1. Introduction

Observations of the dust and gas components in the circumstellar envelopes (CSEs) around asymptotic giant branch (AGB) stars have been carried out at different wavelengths. Observations in the infrared trace the dust as well as the warm molecular layer close to the stellar photosphere (e.g., Justtanont et al. 1996; Aoki et al. 1999; Schöier et al. 2002). Submillimeter and radio observations of trace molecules have been used to study the cooler outer parts of the CSEs (e.g., Knapp & Morris 1985; Schöier & Olofsson 2000; Kemper et al. 2003). To bridge this gap, the ISO was used to observe a large number of AGB stars up to almost 200 μm, and in O-rich stars, a number of H2O emission lines were detected (Barlow et al. 1996; Neufeld et al. 1996). However, the circumstellar far-infrared lines were unresolved. Hence, crucial information about the line profiles remained unknown.

Water is an important molecule in CSEs as it is thought to be one of the main cooling agents in the wind. It is also expected to be a good probe of the inner regions of the CSE where the gas is accelerated. However, to fully explore the potential of H2O lines as a probe of the circumstellar gas, a full radiative transfer model has to be performed. Owing to difficulties in calculating accurate collisional rates coupled with the very high optical depth of the H2O lines in the inner region of the CSE, slow progress has been made. Nevertheless, calculations of the heating and cooling in the CSEs of O-rich stars suggest that H2O dominates the cooling in most parts of the envelope until it is photodissociated by interstellar UV photons (Goldreich & Scoville 1976; Justtanont et al. 1994; Maercker et al. 2008, 2009; Decin et al. 2010a), and that some lines should come mainly from the acceleration zone. Eventually, spectrally resolved circumstellar H2O lines were observed by two space missions dedicated to search for cosmic water line emission: SWAS and Odin. Both missions were able to detect the ground-state line of H2O at 557 GHz in a number of AGB stars (Harwit & Bergin 2002; Melnick et al. 2001; Justtanont et al. 2005; Hasegawa et al. 2006; Maercker et al. 2009). It was shown that not only the line intensity, but also the line profile is crucial for interpreting the data correctly.

In 2009, the ESA-Herschel Space Observatory (Pilbratt et al. 2010) was launched with the Heterodyne Instrument for the Far-Infrared (HIFI, de Graauw et al. 2010), which aims to study H2O line emission in different environments in our Galaxy and beyond.
Table 1. HIFI observations of χ Cyg.

<table>
<thead>
<tr>
<th>Molecule</th>
<th>Transition</th>
<th>ν (GHz)</th>
<th>$E_{up}$ (K)</th>
<th>$I^*$(K km s$^{-1}$)</th>
</tr>
</thead>
<tbody>
<tr>
<td>CO</td>
<td>$J = 6$–$5$</td>
<td>691.473</td>
<td>116</td>
<td>15.3</td>
</tr>
<tr>
<td>CO</td>
<td>$J = 10$–$9$</td>
<td>1151.985</td>
<td>304</td>
<td>13.6</td>
</tr>
<tr>
<td>CO</td>
<td>$J = 16$–$15$</td>
<td>1841.346</td>
<td>752</td>
<td>13.5</td>
</tr>
<tr>
<td>$o$-H$_2$O</td>
<td>$1_{10}$–$1_{10}$</td>
<td>556.936</td>
<td>61</td>
<td>3.0</td>
</tr>
<tr>
<td>$p$-H$_2$O</td>
<td>$2_{11}$–$2_{10}$</td>
<td>752.033</td>
<td>137</td>
<td>4.2</td>
</tr>
<tr>
<td>$p$-H$_2$O</td>
<td>$2_{01}$–$1_{11}$</td>
<td>987.927</td>
<td>101</td>
<td>7.3</td>
</tr>
<tr>
<td>$p$-H$_2$O</td>
<td>$1_{10}$–$0_{0}$</td>
<td>1113.343</td>
<td>53</td>
<td>8.0</td>
</tr>
<tr>
<td>$o$-H$_2$O</td>
<td>$3_{12}$–$2_{21}$</td>
<td>1153.127</td>
<td>249</td>
<td>7.5</td>
</tr>
<tr>
<td>$o$-H$_2$O</td>
<td>$3_{21}$–$3_{12}$</td>
<td>1162.912</td>
<td>305</td>
<td>1.5</td>
</tr>
<tr>
<td>$o$-H$_2$O</td>
<td>$3_{30}$–$2_{12}$</td>
<td>1716.770</td>
<td>197</td>
<td>18.1</td>
</tr>
</tbody>
</table>

Notes. (a) The absolute calibration accuracy is between 10% to 30% (see text).

Fig. 1. The derived gas temperature structure (solid line) and the gas expansion velocity (dashed line) of the CSE of χ Cyg.

3. Modelling HIFI lines

We started the analysis by fitting the spectral energy distribution of the CSE (assumed to be spherically symmetric) using Dusty (Ivezic & Elitzur 1997) to derive the dust mass-loss rate. Based on this, we solved the gas-dust drag equation, assuming that both are momentum coupled, to derive the velocity structure using the observed terminal gas velocity as a constraint, as shown in Fig. 1. This also provides a so-called dynamical mass-loss rate estimate, $4.9 \times 10^{-7} M_\odot yr^{-1}$ in the case of χ Cyg (see e.g., Ramstedt et al. 2008). The gas velocity law is fed into the CO radiative transfer model, where the line fluxes and shapes are computed and compared to the observations. At the same time, the heating by dust grains and the cooling by CO lines are calculated and the resulting temperature structure obtained (e.g., Justtanont et al. 1994; Closn & Menten 1997; Decin et al. 2006; van der Tak et al. 2007; Ramstedt et al. 2008). This provides a circumstellar model, including a mass-loss-rate estimate based on the CO line modelling, which is used to model the H$_2$O line emission and its contribution to the cooling. The H$_2$O cooling is then used to recalculate the gas kinetic temperature in the CO model, and the process is iterated until good fits to the observed CO and H$_2$O lines are obtained.

3.1. Modelling of the CO lines

The Monte Carlo code developed by Schöier & Olofsson (2001) was used to model the observed CO lines. The molecular data were taken from the Hitran database (Rothman et al. 2009) and the collisional rate coefficients from Yang et al. (2010) for the 41 lowest rotational levels in the $ν = 0$ vibrational state. We fitted the line shapes and fluxes for the low-$J$ lines obtained from ground-based observations, as well as the interferometric data for the $J = 1$–0 and 2–1 lines to more tightly constrain the size of the CO envelope (Schöier et al, in prep.). The parameters are listed in Table 2. The gas kinetic temperature distribution is shown in Fig. 1.

We assume a distance of 150 pc (Knapp et al. 2003), and the resulting mass-loss rate is $7 \times 10^{-7} M_\odot yr^{-1}$ using an adopted CO abundance of $6 \times 10^{-4}$ (relative to H$_2$, see Table 2), which provide the best fits to both the line profile and line intensities. The uncertainty in the mass-loss rate is of the order 50%. This mass-loss rate agrees well with the dynamical mass-loss rate (certainly to within the errors in the input parameters). A comparison of the model fits and the HIFI observations can be further discussed in the context of the circumstellar model.
Table 2. Parameters used in the modelling of the line emission.

<table>
<thead>
<tr>
<th>Parameter</th>
<th>Value</th>
</tr>
</thead>
<tbody>
<tr>
<td>Distance</td>
<td>150 pc</td>
</tr>
<tr>
<td>Stellar effective temperature ($T_{\text{eff}}$)</td>
<td>2600 K</td>
</tr>
<tr>
<td>Stellar luminosity ($L_*$)</td>
<td>$7.5 \times 10^3 L_\odot$</td>
</tr>
<tr>
<td>Gas terminal velocity ($v_{\text{gas}}$)</td>
<td>8.5 km s$^{-1}$</td>
</tr>
<tr>
<td>Inner radius of the shell ($R_\infty$)</td>
<td>$2 \times 10^{15}$ cm</td>
</tr>
<tr>
<td>Gas mass-loss rate ($\dot{M}$)</td>
<td>$7 \times 10^{-7} M_\odot$ yr$^{-1}$</td>
</tr>
<tr>
<td>CO abundance (CO/H$_2$)</td>
<td>$6 \times 10^{-4}$</td>
</tr>
<tr>
<td>ortho-H$_2$O abundance (ortho-H$_2$O/H$_2$)</td>
<td>$7.5 \times 10^{-6}$</td>
</tr>
<tr>
<td>para-H$_2$O abundance (para-H$_2$O/H$_2$)</td>
<td>$3.6 \times 10^{-6}$</td>
</tr>
</tbody>
</table>

3.2. Modelling of the H$_2$O lines

To calculate the strength and shape of the circumstellar H$_2$O lines, we apply the parameters derived from our CO line modelling (mass-loss rate, gas temperature and density structure, and gas velocity law) to the radiative transfer model for H$_2$O based on an accelerated lambda iteration (ALI) code (Justtanont et al. 2005; Maercker et al. 2008, 2009). We used the molecular data from Rothman et al. (2009) and the collisional cross-sections from Faure et al. (2007) for the lowest 45 levels of ortho- and para-H$_2$O. The radiative excitation due to the absorption in the $v_3$ bending and $v_3$ stretching modes is included. The latter has been found to have a non-negligible effect in the low mass-loss-rate case (Maercker et al. 2009). The outer radius of the H$_2$O shell was derived using the model results of Netzer & Knapp (1987), i.e., $3.6 \times 10^{15}$ cm. Since the ortho- and para-species are expected to be independent, we model the two species separately, using the same circumstellar input values and estimate the two abundances independently. Using the temperature and velocity structure from the CO modelling (Fig. 1), we calculate the best-fit model to the H$_2$O lines. All the lines are found to be sub-thermally excited. As mentioned above, both CO and H$_2$O line cooling is taken into account.

We present the fits to four lines of ortho-H$_2$O and three lines of para-H$_2$O observed with HIFI in Fig. 3. These lines span upper energies from 60 to 300 K so the lines probe the cool part of the CSE as well as the warmer inner part. From Figs. 1 and 4, it can be seen that H$_2$O lines originate well within the acceleration zone. All lines are of reasonably high signal-to-noise ratio, including the $3_2^0-2_1^2$ line, the highest frequency line, where the spectrum is affected by standing waves inside HIFI. For this line, we used a higher order Chebyshev polynomial for the baseline subtraction (bottom right panel of Fig. 3).

Our results are consistent with the velocity structure of a dust-driven wind as can be seen in good fits to both the CO and H$_2$O line profiles (Figs. 2 and 3, respectively). Unlike the case for IK Tau (Decin et al. 2010a,b), no modification to the dynamical calculation is required.

4. Results and conclusions

The observed HIFI lines are reliable probes of the inner CSE as the high-energy lines probe the wind in the acceleration zone. Both the velocity and density structures are tightly constrained using the HIFI lines. The abundance of CO used is $6 \times 10^{-4}$, intermediate to those usually adopted for O- and C-rich CSEs. The derived ortho- and para-H$_2$O abundances are significantly lower, $(7.5 \pm 1.4) \times 10^{-6}$ and $(3.6 \pm 0.5) \times 10^{-6}$, respectively (Table 2). These values are well below the limits for O-rich AGB stars of $>10^{-4}$ (Justtanont et al. 2005; Maercker et al. 2008, 2009) consistent with $\chi$ Cyg being an S-star of C/O very close to unity.

From our modelling, assuming that all carbon is locked up in CO (i.e., C/H = $3 \times 10^{-4}$) and the oxygen is locked up in both CO and H$_2$O (i.e., O/H = $3 \times 10^{-4}$ + $5.5 \times 10^{-5}$), our derived C/O is $<0.98$, given that a small fractional abundance of the oxygen is in dust grains. This value is slightly higher than that of 0.95, assumed by Duari & Hatchell (2000). A non-thermal equilibrium chemistry model for S-stars (C/O = 0.98) predicts an H$_2$O abundance of $10^{-6}$ at the stellar photosphere, falling off to a few $10^{-6}$ at 5 $R_*$ (Cherchneff 2006), an order of magnitude lower than our value.

In the thermal equilibrium (TE) limit at high temperature, the expected ortho-to-para ratio is 3. Our derived ortho-to-para ratio is $2.1 \pm 0.6$, close to the high-temperature TE value. The reported ortho-to-para ratio in CSEs of O-rich stars vary from 1 in W Hya with a large uncertainty (Barlow et al. 1996) to 3 in IK Tau (Decin et al. 2010b). Our result is consistent with H$_2$O molecules being formed under thermal equilibrium conditions in the warm and dense stellar photosphere.

Given the low total H$_2$O abundance of $(1.1 \pm 0.2) \times 10^{-5}$, it is clear from our analysis that the dominating cooling agent in the CSE of $\chi$ Cyg is CO (Fig. 4). Vibrationally excited H$_2$ and rotationally excited H$_2$O contribute only in the innermost ($r < 10^{15}$ cm) part of the CSE while CO line cooling extends further out until CO is photodissociated by the external UV field. Adiabatic cooling dominates only in the outermost part of the CSE. The derived H$_2$O abundance, although much lower than in O-rich stars, is higher than that observed in C-stars, IRC+10216 (Melnick et al. 2001) and V Cyg (Neufeld et al. 2010), consistent with AGB stars of S-type being chemically intermediate between O-rich and C-rich AGB stars.
Fig. 3. The model fits (smooth lines) to the HIFI ortho- and para-H$_2$O lines (histogram). The scaling factor (used to scale the model result to fit the observed integrated intensity) is given for each line, showing the goodness of the fit. The pipeline-reduced data of the $3_{03} - 2_{11}$ line are plotted along with the fitted baseline at the bottom right. This line is affected by standing waves within HIFI.

Fig. 4. The cooling rates due to different processes: rotational cooling by H$_2$O (solid), CO (dashed) lines, vibrational cooling by H$_2$ (dotted) line, and adiabatic cooling (dot-dash).

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