

Erratum

**A search for near-infrared molecular hydrogen emission
in the CTTS LkH α 264 and the debris disk 49 Ceti**

A. Carmona^{1,2}, M. E. van den Ancker², Th. Henning¹, M. Goto¹, D. Fedele^{2,3}, and B. Stecklum⁴

¹ Max Planck Institute for Astronomy, Königstuhl 17, 69117 Heidelberg, Germany
 e-mail: carmona@mpia.de

² European Southern Observatory, Karl Schwarzschild Strasse 2, 85748 Garching bei München, Germany

³ Dipartimento di Astronomia, Università di Padova, Vicolo dell'Osservatorio 2, 35122 Padova, Italy

⁴ Thüringer Landessternwarte Tautenburg, Sternwarte 5, 07778 Tautenburg, Germany

A&A 476, 853–862 (2007), DOI: 10.1051/0004-6361:20078536

Key words. stars: emission-line, Be – stars: pre-main sequence – planetary systems: protoplanetary disks – errata, addenda

In the text (end of Sect. 4.1) and in Table 4 of the original paper, the $U - V$ colour of LkH α 264 was mistakenly taken as -0.46 . The $U - V$ colour of LkH α 264 should be 0.37 (Bastian & Mundt 1979). Therefore, in Table 4 of the original paper, in the fields corresponding to LkH α 264, $(U - V)_{\text{obs}}$ should be 0.37 instead of -0.46 , $(U - V)_{\text{dered}}$ should be 0.08 instead of -0.74 , and $(U - V)_{\text{ex}}$ should be -2.10 instead of -2.92 . In Table 4 we present the correct input of LkH α 264 for Table 4 of the original paper. In Fig. 3 we present the corrected version of Fig. 3 of the original paper, here taking into account the correct value of the $U - V$ excess for LkH α 264. This correction on the $U - V$ excess of LkH α 264 affects no conclusion in the original paper, therefore, the conclusions presented in the original paper remain unchanged.

Acknowledgements. We thank G. Meeus for bringing the error on the $U - V$ colour of LkH α 264 to our attention.

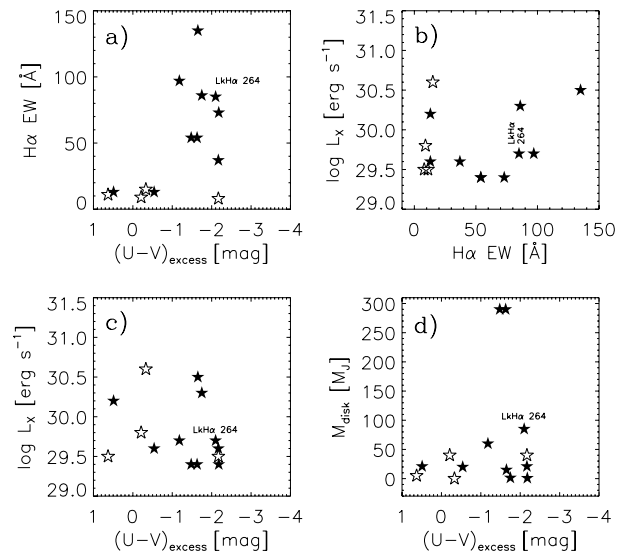


Fig. 3. Physical properties of classical T Tauri stars in which a search for the H $_2$ $v = 1-0$ S(1) line was performed. Filled stars represent detections, non-filled stars represent non-detections. **a)** H α EW versus $(U - V)_{\text{excess}}$. **b)** $\log L_X$ [erg s $^{-1}$] versus H α EW. **c)** $\log L_X$ [erg s $^{-1}$] versus $(U - V)_{\text{excess}}$. **d)** M_{disk} versus $(U - V)_{\text{excess}}$.

Table 4. Physical properties of LkH α 264.

Star	Sp.T.	\dot{M} [$\times 10^{-7} M_{\odot} \text{ yr}^{-1}$]	EW H α [\AA]	A_V [mag]	$(U - V)_{\text{obs}}$ [mag]	$(U - V)_{\text{dered}}$ [mag]	$(U - V)_{\text{ex}}$ [mag]	$\log L_X$ [erg s $^{-1}$]	M_{disk} [M_J]	Ref. H $_2$
LkH α 264	K5 Ve	1–10	85	0.52	0.37	0.08	-2.10	29.7	85	1, 2

References: (1) Carmona et al. (2007); (2) Itoh et al. (2003).

References

- Bastian, U., & Mundt, R. 1979, A&AS, 36, 57
 Carmona, A., van den Ancker, M. E., Henning, T., et al. 2007, A&A, 476, 853
 Itoh, Y., Sugitani, K., Ogura, K., & Tamura, M. 2003, PASJ, 55, L77