

XMMU J134736.6+173403: an eclipsing LMXB in quiescence or a peculiar AGN? (Research Note)

S. Carpano¹, B. Altieri², A. R. King³, A. Nucita¹, and P. Leisy⁴

¹ XMM-Newton Science Operations Centre, ESAC, ESA, PO Box 50727, 28080 Madrid, Spain
e-mail: scarpano@sciops.esa.int

² Herschel Science Centre, ESAC, ESA, PO Box 50727, 28080 Madrid, Spain

³ Department of Physics and Astronomy, University of Leicester, Leicester LE1 7 RH, UK

⁴ Isaac Newton Group of Telescopes, Apartado de Correos 321, 38700 Sta. Cruz de La Palma, Spain;
Instituto de Astrofísica de Canarias, 38205 La Laguna, Tenerife, Spain

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ABSTRACT

Aims. We report the discovery of a peculiar object observed serendipitously with *XMM-Newton*. We present its timing and spectral properties and investigate its optical counterpart.

Methods. The light curve of the X-ray source, its spectrum, and the spectrum of the best optical counterpart are presented and analyzed.

Results. The X-ray flux decreases by a factor of 6.5 within 1 h and stays in a low state for at least 10 h, thereby suggesting the presence of an eclipse. The spectrum is very soft, a power law with a slope of $\Gamma \sim 2.8$, and does not change significantly before and after the flux drop. The source is spatially coincident within few arc-seconds with a Seyfert 2 galaxy belonging to a galaxy pair.

Conclusions. Although the background AGN seems the best counterpart, neither the temporal nor the spectral properties of the X-ray source are compatible with it. We investigate the possibility of having a foreground low-mass X-ray binary in quiescence, where the companion is not detected in the optical wavelength.

Key words. X-rays: general – X-rays: binaries – galaxies: Seyfert

1. Introduction

In this paper we report the discovery of an X-ray source for which the flux drops abruptly by a factor of 6.5 within one hour, and stays in this low state for at least 10 h. The nature of the source described in this paper is very ambiguous. The X-ray light curve suggests the presence of an eclipse and hence a binary system. The detected optical counterpart is a pair of galaxies, one of which is a Seyfert 2 AGN. The X-ray spectrum is also very soft so as to be compatible with such an AGN. We suggest the possibility that we have a foreground low-mass X-ray binary in quiescence for which the optical counterpart is below detectability, but we do not exclude the presence of a peculiar AGN. In this section we introduce the general X-ray properties both of low-mass X-ray binaries in quiescence and of Seyfert 2 galaxies.

Low-mass X-ray binaries (LMXBs) are composed of a compact object, either a neutron star (NS) or a black hole (BH), accreting matter from a companion star with masses $\leq 1 M_{\odot}$. These objects have typical X-ray luminosities of about 10^{37} – 10^{38} erg s⁻¹ but can drop down to 10^{31} – 10^{33} erg s⁻¹ during quiescence. At high luminosities, the disk is optically thick and geometrically thin, while a hot optically thin, advection-dominated accretion disk is expected for lower luminosities. In the advection-dominated accretion flow (ADAF) model, the radiative efficiency is very low ($\sim 10^{-4}$ – 10^{-3}) and most of the gravitational energy is stored as internal energy and advected towards the compact object with little X-ray emission

(Narayan et al. 1996). A hot quasi-spherical flow is expected in the inner regions producing X-rays (Narayan et al. 1996).

The spectrum of NS-LMXBs in quiescence displays a soft ($kT \sim 0.1$ keV) thermal component (McClintock et al. 2004), probably associated with the stellar surface. Indeed, the observed X-ray luminosities, temperatures, and distance of these systems indicate that the thermal emission comes from a source with a radius of ~ 10 km. On the other hand, for BH-LMXBs in quiescence, no thermal emission has been reported for the 15 sources reported by McClintock & Remillard (2004), since they possess an event horizon. Their spectra are well-fitted by a single power law with a photon index $1.5 < \Gamma < 2.1$ (McClintock & Remillard 2004).

The brightest and best-studied NS-LMXBs in quiescence are Cen X-4 and Aql X-1. Spectral and temporal analysis of these sources during quiescence are reported, for example, by Campana et al. (2004) and Campana & Stella (2003), respectively. Both sources display a thermal component and a power law, as well as short-timescale temporal variability (Campana & Stella 2003; Rutledge et al. 2002). The X-ray Nova V 404 Cygni is, on the other hand, the brightest stellar-mass BH in quiescence. A recent review of its properties is reported by Bradley et al. (2007): the spectrum in quiescence is typically a power law with a $\Gamma \sim 2$, consistent with what is expected for the advection-dominated accretion flow. In the low state, the source is also variable on a short time scale up to a factor of ≥ 20 during a 60 ks observation (Hynes et al. 2004).

We review in this paragraph some of the X-ray properties of Seyfert 2 galaxies. In the unified model for AGN (Antonucci 1993), Seyfert 2 galaxies are identical to the Seyfert 1, but observed at high inclination angle. The nucleus of Seyfert 2 galaxies is therefore not directly visible, and shows much higher column densities than in Seyfert 1 objects, blocking the soft X-ray flux below 2 keV. Risaliti et al. (1999) show that all sources of their sample of Seyfert 2 galaxies have $N_{\text{H}} \geq 10^{22} \text{ cm}^{-2}$, and 75% $N_{\text{H}} \geq 10^{23} \text{ cm}^{-2}$. Their spectra are relatively hard. Moran et al. (2001) show a composite spectrum for 29 Seyfert 2 galaxies observed with ASCA. The continuum is described by a double power law component: an unabsorbed component (energy index $\alpha = 0.84$) dominates below 3 keV, and at higher energies the spectrum is dominated by a heavily absorbed component ($N_{\text{H}} = 3 \times 10^{23} \text{ cm}^{-2}$, $\alpha = 0.46$). An Fe K α line is present at 6.4 keV, with an equivalent width of 420 eV. Small-amplitude, short-timescale variability is observed for Seyfert 2 galaxies, as Awaki et al. (2006). These authors report variation of a factor less than two for 13 Seyfert 2 galaxies observed with XMM-Newton, on a time scale of a few thousand seconds.

The paper is organized as follows. Section 2 describes the XMM-Newton and optical observation and data reduction. In Sect. 3, we present timing and spectral analysis of the source, reserving discussion for Sect. 5.

2. Observation and data reduction

2.1. XMM-Newton data

XMMU J134736.6+173403 was observed serendipitously by XMM-Newton on 2003 June 24 for 64 ksec, the target being Tau Boo. The EPIC-MOS (Turner et al. 2001) were operated in large window mode and EPIC-pn (Strüder et al. 2001) cameras in full frame mode with the thick filter for all instruments. The optical monitor was blocked throughout the full observation. After screening the MOS data for proton flares using standard procedures¹, a total of 60 and 56 ksec of low-background emission remained for the MOS and pn, respectively.

Using the XMM-Newton Software Analysis System (SAS) edetect_chain task, which performs maximum-likelihood source detection, XMMU J134736.6+173403, observed at $\sim 8'$ off-axis, was detected with a maximum likelihood of 3×10^4 . The best-fit coordinates were $\alpha_{J2000} = 13^{\text{h}}47^{\text{m}}36.6$ and $\delta_{J2000} = +17^{\circ}34'02''.8$ with a statistical error of $0''.1$. Since the source is off-axis, however, the uncertainty on the statistical error might be much greater than what is provided by edetect_chain. Combining this error to the systematic shift expected between X-ray data and optical counterpart (typically few arcsec), the total estimated error is about $3''$.

2.2. Optical counterpart spectra

Low-resolution spectroscopy was performed with ISIS at the William Herschel Telescope (ING, La Palma) in the night 2007 June 14 in service mode. The slit was aligned east-west to get the spectra of the galaxy and the AGN in the same frames. The blue and red spectra were acquired simultaneously on two different CCDs, thanks to a dichroic system. Two short exposures of 10 mn each were made to avoid cosmic rays.

The data reduction was performed in a standard way for both the blue and red spectrum parts. We wrote a dedicated and

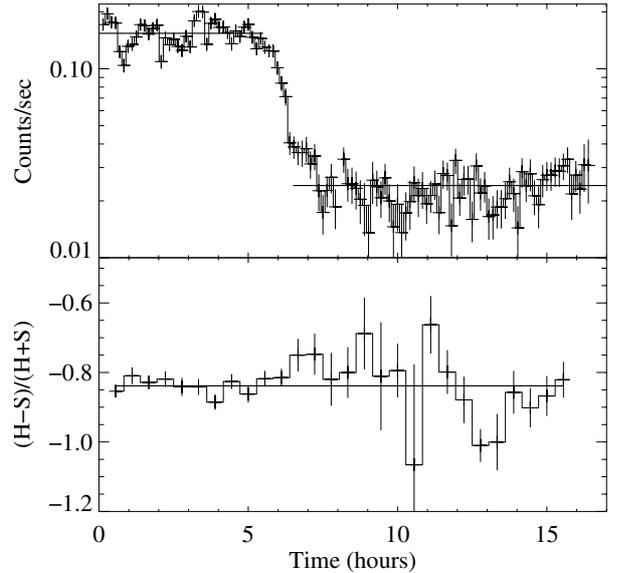


Fig. 1. Mean XMM-Newton EPIC-MOS and pn background-subtracted light curve of XMMU J134736.6+173403 (top), and the corresponding hardness-ratio light curve (bottom). The straight lines show the average of the respective portion of the curves.

semi-automatic MIDAS batches for ISIS long-slit spectra essentially using the LONG package. We first performed the standard reductions, i.e. a bias subtraction, a flat-fielding, and a cosmic rejection. Then we extracted the arcs at the object position, summed them if more than one, and from them we identified some lines and automatically performed a wavelength calibration. All the objects were calibrated in wavelength, corrected from the mean extinction, and calibrated in fluxes with 3 standards stars taken during the night. This instrumental response is very stable in time, regular checks are done monthly. Finally, from these 2D images we extracted the objects with an optimized spectrum extraction batch and a line-fitting to measure the line intensities and the redshift.

3. Timing and spectral analysis of XMMU J134736.6+173403

The combined EPIC-MOS and pn background-subtracted light curve is shown in Fig. 1 (top), with a time bin size of 500 s. The corresponding hardness-ratio light curve is given at the bottom, with a time bin size of 2000 s. Hard and soft bands are given by $H = 2\text{--}10 \text{ keV}$ and $S = 0.2\text{--}2 \text{ keV}$, respectively. Times are given in hours from the start of the observation. Periods of high background at the end of the observation were excluded from the data. The straight lines show the average of the respective portion of the curves. In the first curve the flux drops by a factor of 6.5 within 1 h. We searched for periodicities between 10 s and 6 h in the light curve before the dip where the flux is visibly variable. We found a significant modulation at 1.8 h, but more data are needed to confirm the periodicity. On the other hand, no periodic signal was found in the low state. The fractional root mean-square variability amplitude, defined for example in Vaughan et al. (2003), F_{var} , is $13.7 \pm 1.2\%$ before the flux drop and $13.2 \pm 3.8\%$ after it, showing that the short-term variability is roughly constant during the observation. The hardness-ratio light curve does not vary after the flux decline either.

Figure 2 shows the pn and MOS spectra of the source. The data are binned to have at least 25 counts in each energy bin. The

¹ http://xmm.esac.esa.int/external/xmm_user_support/documentation/sas_usg/USG/

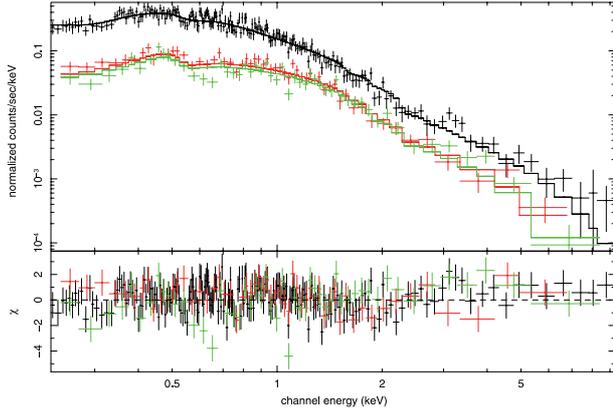


Fig. 2. *XMM-Newton* EPIC-pn and MOS spectra of XMMU J134736.6+173403 in the high state, fitted with an absorbed power law. *Bottom:* residuals expressed in σ .

Table 1. Results of the spectral fits for XMMU J134736.6+173403, using an absorbed power law model (phabs*power, in XSPEC).

	High state	Low state
$N_{\text{H}} (\times 10^{20} \text{ cm}^{-2})$	$4.73^{+0.90}_{-0.83}$	$3.13^{+2.12}_{-1.75}$
Γ	$2.81^{+0.07}_{-0.07}$	$2.72^{+0.19}_{-0.15}$
$F_{0.2-10 \text{ keV}} \times 10^{-12} \text{ (cgs)}$	$1.33^{+0.05}_{-0.05}$	$0.24^{+0.03}_{-0.02}$

source is described well by a power law model, yielding $\chi^2_{\nu}/\nu = 1.18$ in the high state and $\chi^2_{\nu}/\nu = 0.85$ in the low. Adding another component does not improve the fit. The best-fitting parameters of the absorbed power law model are shown in Table 1. Here, N_{H} is the equivalent column density of neutral hydrogen and Γ the photon index. The corresponding 0.2–10 keV flux is shown in the last row. Uncertainties are given at a 90% confidence level.

We searched for X-ray data for the source from other X-ray satellites: no observation has been performed with *Chandra*, and the source is detected in neither the all-sky survey nor in a short (600 s) PSPC observation from the *ROSAT* satellite, where the upper limit flux was of 3×10^{-13} and $8 \times 10^{-13} \text{ erg s}^{-1} \text{ cm}^{-2}$, respectively. The source was, however, visible in a ~ 80 ks *ASCA* observation and reported in the catalogue of *ASCA* sources (Ueda et al. 2005). Their flux and hardness-ratios are compatible to the mean value measured by *XMM-Newton*. We extracted the *ASCA* light curve, and there does not seem to be any flux drop like the one observed in the *XMM-Newton* data, although the count rate is very low (5 cts ks^{-1}).

4. The optical counterparts

Figure 3 shows the optical counterpart image of the X-ray source. The circle is centered on the position of the X-ray source and the radius is $3''$. Both sources are extra-galactic objects with a redshift of $z = 0.045$. From the image and the spectrum, the source on the left is a normal galaxy with some emission lines (not shown in this paper). The object on the right has a spectrum typical of a Seyfert 2 galaxy (narrow H_{β} line and $O_{\text{III}}/H_{\beta} < 3$, see Fig. 4). At this distance the sources are separated by 10 kpc (corresponding to $10''$), suggesting that they form a pair of galaxies, maybe in interaction. The Seyfert 2 galaxy is also associated with a radio source with an intensity of 17.91 mJy at 1.4 GHz (Becker et al. 2003).

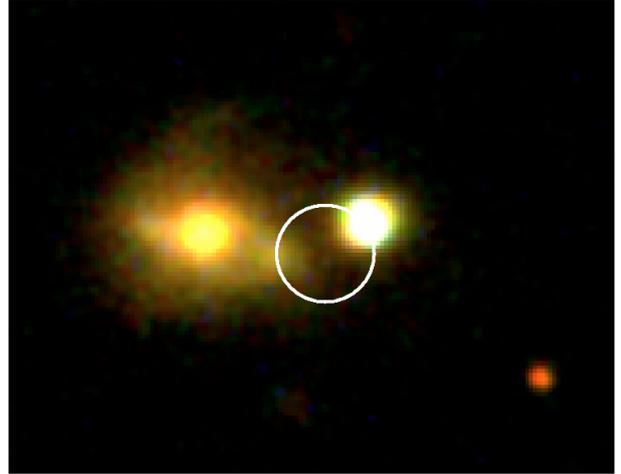


Fig. 3. Color-composite image of the optical counterpart from the Sloan Digital Sky Survey (SDSS). The red color is associated with the *R* band, the green color with the *G* band, and the blue color with the *U* band. The circle is centered on the X-ray source and the radius is $3''$.

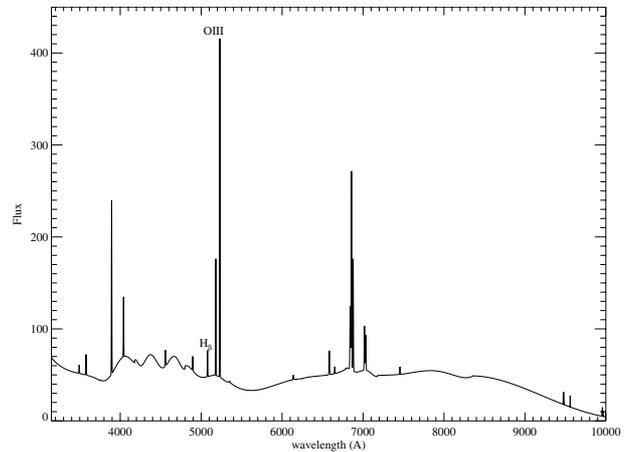


Fig. 4. Optical spectrum of the best optical counterpart of XMMU J134736.6+173403. The width of the H_{β} line, and the ratio between O_{III} over H_{β} (< 3) is consistent with the spectrum of a Seyfert 2 galaxy.

5. Discussion

We reported the discovery of an X-ray source that presents a flux drop of a factor of 6.5 within 1 h and stays in a low state for at least 10 h, suggesting an eclipse. From optical archived images, there are two counterparts: the first an extended normal galaxy and the second a point source. A radio source is also associated to this last one. We got the optical spectrum of the two objects from the ING Telescope in La Palma. The extended object is indeed a redshifted object with emission lines consistent with a normal or slightly active galaxy. Surprisingly, the second object appeared at the same redshift ($z = 0.045$) and has a spectrum typical for a Seyfert 2 galaxy. At that distance the objects are separated by 10 kpc, suggesting that they form a pair of galaxies, maybe in interaction.

Within this context, the X-ray source would be associated with the Seyfert 2 galaxy (generally strong X-ray sources) rather than the normal galaxy. At this distance, the observed luminosity in the 0.2–10 keV band is $6.5 \times 10^{42} \text{ erg s}^{-1}$ in the high state and $1 \times 10^{42} \text{ erg s}^{-1}$ in the low state. This kind of luminosity is typical of Seyfert 2 galaxies. The temporal and spectral properties of the

X-ray object are unexpected, however, for such an AGN. A flux drop of a factor of 6.5 within 1 h has never been reported so far, especially for Seyfert 2 galaxies where the galaxy is observed edge-on. The spectrum, furthermore, is extremely soft for such a galaxy and does not present any $K\alpha$ line at 6.4 keV.

We thus believe that the source is a foreground object within the line-of-sight of this pair of galaxies. The detection limit of the Sloan Digital Sky survey is 23.3 mag and 23.1 mag in the green and red bands, respectively (Willman et al. 2002). Since the size of the Galaxy towards the source is about 5 kpc, the companion must be type M2, $M_V = 9.9$ (Cox 2000) or later to explain the non-detectability. Assuming the object is at 5 kpc, the luminosity is 4×10^{33} erg s⁻¹ and becomes 6×10^{32} erg s⁻¹ if the source is at only 2 kpc (implying therefore a M4/M5 dwarf companion). This is well within the range of what is expected for low-mass X-ray binaries in quiescence and is consistent with these objects having K or M companion stars. We exclude the presence of a cataclysmic variable, since their orbital periods are only a few hours, while here, if the flux drop is associated with an eclipse, it must be at least ~ 1 day. We note that the spectrum is a bit too soft ($\Gamma \sim 2.8$) for an BH-LMXB in quiescence and that a power law continuum is unexpected for a thermal emission from a neutron star surface.

If the X-ray source is a foreground object, then the X-ray flux of the background AGN should also be detected and both are probably confused. A deep *Chandra* observation of the system might represent the only solution for disentangling the LMXB picture or the peculiar AGN scenario. Thanks to its better spatial resolution, we might be able to resolve 2 hypothetical point sources in the first scenario, while one single source, associated with the AGN, should be observed in the second case.

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