

## **BeppoSAX/PDS serendipitous detections at high galactic latitudes**

R. Landi<sup>1,2</sup>, A. Malizia<sup>1</sup>, and L. Bassani<sup>1</sup>

<sup>1</sup> INAF – Istituto di Astrofisica Spaziale e Fisica Cosmica, Sezione di Bologna, via Gobetti 101, 40129 Bologna, Italy  
e-mail: landi@bo.iasf.cnr.it

<sup>2</sup> Dipartimento di Fisica, Università degli Studi di Bologna, Viale C. Berti Pichat 6/2, 40127 Bologna, Italy

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**Abstract.** At a flux limit of  $\sim 10^{-11}$  erg cm<sup>-2</sup> s<sup>-1</sup> in the 20–100 keV band, the PDS instrument on-board *BeppoSAX* offers the opportunity to study the extragalactic sky with an unprecedented sensitivity. In this work we report on the results of a search in the *BeppoSAX* archive for serendipitous high energy sources at high galactic latitudes ( $|b| \geq 13^\circ$ ). We have defined a set of twelve regions in which the PDS/MECS cross-calibration constant is higher than the nominal value. We attribute this mismatch to the presence of a serendipitous source in the PDS field of view. In four cases the likely high energy emitter is also present in the MECS field of view. In these cases, we have performed a broad band spectral analysis (1.5–100 keV) to understand the source spectral behaviour and compare it with previous *BeppoSAX* observations when available. In eight cases the identification of the source likely to provide the PDS spectrum is based on indirect evidence (extrapolation to lower energies and/or comparison with previous observations). This approach led to the discovery of six new hard X-ray emitting objects (PKS 2356–611, 2MASX J14585116–1652223, NGC 1566, NGC 7319, PKS 0101–649 and ESO 025–G002) and to the presentation the PDS spectrum of NGC 3227 for the first time. In the remaining five cases we provide extra *BeppoSAX* observations that can be compared with measurements already published and/or in the archive.

**Key words.** X-rays: general – X-rays galaxies – methods: data analysis

### **1. Introduction**

The hard X-ray sky is still poorly explored at high galactic latitudes and the only truly all sky survey performed so far above  $\sim 15$  keV dates back to the 1980s (Levine et al. 1984). This pioneering work was performed with the *HEAO 1* A4 experiment and detected  $\sim 70$  sources in the 13–80 keV band down to a flux limit of  $2 - 3 \times 10^{-10}$  erg cm<sup>-2</sup> s<sup>-1</sup> with an angular resolution of  $\sim 3^\circ$ ; of these sources sixteen were located at high galactic latitudes ( $|b| \geq 13^\circ$ ) and only seven were of extragalactic nature.

A step forward will be provided by the imager on-board *INTEGRAL* which is surveying a large fraction of the sky with a sensitivity better than a few mCrab<sup>1</sup> in the 20–100 keV energy range and an angular resolution of a few arcmin (Bassani et al. 2004; Bird et al. 2004); this exploratory work will be followed by the *SWIFT* mission which is expected to survey the hard X-ray sky with a sensitivity of  $\sim 1$  mCrab at high galactic latitudes (Gehrels et al. 2004). In the meantime, the *BeppoSAX* archive can be used to probe the extragalactic X-ray sky in the same 20–100 keV energy band. Pointed observations of *BeppoSAX*/PDS have unveiled many hard X-ray emitting AGN and provided the best spectroscopy of this type of object above 10 keV. However, observations were sometimes limited by the

lack of imaging capability: contaminating sources were found inside the target field of view as well as in the offset fields used for background measurements. These data have often been neglected by the original observers although they provide another powerful tool with which to study the extragalactic sky above 10 keV: new sources can be found and spectroscopically measured, while known objects can be re-observed and compared to previous measurements in a search for variability, so far poorly explored. Furthermore, the argument put forward by Fabian (2001) that two of the three nearest AGN have very high column densities ( $> 10^{24}$  cm<sup>-2</sup>) indicates the need to constrain the statistics of highly absorbed AGN: increasing the number of such objects known is therefore an important task. In view of these facts, we have carried out a program to search systematically in the *BeppoSAX*/PDS archive (which at the moment is being reanalysed using the *XAS* software package) for observations that indicate the presence of a contaminating source either in the pointed field of view (search mode 1) or in the offset fields used for background measurements (search mode 2). Results related to search mode 1 are presented in this work, while a future work will be devoted to results found in search mode 2.

The paper is organized as follows: in Sect. 2 the selected sample is described, in Sect. 3 the observations, data reduction techniques and analysis procedures are presented, while Sect. 4

<sup>1</sup> For a Crab-like spectrum 1 mCrab corresponds to  $2 \times 10^{-11}$  erg cm<sup>-2</sup> s<sup>-1</sup> in the 20–100 keV energy band.

is devoted to the discussion of the results. Conclusions and future work are summarized in Sect. 5.

## 2. Sample selection and contaminating source search

Although *BeppoSAX* was not designed to perform an X-ray survey, a systematic search in the entire PDS data archive could allow the discovery of new sources, down to a flux limit of  $\sim 10^{-11}$  erg cm $^{-2}$  s $^{-1}$ . In particular, the simultaneous monitoring of three 1:3 sky regions (target field + two offset fields) allowed the PDS to survey a substantial fraction of the sky over the 15–100 keV energy range.

The PDS instrument has a hexagonal field of view of 1:3  $\times$  1:3 *FWHM* and no imaging capability; its positional uncertainty can be approximated by an error box circle of 1:3 in radius. The MECS has instead a field of view of 30' radius and so covers about 25% of the PDS area. It is therefore likely that a serendipitous source can be undetected by the MECS, but still observed by the PDS.

In this paper we concentrate on the study of those sources discovered in the pointed fields of view or detected in search mode 1. The search in this mode was performed in the following way: first, we visually inspected in the *BeppoSAX* archive ([www.asdc.asi.it](http://www.asdc.asi.it)) all observations performed above 13° in galactic latitude and available to the public as of October 2001 (635); from this set of data we extracted a sample containing those sources which clearly show a mismatch between the MECS and PDS spectra in the standard archive analysis: this mismatch was taken as strong evidence for the presence of a contaminating source in the PDS field of view. From this preliminary list we excluded all those sources that were likely to be Compton thick on the basis of various considerations (i.e. large iron line equivalent width, X-ray to [O<sub>III</sub>] ratio less than 1, extreme absorption): in the spectrum of these types of sources a MECS/PDS spectral mismatch is simply due to a more complex spectral shape than used for the quick look analysis. We then performed a cut in the signal to noise ratio, accepting a source only if it had at least 3 $\sigma$  detection in the PDS. We then confirmed with our own analysis the MECS/PDS mismatch by fitting the MECS and PDS data with a simple model, generally a power law either unabsorbed or absorbed: only when the cross-calibration constant between these two instruments was significantly outside the nominal range of 0.75–0.95 (Fiore et al. 1999a), i.e. greater than 2, was a source maintained. Overall, this analysis provided the sample of twelve regions discussed in the present paper. Although this sample is not complete (not all observations performed by *BeppoSAX* were screened), this search gives an idea of the extragalactic sources that populate the hard X-ray sky; in this sense the sources serendipitously found in this work can be “loosely” taken as representative of the AGN population in the 20–100 keV band.

An important step in the search described above is the reduction of the PDS spectra extracted using the *XAS* v2.1 package (Chiappetti & Dal Fiume 1997): this provides slightly different results than *SAXDAS*, i.e. the package usually used to perform the archive analysis. A preliminary comparison

between PDS spectra extracted by means of these two software packages is under way<sup>2</sup> and will be presented in a future work. The analysis performed on sources of different intensity show that the spectral parameters do not change when computed with the different packages, but the associated errors are smaller when using *XAS*. A significant improvement in the signal to noise ratio is obtained by means of *XAS*. Furthermore, the *XAS* package allows a more reliable check of the background fields by taking advantage of the rocking technique (Frontera et al. 1997b); this is important when the source is faint in the PDS as is often our case. When one collimator is pointing ON source, the other collimator is pointing in one of the two OFF positions. The standard stay time in each position of either collimator was 96 s. At each cycle the two collimators were swapped: the one pointing to the source was moved to monitor the background and vice versa. In this way we can obtain, in addition to the target observation, two independently accumulated spectra of the two +OFF and –OFF fields offset by 210' with respect to the main pointing; these offsets are used as background in the computation of the spectrum of the target source. The comparison between the spectra of these two offset fields and, in particular, the difference between the +OFF and –OFF spectra in count rate is a good diagnostic tool to investigate the presence of contamination.

If no contamination is present in either of the offset fields, we expect the difference of their count rates to be compatible with zero; on the contrary, a positive excess of counts in the difference indicates the presence of contamination in the +OFF field; a negative count rate provides evidence for contamination in the –OFF field.

Applying this method to our sample, we found the presence of a contaminating source in the +OFF field of the first observation of AD Leonis (excess:  $0.189 \pm 0.070$  counts s $^{-1}$ , 2.7 $\sigma$ ) and in part of the first observation of VW Cephei (excess:  $0.139 \pm 0.043$  counts s $^{-1}$ , 3.2 $\sigma$ ), while in part of the observations of MKN 1073 and NGC 7552 the –OFF fields show an excess of  $0.159 \pm 0.060$  counts s $^{-1}$  (2.7 $\sigma$ ) and  $0.203 \pm 0.056$  counts s $^{-1}$  (3.6 $\sigma$ ), respectively. To give an idea of the effect of contamination in an offset field on the PDS data, we have considered two extreme cases: a  $\sim 3\sigma$  excess reduces the source count rate by  $\sim 5\%$ , while a 10 $\sigma$  contamination provides a reduction of around 15%.

In order to extract the uncontaminated source spectra, we excluded the contaminated fields and considered only the uncontaminated ones in the computation of the background for these particular sources.

After assembling the sample, we search for the likely contaminating sources, adopting the following strategy. First, we considered all the sources present in the MECS image in addition to the target source and checked their consistency with the PDS data, i.e. if the cross-calibration constant fell within the nominal interval. If no sources in the MECS field of view matched the PDS data, then, we searched for likely high energy emitters located inside the PDS field of view but not observed

<sup>2</sup> On behalf of the PDS group, see <ftp://ftp.tesre.bo.cnr.it> in the directory /pub/sax/doc/software\_docs/xas\_vs\_saxdas.ps

**Table 1.** The list of serendipitous PDS detections found in the *BeppoSAX* archive.

Pointed Source	Obs. date	Type	Calib <sup>(a)</sup>	Contaminating Source	Type
IRAS 01025–6423	24/03/2001	Seyfert 2	18–1324	PKS 0101–649	Quasar
MKN 1073	15/02/1999	Seyfert 2	485–2035	Perseus/NGC 1275 (?)	Cluster/Seyfert 2
NGC 1553	16/01/1997	Normal Galaxy	14–46	NGC 1566	Seyfert 1
	16/11/1997				
AD Leonis	23/04/1997	Flare Star	152–230	NGC 3227	Seyfert 1
	01/05/1999				
	08/05/1999				
	12/05/1999				
ON 325	23/12/1998	Blazar	11–57	MKN 766 <sup>(b)</sup>	Seyfert 1
NGC 5793	25/07/2001	Seyfert 2	32–1384	2MASX J14585116–165 <sup>(b)</sup>	Galaxy
1E 1839.6+8002	16/10/2000	Flare Star	83–181	3C 390.3 <sup>(b)</sup>	Seyfert 1
	03/02/2001				
H1846–786	08/03/2001	Seyfert 1	2–4	ESO 025–G002	Seyfert 1
VW Cephei	07/05/1998	Eclipsing Binary	164–417	4C +74.26	Quasar
	07/10/1998				
NGC 7331	10/06/2000	LINER	28–121	NGC 7319	Seyfert 2
NGC 7552	23/12/1999	Galaxy	150–309	NGC 7582	Seyfert 2
SCG 2353–6101	28/11/1996	Cluster	8–28	PKS 2356-611 <sup>(b)</sup>	Seyfert 2

<sup>(a)</sup> PDS/MECS cross-calibration constant.

<sup>(b)</sup> Source located within the MECS field of view (see text).

by the MECS given the significantly different fields of view. In particular, when adopting this second approach, we focused our search on bright sources in the 2–10 keV band as these are the most likely to contaminate the PDS observation and are expected to appear in the HEASARC X-ray archives. In Table 1 we list the twelve cases we found, reporting in each case the *BeppoSAX* observation target with its relative observation date and object type, the PDS/MECS cross-calibration constant obtained by fitting the data with a simple power law and finally the name and type of the contaminating source found. It is evident from the Table that the cross-calibration constant is always outside the nominal range, confirming the presence of one or more contaminating objects in the PDS field of view. Extrapolation of a more complex model from the MECS to the PDS energy range still provides in all cases a high cross calibration constant. We found only four fields where the contaminating source is so close ( $\sim 25'$ ) to the *BeppoSAX* target that it is also detected by the MECS instrument.

### 3. Observations and data analysis

In this work we made use of data from three of the *Narrow Field Instruments* (NFIs) on-board the Italian–Dutch satellite *BeppoSAX* (Boella et al. 1997a): the Low Energy Concentrator Spectrometer (LECS, 0.1–10 keV, Parmar et al. 1997), the Medium Energy Concentrator Spectrometer (MECS, 1.3–10 keV, Boella et al. 1997b) and the Phoswich Detection System (PDS, 15–300 keV, Frontera et al. 1997b).

For all sources, the LECS and MECS spectra were downloaded from the *BeppoSAX* archive; the analysis of the on-axis source was standard. For all the off-axis sources detected in the MECS field of view, contaminating or not, the MECS spectra were extracted from a region centered on the source and having a radius chosen according to the criteria suggested by Fiore et al. (1999a). For these sources, because of the lack of the appropriate ancillary response files, the LECS spectral analysis could not be performed. The background subtraction for the on-axis sources was performed using blank sky spectra extracted from the same region of the source, while for the off-axis sources we used a local background spectrum (extracted from a region with a radius equal to the source extraction radius) to account for possible contaminating effects inside the MECS field of view.

For the PDS data reduction, source visibility windows were selected following the criteria of no Earth occultation and high voltage stability during the exposure. In addition, the observations closest to the South Atlantic Anomaly were discarded from the analysis.

The LECS and MECS spectra were rebinned in order to sample the energy resolution of the detectors with an accuracy proportional to the count rate. The PDS data were instead rebinned to have logarithmically equal energy intervals. The data rebinning also required that there were at least 20 counts in each bin so that the  $\chi^2$  statistic could reliably be used.

The energy bands used for spectral fitting were limited to those where the response functions are well known, i.e. 0.1–4.5 keV, 1.5–10.5 keV and 15–100 keV, for the LECS MECS and PDS, respectively.

**Table 2.** Contribution to the PDS flux of the sources detected in the MECS field of view.

Field	Source	Extrapolated flux <sup>(a)</sup> [20–100] keV	Contribution (%)	Offset <sup>(b)</sup> (arcmin)	$R$ <sup>(c)</sup> (%)
ON 325	ON 325	1.2	10	–	–
	1AXG J121854+2957	0.56	3	16.2	78
	MKN 766	17.5	87	19.7	73
NGC 5793	NGC 5793	0.072	1	–	–
	NGC 5796	0.17	2	4.2	92
	2MASX J1458–165	7.72	95	13.4	81
1E 1839.6+8002	1E 1839.6+8002	0.017	<1	–	–
	3C 390.3	47.8	98	25.3	65
SCG 2353–6101	SCG 2353–6101	0.03	<1	–	–
	PKS 2356–611	25.2	99	16.1	78

<sup>(a)</sup> In units of  $10^{-12}$  erg  $\text{cm}^{-2}$   $\text{s}^{-1}$ .

<sup>(b)</sup> Distance between the pointing target and the off-axis sources.

<sup>(c)</sup> Source flux correction factor due to the PDS response to source an off-axis source.

The spectral analysis was performed using the XSPEC v11.2.0 software package (Arnaud 1996) and the instrument response matrices released by the *BeppoSAX* ASI Science Data Center. For the off-axis sources, we used the appropriate MECS ancillary response files to correct for the effects of vignetting due to the mirrors. In addition, we introduced a flux correction factor in the PDS band for each contaminating source in order to estimate the real flux at the source. This correction is simply a function of the distance of the source from the main target and is related to the PDS response: the reduction in sensitivity is of a factor of  $\sim 2$  at  $38'$  from the main pointing coordinates while at  $78'$  the response is zero (Frontera et al. 1997a).

Normalization constants have also been introduced to allow for known differences in the absolute cross-calibration between the detectors. The LECS/MECS cross-calibration constant was allowed to lie within the nominal range 0.7–1.0 (Fiore et al. 1999a), while the PDS/MECS cross-calibration was left free to vary in order to search for the presence of contaminating sources and then to estimate their broad band spectrum. The absorption of X-rays due to our galaxy in the direction of each object (Dickey & Lockman 1990) is added in all models of the spectral analysis. All quoted errors correspond to the 90% confidence interval for one interesting parameter ( $\Delta\chi^2 = 2.71$ ).

The main aim of this work is to estimate the PDS spectrum of each serendipitous source found. The method we used to achieve this was to fit simultaneously the LECS (not always available) and MECS data of the target source to determine the best-fit model to these data; afterwards, we performed an extrapolation of this best-fit model to the PDS energy band in order to estimate the contribution of the target source to the total PDS flux. Hence, by subtracting from the total PDS spectrum the contribution due to the on-axis source, we were able to estimate the high energy spectrum of the serendipitous/contaminating object.

Four sources in our sample (NGC 1553, 1E 1839.6+8002, VW Cephei and AD Leonis), have been observed by

*BeppoSAX* at different epochs (see Table 1). As a first step, our approach was to analyse each single pointing, for each source, individually. Although in some cases we found evidence for flux variability in the target and/or the serendipitous sources, the spectral parameters were always consistent within the respective uncertainties; therefore, we decided to sum, in all cases, all available pointings to improve the statistical quality of the data, reporting only the averaged spectral parameters.

## 4. Results

In the following, we briefly describe the main results obtained for each field. We discuss separately the four fields in which the contaminating object is detected in the MECS instrument because for these objects we can perform a broad band spectral analysis (1.5–100 keV), while in the remaining eight cases we can only present and discuss the PDS spectrum.

### 4.1. Sources detected in the MECS field of view

In this section, we describe the four fields where the object likely to be responsible for the PDS spectrum is also observed by the MECS instrument. In Table 2 (and also in Table 5 afterwards) we report for each of these fields all sources detected by the MECS, their fluxes extrapolated from the MECS to the PDS energy range (15–100 keV) assuming their best fit model, their relative contribution (in %) to the PDS flux, their distance from the target source and the relative flux correction factor,  $R$ . The observed flux is divided by  $R$  to reconstruct the flux at the source. It is evident from the values reported in Col. 4 (Contribution) which source is in each case responsible for most of the PDS flux. In Table 3 we list for each of these likely associations the MECS and PDS exposures and count rates, the 2–10 keV flux, the observed high energy flux (20–100 keV) and the effective flux (20–100 keV) at the source. For these four sources we jointly fitted the MECS and PDS data to provide a broad band (20–100 keV) spectrum. In Table 4 we summarize the best-fit parameters and also report the range of the

**Table 3.** MECS and PDS observation log.

Source	MECS		PDS		$F^{(a)}$	$F_{\text{Obs}}^{(a)}$	$F_{\text{Source}}^{(a)}$
	EXPO (ks)	cts/s	EXPO (ks)	cts/s	[2–10] keV	[20–100] keV	[20–100] keV
MKN 766	31.4	$0.0153 \pm 0.0011$	15.7	$0.152 \pm 0.050$	16.0	17.0	23.3
2MASX J1458–165	78.7	$0.0042 \pm 0.0003$	38.5	$0.086 \pm 0.029$	0.93	8.1	11.1
3C 390.3 <sup>(b)</sup>	157.3	$0.078 \pm 0.001$	80.0	$0.447 \pm 0.002$	23.0	22.0	43.0
PKS 2356–611	23.4	$0.029 \pm 0.001$	25.4	$0.205 \pm 0.041$	13.0	25.5	32.7

Note: Count rates are referred to the energy bands used in the spectral analysis (see text).

<sup>(a)</sup> In units of  $10^{-12}$  erg cm<sup>-2</sup> s<sup>-1</sup>.

<sup>(b)</sup> For this source the parameters of the average spectrum are reported.

**Table 4.** MECS and PDS spectral analysis of the serendipitous sources detected within the MECS field of view.

Source	$\Gamma$	$N_{\text{H}}^{(a)}$	$E_{\text{Line}}$ (keV)	$EW$ (eV)	Calib <sup>(b)</sup>	$\chi^2/\nu$
MKN 766	$1.93 \pm 0.10$	–	–	–	0.5–1.3	36.4/36
2MASX J1458–165	$1.20^{+0.26}_{-0.23}$	–	–	–	0.8–3.2	16.1/24
3C 390.3 <sup>(c)</sup>	$1.72^{+0.04}_{-0.05}$	<0.08	$6.37^{+0.29}_{-0.25}$	$115^{+86}_{-84}$	0.89–1.22	16.5/22
PKS 2356–611	$1.70^{+0.18}_{-0.17}$	$10.2^{+2.91}_{-2.31}$	$6.46^{+0.21}_{-0.17}$	$523^{+238}_{-246}$	0.75–0.95	48.5/41

<sup>(a)</sup> In units of  $10^{22}$  cm<sup>-2</sup>.

<sup>(b)</sup> PDS/MECS cross-calibration constant.

<sup>(c)</sup> For this source the parameters of the average spectrum are reported.

**Table 5.** Contribution to the PDS flux of all the MECS sources in the eight fields in which the contaminating source was not detected in the MECS field of view.

Field	Source	Extrapolated flux <sup>(a)</sup> [20–100] keV	Contribution (%)	Offset <sup>(b)</sup> (arcmin)	$R^{(c)}$ (%)
IRAS 010125–6423	IRAS 010125–6423	0.002	<1	–	–
MKN 1073	MKN 1073	–	–	–	–
NGC 1553	NGC 1553	1.5	7	–	–
Ad Leonis	AD Leonis	0.0002	<1	–	–
H1846–786	H1846–786	10	50	–	–
VW Cephei	VW Cephei	0.0004	<1	–	–
NGC 7331	NGC 7331	0.59	4	–	–
NGC 7552	NGC 7552	0.12	<1	–	–
	FRL 1041	1.78	4	18.9	76
	Sersic 159–03	0.003	<1	25.8	65

<sup>(a)</sup> In units of  $10^{-12}$  erg cm<sup>-2</sup> s<sup>-1</sup>.

<sup>(b)</sup> Distance between the pointing target and the off-axis sources.

<sup>(c)</sup> Source flux correction factor due to the PDS response to source not on-axis.

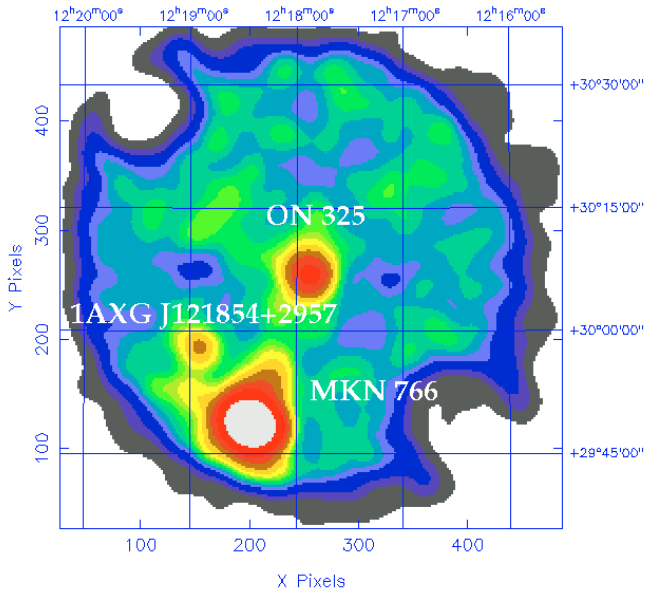
cross-calibration constant PDS/MECS (Col. 6). As is evident from the values of Col. 6, the normalization constants obtained are always compatible with the suggested value of 0.75–0.95. This is a further indication that these sources are indeed responsible for the high energy emission detected by the PDS instrument.

#### 4.1.1. The ON 325 field

Within the MECS field of view (see Fig. 1) of this pointing we find two other sources besides the target. The first, located at  $\alpha(2000) = 12^{\text{h}}18^{\text{m}}54^{\text{s}}.4$  and  $\delta(2000) = +29^{\circ}58'10''.5$ , is identified with the Seyfert 1.9 galaxy 1AXG J121854+2957, which

belongs to the *ASCA* Medium Sensitivity Survey (AMSS) (Ueda et al. 2001) and to the *BeppoSAX* High Energy Large Area Survey (HELLAS) (Fiore et al. 1999b). The second is located at  $\alpha(2000) = 12^{\text{h}}18^{\text{m}}25^{\text{s}}.2$  and  $\delta(2000) = +29^{\circ}48'48''.9$ ,  $\sim 20'$  from ON 325 and corresponds to the Seyfert 1 galaxy MKN 766.

Following Perri et al. (2003), the LECS and MECS data of the blazar ON 325 are best fitted with a broken power law, for which we obtain a  $\chi^2/\nu = 32.4/25$  and a 2–10 keV flux of  $\sim 8 \times 10^{-13}$  erg cm<sup>-2</sup> s<sup>-1</sup>. The two power laws intersect at  $E_B = 4.12^{+1.10}_{-0.83}$  keV and have photon indices  $\Gamma_1 = 2.41^{+0.18}_{-0.27}$  and  $\Gamma_2 = 0.44^{+0.79}_{-0.85}$  in agreement with Perri et al. (2003). Adopting this



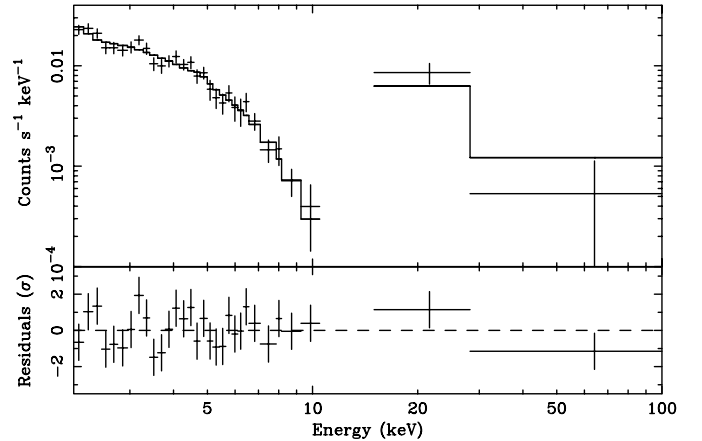
**Fig. 1.** MECS image of the sky region surrounding ON 325.

model, we estimate a contribution of ON 325 to the PDS data of the order of 10% (see Table 2).

The MECS spectra of 1AXG J121854+2957 is instead well fitted ( $\chi^2/\nu = 10.8/14$ ) with an absorbed power law having a photon index  $\Gamma \sim 2.3$ , a column density  $N_{\text{H}} < 12 \times 10^{22} \text{ cm}^{-2}$  and a 2–10 keV flux of  $8.5 \times 10^{-13} \text{ erg cm}^{-2} \text{ s}^{-1}$ . These values are in agreement with those found by Loaring et al. (2003) during a *XMM-Newton* observation of the source. Extrapolation of this spectrum to the PDS band indicates a small contribution to the high energy flux, while a simultaneous fit to the MECS/PDS data provides a PDS/MECS cross-calibration constant in the range 2–30, i.e. well outside the nominal interval. A good fit ( $\chi^2/\nu = 30/34$ ) of the MECS data of MKN 766 is also described by a simple power law with a photon index  $\Gamma = 1.92 \pm 0.13$  and a 2–10 keV flux of  $1.7 \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$ . The high contribution to the PDS flux (see Table 2) found with this model indicates that MKN 766 is the best candidate to account for the PDS flux. In fact, the fit of the MECS spectrum of MKN 766 with the PDS data gives a PDS/MECS cross-calibration constant well inside the nominal interval. The 2–10 keV flux (see Table 3), found during this pointing, is slightly lower than the value of  $2.05 \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$  found by Matt et al. (2000a) during a previous (May 1997) *BeppoSAX* dedicated observation. We find no evidence for an iron line at around 6.4 keV (see Fig. 2), consistent with the fact that this feature seems to be present only when MKN 766 is in a high state (Leighly et al. 1996).

To confirm if MKN 766 is really the contaminating source in the ON 325 field, we also reanalysed the PDS data of the two pointings of MKN 766 performed by *BeppoSAX* in May 1997 and 2001. In the first observation the comparison between the +OFF and –OFF field count rates shows an excess of  $3\sigma$  in the –OFF field, while in the second there is no evidence for contamination.

Searching in the HEASARC archive for possible objects which could contaminate the –OFF field, we found two



**Fig. 2.** Broad band spectrum of MKN 766 fitted with a simple power law (upper panel) and residuals to this model in units of  $\sigma$  (lower panel). The data are graphically rebinned.

sources belonging to the *ROSAT* Bright Source Catalog: a quasar (1RSX J121320.3+270841) and an unidentified object (1RSX J121258.3+272653), at a distance of  $\sim 40'$  and  $\sim 59'$  from the field centre. Fitting the –OFF data with a simple power law, we obtain  $\Gamma \leq 2$  and a 20–100 keV flux of  $3.6 \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$ . Since both sources have similar count rates in *ROSAT*, we can only state that either or both could be responsible for this excess.

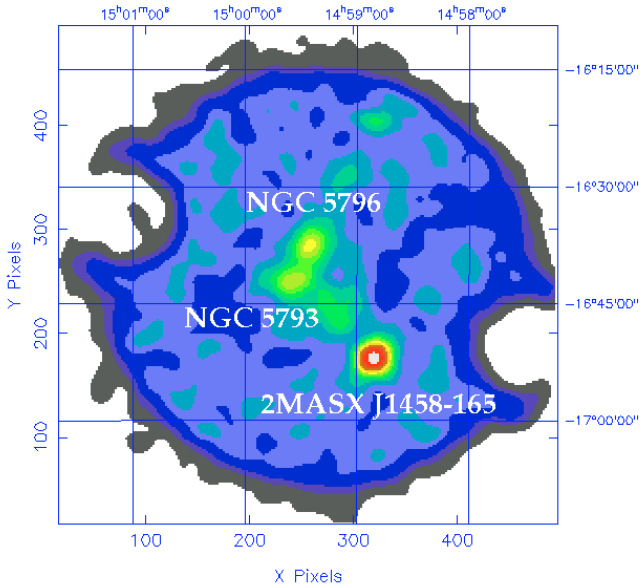
This provides a good example of the capability of our analysis: contamination in the offset fields can be found and identified, while at the same time a more correct evaluation of the target spectrum is provided.

The PDS spectrum of MKN 766 in the May 1997 observation can be corrected and reanalysed together with that of May 2001. In both cases the photon index  $\Gamma$  is  $\sim 2$ , in agreement with that found in December 1998 (i.e. during the ON 325 observation), while the 20–100 keV flux is roughly  $\sim 3 \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$ , higher than during the ON 325 measurement; this finding indicates that the flux variability always seen in this source at soft X-ray energies (Matt et al. 2000a and references therein) extends to above  $\sim 10 \text{ keV}$ .

#### 4.1.2. The NGC 5793 field

NGC 5793 has been detected at  $\sim 5\sigma$  level in the MECS instrument, while it is insignificant ( $< 1\sigma$ ) in the LECS energy range. Fitting the MECS data with a simple power law of photon index  $\Gamma = 1.9$  we obtain a very low flux of  $\sim 7 \times 10^{-14} \text{ erg cm}^{-2} \text{ s}^{-1}$ . This model, extrapolated to the high energy band, gives a negligible contribution to the PDS data (see Table 2).

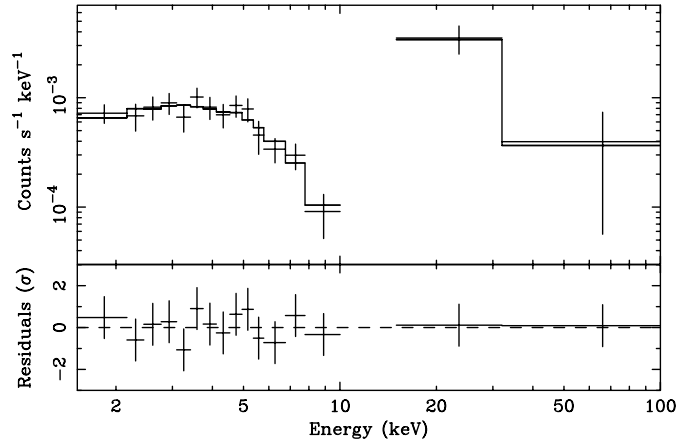
Inspection of Fig. 3 indicates the presence of two extra sources in the MECS field of view. The first located at  $\alpha(2000) = 14^{\text{h}}58^{\text{m}}50^{\text{s}}.7$  and  $\delta(2000) = -16^{\circ}51'50''.7$ , corresponds to a bright irregular spiral galaxy belonging to the 2 Micron All Sky Survey eXtended Source Catalog (i.e. 2MASX J14585116–1652223). The second is located at  $\alpha(2000) = 14^{\text{h}}59^{\text{m}}23^{\text{s}}.8$  and  $\delta(2000) = -16^{\circ}37'13''.4$  and is identified with the normal galaxy NGC 5796.



**Fig. 3.** MECS image of the sky region surrounding NGC 5793.

We extracted the MECS spectra of both sources and extrapolated them to high energies. We find that the MECS spectrum of NGC 5796 is well fitted ( $\chi^2/\nu = 15/20$ ) with a simple power law having a photon index  $\Gamma = 1.95^{+0.67}_{-0.56}$  and a 2–10 keV flux of  $1.5 \times 10^{-13}$  erg cm $^{-2}$  s $^{-1}$ . The contribution of NGC 5796 to the PDS flux is only a few percent and the PDS/MECS cross-calibration constant is extremely high (20–103).

The MECS best-fit model in the case of 2MASX J14585116–1652223 is a simple power law with photon index  $\Gamma = 1.15^{+0.33}_{-0.18}$  and a 2–10 keV flux of  $9.3 \times 10^{-13}$  erg cm $^{-2}$  s $^{-1}$  ( $\chi^2/\nu = 15/20$ ). The extrapolation of this model to high energies (see Table 2) suggests that 2MASX J14585116–1652223 is likely to be the contaminating object. The simultaneous fit of the MECS/PDS data (see Table 3 and Fig. 4) shows a flat spectrum ( $\Gamma = 1.20^{+0.26}_{-0.23}$ ), a more suitable value (0.7–3.5) for the PDS/MECS cross-calibration constant and no evidence for extra absorption. Fixing  $\Gamma = 1.9$  and adding absorption to the power law provides a similar fit and an upper limit to the column density of  $\sim 4 \times 10^{22}$  cm $^{-2}$ . As displayed in Table 3, the 20–100 keV flux turns out to be of  $1.11 \times 10^{-11}$  erg cm $^{-2}$  s $^{-1}$ , one order of magnitude higher than the 2–10 keV flux, confirming that 2MASX J14585116–1652223 emits mostly in the high energy band. Very little is known about this object except for its near infrared properties: the source is fairly bright with a total photometry of 14.4, 13.6 and 13 mag in the *J*, *H* and *K* bands respectively and an extent of 11″; the optical counterpart has magnitudes *B* = 16 and *R* = 15. No previous X-ray data are reported, including the lack of detection by *ROSAT*, which would be consistent with the presence of strong absorption. The *R* – *K* colour is  $\sim 2$ , i.e. quite red which again suggests an object with considerable extinction. No detection in the radio band was found in the literature or in the HEASARC archive.



**Fig. 4.** Broad band spectrum of 2MASX J1458–165 fitted with a simple power law (*upper panel*) and residuals to this model in units of  $\sigma$  (*lower panel*). The data are graphically rebinned.

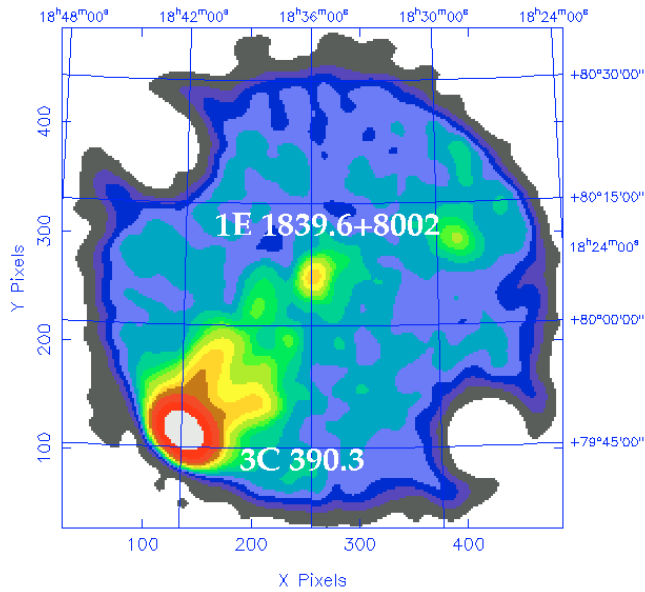
#### 4.1.3. The 1E 1839.6+8002 field

The M4 Ve star 1E 1839.6+8002 was observed by *BeppoSAX* twice (October 2000 and February 2001). This source was detected for the first time during the *Einstein* observation of the Broad Line Radio Galaxy 3C 390.3 in the 1980s. A detailed study of the spectral behaviour of this source is beyond the aim of this paper. Nevertheless, following Pan et al. (1997), the LECS and MECS spectra of both observations are well described by a thermal model (RAYMOND–SMITH model in XSPEC, Raymond & Smith 1977), with a temperature  $kT \sim 1.8$  keV and a 2–10 keV flux of  $9.5 \times 10^{-14}$  erg cm $^{-2}$  s $^{-1}$  (first observation) and  $7.2 \times 10^{-14}$  erg cm $^{-2}$  s $^{-1}$  (second observation).

The source appears to be faint and variable ( $\sim 25\%$ ) in the 2–10 keV range, with a negligible contribution to the high energy emission in both observations (in Table 2 only the averaged contribution of the source is reported).

Within the MECS field of view (see Fig. 5) of both measurements, a bright object located at  $\alpha(2000) = 18^{\text{h}}41^{\text{m}}51^{\text{s}}.5$  and  $\delta(2000) = +79^{\circ}47'01''.3$  is clearly visible: it corresponds to the well-known radio galaxy 3C 390.3. The *Einstein* IPC data of 3C 390.3 revealed the presence of strong intrinsic absorption (Kruiper et al. 1990), while at higher energies *Ginga* data showed the presence of an iron line at 6.4 keV (Inda et al. 1994) and a reflection component (Nandra & Pounds 1994). *ASCA* data confirmed the presence of the iron *K* emission line but could not constrain the reflection hump, most probably because of the limited energy range (Eracleous et al. 1996; Leighly et al. 1997). 3C 390.3 has also been detected by *OSSE* in the soft- $\gamma$  energy range, above 50 keV (Dermer & Gherels 1995). All these results have been confirmed by Grandi et al. (1999), on the basis of a dedicated *BeppoSAX* observation performed in January 1997. The above authors also found a column density variability when comparing *BeppoSAX* results with previous measurements.

No evidence for variability is present in the MECS data of 3C 390.3 nor in the PDS analysed in this work. Hence, in order



**Fig. 5.** MECS image of the sky region surrounding 1E 1839.6+8002.

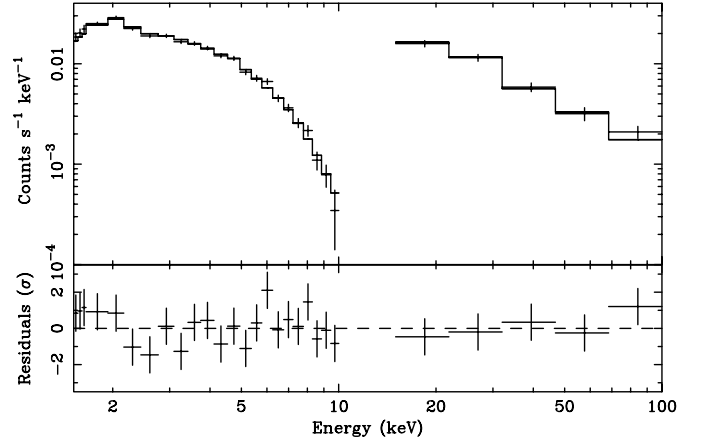
to improve the statistics, we summed the two data sets to perform our spectral analysis. First, we analysed only the MECS data of 3C 390.3, finding as best-fit model ( $\chi^2/\nu = 13.5/17$ ) a power law with a slope of 1.7 plus a narrow Gaussian line having a centroid energy of 6.4 keV and an equivalent width of  $EW = 115$  eV, compatible with  $K_\alpha$  line emission from neutral iron located at the source redshift. The contribution of the 3C 390.3 flux to the total PDS flux is significant (see Table 2) and clearly indicates that this source is indeed responsible for the contamination.

The joint fit of the MECS data of this source with the PDS is also well reproduced by a power law plus a narrow Gaussian line ( $\chi^2/\nu = 16.5/22$ ) having parameters similar to those found fitting only the MECS data. As for the column density, we can only provide an upper limit of  $\sim 8 \times 10^{20} \text{ cm}^{-2}$ , compatible with the values found and reported by Grandi et al. (1999). The broad band spectrum of the source is shown in Fig. 6. The flux turns out to be  $2.2 \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$  and  $4.3 \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$  in the 2–10 and 20–100 keV band respectively, again in agreement with that found by Grandi et al. (1999). The PDS/MECS cross-calibration constant turns out to be consistent (see Table 4) with the suggested values, confirming that 3C 390.3 is the serendipitous source contaminating the PDS emission during the observations of the 1E 1839.6+8002 field of view (see Table 2). There is no evidence for variability between the present (2000–2001) *BeppoSAX* observations and that reported by Grandi et al. (1999) in the 1997 observation.

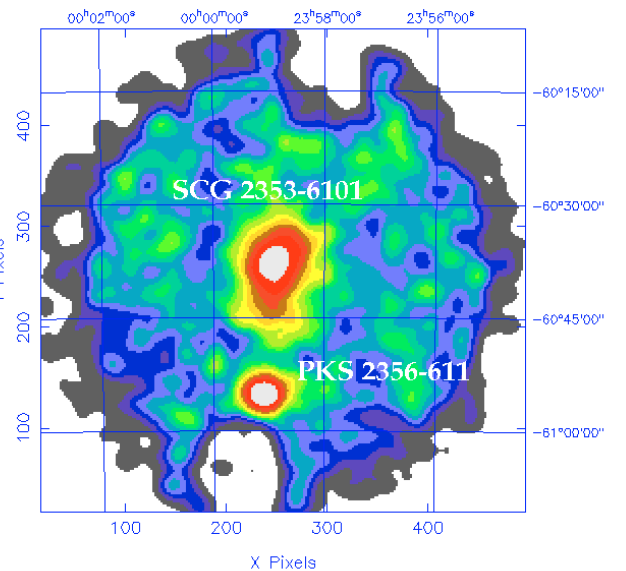
#### 4.1.4. The SCG 2353–6101 field

In the case of the cluster of galaxies SCG 2353–6101 (Abell 4067) there is another source in the MECS field of view, PKS 2356–611, located at  $\alpha(2000) = 23^{\text{h}}59^{\text{m}}04^{\text{s}}.1$  and  $\delta(2000) = -60^{\circ}55'01''.5$  (see Fig. 7).

PKS 2356–611 ( $z = 0.096$ ,  $B = 12.3$ ,  $R = 12.8$ , as reported in the USNO–B1 Catalog (Monet et al. 2003)) is one of



**Fig. 6.** Broad band spectrum of 3C 390.3 fitted with an absorbed power law plus a narrow Gaussian line (*upper panel*) and residuals to this model in units of  $\sigma$  (*lower panel*).

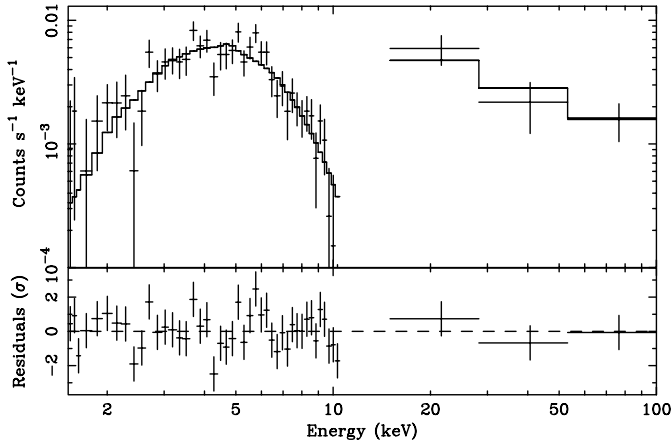


**Fig. 7.** MECS image of the sky region surrounding SCG 2353–6101.

the strongest southern FR II sources, with a total radio power  $P_{1.4 \text{ GHz}} \sim 10^{25.8} \text{ W Hz}^{-1}$  (Koekemoer et al. 1998), while optically it shows strong high-excitation narrow line emission with  $[\text{O III}]\lambda 5007/\text{H}\beta > 10$ . Lipovetsky et al. (1988) have cataloged this source as a Seyfert galaxy of type 2. It is also listed in the 2 Micron All Sky Survey eXtended Source Catalog as 2MASX J23590436–6054594 (with magnitudes  $J = 13.7$ ,  $H = 12.7$ ,  $K = 12.9$ ) and was not detected by *ROSAT*.

The LECS and MECS spectra of SCG 2353–6101 are well fitted ( $\chi^2/\nu = 32/37$ ) with a thermal bremsstrahlung model, giving a temperature of  $4.74^{+1.70}_{-2.26} \text{ keV}$ . The extrapolation of this model to the PDS energy band provides negligible flux (see Table 2). This, combined with the high value (8–28) of the PDS/MECS cross-calibration, indicates the presence of a high energy emitting source and PKS 2356–611 is a good candidate.

As shown in Fig. 7, PKS 2356–611 is located near one of the two  $^{55}\text{Fe}$  calibration sources of the MECS instrument. In this case a careful choice of the background region is required



**Fig. 8.** Broad band spectrum of PKS 2356–611 fitted with an absorbed power law (*upper panel*) and residuals to this model in units of  $\sigma$  (*lower panel*).

to avoid contamination by 5.95 keV photons produced by the calibration source: this is done taking the background region as near as possible to the calibration sources itself.

Fitting the MECS data of PKS 2356–611 and the PDS points with a single power law plus intrinsic absorption provides an acceptable fit ( $\chi^2/\nu = 60/43$ ), a steep spectrum ( $\Gamma = 1.75$ ) and a column density of  $N_{\text{H}} \sim 1 \times 10^{23} \text{ cm}^{-2}$  (Fig. 8). Both spectral parameters are typical of Seyfert 2 galaxies, and seem to confirm the source classification made by Lipovetsky et al. (1988).

The cross-calibration constant in this case turns out to be lower (0.2–0.5) than generally observed. This could be due to imperfect background correction in the MECS because of the location of the source near the calibrator. However, the spectral parameters do not change significantly when this constant is constrained to vary within the nominal range of values.

The broad band spectrum of PKS 2356–611 (see Fig. 8) suggests the possible presence of an excess around  $\sim 6$  keV. Although the addition of an extra component in the form of a narrow Gaussian line at around 6.4 keV with  $EW \sim 500$  eV provides an improvement in the fit ( $\Delta\chi^2 = 11.5$  for two additional parameters), its reality is questioned by the location of the source near the  $^{55}\text{Fe}$  calibrator: residuals at this line energy could still be due to a non-perfect subtraction of the calibration line. Also in this case the cross-calibration constant is low ( $< 0.5$ ) and the  $\chi^2$  values as well as the spectral parameters do not change significantly if we vary the cross-calibration constant within the allowed range. The absorption measured would produce an iron line of 100–200 eV in transmission, while reflection in the torus would increase the  $EW$  to about 400 eV (Turner et al. 1997) possibly indicating the presence of a reflection component. However, inclusion in the fit of this component via the PEXRAV model in XSPEC (Magdziarz & Zdziarski 1995) is not required by the data and does not improve the cross-calibration constant, which turns out to be very small ( $< 0.4$ ). A large  $EW$  is also possible if the source is Compton thick; this type of object is generally characterized by a low  $F_{\text{X}}/F_{\text{OIII}}$  ratio (Bassani et al. 1999). In PKS 2356–611 this ratio is  $\sim 67$ , when  $F_{\text{OIII}}$  is corrected

for the reddening in the galaxy using the observed value of  $H_{\alpha}/H_{\beta} = 4.32$  (Koekemoer et al. 1998). This result suggests that PKS 2356–611 is a Compton thin Seyfert 2 galaxy rather than a Compton thick one and so the large  $EW$  is probably not due to strong absorption ( $N_{\text{H}} > 1.5 \times 10^{24} \text{ cm}^{-2}$ ).

#### 4.2. Sources not detected in the MECS field of view

In this section we describe the remaining eight fields where the source responsible for the PDS spectrum is not observed by the MECS. In Table 5 we report for each of these fields the source/s observed in the MECS, their flux extrapolated to the 20–100 keV band and their contribution to the PDS flux. In Table 6 instead we list for each field the PDS exposure and count rate as well as the photon index  $\Gamma$  and observed 20–100 keV flux obtained assuming a simple power law fit. We also report the effective flux at the source, estimated from that observed by applying the correction factor  $R$ .

##### 4.2.1. The IRAS 01025–6423 field

In this field the target source, IRAS 01025–6423, a Seyfert 2 galaxy, is barely detected by MECS at a  $\sim 2\sigma$  level, while the LECS data are not available. Due to the low signal to noise ratio, we only attempted to put an upper limit on the flux assuming a power law model with photon index  $\Gamma = 2$ . The source appears to be faint in the 2–10 keV range with a flux  $\leq 3.1 \times 10^{-14} \text{ erg cm}^{-2} \text{ s}^{-1}$  and its contribution to the PDS signal is obviously null (see Table 5). Searching in the HEASARC archive for possible contaminating sources within the PDS field of view, we found two possible candidates belonging to the *ROSAT* Bright Source Catalog: 1RXS J011316.2–641142 (CPD–64 120), classified as a K1 Ve star, and 1RXS J010333.7–643925 (PKS 0101–649), identified as a radio source, at distances from the target source of  $\sim 58'$  and  $\sim 33'$ , respectively. PKS 0101–649, which is the closest to the target source, is more likely to be the object responsible for the high energy emission, as a star is not expected to emit above 10 keV. PKS 0101–649 belongs to the 2 Micron All Sky Survey eXtended object (i.e. 2MASX J01033376–6439079 in NED), with a total photometry of  $J = 15$ ,  $H = 14$  and  $K = 13.4$  mag; the USNO–A2 R mag is 16.5 (Monet et al. 1999) thus providing an  $R-K$  value of 3.1 which suggests a quite red object; at radio frequencies it is fairly bright with a 6 cm flux of  $\sim 270$  mJy.

The PDS spectrum is characterized by a photon index  $\Gamma \sim 1.5$  (see Table 6) compatible, within uncertainties, with an AGN spectrum; the 20–100 keV flux is  $2.3 \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$ , making PKS 0101–649 a bright source at these energies. If we extrapolate this power law to soft X-ray energies, we find a 0.1–2.4 keV flux of  $0.22 \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$  which is consistent with the *ROSAT* value of PKS 0101–649. Alternatively, the X-ray source spectrum could be steeper but highly absorbed even above 10 keV as observed in Compton thick objects (Matt et al. 2000b). If we add absorption to the power law, fixing the photon index to 1.9, we find an upper limit to the column density  $< 2 \times 10^{25} \text{ cm}^{-2}$ . Obviously, an optical spectrum

**Table 6.** PDS spectral analysis of serendipitous sources not detected in the MECS instrument.

Source	Offset <sup>(a)</sup> (arcmin)	EXPO (ks)	cts/s	$\Gamma$	$R^{(b)}$ (%)	$F_{\text{Obs}}^{(c)}$ [20–100] keV	$F_{\text{Source}}^{(c)}$ [20–100] keV
PKS 0101–649	32.5	26.4	$0.115 \pm 0.035$	$1.51^{+1.45}_{-1.03}$	56	12.8	23.0
Perseus/NGC 1275 (?)	62.0	25.9	$0.438 \pm 0.043$	$3.27^{+0.53}_{-0.43}$	19	25.0	136.0
NGC 1566 <sup>(d)</sup>	60.2	26.6	$0.192 \pm 0.041$	$2.50^{+1.32}_{-0.84}$	19	15.4	81.1
NGC 3227 <sup>(d)</sup>	55.1	176.0	$0.247 \pm 0.016$	$1.85^{+0.24}_{-0.22}$	26	25.5	98.1
ESO 025–G002	31.4	26.1	$0.138 \pm 0.035$	$2.96^{+0.91}_{-0.73}$	55	11.1	20.2
4C +74.26 <sup>(d)</sup>	34.2	66.2	$0.250 \pm 0.024$	$2.03^{+0.37}_{-0.31}$	53	24.4	46.0
NGC 7319	33.7	29.2	$0.137 \pm 0.031$	$2.07^{+0.67}_{-0.53}$	59	14.1	24.0
NGC 7582	27.6	56.8	$0.490 \pm 0.027$	$1.73^{+0.18}_{-0.16}$	61	49.6	81.3

Note: Count rates are referred to the energy bands used in the spectral analysis (see text).

<sup>(a)</sup> Distance between the pointing target and the off-axis sources.

<sup>(b)</sup> Source flux correction factor due to the PDS response to sources not on-axis.

<sup>(c)</sup> In units of  $10^{-12}$  erg cm<sup>-2</sup> s<sup>-1</sup>.

<sup>(d)</sup> For this source the parameters of the average spectrum are reported.

would be highly desirable to provide more information and establish the true nature of this source.

#### 4.2.2. The MKN 1073 field

The Seyfert 2 galaxy MKN 1073 was not clearly detected by the LECS and the MECS instruments. In both detections there is evidence for diffuse emission probably associated with the Perseus cluster, which is at 62' from the centre of the MECS pointing. The signal in the PDS is high ( $>10\sigma$  level), with an exceptionally steep spectrum ( $\Gamma \sim 3.3$ , see Table 6). We ascribe the PDS emission to the Perseus cluster, which has also been targetted by *BeppoSAX* on September 1996. The spectral analysis of these PDS data also shows a very steep spectrum ( $\Gamma = 3.80^{+0.17}_{-0.18}$ ) and a 20–100 keV flux of  $\sim 5 \times 10^{-11}$  erg cm<sup>-2</sup> s<sup>-1</sup>, lower than the flux seen during the MKN 1073 observation (see Table 6). Since flux variability in the cluster is unlikely, the change in the emission could be due to the Seyfert 2 galaxy NGC 1275, often taken as responsible for the high energy ( $>20$  keV) emission observed from this region; comparison of data from high energy instruments over years shows a large variation in the flux above 20 keV (Osako et al. 1994). It is also possible that both cluster and AGN contribute to the high energy emission, but it is difficult at present to disentangle one contribution from the other.

#### 4.2.3. The NGC 1553 field

*Chandra* observations of the X-ray faint SO galaxy NGC 1553 (Blanton et al. 2001) has spatially and spectrally resolved the source of the X-ray emission. A significant fraction of this ( $\sim 70\%$ ) is detected as diffuse flux while the remainder is due to 49 objects. The strongest source in the field is located at the centre of NGC 1553 to within 0'5 and shows a hard spectrum, typical of an AGN. NGC 1553 was observed twice by *BeppoSAX* (January 1997 and November 1997). Trinchieri et al. (2000) performed the spectral analysis of the LECS and MECS data finding as best-fit a RAYMOND–SMITH model ( $kT \approx 0.26$  keV) for the soft component and a

thermal bremsstrahlung ( $kT \approx 4.8$  keV) for the hard component. Our analysis of the LECS and MECS data from both observations is in agreement with that reported by these authors; the extrapolation of this model to high energies provides a very low contribution to the PDS flux (see Table 5). Searching in past X-ray mission archives, we find that the Seyfert 1 galaxy NGC 1566 is the best candidate for the PDS emission as also suggested by Trinchieri et al. (2000). This source is at a distance of 60'2 from the target object and is therefore just inside the PDS field of view. NGC 1566 ( $z = 0.005$ ,  $B = 9.9$ ,  $R = 9.6$ ) is a very bright X-ray source known since the *HEAO 1* X-ray Source Catalog (Wood et al. 1984). The source belongs also to the 2 Micron All Sky Survey eXtended Source Catalog (2MASX J04200041–5456161), with magnitudes  $J = 7.8$ ,  $B = 7.2$  and  $K = 6.9$ , it is reported by *ROSAT* in the Bright Source Catalog and it is a radio source of  $\sim 100$  mJy at 6 cm. The 2–10 keV flux reported in the literature ranges from  $<0.5$  to  $1.7 \times 10^{-11}$  erg cm<sup>-2</sup> s<sup>-1</sup> (Halpern 1982) indicating variability by around a factor of 3. The PDS emission shows a variation of  $\sim 70\%$  during our two observations, but due to the poor statistics in both measurements (the signal to noise ratio in the PDS is around  $3\sigma$  level in each measurement), we summed the observations to improve the statistics and performed an averaged spectral analysis. As shown in Table 6, this analysis indicates a photon index  $\Gamma$  of  $\sim 2.5$  and an average 20–100 keV flux, corrected for the offset, of  $\sim 8 \times 10^{-11}$  erg cm<sup>-2</sup> s<sup>-1</sup>. By extrapolating the flux down to the 2–10 keV range, we find a value of  $\sim 2 \times 10^{-10}$  erg cm<sup>-2</sup> s<sup>-1</sup>, much higher with respect to previous measurements; however, if the photon index is restricted to values more appropriate for a Seyfert 1 galaxy (i.e. 1.7–1.9), then the 2–10 keV flux reduces to  $\sim 6 \times 10^{-11}$  erg cm<sup>-2</sup> s<sup>-1</sup>, only a factor of 3 higher. This result indicates that NGC 1566 could indeed dominate the PDS emission and that flux variability is likely.

#### 4.2.4. The AD Leonis field

The M-dwarf Ad Leonis was the subject of an observational campaign performed by *BeppoSAX* (April 1997, 1, 8

and 12 May 1999). We adopt the multi-temperature model (MEKAL code in XSPEC, Mewe et al. 1985) proposed by van den Besselaar et al. (2003) for the *XMM-Newton* and *Chandra* observations, and find that the LECS and MECS data of the four observations are best described by a two-temperature model, in which the metal abundances are left free to vary. The 2–10 keV flux varies in the range  $1.4\text{--}2.3 \times 10^{-12}$  erg cm $^{-2}$  s $^{-1}$ , corresponding to a flux variation of  $\sim 60\%$ , while the spectral parameters do not show significant variation. In view of these indications, we summed the data from different pointings and estimated the average spectral parameters. The two-temperature model remains the best fit ( $\chi^2/\nu = 97/82$ ), giving  $kT_1 = 0.65^{+0.03}_{-0.02}$  keV and  $kT_2 = 2.04^{+0.24}_{-0.17}$  keV metal abundances of  $0.14 \pm 0.02$  and  $0.56^{+0.25}_{-0.21}$  respectively and an averaged 2–10 keV flux of  $\sim 2 \times 10^{-12}$  erg cm $^{-2}$  s $^{-1}$ ; this spectrum if extrapolated to high energies provides almost no contribution to the PDS flux (see Table 5).

At a distance of  $\sim 55'$  from AD Leonis, we find the Seyfert 1 galaxy NGC 3227, which could be the contaminating source we are searching for. NGC 3227, first detected in the *Ariel V* all sky survey, is one of the bright sources in the Piccinotti sample of AGN (Piccinotti et al. 1982). NGC 3227 is known to be variable (on time scales of hours–days) and characterized by a spectrum with a photon index  $\Gamma \sim 1.6$ , flatter than that typically observed in Seyfert galaxies (George et al. 1998). *Ginga* observations performed in 1988 indicate a 2–10 keV flux of  $\sim 4 \times 10^{-11}$  erg cm $^{-2}$  s $^{-1}$  (Pounds et al. 1989), while *ASCA* observations performed during 1993 and 1995 provide 2–10 keV flux measurements in the range  $2.4\text{--}2.6 \times 10^{-11}$  erg cm $^{-2}$  s $^{-1}$ . Recently, Lamer et al. (2003), analyzing the *Rossi X-ray Timing Explorer (RXTE)* data of NGC 3227, confirmed the source variability, explaining it in terms of transient absorption by a gas cloud (neutral or weakly ionized) of column density of  $\sim 3 \times 10^{23}$  cm $^{-2}$  moving across the line of sight to the X-ray source. This interpretation has been confirmed also by *XMM-Newton* observations (Gondoin et al. 2003). NGC 3227 has been detected at higher energies by *OSSE* and *BATSE* on-board the *Compton Gamma Ray Observatory (CGRO)*. The *OSSE* data when fitted in the 50–200 keV range, show a spectral index  $\Gamma = 1.86^{+0.36}_{-0.38}$ , softer with respect to that found in the 2–10 keV band and a photon flux of  $(3.40 \pm 0.55) \times 10^{-4}$  cm $^{-2}$  s $^{-1}$  in the 50–150 keV band (Zdziarski et al. 2000). Based on these measurements, we estimate a 20–100 keV flux measurement in the range  $5.0\text{--}7.8 \times 10^{-11}$  erg cm $^{-2}$  s $^{-1}$ . In addition, the *BATSE* data (Malizia et al. 1999) show a 20–100 keV flux in the range  $9.7\text{--}12.8 \times 10^{-11}$  erg cm $^{-2}$  s $^{-1}$ . The spectral analysis of our PDS data indicate that, during the four observations of AD Leonis, the 20–100 keV flux measurements of NGC 3227 ranged in the interval  $6.3\text{--}14.7 \times 10^{-11}$  erg cm $^{-2}$  s $^{-1}$ , with an averaged value of  $9.8 \times 10^{-11}$  erg cm $^{-2}$  s $^{-1}$  and a photon index  $\Gamma \sim 1.85$  (see Table 6). These results turn out to be in perfect agreement with both *OSSE* and *BATSE* data. We want to stress that our measurement contains the only spectral information about NGC 3227 obtained by *BeppoSAX* (the source has never been observed by the satellite) and so represent an interesting addition to the list of Piccinotti sample sources for which *BeppoSAX* high energy data are available.

#### 4.2.5. The H1846–786 field

H1846–786 is a Seyfert 1 galaxy belonging to the Piccinotti sample (Piccinotti et al. 1982). Despite its X-ray brightness, the source has been poorly studied. A simple power law, with a photon index  $\Gamma \sim 1.95$  is the best-fit model to the LECS and MECS data ( $\chi^2/\nu = 125/100$ ) as obtained by Quadrelli et al. (2003). This result indicates that no extra absorption and/or an iron line are present in the spectrum. The extrapolation of this best-fit model to the PDS energy range shows a significant contribution from H1846–786 (see Table 5). In fact, as can be seen in Table 1, this case represents the lowest range for the PDS/MECS normalization constant, indicating that both the target (H1846–786) and another object, located within the PDS field of view, contribute to the PDS flux. This excess emission could be due to the near ( $z = 0.0285$ ,  $B = 8.8$ ,  $R = 10.4$ ) bright Seyfert 1 galaxy ESO 025–G002, which has a *ROSAT* flux only 5 times lower than the target source. Also this source belongs to the 2 Micron All Sky Survey eXtended Source Catalog (i.e. 2MASX J18544039–7853544), with magnitudes  $J = 11.4$ ,  $H = 10.6$  and  $K = 10.3$ , and shows weak radio emission ( $\sim 10$  mJy at 36 cm). The spectral analysis of the PDS data yields a contribution from this object of  $\sim 60\%$ . The spectrum is steep ( $\Gamma \sim 3$ ) but still compatible with the canonical AGN value of 2 and the 20–100 keV flux turns out to be of the order of  $2.0 \times 10^{-11}$  erg cm $^{-2}$  s $^{-1}$  (see Table 6 for more details). Thus, ESO 025–G002 shows at high energies a flux of the same order of that of H1846–786, while at soft energies is 5 time lower. This can be explained if the source is absorbed at soft energies; in fact, fixing the PDS power law index to a value more appropriate to an AGN and allowing for absorption, we find an upper limit to the column density of a few  $10^{22}$  cm $^{-2}$ .

Finally, we can conclude that in this particular case both target and serendipitous source are likely to give roughly the same contribution to the PDS emission.

#### 4.2.6. The VW Cephei field

VW Cephei, a W UMa-type binary system, has been observed by *BeppoSAX* twice (May 1998 and October 1998). Following the investigations performed by Gondoin (2004), we fit separately the LECS and MECS data of both data sets with the MEKAL optically thin plasma model. In each observation, a three-component model, having different temperatures but the same metal abundance, gives the best representation of the data. The 2–10 keV flux turns out to be  $1.6 \times 10^{-12}$  erg cm $^{-2}$  s $^{-1}$  and  $0.8 \times 10^{-12}$  erg cm $^{-2}$  s $^{-1}$  in the first and second pointings, respectively. Although the flux varies ( $\sim 50\%$ ) from the first to the second observation, the spectral shape shows, within the uncertainties, no evidence of variability; therefore, we summed the two data sets. The three-temperature model applied to the combined data still provides a good fit ( $\chi^2/\nu = 37/65$ ) and gives temperature values of 2.4, 0.69 and 0.13 keV; extrapolation of this model to the PDS band indicates that the VW Cephei contribution to the PDS data is negligible (see Table 5).

As a possible contaminating source, we find at around  $34'$  from VW Cephei an object identified with 4C +74.26, a radio-loud active galaxy. This is a particularly interesting

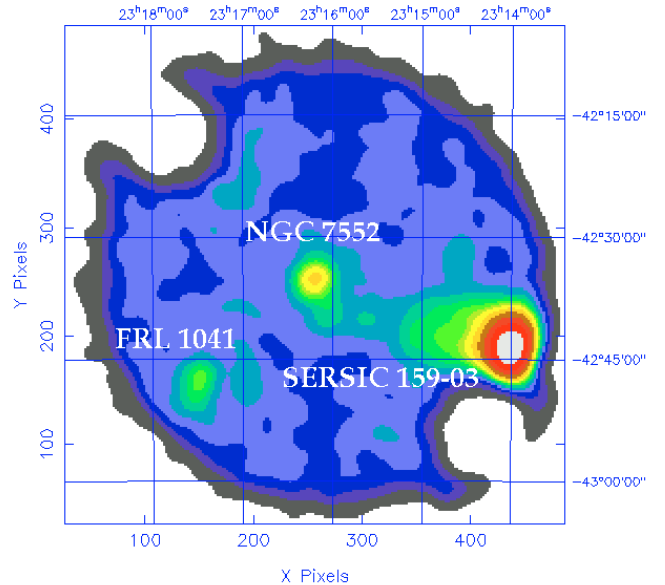
object because its *ASCA* X-ray spectrum (Brinkmann et al. 1998; Sambruna et al. 1999) shows features that are typical of Seyfert galaxies more than of Broad Line Radio Galaxies: a warm absorber and a significant Compton reflection hump. It is also the only quasar in the collection of Sambruna et al. (1999) with a detectable  $\text{Fe}_{\text{K}\alpha}$  line. 4C +74.26 was targetted by *BeppoSAX* on May 1999. Analysis of these data by Hasenkopf et al. (2002) indicated as the best-fit model a power law continuum modified by Compton reflection at high energies and by absorption at low energies. An iron line was also detected, but its energy could not be tightly constrained, falling between 6.4 and 6.9 keV. Our spectral analysis of the average PDS emission obtained by combining the two observations of VW Cephei yields a photon index  $\Gamma \sim 2$  and a 20–100 keV flux, corrected for the off-axis effect, of  $4.6 \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$  (see Table 6). If we fit the PDS spectrum with the *PEXRAV* model, fixing the inclination angle to the value indicated by Hasenkopf et al. (2002), we find a satisfactory fit ( $\chi^2/\nu = 1.2/2$ ), but the reflection coefficient of  $R \sim 1.3$  and the cut-off energy  $E_C > 148 \text{ keV}$  are not well constrained, although in agreement with what reported by Hasenkopf et al. (2002). In order to check the self-consistency of our spectral analysis, we reanalysed the PDS data of the May 1999 observation with a simple power law and obtained spectral parameters for this epoch (a photon index  $\Gamma \sim 2$  and a 20–100 keV flux of  $4.0 \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$ ) consistent with those obtained during the VW Cephei pointings.

Overall, the data indicate that 4C +74.26 is the most likely contaminating source present in the PDS field of view of VW Cephei and further suggest no variability in the high energy flux.

#### 4.2.7. The NGC 7331 field

A good fit to the MECS and LECS data of the LINER NGC 7331 is provided by a single power law having a photon index  $\Gamma = 1.95^{+0.12}_{-0.18}$  and 2–10 keV flux of  $5.1 \times 10^{-13} \text{ erg cm}^{-2} \text{ s}^{-1}$ . This result confirms the non-thermal X-ray spectrum of NGC 7331 reported by Stockdale et al. (1998) on the basis of both *ROSAT* and radio observations. The above model however gives a small contribution (see Table 5) to the PDS emission when extrapolated to high energies.

Searching for potential high energy emitters in the HEASARC archives, we find a source located at about 30' from the centre of the MECS pointing and identified with the Seyfert 2 galaxy NGC 7319, which belongs to the Stephan's Quintet, a compact group of galaxies. Recent *Chandra* data of NGC 7319 (Trinchieri et al. 2003) describe the source emission as due to the superposition of a strong and heavy absorbed nuclear source plus diffuse softer emission. Modeling the nuclear source with a combination of absorbed and unabsorbed power law components having the same photon index, plus a narrow 6.4 keV emission line, the above authors found a photon index  $\Gamma \sim 1.7$ , a column density  $N_{\text{H}} = 4 \times 10^{23} \text{ cm}^{-2}$ , an  $EW \sim 110 \text{ eV}$  and a 2–10 keV flux of  $8.2 \times 10^{-12} \text{ erg cm}^{-2} \text{ s}^{-1}$ . The spectral fit was similar to that derived from a previous *ASCA* observation (Awaki et al. 1997), except for the smaller  $EW$  and higher 2–10 keV flux. The X-ray spectral parameters



**Fig. 9.** MECS image of the sky region surrounding NGC 7552.

as well as the  $F_X/F_{\text{OIII}}$  ratio (Bassani et al. 1999) indicate that the source is likely to be Compton thin. NGC 7319 ( $z = 0.0225$ ,  $B = 14$ ,  $R = 12.7$ ) belongs to the 2 Micron All Sky Survey extended Source Catalog (2MASX J22360355+3358327), with magnitudes  $J = 11.1$ ,  $H = 10.3$  and  $K = 10.1$ , and is also detected in radio ( $\sim 7 \text{ mJy}$  at 6 cm).

The PDS data are well described by a power law with a photon index  $\Gamma \sim 2.1$  (see Table 6) and a 20–100 keV flux of  $2.4 \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$  at the source. By extrapolating the PDS spectrum to the 2–10 keV band, we find a flux of  $1.7 \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$ , which is higher by a factor of  $\sim 9$  and  $\sim 2$  with respect to the *ASCA* and *Chandra* measurements, respectively. However, if we fit the PDS spectrum with an absorbed power law, fixing the column density  $N_{\text{H}}$  and the photon index  $\Gamma$  to the best-fit values found by Trinchieri et al. (2003) and extrapolate to the 2–10 keV energy range, we find a 2–10 keV flux of  $\sim 7 \times 10^{-12} \text{ erg cm}^{-2} \text{ s}^{-1}$  slightly lower than the *Chandra* value, but still compatible with it. This result strongly indicates that NGC 7319 is a good candidate to explain the PDS emission in the NGC 7331 observation.

#### 4.2.8. The NGC 7552 field

In this field, although two extra sources are present within the MECS image (see Fig. 9), neither seems to be responsible for the high energy emission seen in the PDS.

The LECS and MECS data of the starburst galaxy NGC 7552 are well described by a two-component model, consisting of a thermal part (RAYMOND-SMITH model in XSPEC) with  $kT = 0.92^{+0.28}_{-0.34} \text{ keV}$  and a power law with a steep spectrum ( $\Gamma = 2.74^{+0.91}_{-0.75}$ ). The 2–10 keV flux is  $1.6 \times 10^{-13} \text{ erg cm}^{-2} \text{ s}^{-1}$ , and the extrapolation of this model to high energies shows an almost null ( $< 1\%$ ) contribution to the PDS emission.

As displayed in Fig. 9, within the MECS field of view we find two other sources. The first, located at  $\alpha(2000) = 23^{\text{h}}17^{\text{m}}28^{\text{s}}.8$  and  $\delta(2000) = -42^{\circ}47'36''.2$  is associated to the

Seyfert 1 galaxy FRL 1041, while the second at  $\alpha(2000) = 23^{\text{h}}14^{\text{m}}00^{\text{s}}.7$  and  $\delta(2000) = -42^{\circ}43'17''.3$  is identified with the cluster of galaxies Sersic 159–03. Fitting the MECS data of FRL 1041 with a simple power law, provides a satisfactory fit ( $\chi^2/\nu = 14.9/15$ ), a photon index  $\Gamma = 1.63^{+0.67}_{-0.60}$ , compatible with values found for Seyfert galaxies, and a 2–10 keV flux of  $5.0 \times 10^{-13} \text{ erg cm}^{-2} \text{ s}^{-1}$ . The extrapolation of this model to high energies gives a low (see Table 5) contribution and a joint fit of the MECS/PDS data gives a cross-calibration constant (37–82) well outside the nominal range. Concerning the galaxy cluster, Sersic 159–03, we find that the best-fit model ( $\chi^2/\nu = 26.4/32$ ), for the MECS data only, is a thermal bremsstrahlung having a temperature value ( $kT = 2.54^{+0.24}_{-0.22}$  keV), compatible with that found by *XMM-Newton* observations (Kaastra et al. 2001), plus a narrow Gaussian line centered at  $6.47^{+0.24}_{-0.21}$  keV with an *EW* of  $552^{+229}_{-327}$  eV. The 2–10 keV flux is  $7.0 \times 10^{-12} \text{ erg cm}^{-2} \text{ s}^{-1}$ . The quality of the fit is rather worse when we fit simultaneously the MECS and PDS data; in fact, we cannot put any constraint on the spectral parameters and, in addition, the PDS/MECS cross-calibration constant varies within the interval 91–303. The above results indicate that neither source is likely to be responsible for the high energy emission (see also Table 5), even in the case of the cluster which is strongly detected in the MECS instrument ( $\sim 50\sigma$  level), but has an extremely soft spectrum.

In view of these findings, we performed a further search in order to understand if there was any other nearby object which could produce the PDS result. Just outside the MECS field of view we find NGC 7582, a Seyfert 2 galaxy, belonging to the Piccinotti sample, i.e. a very bright source in hard X-rays (Piccinotti et al. 1982; Malizia et al. 1999).

NGC 7582 has already been observed by *BeppoSAX* in November 1998 (Turner et al. 2000). In that case the broad spectrum (2–100 keV) was described by a power law of photon index  $\Gamma = 1.95^{+0.09}_{-0.18}$  (steeper than the index found during a previous *ASCA* observation) transmitted through a dual absorber with column densities  $N_{\text{H}} \sim 1.6 \times 10^{24} \text{ cm}^{-2}$  (covering 60% of the nucleus) and  $\sim 1.44 \times 10^{23} \text{ cm}^{-2}$  (fully covering the source). The authors reported a 10–100 keV flux of  $\sim 1.2 \times 10^{-10} \text{ erg cm}^{-2} \text{ s}^{-1}$  and the presence of variability on time-scales down to a few thousand of seconds. In addition, Turner et al. (2000) found that NGC 7582 was significantly brighter than the average level sampled by *OSSE* ( $\sim 4 \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$  in the 50–150 keV band) and *BATSE* ( $\sim 8.9 \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$  in the 20–100 keV band) (Johnson et al. 1997; Malizia et al. 1999), both estimated using a power law model.

By subtracting from the PDS the contribution due to each of the three sources present in the MECS image of Fig. 9 and fitting the remaining emission with a simple power law, we find a photon index of 1.72 (see Table 6) and a 20–100 keV flux (50–150 keV) at the source of  $8.13 \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$  ( $\sim 5 \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$ ), which is in perfect agreement with the values reported by *BATSE* and *OSSE*. If we fit our data, taking into account the double absorber proposed by Turner et al. (2000), we find a photon index  $\Gamma \sim 1.9$  and a 10–100 keV flux of  $2.3 \times 10^{-10} \text{ erg cm}^{-2} \text{ s}^{-1}$ , a factor of 2 higher than the value reported for the 1998 observation by Turner et al. (2000).

To conclude, NGC 7582 is very likely the serendipitous source responsible for the emission detected above 10 keV in the pointing of NGC 7552. Furthermore, we provide evidence for variability ( $\sim 90\%$ ) in the flux between the two *BeppoSAX* measurements taken one year apart.

## 5. Conclusions

In this work we have shown how a careful search in the *BeppoSAX* public archive can provide evidence for new hard X-ray emitting sources, most likely associated with active galaxies, and/or confirm the high energy emission of known objects. In particular, we report the detection of six new hard X-ray emitters (two type 1 Seyfert galaxies, two type 2, one quasar and a galaxy not yet classified as active); for two of these sources broad band spectra are presented, while in four cases only results above 10 keV are reported.

For the remaining 6 objects in our sample, emission above 10 keV was known before this work, mainly from previous dedicated *BeppoSAX* pointings. In the case of NGC 3227, an object known to emit at high energies from past *OSSE* and *BATSE* measurements, a spectrum is published here for the first time. Comparison of the present data with previous observations either with *BeppoSAX* and/or with other high energy missions indicates that flux variability is present in three (MKN 766, NGC 7582 and NGC 3227) possibly four (NGC 1275) objects.

As the AGN reported in this paper are “loosely” representative of the population of extragalactic objects with emission above 10 keV, it is worth examining their overall characteristics. The first consideration to make is that most of our objects (6 out of 10 with optical classification) are broad line emitting or unabsorbed AGN, contrary to the expectation of finding more absorbed than unabsorbed objects. The other four (possibly five if we include PKS 0101–649) show intrinsic absorption which, however, is compatible with a Compton thin nature ( $N_{\text{H}} < 10^{24} \text{ cm}^{-2}$ ). Only in the case of PKS 0101–649 could the source be Compton thick. This is probably due to the limited capability of current hard X-ray detectors which allow just the brightest and nearest extragalactic objects to be found. In fact, all but two (Perseus and 4C +74.26) or even one (if NGC 1275 is responsible for the emission detected by the PDS) of our AGNs are nearby ( $z < 0.1$ ) and belong to the 2 Micron All Sky Survey eXtended Source Catalog (i.e. the surrounding galaxy is visible in the near-infrared). At least eight (possibly nine if we also consider NGC 1275) of our objects are radio loud, according to the Terashima & Wilson (2003) definition.

Overall, we can conclude that our objects are bright, nearby, with a detectable emission in all wavebands and only a fraction ( $\sim 30\text{--}40\%$ ) have a column density in excess of  $10^{22} \text{ cm}^{-2}$ .

As shown here it is also possible to encounter offset fields contaminated by the presence of serendipitous sources. We are now performing a systematic analysis of all PDS data present in the archive (above  $15^\circ$  in galactic latitudes) evaluating the relative offset fields; the aim of this work is to provide a list of positive detections in these background measurements. We are confident that this analysis will provide new high energy sources as highlighted here in three cases.

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## References

- Arnaud, K. A. 1996, in *Astronomical Data Analysis Software and Systems V*, ed. G. H. Jacoby, & J. Barnes (San Francisco: ASP), ASP Conf. Ser., 101, 17
- Awaki, H., Koyama, K., Matsumoto, H., et al. 1997, *PASJ*, 49, 445
- Bassani, L., Dadina, M., Maiolino, R., et al. 1999, *ApJS*, 121, 473
- Bassani, L., Malizia, A., Stephen, J. B., et al. 2004, *Proc. Fifth INTEGRAL Workshop*, ESA SP-552
- Bird, A. J., Barlow, E. J., Bassani, L., et al. 2004, *ApJ*, 607, L33
- Blanton, E. L., Sarazin, C. L., & Irwin, J. A. 2001, *ApJ*, 552, 106
- Boella, G., Butler, R. C., Perola, G. C., et al. 1997a, *A&AS*, 122, 299
- Boella, G., Chiappetti, L., Conti, G., et al. 1997b, *A&AS*, 122, 327
- Bregman, J. N., & Pildis, R. A. 1994, *ApJ*, 420, 570
- Brinkmann, W., Otani, C., Wagner, S. J., & Siebert, J. 1998, *A&A*, 330, 67
- Chiappetti, L., & Dal Fiume, D. 1997, in *Data Analysis in Astronomy*, ed. V. Di Gesu, M. J. B. Duff, A. Heck, et al. (Singapore: World Scientific), 111
- Dermer, C. D., & Gherels, N. 1995, *ApJ*, 447, 103
- Dickey, J. M., & Lockman, F. J. 1990, *ARA&A*, 28, 215
- Eracleous, M., Halpern, J. P., & Livio, M. 1996, *ApJ*, 459, 89
- Fabian, A. C. 2001, in *X-Ray Astronomy 2000*, ed. R. Giacconi, S. Serio, & L. Stella (San Francisco: ASP), ASP Conf. Proc., 234, 545
- Fiore, F., Guainazzi, M., & Grandi, P. 1999a, *SAXabc*, vs. 1.2, *Cookbook for BeppoSAX NFI Spectral Analysis*
- Fiore, F., La Franca, F., Giommi, P., et al. 1999b, *MNRAS*, 306, L55
- Frontera, F., Cinti, M.N., Dal Fiume, D., et al. 1997a, *Nuovo Cimento C*, 20, 797
- Frontera, F., Costa, E., Dal Fiume, D., et al. 1997b, *A&AS*, 122, 357
- Gehrels, N., Chincarini, G., Giommi, P., et al. 2004, *ApJ*, 611, 1005
- George, I. M., Mushotzky, R., Turner, T. J., et al. 1998, *ApJ*, 509, 146
- Gondoin, P. 2004, *A&A*, 415, 1113
- Gondoin, P., Orr, A., Lumb, D., & Siddiqui, H. 2003, *A&A*, 397, 883
- Grandi, P., Guainazzi, M., Haardt, F., et al. 1999, *A&A*, 343, 33
- Halpern, J. P. 1982, Ph.D. Thesis Harvard University, Cambridge, MA
- Hasenkopf, C. A., Sambruna, R. M., & Eracleous, M. 2002, *ApJ*, 575, 127
- Johnson, W. N., Zdziarski, A. A., Madejski, G. M., et al. 1997, in *Proc. Fourth Compton Symp., Part 1*, ed. C. D. Dermer, M. S. Strickman, & J. D. Kurfess (New York: AIP), AIP Conf. Proc., 410, 283
- Inda, M., Makishima, K., Kohmura, Y., et al. 1994, *ApJ*, 420, 143
- Kaastra, J. S., Ferrigno, C., Tamura, T., et al. 2001, *A&A*, 365, 99
- Koekemoer, A. M., & Bicknell, G. V. 1998, *ApJ*, 497, 662
- Kruper, J. S., Canizares, C. R., & Urry, C. M. 1990, *ApJS*, 74, 347
- Lamer, G., Uttley, P., & Hardy, I. M. M. 2003, *MNRAS*, 342, L41
- Leighly, K. M., Mushotzky, R. F., Yaqoob, T., Kunieda, H., & Edelson, R. 1996, *ApJ*, 469, 147
- Leighly, K. M., O'Brien, P. T., Edelson, R., et al. 1997, *ApJ*, 483, 767
- Levine, A. M., Lang, F. L., Lewin, W. H. G., et al. 1984, *ApJS*, 54, 581
- Lipovetsky, V. A., Neitzvestny, S. I., & Neitzvestnaya, O. M. 1988, *Soobshch. Spets. Artrof. Obs.*, 55, 5
- Loring, N. S., Page, M. J., & Ramsay, G. 2003, *MNRAS*, 345, 865
- Magdziarz, P., & Zdziarski, A. A. 1995, *MNRAS*, 273, 837
- Malizia, A., Bassani, L., Zhang, S. N., et al. 1999, *ApJ*, 519, 637
- Matt, G., Perola, G. C., Fiore, F., et al. 2000a, *A&A*, 363, 863
- Matt, G., Fabian, A. C., Guainazzi, M., et al. 2000b, *MNRAS*, 318, 173
- Mewe, R., Gronenschild, E. H. B. M., & van den Oord, G. H. J. 1985, *A&AS*, 62, 197
- Monet, A. K. B., Levine, S. E., Monet, D. G., et al. 1999, *BAAS*, 31, 1532
- Monet, D. G., Levine, S. E., Canzian, B., et al. 2003, *AJ*, 125, 984
- Nandra, K., & Pounds, K. A. 1994, *MNRAS*, 268, 405
- Osako, C. Y., Ulmer, M. P., Grabelsky, D. A., et al. 1994, *ApJ*, 435, 181
- Pan, H. C., Jordan, C., Makishima, K., et al. 1997, *MNRAS*, 285, 735
- Parmar, A. N., Martin, D. D. E., Bavdaz, M., et al. 1997, *A&AS*, 122, 309
- Perri, M., Massaro, E., Giommi, P., et al. 2003, *A&A*, 407, 453
- Piccinotti, G., Mushotzky, R. F., Boldt, E. A., et al. 1982, *ApJ*, 253, 485
- Pounds, K. A., Nandra, K., Stewart, G. C., & Leighly, K. 1989, *MNRAS*, 240, 769
- Quadrelli, A., Malizia, A., Bassani, L., & Malaguti, G. 2003, *A&A*, 411, 77
- Raymond, J. C., & Smith, B. W. 1977, *ApJS*, 35, 419
- Sambruna, R. M., Eracleous, M., & Mushotzky, R. F. 1999, *ApJ*, 526, 60
- Schwöpe, A. D., Hasinger, G., Lehmann, I., et al. 2000, *AN*, 321, 1
- Stockdale, C. J., Romanishin, W., & Cowan, J. J. 1998, *ApJ*, 508, 33
- Terashima, Y., & Wilson, A. S. 2003, *ApJ*, 583, 145
- Trinchieri, G., Pellegrini, S., Wolter, A., Fabbiano, G., & Fiore, F. 2000, *A&A*, 364, 53
- Trinchieri, G., Sulentic, J., Breitschwerdt, D., & Pietsch, W. 2003, *A&A*, 401, 173
- Turner, T. J., George, I. M., Nandra, K., & Mushotzky, R. F. 1997, *ApJ*, 488, 164
- Turner, T. J., Perola, G. C., Fiore, F., et al. 2000, *ApJ*, 531, 245
- Ueda, Y., Ishisaki, Y., Takahashi, T., Makashima, K., & Ohashi, T. 2001, *ApJS*, 133, 1
- van den Besselaar, E. J. M., Raassen, A. J. J., Mewe, R., et al. 2003, *A&A*, 411, 587
- Wood, K. S., Meekins, J. F., Yentis, D. J., et al. 1984, *ApJS*, 56, 507
- Zdziarski, A. A., Poutanen, J., & Johnson, W. N. 2000, *ApJ*, 542, 703