

Fast transition of type-B quasi-periodic oscillation in the black hole transient XTE J1817-330

K. Sriram¹, A. R. Rao², and C. S. Choi¹

¹ Korea Astronomy and Space Science Institute, 305-348 Daejeon, Republic of Korea
e-mail: astrosriram@yahoo.co.in

² Tata Institute of Fundamental Research, 400005 Mumbai, India

Received 10 January 2012 / Accepted 6 March 2012

ABSTRACT

Context. The evolution of different types of quasi-periodic oscillations (QPOs) and the coupled radiative/physical changes in the accretion disk are still poorly understood. In a few black hole binaries it was found that fast evolution of QPOs is associated with spectral variations. These studies in other black hole binaries are important to understand the QPO phenomenon.

Aims. We study fast QPO transitions and accompanying spectral variations of the black hole transient XTE J1817-330 to investigate the causes of the spectral variation during the QPO transition.

Methods. Recently, QPOs in ten RXTE observations of XTE J1817-330 were found. We found that among these, only one observation shows erratic dips in its X-ray light curve. The power density spectra and the corresponding energy spectra were extracted and analyzed for the dip and non-dip sections of the light curve.

Results. We found that type-B ~ 6 Hz QPO changes into type-A QPO in a few tens of seconds along with a flux decrease. This transient evolution is accompanied by a significant spectral variation.

Conclusions. We report a transient QPO feature and accompanying spectral variation in XTE J1817-330. Based on our findings, we discuss the origin of fast evolution of QPOs and spectral variations.

Key words. accretion, accretion disks – black hole physics – stars: oscillations – X-rays: stars

1. Introduction

The *RXTE* era of observations of black hole binaries (BHBs) has produced a wealth of information on the radiative and dynamical astrophysical processes associated with the accretion disk in these objects. The black hole binaries often experience an outburst during which the source undergoes in three different states, viz. low-hard (LH) or hard state, steep power-law (SPL) state/very high (VH) state or intermediate (IM) state, and thermal dominated (TD)/high-soft state (McClintock & Remillard 2004; Belloni 2010). The LH state is often observed during the rise and decay phase of the outburst and is characterized by a strong power-law component ($\Gamma \sim 1.5$), strong band-limited noise, high radio flux and evidence to suggest that the inner disk front is truncated at a large radius (Done et al. 2007). The TD state is generally observed at the peak of the outburst associated with a strong disk component, weak band-limited noise, low or no radio flux and the accretion disk is believed to reach the last stable orbit. The SPL state is observed at the peak of the outburst in a few sources, e.g., XTE J1550-564, characterized by a steep power-law component, band limited noise with low- and high-frequency quasi-periodic oscillations (QPOs). The jet is assumed to switch on/off in this state (Fender et al. 2004) and the disk is assumed to be truncated at a small radius (Done et al. 2007). This state spans a comparatively wider area in the hardness intensity diagram (HID) plane (Remillard & McClintock 2006) and is furthermore classified to have hard and soft-intermediate states (HIMS and SIMS) (Belloni 2010; Belloni et al. 2011).

The generating site and mechanism for the low-frequency quasi-periodic oscillations (LFQPOs) in black hole binaries/transients are still not properly understood. The possible

regions in the accretion disk responsible for the QPO modulation are thought to be the inner accretion disk or the Compton cloud in the hot inner region (Done et al. 2007). Casella et al. (2005) summarized the various types of LFQPOs seen in black hole binaries and provided the basis for the classification in three broad categories, viz. type A, B, and C, primarily depending on the rms amplitude and the quality factor ($\nu/\delta\nu$). The properties and occurrence of the various types of QPOs are most often studied using the hardness intensity diagram (HID).

The type-C QPOs have variable centroid peak positions in the range of 0.1–15 Hz along with a harmonic. They are narrow ($\nu/\delta\nu = 7-12$) and have a considerable variability strength with a maximum fractional rms variability of $\sim 30\%$. Spectrally they are often associated with the LH state or radio-loud HIMS state. Recently, Miller-Jones et al. (2012) observed in H1743-322 that the disappearance of type-C QPO is associated with a jet ejection event. Type-B QPOs are observed at ~ 6 Hz, are relatively weaker (rms $\sim 4\%$) and moderately narrow ($\nu/\delta\nu \sim 6$). Sometimes a weak harmonic is also observed, and it is interesting that they are sometimes accompanied by a rapid transition on a timescale of a few seconds (Nespoli et al. 2003). Type-B QPOs are observed as the source traverses from the HIMS to the SIMS. A few studies also show that the type-B QPO are connected to the emission and collimation of radio jets (Casella et al. 2004; Fender et al. 2004). The type-A QPOs are weaker and broader than type-B QPOs with a centroid frequency at ~ 8 Hz and have similar spectral properties as the type-B QPO spectral properties (Belloni et al. 2011; Belloni 2010; Homan et al. 2005).

Transitions among different types of QPOs have been observed in different black hole sources. Nespoli et al. (2003)

found a transient QPO at 6 Hz in GX 339-4 and related it to a spectral hardening during the HIMS to SIMS transition. Casella et al. (2004) observed a fast transition in the properties of power density spectra (PDS) and the corresponding energy spectra in the black hole transient XTE J1859+226. Soleri et al. (2008) found type-B QPOs in GRS 1915+105 along with the transitions from type-B to A and B to type-C as the source moves from the HIMS to SIMS and vice versa. Recently, Stiele et al. (2011) reported that during the soft-to-hard-state transition in SIMS (GX 339-4), the low frequency of type-B QPO is characterized by a low power-law index when compared to the hard-to-soft-state transition. These studies suggest that the associated physical changes in accretion disk are complex and could be related to the onset of radio jets, which are often associated with the HIMS/SIMS transition (the so-called “jet-line”) (Fender et al. 2004; Migliari & Fender 2006; Fender et al. 2009).

The source XTE J1817-330, an X-ray transient, was discovered on 2006 January 26 with the All-Sky Monitor onboard the *RXTE* satellite. The observations in radio (Rupen et al. 2006), optical (Torres et al. 2006), and near infrared/ultraviolet (D’Avanzo et al. 2006; Sala et al. 2007) wavelengths suggest that it is a strong candidate for a black hole binary (BHB) system. The spectral studies show that during the peak of the outburst, the source was in the high-soft (HS) state (see Remillard & McClintock 2006, for spectral classification) and later went to the low-hard state as the source intensity gradually decreased (Gierlinski et al. 2008; Roy et al. 2011). Only in ten observations QPOs were observed with a limited frequency domain of 4–6 Hz (Roy et al. 2011) and in one occasion a QPO was observed at 8–9 Hz (Homan et al. 2006; Roy et al. 2011). It was reported that all QPOs were found in the HS state, though in the HID plane the corresponding observations could be associated to the HIMS/SIMS region (Roy et al. 2011). The observed HS state in XTE J1817-330 has a unique signature, a high disk temperature (~ 0.8 – 0.9 keV) along with a relatively hard spectral component ($\Gamma \sim 2.1$ – 2.3), which is not often seen in other BHBs.

In the present study, we report a transient QPO feature in XTE J1817-330, which shows a sudden transition from type-B to type-A QPO. The spectral change during the observed QPO transition is discussed.

2. Data reduction and analysis

We used the archival data obtained by the *RXTE* satellite (Bradt et al. 1993). There are 160 observations for XTE J1817-330 of which ten observations show QPOs in their PDS (Roy et al. 2011). We used the 16 s bin data from the Proportional Counter Array (PCA, Jahoda et al. 2006) to study intensity variations in the ten observations. We found that only one observation, ObsId 91110-02-30-00 (MJD 53 790), shows erratic dips in the light curve. In Fig. 1 we show the *RXTE* PCA count rates during the outburst. The observation showing the dips lies on the decay phase of the outburst when a flare-like feature was observed in the higher energy bands (Fig. 1) and was not seen in the lower energy bands (Roy et al. 2011). The PDS was extracted using the single-bit mode data SB_125us_8-13_1s (3.68–5.71 keV) and SB_125us_14-35_1s (6.12–14.76 keV) with a time resolution of 1/1024 s. We investigated the PDS behavior separately for the dip and non-dip section of the light curve and the details of the observations are given in Table 1. We extracted the energy spectra for the respective sections from PCU2 unit (3–25 keV) and added a systematic error of 0.5%. We also extracted HEXTE cluster B (High-Energy X-ray Timing Experiment, Rothschild et al. 1995) spectra (15–100 keV) excluding the detector 2 for

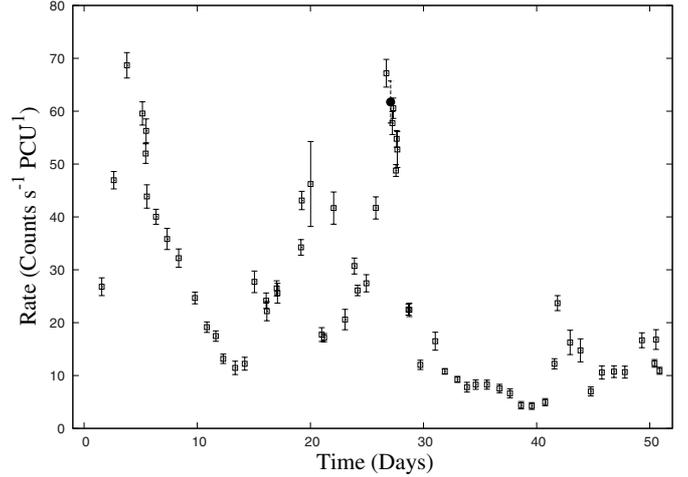


Fig. 1. 9–20 keV PCA light curve of the outburst for initial 50 days (53 764–53 814) is shown. The filled circle marks the observation (MJD 53 790) during which the transient QPO is observed.

dip and non-dip sections. The source and background data were obtained using the `hxtback` command of `FTOOLS` and extracted the corresponding spectra. The dead-time correction was applied to the obtained spectra using `hxtdead` command (see, e.g., Sriram et al. 2007). We used `HEASOFT` v6.8 for the data filtering and used spectral models available in `XSPEC` v12.5.0 (Arnaud 1996).

2.1. Temporal analysis

The X-ray light curve for ObsId 91110-02-30-00 is shown in Fig. 2 (top panel) for the 2–20 keV energy band. A series of dips, lasting for 10–100 s, can be seen in the light curve. For detailed investigations we segregated the data into non-dip and dip sections. Samples are indicated as rectangular boxes in Fig. 2. The PDSs for the non-dip and dip sections in the high- (6.12–14.76 keV) and low-energy band (3.68–5.71 keV) are shown in the middle and bottom panels of Fig. 2 along with the best-fit models consisting of a power-law and a Lorentzian. The PDS in higher energy band of the non-dip section clearly shows a type-B QPO that has a centroid frequency of 5.57 ± 0.06 Hz, a quality factor ($Q = \nu/FWHM$) of 5.74, and an rms of $12 \pm 1\%$. The PDS of the dip section shows a type-A QPO with a centroid frequency of 6.00 ± 0.73 Hz, a quality factor $Q \leq 1$, and an rms of $14 \pm 3\%$. The important signature of type-B QPO is its harmonics, which are seen in the PDS of the non-dip section and they are absent in the type-A QPO (dip section). We have examined the energy dependency of the QPOs and found that the type-B QPO seen in the non-dip section has the same frequency (5.54 ± 0.07 Hz) and quality factor (5.32) in the lower energy band (3.68–5.71 keV), but the rms value is lower ($6.1 \pm 0.5\%$). Interestingly, the dip section PDS shows a QPO at 2.67 ± 0.28 Hz with a $Q \leq 1$ and rms of $7 \pm 2\%$ in the lower energy band. Clearly, the centroid frequency of QPO has significantly changed in the energy domain during the dip section. We examined the dynamical PDS of the light curve and found that in the non-dip section the centroid of the QPO often varied from 5.0 to 6.5 Hz.

2.2. Spectral analysis

To see whether there is a significant spectral difference between the non-dip and dip sections, we first checked the residuals by

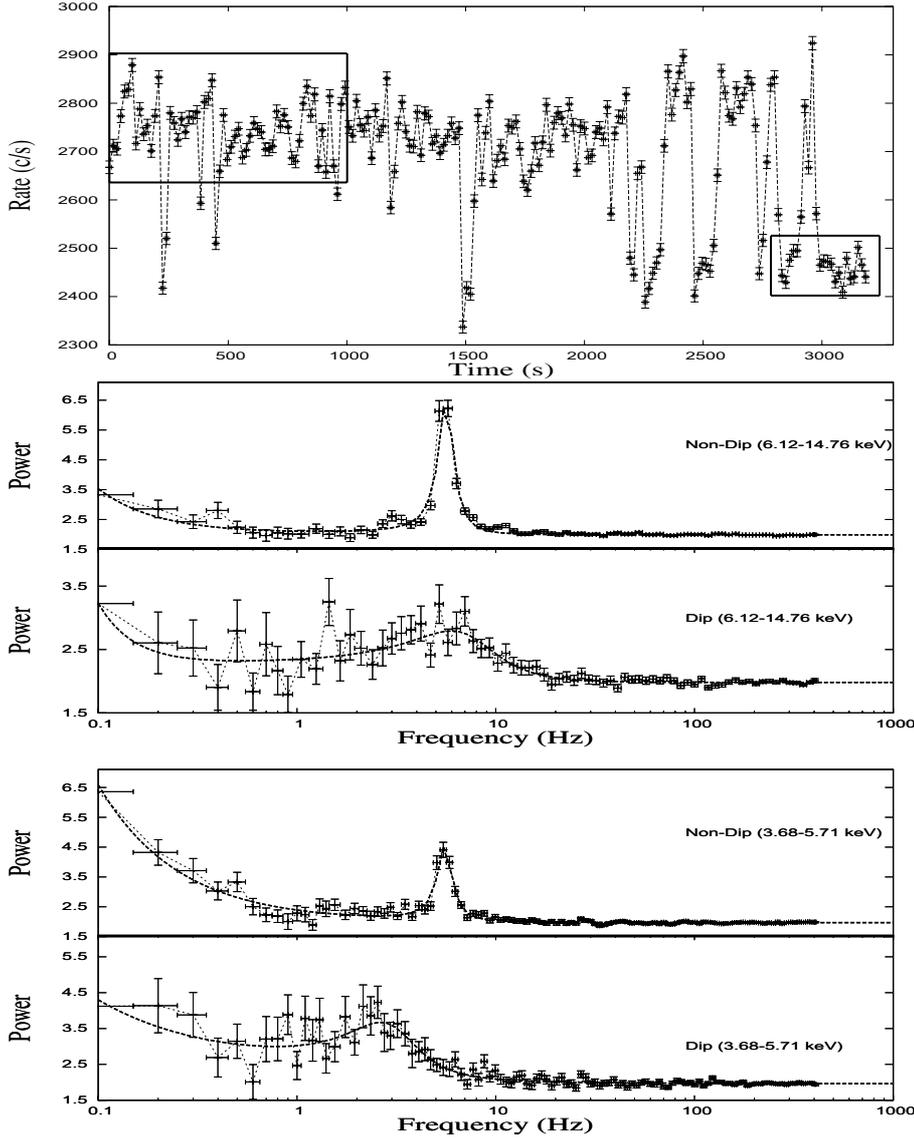


Fig. 2. *Top:* PCA 16 s bin light curve in the 2–20 keV band. The boxes show the sections of non-dip and dip used for PDS and spectral analysis. *Middle:* the PDS of non-dip and dip sections in the energy band 6.12–14.75 keV along with their best-fits using a *power-law+Lorentzian+constant* model (shown in thick dashed lines). The wings or harmonics in the non-dip PDS are clearly visible but absent in the dip PDS, which is a characteristic feature observed when a type-B QPO moves into a type-A QPO. *Bottom:* similar PDSs in the lower energy band (3.68–5.71 keV) are shown. The Poisson noise is not subtracted.

fixing the non-dip spectral parameters to the dip spectrum. The residuals are shown in the top panel of Fig. 3 and the high $\Delta\chi$ values of the dip spectrum suggest that the spectrum has changed significantly during the type-B to A transition. We then fitted the non-dip and dip spectra (3–25 keV) using the model *wabs(diskbb+Gaussian+power-law)* where the hydrogen equivalent column density was fixed at N_{H} (wabs model) = $1.2 \times 10^{21} \text{ cm}^{-2}$ (Rykoff et al. 2007) and the Gaussian line centroid energy as 6.4 keV (for iron line). The best-fit spectral parameters are shown in Table 2 and the unfolded spectra are shown in the bottom panels of Fig. 3. We found that the power-law index has decreased from $\Gamma = 2.46 \pm 0.03$ (non-dip) to $\Gamma = 2.29 \pm 0.03$ (dip), indicating that a spectral hardening occurs during the type-B to A transition. We also fitted the PCA+HEXTE spectrum (3–100 keV) using the same model and observed a similar hardening of power-law index. Both PCA and PCA+HEXTE spectral fit results show that the power-law flux decreased from non-dip to dip section and the flux is decreased by $\sim 17\%$ in the energy band 3–100 keV, whereas no appreciable change is observed in the diskbb model component. The total flux obtained by Roy et al. (2011) (i.e. $1.02 \times 10^{-8} \text{ erg cm}^{-2} \text{ s}^{-1}$ in 3–25 keV) and in the present work ($1.04 \times 10^{-8} \text{ erg cm}^{-2} \text{ s}^{-1}$ in 3–25 keV) are similar. We note that the results do not vary if we free the hydrogen equivalent column density and Gaussian line energy.

An attempt was made to use the CompTT model instead of the power-law, but it resulted in unacceptable fits to the spectra. The spectral hardening along with the flux decrease in the dip suggest that a relatively cooler portion of hard X-ray emitting region is ejected away, leaving behind a hotter hard X-ray emitting region.

3. Discussion and conclusion

The low-frequency QPOs (LFQPOs) are the most common temporal features observed in BHBs and play a key role in the spectral classification schemes (Remillard & McClintock 2006; Belloni 2010; Belloni et al. 2011). During the outburst, the frequency of LFQPO increases, it disappears at HS state, and again re-appears in the VHS/IM state. However, the exact production mechanism of oscillation and the location in the accretion disk is not yet clear but it is thought to occur in the inner regions of the accretion disk. The classification of the type C, B, and A QPOs (Casella et al. 2004, 2005) shows that there is a unique but complex connection with the soft and hard X-ray components and indicates a sudden state transition from HIMS to SIMS or vice-versa (Belloni et al. 2005; Belloni 2010). In general, type-B and A QPOs and their transitions are often observed in BHBs when they are in very high (VH)/steep power-law (SPL) state or intermediate state (IM) (Belloni 2010). In this state the accretion disk

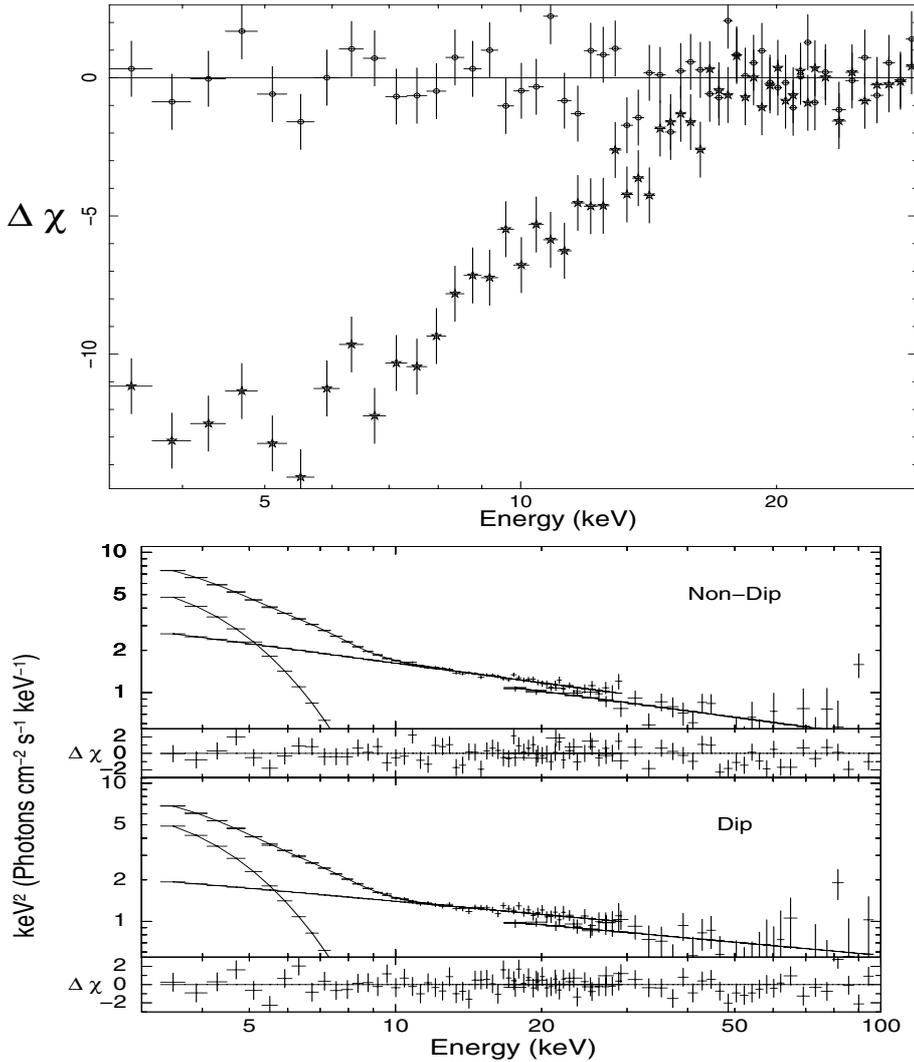


Fig. 3. *Top:* residuals of the non-dip and dip spectra obtained by applying the best-fit spectral parameters of the non-dip spectrum. The high residuals of the dip spectrum (star symbols) show that the spectrum has appreciably changed, which is responsible for the observed type-A QPO in this section. *Bottom:* the spectral fit to the broadband spectra (3–100 keV) of the non-dip and dip sections.

is considered to contain a compact corona (low temperature with high optical depth) along with a high-temperature Keplerian disk located close to the last stable orbit (Done & Kubota 2006; Sriram et al. 2007, 2009, 2010).

We found one such observation in XTE J1817-330, where a type-B QPO of ~ 6 Hz is transitioning to a type-A QPO in a few tens of seconds. During this observation, the source was in the declining phase of the outburst, showing a sudden flaring in the higher energy bands (Fig. 1). It is evident from the PDS that the characteristic harmonics have also disappeared during the transition (middle panel of Fig. 2) and the energy-dependent PDS study of the dip section shows that the centroid frequency of QPO is shifted to lower frequency from higher to lower energy band, whereas no such variation was seen in the non-dip section. This could be caused by some ejection of low-temperature material from the accretion disk. The spectral results during the type-B and type-A QPO show that the power-law component became harder along with a decrease in its flux (Table 2). This observation falls in the HS classification but occupies SIMS location in the HID plane (Roy et al. 2011). This source has the unique property that in its HS state, the hard component is relatively stronger with the occasional presence of QPOs, which is generally not observed in other BH sources. However, in the transient QPO observation, we suggest the dips may occur because of a sudden transition from the soft to SIMS state, which is inferred from the variation in the power-law index.

The transition of QPOs among different types, viz. type C-B and type B-A and viceversa were observed in a few other black hole sources. Homan et al. (2001) found broad QPO features around 6 Hz in XTE J1550-564 when the source was in a low-intensity state and it became narrow in a high-intensity state. For XTE J1859+226, a broad 6 Hz feature was observed when the source was in a high-intensity state and became narrow in a low-intensity state (Casella et al. 2004). In GX 339-4, Nespoli et al. (2003) observed a transition of a type-A 6 Hz QPO (low intensity state) to a type-B along with a harmonic associated with the hardening of the spectrum. They found that during the transition, the disk flux decreased by 9%, power-law component flux increased by 30% and total flux increased by 9%. Belloni et al. (2005) found that this transient feature is associated with the spectral transition from HIMS to SIMS. Similar transient variation (type-B to A QPO) was also observed in H1743-322 (Homan et al. 2005, see their Fig. 4 panel C and D), but no appreciable changes were observed in disk and power-law components. However, during this transition *Chandra* observations revealed an absence of Fe lines possibly caused by the high ionization or by absorbers (Miller et al. 2006). GRS 1915+105 also shows such transient type A-B QPO features and their properties are dependent on the location in HID and were discussed in the context of the onset of a variable jet (Soleri et al. 2008). Recently, Stiele et al. (2011) found that for GX 339-4, the power-law index increases as the source moves from the soft to hard

Table 1. The details of the observation analyzed for the source XTE J1817-330.

ObsId	MJD (Date)	Duration (s)		
		Total observation	Non-dip	Dip
91110-02-30-00	53 790 (2006-02-24)	4132	990	350

Table 2. Best-fit spectral parameters for the non-dip and dip spectra using the model wabs (diskbb+Gaussian+power-law) (see text).

Parameters	ObsId 91110-02-30-00	
	Non-dip (type-B QPO) For PCA	Dip (type-A QPO)
kT_{in} (keV) ^a	0.90 ± 0.01	0.91 ± 0.01
N_{diskbb} ^b	1289 ± 115	1411 ± 150
Γ ^c	2.46 ± 0.03	2.29 ± 0.03
Γ_{Norm} ^d	4.59 ± 0.37	2.74 ± 0.45
diskbb flux	4.37	4.41
power-law flux	6.10	5.00
Total flux	10.43	9.41
$\chi^2/\text{d.o.f.}$	48/48	37/48
For PCA+HEXTE		
kT_{in} (keV)	0.92 ± 0.008	0.91 ± 0.01
N_{diskbb}	1271 ± 66	1413 ± 155
Γ	2.46 ± 0.02	2.31 ± 0.04
Γ_{Norm}	4.71 ± 0.37	2.89 ± 0.42
diskbb flux	4.40	4.45
power-law flux	7.20	6.14
$\chi^2/\text{d.o.f.}$	104/79	71/79

Notes. The quoted errors are at a 90% confidence level and fluxes are in units of 10^{-9} erg cm^{-2} s^{-1} for PCA (3–25 keV) and HEXTE (15–100 keV). The fluxes for PCA+HEXTE spectral fits are obtained in the energy band 3–100 keV. ^(a) Inner disk temperature. ^(b) Normalization of diskbb model. ^(c) Power-law index. ^(d) Normalization of the power-law model (photons keV^{-1} cm^{-2} s^{-1} at 1 keV).

state in SIMS and weakens as the source moves from the hard to soft state.

The spectral result shows the spectral hardening of the power-law component in the dip (Table 2), suggesting that the source was traversing from the soft to hard SIMS. The overall power-law component variation can be explained in the following scenario. Assume a Compton cloud with a radial distribution of electron temperature i.e., hotter at the center and cooler at the outer region of the Compton cloud. From this configuration, as the source moves to the low-intensity state (i.e., dips) the outer region of Compton cloud is ejected away in the form of a jet or an outflow from the disk and leaves behind a hotter part of the Compton cloud. This physical picture explains both the spectral hardening of the power-law component and the flux decrease in the dip. Moreover, this scenario is also favored by the shifting of the centroid frequency of the QPO (low energy band) in the dip section, which suggests that the low-temperature component is relatively more affected than the high-temperature component.

Type-B QPOs are relatively rare and tend to occur during the local peaks of the outbursts (see Motta et al. 2011). These fast transitions from type-B to type-A are still more rare. In GX 339-4 such “flip-flop” transitions were seen in the *Ginga* data (Miyamoto et al. 1991) and later in the *RXTE* data (Motta et al. 2011). Takizawa et al. (1997) reported “flip-flop” transitions in GS 1124-683 and Casella et al. (2004) reported it in the *RXTE* data of XTE J1859+226. The fast transitions seen in GRS 1915+105 in the μ and β classes are also of similar nature (Soleri et al. 2008). As discussed by Miyamoto et al. (1991),

a jet origin for the origin of some part of the hard component (and possibly for the type-B QPOs) is a feasible scenario.

In summary, we found a transient 6 Hz QPO in the source XTE J1817-330 during the decay phase of the outburst. During this evolution the type-B QPO characteristics in non-dip of the light curve were changing into type-A QPO in dip. Our spectral study provides one of the strongest piece of evidence for spectral variations during the fast transition of type-B to type-A QPO. This reported result for the source XTE J1817-330 again give credence to the physical scenario that the type-B to A QPO transition occurs on a short timescale and could be connected to a variable onset of jet or an outflow in the accretion disk.

Acknowledgements. We thank the anonymous referee for the very useful comments. This research has made use of data obtained through the HEASARC Online Service, provided by NASA/GSFC in support of the NASA High Energy Astrophysics Programs. K.S. thanks the hospitality provided at TIFR, where some part of the work was done.

References

- Arnaud, K. A. 1996, ASP Conf. Ser., 105, 17
D’Avanzo, P., Goldoni, P., Covino, S., et al. 2006, *Astron. Tel.*, 724
Belloni, T. M. 2010, in *The jet paradigm: from microquasars to quasars*, Lecture Notes in Physics 794, ed. T. Belloni (Berlin: Springer Verlag), 53
Belloni, T., Homan, J., Casella, P., et al. 2005, *A&A*, 440, 207
Belloni, T. M., Motta, S. E., & Munoz-Darias, T. 2011, *BASI*, 39, 409
Bradt, H. V., Rothschild, R. E., & Swank, J. H. 1993, *A&AS*, 97, 355
Casella, P., Belloni, T., Homan, J., & Stella, L. 2004, *A&A*, 426, 587
Casella, P., Belloni, T., & Stella, L. 2005, *ApJ*, 629, 403
Done, C., & Kubota, A. 2006, *MNRAS*, 371, 1216
Done, C., Gierlinski, M., & Kubota, A. 2007, *A&ARv*, 15, 1
Fender, R. P., Belloni, T., & Gallo, E. 2004, *MNRAS*, 355, 1105
Fender, R. P., Homan, J., & Belloni, T. M. 2009, *MNRAS*, 396, 1370
Gierlinski, M., Done, C., & Page, K. 2008, *MNRAS*, 388, 753
Homan, J., Wijnands, R., van der Klis, M., et al. 2001, *ApJS*, 132, 377
Homan, J., Miller, J. M., Wijnands, R., et al. 2005, *ApJ*, 623, 383
Homan, J., Miller, J. M., & Wijnands, R. 2006, *Astron. Tel.*, 752
Jahoda, K., Markwardt, C. B., Radeva, Y., et al. 2006, *ApJS*, 163, 401
McClintock, J. E., & Remillard, R. A. 2004, in *Compact Stellar X-ray Sources*, ed. W. H. G. Lewin, & M. van der Klis (Cambridge: Cambridge Univ. Press), 157
Migliari, S., & Fender, R. P. 2006, *MNRAS*, 366, 79
Miller, J. M., Raymond, J., Homan, J., et al. 2006, *ApJ*, 646, 394
Miller-Jones, J. C. A., Sivakoff, G. R., Altamirano, D., et al. 2012, *MNRAS*, 421, 468
Miyamoto, S., Kimura, K., Kitamoto, S., Dotani, T., & Ebisawa, K. 1991, *ApJ*, 383, 784
Motta, S., Munoz-Darias, T., Casella, P., Belloni, T., & Homan, J. 2011, *MNRAS*, 418, 2292
Nespoli, E., Belloni, T., Homan, J., et al. 2003, *A&A*, 412, 235
Remillard, R. A., & McClintock, J. E. 2006, *ARA&A*, 44, 49
Rothschild, R. E., Matteson, J. L., Heindl, W., Peterson, L. E., & Hink, P. L., et al. 1995, *Proc. SPIE*, 2518, 13
Roy, J., Agrawal, P. C., Paul, B., & Duorah, K. 2011, *MNRAS*, 412, 1011
Rupen, M. P., Dhawan, V., & Mioduszewski, A. J. 2006, *Astron. Tel.*, 717
Rykoff, E. S., Miller, J. M., Steeghs, D., & Torres, M. A. P. 2007, *ApJ*, 666, 1129
Sala, G., Greiner, J., Ajello, M., Bottacini, E., & Haberl, F. 2007, *A&A*, 473, 561
Sriram, K., Agrawal, V. K., Pendharkar, J. K., & Rao, A. R. 2007, *ApJ*, 661, 1055
Sriram, K., Agrawal, V. K., & Rao, A. R. 2009, *RAA*, 9, 901
Sriram, K., Rao, A. R., & Choi, C. S. 2010, *ApJ*, 725, 1317
Soleri, P., Belloni, T., & Casella, P. 2008, *MNRAS*, 383, 1089
Takizawa, M., Dotani, T., Mitsuda, K., et al. 1997, *ApJ*, 489, 272
Stiele, H., Motta, S., Munoz-Darias, T., & Belloni, T. M. 2011, *MNRAS*, 418, 1746
Torres, M. A. P., Steeghs, D., Jonker, P. G., et al. 2006, *Astron. Tel.*, 733