A&A 412, 399–403 (2003)
DOI: 10.1051/0004-6361:20034225

Astronomy & Astrophysics

A catalogue of quasars and active nuclei: 11th edition⋆

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Received 25 August 2003 / Accepted 15 September 2003

Abstract. The recent release of the final installement of the 2dF quasar catalogue and of the first part of the Sloan catalogue, almost doubling the number of known QSOs, led us to prepare an updated version of our *Catalogue of quasars and active nuclei* which now contains 48 921 quasars, 876 BL Lac objects and 15 069 active galaxies (including 11 777 Seyfert 1s). Like the tenth edition, it includes position and redshift as well as photometry (U, B, V) and 6 and 11 cm flux densities when available. We also give a list of all known lensed and double quasars.

Key words. galaxies: quasars: general – galaxies: BL Lacertae objects: general – galaxies: active

1. Introduction

The first catalogue of quasars was published in 1971 by De Veny et al. It contained 202 objects. The number of known quasars has since steadily increased until the year 2000. But the release of the first part of the "2dF QSO redshift survey" (Croom et al. 2001) almost doubled this number (see Table 1), The recent release of both the final installement of the 2dF catalogue (Croom et al. 2003) and of the first part (Abazajian et al. 2003) of the "Sloan Digital Sky Survey" (Fan et al. 1999) has again more than doubled the number of known QSOs justifying the present edition.

In this edition, containing quasars with measured redshift known to us prior to August 1st, 2003, as in the preceding editions, we do not give any information about absorption lines or X-ray properties. But we give the absolute magnitude for each object and, when available, the 11 and 6 cm flux densities.

This catalogue should not be used for any statistical analysis as it is not complete in any sense, except that it is, we hope, a complete survey of the literature.

2. Description of the catalogue

The quasars are listed in Table_QSO. A sample page is shown in Fig. 1. We have arbitrarily defined a quasar as a starlike object, or an object with a starlike nucleus, with broad emission lines, brighter than absolute magnitude $M_B = -23$.

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⋆ The catalogue (Table_QSO, Table_BL, Table_AGN and Table_reject) and the list of references are only available in electronic form at the CDS via anomymous ftp to cdsarc.u-strasbg.fr (130.79.128.5) or via http://cdsweb.u-strasbg.fr/cgi-bin/qcat?J/A+A/412/399 or at the Observatoire de Haute Provence (http://www.obs-hp.fr/).

Table 1. Increase with time of the number of known QSOs, BL Lacs and Seyfert 1s.

QSO	BL Lac	Seyfert 1	reference
202			De Veny et al. (1971)
2251		190	Véron-Cetty & Véron (1984)
2835	73	236	Véron-Cetty & Véron (1985)
3473	84	258	Véron-Cetty & Véron (1987)
4169	117	358	Véron-Cetty & Véron (1989)
6225	162	575	Véron-Cetty & Véron (1991)
7383	171	695	Véron-Cetty & Véron (1993)
8609	220	888	Véron-Cetty & Véron (1996a)
11 358	357	1111	Véron-Cetty & Véron (1998)
13 214	462	1711	Véron-Cetty & Véron (2000a)
23 760	608	2765	Véron-Cetty & Véron (2001)
48 921	876	11 777	Present edition

In Table_BL, we list all confirmed, probable or possible BL Lac objects with or without a measured redshift, without consideration of their absolute magnitude. As better spectra are becoming available, broad emission lines have been detected in a number of objects formerly classified as BL Lac; they have usually been moved to Table_QSO (Véron-Cetty & Véron 2000b).

Table_AGN lists "active galaxies": Seyfert 1s, Seyfert 2s and Liners fainter than $M_B = -23$. A number of galaxies with a nuclear H II region are also included, the reason being that they have been called Seyfert in the past and later reclassified; we consider it useful to keep track of these reclassifications to avoid further confusion. Seyfert 1s have broad Balmer and other permitted lines; Seyfert 2s have Balmer and forbidden lines of the same width. Osterbrock (1977, 1981) has

Name	Alpha (J2000)	Delta (J2000)		S 6		S11		Z		V	B-V	U-B	Mabs	References			Alpha (B1950)	Delta (B1950)
FIRST J00000-0202	0 0 1.3	- 2 2 0	R		165			1.356		19.64			-25.3			165	23 57 27.5	- 2 18 42
*2QZ J000001-3036	0 0 1.4	-30 36 27	O					1.143		*20.10		-0.80	-24.4		489	489	23 57 27.4	-30 53 9
*2QZ J000001-3122	0 0 1.7	-31 22 26	O					1.331		*20.69		-0.99	-24.2		489	489	23 57 27.7	-31 39 8
*XMM J00000-2511	0 0 2.7	-25 11 37	O					1.314	S1	R21.			-23.4			674	23 57 28.8	-25 28 19
*MS 23574-3520	0 0 2.8	-35 3 33	O		1960			0.508		O17.0			-25.5	1235		1960	23 57 28.8	-35 20 15
*2QZ J000005-2725	0 0 5.6	-27 25 10	O					1.930		*19.43		-1.07	-26.4		487	487	23 57 31.6	-27 41 52
*PSS J0000+2357	0 0 9.4	23 57 16						4.030		R18.93			-28.6			557	23 57 35.8	23 40 34
*SDSS J00001-1027	0 0 9.4	-10 27 52	O		1			1.845		19.18			-26.5			1	23 57 35.6	-10 44 34
*2QZ J000009-3116	0 0 9.7	-31 16 48	O					1.727		*19.05		-0.65	-26.5		489	489	23 57 35.7	-31 33 30
*2QZ J000009-3055	0 0 9.9	-30 55 30	O					1.787		*19.12		-1.19	-26.5		487	487	23 57 35.9	-31 12 12
GB6 23576+3039	0 0 10.1	30 56 0	R	0.049	778			1.801		I19.3			-25.4			1427	23 57 36.5	30 39 18
*2QZ J000010-3159	0 0 10.2	-31 59 50	O					1.638		*20.44		-0.16	-24.9		487	487	23 57 36.2	-32 16 32
*2QZ J000011-3138	0 0 11.7	-31 38 40	O					2.680		*20.27		-0.61	-26.7		489	489	23 57 37.7	-31 55 22
*PB 5669	0 0 12.0	0 2 24	O					0.479		O18.1			-24.3			892	23 57 38.2	- 0 14 18
*Q 2357-024	0 0 12.9	- 2 10 25						*1.45		18.5			-26.7	2359		2359	23 57 39.1	- 2 27 7
*2QZ J000015-2738	0 0 15.9	-27 38 56	O					1.516		*18.38		-0.46	-26.8		489	489	23 57 42.0	-27 55 38
*2QZ J000016-3144	0 0 16.3	-31 44 38	O					1.452		*19.08		-1.08	-26.1		487	487	23 57 42.3	-32 1 20
FIRST J00002-0851	0 0 17.5	- 8 51 23	O		165			1.250		18.93			-25.8			165	23 57 43.7	- 9 8 5
PKS 2357-326	0 0 20.2	-32 21 1	O	0.32	1835	0.54	1835	1.275		18.7			-25.8	1564	1261	1751	23 57 46.3	-32 37 43
*Q 2357-027	0 0 22.9	- 2 27 15						*0.59		18.0			-24.8	2359		2359	23 57 49.1	- 2 43 57
*XMM J00003-2512	0 0 22.9	-25 12 22	O					1.610	S1	R19.7			-25.2			674	23 57 49.0	-25 29 4

Fig. 1. Sample page of the QSO catalogue.

divided the Seyfert 1s into five subgroups: Seyfert 1.0, 1.2, 1.5, 1.8 and 1.9 on the basis of the appearance of the Balmer lines. Seyfert 1.0s are "typical" members of the class, as described by Khachikian & Weedman (1971, 1974), while Seyfert 1.5s are objects intermediate between typical Seyfert 1s and Seyfert 2s, with an easily apparent narrow Hβ profile superimposed on broad wings. The classes Seyfert 1.2 and 1.8 are used to describe objects with relatively weaker and stronger narrow Hβ components, intermediate between Seyfert 1.0 and 1.5 and Seyfert 1.5 and 2 respectively. In Seyfert 1.9, although the broad Hα emission is clearly evident, broad Hβ cannot be detected with certainty by mere visual inspection of the spectra. We have adopted the more quantitative classification introduced by Winkler (1992):

S1.0	5.0	$< R$	
S1.2	2.0	$< R < 5.0$	
S1.5	0.333	$< R < 2.0$	
S1.8		$R < 0.333$	broad component visible in Hα and Hβ
S1.9			broad component visible in Hα but not in Hβ
S2			no broad component visible

where R is the ratio of the total Hβ to the [OIII]λ5007 fluxes. Several objects have been found to show extreme spectral variability, changing from Seyfert 1.8 or 1.9 to Seyfert 1.0. In some cases these changes are consistent with changes in the reddening to the BLR while, in others, they are probably due to real changes in ionizing flux (Goodrich 1989a, 1995; Tran et al. 1992b). In some Seyfert 2s, a broad Pa β line has been detected, indicating the presence of a highly reddened broad line region (Goodrich et al. 1994); we call these objects S1i. A number of Seyfert 2s have, in polarized light, the spectra of Seyfert 1s (Antonucci & Miller 1985; Miller & Goodrich 1990; Tran et al. 1992a); we call them S1h. Typical full widths at half-maximum of the Balmer lines in Seyfert 1s lie in the range 2000–6000 km s^{-1}; however, there is a group of active galactic nuclei with all the properties of Seyfert 1s, but with unusually narrow Balmer lines (Osterbrock & Pogge 1985; Goodrich 1989b); they are defined as having the broad component of the Balmer lines narrower than 2000 km s^{-1} FWHM (Osterbrock 1987); we call them S1n. Liners (as defined by Heckman 1980) are called S3. If broad Balmer lines are observed, they are called S3b; if these broad Balmer lines are only seen in polarized light, they are called S3h. Only objects brighter than $M_B = -23$ appear in Table_QSO, but, clearly, some objects would move from Table_QSO to Table_AGN and vice versa if other values for q_0 and the spectral index were used or if an accurate B apparent magnitude was available for all objects. The variability may have a similar effect, as well as the size of the diaphragm used for the measurement as the contribution of the underlying galaxy for weak quasars may not be negligible.

Table_reject lists the objects which once were believed to be quasars or BL lac objects and are now known to be either stars or normal galaxies.

Table_QSO contains 48 921 objects, Table_BL, 876, Table_AGN, 15 069 and Table_reject, 76. The catalogue is believed to contain all known quasars, BL Lac objects and Seyfert 1s.

Table_QSO, Table_BL and Table_AGN give:

1) Columns 1 and 2. The most common name of the object. For the meaning and the sources of the designations see Hewitt & Burbidge (1987), Fernandez et al. (1983) and Kesteven & Bridle (1977). For the sources discovered by the ROSAT X-ray satellite, we have used the following acronyms: RXS for the sources appearing in the All-Sky Bright Source Catalogue (Voges et al. 1999), RX for the sources in the Faint Source Catalogue (Voges et al. 2000) and 1WGA for the sources published in the WGACAT catalogue (White et al. 1994).

When the name is preceded by an *, the object has not been explicitely associated with a radio source. In fact many of these objects may have been detected by large deep radio surveys such as the FIRST (Becker et al. 1995; White et al. 1997) and NVSS (Condon et al. 1998) surveys.

2) Columns 3 to 10. The best available J2000 optical or radio coordinates. The J2000 positions have been converted from the B1950 positions using the matrix given by Aoki et al. (1983). An O or an R following the coordinates

means that the position is either an optical or a radio position measured with an accuracy better than one arcsec. An *A* means that it is only an approximate position which may be wrong by several arcminutes. No reference is given for the source of the positions. The availability of the Digitized Sky Survey (DSS) allows quick measurements of the optical position of any object brighter than ≈19.5 mag. It has already been used to measure the position of several hundreds QSOs (Schneider et al. 1992; Bowen et al. 1994; Kirhakos et al. 1994; Véron-Cetty & Véron 1996b). Optical positions with an accuracy better than 2″ have also been measured for the 19 369 galaxies in the Zwicky catalogue (Falco et al. 1999) and for the 12 921 UGC galaxies (Cotton et al. 1999).

3) Columns 11 to 14. The 6 and 11 cm flux densities (in Jy) with references to the literature. When several measurements are available we took arbitrarily one of them. When a reference is given for the 6 cm flux density but the value of the flux density itself is left blank and there is an * in Col. 1, only an upper limit is available and this upper limit is not much greater than 1 mJy; in case there is no * in Col. 1, the reference refers to a detection but at a wavelength other than 6 cm.

4) Columns 15 and 16. The redshift as published. An * in front of the redshift means that it has been estimated from a low dispersion slitless spectrum and is of lesser accuracy or even plainly wrong as the emission lines may easily be misidentified. We have given only those values which are described as probable in the original sources and not the possible values.

5) Column 17. In this column an attempt has been made to classify the objects as S1, S1.0, S1.2, S1.5, S1.8, S1.9, S1i, S1h, S1n, S2, S3, S3b, S3h, S, S? or H2. Low redshift quasars are classified as S1 when a good spectrum shows that they are similar to Seyfert 1 galaxies.

In Table_BL, we find in this column:

BL	for a confirmed BL Lac object.
BL?	for a probable BL Lac.
blank	for a possible BL Lac.
?	for a questionable BL Lac
HP	for a Highly Polarized object.

6) Columns 18 to 21. The *V*, $B-V$ and $U-B$ photoelectric or photographic magnitude and colours, when available (the survey of the literature for photographic colours may be incomplete) (an * in front of the magnitude indicates that the colours and the magnitude are photographic, while an *R* or an *I* indicates a red or an infrared magnitude). The column labelled "*V*" gives the *V* magnitude when $B-V$ is also given. When $B-V$ is not given, this column usually gives the *B* magnitude, unless it is preceded by an *R* or an *I*. Maoz et al. (1993) have measured homogeneous *V* magnitudes for 354 QSOs with an accuracy of ±0.1 mag; they have been included. For a few objects, the *O* magnitude, measured on the blue Palomar Sky Survey plates, or the UK Science Research Council SRC-J Survey plates, believed to be accurate within ±0.2 mag, has been extracted from the APS database (Pennington et al. 1993). For a number of objects we give the *O* magnitude, extracted from the USNO-A2 catalogue (Monet et al. 1996) or the Cambridge Automated Plate Measuring Machine (APM) catalogue (Irwin et al. 1994), recalibrated by E. Flesch (private communication); these magnitudes are flagged with an *O*. The *O* and Johnson *B* magnitudes are related by $B-O = -(0.27 \pm 0.06) \times (B-V)$ (Evans 1989).

In the other cases, the magnitude given is an estimate as found in the original publications. These magnitudes are generally quite inaccurate and inhomogeneous; they are most often m_{pg} or *B* magnitudes instead of the Johnson *V* magnitude. Much care should be taken when using them for any purpose. Anyway, even when a photoelectric *V* magnitude is given, it is not very meaningful as most quasars are variable. On the other hand, the colours of quasars vary little, so the listed colours should be accurate. Again, it should be noted that some of the colours listed are photographic and, therefore, less accurate; moreover, in each catalogue of photoelectric measurements, the faintest objects measured are affected by relatively large errors; this too should not be overlooked. For the galaxies, in Table_AGN, we have chosen the magnitudes and colours measured in the smallest possible diaphragm (preferentially 16 arcsec) as we are interested in the nucleus rather than in the galaxy itself.

7) Column 22. The absolute magnitude M_B computed assuming $H_0 = 50$ km s^{-1} Mpc^{-1}, $q_0 = 0$, and an optical spectral index α equal to 0.3 (defined as $S \propto \nu^{-\alpha}$) (Francis et al. 1991), as follows:

$$M = m + 5 - 5 \times \log D - k + \Delta m(z)$$

where *m* is the *B* magnitude, $D = c/H_0 \times A$, with *A* the photometric distance (Terrell 1977):

$$A = z\left[1 + \frac{z(1-q_0)}{(1+2q_0 z)^{0.5} + 1 + q_0 z}\right]$$

z is the redshift; $k = -2.5 \log(1+z)^{1-\alpha}$ is the *k* correction, $\Delta m(z)$ is a correction to *k* taking into account the fact that the spectrum of quasars is not strictly a power law of the form $S \propto \nu^{-\alpha}$, but is affected by emission lines and by the Ly α forest depleting the continuum to the blue of Ly α. Assuming that the spectrum is a power law with $\alpha = 0.3$ may not give the best possible estimate of the *k* correction (Wisotzki 2000). The *R* magnitudes have been transformed into the *B* system by using an average $< B-R > = 0.57$ and the *I* magnitudes by using $< B-I > = 1.1$ for low z QSOs. When the reference for the magnitude is Maoz et al. (1993), the magnitude is *V* and we have used $< B-V > = 0.40$.

8) The next three columns (23 to 25) give the reference for the finding chart, the photometry and the redshift respectively. In many cases, the last reference in Table_AGN is that of the classification of the object (as a Seyfert or otherwise); in these cases the redshift can usually be found in Palumbo et al. (1983).

9) The B1950 position (Cols. 26 to 32).

Since the discovery in 1979 by Walsh et al. of the first gravitationally lensed quasar, Q 0957+561, a number of such objects (52) and of physical pairs with separation less than 10″ (14) have been found. They are listed in Tables 2 and 3 respectively. Mortlock et al. (1999) have stressed the difficulty sometimes encountered in distinguishing lensed quasars from physical pairs.

Table 2. Gravitationally lensed quasars. Column 1: name, Col. 2: short 1950 position, Col. 3: redshift of the quasar, Col. 4: redshift of the lens, Col. 5: separation in arcsec, Col. 6: references (see Table 3).

Name	Position	z_{quasar}	z_{lens}	sep($''$)	Ref.
PKS 0132-097	0132−09	2.216	0.764	0.7	13, 18, 68
UM 673	0142−10	2.719		2.2	60
CTQ 414	0156−43	1.29		1.2	46
B2 0218+35	0218+35	0.936		0.33	15
HE 0230−2130	0230−21	2.162		2.0	72
Q J0240−343	0238−34	1.406		6.1	62
PKS 0411+05	0411+05	2.639	0.958	2.2	34, 64
HE 0435−1223	0435−12	1.689		2.6	73
HE 0512−3329	0512−33	1.565	0.931	0.6	12
B 0712+472	0712+47	1.339		1.27	8
MG 0751+2716	0748+27	3.200	0.350		64
HS 0810+25	0810+25	1.500		0.25	55
HS 0818+1227	0818+12	3.115		2.1	16
CLASS B0827+525	0827+52	2.064		2.8	31
APM 08279+5255	0827+52	3.87		0.4	35
SDSS 09035+5028	0900+50	3.584	0.388	2.8	28
RX J0911.4+0551	0908+06	2.800		0.8	1
SBS 0909+532	0909+53	1.377	0.830	1.11	30, 39
1WGA J09212+4528	0917+45	1.66	0.31	6.93	51
SDSSp J09249+0219	0922+02	1.524		1.8	26
FBQS J0951+2635	0948+26	1.24		1.1	56
BRI 0952−01	0952−01	4.43		0.95	43
Q 0957+561	0957+56	1.414	0.355	6.1	11, 65
FIRST J10044+1229	1001+12	2.65		1.54	32
Q 1009−0252	1009−02	2.74		1.55	21
J 13.03	1015−20	2.55		0.84	61
IRAS F10214+4724	1021+47	2.286			57
B 1030+074	1030+07	1.535		1.56	8
HE 1104−1805	1104−18	2.303	0.729	3.0	37, 70
PG 1115+080	1115+08	1.722	0.311	2.3	63, 66
UM 425	1120+01	1.465		6.5	44
1RXS J11319−1231	1129−12	0.658	0.295	4.2	58
TEX 1152+199	1152+19	1.019	0.439	1.6	52
Q 1208+1011	1208+10	3.803		0.45	41
87GB 1359+1527	1359+15	3.235		1.7	52
H 1413+117	1413+11	2.546		1.4	40
HST J14176+5226	1415+52	3.4		3.2	5
B 1422+231	1422+23	3.62	0.339	1.3	53, 63
SBS 1520+530	1520+53	1.855	0.717	1.6	3, 4
Q 1600+434	1600+43	1.61		1.38	27
FIRST J1633+3134	1631+31	1.516		0.66	48
PMN J1632-0033	1630−00	3.424		1.46	69
Q 1634.9+26.7	1634+26	1.961		3.8	59
SDSSp J16507+4251	1649+42	1.541		1.16	49
MC 1830−211	1830−21	2.507	0.885	0.60	36, 38
TEX 1835−345	1835−34	2.78		1.0	67
MG 2019+1127	2016+11	3.273		3.4	33
87GB 20451+2632	2045+26	1.28	0.867	1.9	9
Q 2138−431	2138−43	1.641		4.5	19
HE 2149-2745	2149−27	2.033		1.7	71
Q 2237+0305	2237+03	1.695	0.039	1.8	24
Q 2345+007	2345+00	2.15		7.1	54

Acknowledgements. This research has made use of the APS catalogue of POSS I database which is supported by the National Science Foundation, the National Aeronautics and Space Administration, and the University of Minnesota. We are very grateful to E. Flesch

Table 3. Quasar pairs. Column 1: name, Col. 2: short 1950 position, Col. 3: redshift of the quasar, Col. 4: separation in arcsec, Col. 5: references.

Name	Position	z	sep($''$)	Ref.
LBQS 0015+02	0015+02	2.469	2.2	25
Q 0023+171	0023+17	0.945	4.8	23
CT 344	0103−27	0.848	0.3	29
PHL 1222	0151+04	1.910	3.3	45
CTQ 839	0250−33	2.24	2.1	47
Q 1145−071	1145−07	1.34	4.2	7
HS 1216+5032	1216+50	1.450	8.9	17
RRS IV 26,27	1343+26	2.030	9.5	6
Q 1429−008	1429−00	2.076	5.1	10, 20
RX J16290+3724	1627+37	0.923	4.3	42
Q J1643+31	1641+32	0.586	2.3	2
Q 2153−2056	2153−20	1.845	7.8	22
MGC 2214+3550	2212+35	0.877	3.0	50
SDSSp J23365−0107	2334−01	1.285	1.7	14

References: (1) Bade et al. (1997); (2) Brotherton et al. (1999); (3) Burud et al. (2002); (4) Chavushyan et al. (1997); (5) Crampton et al. (1996); (6) Crotts et al. (1994); (7) Djorgovski et al. (1987); (8) Fassnacht & Cohen (1998); (9) Fassnacht et al. (1999); (10) Faure et al. (2003); (11) Garrett et al. (1992); (12) Gregg et al. (2000); (13) Gregg et al. (2001); (14) Gregg et al. (2002); (15) Grundahl & Hjorth (1995); (16) Hagen & Reimers (2000); (17) Hagen et al. (1996); (18) Hall et al. (2002); (19) Hawkins et al. (1997); (20) Hewett et al. (1989); (21) Hewett et al. (1994); (22) Hewett et al. (1998); (23) Hewitt et al. (1987); (24) Huchra et al. (1985); (25) Impey et al. (2002); (26) Inada et al. (2003); (27) Jackson et al. (1995); (28) Johnston et al. (2003). (29) Junkkarinen et al. (2001); (30) Kochanek et al. (1997); (31) Koopmans et al. (2000); (32) Lacy et al. (2002); (33) Lawrence et al. (1984); (34) Lawrence et al. (1995); (35) Ledoux et al. (1998); (36) Lidman et al. (1999); (37) Lidman et al. (2000); (38) Lowell et al. (1998); (39) Lubin et al. (2000); (40) Magain et al. (1988); (41) Magain et al. (1992); (42) Mason et al. (2000); (43) McMahon et al. (1992); (44) Meylan & Djorgovski (1989); (45) Meylan et al. (1990); (46) Morgan et al. (1999); (47) Morgan et al. (2000); (48) Morgan et al. (2001); (49) Morgan et al. (2003); (50) Muñoz et al. (1998); (51) Muñoz et al. (2001); (52) Myers et al. (1999); (53) Patnaïk et al. (1992); (54) Pello et al. (1996); (55) Reimers et al. (2002); (56) Schechter et al. (1998); (57) Serjeant et al. (1995); (58) Sluse et al. (2003); (59) Steidel & Sargent (1991); (60) Surdej et al. (1987); (61) Surdej et al. (1997); (62) Tinney (1995); (63) Tonry (1998); (64) Tonry & Kochanek (1999); (65) Tytler & Fan (1992); (66) Weymann et al. (1980); (67) Winn et al. (2000); (68) Winn et al. (2001); (69) Winn et al. (2002); (70) Wisotzki et al. (1993); (71) Wisotzki et al. (1996); (72) Wisotzki et al. (1999); (73) Wisotzki et al. (2002).

and F. Ochsenbein for checking and improving the catalogue and we thank R. Monella for having brought to our attention a number of errors and omissions in previous editions of this catalogue.

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