

Proton acceleration during coalescence of two parallel current loops in solar flares

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Abstract. We investigate the plasma dynamics during coalescence of two parallel current loops in solar flares by using a resistive three-dimensional MHD simulation code, paying particular attention finding the most effective electromagnetic fields for the production of high-energy protons. Next we investigate the orbit of many protons in the electromagnetic fields obtained from the MHD simulations. We investigate two cases of the coalescence process: (1) co-helicity reconnection where only the poloidal magnetic field produced from the axial currents dissipates and (2) counter-helicity reconnection where both poloidal and axial magnetic fields dissipate. We found a bump-on-tail distribution in the same direction as the original loop current for both the cases of co-helicity and counter-helicity. It is shown that the maximum proton energy exceeds the energy (2.223 MeV) of the observed prompt nuclear de-excitation lines of gamma-ray, and that proton energy spectrum is neither a pure power-law type nor a pure exponential type. We may conclude that anisotropic proton acceleration along the loop can be realized both for co-helicity and counter-helicity reconnection during the coalescence of two parallel loops.

Key words. acceleration of particles – magnetohydrodynamics (MHD) – Sun: flares

1. Introduction

When energetic protons accelerated during impulsive flares collide with the solar atmosphere, they produce excited nuclei, which emit prompt nuclear de-excitation lines, as well as secondary neutrons and positrons, which results in the delayed 2.223 MeV neutron-capture and 511 keV positron-annihilation line emission. In most strong flare events the time profile of the prompt gamma-ray line emission is observed to be very similar to that of the bremsstrahlung hard X-rays emitted by energetic electrons. This suggests that the acceleration and propagation of the flare-accelerated protons and electrons are closely related. The most typical event among these events is the 1980 June 7 flare observed by the SMM (Forrest & Chupp 1983), which was explained by the current loop coalescence model (Tajima et al. 1982; Sakai & Ohsawa 1987; Sakai & De Jager 1996; Aschwanden 2002). Sakai & de Jager (1996) gave a review of the high-resolution observations of solar plasma loops with simulations of current-carrying loops and tried to arrive at an understanding of solar flare phenomena. In Chapter 7 in their review, they summarized the understanding of high-energy particle acceleration processes during and after solar flares. There are two types of solar flares: Confined/Impulsive and Eruptive/Dynamic flares. Confined/Impulsive flares are spatially compact and as a fairly general rule they show impulsive hard X-ray and/or microwave bursts. Eruptive/Dynamic flares, on the other hand, are spatially more extended and last longer. As proposed by de Jager (1988)

and Sakai & de Jager (1991), there are three phases of acceleration in a fully developed eruptive/dynamic flare. The relevant observations regarding the first two phases are:

- (1) the acceleration times of electrons and protons to energies of the order of one MeV are of the order of a second;
- (2) MeV protons are accelerated nearly simultaneously with MeV electrons.

Many theoretical studies have been made to explain the above observational requirements (Forman et al. 1986; Simnett 1995; Sakai & de Jager 1996; Miller et al. 1997; Aschwanden 2002).

However, the location, size and geometry of the accelerated proton collision region has remained unknown until now. A recent paper by Hurford et al. (2003) presents the first gamma-ray images of a solar flare taken from the Reuven Ramaty High Energy Solar Spectroscopic Imager (RHESSI) for the X4.8 flare of 2002 July 23. The result shows that the centroid of the 2.223 MeV image was found to be displaced by 20 ± 6 arcsec from that of the 0.3–0.5 MeV, implying a difference in acceleration and /or propagation between the accelerated electron and proton population near the Sun. The fact that the proton-associated gamma-ray source does not coincide with the electron-bremsstrahlung sources suggests that the protons are accelerated in one direction by the DC electric field and can subsequently interact in spatially separate sources. Therefore it is important to investigate in detail the proton acceleration processes for different types of flares.

Sakai (1990) showed that, during 3D X-type current loop coalescence and under suitable assumptions for the size and other physical parameters in the region of acceleration, protons and electrons may be accelerated promptly (i.e., within less than 1 s) to ≈ 100 GeV and ≈ 100 MeV, respectively. De Jager & Sakai (1991) showed that the duration of impulsive phase bursts (5–25 s) observed during the impulsive phases of flares can be explained quantitatively by the mechanism of X-type current-loop coalescence. Sakai (1992) developed a model for long-duration gamma-ray/proton flares (the “gradual GR/P flares”) in order to explain prompt proton and electron acceleration during the impulsive phase. He determined the electric fields during the implosion phase of the current sheet from the MHD equations, and investigated the motion of test protons and electrons. The results showed that, under reasonable assumptions for the size and velocities in the reconnection area, both protons and electrons can be accelerated promptly within 1 s to ~ 70 MeV and ~ 200 MeV, respectively. However, the above studies did not give the energy spectrum of the accelerated particles. In the first phase of acceleration, current loop coalescence is one of the most plausible models of solar flares. Although the triggering mechanisms of most of the complex flares are still puzzling so far, several examples of solar flares triggered by current loop coalescence were observed by Yokoh (see for a review, Sakai & de Jager 1996). In the case of partial reconnection of kink-unstable loops coalescence (Nishikawa et al. 1994), it seems that electrons can be promptly accelerated to relativistic energies, while the problem of proton acceleration has not been investigated because of the limitation of full proton dynamics by using a particle simulation model.

Bulanov (1980) first showed the energy spectrum of accelerated particles in a simple X-type configuration near the reconnection region, which is given by $f(E) \sim \text{Exp}(-(E/E_0)^{3/4})$ for $E \ll mc^2$. Deeg et al. (1991) investigated the same problem using numerical simulation. Their results were in good agreement with the theory by Bulanov (1980).

Mori et al. (1998) investigated the behavior of protons near an X-type magnetic reconnection region by numerical simulations. The magnetic field is taken to be hyperbolic and time stationary with a uniform electric field perpendicular to the magnetic field. They also studied the effects of the magnetic field along the uniform electric field. From many parametric runs, they found that the energy spectrum of accelerated protons near an X-type magnetic reconnection region is universal with a power-law spectrum $E^{-\gamma}$, where the power-law index γ is about 2.0–2.2. The acceleration time of protons with an energy range of 1–20 MeV is very rapid and within $\sim 10^2 \omega_{ci}^{-1}$ which is much less than 1 s for solar flare plasmas. Heerikhuisen et al. (2002) and Craig et al. (2002) investigated the proton dynamics under a self-consistent MHD reconnection solution and found that the proton energy spectrum approximates a power law $E^{-1.5}$ nonrelativistically, but steepens slightly at higher energies. Hamilton et al. (2003) studied the proton energy spectra for the model fields obtained from solutions of the linearized MHD equations for the reconnecting region. They showed that in some cases the energy distributions exhibit a bump-on-tail in the MeV range, but in general the shape of the distribution is sensitive to the model parameters. The previous model

fields used to obtain the proton energy spectrum during coalescence process were limited by the simple spatial configurations near the reconnection region and by the lack of self-consistent MHD solutions.

In this paper we investigate the plasma dynamics during the coalescence of two parallel current loops using a resistive three-dimensional MHD simulation code developed by Sokolov et al. (1999, 2002), paying particular attention finding the most effective electromagnetic fields for the production of high-energy protons. This approach is very important for obtaining the proton energy spectrum that depends on the model electromagnetic fields, as shown by Mori et al. (1998) and Hamilton et al. (2003). Next we investigate the orbit of many protons in the electromagnetic fields obtained from the MHD simulations. We investigate two cases of the coalescence process: (1) co-helicity reconnection where only the poloidal magnetic field produced from the axial currents dissipates; and (2) counter-reconnection where both poloidal and axial magnetic fields dissipate. We found that for the proton acceleration co-helicity reconnection is more efficient than counter-helicity reconnection. The protons can be accelerated mostly in one direction along the loop axis near where magnetic reconnection occurs. There appears a weak bump-on-tail proton distribution in the middle energy region, not in the high-energy region as shown by Nagao et al. (2000) and Hamilton et al. (2003). Therefore the proton energy spectrum is neither pure power-law type nor pure exponential type.

In Sect. 2 we present our methods to obtain the proton energy spectra during coalescence of two parallel current loops. In Sect. 3 we present our MHD simulation results to obtain the most effective electric fields for accelerating the protons. In Sect. 4 we investigate many proton orbits under the electromagnetic fields during two types of coalescence to find their energy spectra. In the final section we summarize our results.

2. Simulation methods

In this section we describe two simulation methods for obtaining the energy spectrum of the protons accelerated during the coalescence of two parallel current loops. Firstly we briefly present a model of a two-loop coalescence flare by using a three-dimensional resistive MHD simulation to obtain the electromagnetic fields during the loop coalescence. Next we present the simulation method to obtain the orbits of many protons in the electromagnetic fields obtained from the MHD simulations.

2.1. Resistive MHD simulation

We use a resistive 3-D MHD code with the recently proposed *Artificial Wind* (AW) numerical scheme (Sokolov et al. 1999, 2002) with splitting over the spatial coordinates. The basis of the AW scheme is the fact that fundamental physical invariance, Galilean (or more general, Lorentz) invariance, make it possible to solve the governing equation in different steadily moving frames. The principle of the AW scheme is that the frame of reference may be chosen in such a way that the flow under simulation is supersonic. The problem of

upwinding becomes trivial and considerably simplified versions of discontinuity-capturing schemes may be employed. An extra velocity (*Artificial Wind*) is added to the velocity of the flow under simulation when the system of coordinates is changed.

The MHD equations are numerically solved in a conservative form as follows:

$$\frac{\partial \rho}{\partial t} + \frac{\partial}{\partial x_i}(\rho V_i) = 0, \quad (1)$$

$$\frac{\partial(\rho V_i)}{\partial t} + \frac{\partial}{\partial x_j} \left[\rho V_i V_j + (p + B^2) \delta_{ij} - 2B_i B_j \right] = 0, \quad (2)$$

$$\frac{\partial B_i}{\partial t} + \frac{\partial}{\partial x_j} (V_j B_i - V_i B_j) = \frac{1}{R_m} \frac{\partial^2 B_i}{\partial x_j^2}, \quad (3)$$

$$\frac{\partial}{\partial t} \left(\frac{\rho V^2}{2} + \frac{p}{\gamma - 1} + B^2 \right) + \frac{\partial}{\partial x_i} \left[V_i \left(\frac{\rho V^2}{2} + \frac{\gamma p}{\gamma - 1} + 2B^2 \right) - 2B_i B_j V_j + q_i \right] = 0, \quad (4)$$

where ρ , V_i , p and B_i are the density, velocity, pressure, and magnetic field, γ is the adiabatic constant, which is taken to be $\gamma = 5/3$. $R_m = 4\pi\sigma c_s L/c^2$ is the magnetic Reynolds number, where $c_s = \sqrt{p_0/\rho_0}$, σ the electric conductivity and L the grid size. δ_{ij} is a unity tensor and q_i is the dissipative energy flux. The density, pressure, velocity and magnetic field are normalized by ρ_0 , p_0 , c_s and $B_0 = \sqrt{8\pi p_0}$, respectively.

Magnetic reconnection takes place through finite conductivity. The conductivity due to microturbulence (Vasyliunas 1975) may be written by the expression

$$\sigma = \omega_{pe}^2 / 4\pi\nu_{\text{eff}}, \quad (5)$$

where ν_{eff} is the effective collision frequency between some high-frequency turbulent fields and the electrons. For strong turbulence driven in the current sheet due to electric currents like Buneman-type instability the effective collision frequency is of the order of the electron plasma frequency, namely $\nu_{\text{eff}} \approx \omega_{pe}$. Therefore, if we take the typical density in the corona as 10^8 , $c_s \approx 10^7$ cm s⁻¹ and $L \approx 10^7$ cm, we obtain $R_m \approx 10^{3-4}$.

For resistive MHD with a large but finite value of R_m the energy equation Eq. (4) should be obtained as the sum of the equation for plasma energy, in which the Joule heating term is present as follows: $\frac{(c \text{rot} \mathbf{B})^2}{(4\pi)^2 \sigma}$, and the equation for magnetic energy which is given by Eq. (3) multiplied by the magnetic field $B_i/(4\pi)$ (none of the variables is normalized here). The dissipation of the magnetic field energy is hence given by the following term: $\frac{c^2}{(4\pi)^2 \sigma} B_i \Delta B_i$. So in Eq. (4) for the total energy the Joule heating is compensated by the magnetic energy dissipation in the following way:

$$\frac{c^2}{(4\pi)^2 \sigma} [B_i \Delta B_i + (\text{rot} \mathbf{B})^2] = \frac{c^2}{(4\pi)^2 \sigma} \text{div}[\mathbf{B} \times \text{rot} \mathbf{B}].$$

So the resistive dissipation in the equation of energy appears only in the form of an additional dissipative energy flux, which, in normalized variables, may be written as

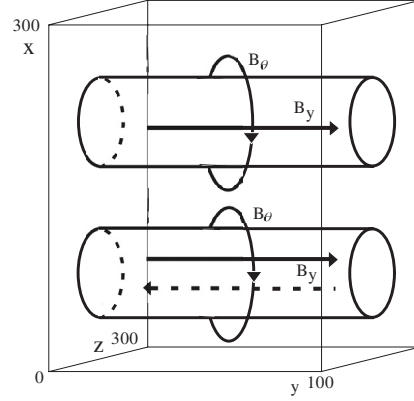


Fig. 1. A schematic picture of two parallel current loops and the coordinate system.

$q_i^{(m)} = R_m^{-1} \frac{\partial}{\partial x_j} (2B_i B_j - B^2 \delta_{ij})$. The dissipative energy flux due to heat transfer may also be taken into account in the usual form: $q_i^{(h)} = \lambda_{ij} \frac{\partial}{\partial x_j} \frac{p}{\rho}$. For the sake of simplicity here we fully neglect the non-diagonal part of the dissipative energy transfer tensor and substitute all the tensors in q_i for unity tensors: $B_i B_j = B^2 \delta_{ij}/3$, $\lambda_{ij} = \lambda_{||} \delta_{ij}/3$. We used the numerical scheme of the Alternating Direction method to solve the above three-dimensional MHD equations. The numerical simulation shows that the influence of the dissipative energy flux is insignificant, so we do not try to take it more carefully into account. Finally we let the dissipative hydrodynamic flux be as follows:

$$q_i = \frac{\partial}{\partial x_i} \left(-\frac{B^2}{3R_m} - \lambda \frac{p}{\rho} \right). \quad (6)$$

The magnetic Reynolds number $R_m = 1.3 \times 10^3$ and the heat transfer constant $\lambda = 2.5 \times 10^{-4}$ are used in the simulation. Radiating boundary conditions were used for all directions.

Magnetic reconnection during the coalescence of two current loops produces strong electric fields in the plasma and, as a result, charged particles can be accelerated. In this paper we concentrate on the proton dynamics around an X-point. During two current loops coalescence, there are two types of magnetic reconnection possible: one is complete reconnection (we call here counter-helicity case), where the B_y component is null, while for partial magnetic reconnection (we call here co-helicity case) the B_y component is not zero. Here we take a system size of $N_y = 100$ and $N_x = N_z = 300$, as shown in Fig. 1.

Each current loop ($i = 1, 2$) located parallel to the y -axis is assumed to be in an equilibrium state ($\nabla p_i = \mathbf{J}_i \times \mathbf{B}_i/c$). This means that the initial state with two current loops is not in equilibrium. The magnetic field and pressure for each loop are given as

$$B_x = q_i B_y (z - z_{ci})/a, \quad (7)$$

$$B_y = B_{0i} e^{-(z/a)^2}, \quad (8)$$

$$B_z = -q_i B_y (x - x_{ci})/a, \quad (9)$$

$$p_i = \left(\frac{q_i^2}{2} - \frac{q_i^2 r_i^2}{a^2} - 1 \right) e^{-2(z/a)^2} + 0.55, \quad (10)$$

where $r_i = \sqrt{(x - x_{ci})^2 + (z - z_{ci})^2}$, and the centers of the two flux tubes with radius $a = 30$ are $(x_{c1}, z_{c1}) = (105, 150)$ and $(x_{c2}, z_{c2}) = (195, 150)$. The density profile is the same as the pressure profile. The twist parameter q_i is $q_1 = q_2 = 1$ and $B_{01} = B_{02} = 1.0$ for co-helicity, while $B_{01} = 1.0, B_{02} = -1.0$ for counter-helicity. The plasma beta β in the center of the two loops is 0.06. Radiating boundary conditions were used for all directions. We tried to do several different simulations by changing the system size from $N_x = N_y = N_z = 200$ to $N_x = N_z = 300$ and $N_y = 100$ as well as the loop radius from 20 to 30. We found that the results are almost the same.

2.2. Simulation of proton dynamics

Using the electromagnetic fields obtained from the above MHD equations during the coalescence of two parallel current loops, we investigate the proton dynamics to obtain their energy spectrum. The normalized relativistic equations of the motion of a proton are given by

$$\frac{d\mathbf{u}}{dt} = \mathbf{E} + \frac{\mathbf{u} \times \mathbf{B}}{\gamma}, \quad (11)$$

$$\frac{d\mathbf{x}}{dt} = \frac{R\mathbf{v}}{\gamma}, \quad (12)$$

where the proton velocity $\mathbf{v} = \gamma^{-1}\mathbf{u}$, the electric field \mathbf{E} , and the magnetic field \mathbf{B} are normalized by Alfvén velocity V_A , $E_0 = V_A B_0/c$, and B_0 , respectively. The Alfvén velocity V_A is calculated by using the initial magnetic field (B_0) and initial density at the center of the loop. The time is normalized by the proton cyclotron frequency $\omega_{ci} = eB_0/m_i c$, and the length is normalized by the distance over which the fields are changed across the loop radius (a), respectively. The Lorentz factor γ is given by $\gamma = (1 + A^2\mathbf{u}^2)^{1/2}$, where $A = V_A/c$. The parameter R in Eq. (12) is defined as $R = V_A/(\omega_{ci}a)$. In our simulation we take $A = 1/300$ and $R = 10^{-8}$.

To solve numerically the equation of motion of a test particle in an electric and magnetic field, we used a centered-difference form of the Newton-Lorentz equations of motions (Buneman 1993; Birdsall & Langdon 1991), namely

$$\mathbf{u}^{\text{new}} - \mathbf{u}^{\text{old}} = \frac{q\delta t}{m} \left[\mathbf{E} + \frac{1}{2\Gamma} (\mathbf{u}^{\text{new}} + \mathbf{u}^{\text{old}}) \times \mathbf{B} \right], \quad (13)$$

where q is the charge of the particle, m is the mass of the particle, and δt is time step in the calculations, respectively. As it stands, the equation for \mathbf{u}^{new} is implicit. We choose the method proposed by Boris (1970) to obtain a simpler explicit solution using several steps.

The protons are randomly distributed with two protons per cell. We changed the test particle number just twice, namely from one to two particles per each cell. We found that the results are almost the same. The initial velocity distribution function for protons is Maxwellian with thermal velocity $v_{\text{th}} = 0.4V_A$. The simulation time step is $\omega_{ci}\delta t = 0.05$. The boundary conditions for protons are periodic in the y -direction and open in the x - and z -directions.

3. Simulation results of two loops coalescence

The basic physical evolution of reconnection between two loops with parallel axial currents is known to be as follows (see for a review, Sakai & de Jager 1996). Each of the two current loops is in an equilibrium state as long as the distance between them is much greater than the loop radius a . While the loops approach to each other, they are no longer in the equilibrium state and the whole non-equilibrium system tends to a new equilibrium state. In more detail, due to attraction, the two initially static current loops begin to move and approach to each other. Then they meet, merge, and form a new single loop. Our present simulation is performed to investigate the phenomena accompanying this reconnection process.

The simulation result for the co-helicity reconnection can be described as follows. In the early stage, while the two loops begin to approach each other, the reconnection of the poloidal magnetic field is important. As the two current loops approach each other, the original magnetic field lines break and join to new ones, the reconnected magnetic field lines being strongly bent. Due to magnetic stress, the magnetic field lines tend to shorten and drive the surrounding plasma moving towards in the z and $-z$ directions. Thus a pair of jet-like flows is formed.

On the contrary, for the counter-helicity which proceeds in the form of a comparatively fast oscillation, the period of the oscillations is governed by MHD processes and is comparable with the Alfvén transient time. This means that the dynamics of the system enter the axial field reconnection phase.

First we present the simulation results of coalescence of the two current loops for the co-helicity and counter-helicity cases, paying particular attention to the time period on which the induced electric field becomes strongest in order to find the most efficient proton acceleration phase. In Fig. 2 we show the time evolution of magnetic fields ((a), (b), (c)) and electric fields ((d), (e), (f)) in the $x - z$ plane for $y = 50$ for co-helicity at three different times: $3.8 \tau_A$ ((a) and (d)), $4.7 \tau_A$ ((b) and (e)) and $6.6 \tau_A$ ((c) and (f)). The gray scale in Figs. 2a–c shows the spatial distribution of the magnetic field B_y along the current loop, while the gray scale in Figs. 2d–f shows the spatial distribution of the electric field E_y along the current loop. The vectors in the $x - z$ plane show the $B_x - B_z$ in Figs. 2a–c and the $E_x - E_z$ in Figs. 2d–f, respectively.

As seen in Figs. 2a–c, the reverse current that flows against the original loop currents (positive y -direction) is generated in the middle region of the two loops as they approach each other. The magnetic reconnection is triggered in the induced current region whose shape becomes elongated in the z -direction, as seen from the structure of the magnetic vector fields in Fig. 2c. If the original loop currents were stronger than at present case, the elongated current structure in the x -direction is again formed with some oscillation period. As seen in Fig. 2e, the strongest electric fields can be induced just where the two loops begin to merge with magnetic reconnection at about $t = 4.7\tau_A$. The electric field in the $x - z$ plane is characterized by a quadrupole like structure. Therefore we will use the electromagnetic fields (Figs. 2b and 2e) at $t = 4.7\tau_A$ to study the proton acceleration for co-helicity reconnection.

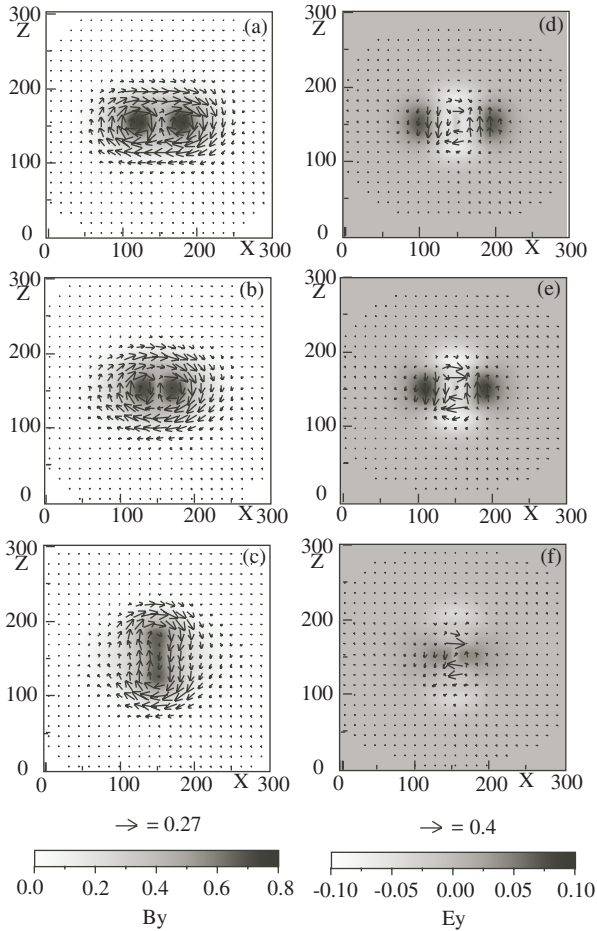


Fig. 2. The time evolution of magnetic fields (a, b, c) and electric fields (d, e, f) in the $x-z$ plane for $y = 50$ for the co-helicity case at three different times: $3.8 \tau_A$ (a) and (d), $4.7 \tau_A$ (b) and (e) and $6.6 \tau_A$ (c) and (f). The gray scale in (a)–(c) shows the spatial distribution of the magnetic field B_y along the current loop, while the gray scale in (d)–(f) shows the spatial distribution of the electric field E_y along the current loop. The vectors in the $x-z$ plane show the $B_x - B_z$ in (a)–(c) and the $E_x - E_z$ in (d)–(f), respectively.

A remarkable phenomenon in the counter-helicity case (Zhang et al. 2001) is the oscillation of the axial component of the magnetic field B_y , and the velocity V_y . The reconnection and oscillation mechanism can be clarified as follows. As the two current loops approach each other, the B_y reconnection begins. After reconnection (the original magnetic field lines break and join new ones) the magnetic field lines appear to be strongly bended. Due to magnetic stress, the magnetic field lines tend to shorten and drive the surrounding plasma motion in the y or $-y$ directions. As the magnetic field lines shorten, the surrounding plasma gets momentum in the y -direction. After the magnetic field lines have reached their balance point, the moving surrounding plasma begins to drive and stretch the magnetic field lines. Thus, the interaction between magnetic field and surrounding plasma leads to the oscillation.

In Fig. 3 we show the time evolution of magnetic fields (Figs. 3a–c) and electric fields (Figs. 3d–f) in the $x-z$ plane on the plane $y = 50$ for the counter-helicity case at three different times: $3.8 \tau_A$ (Figs. 3a and d), $4.7 \tau_A$ (Figs. 3b and e)

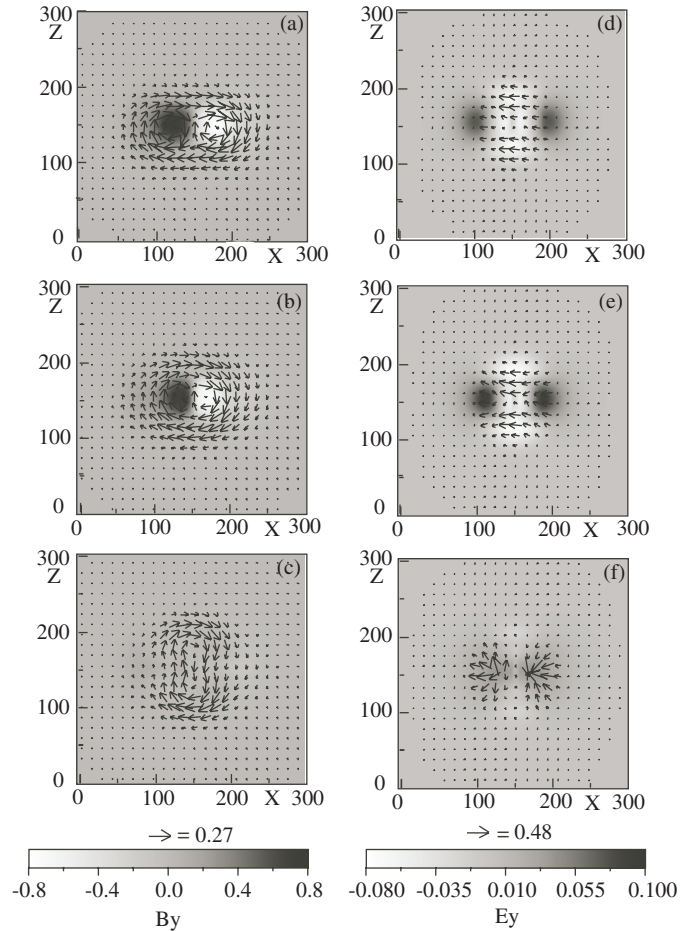


Fig. 3. The time evolution of magnetic fields (a, b, c) and electric fields (d, e, f) in the $x-z$ plane for $y = 50$ for the counter-helicity case at three different times: $3.8 \tau_A$ (a) and (d), $4.7 \tau_A$ (b) and (e) and $6.6 \tau_A$ (c) and (f). The gray scale in (a)–(c) shows the spatial distribution of the magnetic field B_y along the current loop, while the gray scale in (d)–(f) shows the spatial distribution of the electric field E_y along the current loop. The vector in the $x-z$ plane show the $B_x - B_z$ plot in (a)–(c) and the $E_x - E_z$ plot in (d)–(f), respectively.

and $6.6 \tau_A$ (Figs. 3c and f). The gray scale in Figs. 3a–c shows the spatial distribution of the magnetic field B_y along the current loop, while the gray scale in Figs. 3d–f shows the spatial distribution of the electric field E_y along the current loop. The vectors in the $x-z$ plane show the $B_x - B_z$ in Figs. 3a–c and the $E_x - E_z$ in Figs. 3d–f, respectively. As seen in Fig. 3e, the strongest electric fields can be induced just where the two loops begin to merge with magnetic reconnection at about $t = 4.7 \tau_A$. The electric field in the $x-z$ plane perpendicular to the loops is dominant in the negative x -direction. Therefore we will use the electromagnetic fields (Figs. 3b and 2e) at $t = 4.7 \tau_A$ to study the proton acceleration for counter-helicity reconnection. The electric field strength (E_y) along the loops for the counter-helicity case is slightly weaker than that for the co-helicity case.

4. Proton dynamics and its energy spectra

In this section we investigate the proton dynamics in the electromagnetic fields obtained from the resistive MHD simulation

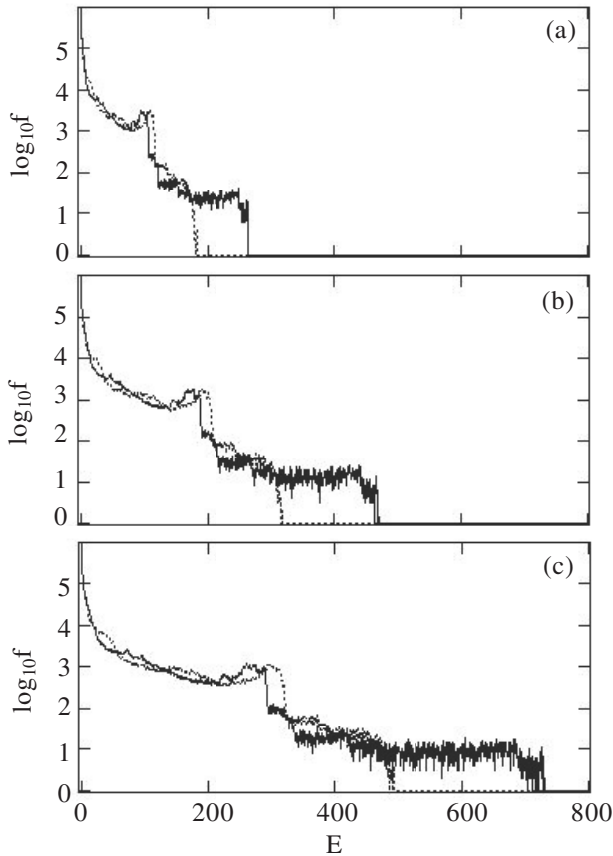


Fig. 4. Proton energy spectra at three different times, **a)** at $\omega_{ci}t = 1500$, **b)** at $\omega_{ci}t = 2000$ and **c)** at $\omega_{ci}t = 2500$. The solid and dotted curves show the co-helicity and counter-helicity case, respectively. $E = (V_x^2 + V_y^2 + V_z^2)/V_A^2$.

for both cases of co-helicity and counter-helicity reconnection. In Fig. 4 we show the proton energy spectra at three different times: at $\omega_{ci}t = 1500$ (Fig. 4a), at $\omega_{ci}t = 2000$ (Fig. 4b) and at $\omega_{ci}t = 2500$ (Fig. 4c). The solid and dotted curves show the co-helicity and counter-helicity case, respectively. $E = (V_x^2 + V_y^2 + V_z^2)/V_A^2$.

In contrast with the result shown by Mori et al. (1998) who showed that the energy distribution of the accelerated protons has a power-law spectrum, we find that the proton energy spectrum is more complicated. There appears a weak bump-on-tail in the medium energy region, not in the high-energy tail shown by Hamilton et al. (2003). The reason may be related to the fact that the induced electric fields are not simple as compared with the model fields used by Mori et al. (1998) and Hamilton et al. (2003). Except for the weak bump-on-tail, the global proton spectrum seems to be exponential as predicted by Bulanov (1980). The fact that the production of high-energy protons in the counter-helicity case is less weak than in the co-helicity case is mainly due to the following reason. As shown by Mori et al. (1998) and Hamilton et al. (2003), the magnetic field component along the reverse current is more effective for high-energy proton production, because the guiding magnetic field plays a role in keeping the protons within the accelerating region for a long time, while in a weak guiding field like in the counter-helicity case the accelerated protons may escape from

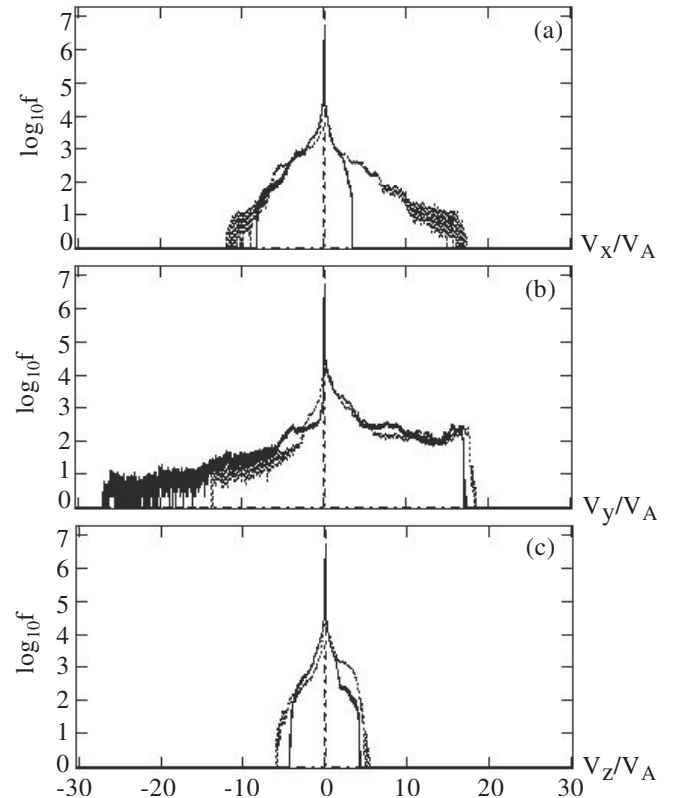


Fig. 5. The proton velocity distribution functions of three components, **a)** V_x , **b)** V_y and **c)** V_z . The dot-dashed, dotted and solid curves show the initial proton velocity distribution, and the proton velocity distribution for the counter-helicity case and for the co-helicity case at $\omega_{ci}t = 2500$, respectively.

the strong electric field region. The production of the proton bump-on-tail in the medium energy region is also related to the electric fields perpendicular to the reverse current, as shown by the test proton simulation (Nagao et al. 2000) including both parallel electric field (E_y) and perpendicular electric fields (E_x and E_z).

Next we show the proton velocity distributions to see the details of the direction of acceleration of the protons. In Fig. 5 we present the proton velocity distribution functions for the three components, (a) V_x , (b) V_y and (c) V_z . The dot-dashed, dotted and solid curves show the initial proton velocity distribution, proton velocity distribution for the counter-helicity case and for the co-helicity case at $\omega_{ci}t = 1500$, respectively. As seen in Fig. 5b, strong anisotropic acceleration in the negative y -direction is observed for the co-helicity case, while for the counter-helicity case, almost symmetric acceleration in the y -direction occurs and the maximum proton velocity in the y -direction is small compared with what it is in the co-helicity case. The proton bump-on-tail appears in the positive y -direction for both co-helicity and counter-helicity, as seen in Fig. 5b. Therefore the origin of the bump-on-tail may be related to the acceleration by the combination of three electric field components as shown by Nagao et al. (2000), not simply only to the acceleration by the electric field parallel to the reverse current.

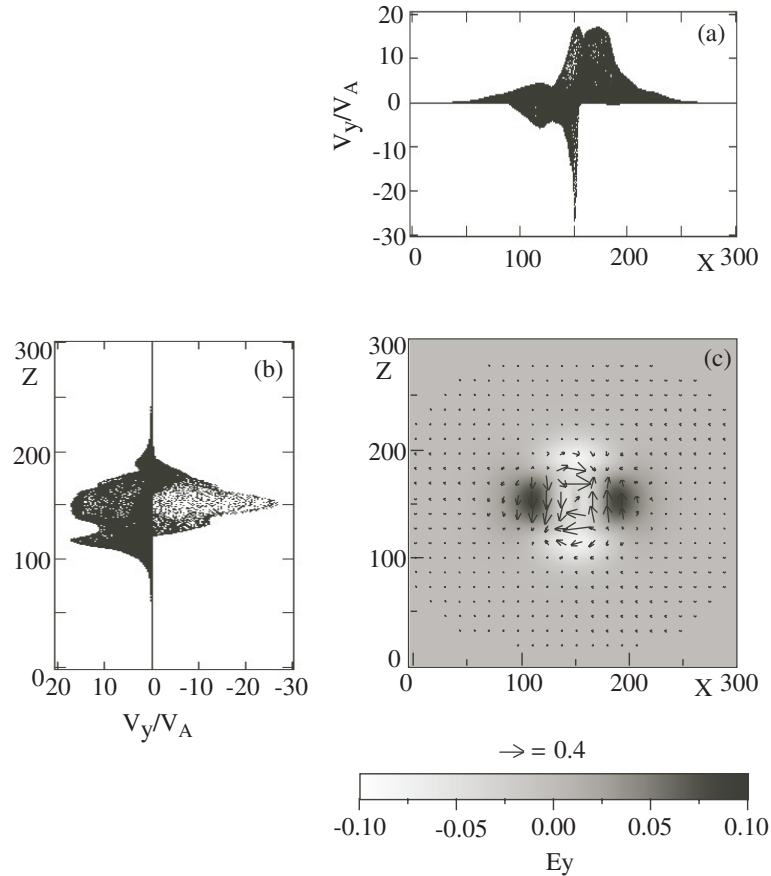


Fig. 6. The phase-diagrams of protons: **a)** the $x - V_y$ plane and **b)** the $z - V_y$ plane at $\omega_{ci}t = 2500$. **c)** The spatial distribution of the electric field E_y with gray scale and the vector plots $(E_x - E_z)$ at $4.7 \tau_A$ for the co-helicity case.

Next we investigate the location where the high-energy protons can be accelerated. We show the phase-diagrams of protons of the $x - V_y$ plane in Fig. 6a, and the $z - V_y$ plane in Fig. 6b at $\omega_{ci}t = 2500$ for the co-helicity case. In Fig. 6c we show the spatial distribution of the electric field E_y with gray scale and vector plots $(E_x - E_z)$ at $4.7 \tau_A$ for the co-helicity case. We found that the high-energy protons in the negative y -direction can be produced near the center of the reversed current sheet and the medium energy proton in the positive y -direction can be accelerated near the region where the electric field E_z is strong in the positive z -direction. Therefore the protons with medium energy can be produced mostly in one loop such as the right-hand loop seen in Fig. 6c.

In Figs. 7a and b we show the phase-diagrams of protons of the $x - V_y$ plane and the $z - V_y$ plane at $\omega_{ci}t = 2500$ for the counter-helicity case, respectively. Fig. 7c shows the spatial distribution of the electric field E_y with gray scale and vector plots $(E_x - E_z)$ at $4.7 \tau_A$ for the counter-helicity case. As in Fig. 7a, the proton acceleration near the center of the reverse current sheet is weak compared with the previous co-helicity case. The reason is due to the fact that the guiding magnetic field B_y near the reverse current sheet is weak by its magnetic reconnection. As seen in Figs. 7a and b, the high-energy protons can be accelerated in the positive y -direction mainly near the inner part of the right-handed loop. Therefore, for the

counter-helicity case there is no strong anisotropic proton acceleration in the y -direction along the loop.

Since the aim of this study is to investigate the acceleration of protons by magnetic reconnection in solar flares, we establish physical links between the scaling of the basic parameters in the simulations and solar flare conditions. As the initial given value of $A = V_A/c = 1/300$, the Alfvén velocity is about 1000 km s^{-1} . Also, the proton thermal velocity is assumed to be equal to $0.4V_A$, or 400 km s^{-1} , which corresponds to the $\sim 1.6 \text{ keV}$ ambient ions. The temperature is acceptable for a well-developed flare in the preflare phase. Those $\sim 1.6 \text{ keV}$ ions can be accelerated to the MeV energy range, whose maximum is about $\sim 7 \text{ MeV}$ for co-helicity reconnection and about $\sim 5 \text{ MeV}$ for counter-helicity reconnection, respectively. Therefore protons accelerated during co-helicity reconnection can produce the observed prompt nuclear de-excitation lines of gamma-ray. The bump-on-tail proton energy corresponds to about 1 MeV .

5. Conclusions

We have investigated the behavior of protons near an X-type magnetic reconnection region during the coalescence of two parallel loops for the two cases of co-helicity and counter-helicity. The electromagnetic fields during the coalescence

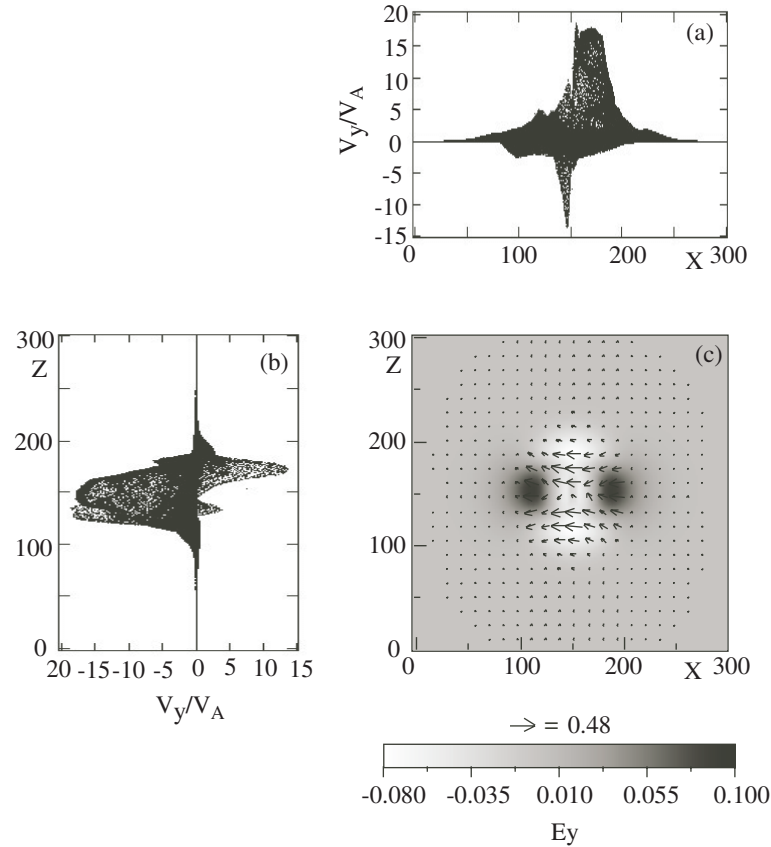


Fig. 7. The phase-diagrams of protons: **a)** the $x - V_y$ plane and **b)** the $z - V_y$ plane at $\omega_{ci}t = 2500$. **c)** The spatial distribution of the electric field E_y with gray scale and the vector plots ($E_x - E_z$) at $4.7 \tau_A$ for the counter-helicity case.

process were calculated from a three-dimensional resistive MHD simulation, to find the most effective electric fields for the proton acceleration. In contrast to the result by Mori et al. (1998) who showed that the energy spectrum of accelerated protons near X-type magnetic reconnection regions is universal with a power-law spectrum $E^{-\gamma}$, where the power-law index γ is about 2.0–2.2, the proton energy spectrum found by us is neither purely exponential nor purely power-law, because of the more complicated structure of the electromagnetic fields during coalescence process. We found a bump-on-tail distribution in the positive y -direction along the original loop current for both the cases of co-helicity and counter-helicity. While the high-energy protons producing the observed prompt nuclear de-excitation gamma-ray lines can be produced mostly in the direction opposite to the original loop current near the reverse current sheet only for co-helicity reconnection. Therefore we may conclude that the anisotropic proton acceleration along the loop can be realized both for co-helicity and counter-helicity reconnection during two parallel loops coalescence. This result should be very important for understanding the recent gamma-ray images taken from the RHESSI observation (Hurford et al. 2003). The observation shows that the centroid of the 2.223 MeV image was found to be displaced by 20 ± 6 arcsec from that of the 0.3–0.5 MeV, implying a difference in acceleration and/or propagation between the accelerated electron and proton population near the Sun. The fact that the proton-associated gamma-ray source does not coincide

with the electron-bremsstrahlung sources suggests that the protons are accelerated in one direction by the DC electric field and could subsequently interact in spatially separated sources. One possible scenario to explain the above observation is as follows. The single loop with $\beta = 0.5$ in the center of the loop is supposed to be disrupted by the mechanism recently proposed by Sakai et al. (2002), then the disrupted part of the loop with high energy protons as well as hot thermalized protons could move up and interact with the overlying other loop. The interaction between the ascending magnetized plasma blobs and the other loop can lead to magnetic reconnection in the interaction region, where the proton could be accelerated further by the inductive electric field associated with the magnetic reconnection mostly in one direction along the guiding magnetic field.

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References

- Aschwanden, M. J. 2002, *Space Sci. Rev.*, 101, Nos. 1–2, 1
- Birdsall, C. K., & Langdon, A. B. 1991, *Plasma Physics via Computer Simulation* (Adam Hilger), 58

- Boris, J. P. 1970, Proc. Fourth Conf. Num. Sim. Plasmas, Naval Res. Lab., Wash. DC, 3–67, 2–3 November
- Bulanov, S. V. 1980, SvA Lett., 6, 206
- Buneman, O. 1993, in Computer Space Plasma Physics, Simulation Techniques and Software, ed. H. Matsumoto, & Y. Omura (Tokyo: Terra Scientific), 67
- Craig, I. J. D., & Litvinenko, Y. E. 2002, ApJ, 570, 387
- Deeg, H. J., Borovsky, J. E., & Duric, N. 1991, Phys. Fluids B, 3, 2660
- De Jager, C. 1988, Proc. 20th Cosmic Ray Conf., Moscow, 7, 66
- De Jager, C., & Sakai, J. I. 1991, Sol. Phys., 133, 395
- Forman, M. A., Ramaty, R., & Zweibel, E. G. 1986, Chapter 13 in Physics of the Sun, Vol. 2, ed. P. A. Sturrock, T. E. Holzer, D. M. Mihalas, & R. K. Ulrich (D. Reidel Publishing Co.)
- Forrest, D. J., & Chupp, E. L. 1983, Nature, 305, 291
- Hamilton, B., McClements, K. G., Fletcher, L., & Thyagaraja, A. 2003, Sol. Phys., 214, 339
- Heerikhuisen, J., Litvinenko, Y. E., & Craig, I. J. D. 2002, ApJ, 566, 512
- Hurford, G. J., Schwartz, R. A., Krucker, S., et al. 2003, ApJ, 595, L77
- Miller, J. A., Cargill, P. J., Emslie, G., et al. 1997, J. Geophys. Res., 102/A7, 14631
- Mori, K., Sakai, J. I., & Zhao, J. 1998, ApJ, 494, 430
- Nagao, W., Sakai, J. I., & Bulanov, S. 2000, Conference 2000 on Magnetic Reconnection, Tokyo
- Nishikawa, K. I., Sakai, J. I., Zhao, J., Neubert, T., & Buneman, O. 1994, ApJ, 434, 363
- Sakai, J. I. 1990, ApJS, 73, 321
- Sakai, J. I. 1992, Sol. Phys., 140, 99
- Sakai, J. I., & de Jager, C. 1991, Sol. Phys., 134, 329
- Sakai, J. I., & de Jager, C. 1996, Space Sci. Rev., 77, 1
- Sakai, J. I., Nishi, K., & Sokolov, I. V. 2002, ApJ, 576, 519
- Sakai, J. I., & Ohsawa, Y. 1987, Space Sci. Rev., 46, 113
- Simnett, G. M. 1995, Space Sci. Rev., 73, 387
- Sokolov, I. V., Timofeev, E. V., Sakai, J. I., & Takayama, K. 1999, Shock waves, 9, 423
- Sokolov, I. V., Timofeev, E. V., Sakai, J. I., & Takayama, K. 2002, J. Computational Phys., 181, 354
- Tajima, T., Brunel, F., & Sakai, J. I. 1982, ApJ, 245, L45
- Vasylinuas, V. M. 1975, Geophys. Space Phys., 13, 303
- Zhang, M. H., Sokolov, I. V., & Sakai, J. I. 2001, Plasma Phys. Rep. 27(4), 303