

# Kinetics of $\text{OCN}^-$ formation from the $\text{HNCO} + \text{NH}_3$ solid-state thermal reaction

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## ABSTRACT

**Context.** Solid-state features in infrared astronomical spectra can provide useful information on interstellar ices within different astrophysical environments. Solid  $\text{OCN}^-$  has an absorption feature at  $4.62 \mu\text{m}$ , which is observed in star formation regions only with a large source-to-source abundance variation.

**Aims.** We aim to investigate the thermal formation mechanism of solid  $\text{OCN}^-$  from  $\text{HNCO}$  on the basis of kinetic arguments.

**Methods.** We experimentally studied the kinetics of the low-temperature  $\text{OCN}^-$  formation from the purely thermal reaction between  $\text{HNCO}$  and  $\text{NH}_3$  in interstellar ice analogs using Fourier transform infrared spectroscopy. We used a rate equation approach, a kinetic Monte Carlo approach and a gamma probability distribution approach to derive kinetic parameters from experimental data.

**Results.** The kinetics can be divided into two-processes, a fast process corresponding to the chemical reaction, and a slow process that we interpret as the spatial orientation of the two reactants within the ice. The three approaches give the same results. The  $\text{HNCO} + \text{NH}_3 \rightarrow \text{OCN}^- + \text{NH}_4^+$  reaction rate follows an Arrhenius law with an activation energy of  $0.4 \pm 0.1 \text{ kJ mol}^{-1}$  ( $48 \pm 12 \text{ K}$ ) and a pre-exponential factor of  $0.0035 \pm 0.0015 \text{ s}^{-1}$ .

**Conclusions.** The present experiment has the important implication that the  $\text{HNCO} + \text{NH}_3$  reaction can account for the observed abundances of solid  $\text{OCN}^-$  and the  $\text{HNCO}$  non detection in young stellar objects.

**Key words.** astrochemistry – ISM: molecules – molecular processes

## 1. Introduction

The strong absorption feature at 4.62 microns ( $2165 \text{ cm}^{-1}$ ), called the XCN band, is commonly attributed to the solid-phase cyanate ion  $\text{OCN}^-$  (Soifer et al. 1979; Grim & Greenberg 1987; Lowenthal et al. 2000; Novozamsky et al. 2001). Although absent in starless dark molecular clouds (Knez et al. 2005), this feature appears in the spectra of several protostellar objects with large source-to-source abundance variations (van Broekhuizen et al. 2005). The  $\text{OCN}^-$  band is therefore often seen as a manifestation of energetic processes, thermal or non-thermal, in a protostellar environment.

Solid-state  $\text{OCN}^-$  can originate from different processes: proton bombardment (Moore et al. 1983), electron bombardment (Bennett et al. 2010), vacuum ultraviolet photolysis of  $\text{NH}_3:\text{CO}:\text{H}_2\text{O}$  ices where the photochemical formation of isocyanic acid  $\text{HNCO}$  is followed by a proton transfer to ammonia  $\text{NH}_3$  (d'Hendecourt et al. 1986b; Grim et al. 1989), or directly from an acid-base reaction of  $\text{HNCO}$  with solid  $\text{NH}_3$  or  $\text{H}_2\text{O}$  (Demyk et al. 1998; Raunier 2003b,a; van Broekhuizen et al. 2004). The quantitative comparison of these different processes made in vanBroekhuizen et al. (2004) tends to show that the UV photoprocessing of  $\text{NH}_3:\text{CO}:\text{H}_2\text{O}$  ices may not account for the observed  $\text{OCN}^-$  abundances, and that  $\text{OCN}^-$  is most likely formed thermally from  $\text{HNCO}$ . The thermal formation from the reaction between  $\text{HNCO}$  and either  $\text{NH}_3$  or  $\text{H}_2\text{O}$  could explain the large variations in abundances observed in VLT studies of low-mass young stellar objects (YSOs) (van Broekhuizen et al. 2005), which are difficult to explain by photochemistry. However, the validity of these mechanisms to reproduce or not reproduce, the observed abundances of  $\text{OCN}^-$  depends on their respective efficiency and kinetics and on the initial abundance

of the  $\text{HNCO}$  reactant. The kinetic study of the  $\text{HNCO} + \text{H}_2\text{O}$  solid-state reaction has been investigated (Theule et al. 2011) and an activation barrier of  $26 \text{ kJ mol}^{-1}$  ( $3127 \text{ K}$ ) has been found for this reaction. A barrier like this is too high for this reaction to occur on YSO timescales or to counterbalance faster reactions such as hydrogenation of  $\text{HNCO}$  in molecular clouds.  $\text{HNCO} + \text{NH}_3$  and  $\text{HNCO} + \text{H}_2\text{O}$  are two competitive reactions, but the timescale of the latter one is much longer than the timescale of the former, so we only need to consider the  $\text{HNCO} + \text{NH}_3$  reaction on our experimental timescales.

In this paper we measure the kinetics of the  $\text{HNCO} + \text{NH}_3 \rightarrow \text{OCN}^- + \text{NH}_4^+$  isolated solid-state reaction on the basis of laboratory experiments using Fourier transform infrared (FTIR) spectroscopy between 8 K and 60 K. Using a rate-equation approach, a kinetic Monte Carlo (KMC) approach and a gamma probability distribution approach, we show that the kinetics can be decomposed into two successive steps. One fast step, corresponding to the reaction itself, follows a second-order kinetic law, and its reaction rate follows an Arrhenius law with an activation energy of  $0.4 \pm 0.1 \text{ kJ mol}^{-1}$  ( $48 \pm 12 \text{ K}$ ) and a pre-exponential factor of  $0.0035 \pm 0.0015 \text{ s}^{-1}$ . One slow step, corresponding to the reorientation of the molecules position within the molecular solid, follows a first-order kinetic law. This experimental work provides the reference input data to a gas-grain chemistry model that we will use to discuss the observed solid-state  $\text{OCN}^-$  abundances in different astrophysical objects.

## 2. Experimental

The  $\text{HNCO}$  monomer was synthesized from the thermal decomposition of the cyanuric acid polymer (Aldrich Chemical Co., 98%) at  $650 \text{ }^\circ\text{C}$  under a primary vacuum. Ammonia is

commercially available as a 99.999% pure gas from Air Liquide. HNCO and NH<sub>3</sub> were inserted into two separated vacuum lines using standard manometric techniques, outgassed and then co-deposited. The two gases were sprayed onto a gold plated copper surface maintained at 10 K using a closed-cycle helium cryostat (ARS Cryo, model DE-204 SB) within a high vacuum chamber at few 10<sup>-9</sup> mbar pressure. Infrared spectra were recorded at different temperatures using Fourier transform reflection absorption infrared spectroscopy (FT-RAIRS) using an Vertex 70 spectrometer with a MCT detector. A typical spectrum has a 1 cm<sup>-1</sup> resolution and is averaged over one hundred interferograms.

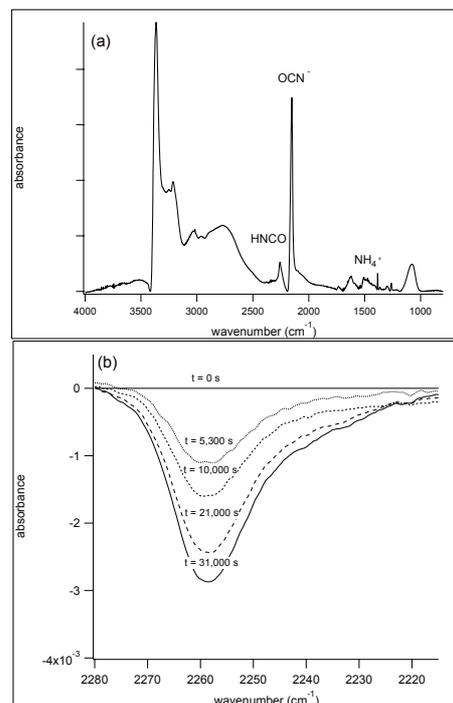
The concentration ratio of the different mixtures is derived after deposition at 10 K from the IR spectra by integrating the HNCO and NH<sub>3</sub> bands to estimate their column densities according to their band strengths. For NH<sub>3</sub> we used the value of 1.3 × 10<sup>-17</sup> cm molecule<sup>-1</sup> for the band strength of the wagging mode at 1070 cm<sup>-1</sup> (Kerkhof et al. 1999) in its pure solid form. For HNCO, we adopted the value of 7.8 × 10<sup>-17</sup> cm molecule<sup>-1</sup> for the band strength of the ν<sub>as</sub>NCO mode at 2259 cm<sup>-1</sup>, which is the value given for a H<sub>2</sub>O-poor ice (van Broekhuizen et al. 2004). For OCN<sup>-</sup> we used the value of 1.3 × 10<sup>-16</sup> cm molecule<sup>-1</sup> for the band strength of the ν<sub>as</sub>NCO mode (van Broekhuizen et al. 2004). However, the calculated concentration ratio depends on the band strengths provided by the literature. These values depend on the nature, composition, and temperature of the ice in which they are found, and this dependence is a major source of uncertainties when evaluating the column densities of frozen molecules. A 20% uncertainty on the band strengths and accordingly on concentration ratios is probably conservative. A typical ice mixture has a few hundreds of monolayers thickness according to integrated molecules IR bands and Table 1 band strengths, assuming a 10<sup>15</sup> molecule cm<sup>-2</sup> surface density. We used these bands to monitor the evolution of the HNCO and OCN<sup>-</sup> abundances as they are consumed or produced during the reaction to estimate the different kinetic parameters. The uncertainty on the band strengths has no influence on the time evolution of the molecule abundances, and accordingly on the derived kinetic parameters. In these experiments, only bulk effects are concerned since our set-up is fairly insensitive to surface molecules.

In a typical kinetic experiment, the temperature is set to a fixed value using a Lakeshore Model 336 temperature controller and measured using a DTGS 670 Silicon diode with a 0.1 K uncertainty. The temperature is set at a fixed value between 8 K and 300 K using a heating resistance and the temperature controller. Infrared spectra are recorded at this fixed temperature at fixed time intervals.

### 3. Results

#### 3.1. Kinetic measurements of the HNCO + NH<sub>3</sub> → OCN<sup>-</sup> + NH<sub>4</sub><sup>+</sup> reaction at a fixed temperature

An HNCO:NH<sub>3</sub> ice mixture where NH<sub>3</sub> is in excess with respect to HNCO is deposited at *T* = 10 K, as shown in Fig. 1a. An excess of NH<sub>3</sub> with respect to HNCO is chosen so that one HNCO molecule is surrounded by NH<sub>3</sub> molecules only. An excess of NH<sub>3</sub> is consistent with observations of interstellar ices against high-mass (Dartois 2005) and low-mass (Bottinelli et al. 2010) stars, which indicate that NH<sub>3</sub> is present within few percents with respect to H<sub>2</sub>O, and that HNCO is not observed (estimated upper limit of 0.1%). We typically used an HNCO:NH<sub>3</sub> ice mixture in a 1:20 ratio to have an excess of NH<sub>3</sub> with respect to HNCO. Different mixing ratios will give the same results because the probability of the reaction is set by the less abundant



**Fig. 1.** **a)** Infrared spectrum of an HNCO:NH<sub>3</sub> ice mixture in a 1:20 ratio. OCN<sup>-</sup> has a band at 2151 cm<sup>-1</sup>. HNCO has a band at 2259 cm<sup>-1</sup> corresponding to HNCO in a NH<sub>3</sub> environment. **b)** Difference spectrum with respect to *t* = 0 s of the HNCO 2259 cm<sup>-1</sup> band at *T* = 15 K at *t* = 5300 s, 10 000 s, 21 000 s, and 31 000 s.

species, HNCO in this case. How many NH<sub>3</sub> molecules are spectators and do not participate to the reaction has no effect on the kinetics.

At 10 K the ammonia spectrum is characterized by its strong band at 1110 cm<sup>-1</sup> assigned to the inversion mode. Isocyanic acid HNCO is characterized by strong bands located at 3373, 3233, 2259 and 858 cm<sup>-1</sup>. The most intense band at 2259 cm<sup>-1</sup> is assigned to the NCO asymmetric stretch vibration. The isocyanate ion OCN<sup>-</sup> is characterized by its intense band at 2151 cm<sup>-1</sup> and the ammonium ion NH<sub>4</sub><sup>+</sup> is characterized by its band at 1495 cm<sup>-1</sup>. Since the HNCO + NH<sub>3</sub> reaction is possible during the deposition step when the gases at 300 K are co-deposited onto the cold metal surface, and thermalized from 300 K to 10 K, a significant amount of OCN<sup>-</sup> is initially formed as seen from its band at 2151 cm<sup>-1</sup> immediately after deposition at 10 K. The amount of initially formed OCN<sup>-</sup> slightly varies from one deposition to another. However, an initial amount of solid HNCO remains after deposition, and will react with solid NH<sub>3</sub> at the metal surface temperature. An extensive vibrational assignment of the OCN<sup>-</sup> and NH<sub>4</sub><sup>+</sup> products is given in Raunier et al. (2003a). All observed bands are listed in Table 1 along with their corresponding assignments and band strengths. There are no unassigned bands, which means that there are no other reactions and that NH<sub>4</sub><sup>+</sup> and OCN<sup>-</sup> cannot react any further.

After deposition the HNCO:NH<sub>3</sub> ice mixture is warmed as fast as possible, in few tens of seconds approximately, to a fixed temperature, between 10 K and 60 K, without overshooting the fixed temperature value. This defines the initial time (*t* = 0 s) of the kinetic experiment. When the temperature is kept to this fixed value, the OCN<sup>-</sup> band at 2151 cm<sup>-1</sup> and the NH<sub>4</sub><sup>+</sup> band at 1495 cm<sup>-1</sup> are progressively growing, while the bands corresponding to HNCO and NH<sub>3</sub> are decreasing along with time.

**Table 1.** Positions and integrated band strengths of an HNCO:NH<sub>3</sub> ice mixture at 10 K.

HNCO:NH <sub>3</sub> ice mixture at 10 K			
Position (cm <sup>-1</sup> )	A (cm molecule <sup>-1</sup> )	Mode	Molecule
3377	–	$\nu$ NH	NH <sub>3</sub>
3218	–	$\nu$ NH	HNCO
3296	–	$\nu$ NH	NH <sub>3</sub>
2259	$7.8 \times 10^{-17a}$	$\nu_{as}$ NCO	HNCO
2151	$1.3 \times 10^{-16a}$	$\nu_{as}$ NCO	OCN <sup>-</sup>
1628	–	$\delta$ NH	NH <sub>3</sub>
1495	$4.6 \times 10^{-17a}$	$\delta$ NH	NH <sub>4</sub> <sup>+</sup>
1313–1264	–	$\nu_s$ NCO	HNCO
1110	$1.3 \times 10^{-17b}$	$\omega$ NH	NH <sub>3</sub>
858	–	$\delta$ HNC	HNCO

Notes. <sup>(a)</sup> van Broekhuizen et al. (2004); <sup>(b)</sup> Kerkhof et al. (1999).

This indicates that a thermally induced reaction between HNCO and NH<sub>3</sub> is occurring to give the NH<sub>4</sub><sup>+</sup>OCN<sup>-</sup> product, as previously studied (Raunier 2003a). We will not use the OCN<sup>-</sup> asymmetric stretch at 2151 cm<sup>-1</sup> to monitor the NH<sub>4</sub><sup>+</sup>OCN<sup>-</sup> product formation since OCN<sup>-</sup> is reorganizing between 10 K and 40 K within a time scale similar to the reaction time scale, as shown by a characteristic crystallization band profile. We will instead use the HNCO band at 2259 cm<sup>-1</sup> to monitor the kinetics of the HNCO transformation into OCN<sup>-</sup>. Figure 1b shows the evolution along with time of a HNCO:NH<sub>3</sub> ice mixture in a 1:20 ratio at a fixed 15 K temperature.

During this reaction a proton is transferred from the HNCO nitrogen atom to the NH<sub>3</sub> nitrogen atom to form the OCN<sup>-</sup>NH<sub>4</sub><sup>+</sup> species as expressed in

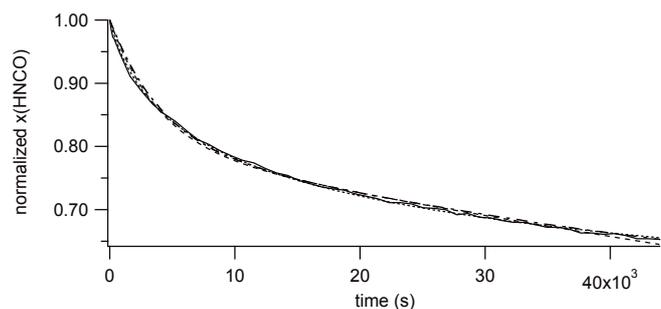


There is no desorption of the different species at temperatures below 60 K, because the HNCO, NH<sub>3</sub> and OCN<sup>-</sup> desorption energies are high (Burke & Brown 2010; Theule et al. 2011) and desorb around 130 K within our experimental conditions. Moreover, no HOCN that may come from the reaction NH<sub>4</sub><sup>+</sup> + OCN<sup>-</sup> → HOCN + NH<sub>3</sub> is observed from its band at 2295 cm<sup>-1</sup> (Raunier 2003a; Theule et al. 2011). NH<sub>3</sub> is a much stronger base than H<sub>2</sub>O, consequently a proton transfer from NH<sub>4</sub><sup>+</sup> is more difficult to achieve.

To derive a kinetic law for the formation of NH<sub>4</sub><sup>+</sup>OCN<sup>-</sup> from the HNCO + NH<sub>3</sub> reaction, we carried out experiments at different fixed temperatures. The temperature must be sufficiently high for the formation to be measured in a realistic time scale of a few hours, which corresponds to a temperature above 8 K. However, for temperatures above 60 K, the reaction is too fast to be able to measure its kinetics, so the measurement range is limited between approximately 8 K and 60 K. Figure 2 shows the decrease of the HNCO normalized column density along with time, as derived from its integrated band at 2259 cm<sup>-1</sup>, at different fixed temperatures. Not 100% of HNCO is consumed during a *c.a.* 12 h time scale. Although the reaction is total and not partial (Theule et al. 2011), only 30% of HNCO is consumed for all experiments on this limited time scale.

### 3.2. Kinetic studies of the HNCO + NH<sub>3</sub> reaction using a rate equation approach

We first used a deterministic approach based on rate equations to express the time evolution of HNCO. The kinetics of solid-state reactions cannot generally be assumed to follow



**Fig. 2.** Time evolution of the normalized HNCO molar fraction recorded by FTIR spectroscopy (full line) at  $T = 20$  K for a HNCO:NH<sub>3</sub> mixture where NH<sub>3</sub> is largely in excess (1:20 concentration ratio). Fitting the experimental data against the analytical expressions for the HNCO time evolution using a rate equation approach (dashed line), a Monte Carlo approach (dotted-dashed line) and a gamma probability distribution (dotted line) gives the value for the reaction rate at a fixed temperature. The three approaches perfectly fit the experimental HNCO decay curve.

simple rate laws, as is the case for gas-phase or liquid-phase reactions. Different solid-state models have been developed to extract Arrhenius parameters from thermal analysis experiments (Viazovkin & Wight 1997). We write the kinetic equations under the general power law form

$$\begin{cases} \frac{dx(\text{HNCO})}{dt} = -k(T) \times x(\text{NH}_3)^\alpha \times x(\text{HNCO})^\beta \\ \frac{dx(\text{OCN}^-)}{dt} = k(T) \times x(\text{NH}_3)^\alpha \times x(\text{HNCO})^\beta, \end{cases} \quad (2)$$

where  $x(\text{NH}_3)$  and  $x(\text{HNCO})$  are the unitless molar fraction of NH<sub>3</sub> and HNCO,  $\alpha$  and  $\beta$  the partial orders of the reaction for NH<sub>3</sub> and HNCO, respectively, and  $k(T)$  is the rate of the reaction. To obtain a comprehensive expression for the rate of the reaction, we need to derive  $k(T)$ ,  $\alpha$  and  $\beta$  from the experimental data.  $\alpha$  and  $\beta$  are not necessarily equal to unity.

Figure 2 shows that the HNCO decay as a function of time does not reach zero. Because HNCO is highly diluted into NH<sub>3</sub>,  $x(\text{NH}_3)^\alpha \approx 1$ , and this factor is almost constant throughout the reaction.

The HNCO decay cannot be satisfactorily fitted by a first- ( $\beta = 1$ ) or second- order decay ( $\beta = 2$ ). To explain the peculiar decay curve shape and that the reaction yield is not 100%, we break down the HNCO decay into two processes: a fast process and a slow process.

This two-consecutive-step model allows us to isolate two phenomena occurring on two different timescales. The first step of the reaction is a slow process, which is interpreted as a change on geometrical configuration of HNCO within its NH<sub>3</sub> environment, since HNCO can be orientated in different configurations with respect to NH<sub>3</sub> (Benderskii et al. 1994). This slow process consists of HNCO rotating around its center of mass within its NH<sub>3</sub> environment, to achieve a configuration state, noted HNCO<sup>a</sup>, which corresponds to the configuration enabling the chemical reaction. This steering motion may be purely entropic, similar to a phase change, because the amorphous solid may be out of equilibrium and going toward a lower free enthalpy  $G$  corresponding to a more organized state. Reorientation (or isomerization) effects are well-known in rare gas cryogenic matrices (Benderskii et al. 1994). HNCO<sup>a</sup> is a chemically active configuration that corresponds to the orientation of HNCO with the hydrogen atom pointing toward the ammonia-free electronic doublet, so that the NH<sub>3</sub><sup>a</sup> actually reacting with HNCO<sup>a</sup> among the other NH<sub>3</sub> molecule will be this proton acceptor NH<sub>3</sub><sup>a</sup>. The activated complex is therefore the pair HNCO<sup>a</sup>-NH<sub>3</sub><sup>a</sup>. The other

non-chemically active configurations for HNCO are condensed in the notation  $\text{HNCO}^{\text{na}}$ . This change of orientation from a non-active  $\text{HNCO}^{\text{na}}$  configuration to the active  $\text{HNCO}^{\text{a}}$  configuration is modeled by a single barrier and a single rate  $k_o$ , although there must be a distribution of non-active configurations and a distribution of reorientation barriers. Since this change can be described by an unimolecular elementary process, its order is one. The fast process corresponds to a low activation barrier reaction and has a rate constant  $k_r$ .

This two-consecutive-step model is expressed in Eq. (3):



Because the  $\text{HNCO}^{\text{a}}$  proton donor is necessarily reacting with the  $\text{NH}_3$  proton acceptor in a one-on-one way, the decay rate of  $x(\text{HNCO}^{\text{a}})$  is proportional to the molar fraction of the  $\text{NH}_3$ - $\text{HNCO}^{\text{a}}$  pair, and thus to  $x(\text{HNCO}^{\text{a}})$ . We then can write the kinetic equations for the two consecutive phenomena as two consecutive first-order reactions:

$$\begin{cases} \frac{dx(\text{HNCO}^{\text{na}})}{dt} = -k_o(T) \times x(\text{HNCO}^{\text{na}}) \\ \frac{dx(\text{HNCO}^{\text{a}})}{dt} = k_o(T) \times x(\text{HNCO}^{\text{na}}) - k_r(T) \times x(\text{HNCO}^{\text{a}}). \end{cases} \quad (4)$$

There is no spectroscopic difference between  $\text{HNCO}^{\text{na}}$  and  $\text{HNCO}^{\text{a}}$ , since experimentally we record  $\text{HNCO}(t) = \text{HNCO}^{\text{na}}(t) + \text{HNCO}^{\text{a}}(t)$ . Solving Eq. (4) we can derive the analytical expression for the variation of the HNCO abundance as a function of time:

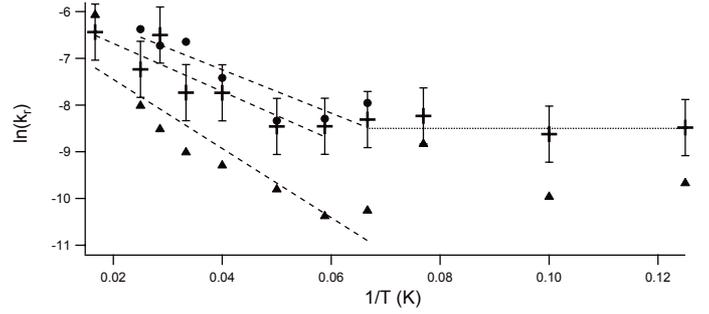
$$x(\text{HNCO})(t) = x(\text{HNCO}^{\text{na}})_{(t=0)} \left( 1 + \frac{k_o}{k_r - k_o} \right) \times (\exp(-k_o t) - \exp(-k_r t)) + x(\text{HNCO})_{(t=0)} \exp(-k_r t), \quad (5)$$

where  $x(\text{HNCO}^{\text{na}})_{(t=0)}$  and  $x(\text{HNCO}^{\text{a}})_{(t=0)}$  correspond to the initial population of HNCO in a chemically inactive or active configuration, and accordingly we have  $x(\text{HNCO}^{\text{na}})_{(t=0)} + x(\text{HNCO}^{\text{a}})_{(t=0)} = x(\text{HNCO})_{(t=0)} = 1$  because we normalized the decay curve to 1.

The activation energy of diffusion of HNCO within a  $\text{NH}_3$  ice should be on the order of magnitude of a few 10  $\text{kJ mol}^{-1}$ , the energy required to break few intermolecular hydrogen bonds (Park et al. 2010). As a comparison the activation energy of the self-diffusion in bulk ice is  $E_a^{\text{bulk}} = 71 \pm 4 \text{ kJ mol}^{-1}$  (George et al. 1997; Livingston et al. 1997). Therefore the diffusion of the molecules can be neglected between 8 and 60 K, on the time scale of our experiments, around 12 h typically. Moreover, because HNCO is a minor constituent in an  $\text{NH}_3$  ice, the two reactants are compulsorily in contact, and no diffusion is needed to make pairs of reactants.

Fitting the experimental data against the analytical expression of Eq. (5) at a fixed temperature, as shown in Fig. 2, allows one to extract  $k_o$ ,  $k_r$  and  $\text{HNCO}^{\text{na}}_{(t=0)}$  for this temperature. We derived these three parameters from several kinetic experiments performed at different fixed temperatures, as summarized in Table 2. Several measurements were made at a same temperature, 20 K, to estimate the dispersion of the results, which might have been caused by the inhomogeneity of our ice mixture deposition. The dispersion measured at 20 K was then assumed for all other temperatures.

We note that within the measurement dispersion, at the initial time the proportion of HNCO in a chemically active or non-active state, the  $x(\text{HNCO}^{\text{a}})_{(t=0)}$  and  $x(\text{HNCO}^{\text{na}})_{(t=0)}$  molar fractions do not vary with the temperature of the isothermal kinetic experiment as they should at the equilibrium. Because the



**Fig. 3.**  $\ln(k_r^{\text{RE}})$  (crosses),  $\ln(k_r^{\text{MC}})$  (circles) and  $\ln(\bar{k}_r)$  (triangles) as a function of the inverse of the temperature.  $k_r^{\text{RE}}$ ,  $k_r^{\text{MC}}$  and  $\bar{k}_r$  are the rates for the  $\text{HNCO} + \text{NH}_3 \rightarrow \text{NH}_4^+ + \text{OCN}^-$  reaction using a rate equation approach, a Monte Carlo approach, and a gamma distribution approach, respectively. The uncertainty on a point is given by the dispersion of the measurements at 20 K. The dashed lines represent the fit of the experimental points against an Arrhenius law for each approach.

$x(\text{HNCO}^{\text{a}})_{(t=0)}/x(\text{HNCO}^{\text{na}})_{(t=0)}$  ratio at the equilibrium is a function of the ratio between the difference of  $\text{HNCO}^{\text{a}}_{(t=0)}$  and  $\text{HNCO}^{\text{na}}_{(t=0)}$  free Gibbs energy and the equilibrium temperature, the relevant temperature ought to be constant for all experiments to get a constant  $x(\text{HNCO}^{\text{a}})_{(t=0)}/x(\text{HNCO}^{\text{na}})_{(t=0)}$  ratio. This temperature may be the room temperature at which the deposition of the two gases is performed, the room temperature proportion between  $\text{HNCO}^{\text{a}}_{(t=0)}$  and  $\text{HNCO}^{\text{na}}_{(t=0)}$  is frozen when the molecular solid is condensed onto the cold surface. At a 300 K room temperature, we found a difference in Gibbs energy of 4.3  $\text{kJ mol}^{-1}$ , which is on the order of magnitude of a molecular reorientation (Mielke et al. 2006, 2008). This means that our ice analogs, made from a room temperature gas quenching into an amorphous solid, are out of equilibrium. Therefore they must tend toward the equilibrium corresponding to the temperature at which each experiment is performed within a certain amount of time. This time scale is not astrophysically relevant since the making process of an interstellar ice in a dense molecular cloud is different. The orientation phenomenon is of the same physical nature as the reorganization of the out of equilibrium ice mixture, i.e. the breaking and reorganization of a hydrogen bond network, the hindered rotation of molecules within a solid. Consequently, the  $k_o$  rate constant should be temperature-dependent, although not necessarily equal to the hydrogen bond network reorganization rate. Although not astrophysically relevant, the  $k_o$  rate constant must be taken into account to correctly derive the  $k_r$  reaction rate.

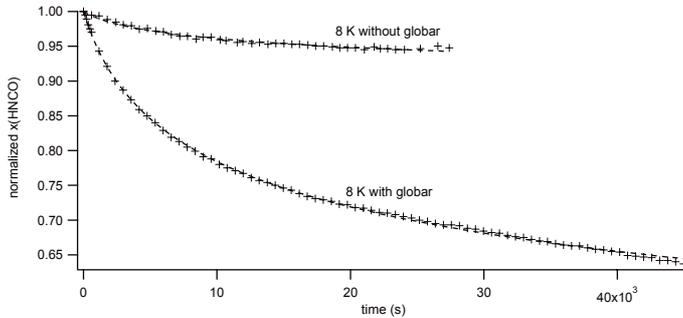
We now determine the temperature dependence of  $k_r(T)$ . Figure 3 shows that  $k_r(T)$  has two temperature domains. In the first domain, between 60 K and 15 K, the reaction rate can be fitted against an Arrhenius law  $k(T) = \nu_r^{\text{RE}} \times \exp\left(\frac{-E_r^{\text{RE}}}{RT}\right)$ . We found an activation energy  $E_r^{\text{RE}} = 0.4 \pm 0.1 \text{ kJ mol}^{-1}$  ( $48 \pm 12 \text{ K}$ ) and a pre-exponential factor  $\nu_r^{\text{RE}} = 0.0035 \pm 0.0015 \text{ s}^{-1}$  for the reaction between 60 K and 15 K. In the second domain between 15 K and 8 K, the reaction rate levels off, and the temperature dependence shows a plateau. A plateau like this in the reaction rate temperature dependence has been seen in many low-temperature systems (Benderskii et al. 1994) and can be interpreted as caused by a quantum tunneling reaction. This quantum tunneling regime coexists with the thermal regime, but becomes predominant at low temperature when the thermally induced reaction becomes minor.

The orientation phenomenon does not exhibit a dependence in temperature and has an average value of  $k_o = 3 \pm 1 \times 10^{-6} \text{ s}^{-1}$ , and consequently we cannot determine an activation energy  $E_o^{\text{RE}}$

**Table 2.** Experimental reaction rates (see text) for the HNCO + NH<sub>3</sub> → NH<sub>4</sub><sup>+</sup> + OCN<sup>-</sup> chemical reaction as a function of temperature.

$T(K)$	$k_o^{\text{RE}}$ (s <sup>-1</sup> )		Reaction rates					$\bar{k}_r$	$\beta$
			HNCO <sup>naRE</sup>	$k_o^{\text{MC}}$ (s <sup>-1</sup> )	$k_r^{\text{MC}}$ (s <sup>-1</sup> )	HNCO <sup>naMC</sup>			
8	$4.8 \times 10^{-6}$	$2.1 \times 10^{-4}$	0.77	–	–	–	$6.3 \times 10^{-5}$	0.15	
10	$4.9 \times 10^{-6}$	$1.8 \times 10^{-4}$	0.77	–	–	–	$4.7 \times 10^{-5}$	0.18	
13	$3.6 \times 10^{-6}$	$2.7 \times 10^{-4}$	0.77	–	–	–	$1.5 \times 10^{-4}$	0.09	
15	$7.6 \times 10^{-6}$	$2.5 \times 10^{-4}$	0.86	$2.7 \times 10^{-5}$	$3.5 \times 10^{-4}$	0.88	$3.5 \times 10^{-5}$	0.17	
17	$2.9 \times 10^{-6}$	$2.1 \times 10^{-4}$	0.89	$3.2 \times 10^{-6}$	$2.5 \times 10^{-4}$	0.89	$3.1 \times 10^{-5}$	0.07	
20	$8.3 \times 10^{-6}$	$3.1 \times 10^{-4}$	0.85	$1.0 \times 10^{-5}$	$2.4 \times 10^{-4}$	0.72	$5.5 \times 10^{-5}$	0.13	
20	$4.6 \times 10^{-6}$	$2.1 \times 10^{-4}$	0.77	–	–	–	$3.4 \times 10^{-4}$	0.07	
20	$2.8 \times 10^{-6}$	$2.4 \times 10^{-4}$	0.87	–	–	–	$2.4 \times 10^{-4}$	0.10	
20	$5.6 \times 10^{-6}$	$3.1 \times 10^{-4}$	0.85	–	–	–	$2.7 \times 10^{-4}$	0.13	
20	$5.5 \times 10^{-6}$	$2.3 \times 10^{-4}$	0.77	–	–	–	$2.0 \times 10^{-4}$	0.14	
25	$7.6 \times 10^{-6}$	$4.4 \times 10^{-4}$	0.85	$1.9 \times 10^{-5}$	$6.0 \times 10^{-4}$	0.85	$9.2 \times 10^{-5}$	0.09	
30	$2.7 \times 10^{-6}$	$4.4 \times 10^{-4}$	0.87	$3.8 \times 10^{-5}$	$1.3 \times 10^{-3}$	0.87	$1.2 \times 10^{-4}$	0.05	
35	$1.1 \times 10^{-6}$	$1.5 \times 10^{-3}$	0.91	$2.5 \times 10^{-5}$	$1.2 \times 10^{-3}$	0.88	$2.0 \times 10^{-4}$	0.04	
40	$3.1 \times 10^{-6}$	$7.2 \times 10^{-4}$	0.86	$1.5 \times 10^{-5}$	$1.7 \times 10^{-3}$	0.85	$3.3 \times 10^{-4}$	0.04	
60	$6.4 \times 10^{-6}$	$1.6 \times 10^{-3}$	0.59	$1.5 \times 10^{-5}$	$1.7 \times 10^{-3}$	0.85	$2.3 \times 10^{-3}$	0.10	

**Notes.** We were unable to fit some experimental data with the MC approach.



**Fig. 4.** Decay curves of  $x(\text{HNCO})(t)$  at 8 K, with and without the ice sample irradiated with the FTIR spectrometer beam in between two measures. The FTIR spectrometer clearly has an effect on the slow orientation process, while the fast reaction process is not affected.

and a pre-exponential factor  $\nu_o^{\text{RE}}$  for  $k_o$ . This indicates that the orientation phenomenon is non-thermal. This orientation phenomenon should be temperature-dependent since it is of the same physical nature as the reorganization of the out of equilibrium ice mixture. However, as shown in Fig. 4, the orientation phenomenon we observed in our experiments is non-thermal and induced by the mid-infrared photons from the mid-IR source of our FTIR spectrometer. This mid-IR source has a temperature of few thousand Kelvin and emits as a black-body in the mid-infrared. Figure 4 shows the difference between the HNCO band decay at 8 K if the ice mixture is irradiated in between two measurements by the mid-IR source beam and if the mid-IR source beam is blocked in between two measurements. Fitting these two decay curves with Eq. (5) gives the same  $k_r^{\text{RE}}$ , which means that the plateau in the  $k_r^{\text{RE}}$  temperature dependence is not induced by the IR photons, but is a real quantum tunneling. However,  $k_o^{\text{RE}}$  is extremely low at 8 K when the mid-IR source is blocked. When the mid-IR source is not blocked, the observed  $k_o^{\text{RE}}$  is non-thermal and induced by the IR photons of the spectrometer, and consequently does not have a temperature dependence. The FTIR photons are changing the distribution of populations from a Maxwell-Boltzmann thermal distribution to a distribution where IR photons can populate higher energy states of the complex, because the reorientation potential energy

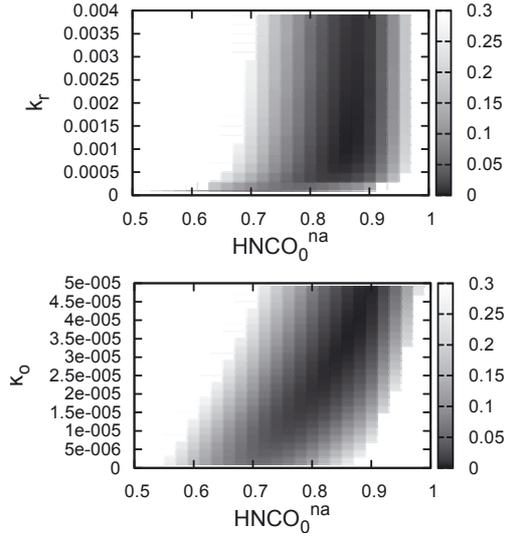
surface is probably very shallow. These states are closer to the reorientation barrier and consequently the kinetics is enhanced. Therefore, the spectrometer IR photons both cloud the temperature dependence of the orientation process and accelerate it. This IR-induced orientation is an experimental artifact and is not astrophysically relevant, but must be taken into account to correctly derive  $k_r$ .

This simple two-step model allows us to explain the HNCO time decay curve, especially the fact that the reaction does not reach completion, while irreversible (total). The activation barrier  $E_r^{\text{RE}}$  is small,  $0.4 \pm 0.1$  kJ mol<sup>-1</sup> ( $48 \pm 12$  K), accordingly the thermal reaction HNCO + NH<sub>3</sub> → NH<sub>4</sub><sup>+</sup> + OCN<sup>-</sup> needs to be taken into account in the interstellar ice reactions network.

### 3.3. Kinetic studies of the HNCO + NH<sub>3</sub> reaction using a kinetic Monte Carlo approach

In addition to using a deterministic rate equation approach as in Sect. 3.2, we used a kinetic Monte Carlo approach to fit the experimental kinetic data. Chemistry on grain mantle is by nature stochastic, because interstellar grains are cell-like objects and the non-abundant molecules (complex organic molecules mainly) population is small (sometimes  $\leq 1$  per grain) (Charnley 1998). Much work has been done to apply Monte Carlo methods to gas-grain chemistry (Charnley 1998; Lipshtat & Biham 2003; Chang et al. 2007) or to modify the rate equation formalism to take into account the stochastic effects (Garrod et al. 2009). While our experiments are not in the low-abundance limiting case (HNCO is diluted into NH<sub>3</sub> in a 1:20 concentration ratio within a single macroscopic sample and thus we are definitely within the thermodynamic limit), we used the Monte Carlo formalism applied to coupled chemical reactions (Gillespie 1976) to compare it with the rate equation formalism in the simplest case possible: only one two-body solid-state reaction with no diffusion effect.

To do this, we used the PIZZAZZ code (Jansen 1995). PIZZAZZ is a general purpose program, developed by Dr. Jansen, and designed to simulate reactions on 3D regular grids. It is an implementation of the “first reaction method” (Gillespie 1976). Our simulations were performed using a 3D grid of  $100 \times 100 \times 100$  sites to consider bulk reactions in



**Fig. 5.** Chi square surface in grayscale, resulting from a fit of experimental HNC0 decay at 20 K as a function of the activation rate  $k_r$  and  $\text{HNC0}_{(t=0)}^{\text{na}}$  (top) and  $k_0$  and  $\text{HNC0}_{(t=0)}^{\text{na}}$  (bottom). There is no local minimum, the uncertainty on the optimized parameters is small, and little correlation is observed between  $k_0$  and  $\text{HNC0}_{(t=0)}^{\text{na}}$ .

the ice volume. The initial configuration is set randomly according to an initial reactants ratio  $\text{HNC0}/\text{NH}_3$  that was measured on the experimental spectra, and we completely filled the grid according to this ratio. We saw in Sect. 3.2 that the evolution of HNC0 abundance can be explained by two processes, a slow process that activates the HNC0 molecules and a fast process corresponding to the reaction of activated HNC0 with  $\text{NH}_3$ . We also saw that there was a non-trivial ratio of  $\text{HNC0}^{\text{a}}/\text{HNC0}^{\text{na}}$  at the beginning of the reaction. Because this quantity cannot be measured, it was kept unconstrained. Therefore, we have three parameters to explore: the reaction rate  $k_r$ , the activation rate  $k_0$ , and  $\text{HNC0}_{(t=0)}^{\text{na}}$  since  $\text{HNC0}_{(t=0)}^{\text{na}} + \text{HNC0}_{(t=0)}^{\text{a}} = 1$ . These three parameters were varied to fit the experimental data. The chi square surface resulting from the fit optimization was explored over the parameters space as shown in Fig. 5. We can see that there is a global minimum and that the uncertainty caused by the stochastic fitting procedure is smaller than the experimental dispersion, which is assumed to be the same as the dispersion at 20 K determined using the rate equation approach, because it is mainly caused by an experimental reproducibility due to the heterogeneous nature of the ice samples. We can also see that there is a defined global minimum between  $k_r$  and  $\text{HNC0}_{(t=0)}^{\text{na}}$ , which fixes  $\text{HNC0}_{(t=0)}^{\text{na}}$ , while there is a correlation between  $k_0$  and  $\text{HNC0}_{(t=0)}^{\text{na}}$ , which leads to a less reliable value of  $k_0$ . The results of the fits using the Monte Carlo approach are summarized in Table 2.

We can now determine the temperature dependence of  $k_0^{\text{MC}}(T)$  and  $k_r^{\text{MC}}(T)$  by fitting the reaction rates at different temperatures against an Arrhenius law, as shown in Fig. 3. As observed previously in the rate equation approach, there is no noticeable variation of  $k_0^{\text{MC}}$  with temperature and we find a rate  $k_0^{\text{MC}} = 5 \pm 4 \times 10^{-5} \text{ s}^{-1}$  for the orientation phenomenon. We find an activation energy  $E_r^{\text{MC}} = 0.4 \pm 0.2 \text{ kJ mol}^{-1}$  ( $48 \pm 24 \text{ K}$ ) and a pre-exponential factor  $\nu_r^{\text{MC}} = 0.005 \pm 0.003 \text{ s}^{-1}$  for the reaction rate.

Within the uncertainty, the same activation energy and pre-exponential factor are found for the reaction using the rate equation approach and the Monte Carlo approach. This equality of

rates on one thermal reaction within a macroscopic ice sample is significant because it is a first step to validate the equivalence between using a rate equations approach or a Monte Carlo approach in a gas-grain chemistry code.

### 3.4. Kinetic studies of the $\text{HNC0} + \text{NH}_3$ reaction using a gamma probability distribution of orientations

To go beyond the simple model expressed in Eq. (3) and to account for the fact that there is no single  $\text{HNC0}^{\text{na}}$  state but rather a distribution of  $\text{HNC0}^{\text{na}}$  states, we considered a distribution of HNC0 configurations, each one of them with its own  $k_r$  reaction constant. The difference between  $k_r$  reaction constants stems from the difference in the orientation energy of each site, the reaction rate of the activated complex being the same, following the two-step process approach described in Eq. (3). However, we considered that a distribution of  $k_0$  that is translated into a distribution of  $k_r$ . Accordingly, we can write the HNC0 decay rate as a sum of exponential decays:

$$x(\text{HNC0}_{(t)}) = x(\text{HNC0}_{(t=0)}) \times \int_0^\infty e^{-k_r \times t} g_{\bar{k}_r, \beta}(k_r) dk_r \quad (6)$$

$$= \frac{x(\text{HNC0}_{(t=0)})}{\left(1 + \frac{\bar{k}_r}{\beta} \times t\right)^\beta}, \quad (7)$$

where the  $k_r$  rates are distributed according to a gamma probability distribution function with a  $\bar{k}_r$  mean value

$$g_{\alpha, \beta}(k_r) = k_r^{(\beta-1)} \times \frac{\exp\left(-\frac{k_r}{\alpha}\right)}{\alpha^\beta \Gamma(\beta)}, \quad (8)$$

where  $\alpha$  and  $\beta$  are the scale and the shape factors respectively, and we have  $\bar{k}_r = \alpha \times \beta$  (Johnson et al. 1995).

The gamma probability distribution function is a generalization of the chi-square probability distribution function, i.e.  $\beta = n/2$  can be non-integer and lower than 1,  $n$  being the degree of freedom of the chi-square distribution (Johnson et al. 1995). The integration of this distribution gives a power law, as seen in Eq. (6).

The chi-square distribution has been previously used to describe dynamical processes in many different systems, such as intramolecular vibrational redistribution (Uzer & Miller 2004) or relaxation processes in amorphous solids (Cannizzo et al. 2008). These processes are described by the random-matrix theory, which applies well to chaotic or disordered systems. We tentatively give a physical explanation for the good agreement of our experimental curve with a gamma probability distribution function, based on two assumptions. The first assumption is that each individual non-active  $\text{HNC0}_j^{\text{na}}$  site corresponds to one particular geometrical orientation of the molecules, with one possible realization of  $c_{\text{na/a}}$ , the coupling strength between the  $\text{HNC0}_j^{\text{na}}$  non-active state and the  $\text{HNC0}^{\text{a}}$  pre-reaction complex. At high energy, close to the top of the reaction barrier, the  $c_{\text{na/a}}$  coupling strengths are independently distributed, according to the central limit theorem, with a Gaussian probability  $P(c_{\text{na/a}})$  around  $c_{\text{na/a}}^j$ , the average coupling strength, and we have  $k_r^j \propto |c_{\text{na/a}}^j|^2$ . The average coupling strength  $c_{\text{na/a}}^j$  can be physically seen as the average overlap between  $\text{HNC0}_j^{\text{na}}$  and  $\text{HNC0}^{\text{a}}$  wavefunctions. The second assumption is that the inhomogeneous ensemble of the sites corresponds to a large and statistically independent sample of all possible random realizations of  $k_r$ . Thus, the orientation rate  $k_r$  is the sum of the squares

of the independent, standard normal random variables, and as such is distributed according to a gamma distribution with  $2\beta$  degrees of freedom and an average value  $\bar{k}_r$ .  $2\beta$  is not directly the number of non-active sites, but rather the degree of disorder of the system. In the limit  $\beta \gg 1$ , the distribution tends to a narrowly peaked Gaussian, which would correspond to a highly disordered system, where all configurations are explored during the measurement process. In our case, the number of orientational configurations HNCO<sub>*j*</sub><sup>na</sup> is large, but finite; the system is disordered, but there is no fast conversion between the different configurations like in liquids or solids, which corresponds to a low- $\beta$  case. This case corresponds to high fluctuations around the  $\bar{k}_r$  mean value, which is shown by the high fluctuations in the results measured at 20 K.

Fitting the experimental HNCO decay curve against Eq. (6), as seen in Fig. 2, gives  $\beta$  and  $\bar{k}_r$  at each temperature, as shown in Table 2. We find  $\beta$  around 0.10, varying between 0.05 and 0.18, since the population of the different HNCO<sub>*j*</sub><sup>na</sup> sites can vary from one ice mixture deposition to another. However, Fig. 4 shows how the populations can be altered by the spectrometer IR photons. Fitting Fig. 4 with Eq. (6) gives the same  $\bar{k}_r$  but a  $\beta$  parameter different by a factor of seven. We find a thermal regime between 60 K and 15 K and a quantum regime between 15 K and 8 K for the  $\bar{k}_r$  value, like in the rate equation and Monte Carlo approaches. Fitting the  $\bar{k}_r$  reaction rate temperature dependence against an Arrhenius law between 60 K and 15 K, as shown in Fig. 3, gives an activation energy  $E_r = 0.61 \pm 0.02$  kJ mol<sup>-1</sup> (73 ± 3 K) and a pre-exponential factor  $\nu_r^{\text{MC}} = 0.003 \pm 0.002$  s<sup>-1</sup>. Between 15 K and 8 K,  $\bar{k}_r = 2 \times 10^{-4}$  s<sup>-1</sup>. These values are similar, though slightly lower, to the values found using the rate equation and Monte Carlo approaches. This simple statistic approach has the advantage of taking into account the way an ice mixture is grown into the  $\beta$  parameter and extracting a  $\bar{k}_r$  parameter insensitive to experimental procedures. It also underlines the sensitivity of the derived kinetic constants to experimental conditions and the care that needs to be taken to derive a reliable kinetic rate.

#### 4. Astrophysical implications

We have derived the expressions for the kinetic laws and reaction rates of NH<sub>4</sub><sup>+</sup> OCN<sup>-</sup> formation from the HNCO + NH<sub>3</sub> isolated reaction from experimental kinetic data for temperatures between 8 K and 60 K using a rate equation approach, a Monte Carlo approach and a distribution approach, as summarized in Table 2. Although conceptually very different, the deterministic and the stochastic approaches give similar results given the uncertainties in the fitting procedures and the dispersion of the experimental results. This shows that in ice analogs as well the master equation probabilities can be approximated by rate constants, derived from a simple Arrhenius law. The former approach gives empirically derived rate constants, while the latter approach gives physical insight and explains why the reaction is not complete. We also tried to describe the distribution of HNCO orientation configurations with respect to NH<sub>3</sub> using a gamma probability distribution, which gives similar results to the previous approaches.

Whether or not this purely thermal mechanism is relevant to OCN<sup>-</sup> formation in the interstellar medium is a question, first, of the kinetics of this reaction and of other competing reactions, such as photodissociations or hydrogenations of HNCO at a particular temperature; and second, of the physical environment the ice is located in. Indeed, in different interstellar environments,

the grains are subjected to different temperatures, photon fluxes, and density conditions, and solid-state OCN<sup>-</sup> can be formed or destroyed accordingly. We will discuss the production of OCN<sup>-</sup> from the HNCO + NH<sub>3</sub> reaction first in a dense molecular cloud and then in a YSO, using a gas-grain model.

To take into account the influence of H<sub>2</sub>O and all competing reactions in a dense molecular cloud, we used a rate equation gas-grain model (Hasegawa et al. 1992) along with the physical parameters used in Hasegawa et al. (1992): hydrogen density  $N(\text{H}) = n(\text{H}) + n(\text{H}_2) = 2 \times 10^4$  cm<sup>-3</sup>, dust-to-gas mass ratio 1%, grain density 3 g cm<sup>-3</sup>, 0.1 micron grains, surface sites of adsorption 10<sup>6</sup>, a high visual extinction  $A_v = 10$ , and the initial gas-phase conditions with all hydrogen into its molecular form at the initial time. We used the most recent reactions database for the gas phase (OSU database) and the solid phase (Garrod et al. 2008). The HNCO + H<sub>2</sub>O and H<sub>3</sub>O<sup>+</sup> + OCN<sup>-</sup> reactions measured in an earlier work (Theule et al. 2011), along with the HNCO + NH<sub>3</sub> reaction measured in this work, were added to this reaction network with their measured activation energies. We assumed the  $E_r = 0.4 \pm 0.1$  kJ mol<sup>-1</sup> (48 ± 12 K) activation energy value for the HNCO + NH<sub>3</sub> reaction, excluding the slow phenomenon, which is of the same nature as diffusion, so that we did not account for it twice.

The HNCO + H<sub>2</sub>O and HNCO + NH<sub>3</sub> reactions are the only ways of producing solid-state OCN<sup>-</sup> in the gas-grain models. We have measured that the HNCO + NH<sub>3</sub> reaction rate at 10 K, the typical temperature of a molecular cloud, is  $k(10 \text{ K}) \approx 1.8 \times 10^{-4}$  s<sup>-1</sup>. When accounting for the diffusion of HNCO and NH<sub>3</sub>, following the usual expression for the thermal diffusion of the reactants species on water ice, expressed in Eq. (9) from Hasegawa et al. (1992), this rate decreases to  $2 \times 10^{-37}$  cm<sup>-3</sup> s<sup>-1</sup>. The reaction is therefore very slow, and not efficient enough to counterbalance other reactions such as the hydrogenation of HNCO, as previously noticed in Theule et al. (2011). OCN<sup>-</sup> is not formed in a typical molecular cloud in  $3 \times 10^5$  years. Hydrogenation reactions are orders of magnitude faster, and as such dominate solid HNCO surface chemistry. Therefore OCN<sup>-</sup> cannot be formed in a cold molecular cloud, as confirmed by background star observations (Knez et al. 2005).

We then investigated if OCN<sup>-</sup> can be formed from the HNCO + NH<sub>3</sub> reaction in a YSO. At  $T_{\text{grain}} = 40$  K, which is a typical temperature in a YSO (van Broekhuizen et al. 2004), and a sufficiently low temperature to encounter water ice and HNCO on a YSO grain, we measured  $k(40 \text{ K}) \approx 1 \times 10^{-3}$  cm<sup>-3</sup> s<sup>-1</sup> ( $6 \times 10^{-4}$  cm<sup>-3</sup> s<sup>-1</sup> when corrected from the diffusion). Moreover, because the surface temperature is warmer, the residence time of H on grains is shorter and hydrogenation reactions are not dominant anymore. Then, solid OCN<sup>-</sup> can reach an important value, approximately 40% with respect to solid H<sub>2</sub>O in  $3 \times 10^5$  years. This result is two orders of magnitude above typically observed OCN<sup>-</sup> quantities in YSOs (van Broekhuizen et al. 2005). This quantitative overestimation can be explained by the poor knowledge we have on how to calculate the diffusion of the reactants, by the uncertainty on the reaction rates in a genuine interstellar ice environment, on the solid HNCO quantity, or by the fact that astronomical ice features are integrated along a line of sight where solid OCN<sup>-</sup> abundance can fluctuate. However, we can qualitatively account first for the observations of the OCN<sup>-</sup> ice feature and second for the non-observation of solid-state HNCO in a YSO. Indeed, according to Garrod et al. (2008) solid HNCO is mainly formed from the NH + CO thermal reaction, which has a  $4 \times 10^{-12}$  cm<sup>-3</sup> s<sup>-1</sup> reaction rate at 40 K and  $A_v = 10$ , and is therefore very efficient in forming HNCO. This mechanism has been

qualitatively confirmed from NH<sub>3</sub>:CO ices VUV irradiation experiments (Grim et al. 1989). However, solid HNCO abundance is calculated to remain very low, approximately 10<sup>-10</sup> with respect to solid H<sub>2</sub>O, and therefore cannot be detected in infrared spectra. Once formed, HNCO is an intermediate species that reacts very quickly with NH<sub>3</sub> to give OCN<sup>-</sup>, which explains why OCN<sup>-</sup> is detected, while HNCO is not.

## 5. Conclusion

We experimentally measured the kinetics of the HNCO + NH<sub>3</sub> isolated solid-state reaction in a HNCO:NH<sub>3</sub> ice at different temperatures between 8 K and 60 K. We derived the expressions for the kinetic laws and the temperature dependence of the OCN<sup>-</sup> formation rate using a two-consecutive-step model. The first step is a slow first-order process, where the two reacting molecules are spatially orientating around their center of mass in a configuration that allows the chemical reaction, with a rate  $k_0(T)$ . This orientation process is an experimental artifact because interstellar ice analogs are out of equilibrium amorphous solids. Its barrier is so low that it is influenced by the IR photons of our spectrometer. The second step is a fast first-order process, where the proton is transferred from the HNCO nitrogen atom to the NH<sub>3</sub> nitrogen atom with a rate  $k_r(T)$ . We used a deterministic approach based on a rate equation formalism and a stochastic approach based on a kinetic Monte Carlo formalism and on a gamma probability distribution function to derive the reaction rate. The reaction rate follows an Arrhenius law,  $k_r(T) = \nu_r \times \exp(-E_r/RT)$  with  $\nu_r = 0.0035 \pm 0.0015 \text{ s}^{-1}$  and  $E_r = 0.4 \pm 0.1 \text{ kJ mol}^{-1}$  (48 ± 12 K), between 60 K and 15 K and exhibits a quantum tunneling reaction between 15 K and 8 K with a  $2 \times 10^{-4} \text{ s}^{-1}$  rate independent from the temperature. The three approaches, although conceptually different, give a similar activation energy and pre-exponential factor. This specifically shows that it is equivalent to use a deterministic rate equation approach or a Monte Carlo approach to model a thermal reaction within a macroscopic ice sample. Using the measured reaction rate and a gas-grain model, we qualitatively showed that the HNCO + NH<sub>3</sub> reaction is too slow to compete with hydrogenation reactions in a 10 K molecular cloud, while it is fast enough to produce observed OCN<sup>-</sup> in young stellar object interstellar ices at an average temperature of 40 K without producing observable quantities of HNCO. Achieving a more quantitative agreement to account for the OCN<sup>-</sup> observations in different astrophysical environments will require a better understanding of how solid-state HNCO is formed and hydrogenated and how the diffusion in H<sub>2</sub>O dominated ices furthermore limits chemical reaction rates. Indeed, in a real H<sub>2</sub>O dominated astrophysical ice, H<sub>2</sub>O is dramatically slowing down the kinetics of the HNCO + NH<sub>3</sub> reaction by forcing the reactants to diffuse to meet each other.

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