

Spectral and timing properties of the X-ray emission from the millisecond pulsar PSR B1821–24

T. Mineo¹, G. Cusumano¹, E. Massaro^{2,3}, W. Becker⁴, and L. Nicastro¹

¹ Istituto di Astrofisica Spaziale e Fisica Cosmica, INAF-CNR, Sezione di Palermo, via Ugo La Malfa 153, 90146 Palermo, Italy
e-mail: mineo@pa.iasf.cnr.it

² Dipartimento di Fisica, Università La Sapienza, Piazzale A. Moro 2, 00185 Roma, Italy

³ Istituto di Astrofisica Spaziale e Fisica Cosmica, INAF-CNR, Sezione di Roma, via Fosso del Cavaliere, 00113 Roma, Italy

⁴ Max-Planck-Institut für Extraterrestrische Physik, 85740 Garching-bei-München, Germany

Received 4 February 2004 / Accepted 27 April 2004

Abstract. We report results on the timing and spectral analysis of observations of the millisecond pulsar PSR B1821–24 with RXTE, BeppoSAX and Chandra. The X-ray light curve is characterized by two narrow peaks at a phase distance of 0.452 ± 0.002 . The average pulsed emission, over the range 1.6–20 keV, is well represented by a single power law with a photon index $\alpha = 1.30_{-0.05}^{+0.02}$ and unabsorbed (2–10 keV) pulsed X-ray flux of $3.9 \times 10^{-13} \text{ erg cm}^{-2} \text{ s}^{-1}$. We searched for a possible bunching of X-ray photons to verify if the X ray emission has a time structure similar to that of giant pulses and found a negative result.

Key words. stars: neutron – stars: pulsars: general – stars: pulsars: individual: PSR B1821–24 – X-rays: stars

1. Introduction

PSR B1821–24 is an isolated millisecond radio pulsar discovered in the core of the globular cluster M 28 (Lyne et al. 1987). It is one of the most powerful (spin-down luminosity $2.2 \times 10^{36} \text{ erg s}^{-1}$) and youngest (spin-down age $3 \times 10^7 \text{ yr}$) millisecond pulsars (MSPs). This solitary MSP has a rotation period of 3.05 ms and a period derivative of $\dot{P} = 1.61 \times 10^{-18} \text{ s s}^{-1}$. The inferred magnitude of the dipolar field at the magnetic pole is $4.5 \times 10^9 \text{ G}$, the highest among the known MSPs. The distance of the source, assumed equal to the globular cluster one, ranges from $5.1 \pm 0.5 \text{ kpc}$ (Rees & Cudworth 1991) to 5.7 kpc (Harris 1996). In analogy with Becker et al. (2003) we use 5.5 kpc in our calculations.

Marginal detection of X-ray pulsed emission was reported by Danner, Kulkarni & Thorsett (1994) from a ROSAT (PSPC) observation from which they inferred a luminosity of $2.9 \times 10^{33} \text{ erg s}^{-1}$ in the energy band 0.1–2.4 keV. ASCA observation of PSR B1821–24 (Saito et al. 1997) evidenced a double peak profile with different shapes and a phase separation of 0.44 ± 0.01 .

Studies of the absolute phase in order to establish the phase lag of the X-ray emission with respect to the radio pulses have been performed with RXTE and Chandra observations (Rots et al. 1998; Rutledge et al. 2003). The absolute phase of the higher and narrower of the two X-ray peaks significantly

lags the corresponding radio feature, named P1 by Backer & Sallmen (1997), of $60 \pm 20 \mu\text{s}$.

The total pulsed spectrum was modelled with a single power law of spectral index 1.30 ± 0.07 for the ASCA/GIS (0.8–10 keV) data, 1.15 ± 0.02 for the RXTE/PCA data (2–20 keV) and 1.13 ± 0.02 for the joint ASCA-RXTE data (Kuiper et al. 2003).

Radio giant pulses have been observed from PSR B1821–24 with intensities that reach 80 times the average pulse energy. The brightest pulses are concentrated in a narrow phase window (~ 0.07), as derived from the average pulsed profile, that lags P1 by $80 \mu\text{s}$ and so it is coincident with the main X-ray peak (Romani & Johnston 2001).

Recently, the high spatial resolution of Chandra allowed us to resolve 12 sources within the core of M 28 where the pulsar is located (Becker et al. 2003; Rutledge et al. 2003). This made it possible to measure the unconfused phase averaged X-ray spectrum and the level of the unpulsed emission. The total (pulsed plus unpulsed) X-ray spectrum was modeled by a single power law with a photon index of $1.2_{-0.13}^{+0.15}$ and an unabsorbed (0.5–8 keV) flux of $3.5 \times 10^{-13} \text{ erg cm}^{-2} \text{ s}^{-1}$ with about 15% assigned to the non-pulsed emission (Rutledge et al. 2003). Moreover, the residuals in the Chandra spectrum show a hint of a feature at 3.3 keV, that Becker et al. (2003) fitted with a Gaussian. This line, having a width of 0.8 keV and a flux of $6 \times 10^{-6} \text{ photons cm}^{-2} \text{ s}^{-1}$, was identified as an electron

Table 1. Observation log.

Instrument(ObsId)	Starting time		Exposure (ks)
		(MJD)	
RXTE/PCA(P10421)	16 Sep. 1996	50 342.26	6.4
RXTE/PCA(P20159)	10 Feb. 1997	50 489.87	97.3
RXTE/PCA(P40090)	12 Nov. 1999	51 494.25	38.7
BSAX/MECS	17 Mar. 2000	51 620.57	96.9
Chandra/ACIS	4 Jul. 2002	52 459	12.7
Chandra/ACIS	8 Aug. 2002	52 494	13.5
Chandra/ACIS	9 Sep. 2002	52 526	11.4

cyclotron line formed in a magnetic field of 3×10^{11} G. The value, much higher than that derived from the dipole radiation, was explained either with a possible presence of multipolar components or with a strong displacement of the magnetic dipole center.

In this paper we present the timing and spectral analysis of one BeppoSAX and three RXTE observations of the pulsar covering the energy range 1.6–20 keV, including, for the spectral analysis, data from a Chandra/ACIS observation.

2. Observation and data reduction

RXTE observed the source three times from September 1996 to November 1999 as listed in Table 1. The two older observations (1996 and 1997) were performed with all five PCA units operating, while in the 1999 observation only three units were on. Standard selection criteria were applied to the data excluding time intervals corresponding to South Atlantic Anomaly passage, Earth’s limb lower than 10° and angular distance between the source and the satellite pointing direction larger than $0^\circ.02$. We used only data obtained with the PCA (Jahoda et al. 1996) accumulated in “Good Xenon” telemetry mode, time tagged with a $1 \mu\text{s}$ accuracy with respect to the spacecraft clock and with absolute time accuracy of $5\text{--}8 \mu\text{s}$ with respect to the UTC (Rots et al. 1998). Data from the first detector layer only are selected for the analysis to increase the signal-to-noise ratio.

BeppoSAX observed the source on 17 March 2000 for 196 ks elapsed time. In our analysis we consider only data from the MECS (1.6–10 keV; Boella et al. 1997) whose total exposure was 96.9 ks. MECS events were extracted from a circular region of $2'$ radius which maximize the signal-to-noise ratio of the pulsed component.

In addition the BSAX and RXTE data, we used the pulsar spectrum from Becker et al. (2003) which was taken with Chandra between July and September 2002. For observational and data analysis details see Becker et al. (2003).

The observing and exposure times of each observation are shown in Table 1.

Table 2. Radio ephemeris of PSR B1821–24 (Rutledge et al. 2003).

Parameter	Value
Right Ascension (J2000)	$18^{\text{h}} 24^{\text{m}} 32.0^{\text{s}}$
Declination (J2000)	$-24^\circ 52' 10''.7$
Validity epoch (MJD)	50 351–52 610
DM (pc cm^{-3})	119.873
ν (Hz)	327.40564101150(1)
$\dot{\nu}$ ($10^{-12} \text{ Hz s}^{-1}$)	$-0.1735080(1)$
$\ddot{\nu}$ ($10^{-24} \text{ Hz s}^{-2}$)	0.66
t_0 (MJD)	51 468.0

Table 3. Detected frequencies.

Instrument	Reference time		Frequency (Hz)
		(MJD)	
RXTE/PCA	50 342.31617688	327.4056574(8)	
RXTE/PCA	50 490.95766180	327.40565566(1)	
RXTE/PCA	51 494.73885753	327.40564068(5)	
BSAX/MECS*	51 621.70908612	327.4056429(4)	

*Quoted errors are the statistical ones. A further systematic error of -3.2×10^{-6} Hz must be considered for this observation (see text).

3. Timing analysis

Arrival times were converted to the Solar System Barycenter using the (J2000) pulsar coordinates given by Rutledge et al. (2003) (see Table 2) and the JPL2000 ephemeris (DE2000; Standish 1982). Search for the period was performed with the folding technique in a small frequency range around the value expected from the Nançay radio ephemeris reported in Table 2 (Rutledge et al. 2003).

The best frequency of each RXTE observation was computed fitting the peak of the χ^2 distribution with a Gaussian and it was found compatible with the expected value from radio measurements. Errors at 1σ level were derived by the frequency interval corresponding to a unit decrement with respect to the maximum in the χ^2 curve ($\text{err} = \nu(\chi_{\text{max}}^2) - \nu(\chi_{\text{max}}^2 - 1)$). Table 3 shows the detected frequency for the entire data set at the reference time assumed equal to the central time of each observation; the values in parenthesis are the statistical uncertainty in the last digit corresponding to one Gaussian standard deviation.

A cumulative RXTE light curve has been generated at the radio ephemeris adding the light curves of each observation. The highest significance of the pulsed emission is reached in the energy interval 2–20 keV (PHA 5–54) and the resulting X-ray spulse profile is shown in the top panel of Fig. 1. Two peaks are evident in this light curve: a sharp peak at phase $\Phi = 0.62$ and a wider and lower peak at phase $\Phi = 0.18$. In this paper, we use the nomenclature of Backer & Sallmen (1997)

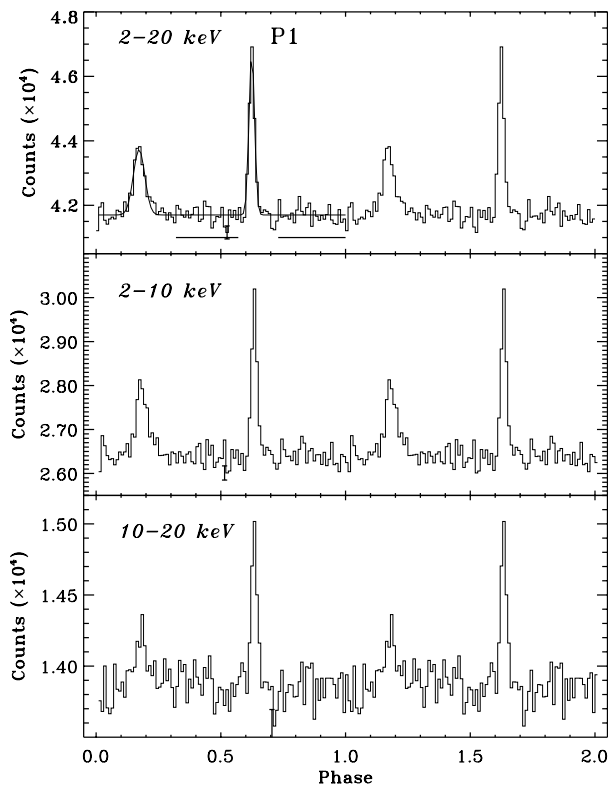


Fig. 1. PSR B1821–24 phase histogram obtained with RXTE/PCA: in the top panel the 2–20 keV light curve is shown with the two Gaussian fitting model. The two partial light curves 2–10 and 10–20 keV are shown in the middle and bottom panel, respectively. The off-pulse region used to extract the spectrum are indicated in the top panel.

referring the peak at phase 0.62 as $P1_X$ and the other as $P2_X$. The statistical significance is 19σ and 10σ for the first and second peak, respectively. The entire pulse profile can be modeled by two Gaussian curves, one for each peak, plus a constant level. The statistical agreement is good and the Gaussian parameters can be used to evaluate some interesting parameters: the phase separation between the two peaks is 0.452 ± 0.002 measured from $P2_X$ to $P1_X$; the best fit parameters for $P1_X$ and $P2_X$ for the FWHM are 0.0259 ± 0.0011 and 0.054 ± 0.004 for $P1_X$ and $P2_X$ respectively, and the normalization are $(2.03 \pm 0.13) \times 10^3$ and $(5.08 \pm 0.19) \times 10^3$, respectively. Two other pulse profiles in narrower energy ranges (2–10 keV and 10–20 keV) are shown in the other panels of Fig. 1. To search for possible morphology changes, we computed the $P2_X/P1_X$ ratios in different energy bands. This calculation was performed from the total counts in the phase intervals 0.60–0.65 and 0.13–0.24 for $P1_X$ and $P2_X$, respectively, after the subtraction of the off-pulse level, estimated in the phase intervals 0.32–0.57 and 0.73–1.00 (see Fig. 1). The $P2_X/P1_X$ values of the two ratios were 0.80 ± 0.06 and 0.65 ± 0.14 in the 2–10 keV and 10–20 keV band respectively, indicating only a marginal variation at 1.5σ significance. We verified these results from the integrated Gaussian profiles, with the best fit parameters given above, and found an agreement within one standard deviation.

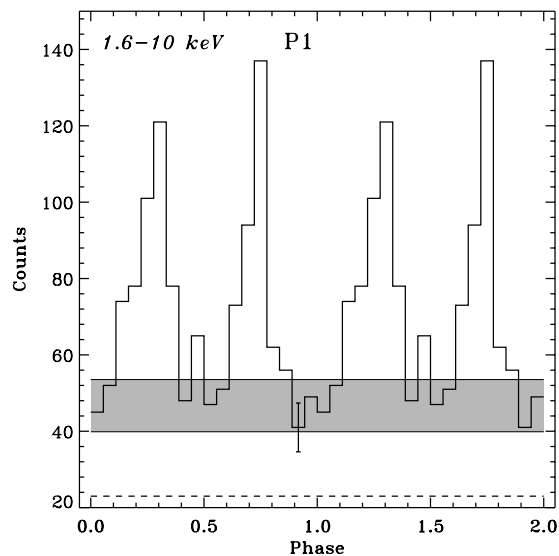


Fig. 2. PSR B1821–24 phase histogram obtained by the BeppoSAX/MECS in the energy range 1.6–10 keV. The grey band indicates the DC level and the dashed line the background.

BeppoSAX data were analyzed with the same method and the resulting χ^2 distribution of the MECS events showed a discrepancy between the measured frequency ν_X and the expected one from radio ephemeris ν_R , $\nu_R - \nu_X = -3.2 \times 10^{-6}$ Hz. This difference is due to a systematic of the onboard clock present in the BeppoSAX observations after January 2000 (Nicastro et al. 2002). For this reason, in the production of MECS light curves we used the proper ephemeris ν_X . The resulting 1.6–10 keV profile is shown in Fig. 2: two peaks with phase separation 0.44 ± 0.02 are evident. The profile is different from the RXTE one because the worse BeppoSAX timing accuracy ($\sim 200 \mu\text{s}$; Mineo et al. 2000) widens the two peaks. We arbitrarily associated the higher and narrower peak to $P1_X$ because the BeppoSAX accuracy does not allow the evaluation of the absolute phase.

The location of PSR B1821–24 near the core of the globular cluster makes difficult the evaluation of the pulsed fraction because of the contamination from these sources not resolved by the large MECS point spread function. Their contribution in the extraction region was computed using the results published in Becker et al. (2003) and subtracted from the light curve. The resulting signal is consistent with being 100% pulsed with a 2σ upper limit on the unpulsed emission of 20% in agreement with the Rutledge et al. (2003) results.

The occurrence of giant pulses (GPs) at the same phase of the X-ray peak $P1_X$ suggests the possibility that energetic photons are emitted during the same events in which the radio GPs are originated. To verify this hypothesis, we searched for a bunching of the X-ray photons, possibly with a rate similar to that of GPs. According to Romani & Johnston (2001) the GP rate from PSR B1821–24 is ~ 5 –6 events per hour, and therefore, if the X-ray flux is more intense in these occasions, one can expect to observe few pulses with a number of photons largely in excess to the poissonian fluctuations. During

the RXTE exposure we should expect about 200 possible X-ray flares correlated with giant pulses. In the search of these events, we applied the same method used in the analysis of the MSP PSR B1937+21 (Cusumano et al. 2003) and produced an X-ray light curve selecting only the events within a narrow phase interval of width $\Delta\Phi = 0.06$ ($180 \mu\text{s}$) centered at $P1_X$ and studied the frequency distribution of these events. This width has been chosen because it is comparable to the full width half maximum of $P1_X$ and of the giant pulse phase window. Since the dead time of the PCA is about $10 \mu\text{s}$, the maximum content of a bin in the presence of an X-ray flare cannot exceed 18–19 counts. The resulting distribution deviates from the Poisson law because we found a number of bins having more than 2 events much higher than expected. In particular, we obtained 1312 bins with 2 counts, 16 bins with 3 counts, 3 bins with 4 counts and 1 bin with 5 counts. However, comparing this distribution with others from several phase intervals in the off-pulse region, we found a similar number of bins with more than two events. Furthermore, comparable excesses were also found in the analysis of PSR B1937+21 (Cusumano et al. 2003), and therefore they seem related to instrumental timing performance rather than some other origin, like the variation of sources in the field of M 28 (Becker et al. 2003). We conclude that there is no evidence for a bunching of the X-ray emission.

4. Spectral analysis

Pulse phase histograms were accumulated for each unit of RXTE/PCA independently for the 256 PHA channels. The same procedure was applied to MECS data accumulating a phase resolved histogram for each energy channel.

RXTE spectra of the two peaks were accumulated in the phase intervals 0.60–0.65 and 0.13–0.24 for $P1_X$ and $P2_X$ respectively; off-pulse counts were extracted from the phase intervals 0.32–0.57 and 0.73–1.00 (see Fig. 1). The total pulsed spectrum was obtained from the sum of the two peaks. Energy channels were rebinned in order to have at least 30 counts per bin and make the χ^2 statistics available. RXTE spectra were combined by summing the individual units and assigning a total exposure time equal to the sum of the individual exposures. RXTE response matrix were derived for each PCA unit and summed together weighting with the background subtracted counts of the correspondent PCU's phase histogram.

A similar procedure was also followed for the MECS spectra. In this case however, because of the broad peak profiles (see Fig. 2), the off-pulse was estimated from the interval 0.40–0.60.

RXTE (energy range 2–20 keV) and MECS (energy range 1.6–10 keV) spectra were simultaneously fitted fixing the intercalibration factor to the value obtained modelling the two spectra in the common range 2–10 keV with a power law. The value 1.12 ± 0.05 is in agreement within two standard deviation with the values expected from the PCA-MECS intercalibration (Kuulkers et al. 2003).

Pulsed spectra were fitted with an absorbed power law with the galactic column density fixed to the values $N_H = 1.6 \times 10^{21} \text{ cm}^{-2}$ obtained by Chandra, in agreement with that computed from the reddening $E(B - V)$ of the globular cluster (Becker et al. 2003).

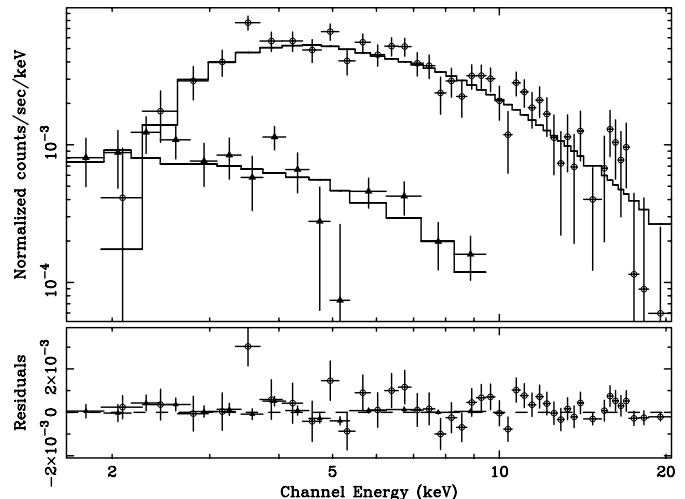


Fig. 3. PSR B1821–24 pulsed spectrum with the best fit power-law model (*top panel*) and the relative residuals (*bottom panel*). Triangles indicate the MECS points and stars indicate RXTE data.

The average pulsed spectral index was $\alpha = 1.30_{-0.12}^{+0.05}$ and the unabsorbed (2–10 keV) pulsed X-ray flux was $3.52 \times 10^{-13} \text{ erg cm}^{-2} \text{ s}^{-1}$ corresponding to a uniform luminosity of $L_X = 1.3 \times 10^{33} (d/5.5 \text{ kpc})^2 \text{ erg s}^{-1}$. This value agrees very well with the luminosity estimated by Saito et al. (1997) after the correction due to the small difference in the distance. The ratio of the observed (2–10 keV) luminosity with the spin-down energy loss is

$$\frac{L_X}{|\dot{E}_r|} = 5.8 \times 10^{-4} \frac{(d/5.5 \text{ kpc})^2}{I_{45}}$$

where I_{45} is the moment of inertia in 10^{45} gr cm^2 . It lies very nicely on the correlation between X-ray and spin-down luminosity given by Becker & Trümper (1999) and Possenti et al. (2002). As already noticed by Becker et al. (2003), the value of \dot{E}_r is essentially intrinsic because the line-of-sight projection of the acceleration in the cluster's gravitational potential can affect $\dot{\nu}$ by an amount less than 10%.

The pulsed spectra with the fitting model (*top panel*) and the relative residuals (*bottom panel*) are shown in Fig. 3. Only one bin around 3 keV shows a significant excess but it is too narrow to be compatible with the instrumental resolution and cannot be due to a true spectral feature. Moreover, adding a Gaussian line at 3.3 keV and having a width of 0.8 keV we obtained a 2σ upper limit of $6.8 \times 10^{-6} \text{ photons cm}^{-2} \text{ s}^{-1}$, which means that neither the RXTE nor the BSAX data are sensitive enough to confirm this spectral feature seen in the Chandra data. Furthermore, a black body spectral distribution did not fit the average pulsed spectra giving an unacceptable reduced $\chi_{\text{red}}^2 = 2.1$ (56 d.o.f.).

Fitting a power law spectrum to the two peaks separately, a marginal difference of the spectral indices has been detected with $\alpha = 1.17_{-0.13}^{+0.06}$ and $\alpha = 1.38_{-0.09}^{+0.18}$ for $P1_X$ and $P2_X$,

Table 4. Best fit parameters of the power-law model.

Component	α	$F(2-10 \text{ keV})^1$	$\chi^2_r(\text{d.o.f.})$
P1 _X	$1.17^{+0.06}_{-0.13}$	1.82 ± 0.24	0.93(56)
P2 _X	$1.38^{+0.18}_{-0.09}$	1.67 ± 0.21	1.22(56)
P1 _X +P2 _X	$1.30^{+0.11}_{-0.05}$	3.52 ± 0.33	1.07(56)
P1 _X +P2 _X +(DC) ²	$1.30^{+0.05}_{-0.02}$	3.52 ± 0.33	0.97(91)

¹ In units of $10^{-13} \text{ erg cm}^2 \text{ s}^{-1}$.

² MECS, RXTE and Chandra (total signal) data.

respectively. Table 3 reports the best fit parameters for the two peaks fitted independently and for the total pulsed emission (P1_X+P2_X).

To improve the fit accuracy, we performed the spectral analysis of the total signal from PSR B1821–24 considering also Chandra/ACIS data. This spectrum is relative to the total pulsed emission and includes an unpulsed component that, according Rutledge et al. (2003), can account for 15% of the total emission; presently there is no spectral information about this component. Taking into account these remarks, we fitted the pulsed RXTE and MECS spectra together with the Chandra total one using a single power law. The Chandra/ACIS intercalibration factor was determined with respect to RXTE by fitting the events in the common energy ranges and resulted equal to 0.91 ± 0.07 . The spectral index was found fully compatible with the previous analysis but with a much smaller uncertainty. The subtraction of the RXTE pulsed spectrum, extrapolated at energies below 2 keV, from the Chandra data did not give any significant low energy residuals, indicating that the DC component reported by Rutledge et al. (2003), if detectable, does not have a spectrum much softer than the pulsed one. An estimate of significant constraints on the spectral shape is practically not possible because of the large uncertainty on the intercalibration factors.

5. Discussion

Our analysis of the available X-ray observations of PSR B1821–24 performed with RXTE, BeppoSAX and Chandra confirms the known picture on this source. In particular, the joint analysis of the whole data set provided an accurate estimate of the phase averaged photon index which resulted $1.30^{+0.02}_{-0.05}$; moreover, RXTE and BeppoSAX gave an indication of a spectral difference between the two X-ray peaks. PSR B1821–24 and PSR B1937+21 are the only two MSPs whose radio emission shows giant pulses concentrated in the

same narrow phase window of X-ray pulses (Kinkhabwala & Thorsett 2000; Romani & Johnston 2001; Cusumano et al. 2003). It is unlikely that this phase coincidence would be casual and therefore it can be an indication that radio GPs and X-rays are emitted in the same region of the magnetosphere. For this reason we focused our attention to verify if the X-ray emission is occasionally enhanced with a rate comparable to that of GPs. A statistical test on the distribution of the number of X-ray events gave a negative result, indicating that, differently from GPs, X-rays photons are emitted in a steady process. Furthermore, the fact that the phase distances between normal radio and X-ray pulses are about the same in these two MSPs indicates that their respective origin regions do not have a large spatial separation.

It is possible that the high-energy emission is enhanced in coincidence with radio GPs. This effect has been found in the Crab pulsar whose optical pulses intensity simultaneous to GPs increases on average of 3% with respect to the normal radio pulses (Shearer et al. 2003). This result has not been verified at γ -ray energies though the present limit cannot exclude it (Lundgren et al. 1995). As pointed out by Shearer et al. (2003) an increase of the pair plasma density of the same order of the optical intensity cannot completely explain the very large flux of radio GPs, which in Crab are orders of magnitude stronger than normal pulses.

In the case of PSR B1821–24 we can derive a crude upper limit of the pulsed signal fraction emitted in correlation with the GPs. Assuming that $\sim 10\%$ of the total pulsed counts are concentrated within the GP phase window, we should expect several bins with 6 or more events in a single pulse; but they are not observed indicating that the fraction must be less than the above value. The same limit applies to the total number of emitting particles, whereas it does not hold for the local fluctuations of their density within a small region of the magnetosphere.

The actual site in the magnetosphere where pulses originate is uncertain. In some models the phase separations close to 0.5 has been considered as an indication that pulses originate from the polar regions of an orthogonal rotator (see e.g. Ruderman 2003). Electrons should then be accelerated in the polar gap and emit coherent and incoherent radiation from this site. Romani & Johnston (2001) proposed for PSR B1821–24 a different scenario in which GPs and X rays are emitted in the outer gaps, where a high local magnetic field should enhance the synchrotron emissivity and the production rate of secondary e^+e^- pairs. X-ray pulses are then emitted by this dense pair plasma that would also produce instabilities which in turn increase the particles coherence and hence the emission of radio GPs. Pulse profiles for the outer gap geometry have been calculated by Wang et al. (2002), who obtained narrow radio and X-ray peaks with a phase separation close to 0.5. Whether these models are appropriate to describe the observed scenario can be decided only when more pulsars showing the same behaviour are detected.

Acknowledgements. This work has been partially supported by INAF (Istituto Nazionale di Astrofisica).

References

- Backer, D. C., & Sallmen, S. 1997, *AJ*, 114, 1539
- Becker, W., & Trümper, J. 1999, *A&A*, 341, 803
- Becker, W., Swartz, D. A., Pavlov, G. G., et al. 2003, *ApJ*, 594, 798
- Boella, G., Chiappetti, L., Conti, G., et al. 1997, *A&AS*, 122, 327
- Cusumano, G., Massaro, E., Mineo, T., et al. 2003, *A&A*, 410, L9
- Danner, R., Kulkarni, S. R., & Thorsett, S. E. 1994, *ApJ*, 436, L153
- Harris, W. E. 1996, *AJ*, 112, 1487
- Jahoda, K., Swank, J. H., Giles, A. B., et al. 1996, *SPIE*, 2808, 59
- Kinkhabwala, A., & Thorsett, S. E. 2000, *ApJ*, 535, 365
- Kuiper, L., Hermsen, W., & Stappers, B. 2003, Proc. of the International Workshop held in Marsala, Pulsars, AXPs and SGRs observed with BeppoSAX and Other Observatories, September 23–25, 2002, ed. E. Cusumano Massaro, & T. Mineo
- Kuulkers, E., den Hartog, P. R., in't Zand, J. J. M., et al. 2003, *A&A*, 399, 663
- Lundgren, S. C., Brinklow, A., Middleditch, J., Kulkarni, S. R., & Backer, D. C. 1995, *ApJ*, 470, 469
- Lyne, A. G., Brinklow, A., Middleditch, J., et al. 1987, *Nature*, 328, 399
- Nicastro, L., Cusumano, G., Kuiper, L., et al. 2002, Proc. of the Seminar on Neutron Stars, Pulsars and Supernova Remnants, Physikzentrum Bad Honnef, ed. W. Becker, H. Lesch, & J. Trümper
- Possenti, A., Cerutti, R., Colpi, M., & Mereghetti, S. 2002, *A&A*, 387, 993
- Mineo, T., Cusumano, G., Kuiper, L., et al. 2000, *A&A*, 355, 1053
- Rees, R. F., & Cudworth, K. M. 1991, *AJ*, 102, 152
- Romani, R. W., & Johnston, S. 2001, *ApJ*, 557, L93
- Rots, A. H., Jahoda, K., Macomb, D. J., et al. 1998, *ApJ*, 501, 749
- Ruderman, M. 2003, X-ray and γ -ray Astrophysics of Galactic Sources, in press [[arXiv:astro-ph/0311228](https://arxiv.org/abs/astro-ph/0311228)]
- Rutledge, R. E., Fox, D. W., Kulkarni, S. R., et al. 2003 [[arXiv:astro-ph/0301453](https://arxiv.org/abs/astro-ph/0301453)]
- Saito, Y., Kawai, N., Shibata, S., Dotani, T., & Kulkarni, S. R. 1997, *ApJ*, 477, L37
- Shearer, A., Stappers, B., O'Connor, P., et al. 2003, *Science*, 301, 493
- Standish, E. M. 1982, *A&A*, 114, 297