

Evolution of the T Tauri star population in the Lupus association[★]

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ABSTRACT

Aims. In a recent study, we derived individual distances for 109 pre-main sequence stars that define the Lupus kinematic association of young stars. Here, we use these new distances to derive the masses and ages of Lupus T Tauri stars with the aim of better constraining the lifetime of their circumstellar disks.

Methods. Using the photometric and spectroscopic information available in the literature, we computed the photospheric luminosity of 92 T Tauri stars in the Lupus association. Then, we estimated their masses and ages from theoretical evolutionary models. Based on Monte Carlo simulations and statistical tests, we compare the mass and age distribution of the classical T Tauri stars (CTTS) and weak-line T Tauri (WTTS) in our sample.

Results. We show that the CTTSs are on average younger than the WTTSs and that the probability that both T Tauri subclasses are drawn from the same mass and age parental distribution is very low. Our results favor the scenario proposed earlier for the Taurus-Auriga association, where the CTTSs evolve into WTTSs when their disks are fully accreted by the star. Based on an empirical disk model, we find that the average disk lifetime for the T Tauri stars in the Lupus association is $\tau_d = 3 \times 10^6 (M_*/M_\odot)^{0.55}$ yr.

Conclusions. We find evidence that the average lifetime of the circumstellar disks in the Lupus association is shorter than in the Taurus-Auriga association and discuss the implications of this result.

Key words. stars: pre-main sequence – circumstellar matter – stars: formation

1. Introduction

T Tauri stars (TTs) are late-type low-mass ($M \lesssim 2 M_\odot$) pre-main sequence stars that are usually associated with molecular clouds. They tend to be found in young associations (Joy 1945; Herbig 1962; Bertout 1989). They are classified into one of two types: (i) the classical T Tauri stars (CTTSs) that show strong emission lines indicative of accretion from a circumstellar disk; and (ii) the weak-line T Tauri stars (WTTSs) that exhibit enhanced magnetic activity and show no signs of accretion. To explain the existence of both subgroups, it has been hypothesized for a long time (Walter et al. 1988; Bertout 1989) that the CTTSs become WTTSs when their disks dissipate. Although this is not a problem for the older population ($t > 10$ Myr) of WTTSs, this picture implies that the timescale of disk dissipation should be a few Myr to justify the CTTSs and WTTSs coexisting (see, e.g., Fig. 15 of Kenyon & Hartmann 1995) in the Hertzsprung-Russell diagram (HRD). Indeed, previous studies (Haisch et al. 2001; Takagi et al. 2014) indicate that the disk lifetime in different clusters is typically in the range 3–6 Myr. However, Fedele et al. (2010) demonstrate in a more detailed study that mass accretion and dust dissipation in protoplanetary disks occur on a typical timescale of 2.3–3.0 Myr (see Fig. 4 of their paper), which implies a rapid disk evolution.

Some progress in determining the lifetime of circumstellar disks of TTs was achieved by Bertout et al. (2007), who used the kinematic parallaxes derived from a kinematic study of the Taurus-Auriga T association (Bertout & Genova 2006) and

calculated the stellar luminosities, masses, and ages to readdress the question of the relationship between CTTSs and WTTSs. As a result, they obtained, for the first time, the lifetime of a circumstellar disk in terms of the mass of the parent star, $\tau_d = 4 \times 10^6 (M/M_\odot)^{0.75}$ yr, and demonstrated that the observed mass and age distributions of the TTs in the Taurus-Auriga association can be explained by assuming that the CTTSs evolve into WTTSs when their disks are fully accreted by the star.

While supporting the current scenario of disk evolution for the Taurus-Auriga T association, this first result needs to be confirmed and compared to other star-forming regions (SFRs). One reason for extending this study to other SFRs is that the origin of the initial mass function remains unclear, but is expected to depend on the intrinsic properties of each SFR, such as clustering, mass segregation, and binarity (Bonnell et al. 2007). Moreover, the evolution of circumstellar (protoplanetary) disks depends on environmental criteria (Rosotti et al. 2014), and a comparison with other SFRs could provide important clues for constraining the timescale of planet formation in different environments. For example, Mariñas et al. (2013) conclude that the dependence of disk fraction on stellar mass in Taurus and the NGC 2264 open cluster is different (see Fig. 10 of their paper), and they argue that the lower disk fractions in NGC 2264 might be due to the much higher stellar density than in Taurus.

In the following, we perform a study similar to the one by Bertout & Genova (2006) for the TTs in the Lupus SFR. The Lupus complex of molecular clouds is located in an environment surrounded by the more massive stars from the nearby Scorpius-Centaurus OB association (de Zeeuw et al. 1999). This makes

[★] Table 1 is available in electronic form at <http://www.aanda.org>

the Lupus SFR an interesting target for comparative studies with Taurus-Auriga, because of the different environmental and dynamical effects (e.g., solar winds and UV radiation) caused by the existence of hot stars in this region. In a recent work, [Galli et al. \(2013, hereafter Paper I\)](#) investigated the kinematics of the Lupus SFR and performed a membership analysis. They identified 109 pre-main sequence stars that define the comoving association of the Lupus SFR (hereafter, the Lupus association) based on a new convergent point search method ([Galli et al. 2012](#)) and a modified version of the k-NN method. Then, they used stellar proper motions and radial velocities to calculate individual distances for all group members. In this paper, we use these newly derived kinematic distances to accurately determine the masses and ages of the TTSs in the Lupus association and constrain the lifetime of their circumstellar disks.

2. Physical properties of the comoving stars

The sample of stars used in this study is based on the 109 comoving stars of the Lupus association identified in Paper I. One star, SStc2dJ154013.7-340142, which was included in our sample as a YSO candidate member of the Lupus clouds (see [Comerón et al. 2009](#)), is excluded from the upcoming analysis, because it turned out to be a background giant star (see [Comerón et al. 2013](#)). Our sample consists of 38 CTTSs (or CTTSs systems), 67 WTTSs (or WTTS systems), and 3 Herbig Ae/Be stars (HAeBe). These numbers correct those given in Sect. 7.1 of Paper I, where some objects were either misidentified or have been reclassified in the meantime. The results and conclusions given in Paper I are obviously not affected by this minor correction. We restricted our analysis in subsequent sections to the TTSs in the Lupus association (i.e., CTTSs and WTTSs) and excluded the HAeBes in our sample that appear in small numbers and are not statistically significant for this study.

The TTS subclass (CTTS or WTTS) from the stars in our sample were taken from the literature (see Table 1 for references), and no attempt has been made to reclassify them. However, we compiled the $H\alpha$ equivalent widths (EW) of our targets and calculated the near-infrared excess emission that is indicative of optically thick accretion disks to investigate whether the given TTS subclass was consistent with other diagnostics. We measure the excess emission by

$$\Delta(H - K) = (H - K) - (H - K)_0, \quad (1)$$

where $(H - K)$ is the dereddened color, and $(H - K)_0$ the intrinsic color of the star (see [Hughes et al. 1994](#); [Wichmann et al. 1997a](#)).

We used the standard limit of 10 \AA (see, e.g., [Appenzeller & Mundt 1989](#)) to distinguish between CTTSs ($EW(H\alpha) \geq 10 \text{ \AA}$) and WTTSs ($EW(H\alpha) < 10 \text{ \AA}$), together with the limit of $\Delta(H - K) \geq 0.09$ found by [Edwards et al. \(1993\)](#) for a sample of TTSs that show spectroscopic evidence of accretion. Figure 1 (upper panel) illustrates our analysis and clearly separates the accreting TTSs (i.e., CTTSs) from the non-accreting TTSs (i.e., WTTSs). Alternatively, [White & Basri \(2003\)](#) use the $H\alpha$ EW with a different cutoff that depends on the spectral type to distinguish between CTTSs and WTTSs. They propose that a TTS is indeed a CTTS if $EW(H\alpha) \geq 3 \text{ \AA}$ for K0-K5 stars, $EW(H\alpha) \geq 10 \text{ \AA}$ for K7-M2.5 stars, $EW(H\alpha) \geq 20 \text{ \AA}$ for M3-M5.5 stars, and $EW(H\alpha) \geq 40 \text{ \AA}$ for M6-M7.5 stars. We applied this classification scheme to the stars with spectral type later than K0 in our sample as shown in Fig. 1 (lower panel). We modified the first spectral type interval slightly to include two K6 stars in our

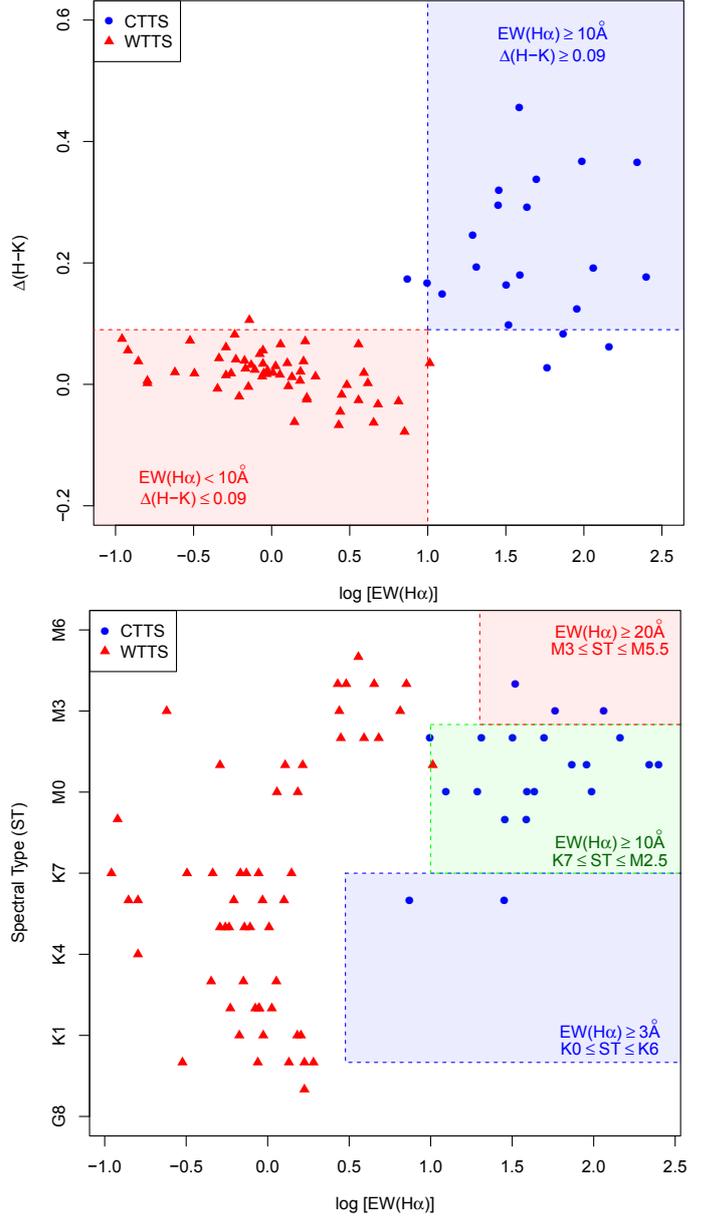


Fig. 1. $H\alpha$ EWs against the infrared excess emission $\Delta(H - K)$ (upper panel) and spectral type (lower panel) to distinguish between CTTSs and WTTSs.

sample and provide a TTS subclass to all stars. Our analysis presented in Fig. 1 not only confirms the TTS subclass found in several sources in the literature, but also gives us more confidence when discussing the different properties of these two subgroups in the rest of this paper. Only one star (Sz 76) would be misidentified by this alternative classification (see Fig. 1), because its $EW(H\alpha) = 10.3 \text{ \AA}$ (see [Hughes et al. 1994](#)) is marginally above the cutoff limit for its spectral type. However, Sz 76 has been classified recently as a WTTS by [Wahhaj et al. \(2010\)](#) and a Class III source by [López Martí et al. \(2011\)](#). Moreover, its excess emission $\Delta(H - K)$ is clearly beyond the cutoff limit of 0.09 between accreting and non-accreting TTSs (see [Edwards et al. 1993](#)). We therefore retain Sz 76 as a WTTSs in our study.

The first step toward accurately computing the photospheric luminosity of stars is to derive their visual extinction. We computed the stellar extinction A_V from the color excess in $(V - I_C)$,

($V - R_C$), and ($J - K$) for the WTTs in our sample, following the procedure described in Sect. 4.2 of [Kenyon & Hartmann \(1995\)](#). We set the extinction to zero when the derived A_V estimates yield a non-physical negative value or a non-significant result (i.e., $A_V < \sigma_{A_V}$). We took the mean of all non-zero A_V estimates as the final result with the standard deviation as the formal uncertainty. We used a total-to-selective extinction ratio R_V equal to 3.1 ([Savage & Mathis 1979](#)) and adopted the reddening relations for the various colors from [Bessell & Brett \(1988\)](#). The extinction for the CTTSs in our sample is estimated from the ($R_C - I_C$) color excess with $A_V = 4.76 \cdot [(R_C - I_C) - (R_C - I_C)_0]$ as described in Sect. 2 of [Cieza et al. \(2005\)](#).

To gain more confidence in our results, we decided to compute our A_V estimates again by using the same dereddening method for CTTSs and WTTs. In this alternative approach we used only the optical color excess in ($V - I_C$), ($V - R_C$), and ($R_C - I_C$) for all stars in the sample (i.e., CTTSs and WTTs). We concluded that both strategies yield compatible results and that our choice of using different methods to deredden CTTSs and WTTs (as described above) has no significant impact later on the determination of stellar masses and ages.

We used the individual parallaxes given in Tables 6 and 7 of Paper I to derive the photospheric luminosity of the Lupus comoving stars. The stellar luminosities for the CTTSs in our sample were calculated from the I_C band, because the standard procedure of using the J flux systematically overestimates their luminosities (see discussion in [Cieza et al. 2005](#)). For the WTTs, we derived the stellar luminosities from the V flux when the I_C flux was not available. We verified that both procedures yield equivalent results within 5% when both measurements are available. We used the intrinsic colors, bolometric corrections, and temperatures for each spectral type given by [Kenyon & Hartmann \(1995\)](#). The uncertainties on the derived luminosities take the various sources of photometric errors into account, but they are strongly dominated by the error budget of parallaxes. No optical photometry is available for two of our WTTs (RXJ1546.7-3459 and RXJ1612.3-4012). We thus estimate their luminosity from the J flux given by the 2MASS catalog ([Cutri et al. 2003](#)). Another star, Sz 102, was removed from the sample because it is believed to be associated with bipolar jets and an edge-on disk that makes our A_V estimate very imprecise ([Krautter 1986](#); [Hughes et al. 1994](#)). Moreover, its spectral type is not well-defined, and our analysis underestimates the luminosity, putting it below the zero-age main sequence (ZAMS).

To better compare our results with those derived for the Taurus-Auriga association, we used the same completeness limits as [Bertout et al. \(2007\)](#) by removing the stars (or stellar components) that have derived luminosities lower than $0.15 L_\odot$ and spectral types later than M4. These values correspond to $V \simeq 16$ mag and $T_{\text{eff}} \simeq 3300$ K. The final sample of stars with known photometry and spectral type that fulfill the above-mentioned criteria contains 30 CTTSs and 62 WTTs. Physical properties of the Lupus stars discussed in this section are given in Table 1. Eighteen stars in our sample were identified as binary systems in the literature. The photometry of these stars was corrected for binarity with the information (mainly the flux ratio of the components) provided by [Reipurth & Zinnecker \(1993\)](#), [Brandner et al. \(1996\)](#), [Ghez et al. \(1997\)](#), [Guenther et al. \(2001\)](#), [Torres et al. \(2006\)](#), and [Romero et al. \(2012\)](#).

3. Revisiting the CTTSs and WTTs in Lupus

In the following we use the stellar luminosities and effective temperatures of the stars to place them in the HRD and discuss the

Table 2. Mean masses and ages of the TTSs in the Lupus association derived from the [Siess et al. \(2000\)](#) models.

Sample	Stars	Mean M/M_\odot	Median M/M_\odot	Mean $\log t$ (t in Myr)	Median $\log t$ (t in Myr)
CTTS	30	0.54 ± 0.04	0.48	6.32 ± 0.24	6.28
WTTs (on-cloud)	33	0.69 ± 0.07	0.59	7.05 ± 0.67	6.55
WTTs (off-cloud)	29	1.09 ± 0.07	1.10	7.19 ± 0.28	7.11
WTTs (all stars)	62	0.88 ± 0.05	0.81	7.12 ± 0.48	6.87

age and mass distributions of the Lupus association. We derive the mass and age of the stars by interpolating between the evolutionary tracks and isochrones of the pre-main sequence stars models computed by [Siess et al. \(2000\)](#). Among the various models that exist in the literature, we chose the [Siess et al. \(2000\)](#) grid of pre-main sequence stars to allow for a direct comparison with the results in the Taurus-Auriga association (see Sect. 5) obtained by [Bertout et al. \(2007\)](#). The uncertainties affecting our mass and age estimates are estimated by defining the upper and lower limits for the stellar luminosities and effective temperatures within their error bars, and interpolating between the tracks and isochrones. We assumed an error of ± 100 K on the effective temperatures that roughly corresponds to an error of one subtype of spectral type for the late-type stars in our sample. Our mass and age results are presented in Table 1.

The HRD of the Lupus association illustrated in Fig. 2 shows that a few WTTs and CTTSs share the same region of the HRD. As already discussed in Paper I (see, e.g., Fig. 15 of that paper), the WTTs in our sample comprise an on-cloud population, which is located in the direction of the main star-forming clouds (Lupus 1-4) of the Lupus SFR, and a more dispersed off-cloud population surrounding the clouds. The question arises whether the two populations of WTTs and the CTTSs of the Lupus association exhibit the same mass and age distributions. Table 2 lists the average mass and age of each population for comparison and reveals the different properties of the CTTSs and WTTs in our sample.

We computed our mass and age estimates again by using the [D'Antona & Mazzitelli \(1997\)](#) tracks and the [Baraffe et al. \(1998\)](#) models to demonstrate that the different age and mass distributions are not an artifact from our choice of evolutionary model (see Table 3). Although evolutionary models computed by different groups yield different mass and age estimates for a given star, we note from Tables 2 and 3 that the results obtained with the various models employed in this work are consistent between themselves. This allows us to confirm that (i) the CTTSs are on average younger than the WTTs; and (ii) the off-cloud stars in the Lupus association are on average older than the on-cloud stars.

We performed a Kolmogorov-Smirnov (KS) and a Wilcoxon rank sum test to investigate the null hypothesis that the age and mass distributions of the CTTSs and WTTs in the Lupus association derive from the same distribution. The results of this analysis are given in Tables 4 and 5. Adopting the significance level of $\alpha = 0.05$ (see, e.g., [Feigelson & Jogesh Babu 2012](#)), we conclude that the null hypothesis can be rejected in each case regarding stellar ages (see Table 4), which means that the age distributions of the CTTSs and WTTs in Lupus are not identical. On the other hand, our results with the statistical tests applied to the mass distribution of CTTSs and on-cloud WTTs are marginally above the adopted significance level in a few cases (see Table 5) and remain inconclusive.

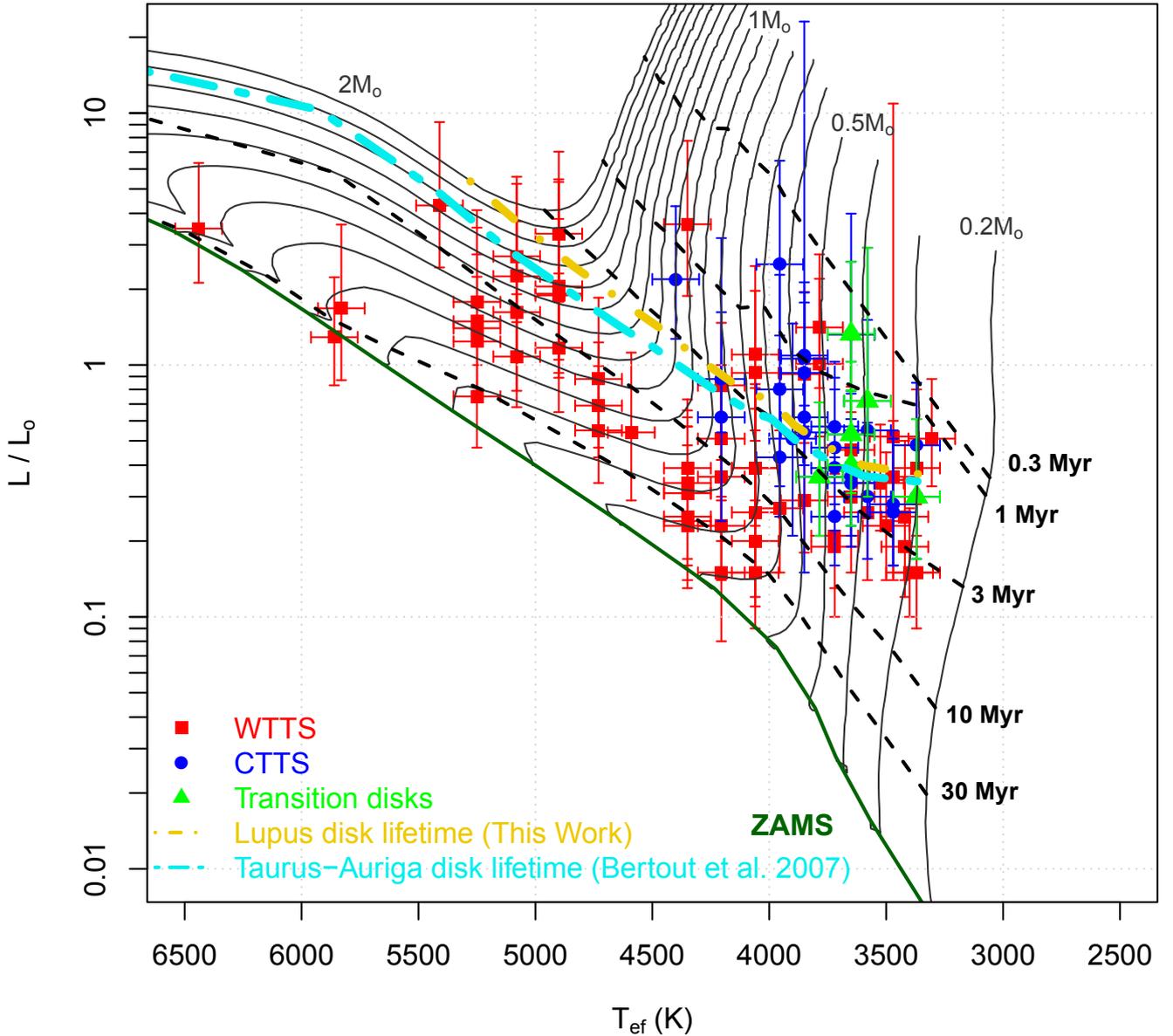


Fig. 2. HRD of the Lupus association. Red squares and blue circles denote, respectively, the WTTSs and CTTSs in our sample. The green triangles represent those stars with transition disks (see Sect. 4). The error bars indicate the 1σ uncertainties on the stellar luminosities and effective temperatures. The black solid lines and dashed lines indicate, respectively, the evolutionary tracks and isochrones computed by Siess et al. (2000) with $Y = 0.277$ and $Z = 0.02$. The green solid line denotes the zero-age main sequence (ZAMS) computed with that model. The dash-dotted lines show the disk lifetime τ_d for the TTSs in the Lupus (this work) and Taurus-Auriga (Bertout et al. 2007) associations.

Furthermore, we performed Monte Carlo simulations to assess the robustness of this finding and the errors that affect our mass and age estimates. To do so, we constructed a total of 1000 synthetic samples of the Lupus association by varying, within their error bars, the effective temperatures and the remaining observables (parallaxes, magnitudes, and visual extinction) used to compute the stellar luminosities of each star (as described in Sect. 2), assuming that these variables are normally distributed. Then, we calculated the mass and age of each star from the Siess et al. (2000) models and performed a Kolmogorov-Smirnov and a Wilcoxon rank sum test on the mass and age distributions of the CTTSs and WTTSs for each synthetic dataset. The result of this analysis is the mass and age distributions for all synthetic realizations of the Lupus comoving stars (see Fig. 3), and the histogram of probabilities that result from the statistical tests (see Fig. 4). It is apparent from Fig. 3

that the on-cloud and off-cloud WTTSs define different age distributions despite exhibiting similar kinematic properties (as discussed in Paper I). The mass distribution of both TTS subclasses shows that there is a lack of CTTSs at higher masses, while the WTTSs display a wide range of masses (see also Wichmann et al. 1997a).

To investigate whether the different mass and age distributions of the TTSs in the Lupus association could be an artefact of the procedure used in Sect. 2 to compute the stellar luminosities, we utilized a different set of tables to convert the magnitudes and spectral types to luminosities and effective temperatures. We built a control sample by using the bolometric corrections and temperatures tabulated by Pecaute & Mamajek (2013) to recompute the stellar luminosities. We then repeated the Monte Carlo simulations with our control sample as described above. The results of the KS- and Wilcoxon statistics applied to the control

Table 3. Mean masses and ages of the TTSS in the Lupus association derived from the D’Antona & Mazzitelli (1997) and Baraffe et al. (1998) pre-main sequence stars evolutionary models.

Sample	D’Antona & Mazzitelli (1997)				Baraffe et al. (1998)			
	Mean M/M_{\odot}	Median M/M_{\odot}	Mean $\log t$ (t in Myr)	Median $\log t$ (t in Myr)	Mean M/M_{\odot}	Median M/M_{\odot}	Mean $\log t$ (t in Myr)	Median $\log t$ (t in Myr)
CTTS	0.40 ± 0.02	0.38	6.15 ± 0.21	6.10	0.63 ± 0.04	0.60	6.38 ± 0.20	6.35
WTTS (on-cloud)	0.61 ± 0.06	0.47	6.79 ± 0.59	6.42	0.72 ± 0.06	0.73	7.05 ± 0.45	6.80
WTTS (off-cloud)	1.00 ± 0.07	1.00	6.93 ± 0.35	6.79	0.99 ± 0.05	0.98	7.17 ± 0.28	7.09
WTTS (all stars)	0.80 ± 0.05	0.80	6.86 ± 0.46	6.63	0.85 ± 0.04	0.86	7.11 ± 0.36	7.04

Table 4. Results of the Kolmogorov-Smirnov (KS) and Wilcoxon rank sum (WRS) statistical tests applied to the age distributions of the TTSS in the Lupus association computed with different sets of isochrones.

Sample 1	Sample 2	Siess et al. (2000)		D’Antona & Mazzitelli (1997)		Baraffe et al. (1998)	
		p_{KS}	p_{WRS}	p_{KS}	p_{WRS}	p_{KS}	p_{WRS}
CTTS	WTTS (on-cloud)	2.13×10^{-3}	1.31×10^{-3}	3.06×10^{-4}	3.52×10^{-4}	1.65×10^{-3}	2.71×10^{-4}
CTTS	WTTS (off-cloud)	1.01×10^{-10}	2.01×10^{-9}	5.66×10^{-9}	4.27×10^{-8}	3.27×10^{-7}	1.24×10^{-5}
CTTS	WTTS (all stars)	4.19×10^{-8}	1.57×10^{-7}	3.95×10^{-8}	2.31×10^{-7}	1.66×10^{-6}	4.25×10^{-6}

Notes. We provide the p -value given by each test.

Table 5. Results of the Kolmogorov-Smirnov (KS) and Wilcoxon rank sum (WRS) statistical tests applied to the mass distributions of the TTSS in the Lupus association computed with different evolutionary tracks.

Sample 1	Sample 2	Siess et al. (2000)		D’Antona & Mazzitelli (1997)		Baraffe et al. (1998)	
		p_{KS}	p_{WRS}	p_{KS}	p_{WRS}	p_{KS}	p_{WRS}
CTTS	WTTS (on-cloud)	3.84×10^{-2}	2.13×10^{-1}	1.95×10^{-2}	7.72×10^{-2}	1.82×10^{-1}	2.95×10^{-1}
CTTS	WTTS (off-cloud)	1.42×10^{-8}	1.10×10^{-7}	9.25×10^{-10}	4.79×10^{-8}	9.88×10^{-6}	1.62×10^{-5}
CTTS	WTTS (all stars)	6.16×10^{-6}	2.39×10^{-4}	1.97×10^{-6}	4.20×10^{-5}	5.45×10^{-4}	2.28×10^{-3}

Notes. We provide the p -value given by each test.

sample are also shown in Fig. 4. These results confirm the different age distributions of the CTTSs and WTTSs in the Lupus association for our data and control samples.

To summarize, the statistical investigations reported above of the mass and age distributions of CTTSs and WTTSs contained in the Lupus association, using different conversions between observed data and stellar parameters, as well as the different evolutionary models, make us confident that these two subgroups are indeed different as far as ages are concerned. A similar conclusion was reached for the TTS population in the Taurus-Auriga association by Bertout et al. (2007).

4. T Tauri disk lifetimes in the Lupus association

In their study of the Taurus-Auriga SFR, Bertout et al. (2007) obtained a relationship between the disk lifetime and the mass of its parent star. We performed a similar study for the Lupus association, with the aim of investigating how the different environment of the Lupus SFR affects the evolution of protoplanetary disks. We thus assumed that evolution from CTTSs to WTTSs is caused by the accretion of their circumstellar disks. Following Bertout et al. (2007), the disk mass of a TTS is written as

$$M_d = \alpha (M_{\star}/M_{\odot})^{\beta}, \quad (2)$$

while the mass accretion rate in units of M_{\odot}/yr is given by

$$\dot{M}_{\text{acc}} = \gamma (M_{\star}/M_{\odot})^{2.1}, \quad (3)$$

where α , β , and γ are unknowns to be determined in our analysis. This last relationship agrees with previous studies (see, e.g., Muzerolle et al. 2003, 2005), but more recent work (Alcalá et al. 2014) indicates that the exponent in Eq. (3) might be somewhat lower for low-mass CTTSs. We will come back to this point below, but for the time being we use Eq. (3) to compare our findings to those of Bertout et al. (2007). The disk lifetime $\tau_d = M_d/\dot{M}_{\text{acc}}$ can then be written as

$$\log \tau_d = \log(\alpha/\gamma) + (\beta - 2.1) \log(M_{\star}/M_{\odot}). \quad (4)$$

We varied these parameters in the range of $5 \leq \log(\alpha/\gamma) \leq 8$ and $-4 \leq \beta \leq 4$ in steps of 0.01 to compute τ_d for each star with the stellar masses given in Table 1. Then, we compared this derived disk lifetime τ_d for each star with its isochronal age t_{\star} , also listed in Table 1. The model developed by Bertout et al. (2007) predicts that the star is a CTTS whenever $t_{\star} \leq \tau_d(M_{\star})$, or a WTTS if $t_{\star} > \tau_d(M_{\star})$. Finally, we compared the mass distributions of the modeled CTTSs (and WTTSs) with the observed CTTSs (and WTTSs) in our data sample and ran a KS-test to find the best match between these populations for a given combination of α , β , and γ . We adopted a probability threshold of 99.95% (i.e., $p_{KS} \geq 0.9995$) as in Bertout et al. (2007). Doing so, we find $\log(\alpha/\gamma) = 6.53 \pm 0.05$ and $\beta = 2.65 \pm 0.10$. Our analysis is consistent with an average disk lifetime given by

$$\tau_d \simeq 3 \times 10^6 (M_{\star}/M_{\odot})^{0.55(\pm 0.10)} \text{ yr}. \quad (5)$$

The position of the thus-derived disk lifetime of Lupus TTSS is plotted in the HRD of Fig. 2, together with the result obtained by Bertout et al. (2007) for the Taurus-Auriga SFR.

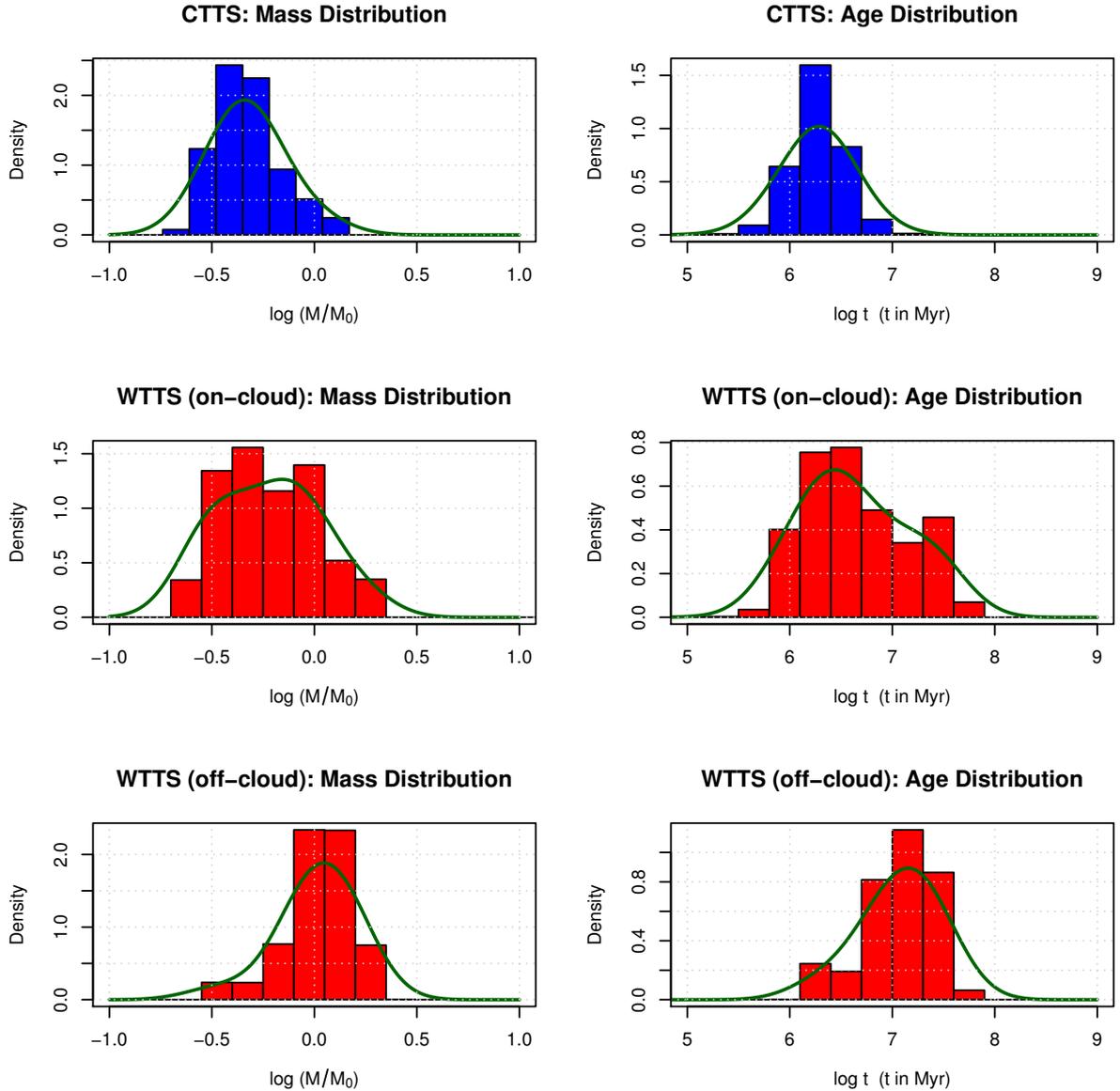


Fig. 3. Mass and age distributions of the CTTs and WTTs of the Lupus association obtained after 1000 Monte Carlo simulations. The green solid line indicates the kernel density estimator.

To assess the significance of this finding, we performed Monte Carlo simulations. We applied the above strategy to the 1000 synthetic realizations of the Lupus association, as well as to the control sample constructed in Sect. 3. The results, shown in Table 6, demonstrate that the simulations are consistent and confirm the values mentioned above. Simulations of the control sample returned lower values for the disk parameters, but these are statistically compatible with our result (within 3σ). Thus, our results remain valid if we use a different set of photometric tables and spectral type to effective temperature calibration for computing the stellar luminosities and the resulting mass and age estimates. The evolution of TTSs of the Lupus association is therefore consistent with the evolutionary scenario proposed by Bertout et al. (2007) for the Taurus-Auriga T association.

As mentioned above, Alcalá et al. (2014) have recently found that $\dot{M}_{\text{acc}} \propto M_{\star}^{1.8(\pm 0.2)}$ in a sample of Lupus low-mass accreting stars. If we use this revised accretion rate estimate in Eqs. (3) and (4) and run our calculations as described above, the previous value of $\beta = 2.65 \pm 0.10$ shifts to $\beta = 2.35 \pm 0.10$, and the

resulting exponent on the stellar mass in Eq. (5) is not affected. Similarly, if we use the mass accretion rate of $\dot{M}_{\text{acc}} \propto M_{\star}^{1.4(\pm 0.3)}$ derived by Venuti et al. (2014) for a sample of CTTs in the NGC 2264 open cluster, we find that $\beta = 2.65 \pm 0.10$ shifts to $\beta = 1.95 \pm 0.10$. In both cases, the different value obtained for β affects the disk mass given by Eq. (2) in our disk model. However, the combination of the different exponent β on the disk mass with a different mass accretion rate \dot{M}_{acc} yields the same exponent on the disk lifetime in Eq. (5). Thus, our result for the average disk lifetime remains the same.

Recently, Daemgen et al. (2013) have investigated the impact of binary companions on the evolution of circumstellar disks and concluded that they evolve faster in binary systems. In this context, we decided to remove the binaries (and multiple systems) from our sample to inspect this point. Doing so, we recalculated the disk lifetime in Lupus and find $\log(\alpha/\gamma) = 6.56 \pm 0.03$ and $\beta = 2.69 \pm 0.06$. Thus, the average disk lifetime excluding binaries from the sample is $\tau_{\text{d}} \simeq 4 \times 10^6 (M_{\star}/M_{\odot})^{0.59(\pm 0.06)}$ yr. This result implies that the average disk lifetime of the association is

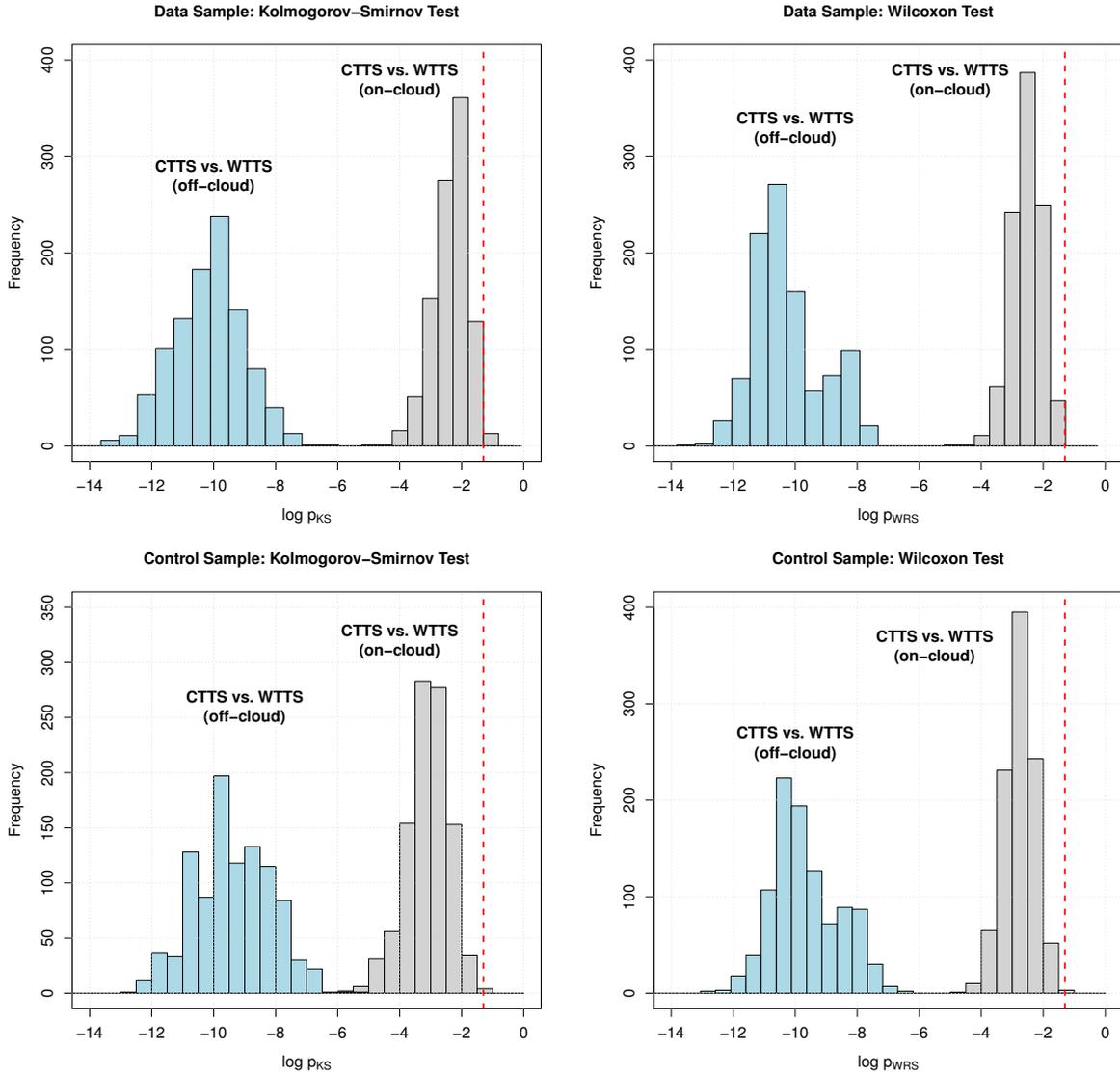


Fig. 4. Histogram of probabilities (i.e., p -value) obtained from the Kolmogorov-Smirnov (*left panels*) and Wilcoxon tests (*right panels*) for the age distributions of the CTTSs and WTTSs in our data sample (*upper panels*) and control sample (*lower panels*) after 1000 Monte Carlo realizations. The red dashed line indicates the significance level of $\alpha = 0.05$.

Table 6. Mean and median values of the disk parameters (α , β , and γ) obtained via Monte Carlo simulations after 1000 realizations of the data and control samples.

	$\log(\alpha/\gamma)$		β	
	Mean	Median	Mean	Median
Data sample	6.52 ± 0.05	6.52	2.51 ± 0.10	2.52
Control sample	6.35 ± 0.05	6.36	2.21 ± 0.09	2.22

likely to be shorter when binaries (and multiple systems) are included in our analysis. However, it is important to mention that both results are still statistically compatible within their error bars. For a more direct comparison with the results obtained by Bertout et al. (2007) in the Taurus-Auriga region (see Sect. 5), we prefer our first solution, because the binary stars were not excluded in their study.

Despite the success of this empirical disk model in explaining the different age distributions of the CTTSs and WTTSs in the Lupus association, we stress that the conclusions are only meaningful in a statistical sense and that our model may lead to a false classification in a few cases. Three

WTTSs (RX J1559.9-3750, RX J1609.4-3850, and RX J1612.0-3840) were misidentified as CTTSs in our analysis because of an inconsistency between the observed isochronal age and the computed lifetime of their disks at the 1σ level. These misidentified WTTSs are among the youngest stars in our sample.

Another interesting point in our study is the existence of six stars (Sz91, Sz95, Sz96, Sz112, RX J1609.9-3923, and RX J1608.5-3847) in our sample that were first identified as CTTSs and later reclassified as transition disk objects (see Romero et al. 2012; Tsukagoshi et al. 2014). We have labeled them with the mention “CTTS/TD” in Table 1 to distinguish them from the remaining CTTSs. Romero et al. (2012) claim that RX J1609.9-3923, the youngest transition disk object in the sample, is a “triple system with tight binary components consistent with two equally bright objects”. Later in the text, they state that the source has an uncertain classification because its spectral energy distribution resembles a CTTS of M spectral type. Taking only the position of RX J1609.9-3923 into account in the HRD, our analysis suggests that this source is more likely a CTTS. The remaining transition disk objects in the sample intermingle among CTTSs and WTTSs in the HRD, and their positions are

fully consistent with the disk lifetime locus that roughly separates both species (see Fig. 2).

5. Comparison with the Taurus-Auriga association

Both the Lupus and Taurus-Auriga associations discussed in this paper have been identified by using the convergent point search method applied to proper motion data in our previous studies (Bertout & Genova 2006; Galli et al. 2013) and the physical properties of the comoving stars were computed in the same completeness limits as described in Sect. 2. In this context, it is interesting to note that the Lupus association seems to be older than the Taurus-Auriga association. It is also apparent (see Fig. 2) that the distribution of CTTs in Lupus is shifted toward lower masses and late spectral types, while in Taurus-Auriga they display a wider range of masses and effective temperatures (see, e.g., Fig. 1 of Bertout et al. 2007). These first results had already been suggested in previous studies by using an average distance to the Lupus SFR to compute the stellar ages (and masses) for all stars and comparing different samples of TTSs in these SFRs (Hughes et al. 1994; Wichmann et al. 1997a). The individual parallaxes from Paper I allowed us to unambiguously confirm this scenario by comparing two samples of TTSs that have been investigated with the same selection criteria.

Bertout et al. (2007) derived an average disk lifetime of $\tau_d \approx 4 \times 10^6 (M_\star/M_\odot)^{0.75}$ yr for the TTSs in the Taurus-Auriga association that contrasts with the result of $\tau_d \approx 3 \times 10^6 (M_\star/M_\odot)^{0.55}$ yr obtained in this paper for the Lupus association. Our heuristic disk model thus predicts different disk lifetimes for the two associations. First, the exponents defining the mass dependence are different, but the two values are still compatible with each other if one considers the uncertainties involved ($\beta = 2.65 \pm 0.10$ in Lupus and $\beta = 2.85 \pm 0.21$ in Taurus-Auriga). Second, regardless of the exact β value, we note that the standard deviation values of $\log(\alpha/\gamma)$ found from the Monte Carlo simulations of Lupus and Taurus-Auriga comoving stars lead to disjointed age distributions at the 1σ significance level for the $1 M_\odot$ stars of both regions (2.95–3.72 Myr for Lupus and 3.80–4.79 Myr for Taurus-Auriga). Although this result will need confirmation with more significant samples statistically, it gives a first indication that disk lifetimes can be different in different star-forming regions.

This is not the only indication that disk lifetimes are likely to be different in Taurus-Auriga and Lupus. For example, a comparative analysis of the HRD of each association reveals that the overlap between CTTs and WTTs is more important in Lupus than in Taurus-Auriga (see Fig. 2 of this paper and Fig. 1 of Bertout et al. 2007). According to the results for the Taurus-Auriga association obtained by Bertout et al. (2007), it is only the evolutionary status of the circumstellar disks that allows us to distinguish between both subgroups. In this context, the existing overlap between both species in the HRD could be indicative of a shorter disk lifetime for the TTS population in the Lupus association as illustrated in Fig. 2.

Another fact that seems to support the different disk lifetimes in each association is the age of the transition disk objects that show signs of disk dissipation and are presumably in an intermediate state between CTTs and WTTs with little or no detectable near-IR excess emission and significant far-IR excess (Calvet et al. 2005; Romero et al. 2012). When excluding RX J1609.9-3923 for the reasons discussed in Sect. 4, the average age of the remaining transition disks included in our sample (Sz91, Sz95, Sz96, Sz112, and RX J1608.5-3847) is only 1.9 ± 0.3 Myr. On the other hand, the average age of

the transition disks (DM Tau, GM Aur, LkCa 15, and DI Tau) in Taurus-Auriga is 5.3 ± 1.4 Myr (see Table 4 of Bertout et al. 2007), thus significantly older than the observed value for Lupus.

Martin et al. (1998) investigated the ratio of WTTs over CTTs in different regions of the ρ Ophiuchi molecular cloud complex, and they argue that the high WTTs/CTTS ratio observed in some fields could be the result of the strong UV radiation, stellar winds, or supernova explosions from nearby massive stars that act to shorten the lifetime of circumstellar disks (see also Walter et al. 1994). A similar argument could apply to the Lupus association, where the WTTs greatly exceed the number of CTTs as compared to the Taurus-Auriga association. The Lupus cloud complex is located in an environment disturbed by the existence of hot stars from the Scorpius-Centaurus OB association and their dynamical effects, while the low-mass SFR of Taurus-Auriga is more isolated and shows no (or weak) interaction with hot stars. In this context, the large number of WTTs in Lupus could therefore be caused by the early disappearance of the circumstellar disks in CTTs.

6. Conclusions

The newly derived individual parallaxes from our previous kinematic study of the Lupus association have been used in this paper to determine the photospheric luminosities and refine the masses and ages of the TTS population in this SFR. We investigated the mass and age distributions of CTTs and WTTs in the Lupus association and demonstrated that the probability of both subgroups being drawn from the same age distribution is very low. We also confirmed the different properties of on-cloud and off-cloud WTTs of the Lupus association using our individual parallaxes, which is one result that had been anticipated by Wichmann et al. (1997a) by using an average distance estimate of the Lupus region for all stars. We concluded that the different mass and age distributions in the Lupus association can be explained by assuming that a CTT evolves into a WTT when the disk is accreted by the central star. Such a result had already been proposed by Bertout et al. (2007) for the Taurus-Auriga association, and it was now confirmed for a different SFR.

Based on an empirical disk evolution model, we derived the average lifetime of a circumstellar disk in the Lupus association in terms of the stellar mass and compared our result with a similar study conducted for the Taurus-Auriga association. We argued that specific properties of the Lupus association, including the more evident overlap of CTTs and WTTs in the HRD, the age of transition disk objects, and the ratio of WTTs/CTTS, could be the result of an earlier disappearance of the circumstellar disks in CTTs implying that the timescale of disk dissipation in the Lupus association is shorter (than in Taurus-Auriga).

This study represents an important step toward a better understanding of the early stages of star formation and disk evolution in different environments. It will be extended to other nearby SFRs in the future to allow for a more detailed comparative analysis of the evolution of TTSs in different associations.

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Table 1. Physical properties of the Lupus comoving stars.

Star	ST	T_{eff} (K)	$EW(H\alpha)$ (Å)	A_V (mag)	L/L_{\odot}	Filter	M/M_{\odot}	$\log t$ (t in Myr)	T Tauri Subclass	$\log \tau$ (τ in Myr)	Ref.
RX J1508.8-3715	K5	4350	1.02	0.25 ± 0.17	$0.23^{+0.16}_{-0.09}$	V	0.80 ± 0.15	7.59 ± 0.26	WTTS	6.48 ± 0.14	1, 2
RX J1518.4-3738	K0	5250	1.91	0.32 ± 0.10	$0.75^{+0.54}_{-0.28}$	Ic	0.92 ± 0.18	7.52 ± 0.19	WTTS	6.51 ± 0.15	1, 2, 3
RX J1524.5-3652	K0	5250	1.35	0.44 ± 0.07	$1.24^{+1.00}_{-0.47}$	Ic	1.17 ± 0.27	7.29 ± 0.32	WTTS	6.57 ± 0.17	1, 2, 3
RX J1525.0-3604	K1	5080	0.94	0.53 ± 0.18	$2.70^{+2.89}_{-1.22}$	Ic	1.68 ± 0.45	6.77 ± 0.39	WTTS	6.65 ± 0.20	1, 2, 3
RX J1525.5-3613	K2	4900	0.59	0.51 ± 0.21	$1.17^{+1.14}_{-0.52}$	Ic	1.29 ± 0.34	7.06 ± 0.42	WTTS	6.59 ± 0.20	1, 2, 4
RX J1531.3-3329	G9	5410	1.68	1.54 ± 0.04	$4.30^{+4.91}_{-1.86}$	Ic	1.73 ± 0.54	6.83 ± 0.34	WTTS	6.66 ± 0.23	1, 2, 3
RX J1534.6-4003	K5	4350	0.78	0.33 ± 0.16	$0.25^{+0.32}_{-0.12}$	Ic	0.82 ± 0.16	7.50 ± 0.36	WTTS	6.48 ± 0.15	1, 2, 3
RX J1540.7-3756	K6	4205	0.93	0.00:	$0.36^{+0.43}_{-0.16}$	Ic	0.89 ± 0.15	7.05 ± 0.58	WTTS	6.50 ± 0.13	1, 2, 5
RX J1544.5-3521	K5	4350	0.51	0.59 ± 0.19	$0.31^{+0.42}_{-0.15}$	V	0.90 ± 0.25	7.36 ± 0.39	WTTS	6.50 ± 0.21	1, 2
Sz73	M0	3850	97.20	2.71 ± 0.11	$0.54^{+22.54}_{-0.39}$	Ic	0.57 ± 0.08	6.34 ± 0.58	CTTS	6.40 ± 0.11	2, 6
GQ Lup*	K7-M0	3955	38.60	0.00:	$2.51^{+3.96}_{-1.20}$	Ic	0.73 ± 0.13	5.82 ± 0.35	CTTS	6.45 ± 0.13	2, 6, 7
RX J1549.9-3629	K2	4900	1.06	0.28 ± 0.18	$1.92^{+3.41}_{-1.03}$	Ic	1.55 ± 0.55	6.80 ± 0.52	WTTS	6.63 ± 0.26	1, 2, 4
RX J1552.3-3819	K7	4060	0.68	0.00:	$0.39^{+0.42}_{-0.16}$	Ic	0.80 ± 0.12	6.87 ± 0.54	WTTS	6.48 ± 0.12	1, 2, 4
RX J1605.7-3905	K0	5250	0.30	0.15 ± 0.07	$1.49^{+2.63}_{-0.75}$	Ic	1.25 ± 0.46	7.20 ± 0.46	WTTS	6.58 ± 0.27	2, 3
F304	K6	4205	1.26	0.00:	$0.23^{+0.19}_{-0.09}$	Ic	0.79 ± 0.15	7.42 ± 0.36	WTTS	6.47 ± 0.14	1, 2, 4
RX J1608.5-3847	M2	3580	5.98	1.00 ± 0.07	$0.72^{+2.20}_{-0.42}$	Ic	0.40 ± 0.05	6.04 ± 0.23	CTTS/TD	6.31 ± 0.11	1, 2, 4, 8
RX J1610.0-4016	K2	4900	0.88	0.52 ± 0.28	$2.05^{+3.41}_{-1.17}$	Ic	1.59 ± 0.57	6.76 ± 0.51	WTTS	6.64 ± 0.26	1, 2, 3
SZ121	M3	3470	6.50	0.98 ± 0.18	$0.52^{+10.39}_{-0.38}$	Ic	0.31 ± 0.08	6.17 ± 0.31	WTTS	6.25 ± 0.20	2, 6
RX J1613.0-4004	K7	4060	0.74	1.34 ± 0.04	$0.26^{+0.58}_{-0.14}$	Ic	0.80 ± 0.13	7.15 ± 0.71	WTTS	6.48 ± 0.13	1, 2, 4
RX J1511.0-3252AB	K6	4205	0.62	0.00:	$0.83^{+0.64}_{-0.30}$	Ic	0.94 ± 0.12	6.53 ± 0.43	WTTS	6.51 ± 0.11	1, 2, 4
RX J1511.6-3550	K5	4350	0.58	0.00:	$0.39^{+0.27}_{-0.13}$	Ic	0.96 ± 0.15	7.23 ± 0.40	WTTS	6.52 ± 0.12	1, 2, 4
RX J1512.6-3417	K5	4350	–	0.56 ± 0.19	$0.31^{+0.31}_{-0.14}$	V	0.90 ± 0.19	7.36 ± 0.44	WTTS	6.50 ± 0.16	2, 9
HD 135127	F5	6440	4.13	0.10 ± 0.09	$3.48^{+2.86}_{-1.36}$	Ic	1.41 ± 0.08	7.30 ± 0.20	WTTS	6.61 ± 0.06	1, 2, 4
RX J1515.7-3332	K0	5250	1.68	0.37 ± 0.18	$1.40^{+1.36}_{-0.61}$	Ic	1.22 ± 0.30	7.22 ± 0.36	WTTS	6.58 ± 0.19	1, 2, 3
RX J1525.6-3537	K6	4205	0.14	0.24 ± 0.15	$0.51^{+0.49}_{-0.22}$	Ic	0.95 ± 0.11	6.88 ± 0.54	WTTS	6.52 ± 0.10	1, 2, 4
RX J1527.3-3603	K7-M0	3955	0.12	0.50 ± 0.21	$0.27^{+0.28}_{-0.12}$	V	0.70 ± 0.11	6.98 ± 0.60	WTTS	6.45 ± 0.12	1, 2
RX J1529.3-3737	M3	3470	0.24	0.00:	$0.36^{+0.24}_{-0.12}$	Ic	0.35 ± 0.06	6.24 ± 0.19	WTTS	6.28 ± 0.14	1, 2, 4
RX J1529.7-3628	K2	4900	0.89	0.46 ± 0.20	$1.89^{+1.91}_{-0.84}$	V	1.54 ± 0.34	6.81 ± 0.41	WTTS	6.63 ± 0.17	1, 2, 3
Sz65*	M0	3850	19.40	0.29 ± 0.11	$1.06^{+1.07}_{-0.46}$	Ic	0.56 ± 0.08	6.00 ± 0.36	CTTS	6.39 ± 0.12	2, 6, 10
RX J1539.7-3450	K4	4590	0.16	0.50 ± 0.21	$0.54^{+0.58}_{-0.25}$	V	1.04 ± 0.25	7.24 ± 0.49	WTTS	6.54 ± 0.18	1, 2
RX J1540.3-3426A	M3.5	3420	3.03	0.00:	$0.19^{+0.16}_{-0.07}$	Ic	0.31 ± 0.05	6.48 ± 0.27	WTTS	6.25 ± 0.14	1, 2, 11
SSTc2dJ154148.3-350145	M2.7	3500	–	0.00:	$0.23^{+0.22}_{-0.09}$	Ic	0.36 ± 0.07	6.45 ± 0.30	WTTS	6.28 ± 0.15	2, 11, 12
RX J1542.0-3601	K7	4060	0.32	0.23 ± 0.18	$0.20^{+0.20}_{-0.09}$	V	0.76 ± 0.12	7.34 ± 0.44	WTTS	6.46 ± 0.12	1, 2
RX J1544.0-3311*	K0	5250	0.87	0.65 ± 0.19	$1.78^{+1.75}_{-0.78}$	V	1.33 ± 0.36	7.11 ± 0.37	WTTS	6.60 ± 0.20	1, 2, 3
RX J1546.6-3618	K1	5080	0.67	0.58 ± 0.33	$1.62^{+2.10}_{-0.83}$	Ic	1.38 ± 0.46	7.03 ± 0.45	WTTS	6.61 ± 0.25	1, 2, 3
RX J1546.7-3459	M0	3850	1.53	0.00:	$0.29^{+0.24}_{-0.11}$	J	0.59 ± 0.09	6.74 ± 0.49	WTTS	6.40 ± 0.12	1, 2
RX J1547.1-3540	K5	4350	0.55	0.20 ± 0.18	$0.24^{+0.24}_{-0.10}$	V	0.80 ± 0.10	7.61 ± 0.31	WTTS	6.48 ± 0.10	1, 2
RX J1547.6-4018	K1	5080	1.60	0.35 ± 0.07	$1.08^{+0.83}_{-0.40}$	Ic	1.17 ± 0.28	7.24 ± 0.33	WTTS	6.57 ± 0.18	1, 2, 3
HM Lup	M3	3470	115.30	0.00:	$0.26^{+0.25}_{-0.10}$	Ic	0.34 ± 0.08	6.37 ± 0.27	CTTS	6.27 ± 0.17	2, 6
HN Lup*	M1.5	3650	49.60	2.14 ± 0.11	$1.32^{+2.67}_{-0.71}$	Ic	0.42 ± 0.09	5.88 ± 0.25	CTTS	6.32 ± 0.17	2, 6, 10
RX J1548.1-3452	M2.5	3525	2.75	0.00:	$0.34^{+0.24}_{-0.12}$	Ic	0.37 ± 0.05	6.29 ± 0.24	WTTS	6.29 ± 0.12	1, 2, 4
RX J1548.9-3513	K6	4205	0.16	0.51 ± 0.18	$0.15^{+0.15}_{-0.07}$	V	0.70 ± 0.11	7.79 ± 0.44	WTTS	6.44 ± 0.12	1, 2
Sz76	M1	3720	10.30	0.38 ± 0.18	$0.19^{+0.20}_{-0.09}$	Ic	0.47 ± 0.05	6.80 ± 0.55	WTTS	6.35 ± 0.10	2, 6, 8
HD 141277*	K1	5080	1.52	0.49 ± 0.34	$2.25^{+2.99}_{-1.17}$	Ic	1.57 ± 0.54	6.87 ± 0.46	WTTS	6.64 ± 0.25	1, 2, 3
RX J1550.7-3828	K7	4060	0.88	0.86 ± 0.30	$1.10^{+1.37}_{-0.55}$	Ic	0.75 ± 0.12	6.20 ± 0.47	WTTS	6.46 ± 0.13	1, 2, 4
Sz77*	M0	3850	12.40	0.00:	$0.62^{+0.46}_{-0.22}$	Ic	0.57 ± 0.08	6.27 ± 0.38	CTTS	6.39 ± 0.12	2, 6, 10
RX J1555.4-3338	K5	4350	–	0.00:	$0.34^{+0.23}_{-0.12}$	Ic	0.91 ± 0.15	7.35 ± 0.41	WTTS	6.51 ± 0.13	1, 2, 4
RX J1556.0-3655	M1	3720	73.60	0.14 ± 0.07	$0.47^{+0.42}_{-0.19}$	Ic	0.48 ± 0.05	6.28 ± 0.37	CTTS	6.35 ± 0.10	1, 2, 4
Sz82 *	M0	3850	39.00	0.00:	$0.93^{+0.84}_{-0.36}$	Ic	0.56 ± 0.08	6.06 ± 0.35	CTTS	6.39 ± 0.12	2, 6, 8, 10
Sz126*	K7-M0	3955	–	0.60 ± 0.13	$0.43^{+0.41}_{-0.18}$	Ic	0.68 ± 0.12	6.66 ± 0.54	CTTS	6.44 ± 0.14	2, 6, 10, 13

Notes. (*) For each star, we provide the spectral type, temperature, $H\alpha$ equivalent width, visual extinction, luminosity, photometric filter used to compute luminosities, the mass and age derived from the [Siess et al. \(2000\)](#) models, disk lifetime computed from Eq. (5), T Tauri subclass, and sources of photometric/spectroscopic information. The A_V estimate of 0.00: denotes a non-physical negative value or a non-significant result. Stars marked with “*” are binary (or multiple) systems (see Sect. 2).

References. (1) [Krautter et al. \(1997\)](#); (2) [2MASS \(Cutri et al. 2003\)](#); (3) [Torres et al. \(2006\)](#); (4) [Wichmann et al. \(1997a\)](#); (5) [Sartori et al. \(2003\)](#); (6) [Hughes et al. \(1994\)](#); (7) [Neuhäuser et al. \(2008\)](#); (8) [Wahhaj et al. \(2010\)](#); (9) [Wichmann et al. \(1997b\)](#); (10) [Ghez et al. \(1997\)](#); (11) [Comerón et al. \(2009\)](#); (12) [López Martí et al. \(2011\)](#); (13) [Valenti et al. \(2003\)](#); (14) [Brandner et al. \(1996\)](#); (15) [Merín et al. \(2008\)](#); (16) [Guenther et al. \(2001\)](#); (17) [Alcalá et al. \(2014\)](#); (18) [Comerón et al. \(2013\)](#); (19) [Cieza et al. \(2007\)](#).

Table 1. continued.

Star	ST	T_{eff} (K)	$EW(H\alpha)$ (Å)	A_V (mag)	L/L_{\odot}	Filter	M/M_{\odot}	$\log t$ (t in Myr)	T Tauri Subclass	$\log \tau$ (τ in Myr)	Ref.
Sz128*	M1.5	3650	9.90	1.86 ± 0.13	$0.36^{+0.41}_{-0.17}$	Ic	0.43 ± 0.10	6.36 ± 0.41	CTTS	6.33 ± 0.17	2, 6, 14
RXJ1558.9-3646	M1.5	3650	2.81	0.00:	$0.46^{+0.36}_{-0.17}$	Ic	0.43 ± 0.10	6.24 ± 0.31	WTTS	6.33 ± 0.17	1, 2, 4
CD-3610569	K3	4730	0.71	0.58 ± 0.08	$0.69^{+0.54}_{-0.26}$	Ic	1.10 ± 0.27	7.24 ± 0.39	WTTS	6.55 ± 0.18	1, 2, 3
RXJ1559.9-3750	M0.5	3785	0.51	0.32 ± 0.28	$1.01^{+1.19}_{-0.49}$	Ic	0.51 ± 0.10	5.98 ± 0.47	WTTS	6.37 ± 0.15	1, 2, 4
SSTc2dJ160000.6-422158	M3.7	3400	–	0.00:	$0.15^{+0.12}_{-0.05}$	Ic	0.30 ± 0.05	6.57 ± 0.31	WTTS	6.24 ± 0.13	2, 11, 12, 15
Sz131	M2	3580	31.90	2.48 ± 0.24	$0.55^{+0.96}_{-0.30}$	Ic	0.41 ± 0.06	6.12 ± 0.77	CTTS	6.31 ± 0.12	2, 6, 15
RXJ1601.9-3613	K3	4730	0.45	0.59 ± 0.21	$0.88^{+0.97}_{-0.41}$	Ic	1.20 ± 0.29	7.10 ± 0.48	WTTS	6.57 ± 0.18	1, 2, 4
EXLup	M0	3850	43.30	0.00:	$1.09^{+0.85}_{-0.40}$	Ic	0.56 ± 0.08	5.99 ± 0.31	CTTS	6.39 ± 0.12	2, 6
RXJ1603.8-3938*	K3	4730	1.13	0.43 ± 0.08	$0.55^{+0.44}_{-0.21}$	Ic	1.01 ± 0.25	7.35 ± 0.39	WTTS	6.53 ± 0.19	1, 2, 3, 16
HD 143978	G2	5860	3.60	0.00:	$1.29^{+0.94}_{-0.46}$	Ic	1.11 ± 0.07	7.88 ± 0.48	WTTS	6.56 ± 0.07	1, 2, 4
RXJ1605.5-3837	M1	3720	1.28	0.00:	$0.21^{+0.16}_{-0.08}$	Ic	0.50 ± 0.05	6.74 ± 0.42	WTTS	6.36 ± 0.09	1, 2, 4
HO Lup*	M1	3720	219.80	0.00:	$0.57^{+0.46}_{-0.21}$	Ic	0.48 ± 0.05	6.19 ± 0.31	CTTS	6.35 ± 0.10	2, 6, 10
Sz90*	K7-M0	3955	28.50	2.12 ± 0.40	$0.80^{+1.48}_{-0.47}$	Ic	0.66 ± 0.13	6.26 ± 0.65	CTTS	6.43 ± 0.15	2, 6, 10
Sz91*	M0.5	3785	95.90	1.93 ± 0.11	$0.36^{+0.35}_{-0.15}$	Ic	0.52 ± 0.10	6.51 ± 0.44	CTTS/TD	6.38 ± 0.15	2, 6, 10
RXJ1607.2-3839	K7	4060	1.40	0.29 ± 0.20	$0.93^{+1.04}_{-0.43}$	Ic	0.75 ± 0.11	6.29 ± 0.47	WTTS	6.46 ± 0.12	1, 2, 4
Sz95 *	M1.5	3650	10.20	1.29 ± 0.11	$0.40^{+0.39}_{-0.17}$	Ic	0.43 ± 0.10	6.30 ± 0.37	CTTS/TD	6.33 ± 0.17	2, 6, 15
RXJ1608.0-3857	M0	3850	1.14	0.24 ± 0.22	$0.92^{+1.03}_{-0.43}$	Ic	0.56 ± 0.08	6.06 ± 0.42	WTTS	6.39 ± 0.12	1, 2, 4
Sz96*	M1.5	3650	11.00	0.43 ± 0.11	$0.53^{+0.50}_{-0.22}$	Ic	0.44 ± 0.10	6.17 ± 0.33	CTTS/TD	6.33 ± 0.17	2, 6, 10
RXJ1608.3-3843	K7	4060	0.11	0.00:	$0.39^{+0.28}_{-0.14}$	Ic	0.80 ± 0.12	6.87 ± 0.44	WTTS	6.48 ± 0.12	1, 2, 4
Sz97	M3	3470	58.20	0.95 ± 0.11	$0.28^{+0.27}_{-0.12}$	Ic	0.34 ± 0.08	6.34 ± 0.28	CTTS	6.28 ± 0.17	2, 6
RXJ1608.6-3922	K6	4205	7.42	1.52 ± 0.07	$0.88^{+0.74}_{-0.34}$	Ic	0.93 ± 0.13	6.49 ± 0.47	CTTS	6.51 ± 0.11	1, 2, 4, 8
SSTc2dJ160853.2-391440	K7.8	3900	–	4.60 ± 0.43	$0.51^{+0.94}_{-0.30}$	Ic	0.61 ± 0.10	6.45 ± 0.70	CTTS	6.41 ± 0.13	2, 11, 12, 15
RXJ1608.9-3905	K2	4900	0.84	0.38 ± 0.18	$3.32^{+3.71}_{-1.54}$	Ic	1.86 ± 0.34	6.47 ± 0.40	WTTS	6.68 ± 0.14	1, 2, 4
RXJ1608.9-3945	M3.5	3420	2.69	0.16 ± 0.14	$0.25^{+0.25}_{-0.11}$	Ic	0.32 ± 0.06	6.36 ± 0.28	WTTS	6.26 ± 0.15	1, 2, 4
Sz111	M1.5	3650	145.20	0.00:	$0.34^{+0.27}_{-0.13}$	Ic	0.43 ± 0.10	6.38 ± 0.33	CTTS	6.33 ± 0.17	2, 6
Sz112	M4	3370	46.70	0.81 ± 0.11	$0.30^{+0.31}_{-0.13}$	Ic	0.30 ± 0.05	6.26 ± 0.29	CTTS/TD	6.24 ± 0.14	2, 6
V908Sco	M4	3370	33.00	0.00:	$0.48^{+0.37}_{-0.17}$	Ic	0.30 ± 0.04	6.18 ± 0.55	CTTS	6.24 ± 0.13	2, 6
Sz115	M4	3370	7.10	0.00:	$0.15^{+0.13}_{-0.06}$	Ic	0.29 ± 0.05	6.55 ± 0.30	WTTS	6.23 ± 0.15	2, 6, 15, 19
Sz134	M1	3720	90.50	0.00:	$0.25^{+0.18}_{-0.09}$	Ic	0.47 ± 0.05	6.64 ± 0.40	CTTS	6.35 ± 0.10	2, 6
RXJ1609.4-3850	M0.5	3785	1.64	0.13 ± 0.12	$1.41^{+1.34}_{-0.60}$	Ic	0.51 ± 0.08	5.86 ± 0.32	WTTS	6.37 ± 0.13	1, 2, 4
Sz116*	M1.5	3650	3.90	0.42 ± 0.24	$0.30^{+0.35}_{-0.15}$	Ic	0.43 ± 0.10	6.45 ± 0.43	WTTS	6.33 ± 0.18	2, 6, 15
Sz117	M2	3580	20.50	0.43 ± 0.11	$0.30^{+0.29}_{-0.13}$	Ic	0.39 ± 0.05	6.38 ± 0.32	CTTS	6.31 ± 0.12	2, 6
Sz118	K6	4205	28.30	4.19 ± 0.11	$0.62^{+2.57}_{-0.38}$	Ic	0.96 ± 0.14	6.73 ± 0.88	CTTS	6.52 ± 0.11	2, 6
RXJ1609.9-3923*	M1.5	3650	13.67	2.33 ± 0.07	$1.33^{+1.24}_{-0.54}$	Ic	0.42 ± 0.09	5.88 ± 0.28	CTTS/TD	6.32 ± 0.17	1, 2, 4
Sz119	M4	3370	4.50	0.53 ± 0.18	$0.39^{+0.41}_{-0.18}$	Ic	0.30 ± 0.05	6.18 ± 0.61	WTTS	6.25 ± 0.13	2, 6
Sz122	M2	3580	4.80	0.37 ± 0.18	$0.26^{+0.28}_{-0.12}$	Ic	0.39 ± 0.05	6.45 ± 0.38	WTTS	6.31 ± 0.12	2, 6
Sz123*	M1	3720	250.60	0.86 ± 0.11	$0.39^{+0.41}_{-0.17}$	Ic	0.48 ± 0.05	6.38 ± 0.43	CTTS	6.35 ± 0.10	2, 6, 10
RXJ1612.0-3840	K5	4350	0.72	0.61 ± 0.26	$3.62^{+4.14}_{-1.74}$	Ic	1.15 ± 0.25	5.87 ± 0.41	WTTS	6.56 ± 0.16	1, 2, 4
SSTc2dJ161207.6-381324	M4.5	3305	3.60	0.00:	$0.51^{+0.37}_{-0.18}$	Ic	0.25 ± 0.05	6.14 ± 0.64	WTTS	6.20 ± 0.15	2, 11, 12, 18
SSTc2dJ161243.8-381503	K4.8	4400	–	1.27 ± 0.12	$2.19^{+3.08}_{-0.92}$	Ic	1.20 ± 0.18	6.18 ± 0.45	CTTS	6.57 ± 0.12	2, 11, 12, 15
RXJ1614.4-3808	K7	4060	0.46	0.23 ± 0.17	$0.15^{+0.14}_{-0.06}$	V	0.70 ± 0.11	7.54 ± 0.35	WTTS	6.45 ± 0.12	1, 2
HD 147402	G3	5830	–	0.37 ± 0.28	$1.68^{+1.93}_{-0.81}$	Ic	1.14 ± 0.16	7.40 ± 0.16	WTTS	6.56 ± 0.11	2, 3, 8