

Size and disk-like shape of the broad-line region of ESO 399-IG20

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Received 6 January 2013 / Accepted 5 February 2013

ABSTRACT

We present photometric reverberation mapping of the narrow-line Seyfert 1 galaxy ESO 399-IG20 performed with the robotic 15 cm telescope VYSOS-6 at the Cerro Armazones Observatory. Through the combination of broad- and narrow-band filters we determine the size of the broad-line emitting region (BLR) by measuring the time delay between the variability of the continuum and the H α emission line. We use the flux variation gradient method to separate the host galaxy contribution from that of the active galactic nucleus (AGN), and to calculate the 5100 Å luminosity L_{AGN} of the AGN. Both measurements permit us to derive the position of ESO 399-IG20 in the BLR size – AGN luminosity $R_{\text{BLR}} \propto L_{\text{AGN}}^{0.5}$ diagram. We infer the basic geometry of the BLR through modeling of the light curves. The pronounced sharp variability patterns in both the continuum and the emission line light curves allow us to reject a spherical BLR geometry. The light curves are best fitted by a disk-like BLR seen nearly face-on with an inclination angle of $6^\circ \pm 3^\circ$ and with an extension from 16 to 20 light days.

Key words. galaxies: active – galaxies: Seyfert – galaxies: distances and redshifts – galaxies: individual: ESO 399-IG20 – quasars: emission lines

1. Introduction

The physical properties of the broad line region (BLR) in active galactic nuclei (AGN) have been extensively studied during the past three decades. Consensus has been reached that the luminosity variations of the continuum emitting region of the hot accretion disk produce variations of the broad emission lines with a delay due to the light travel time across the BLR. To date, the only method independent on spatial resolution is reverberation mapping (Blandford & McKee 1982; Peterson 1993), in which one measures the time delay or “echo” between the continuum variability and the variability observed in the broad emission lines. This method has been used successfully for determining the BLR size, kinematics and black hole mass (M_{BH}) of several Type-1 AGNs. RM provides basic M_{BH} estimates through the virial product $M_{\text{BH}} = f \cdot R_{\text{BLR}} \cdot \sigma_v^2 / G$, where G is the gravitational constant, $R_{\text{BLR}} = c \cdot \tau$ is the BLR size, σ_v is the emission-line velocity dispersion of the BLR gas and the factor f depends on the – so far unknown – geometry and kinematics of the BLR (Peterson et al. 2004 and references therein). Furthermore, as a widely used extrapolation of locally obtained RM results to higher redshift data, M_{BH} has been estimated using the BLR Size-Luminosity $R_{\text{BLR}} \propto L_{\text{AGN}}^{0.5}$ relationship (Koratkar & Gaskell 1991; Kaspi et al. 1996, 2000; Wandel et al. 1999; McGill et al. 2008; Vestergaard et al. 2011). However, a reliable M_{BH} estimation from the luminosity requires a considerable reduction of the $R - L$ dispersion present.

As the size of the BLR ranges from a few to several hundred light days, it is necessary to perform observational monitoring campaigns ranging from months to years in order to sufficiently sample the time domain of the echo. Most observational RM campaigns have used spectroscopic monitoring of the sources, which however requires large amounts of telescope time.

Recently, Haas et al. (2011) proposed photometric reverberation mapping (PRM) as an efficient method to determine the BLR size through the use of broad band filters to trace the AGN continuum variations and narrow-band filters to catch the BLR emission line response. Because the narrow-band collects both the emission line flux and the underlying continuum, the challenge is to extract the pure emission line light curve. In the dedicated case study of 3C 120, Pozo Nuñez et al. (2012) demonstrated that PRM reaches an accuracy similar to spectroscopic RM (Grier et al. 2012).

Another efficient approach to determine the BLR size has been proposed by Chelouche & Daniel (2012), through the analysis of the difference between the cross-correlation (CCF) and the auto-correlation (ACF) functions for suitably chosen broad band filters, one filter covering a bright emission line and one covering only the continuum. While this method does not place a strict requirement on the object’s redshift such that the emission line falls into the narrow-band filter (as in the case proposed by Haas et al. 2011), the use of broad band filters only (as proposed by Chelouche & Daniel 2012) limits the application of PRM to cases with a sufficiently strong emission line contribution in the respective filter used. Despite of this handicap, this method has been successfully applied to determine the BLR size for one low-luminosity AGN NGC 4395 (Edri et al. 2012) and one high- z luminous MACHO quasar (Chelouche et al. 2012) consistent with previous spectroscopic RM results.

A fundamental issue in both these methods is to separate the host-galaxy contribution from the total luminosity to calculate the AGN luminosity. An incorrect determination of the nuclear AGN luminosity (due to the contamination of the host component) causes an overestimation of the linear regression slope α . This has been demonstrated by Bentz et al. (2009), who presented the most recent compilation of host-subtracted

AGN luminosities for several reverberation-mapped Seyfert 1 galaxies obtained through host-galaxy modeling of high-resolution *Hubble* Space Telescope (HST) images. Compared to the previous slope of $\alpha = 0.7$ they determined an improved slope of $\alpha = 0.519^{+0.063}_{-0.066}$ close to $\alpha = 0.5$ expected from photoionization models of the BLR (Davidson & Netzer 1979). A different approach to estimate the host galaxy contribution is the flux variation gradient method (FVG, Choloniewski 1981; Winkler 1997). An advantage is that it does not require high spatial resolution imaging and can be applied directly to the monitoring data. Recently, FVG has been tested on PRM data (Haas et al. 2011; Pozo Nuñez et al. 2012). These tests show that by using a well defined range for the host galaxy slope (Sakata et al. 2010) it is possible to separate the AGN contribution at the time of the monitoring campaigns.

ESO 399-IG20 has been classified as a Narrow-Line Seyfert 1 (NLS1) galaxy (Véron-Cetty & Véron 2010). It has often been speculated that the lack of broad emission lines of – at least – part of the NLS1 population can be explained by face-on disk-like BLR geometries. Single-epoch spectroscopic investigations by Dietrich et al. (2005) show high-ionized gas orbiting at high velocities, producing strong broad profiles ($FWHM(H\beta) = 2425 \pm 121 \text{ km s}^{-1}$, see their Table 3 and Figs. 2–5) like in Broad-Line Seyfert 1 (BLS1) galaxies. Furthermore, they find a strong host-galaxy contribution with respect to the total observed 5100 Å continuum flux (~50%). Here, we present the first measurement of the BLR size and the host-subtracted AGN 5100 Å luminosity for ESO 399-IG20. Both results allow us to infer its position in the BLR size-Luminosity $R_{BLR} \propto L_{AGN}^{0.5}$ diagram. Additionally, pure broadband PRM (Chelouche & Daniel 2012) is applied and compared with the broad- and narrow-band PRM technique (Haas et al. 2011). Finally, the geometry of the BLR, whether spherical or an inclined disk, is inferred via light curve modeling.

2. Observations and data reduction

Broad-band Johnson *B* (4330 Å), Sloan-band *r* (6230 Å) and the redshifted $H\alpha$ (SII 6721 ± 30 Å at $z = 0.0249$) images were obtained with the robotic 15 cm VYSOS6 telescope of the Universitätssternwarte Bochum, located near Cerro Armazones, future location of the ESO Extreme Large Telescope (ELT) in Chile. Monitoring occurred between May 5 and November 18 of 2011, with a median sampling of 3 days.

ESO 399-IG20 lies at redshift $z = 0.0249$, therefore the $H\alpha$ emission line falls into the SII 6721 ± 30 Å narrow-band filter. Figure 1 shows the position of the narrow-band with respect to the $H\alpha$ emission line together with the effective transmission of the other filters used. The $H\alpha$ line is broader than the SII filter, however, simulations with different asymmetric and perfect symmetric cutting of the high-velocity line wings show negligible systematic effect introduced in the time delay (~2%), even considering the fact of only using one part of the wing profile. This effect will be discussed in depth in a forthcoming contribution (Bruckmann et al., in prep.).

Data reduction was performed using IRAF¹, in combination with SCAMP (Bertin 2006) and SWARP (Bertin et al. 2002) routines, in the same manner as described by Haas et al. (2012). Light curves were extracted using different apertures

¹ IRAF is distributed by the National Optical Astronomy Observatory, which is operated by the Association of Universities for Research in Astronomy (AURA) under cooperative agreement with the National Science Foundation.

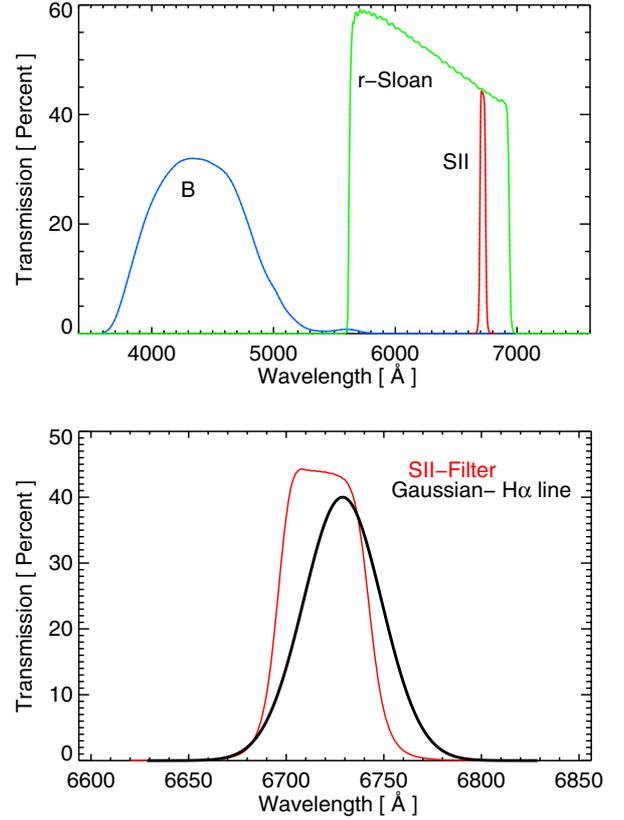


Fig. 1. Effective transmission of the VYSOS6 filters convolved with the quantum efficiency of the ALTA U16M CCD camera (top). The position of the narrow-band [SII] 6721 Å filter (red line) and the simulated redshifted Gaussian-shaped $H\alpha$ emission line using the parameters obtained by Dietrich et al. (2005) (black line) are shown in the bottom panel.

(5'', 7''.5, 15'' and 25'') in order to compare and trace the host and AGN contribution. The light curves for the different apertures are shown in Fig. 2. Additionally, for each aperture we determined the level of variability using the fractional flux variation introduced by Rodríguez-Pascual et al. (1997):

$$F_{\text{var}} = \frac{\sqrt{\sigma^2 - \Delta^2}}{\langle f \rangle} \quad (1)$$

where σ^2 is the flux variance of the observations, Δ^2 is the mean square uncertainty, and $\langle f \rangle$ is the mean observed flux. The variability statistics is listed in Table 2. Selecting a small aperture (5'') we are only considering a small portion of the total flux, which remains heavily dependent on the quality of the PSF, whereas bigger apertures (15'' and 25''), enclose more contribution of the host-galaxy and the results are more sensitive to sky background contamination resulting in larger uncertainties. We found that 7''.5 around the nucleus is the right size aperture which maximizes the signal-to-noise ratio (S/N) and delivers the lowest absolute scatter for the fluxes. Thus, we use this aperture for the further analysis of the BLR size and AGN luminosity. Although the light curves in Fig. 2 have not yet been corrected for the host galaxy contribution, one can see that the fractional variation (F_{var}) and the ratio of the maximum to minimum fluxes (R_{max}) are higher in the *B*-band than in the *r*-band for which the host-galaxy present a greater contribution. A different case can be seen in the SII light curves, which are mostly dominated by the contribution of the $H\alpha$ emission line resulting

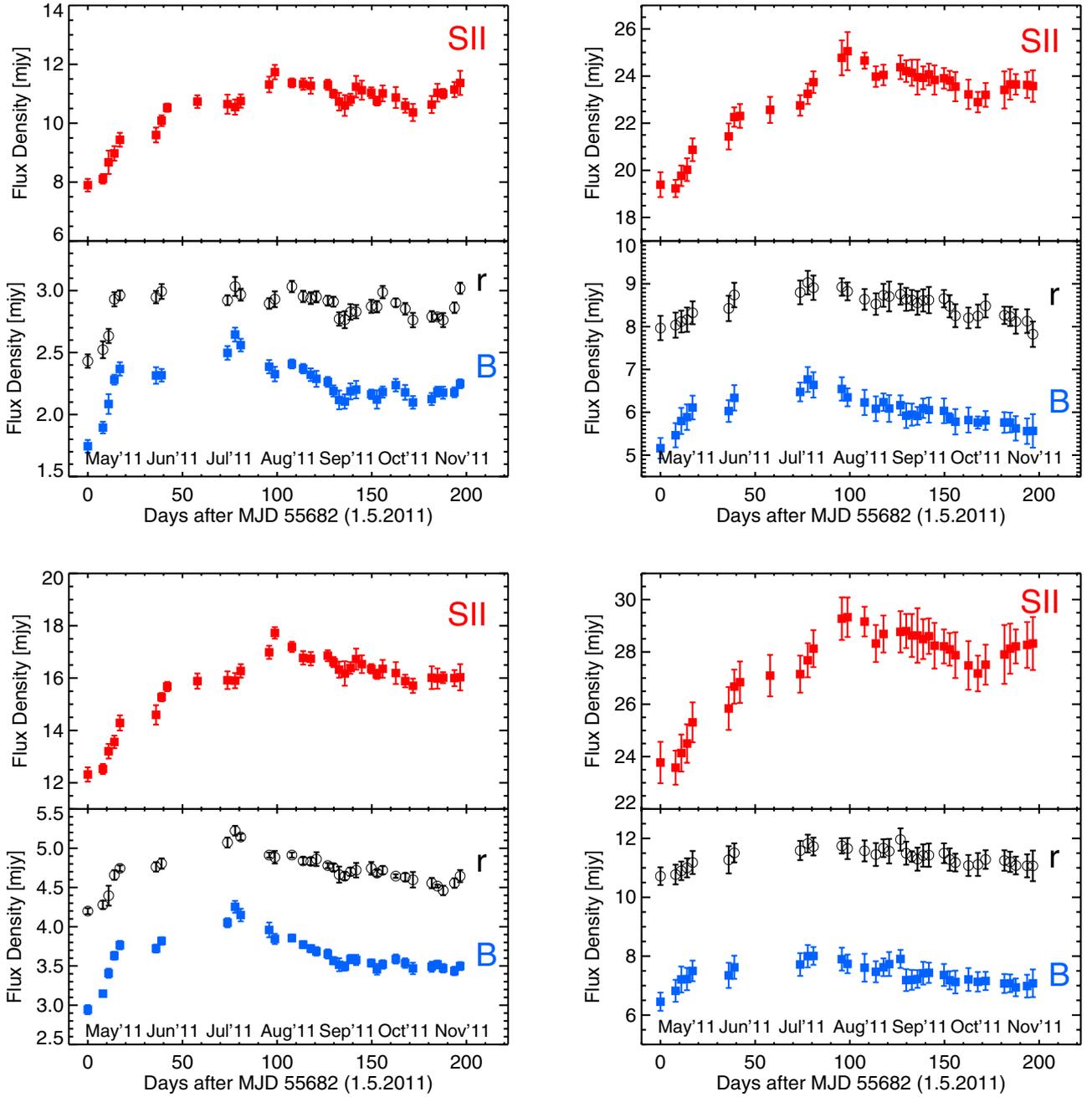


Fig. 2. *B*, *r*-Sloan and SII-bands light curves for different apertures (5'' upper left, 7.5'' lower left, 15'' upper right and 25'' lower right). The light curves have been corrected for foreground galactic extinction using the values given in Table 1.

in a higher amplitude of variability. As expected, for bigger apertures the *B* and *r* light curves become flatter with F_{var} and R_{max} smaller due to the larger contribution from the host-galaxy.

Non-variable reference stars located on the same field and with similar brightness as the AGN were used to create the relative light curves in normalized flux units. For the absolute photometry calibration we used reference stars from Landolt (2009) observed on the same nights as the AGN, considering the atmospheric extinction for the nearby site Paranal by Patat et al. (2011) and the recalibrated galactic foreground extinction presented by Schlafly & Finkbeiner (2011) obtained from the Schlegel et al. (1998) dust extinction maps. In Table 1, we give the characteristics of ESO 399-IG20. A summary of the photometric results and the fluxes in all bands are listed in Tables 3 and 4, respectively.

3. Results and discussion

3.1. Light curves and BLR size

The *B*-band, which is dominated by the AGN continuum, shows a strong flux increase (about 35%) between the beginning and the end of May. Afterwards, the two measurements obtained in June reflect a more gradual increase until a maximum is reached at the end of July. After this maximum, the fluxes decrease gradually (about 20%) until the end of September, and then the light curves become more constant until middle of November. Likewise, the *r*-band, which is dominated by the continuum but also contains a contribution from the strong $H\alpha$ emission line, follows the same features as the *B*-band light curve, although with a lower amplitude. In contrast to the continuum dominated

Table 1. Characteristics of ESO 399-IG20.

α (2000) ¹	δ (2000) ¹	z ¹	D_L ¹ (Mpc)	$B - V$ ² (mag)	M_{abs} ² (mag)	A_B ³ (mag)	A_V ³ (mag)
20:06:57.7	-34:32:58.0	0.024951	102.0	0.96	-19.7	0.409	0.309

Notes. ⁽¹⁾ Values from NED database; ⁽²⁾ Véron-Cetty & Véron (2010); ⁽³⁾ Schlafly & Finkbeiner (2011).

Table 2. Light curve variability statistics.

Aperture diameter (1)	<i>B</i> -band				<i>r</i> -band				SII-band			
	$\langle f \rangle$ (2)	σ (3)	F_{var} (4)	R_{max} (5)	$\langle f \rangle$ (6)	σ (7)	F_{var} (8)	R_{max} (9)	$\langle f \rangle$ (10)	σ (11)	F_{var} (12)	R_{max} (13)
5.0	2.24	0.03	0.02	1.52	2.86	0.02	0.01	1.25	10.56	0.85	0.08	1.49
7.5	3.63	0.06	0.03	1.45	4.71	0.04	0.01	1.24	15.80	1.55	0.10	1.44
15.0	5.98	0.12	0.01	1.31	8.47	0.09	0.01	1.15	23.06	2.33	0.09	1.30
25.0	7.35	0.14	0.01	1.24	11.3	0.09	0.01	1.11	27.51	2.42	0.08	1.24

Notes. Column 1 list the aperture in arc-sec, Cols. 2–5 list the mean flux, the standard deviation, the normalized excess variance and the ratio of the maximum to minimum fluxes for the continuum light curve. Similarly, Cols. 6–13 list the variability for the *r* and SII bands respectively. The standard deviation and the mean fluxes are expressed in units of mJy.

Table 3. Summary of the photometry results for 7"5 aperture.

<i>B</i> (mag)	<i>r</i> (mag)	SII (mag)	fB_{total} ¹ (mJy)	fr_{total} ¹ (mJy)	$f\text{SII}_{\text{total}}$ ¹ (mJy)
15.05–15.46	14.57–14.81	13.21–13.58	3.45 ± 0.06	4.87 ± 0.08	17.17 ± 0.34

Notes. ⁽¹⁾ fB_{total} , fr_{total} and $f\text{SII}_{\text{total}}$ refer to the mean of the total flux ranges during our monitoring.

broad-band light curves, the narrow SII-band light curve exhibits a flux increase which is stretched. For instance the prominent maximum in the *B*- and *r*-band at the end of July occurs in the SII-band in August, giving a first approximation for the H α time delay of 15–20 days. Note that the visual inspection of the *r*- and *B*-band light curves does not allow an approximation for the time delay.

As already discussed in previous PRM studies, the narrow-band contains a contribution of the varying continuum, which must be removed before applying cross correlation techniques (Haas et al. 2011; Pozo Nuñez et al. 2012). To determine this contribution, we used the SII and *r*-band fluxes, previously calibrated to mJy, as is shown with the flux-flux diagram in Fig. 3. The H α line is strong contributing, on average, about 70% of the total flux enclosed in the SII-band, while the continuum contribution (*r*-band) is about 30%. Following the usual practice of PRM, we construct a synthetic H α light curve by subtracting a third of the *r*-band flux ($\text{H}\alpha = \text{SII} - 0.3r$). The H α light curve was used afterwards to estimate the time delay. For this purpose, we used the discrete correlation function (DCF, Edelson & Krolik 1988) to cross-correlate the continuum and the synthetic H α emission line, taking into account possible bin size dependency (Rodríguez-Pascual et al. 1989). The centroid in the cross-correlation of *B*-band and H α show a time delay of 18.1 days, while the centroid in the cross-correlation between *B*-band and SII-band yields a time delay of 15.8 days. However, this is an expected result because of the more pronounced peak at zero lag due to the contamination by the continuum emission in the narrow-band filter. Both cross-correlation functions are shown in Fig. 4. Uncertainties in the time delay were calculated using the flux randomization and random subset selection method (FR/RSS, Peterson et al. 1998, 2004). From the observed light curves we create 2000 randomly selected subset light curves, each containing 63% of the original data points (the other

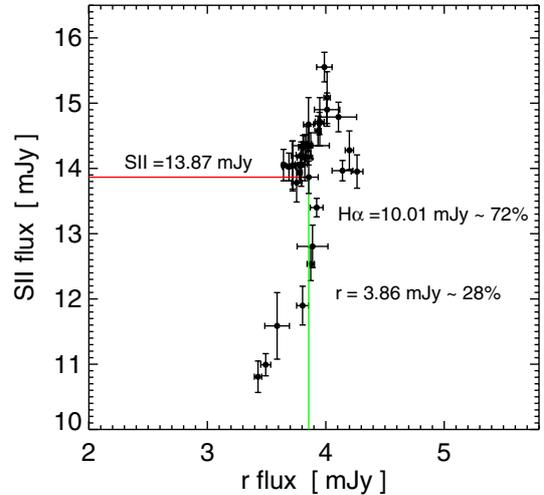


Fig. 3. Flux-flux diagram for the SII and *r* band measured using 7.5" aperture. Black dots denote the measurement pair of each night. The straight red and green lines represent the average flux in the SII and *r* band respectively. The data are as observed and not corrected for galactic foreground extinction.

fraction of points are unselected according to Poisson probability). The flux value of each data point was randomly altered consistent with its (normal-distributed) measurement error. We calculated the DCF for the 2000 pairs of subset light curves and the corresponding centroid (Fig. 5). From this cross-correlation error analysis, we measure a median lag of $\tau_{\text{cent}} = 18.7^{+2.5}_{-2.2}$. Correcting for the time dilation factor ($1+z = 1.0249$) we obtain a rest frame lag of 18.2 ± 2.29 days.

An alternative method to estimate the time delay, called Stochastic Process Estimation for AGN Reverberation (SPEAR), has been worked out recently by Zu et al. (2011).

Table 4. B , r , SII and $H\alpha$ fluxes corrected by extinction.

JD - 2 450 000	F_B (mJy)	F_r (mJy)	F_{SII} (mJy)	$F_{H\alpha}$ (mJy)
55 687.340	2.944 ± 0.058	4.199 ± 0.038	12.318 ± 0.274	11.058 ± 0.276
55 695.316	3.147 ± 0.036	4.279 ± 0.052	12.528 ± 0.194	11.244 ± 0.200
55 698.332	3.407 ± 0.058	4.396 ± 0.127	13.206 ± 0.281	11.887 ± 0.308
55 701.352	3.631 ± 0.051	4.659 ± 0.061	13.561 ± 0.239	12.163 ± 0.246
55 704.281	3.765 ± 0.054	4.744 ± 0.038	14.287 ± 0.291	12.863 ± 0.293
55 723.359	3.722 ± 0.051	4.763 ± 0.060	14.594 ± 0.371	13.165 ± 0.375
55 726.305	3.819 ± 0.047	4.805 ± 0.066	15.272 ± 0.161	13.830 ± 0.174
55 729.289	- ± -	- ± -	15.676 ± 0.178	15.676 ± 0.178
55 745.234	- ± -	- ± -	15.886 ± 0.291	15.886 ± 0.291
55 761.137	4.051 ± 0.058	5.073 ± 0.064	15.918 ± 0.339	14.396 ± 0.344
55 765.148	4.253 ± 0.076	5.223 ± 0.061	15.902 ± 0.291	14.335 ± 0.297
55 768.117	4.149 ± 0.080	5.143 ± 0.042	16.273 ± 0.258	14.730 ± 0.261
55 783.129	3.960 ± 0.094	4.913 ± 0.032	16.983 ± 0.252	15.509 ± 0.254
55 786.133	3.845 ± 0.065	4.885 ± 0.080	17.726 ± 0.220	16.260 ± 0.234
55 795.133	3.856 ± 0.043	4.913 ± 0.033	17.193 ± 0.194	15.719 ± 0.196
55 801.082	3.772 ± 0.043	4.838 ± 0.047	16.774 ± 0.258	15.322 ± 0.262
55 805.109	3.722 ± 0.047	4.833 ± 0.047	16.741 ± 0.258	15.291 ± 0.262
55 814.109	3.686 ± 0.054	4.861 ± 0.088	16.854 ± 0.210	15.395 ± 0.227
55 817.117	3.653 ± 0.058	4.780 ± 0.033	16.612 ± 0.194	15.178 ± 0.196
55 820.105	3.566 ± 0.047	4.755 ± 0.047	16.322 ± 0.307	14.895 ± 0.310
55 823.090	3.516 ± 0.083	4.664 ± 0.103	16.176 ± 0.468	14.776 ± 0.479
55 826.094	3.494 ± 0.054	4.645 ± 0.052	16.386 ± 0.226	14.992 ± 0.231
55 829.094	3.588 ± 0.054	4.701 ± 0.052	16.725 ± 0.404	15.314 ± 0.407
55 832.094	3.577 ± 0.065	4.720 ± 0.093	16.531 ± 0.339	15.115 ± 0.351
55 837.094	3.537 ± 0.047	4.744 ± 0.080	16.354 ± 0.178	14.930 ± 0.195
55 840.102	3.458 ± 0.069	4.683 ± 0.042	16.144 ± 0.178	14.739 ± 0.182
55 843.102	3.519 ± 0.054	4.720 ± 0.052	16.354 ± 0.339	14.938 ± 0.342
55 850.020	3.588 ± 0.058	4.645 ± 0.023	16.192 ± 0.420	14.798 ± 0.420
55 855.000	3.537 ± 0.061	4.631 ± 0.052	15.886 ± 0.242	14.496 ± 0.247
55 859.059	3.469 ± 0.072	4.598 ± 0.099	15.708 ± 0.274	14.328 ± 0.291
55 869.008	3.490 ± 0.065	4.556 ± 0.066	16.015 ± 0.436	14.648 ± 0.440
55 872.008	3.519 ± 0.051	4.518 ± 0.019	15.983 ± 0.387	14.627 ± 0.387
55 875.027	3.472 ± 0.054	4.462 ± 0.061	16.015 ± 0.210	14.676 ± 0.218
55 881.016	3.436 ± 0.054	4.556 ± 0.061	15.999 ± 0.307	14.632 ± 0.313
55 884.016	3.498 ± 0.051	4.645 ± 0.077	16.031 ± 0.500	14.637 ± 0.505

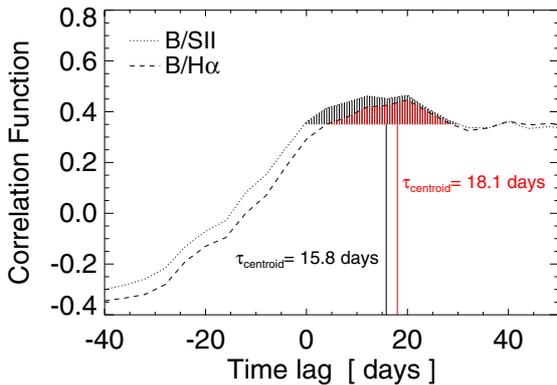


Fig. 4. Cross correlation of B and SII light curves (dotted line) and of B and $H\alpha$ light curves (dashed line). The error range ($\pm 1\sigma$) around the cross correlation was omitted for better viewing. The red and black shaded areas marks the range used to calculate the centroid of the lag (vertical red and black straight lines).

Through the modeling of the AGN light curves as a damped random walk (DRW; Zu et al. 2011, 2013 and reference therein), this method appears to be consistent with previous cross-correlation techniques (e.g. Grier et al. 2012, for spectroscopic reverberation data). Using the SPEAR code² on our light curves,

² <http://www.astronomy.ohio-state.edu/~yingzu/codes.html>

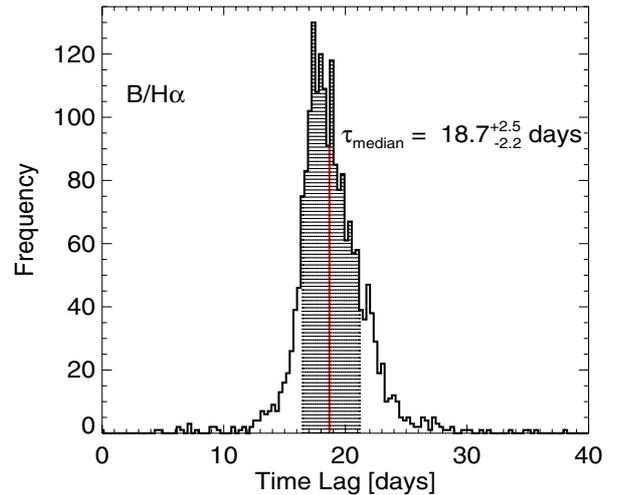


Fig. 5. Results of the lag error analysis. The histogram shows the distribution of the centroid lag obtained by cross correlating 2000 flux randomized and randomly selected subset light curves (FR/RSS method). The black area marks the 68% confidence range used to calculate the errors of the centroid (red line).

we obtain a time delay of $\tau_{\text{spear}} = 18.4^{+1.2}_{-1.0}$. After correction for time dilation the rest frame lag $\tau_{\text{rest}} = 17.9 \pm 1.1$ days, in agreement with the results from the DCF method. Figure 6 shows the SPEAR light curve models.

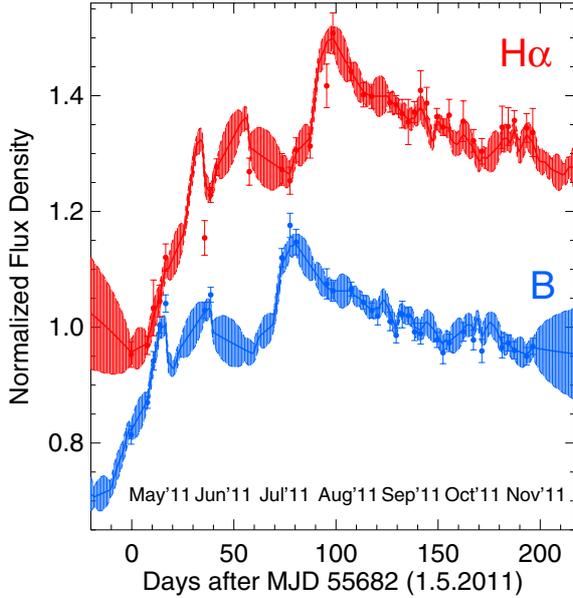


Fig. 6. Synthetic H α and continuum light curves. The solid red and blue lines show the H α and continuum models estimated by SPEAR respectively. The red and blue area (enclosed by the dashed line) represent the expected variance about the mean light curve model. The original H α light curve is computed by subtracting a scaled r curve from the SII curve and re-normalizing it to mean = 1 (red dots). The H α light curve is vertically shifted by 0.2 with respect to the continuum light curve (blue dots) for clarity.

3.2. Broad-band reverberation mapping approach

Recently, pure broad-band photometric RM has been introduced as an efficient alternative to measure the BLR size in quasars (Chelouche & Daniel 2012). In this method, the emission line is also measured using a broad-band filter, and the removal of the underlying continuum is performed in the correlation domain. The centroid of the lag is obtained by the subtraction of the ACF between the continuum (represented by one broad band filter enclosing only continuum emission) from the CCF between the continuum and the emission line (represented by one broad-band filter that contains a sufficiently strong emission line contribution). This method has been applied for two objects; the low-luminosity AGN NGC 4395 (Edri et al. 2012) and the high-redshift ($z = 1.72$) luminous MACHO quasar (13.6805.324) (Chelouche et al. 2012); in both cases a successful recovery of the time delay has been reported. Our ESO 399-IG20 has a strong H α emission line contributing to about 20% to the r -band flux (Fig. 3) and the light curves are well sampled. This makes ESO 399-IG20 an ideal object to perform a further test of the pure broad-band PRM method.

The B -band is mostly dominated by the continuum, and the r -band contains a strong H α emission line. This allows us to calculate the line-continuum cross-correlation function $CCF(\tau) = CCF_{Br}(\tau) - ACF_B(\tau)$ and to obtain the time delay (centroid). As shown in Fig. 7, the B/r cross-correlation exhibits two different peaks, one peak at zero lag (from the auto-correlation of the continuum) and one peak at lag ~ 3 days. Furthermore, a small but extended enhancement can be seen in the auto-correlation of the r -band at about 10–20 days, in comparison with the ACF of the B -band. This feature can be interpreted as a clear contribution from the H α emission line to the broad r -band filter. In fact, the cross-correlation shows a broad peak with a lag of 15.0 days as defined by the centroid in Fig. 7. To determine the lag uncertainties, we applied the FR/RSS method. Again

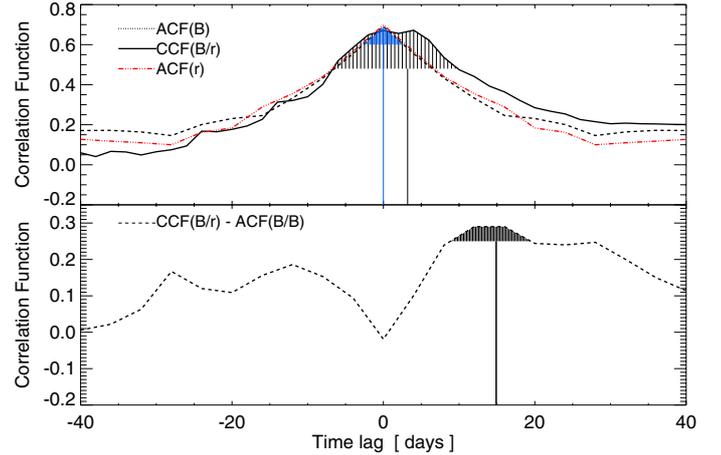


Fig. 7. Broad-band RM results. *Top:* the B -band ACF (dotted line) show a well defined centroid at zero lag (vertical blue line), while the B/r -band CCF shows the same peak at zero lag, a secondary peak at 3 days is also visible (vertical black line). In the same case as the B -band ACF, the r -band ACF (red line) show a peak at zero lag, however, one additionally faint peak is clearly visible between 10–20 days, in comparison with the ACF of the B -band. *Bottom:* the cross-correlation yields a centroid lag of 15.0 days (vertical black line). The red shaded area marks the range used to calculate the centroid.

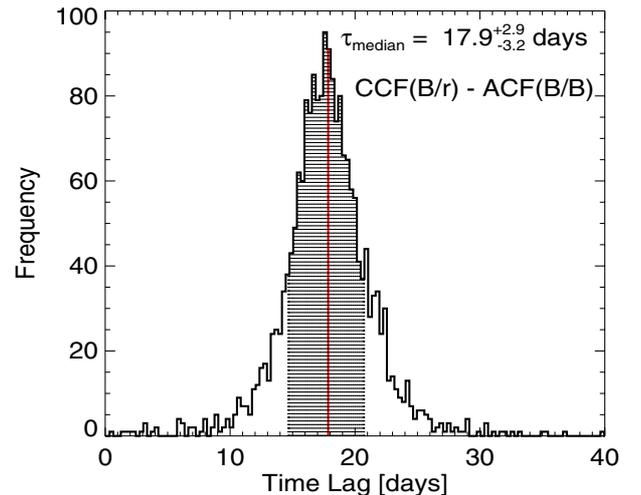


Fig. 8. Results of the lag error analysis of $CCF_{Br}(\tau) - ACF_B(\tau)$. The histogram shows the distribution of the centroid lag obtained by cross correlating 2000 flux randomized and randomly selected subset light curves (FR/RSS method). The black shaded area marks the 68% confidence range used to calculate the errors of the centroid (red line).

from the observed light curves we created 2000 randomly selected subset light curves, each containing 63% of the original data points, and randomly altering the flux value of each data point consistent with its (normal-distributed) measurement error. We calculated the DCF for the 2000 pairs of subset light curves and the corresponding centroid (Fig. 8). This yields a median lag $\tau_{\text{cent}} = 17.9^{+2.9}_{-3.2}$ days. Correcting for time dilation we obtain a rest frame lag $\tau_{\text{rest}} = 17.5 \pm 3.1$ days. This nicely agrees with the results from narrow-band PRM (18.2 ± 2.3 days for the DCF, and 17.9 ± 1.1 days for SPEAR)

3.3. Modeling the geometry of the broad-line region

We have inferred the geometry of the BLR through the direct modeling of the results obtained from PRM. Following

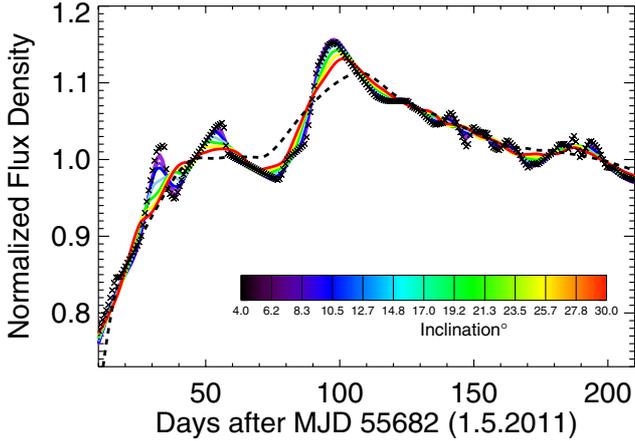


Fig. 9. BLR disk/sphere-model. Simulated $H\alpha$ light curves for a disk-like BLR geometry and different inclinations ($4 \leq i \leq 30^\circ$) are shown in different colors for illustration. The black dotted line is the $H\alpha$ simulated light curve for a spherical BLR model. A disk-like BLR model with inclination of 6° is able to reproduce the features of the original $H\alpha$ light curve (black crosses).

Welsh & Horne (1991), the time delay function for a spherically symmetric geometry of the BLR is:

$$\Delta(\tau) = \frac{r}{c}(1 + \sin \phi \cdot \cos \theta) \quad (2)$$

where r is the radius of the spherical shell, ϕ and θ the respective angles in the spherical coordinates. While for a Keplerian ring/disk structure the time delay function is:

$$\Delta(\tau) = \frac{r}{c}(1 - \sin i \cdot \cos \phi) \quad (3)$$

where r is the radius of the ring, i is the inclination ($0 \leq i \leq 90^\circ$) of the axis of the disk with respect to the observer line of sight ($0 = \text{face-on}$, $90 = \text{edge-on}$) and ϕ is the azimuthal angle between a point on the disk and the projection of the line of sight onto the disk. The observed continuum light curve (previously interpolated by SPEAR) was convolved with the time-delay function for the respective disk and sphere geometry to simulate the expected $H\alpha$ light curve. The results of the simulation are shown in Fig. 9.

The $H\alpha$ light curve for a spherical BLR configuration does not reproduce the observations, however a nearly face-on disk-like BLR geometry with an inclination between 4 to 10° provides an acceptable fit to the original data. To estimate the disk inclination, we performed the χ^2 minimizing fitting procedure for a range of inclination angles (Fig. 9). The best fit yields a value of $i = 6^\circ \pm 3^\circ$ for a disk-like BLR model with an extension from 16 to 20 light days. For such a nearly face-on BLR the velocity of orbiting broad line gas clouds will appear about a factor ten smaller, mimicking a NLS1.

3.4. AGN luminosity and the Host-subtraction process

To determine the AGN luminosity free of host galaxy contributions, we applied the flux variation gradient (FVG) method, originally proposed by Choloniewski (1981) and later modified by Winkler et al. (1992). A detailed description of the FVG method on PRM data is presented in Pozo Nuñez et al. (2012), and here we give a brief outline. In this method the fluxes obtained through different filters and same apertures are plotted in a flux-flux diagram. The fluxes follow a linear slope representing the

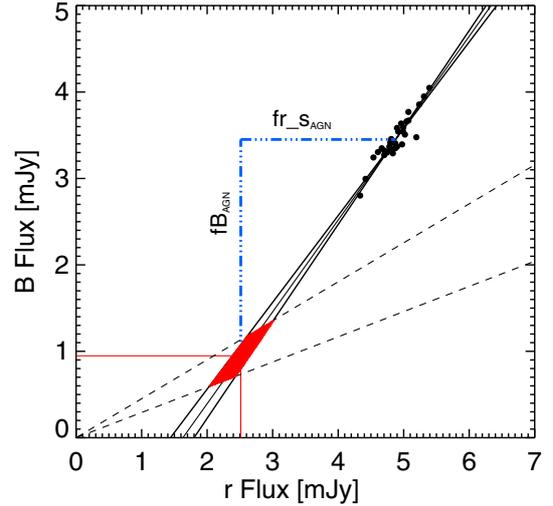


Fig. 10. Flux variation gradient diagram of ESO 399-IG20 for $7''.5$ aperture. The solid lines delineate the ordinary least square bisector regression model yielding the range of the AGN slope. The dashed lines indicate the interpolated range of host slopes obtained from Sakata et al. (2010) for 11 nearby AGN. The intersection between the host galaxy and AGN slope (red area) gives the host galaxy flux in both bands. The dash-dotted blue lines represent the range of the AGN flux in both filters.

AGN color, while the slope of the nuclear host galaxy contribution (including the contribution from the narrow line region (NLR)) lies in a well defined range ($0.4 < \Gamma_{\text{BV}}^{\text{host}} < 0.53$, for $8''.3$ aperture and redshift $z < 0.03$, Sakata et al. 2010). The AGN slope is determined through linear regression analysis. The intersection of the AGN slope with the host galaxy range yields the actual host galaxy contribution at the time of the monitoring campaign. Figure 10 shows the FVG diagram for the B and r fluxes (corrected for galactic foreground extinction) obtained during the same nights and through $7''.5$ aperture. Additionally, FVGs were evaluated for different apertures, as shown in Fig. 11. As already noted in Sect. 2, a big fraction of the total flux is lost in the small $5''$ aperture, hence, the results are more sensitive to a possible underestimation of the real AGN and host galaxy contribution. Furthermore, one may expect that for larger apertures the fluxes will lie closer to the range for the host galaxy slope, however, it appears that the host galaxy is intrinsically strong and very blue, closer to the nucleus and in the outer parts. As for the analysis of the light curves we here use the results for the $7''.5$ aperture.

Averaging over the intersection area between the AGN and the host galaxy slopes, we obtain a mean host galaxy flux of (0.95 ± 0.18) mJy in B and (2.51 ± 0.20) mJy in r . During our monitoring campaign the host galaxy subtracted AGN fluxes range between 2.29 and 2.71 and between 2.14 and 2.58 mJy in the B and r band, respectively. These fluxes are represented by the blue dotted lines in Fig. 10. From this range we interpolate the host-subtracted monochromatic AGN flux at rest-frame 5100 \AA $F_{5100} = 2.66 \pm 0.22 \times 10^{-15} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ \AA}^{-1}$. For the interpolation we assumed a power law spectral energy distribution (SED) ($F_\nu \propto \nu^\alpha$) with an spectral index $\alpha = \log(fB_{\text{AGN}}/fr_{\text{AGN}}) / \log(\nu_B/\nu_r)$, where ν_B and ν_r are the effective frequencies in the B and r bands, respectively. The error was determined by interpolation between the ranges of the AGN fluxes $\pm\sigma$ in both filters. At the distance of 102 Mpc this yields a host-subtracted AGN luminosity at 5100 \AA $L_{\text{AGN}} = (1.69 \pm 0.25) \times 10^{43} \text{ erg s}^{-1}$.

Table 5. Total, host galaxy and AGN fluxes of ESO 399-IG20.

fB_{total} (mJy)	fB_{host} (mJy)	fB_{AGN}^1 (mJy)	fr_{total} (mJy)	fr_{host} (mJy)	fr_{AGN}^1 (mJy)	$f_{\text{AGN}}((1+z)5100 \text{ \AA})$ ($10^{-15} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ \AA}^{-1}$)	$\lambda L_{\lambda, \text{AGN} 5100 \text{ \AA}}$ ($10^{43} \text{ erg s}^{-1}$)
3.45 ± 0.11	0.95 ± 0.18	2.50 ± 0.21	4.87 ± 0.10	2.51 ± 0.20	2.36 ± 0.22	2.66 ± 0.21	1.69 ± 0.25

Notes. ⁽¹⁾ AGN fluxes values $fB_{\text{AGN}} = fB_{\text{total}} - fB_{\text{host}}$ and $fr_{\text{AGN}} = fr_{\text{total}} - fr_{\text{host}}$, with uncertainty range $\sigma_{\text{AGN}} = (\sigma_{\text{total}}^2 + \sigma_{\text{host}}^2)^{0.5}$.

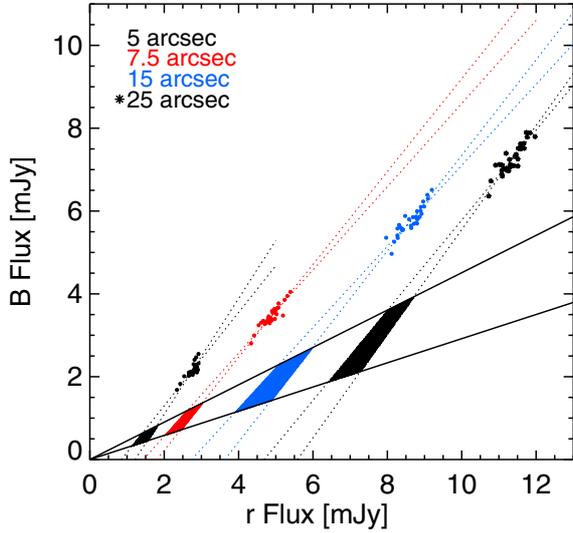


Fig. 11. Flux variation gradient diagram of ESO 399-IG20 for different apertures. The fluxes obtained through 7.5 show a strong correlation, with a correlation coefficient of $r_c = 0.94$. While for 5'', 15'' and 25'' the correlation coefficients are equal to 0.81, 0.90 and 0.89 respectively. This significant decrease in the correlation supports once again our choice of 7.5 as the best aperture.

The total fluxes, host galaxy subtracted AGN fluxes and the AGN luminosity are listed in Table 5.

3.5. The BLR size-luminosity relationship

Estimates of the BLR size and host-galaxy subtracted AGN luminosity in the literature have been derived from several spectroscopic RM campaigns and through host-galaxy modeling using high-resolution images from *HST*. In consequence, the relationship between the $H\beta$ BLR size and the luminosity (5100 \AA) $R_{\text{BLR}} \propto L^\alpha$ (Kaspi et al. 2000) has been improved considerably with the most recent slope of $\alpha = 0.519_{-0.066}^{+0.063}$ (Bentz et al. 2009). Although to date this relationship has been corroborated for 38 AGNs, still there exist objects with large uncertainties in both measurements. In order to improve the statistic, it is of interest to see the position for this new Seyfert 1 galaxy on the BLR-Luminosity relationship. In order to obtain the $H\beta$ BLR radius, we used the weighted mean ratio for the time lag $\tau(H\alpha):\tau(H\beta):1.54:1.00$, obtained recently by Bentz et al. (2010) from the Lick AGN Monitoring Program of 11 low-luminosity AGN. Therefore, the $H\alpha$ lag of 18.2 days translates into an $H\beta$ lag of 11.8 days. Figure 12 shows the position of ESO 399-IG20 on the $R_{\text{BLR}}-L_{\text{AGN}}$ diagram. The data are taken from Bentz et al. (2009) and here we include the most recent results for particular objects obtained from spectroscopic RM by Denney et al. (2010), Doroshenko et al. (2012) and Grier et al. (2012) respectively.

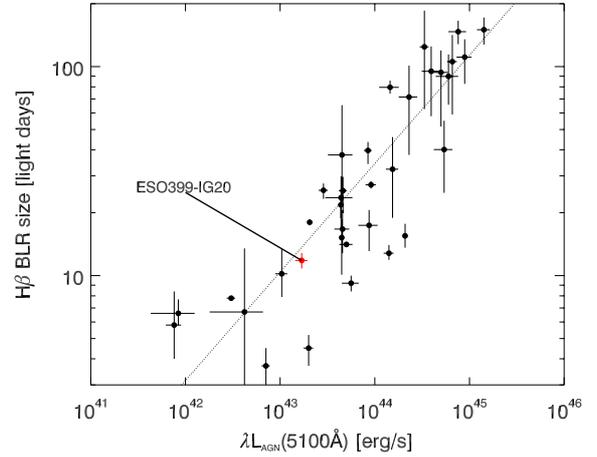


Fig. 12. R_{BLR} versus L , using data from Bentz et al. (2009), Denney et al. (2010), Doroshenko et al. (2012) and Grier et al. (2012). As a red dot we included our results for ESO 399-IG20. Shown is a zoomed portion containing ESO 399-IG20. The solid dotted line show a fitted slope $\alpha = 0.519$. ESO 399-IG20 lies close to the expected slope.

4. Summary and conclusions

We presented new photometric reverberation mapping results for the Seyfert 1 galaxy ESO 399-IG20. We determined the broad line region size, the basic geometry of the BLR and the host-subtracted AGN luminosity. The results are:

1. The cross-correlation of the $H\alpha$ emission line measured in a narrow-band filter with the optical continuum light curve yields a rest-frame time delay $\tau_{\text{rest}} = 18.2 \pm 2.29$ days. We explored the SPEAR method, and – given that $H\alpha$ contributes about 15–20% to the r -band – also the capabilities of pure broad-band photometric reverberation mapping. The SPEAR method yields a rest-frame time delay of $\tau_{\text{rest}} = 17.9 \pm 1.1$ days, while the pure broad-band PRM method yields $\tau_{\text{rest}} = 17.5 \pm 3.1$ days. The results indicate that, within the errors, the three methods are in good agreement.
2. We constrained the basic geometry of the BLR by comparing simulated light curves, using the size determined for the BLR, to the observed $H\alpha$ light curve. The pronounced and sharp variability features in both the continuum and emission line light curves allow us to exclude a spherical BLR geometry. We found that the BLR has a disk-like shape with an inclination of $i = 6 \pm 3^\circ$ and an extension from 16 to 20 light days. This nearly face-on BLR can explain the appearance of ESO 399-IG20 as a narrow-line Seyfert-1 galaxy.
3. We successfully separated the host-galaxy contribution from the total flux through the flux variation gradient method (FVG). The average host-galaxy subtracted AGN luminosity of ESO 399-IG20 at the time of our monitoring campaign is $L_{\text{AGN}} = (1.69 \pm 0.25) \times 10^{43} \text{ erg s}^{-1}$. In the BLR size – AGN luminosity diagram ESO 399-IG20 lies close to the best fit of the relation.

These results document the efficiency and accuracy of photometric reverberation mapping for determining the AGN luminosity, the BLR size and the potential to constrain even the BLR geometry.

Acknowledgements. This publication is supported as a project of the Nordrhein-Westfälische Akademie der Wissenschaften und der Künste in the framework of the academy program by the Federal Republic of Germany and the state Nordrhein-Westfalen. The observations on Cerro Armazones benefitted from the care of the guardians Hector Labra, Gerardo Pino, Roberto Munoz, and Francisco Arraya. This research has made use of the NASA/IPAC Extragalactic Database (NED) which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration. This research has made use of the SIMBAD database, operated at CDS, Strasbourg, France. We thank the anonymous referee for his comments and careful review of the manuscript.

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