate TeV spectra and with morphological measurements will provide information key to understanding particle acceleration in the vicinity of pulsars.

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