The discovery and study of the optical counterparts of the transient X-ray pulsars RX J0052.1–7319 and XTE J0111.2–7317 in the SMC

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Received 3 April 2001 / Accepted 23 May 2001

Abstract. We report on the discovery and confirmation of the optical counterparts of the two transient X-ray pulsars, RX J0052.1–7319 and XTE J0111.2–7317. In the narrow (~3′′ radius) X-ray error circle of RX J0052.1–7319 we found a V = 14.6 (B − V = 0.1) 09.5IIIe (a classification as a B0Ve star is also possible, since the luminosity class depends on the uncertainty on the adopted reddening). Medium resolution spectra for this object show Balmer lines in emission with an equivalent width of H\textsubscript{\alpha} = 12−16 Å. In the X-ray error box of XTE J0111.2–7317 we found a relatively bright object (V = 15.4, B − V = 0.06) which has been classified as a B0.5−B1Ve star and that was later confirmed by Coe et al. (2000) as the most plausible counterpart for XTE J0111.2–7317. Also in this case we easily detect Balmer emission lines with an H\textsubscript{\alpha} equivalent width of about −21 Å. There is also evidence for the presence of a surrounding nebula, possibly a supernova remnant. A further bright B0Ve star was found just outside the X-ray error circle of XTE J0111.2–7317. We discuss the implication of these results in the light of the current knowledge of Be/X-ray binary systems in the Magellanic Clouds and within our Galaxy.

Key words. pulsars: individual (RX J0052.1–7319, XTE J0111.2–7317) – stars: neutron – X-ray: stars

1. Introduction

X-ray pulsars in High Mass X-ray Binaries (HMXBs) consist of an accreting neutron star orbiting a (super)giant or a main-sequence Be-type companion star. The greater part of known neutron stars in Supergiant XBs (from now on SXBs) show permanently high X-ray fluxes driven by accretion of a roughly spherical dense wind from the massive companion (which may be enhanced by Roche lobe overflow). In contrast, the neutron stars in BeXBs often exhibit transient X-ray outbursts which may occur periodically at periastron (type I) or when the companion star undergoes a mass loss episode from the equatorial regions due to its high rotational velocity (up to ~75% of the break-up velocity; type II). In BeXBs, the primary star is an early type star in the range 10–20 M\odot of luminosity class III to V, which often displays Balmer lines, but also He I and metallic lines, in emission. Due to these lines, Be stars are difficult to classify. Moreover, since BeXBs are generally first detected as X-ray sources with positional accuracies which can be as poor as a few arcminutes, the optical/IR follow-up observations and the detection of the possible counterpart is often a difficult task. In addition, the BeXB concentration on the Galactic Plane and in the Magellanic Clouds (MCs) results in a further decrease in the chances of detecting them because of crowding and source confusion. So far only about ~20 optical counterparts of BeXBs have been discovered out of the >100 known and 10\textsuperscript{4}−10\textsuperscript{5} expected Be/X-ray pulsars (Nelson et al. 1993 and 1995). Detailed discussions on the nature and properties of these systems have been recently reported by several authors (Negueruela 1998; Coe 2000; Negueruela & Okazaki 2000).

In this paper we report on the results of optical observations aimed at searching for the counterpart of two transient X-ray pulsars: RX J0052.1–7319 and XTE J0111.2–7317. Both systems are hosted by the Small Magellanic...
Cloud (SMC) which harbors a large number of these systems (Yokogawa et al. 2000a; Mereghetti 2001).

The paper is organized as follows: in Sect. 2 we describe the techniques we have applied for the observations and the data reduction and in Sect. 3 we discuss the result of the observations and derive physical parameters for the proposed optical counterparts of RX J0052.1−7319 and XTE J0111.2−7317. The detailed spectral classification of the observed stars are also provided.

2. Observations

The search for the optical counterparts in the fields of the X-ray sources RX J0052.1−7319 and XTE J0111.2−7317 was carried out by means of several techniques and instruments. Photometry was performed in order to study the stellar field main properties and to select objects with colors and magnitudes expected for the counterparts of a BeXB in the SMC. Spectroscopy of the optical counterpart candidates was carried out to infer the physical parameters (i.e. spectral classification) and to look for signatures of activity to be correlated with the X-ray emission.

Multicolor photometry (B, V, and R bands) was carried out for each field with the 1.5 m Danish telescope of the ESO at La Silla (Chile). Instrumental magnitudes for each stellar object in the images were derived with aperture and profile-fitting photometry by means of the DAOPHOTII package (Stetson 1987). Low and medium resolution spectroscopy was also performed at the 1.5 m Danish and 3.6 m telescopes of the ESO at La Silla for a number of selected bright objects in the X-ray error circle. When the observing conditions were adequate we also performed absolute calibration both for photometry and spectroscopy. Spectroscopic data were analyzed by the "long" context in the ESO−MIDAS package (96NOV and later versions). Medium resolution spectroscopy of XTE J0111.2−7317 was also obtained with the ESO 1.5 m telescope. These observations were reduced with the FIGARO package (Shortridge et al. 1997) of the Starlink suite. The S/N of the spectra obtained at the 3.6 m was always of the order of 150 (or better) while those obtained at the ESO 1.5 m had a S/N around 50. The spectral features discussed in the text were present in at least two spectra or clearly above the noise level.

A full record of the optical observations is reported in Table 1.

3. Discussion

3.1. RX J0052.1−7319

The X-ray transient RX J0052.1−7319 was discovered by Lamb et al. (1999) with the analysis of ROSAT HRI and BATSE data. The object showed a period of 15.3 s (Kahabka 1998, 1999; Sasaki et al. 2000; for PSPC data analysis see also Haberl et al. 2000) and a flux in the 0.1−2 keV band of $2.6 \times 10^{-11}$ erg s$^{-1}$ cm$^{-2}$.

We devoted particular care to the definition of the X-ray error circles of this source. The archival ROSAT HRI images, available for RX J0052.1−7319, were analyzed using both a sliding cell and a wavelet transform−based algorithm (Lazzati et al. 1999; Campina et al. 1999). These techniques, already applied by us to the ROSAT HRI fields of several X-ray sources (Israel et al. 1999a; Israel et al. 2000a; Covino et al. 2000; Israel et al. 2000b; Israel et al. 2000c; Mereghetti et al. 2001), allowed us to infer a positional accuracy in the $5 \div 10''$ range. In particular RX J0052.1−7319 was observed three times in the BMW−HRI source catalog (see Table 2). The source is highly variable with a $\gtrsim 1,000$ variation in the count rate. The source is serendipitously detected at relatively large off−axis angles (> 8′) affecting the source positioning. For each field we are able to perform a boresight correction, registering the X-ray source positions to an optical frame. In the case of the shortest observation, we did not find X-ray sources suitable for the boresight. For the remaining two observations we used for the boresight correction 3 and 1 sources, respectively. Boresight errors are summed up in quadrature with the source location error (see Table 2). We then combined the source positions obtained in the three observations to derive the final position: $RA = 00^{h}52^{m}13^{s}27$,
DEC = −73°19′19.5″ (J2000). The 90% error circle radius amounts to 3.3″.

An optical counterpart for this source was proposed by Israel et al. (1999b), a B-type star with R = 14.54 ± 0.03 and V − R = +0.08 ± 0.04, located at coordinates (J2000, estimated uncertainty 1″) RA = 00°52′13.900, DEC = −73°19′19.5″. In Fig. 1 a R frame for the RX J0052.1−7319 field is shown, together with the X-ray error circle. The proposed counterpart is the object labeled as A. This star was also identified in the OGLE photometric database (labeled as SMC_SC699923, Udalsky et al. 1999) as a long-term variable star with quasiperiodic light variation of amplitude 0.13 mag in the I band. This optical object is within the X-ray error circle of RX J0052.1−7319, lying at 2.8″ from the best X-ray position.

A 1000 s optical spectrum at the 1.5 m Danish telescope of the ESO at La Silla (Fig. 2) for object A shows Hα and Hβ emission (EW ~ −12 and ~3 Å, respectively), supporting the Be nature of the star. Star B, just outside the X-ray error circle, does not show any emission line. Its spectrum is compatible with being a moderately bright young SMC star somewhat later than object A.

### Table 2. Source characteristics of RX J0052.1−7319 in the BMW−HRI catalog.

<table>
<thead>
<tr>
<th>Observation (BOJ)</th>
<th>RA (J2000)</th>
<th>DEC (J2000)</th>
<th>Error (″)</th>
<th>Off-axis (arcmin)</th>
<th>Exp. time (s)</th>
<th>Count rate (c s⁻¹)</th>
</tr>
</thead>
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<td>rh300513a00</td>
<td>00 52 12.60</td>
<td>−73 19 19.7</td>
<td>10.00</td>
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<tr>
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<td>−73 19 19.3</td>
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<td>8.4</td>
<td>27395.0</td>
<td>0.9989</td>
</tr>
<tr>
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<td>16.7</td>
<td>17064.5</td>
<td>0.6863</td>
</tr>
</tbody>
</table>

Fig. 1. 1.5 m ESO Danish telescope RX J0052−7319 R image. The X-ray error circle is also shown with the optical counterpart labeled as A. A nearby object is labeled as B.

Fig. 2. 1.5 m ESO Danish telescope RX J0052−7319 spectrum for objects A and B. Position of the main hydrogen lines is also shown.
we can assume that the interstellar reddening takes a value $E_{B-V} = 0.08$ typical for the SMC and that the rest of the reddening is circumstellar, in which case $M_V = -4.4 \pm 0.15$. This value is compatible with a B0V star with $E_{B-V} = 0.3$, though we note that such circumstellar emission would be rather high. Even though the weakness of all the Si lines (SiIV λ4089 Å is hidden on the blue wing of Hδ) would support a main sequence classification for a Galactic object, we note that the metal content of the SMC is much lower than that of the Milky Way and that even supergiants in the SMC can have metallic lines as weak as those of main-sequence objects in our Galaxy (Walborn 1983; see also Lennon 1997). As a consequence, we find that we cannot decide on the luminosity class of the star without knowing how much of the reddening is circumstellar. If it is a giant, the presence of weak HeII λ4200 Å would indicate spectral type O9.5IIf, while B0Ve would be the spectral classification if it is on the main sequence.

3.2. XTE J0111.2–7317

The X-ray transient XTE J0111.2–7317 was discovered by the RXTE X-ray observatory in November 1998 (Chakrabarty et al. 1998a). Analysis of ASCA observation (Chakrabarty et al. 1998b; Yokogawa et al. 2000b) identified this source as a as a 31 s X-ray pulsar with a flux in the 0.7 – 10 keV band of $3.6 \times 10^{-10}$ erg s$^{-1}$ cm$^{-2}$ and $\sim$45% pulsed fraction. The detection was also confirmed from the BATSE telescope on the CGRO satellite which detected the source in the hard 20 – 50 keV band with a flux ranging from 18 to 30 m Crab (Wilson & Finger 1998).

The best X-ray position for XTE J0111.2–7317 was inferred from the ASCA observation carried out during the 1998 outburst (Yokogawa et al. 2000b). Moreover, we applied the new calibration for the restoration of the ASCA source position accuracy applying the algorithm described by Gotthelf et al. (2000). This yielded the following improved position of the source for the ASCA–SIS observation: RA = 1$^{h}$11$^{m}$11.5$^{s}$, DEC = $-73^\circ$16'44.6" (J2000), with an error radius of 15" (90% confidence). As a reference, we also report the position as inferred in the analysis of the ASCA–GIS observation: RA = 1$^{h}$11$^{m}$12.3$^{s}$, DEC = $-73^\circ$16'47.6" (J2000), with an error radius of 25" (90% confidence).

No ROSAT observations were able to detect this source. Summing up all the 6 available PSPC observations we are able to derive a $3 \sigma$ upper limit of $1.4 \times 10^{-3}$ c s$^{-1}$ (0.1–2.4 keV, for a total exposure time of $\sim$90 ks), which translates to a luminosity limit of $\sim 3 \times 10^{34}$ erg s$^{-1}$ in the 0.1 $\div$ 2.4 keV energy band for a Crab–like spectrum and a column density of $10^{21}$ cm$^{-2}$. These observations clearly indicate the transient nature of this source.

A Be star counterpart was proposed by Israel et al. (1999b) and subsequently confirmed by Coe et al. (2000) as a B0–B2 star with strong evidence for the presence of a surrounding nebula, possibly a SuperNova Remnant (SNR).

Images for this field were taken (see Fig. 4) revealing two bright stars within the 15" error circle. The fainter one (star B) is a B-type star ($R = 15.29 \pm 0.03$, $V - R = +0.06 \pm 0.04$), located at the coordinates (J2000, estimated uncertainty 1") RA = 01$^{h}$11$^{m}$08.4$^{s}$, DEC = $-73^\circ$16'46.9".
A 1,000 s spectrum taken with the 1.5 m Danish telescope (Table 1) revealed strong Hα and Hβ emission lines (EW ~ 21 and 1.5 Å, respectively), indicating that this source is likely the counterpart of the X-ray transient. The brighter star (R = 14.30 ± 0.03) does not show any emission line. No other object brighter than R ~ 17.3 is found within the error circle. One more B-type star (R = 14.55 ± 0.03, V − R = +0.05 ± 0.04), located at the coordinates (2000, estimated uncertainty 1") RA = 01h11m25.9s, DEC = −73°17′27″ (1′ away from the X-ray position, star C) also shows strong Hα and Hβ emission lines (EW = 36 and 3 Å, respectively).

In order to infer an accurate spectral classification of the selected objects we observed again the field and, on 1999, September 15th, we set the slit so that the bright Be star just outside the error circle (star C) was also included in the slit. This star can be seen on the Eastern edge of all the images in Coe et al. (2000) and in our Fig. 4. The orientation of the slit was therefore almost E–W. On 1999 November 1st, on the other hand, we set the slit so that the core of the dark nebulous region that can be seen in Fig. 2 of Coe et al. (2000) to the SW of star B was inside the slit.

The two spectra of star B in the classification region are displayed in Fig. 6. The spectrum from November 1st has a higher resolution, but the spectrum from September 15th has rather higher S/N. Also displayed is the spectrum of star C, which shows Hβ in emission and emission components in other Balmer series lines. The measured Hβ EW = −3.3 Å is typical of Be stars with strong emission. This is confirmed by the red spectrum of the source (see Fig. 7), where Hα is strongly in emission (EW = −40 Å). Typically, Be stars with such strong emission in the Balmer series also display many weak emission lines corresponding to low–ionization metallic transitions (mainly Fe II). The emission spectrum of star C is clearly seen in Fig. 7.

The blue spectrum of star C shows weak He II lines, indicating a spectral type earlier than B0.5. The weakness of all metallic lines (and specifically the Si IV lines in the wings of Hδ) suggests a main sequence object (though see previous section). The weak He II λ 4269, 4541 Å indicate a spectral type O9.5–B0, while He II λ 4686 Å < He I λ 4713 Å favors the latter classification. Assuming a distance modulus (m − M) = 18.8 (Coe et al. 2000), the measured V = 14.6 implies a MV = −4.2, consistent with a B0V star (Vacca et al. 1996). Considering an interstellar reddening A0.14 (Coe et al. 2000), the source must actually have MV = −4.6, which is slightly too high for a main sequence object. However, Be stars tend to be brighter than that corresponds to their spectral type due to the continuum emission from the circumstellar disc. Zorec & Briot (1991) and Fabregat & Torrejón (1998) give an average difference of 0.3 mag between Be stars and main–sequence objects. Using the approximate relationship of Fabregat & Reglero (1990), which is valid for isolated Be stars, an Hα EW = −40 Å implies a circumstellar brightening ΔV ∼ 0.3. Therefore star C has a spectral type (B0Ve) and spectral morphology typical of a counterpart to Be/X–ray binary.

The spectrum of star B, on the other hand, does not show obvious He II lines, indicating a later spectral type. The only photospheric features that could be visible in its spectrum and are not marked on the spectrum of star C are the O II blends at λλ 4146−53 and 4185−90 Å, He I λ 4169 Å and perhaps N II λ 4237−4242 Å. These are possible identifications for features that are visible in both spectra of the source. However, from the two spectra, it is clear that longwards of ≈λ 4200 Å the observed “continuum” actually consists of a forest of weak emission lines, that create apparent absorption features, such as the shallow band at ≈λ 4600 Å in the November spectrum, which is probably only a gap between emission components centered on Fe II λ 4584 Å and Fe II λ 4629 Å.

If the weak feature seen at ≈λ 4689 Å in both spectra is real and can be identified with He II λ 4686 Å (with the radial velocity shift of 165 km s−1 measured by Coe et al. 2000), then the spectral type would be B0.5–B0.7. The presence of the strong O II/C IV blend at ≈λ 4650 Å indicates that it cannot be later than B1 in any case.

Longwards of Hβ the Fe II emission spectrum is very strong (see Fig. 7). Close to λ 5000 Å, apart from Fe II λ 5018 Å, two other lines, which are particularly strong in the November spectrum, can be seen. Their wavelengths suggest identification with the [O III] λ 4959 and 5007 Å lines. Such lines are not generally seen on stellar spectra, but are typical of nebular emission.
Fig. 5. 1.5 m ESO Danish telescope XTE J0111.2–7317 spectrum for objects A, B (left panel), C and D (right panel) in the X-ray error circle. Position of the main hydrogen lines is also shown.

Fig. 6. The two blue spectra of the optical counterpart to XTE J0111.2–7317 (top) compared to that of the B0Ve star just outside the error circle (star C, bottom). The lines marked on the spectrum of star C are characteristic of the spectral type (compare with the spectroscopic standard in Fig. 3). All spectra have been divided by a spline fit to the continuum for normalization and smoothed with a $\sigma = 1.2$ Gaussian filter. All identified features are present in at least two spectra.

Diffuse emission is actually apparent on the region immediately to the East of star B in our September image. When the sky spectrum from that area is extracted and a sky spectrum from a distant area is subtracted, strong emission lines corresponding to [O ii] $\lambda\lambda 3727$ Å, H$\alpha$, H$\beta$, H$\gamma$, [O iii] $\lambda\lambda 4959, 5007$ Å and presumably [S ii] $\lambda 6716$ Å can be seen.

In the image corresponding to the November observation, diffuse emission is present on the region immediately to the SouthWest of star B. [O ii] $\lambda 3727$ Å, H$\beta$, and [O iii] $\lambda\lambda 4959, 5007$ Å are readily seen in the extracted sky spectrum.

Extracting several sky spectra around star B, we come to the conclusion that most of the nebular line emission comes from the immediate vicinity of the star and that it is strongest on the spatial position of the star. Therefore we conclude that a significant portion of the emission seen on star B comes from the nebulosity surrounding it and we are tempted to attribute (at least part of) the obvious differences between the September and November spectra to the different slit orientations.

However, it is clear from the detection of many permitted Fe ii emission lines that star B is also intrinsically an emission line star, i.e., is surrounded by a circumstellar envelope. As a matter of fact, we believe that star B is also a Be star, since other possibilities seem to be ruled out because of the following points: (i) the detected forbidden lines are typical of high excitation nebulae and not of B[e] stars. (ii) The presence of a compact companion precludes the possibility that star B is a Herbig Be star. (iii) The EWs of H$\alpha$ and H$\beta$ are typical of Be stars and very low for any B[e] star. (iv) If star B was a B[e] star, it should show narrow emission lines corresponding to [Fe ii].
Fig. 7. Red spectra of the optical counterpart to XTE J0111.2−7317 (top) and the B0Ve star just outside the error circle (star C, bottom). The many permitted emission lines on the spectrum of the optical counterpart to XTE J0111.2−7317 indicate that it is intrinsically a Be star, while the presence of weak forbidden lines must be due to emission from the surrounding nebula. Both spectra have been divided by a spline fit to the continuum for normalization and smoothed with a $\sigma = 1.2$ Gaussian filter.

<table>
<thead>
<tr>
<th>Field</th>
<th>Star</th>
<th>B</th>
<th>V</th>
<th>R</th>
<th>Spectral type</th>
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<td></td>
<td>B</td>
<td>15.82</td>
<td>15.92</td>
<td>16.05</td>
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</tr>
<tr>
<td>XTE J0111.2−7317</td>
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<td>15.32</td>
<td>14.65</td>
<td>14.30</td>
<td></td>
</tr>
<tr>
<td></td>
<td>B*</td>
<td>15.42</td>
<td>15.36</td>
<td>15.29</td>
<td>B0.5−B1Ve</td>
</tr>
<tr>
<td></td>
<td>C</td>
<td>14.63</td>
<td>14.61</td>
<td>14.55</td>
<td>B0Ve</td>
</tr>
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<td></td>
<td>D</td>
<td>15.05</td>
<td>13.67</td>
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</tr>
</tbody>
</table>

Moreover, the $E(B−V) = 0.2−0.3$ derived by Coe et al. (2000) shows that the obscuration by the nebula is small. Therefore the luminosity of star B is far too low for a B[e] star.

Therefore star B has a spectral type B0.5−B1Ve, compatible with the $M_V = −3.8$ derived by Coe et al. (2000) and it is the likely counterpart to XTE J0111.2−7317. Its association with the nebulous object that apparently surrounds it is almost certain, though further study is necessary in order to assess its nature.

We note that our value $(B−V) = 0.05$ would imply an $E(B−V) = 0.28$, consistent with the value derived by Coe et al. (2000) from Strömgren photometry (though hardly consistent with the value they derived from Johnson photometry, which has $(B−V) = −0.08$. However, the two Coe et al. measurements of the $V$ magnitude reported for this star show a spread that is much larger than the measurement errors, indicating variability.

4. Conclusions

We discuss the optical counterparts of two transient X-ray pulsars in the SMC. In the case of RX J0052.1−7319, ROSAT HRI observations allowed us to derive an error circle of $\sim 3''$ radius (90%). Accurate positioning can then be achieved with past HRI instruments if a boresight correction can be derived. Within the error circle only one relatively bright source is found with $V = 14.6$ ($B − V = 0.1$). The object has been investigated photometrically and spectroscopically leading to classify it as a O9.5IIe star (a classification as a B0Ve star is also possible, since the luminosity class depends on the adopted reddening). Medium resolution spectra show prominent Hα and Hβ emission lines, further strengthening the identification.

The position of XTE J0111.2−7317 is based on ASCA observations analyzed by the new calibration for the restoration of the ASCA source position accuracy described by Gotthelf et al. (2000). The analysis provided a 15'' error box. Within it we found a relatively bright object ($V = 15.4$, $B − V = 0.06$) which can be classified as a B0.5−B1Ve star. Also in this case we easily detect Balmer emission lines. A further bright B0Ve star was found outside the X-ray error circle. Coe et al. (2000) found evidence for the presence of a surrounding nebula, possibly a supernova remnant, around the likely counterpart of XTE J0111.2−7317.

Acknowledgements. SCo thanks the Rome Astronomical Observatory for the kind hospitality. During part of this work, IN was supported by an external ESA fellowship. This work is partially supported through CNAA, ASI and MURST grants. We also thank the anonymous referee for his quick reply and useful comments, suggestions and corrections that increased the readability of the paper.
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