

Variable stars towards the bulge of M 31: The AGAPE catalogue^{★,★★}

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Abstract. We present the AGAPE astrometric and photometric catalogue of 1579 variable stars in a $14' \times 10'$ field centred on M 31. This work is the first survey devoted to variable stars in the bulge of M 31. The R magnitudes of the objects and the $B - R$ colours suggest that our sample is dominated by red long-period variable stars (LPV), with a possible overlap with Cepheid-like type II stars. Fits of the light curves with sinusoids suggest that a large fraction of the stars correspond to periodic or semi-periodic objects with periods longer than 100 days. Twelve nova candidates are identified. Correlations with other catalogues suggest that 2 novae could be recurrent novae and provide possible optical counterparts for 2 supersoft X-ray sources candidates observed with Chandra.

Key words. galaxies: individual: M 31 – cosmology: dark matter – stars: variables: general – stars: novae, cataclysmic variables – stars: variables: Cepheids – X-rays: stars

1. Introduction

AGAPE (Ansari et al. 1997) made a survey of several fields in the central part of M 31 to look for gravitational microlensing events. As a major by-product, this survey discovered 1579 variable objects in a $14' \times 10'$ field centred on M 31. The aim of this paper is to catalogue these objects. Due to the high star background, pixel lensing techniques have to be used to detect flux variations of individual stars. Since we are

observing brightness fluctuations of unresolved stars, variable stars, particularly those that are intrinsically faint, are a significant background in the detection of real microlensing events. On the other hand, variable stars are also of high interest in their own right. Systematic surveys looking for variable stars in M 31 have been first done by Hubble (1929) and Baade & Swope (1963, 1965). Recent surveys on large fields have been used to search for Cepheids (Magnier et al. 1997) and cosmological calibrators (DIRECT project, Kaluzny et al. 1998). However, except for novae and X-ray sources Kaaret (2002), Kong et al. (2002), published work avoid the bulge region. Our study is then the first one done on this central zone and it leads to the identification of some variable stars poorly known in M 31: Miras and long period Cepheid-like objects. Our catalogue is drawn from 3-seasons of measurements and our field has significant intersection with those of other studies (Crotts & Tomaney 1996, DIRECT field E, Riffeser et al. 2001) and the Point AGAPE collaboration using the INT (Kerins et al. 2001).

Section 2 presents AGAPE observations and variable selection. Section 3 presents the astrometry as well as our

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* The full Table 3 is only available in electronic form at the CDS via anonymous ftp to

cdsarc.u-strasbg.fr (130.79.128.5) or via
<http://cdsweb.u-strasbg.fr/cgi-bin/qcat?J/A+A/421/509>

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photometry and period search (Time Series Analysis). Section 4 presents our complete catalogue. Section 5 discusses our type-identified variable stars and Sect. 6 presents some conclusions.

2. AGAPE observations and variable selection

2.1. AGAPE observations and pixel method photometry

The AGAPE observations were made at the 2 m Bernard Lyot telescope (TBL) of the Pic du Midi Observatory with the f/8 spectro-reducer ISARD. A thin Tektronik 1024 × 1024 CCD was used with a useful field of 4' × 4.5' with 0.3'' pixels. Since the field of ISARD is small, 6 fields (called A, B, C, D, E, F) were necessary to cover one 14' × 10' field centred on M 31 and oriented along its main axis. The exposure times were generally 30(20) min in the *B(R)* pass-band. An additional field, Z, centred on the nucleus of M 31 and with large overlap with A and B fields, was taken at the beginning of each night, as a reference to help in the pointing of the telescope. The exposures were as short as 1 min in both the *B* and *R* passbands for the Z central field. Since the short exposures prevent detection of faint or small amplitude variable objects, only novae are presented for the Z field data in this paper. The observing campaign ran from 1994 to 1996. It turned out that it was impossible to monitor all the fields in both colours each night. We put a priority on the first four fields, with an emphasis on red exposures. We obtained 69 observing nights in *R*: 19 in 1994, 43 in 1995 and 7 in 1996. The number of images taken in each field during the whole survey is given in Table 1.

The AGAPE detection procedure is described in Ansari et al. (1997). It is based on super-pixel (7 × 7 pixels) photometry with sides roughly twice the standard seeing. These super-pixels are photometrically normalised to a reference frame and corrected for seeing variations. Then we transform the super-pixel ADU photometry to Johnson-Cousins *B* and *R* magnitudes as described in Sect. 3.2.

2.2. Variable stars selection

During the selection procedure, we excluded an area of 20 pixels around well-resolved stars. Automatic procedures are efficient to detect stars brighter than $R = 19$. Le Du (2000) studies the relation of this influence zone with seeing. These resolved stars should contain no or very few M 31 variable stars but stars from our Galaxy and M 31 globular clusters. The light curves for each uncontaminated super-pixel are analysed. Baseline for background flux is taken as the minimum of a sliding average calculated on 4 to 8 consecutive points depending of the time sampling of the field. One bump is defined as beginning when 3 consecutive points lie 3σ above background and ending when 2 consecutive points fall below 3σ . For each bump, we compute its likelihood function:

$$L = -\ln \left(\prod_{i \in \text{bump}} P(\Phi > \Phi_i | < \Phi >_i, \sigma_i) \right).$$

Table 1. Number of images taken during the AGAPE observations.

Year	1994	1995	1996	total
filter	<i>R/B</i>	<i>R/B</i>	<i>R/B</i>	<i>R/B</i>
A	25/10	51/21	8/2	84/33
B	24/10	41/21	7/2	73/33
C	20/8	40/16	7/0	67/24
D	18/5	34/14	6/0	59/19
E	6/0	34/10	6/0	46/10
F	5/0	27/8	6/0	38/8
Z	32/12	55/20	13/1	100/33

One superpixel is selected as including a variable object when it contains one bump with its likelihood function greater than 100. We choose a high value for L in order to eliminate false variable stars. These criteria lead to the selection of 1579 variable stars.

3. Astrometry, photometry and period search

3.1. Astrometry

Astrometry is based upon the catalogue of Magnier et al. (1992), the most populous one available towards the bulge of M 31 with the most homogeneous coverage of our field. This catalogue is connected to the Space Telescope Guide Star Catalog with an accuracy better than 1'' (Magnier et al. 1993; Haiman et al. 1994). For each of AGAPE fields B to F we have from 6 to 20 stars in common with this catalogue. The final standard deviation on the positions is in the range 0.10''–0.14''. However, for the near-central fields Z and A, less than 3 stars are identified as common and we were obliged to use another reference catalogue. We thus used the one obtained by Aurière et al. (1992) for their globular clusters study, based on the combination of Battistini et al. (1987) and Wirth et al. (1985). For these two fields the standard deviation on the distances between our positions and those on the catalogue for 11 and 10 common stars in fields A and Z respectively is 0.45''. Finally we conclude that we have an astrometric accuracy in the range 1''–1.5'' for our variable star positions, combining our uncertainties with those of Magnier and the other catalogues.

3.2. Photometry

The super-pixel photometry gives the integrated flux in squares of 7 × 7 pixels corresponding to 2.1'' × 2.1'' on the sky. We first have to remove the stellar background which is bright and chaotic in the M 31 bulge. Almost all variable stars are not resolved nor significantly measured at minimum. In this case we can consider the background as being the average minimum flux, which is taken as the minimum of the sliding average on five points along the light curve. However, the star always contributes at some level even at minimum brightness and we also compute the average flux (with 3σ rejection) in a zone around the object. The background is taken as the lesser of these two values and is subtracted from flux values. To obtain *B* and *R* Johnson-Cousins magnitudes, we use the colour equations as

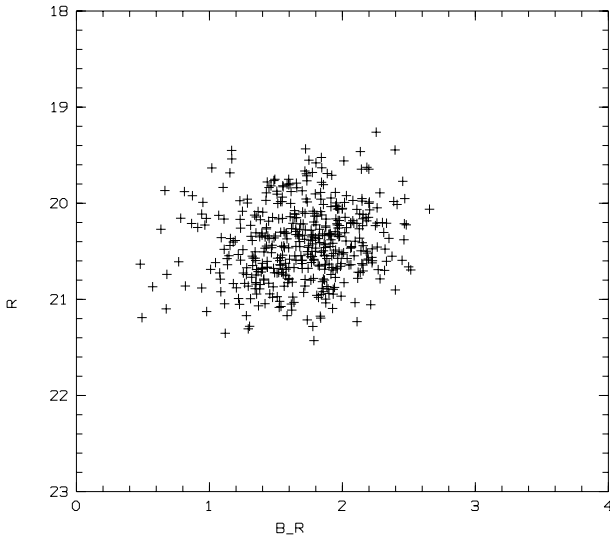


Fig. 1. Colour Magnitude diagram at observed maximum brightness for the 459 variable stars which have one B measurement.

given in Ansari et al. (1997) transforming between super-pixel and PSF photometry. Since our field contains several “bright unvariable” stars measured by Magnier et al. (1992), we use 36 common stars in R and 19 common stars in B (excluding central fields A and B) to derive magnitude scale constants for our data. The standard deviations corresponding to the computation of these constants are 0.15 and 0.24 for R and B respectively. They are dominated by random scatter and include our errors and Magnier et al. (1992) errors. This accuracy is sufficient to characterise the R magnitude at maximum observed brightness of the variable stars and the colour range for our star sample, but is too poor for using individual $B - R$ indices for identification purposes. The B magnitude measurements will be thus used for making a Colour Magnitude Diagram but will not be given in the final catalogue. To compute the R magnitude at maximum we select the 5 brightest R measurements and eliminate the 2 brightest ones to avoid possible residual cosmic rays. R_{\max} is the average of the 3 remaining measurements. To enable a reliable B magnitude measurement and thus a $B - R$ colour, the stars have to emerge significantly from the background. Since there are fewer B measurements, the magnitude at maximum is the average of the 2 brightest measurements. $B - R$ is the difference between the peak B and R in each passband. Figure 1 gives our Colour Magnitude Diagram at observed maximum brightness for the 459 variable stars which have been measured in the B passband, in fields A, B, C, D. This diagram is interpreted in Sect. 5.1.

3.3. Period search (time-series analysis)

Our photometric data, for the time series analysis, are characterised by: – their sampling which is unevenly spaced; – one measurement per night; – yearly season-limited autumn observations. We are able to detect periods in the range 60–400 days (and we will see that this range is appropriate for the study of bulge variables which contain an interesting population of long-period objects). In this work we do not extract an

Table 2. Variable stars with period estimate. Shortest period objects are population II Cepheid-like objects; longest period objects are LPVs.

Ident	R_{\max}	Period (day)	Ident	R_{\max}	Period (day)
B030	20.5	060	C032	20.6	187
B098	21.1	103	B005	20.3	188
A266	21.1	111	A320	20.2	189
B054	19.9	116	A077	19.7	190
A058	19.6	123	B043	20.1	193
C071	20.9	125	A189	20.4	194
C188	21.2	127	A185	20.6	196
A014	20.1	131	A044	20.3	197
C055	20.5	132	A162	20.0	197
C012	20.7	133	A241	20.8	197
C081	20.9	134	A343	20.0	200
C013	20.3	140	A190	20.3	202
A059	20.4	142	A002	19.6	210
A018	20.1	148	B006	20.3	212
A046	20.2	151	C016	20.5	212
B014	20.4	164	C169	21.0	216
A286	20.6	155	A019	20.2	220
B058	20.6	156	A109	20.3	223
C017	20.5	162	B140	20.0	244
B165	20.2	164	C006	20.1	244
A079	19.9	165	B007	20.0	250
B016	20.0	165	A136	20.9	254
C128	20.7	165	A049	20.6	259
A060	20.3	168	C004	19.9	335
B164	20.2	171	C072	20.1	344
C177	20.1	179	A001	19.5	368
C015	20.3	186	A282	20.6	385

extensive sample of periodic variable stars but rather identify a subset to characterise our sample. We use the Scargle method (1982) as implemented in MIDAS for the Time Series Analysis of our data. It is a modification of the periodogram analysis, equivalent to best-squares fitting of sine waves to the data and suited to unevenly spaced data (Scargle 1982). For each variable star, we have fitted the Scargle sinusoid on the light curves from the three seasons and made a visual inspection. A majority of these light curves could correspond to periodic or semi-periodic objects with periods greater than 100 days. The longest observational season (1995) spans about 150 days. Periods shorter than 150 days are thus favoured for giving a convincing visual confirmation of the light-curve fit. For periods greater than 250 days, light curves with a sharp feature which is clearly observed on each year are accepted. For these longest periods, we checked systematically that the half period did not fit. We present in Figs. 2–10 nine of these light curves fitted by their Scargle sinusoid (X -axis is in (J-2449624.5) days, Y -axis in ADU/s).

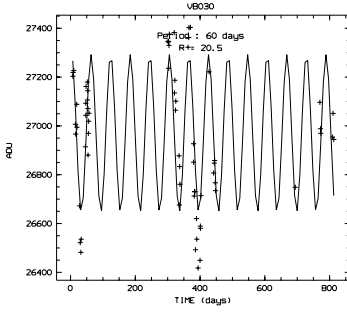


Fig. 2. Variable star B030: $P = 60$ days; $R = 20.5$.

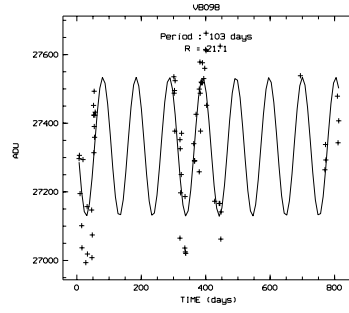


Fig. 3. Variable star B098: $P = 103$ days; $R = 21.1$.

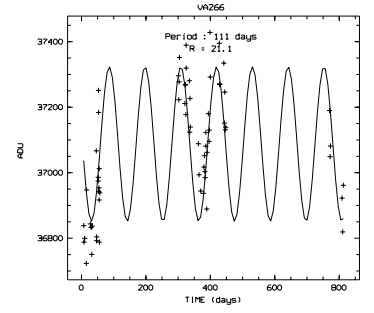


Fig. 4. Variable star A266: $P = 111$ days; $R = 21.1$.

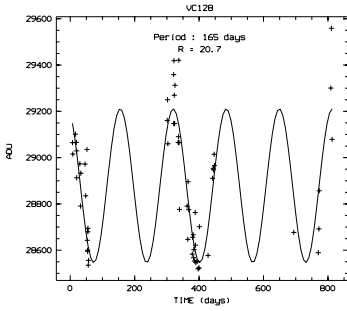


Fig. 5. Variable star C128: $P = 165$ days; $R = 20.7$.

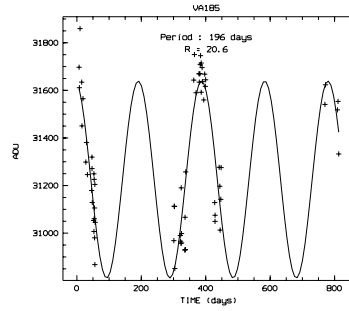


Fig. 6. Variable star A185: $P = 196$ days; $R = 20.6$.

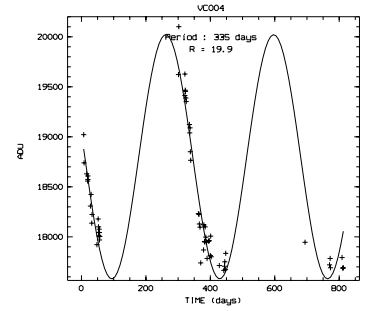


Fig. 7. Variable star C004: $P = 335$ days; $R = 19.9$.

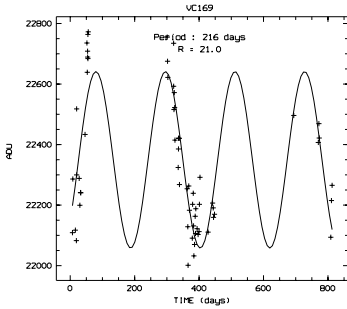


Fig. 8. Variable star C169: $P = 216$ days; $R = 21.0$.

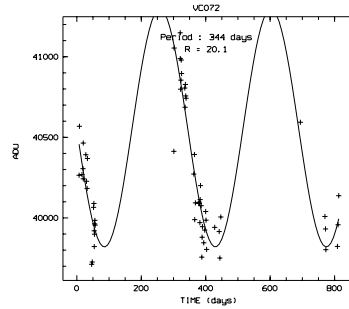


Fig. 9. Variable star C072: $P = 344$ days; $R = 20.1$.

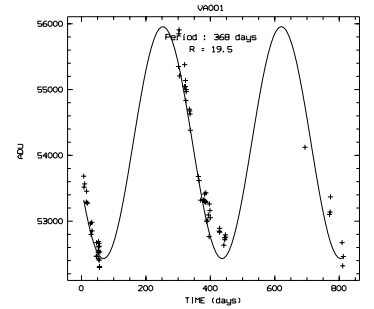


Fig. 10. Variable star A001: $P = 368$ days; $R = 19.5$.

Fig. 2–Fig. 10. Variable light-curves fitted by their Scargle sinusoid. X-axis is in (J-24449624.5) days, Y-axis in ADU/sec. The 3 first objects are of the Cepheid type, the 5 last ones are certainly LPVs.

Table 2 gives R_{\max} and periods for 54 stars for which the visual fit of the Scargle sine was as good as in the cases of Figs. 2–10. The period estimate is given with 10 to 20% accuracy. Since our period determinations are only estimates, we don't give them in our final catalogue.

3.4. Period luminosity diagram

The Period Luminosity Diagram is a classical representation for characterising pulsating variable stars. We thus plot our data presented in Table 2 in such a diagram in Fig. 11. In order to characterise the nature of our variable stars we have drawn fiducial lines for the location of M 31 cepheids (Freedman & Madore 1990) LMC RV Tauri stars (Alcock et al. 1998) and Miras (Clayton & Feast 1969), the two last sets of data being converted to M 31 distance. M 31 cepheids and LMC RV Tauri stars have been observed in the R passband and obtaining these

lines was rather straightforward. On the other hand, as for Miras, almost all recent investigations for P/L studies are made in infrared colours or are transformed in bolometric magnitudes. Thus, we decided to use the V passband absolute magnitude/period diagram of Clayton & Feast (1969), which we transformed to R passband at M 31 distance. Colour equation from V to R is derived from diagrams from MACHO data (Alves et al. 1998). We used also this reference to check the reliability of the Clayton & Feast's data for long period Miras, since these authors said there were uncertain. The astrophysical aspect of this diagram is discussed in Sect. 5.2.

4. Variable star general catalogue

4.1. The catalogue

Each variable star in the catalogue is identified as AGPVM, m running from 1 to 1579. Then are given: $F_n(x, y)$ (F being the

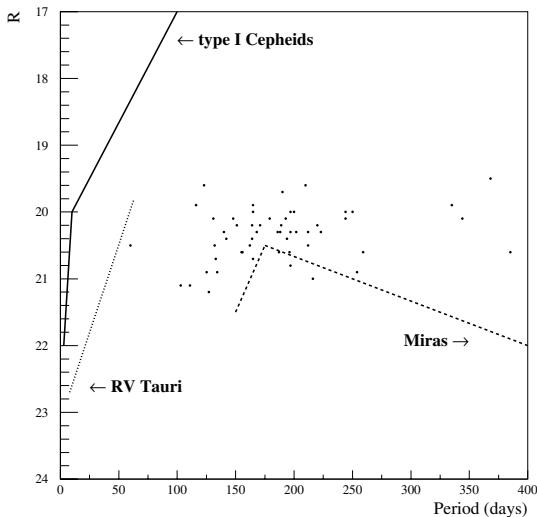


Fig. 11. Period Luminosity Diagram: R magnitude is plotted against the period for the stars selected in Table 2. Locations for Cepheids, RV Tauri stars and Miras are shown. Details are given in the text.

letter of identification of AGAPE field and n the variable star number in this field; x and y being the pixel coordinates in the reference frame), right ascension and declination for 2000.0, R_{\max} magnitude at observed maximum as defined in Sect. 3.2, and L the likelihood function as defined in Sect. 2.2. Table 3 presents an extraction of this catalogue, showing the first and last star for each field and the 12 nova candidates with their numbers added (with respect to the CDS catalog) in the last column. When the novae were discovered by eye, no L value is given.

4.2. Correlation with other catalogues

We have then correlated our catalogue with the General Catalogue of Variable Stars (GCVS, Durlevich et al. 1996) as available at CDS, with Chandra sources (Kaaret 2002; Kong et al. 2002; Di Stefano et al. 2003), as well as 2MASS data (Skrutskie et al. 1997). As to the GCVS, about 200 objects fall in our field, all being classified as novae but one cepheid. 7 objects of GCVS (all novae) fall at less than $2''$ from our variable stars. Among these associations, only one concerns one of our nova candidates and we consider that the other ones are chance associations. Nova 27 from GCVS is at $1.9''$ from NOV12 (AGPV1579, Z002, Sect. 5.3). Nova 27 was observed in 1985 (Ciardullo et al. 1987) and could then be a recurrent nova.

M 31 V0934 of GCVS, which is V38 of (Hubble 1929), is a type I cepheid (period 17.764 days). It falls in our field C but has no counterpart in our catalogue. However, we find on our data one variable object located $2.5''$ from V0934 position with R about 20 mag at maximum. M 31 V0934 may have been discarded in our microlensing survey as being one resolved star on our reference frame (Sect. 2.2).

Chandra made several observations of the bulge of M 31. Tens of X-ray sources were detected with location accuracy better than $1''$ (Kaaret 2002; Kong et al. 2002). These X-ray observations are the only ones which deserve correlation with

our catalogue since they have small error boxes ($1''$ accuracy versus several several $''$ for ROSAT HRI) necessary to avoid too many chance identifications. Among the 4 X-ray sources which fall nearer than $2''$ from our variable objects, 2 deserve peculiar attention. CXOM31 J004242.1+411218 (Kaaret 2002) is located $0.7''$ from NOV09 (AGPV1607, F139). This identification is discussed in Sect. 5.3.

CXOM31 J004247.8+411550. is located $0.6''$ from AGPV0288, in the very central bulge. Kong et al. (2002) observed it again as r1-25. Di Stefano et al. (2003) include this source in their list of super soft X-ray sources candidates. Di stefano et al. (2003) could not find any optical counterpart brighter than magnitude 21 on HST archives, in the R wavelength range. AGPV0288 peaks at $R = 19.7$ in AGAPE observations, and is a small amplitude irregular object. r1-25 is fitted with a blackbody model $kT = 122$ eV, rather hard for a supersoft source. Clearly, this association will deserve a deeper study.

About 50 objects from 2MASS catalogue fall in our field. None are associated with AGAPE variable stars within $2''$ radius, but some are associated with brighter objects (galactic stars or globular clusters). This is in accordance with the fact that the rather short period LPVs of our catalog are expected to be fainter in K than the limiting magnitude of the 2MASS catalog (about $K = 16$), when scaling LMC results (Hughes & Wood 1990; Lebzelter et al. 2002).

4.3. Spatial distribution of variable stars and the extinction in M 31 bulge

The spatial distribution of the variable stars is of importance for two reasons: variable stars are often considered as a possible reference background from which it will be possible to disentangle a microlensing event population which will present a near/far M 31 asymmetry (Crotts 1992; Kerins et al. 2001); part of the variables we have observed really belong to the bulge and we can thus contribute to the disk/halo distribution debate of objects in M 31. An important factor in the star distribution in M 31 comes from extinction (Hatano et al. 1997).

In the AGAPE survey, the six investigated fields are distributed apart from the major axis of M 31, A D F being on the “far side” and B C E on the “near side”. We consider that the E and F fields have too sparse observations to give significant data (cf. Table 1) for this problem. At the near side, which corresponds to the “bottom” of (Hatano et al. 1997), the bulge is in part behind the (dusty) disk. In our study, the most significant anomaly in the variable star distribution is due to B field. It contains half the number of variable stars in field A (190 with respect to 370; almost same ratio for the corresponding numbers of stars per brightness unit). The B field contains extinction zones such as the dust complex D395A (Melchior et al. 2000b). However, using the extinction map of Melchior et al. (2000b), we did not find a correlation between the variable stars and extinction distributions. This could come because extinction can be very patchy in galaxies (Thoraval et al. 1999). Local large extinction has been pointed out in M 31 in the case of globular clusters (Barmby et al. 2002).

Table 3. Extraction of the catalogue of the AGAPE variable stars showing the first and last star for each field and the 12 novae candidates (more information are given in the text). L is the likelihood function. The nova number is added (with respect to CDS catalog) in the last column; more information on the columns is given in Sect. 4.1.

Id	Fn(x, y)	$\alpha(2000.0)$	$\delta(2000.0)$	R_{\max}	L	#
AGPV0001	A001(0697,0392)	0h42m53.36s	41d14'26.8"	19.48	940	
...						
AGPV0390	A397(0834,0352)	0h42m50.26s	41d14'00.8"	20.46	100	
AGPV0391	A398(0365,0380)	0h42m58.52s	41d15'49.8"	17.14		NOV01
AGPV0392	B001(0920,0220)	0h42m26.41s	41d16'56.2"	16.07	32000	NOV03
AGPV0393	B002(0754,0333)	0h42m31.57s	41d17'16.0"	17.24	20000	NOV02
AGPV0394	B003(0422,0305)	0h42m36.32s	41d18'41.7"	19.45	720	
...						
AGPV0581	B190(0446,0372)	0h42m37.40s	41d18'23.7"	20.94	100	
AGPV0582	C001(0597,0350)	0h42m50.65s	41d21'20.4"	19.28	3000	
AGPV0583	C002(0784,0384)	0h42m48.36s	41d20'28.5"	18.47	2800	NOV04
...						
AGPV0842	C272(0771,0504)	0h42m51.18s	41d20'09.8"	21.11	100	
AGPV0843	D001(0910,0609)	0h43m 8.41s	41d16'39.9"	16.80	30000	NOV05
AGPV0844	D002(0927,0567)	0h43m 7.22s	41d16'43.4"	19.39	1800	
...						
AGPV1185	D348(0585,0580)	0h43m13.04s	41d18'04.9"	21.43	100	
AGPV1186	E001(0721,0300)	0h42m17.24s	41d13'35.4"	17.65	4900	NOV08
AGPV1187	E002(0459,0646)	0h42m29.00s	41d13'36.1"	18.22		NOV06
AGPV1188	E003(0317,0264)	0h42m23.07s	41d15'20.6"	20.63	490	
...						
AGPV1438	E259(0337,0392)	0h42m25.51s	41d14'52.3"	21.00	100	
AGPV1439	E260(0724,0490)	0h42m21.30s	41d20'28.7"	18.41		NOV07
AGPV1440	F001(0905,0390)	0h42m33.27s	41d09'59.6"	20.35	520	
...						
AGPV1575	F138(0425,0400)	0h42m41.22s	41d11'55.2"	20.89	100	
AGPV1576	F139(0346,0381)	0h42m42.08s	41d12'18.0"	17.56		NOV09
AGPV1577	F140(0430,0650)	0h42m46.60s	41d11' 8.3"	17.82		NOV10
AGPV1578	Z001(0528,0266)	0h42m43.75s	41d16'52.0"	16.35		NOV11
AGPV1579	Z002(0658,0411)	0h42m44.81s	41d15'53.6"	14.98		NOV12

Thus, because of extinction and perhaps threshold effects, the variable star distribution towards the bulge of M 31, might be not symmetric with respect to the major axis of the galaxy.

5. Identified variables

5.1. Introduction

Our observational data enable us to characterise variable stars with their absolute magnitudes, colours and periods. Observational conditions of the AGAPE experiment bring some biases and thus some limitations on the identification of the variable stars which were discussed in previous sections. Our robust variability criteria lead to almost all measured R_{\max} of variable stars being between 20 and 21 mag. Inspection of

tables of data published by the DIRECT collaboration shows that the bulk of their detected variable stars is in the same magnitude range as ours.

Our search detects M 31 variable stars brighter than $R = 21$ or $M_R < 3.3$. Colour magnitude diagrams such as that of Kukarkin (1975) show that this limit enables the study of only the brightest part of the variable star population including long period red stars and Cepheids, and almost all the kinds of stars observed by Baade et al. (1963, 1965). On the other hand, RR Lyrae stars, which are difficult to observe from the ground in M 31 because of their faintness (Pritchett & van den Bergh 1987), are clearly out of reach.

We are able to study the population II variable stars which are tracers of the bulge in a region where by projection bulge

and disk populations are not easy to disentangle. Few population II objects like RV Tauri stars have been studied, and Miras have not been unambiguously detected in M 31. Up to now variable surveys were done more in the disk population than in the bulge. Studies have shown that when approaching the center of M 31 a number of new star classes appeared. This is first the case for bright red stars, some of them being possibly very long period Miras (Rich & Mould 1991). Hodge (1992) remarks that from the work of Swope & Baade (1963, 1965) the number of irregular variables increases when approaching the bulge. Also, the same author shows that the ratio of type I Cepheids to type II Cepheids has a maximum in medium parts of the disk. Thus one expects to find very few type I Cepheids in the bulge direction. Apart from novae which are outstanding bright objects in our sample, other objects can be characterised by the CMD of the population resolved in *B*. Figure 1 shows that this population is peaked on *R* around 20.4 ± 0.5 mag. ($M_R = -4 \pm 0.5$) and *B*–*R* around 1.7 ± 0.5 mag. The photometric properties of the variable stars when compared to the HR diagram of Kukarkin (1975), are compatible with faint RV Tauri variables or bright semi-regular ones. The bulk of the measured periods is in the range 100–250 days and thus they are compatible with this hypothesis. It is also remarkable that the bulk of the CMD of Fig. 1 is very similar (taking into account the distance moduli) to that obtained by Melchior et al. (2000a) for variables detected using the pixel method in the LMC, which are interpreted as “Long Timescale and Long Period Variables”.

In the two following sections, we describe work on variable stars for which we could identify the type, namely rather long period pulsating objects and novae.

5.2. Periodic pulsating stars

Bright pulsating stars are generally divided into Cepheid-like and long period red variables (called hereafter LPVs). In the part of the colour magnitude diagram we investigate here, as well as in the 100–200 days period interval, these two kinds of variable stars overlap. However, since LPVs have generally periods greater than 100 days, we can expect the shortest periods variable stars in our catalogue to be Cepheid-like (example Fig. 2). Actually, some of the light curves are reminiscent of the asymmetric shape observed for cepheids (Stetson 1996). On the other hand, stars with periods greater than 200 days are presumably LPVs (e.g., Figs. 8–10).

Early work on M 31 Cepheids was done by Baade & Swope (1963, 1965) with the Palomar telescope to calibrate the M 31 distance. Modern CCD studies were made (Freedman & Madore 1990; Magnier et al. 1997; Mochejska et al. 2000) and the period luminosity for M 31 population I Cepheids is by now well established for several colours including *R*. Figure 11 shows that our shorter period variable stars are at least 2 mag fainter than M 31 population I cepheids (1990). They thus clearly belong to the population II type (Wallerstein & Cox 1984). Actually, the shortest period star of Table 2 lies very near the location of LMC RV Tauri stars, for which, unfortunately the long period range is not known (Alcock et al. 1998). The

light-curve of one star of this type in M 31 has been recently presented (Riffeser et al. 2001); however, our data sampling is not good enough to uncover the characteristic double-wave shape of RV Tauri stars in our light curves.

LPV (Period > 100 days) are classically divided into Miras and semi-regular variables (Habbing 1996); the Miras are classically those with amplitude larger than 2.5 mag. Microlensing surveys have brought an impressive number of new variables, infrared data (both near-infrared from the ground and far-infrared from space) Hipparcos and HST observations have completely renewed our knowledge on pulsating stars and the two classes are now considered as overlapping (Bedding & Zijlstra 1998; Lebzelter et al. 2002). Optical light curves and optical colours of LPVs have instigated new studies (Delaverny et al. 1997; Mennessier et al. 1997). Figure 11 shows that our stars are generally brighter than what is expected for a Mira (Clayton & Feast 1969). Since semi-regular variables are rather brighter than Miras, they should be the more frequent in our catalog and Table 2. Yearly windowing has certainly biased our selection towards near-one-year periods. We checked that half the derived periods did not give plausible fits. C004 (Fig. 8), C072 (Fig. 9) and A001 (Fig. 10) are particularly convincing. A001 presents a 3 mag amplitude light curve and is very likely to be a Mira. However, these objects are more than one magnitude brighter than what is expected for galactic or LMC Miras with the same period. We cannot exclude neither peculiar LPVs nor that our phase undersampling led us to a false (too long) period determination.

5.3. Novae

M 31 novae have been systematically searched for more than 70 years (Hubble 1929; Arp 1956; Rosino et al. 1989; Ciardullo et al. 1987; Shafter & Irby 2001). The relation between their magnitude and rate of decline is now established (Capaccioli et al. 1989) and can be compared to those of other galaxies (Della Valle & Livio 1995). In addition to being interesting in their own right, novae are of cosmological interest in the context of Type Ia supernova precursors and the nature of supersoft sources (Kahabka & van den Heuvel 1997).

Twelve nova candidates were detected by AGAPE. These objects are so bright at maximum that some were found visually and we can expect detection of all novae which occurred in our field during AGAPE observation periods and for which 3 points of the light curve were measured.

Table 4 gives the data for our 12 nova candidates named as NOV n (first column) where n goes from 01 to 12. Column 2 gives the catalogue identification, so that coordinates can be found in Table 3 (see also Sect. 4.1). The next three columns give R_{\max} , the date of the observed maximum brightness given in (J-2449624.5) days, and the year of occurrence. Figure 12 shows the light curves for these twelve nova candidates. Only six of the novae were observed at maximum. NOV02, 04, 06, 08, 11 were observed by Shafter & Irby (2001). NOV08 and NOV09 present very small brightness decreases and are peaked around $R = 18$. They could correspond to the wing part of bright nova light-curves, as NOV06 could be as well.

Table 4. Characteristics of the 12 novae candidates named as NOV n (first column) where n goes from 01 to 12.

#	Id	R_{\max}	Date	year
NOV01	AGPV0391	17.14	006.88	1995
NOV02	AGPV0393	17.24	399.83	1995
NOV03	AGPV0392	16.07	303.10	1995
NOV04	AGPV0583	18.47	771.89	1996
NOV05	AGPV0843	16.80	335.14	1995
NOV06	AGPV1187	18.22	301.12	1995
NOV07	AGPV1439	18.41	006.96	1994
NOV08	AGPV1186	17.65	443.93	1995
NOV09	AGPV1576	17.56	006.97	1994
NOV10	AGPV1577	17.82	321.10	1995
NOV11	AGPV1578	16.35	320.94	1995
NOV12	AGPV1579	14.98	444.82	1995

Actually AGAPE observed NOV08 during 1995's season and Shafter & Irby (2001) in 1996, when it was the brightest object of their survey. NOV08 could thus be a recurrent nova. Also one nova appears to be observed at 1.9'' from NOV12 by Ciardullo et al. (1987) in 1985 (their nova 27, see Sect. 4.2). NOV12 could thus be a recurrent nova.

Now, the Chandra source CXOM31 J004242.1+411218. (Kaaret 2002) is located 0.7'' from NOV09 (see Sect. 4.2). NOV09 was observed in 1994's AGAPE runs and Chandra observations were performed in 2001 October. Origo et al. (2001) describe which can be expected as X-ray emission from classical and recurrent novae. Most classical and recurrent novae emit hard X-rays in the first months after the outburst and during 2–10 years, with peak X-ray luminosity of a few times 10^{33} erg/s. On the other hand, the supersoft X-ray phase of novae which is also relatively short lived (up to 10 years) is observed only for up to 20% of novae according to 108 objects analysed contained in the ROSAT archive. This phase corresponds to very soft spectra (equivalent temperatures in the range 15–80 eV) and high luminosity (10^{36} – 10^{38} erg/s bolometric luminosity) (Kahabka & van den Heuvel 1997). CXOM31 J004242.1+411218. has a luminosity of 0.4×10^{36} erg/s and thus cannot be associated with the hard emission of an M 31 nova. On the other hand, scaling an estimate presented by Kaaret (2002), a blackbody spectrum with a temperature of 80 eV would imply an unabsorbed luminosity of 1.2×10^{36} erg/s in the 0.1–10 keV band for HRC count rate of this source. Lower temperatures would require higher unabsorbed luminosities to produce the same observed count rate. The luminosity of CXOM31 J004242.1+411218 is thus consistent with that of a supersoft source. Unfortunately it is a rather faint source and no spectral information is available. No other observations with Chandra (Kong et al. 2002) nor XMM Newton (Di Stefano et al. 2003) report this source. It may be because it is near the limiting sensitivity of these observatories, but supersoft sources are also known to be transient sources (Trudolyubov et al. 2001). Up to now only 2 X-ray sources were suggested as being possible supersoft ones linked with M 31 nova outbursts (Nedialkov et al. 2002; Kaaret 2002;

Di Stefano et al. 2003). In Sect. 4.2 we discuss another possible identification of a supersoft source with an AGAPE catalogue variable object.

An important question concerning novae is about their location; both their radial distribution and population to which they belong is still in debate. First Ciardullo (1987) showed that there is not a central hole in their distribution as could be derived from (Arp 1956) observations. The debate is to know whether novae belong to the bulge or to the disk. Ciardullo et al. (1987) concluded that the novae were a bulge population and suggested that the binary central excess in M 31 (novae, X-ray sources) could come from ejection from globular clusters. Aurière et al. (1992) argued that actually all X-ray sources and novae in the bulge of M 31 were sufficiently near globular clusters that they could have been ejected from. Shafter & Irby (2001) also support a main bulge origin though they cannot exclude a bulge fraction as low as 50%. On the other hand, Hatano et al. (1997) propose that the M 31 novae belong mainly to the disk population. They made a study of the positions of 191 novae observed at Mount Wilson and Asiago for which B photometry at maximum was available, and used a model of extinction in order to interpret asymmetry along the minor axis. Studies of other systems show that young populations may be more efficient than old populations to produce novae. Now, it is classically considered that disk dominated galaxies produce more novae, which are brighter and with faster decline, than bulge dominated galaxies which produce fainter and slower ones (Della Valle et al. 1994; Della Valle & Livio 1998). This scenario has population-synthesis code support (Yungelson et al. 1997). Our sample of 12 novae candidates does not show any asymmetric distribution along the minor axis of M 31. However, it is too small to bring a significant contribution to the debate and it will be useful when joined to those from other M 31 surveys.

6. Conclusions

We have presented the first catalogue of variable stars located in (or at least towards) the bulge of M 31. The 1579 objects include 12 nova candidates, and mainly semi-periodic or periodic objects. Their magnitudes, colours and long period of variation suggest that our sample is dominated by a mixture of Long Period Variables (LPV) and type 2 Cepheid like variables (RV Tauri). The AGAPE data thus can address the question of the M 31 bulge variable star population and particularly the nature of its bright component as well as its spatial distribution.

It is worth mentioning that both novae and the pulsating periodic variable stars are of cosmological interest as luminosity candles and distance calibrators. Also, the nova population is related to supersoft X-ray sources which may be progenitors of Type I supernovae. We propose both the detection of two recurrent novae and the possible optical identification of two supersoft X-ray sources candidates observed with Chandra. This work will be complemented and extended using data from the new large field survey POINT-AGAPE performed with the Isaac Newton Telescope (An et al. 2004).

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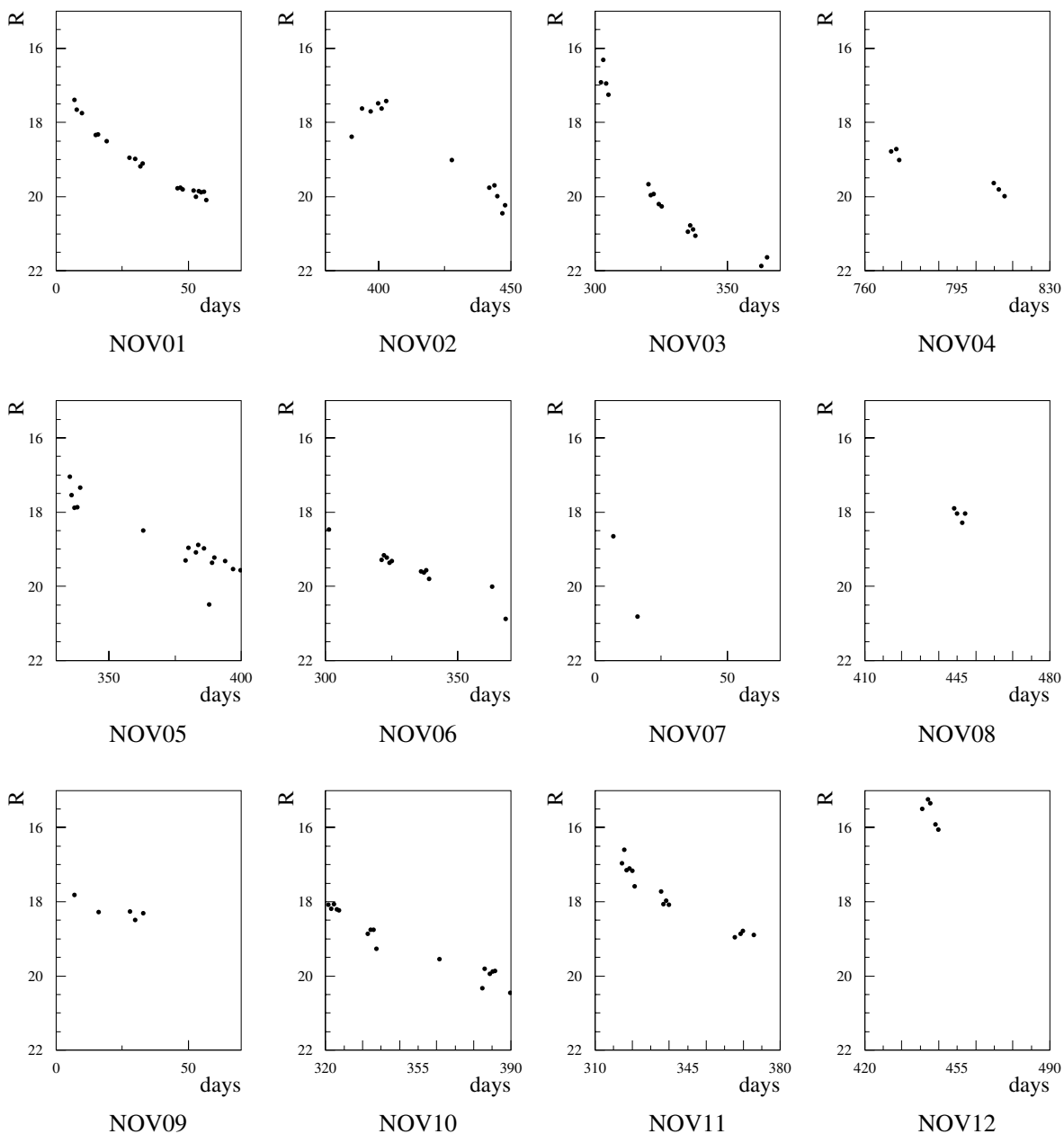


Fig. 12. Light curves of the 12 novae candidates detected in the Pic-du-Midi AGAPE survey.

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